CHAPTER 2

BACKGROUND

# 2.1 Introduction

Stratospheric aerosol plays an important role in the global radiative forcing balance by scattering solar irradiation and causing an overall cooling effect that depends on the particle size distribution and the concentration (*Kiehl and Briegleb*, 1993; *Stocker et al.*, 2013). These climate effects are an important and recent focus of research due to the potential contribution of stratospheric aerosol to the so-called global warming hiatus (*Solomon et al.*, 2011; *Haywood et al.*, 2014; *Fyfe et al.*, 2013) and efforts to quantify the variability and trends in the global stratospheric aerosol load are underway with various ground-based and satellite data sets (*Rieger et al.*, 2015; *Ridley et al.*, 2014).

Since its discovery with stratospheric balloon observations (*Junge et al.*, 1961), stratospheric aerosol has been measured with various techniques, although due to the variability of physical composition and particle size, the observations are always limited to some degree and no single measurement technique can fully determine the full range of aerosol properties unambiguously. In-situ balloon observations continue to be used and have provided highly valuable data sets, including most notably the long time series of Optical Particle Counter (OPC) measurements from Laramie, WY (*Deshler et al.*, 2003; 2008; *Kovilakam et al.*, 2015). Aircraft-borne nephelometers (*Beuttell and Brewer*, 1949; *Charlson et al.*, 1969) acquire detailed in-situ measurements, providing, for example, plume composition (*Murphy et al.*, 2014), but are spatially limited to the aircraft track. Ground based lidars have been used to do detailed studies of the extent of volcanic aerosol plumes (*Chazette et al.*, 1995; *Sawamura et al.*, 2012) and provide valuable insight into long term local variability and trends in the aerosol layer. For example, lidar observations were used by *Hofmann et al.* (2009) to first report the observed increase in stratospheric aerosol over approximately the last decade. However, the global distribution, which can only really be obtained with satellite observations, provides invaluable insight into aerosol processes and variability. A good example of this is the use of satellite observations by *Vernier et* (2011b) to determine that the increased stratospheric aerosol load reported by *Hofmann et al.* (2009) was in fact due to a series of relatively minor, mostly tropical, volcanic eruptions.

Several recent studies have highlighted the requirement for continued global stratospheric aerosol observations and especially the need to resolve, both vertically and horizontally, aerosol in the lowermost stratosphere and the upper troposphere. This is the case for tracking the evolution of aerosol from volcanic eruptions, which can have a substantial effect on the aerosol optical depth in the lowermost stratosphere (*Ridley et al.*, 2014; *Andersson et al.*, 2015). Furthering the understanding of the transport of aerosol near and across the tropopause would also benefit from higher spatial and temporal resolution observations. This is evident in the case of volcanic plumes, such as that from Nabro in 2011, the transport and origin of which has been studied extensively and the conclusions are somewhat controversial (*Bourassa et al.*, 2012c; 2013; *Vernier et al.*, 2013; *Fromm et al.*, 2013; 2014; *Fairlie et al.*, 2014; *Clarisse et al.*, 2014). However, this is also the case for the formation of background-level aerosol, particularly in the region of the Asian and North American monsoons, which have been identified as a source of substantial, seasonal and highly structured aerosol formation from precursor tropospheric source gases (*Vernier et al.*, 2011a; *Neely et al.*, 2014; *Thomason and Vernier*, 2013).

Continued stratospheric aerosol observations from space are drastically needed though few, if any, planned missions with such capability are underway. In this work, we present the design and test of a prototype instrument for potential future satellite-based stratospheric aerosol observation. The Aerosol Limb Imager (ALI) concept is a relatively small, low-cost, low-power, passive instrument, suitable for microsatellite deployment with the capability to provide high spatial resolution measurements, both vertically and horizontally, of the visible/NIR aerosol extinction coefficient. The basic idea is to leverage the clear advantages of the limb scatter technique as a passive, and therefore low mass and low power, means to obtain daily global coverage, with a two dimensional hyperspectral imager for filling cross-track observation.

The ALI instrument concept is built around the use of an Acousto-Optic Tunable Filter (AOTF), which is a novel filtering technology that provides the ability to rapidly select the central wavelength of an image with no moving parts. These filters, which have recently been developed as large aperture imaging quality devices, operate very efficiently in the red and near infrared spectral range, which is a well matched spectral range for limb scatter sensitivity to aerosol and cloud (*Rieger et al.*, 2014). Additionally, the spectral bandpass of the AOTF, which is typically between 3-6 nm at these wavelengths, is very suitable for the broadband scattering characteristics of the aerosol limb signal. The two dimensional imaging nature of the design provides the capability to achieve at least sub-kilometer resolution at the tangent point, which is on the order of the scale size of the upper troposphere and lower stratosphere (UTLS) aerosol features mentioned above.

It should be noted that the basic instrument design concept of ALI is very similar to that of the Atmospheric Limb Tracker for the Investigation of the Upcoming Stratosphere (ALTIUS) (*Dekemper et al.*, 2012), which is a Belgian instrument concept from the Belgian Institute for Space Aeronomy (BIRA). ALTIUS is designed to measure limb scattered sunlight; however, it also has solar, stellar, and planetary occultation modes and is scientifically focused on trace gas measurements, particularly for ozone, whereas ALI is optimized for aerosol observation.

# 2.2 Stratospheric Aerosol

In the late 18th century, it was known that atmospheric temperature decreased with altitude and a theory had been raised that at a specific altitude the temperature must eventually go to absolute zero (*Hoinka*, 1997). This lead to a series of balloon campaigns, which were noisy and unreliable due to the technology available (*Hoinka*, 1997), to discover this mysterious altitude in the atmosphere. However, in the late 19th century the technology used in these sounding balloons had improved to a point where the atmospheric temperature could be accurately measured and it was found that at approximately 12 km an inversion point was discovered where the temperature started to increase and thus the tropopause, which separates troposphere and the stratosphere, was discovered. The stratosphere is the region of the atmosphere above the temperature inversion of the troposphere, where atmospheric temperature increases, and the lower bound of the stratosphere is in-between 10 and 16 km from the high latitudes to the tropics (*Andrews*, 1987). The stratosphere, which extends up to approximately 50 km, is thermodynamically stable and fairly dry (*Boucher*, 2015). The characteristic stability of the region limits vertical transport of the stratosphere, leading to long lifetimes, spanning from months to years, for non-volatile species (*Volk et al.,* 1997; *Brasseur and Solomon*, 2005)

The stratosphere undergoes exchange of air with the troposphere though a series of dynamical processes including tropical convection, polar vortices, tropopause folding (*Holton et al.,* 1995) and meridional circulation within the stratosphere is dominated by the slow Brewer-Dobson circulation, although zonal circulation is much faster (*Plumb and Eluszkiewicz*, 1999). Some chemicals can cross the tropopause thermal barrier into the stratosphere which allows for chemicals reactions to occur. One such reaction forms stratospheric sulfate aerosol, discovered by *Junge et al.* (1961) through stratospheric balloon sondes. Sulfate aerosols are droplets of hydrated sulfuric acid (H­2­SO­4) formed from the oxidation of sulfur compounds, primary OCS and SO2 (*Brock et al.*, 1995). This stable layer of aerosol exists in the stratosphere from the altitude of the tropopause to approximately 30 km.

## 2.2.1 Aerosol Sources

The sources for sulfate aerosol are primarily produced in the troposphere and formed though both natural and anthropogenic processes. One primary source of this sulfur is OCS which originates from marine processes, biomass burnings, and industry (*Kettle et al.*, 2002; *Notholt et al.*, 2003). OCS has a long lifetime in the troposphere and low solubility allowing for a significant portion to reach the stratosphere where some of it oxidizes and hydrates to form sulfate aerosol (H2SO4) and adds to the background aerosol layer (*Crutzen*, 1976).

Another source of sulfur is SO2, which originates in the troposphere through industry from the burning of fossil fuels. Sulfur dioxide (SO2) has a short lifetime in the troposphere, and its concentration varies regionally. This sulfur hydrates into sulfate aerosol, which can enter the stratosphere though atmospheric processes (*Thomason and Peter*, 2006). A second source of SO2 is from volcanic eruptions; which, although highly variable in location and time, can inject a large amount of sulfur directly into the stratosphere which undergoes hydroxyl chemistry and is converted into sulfate aerosol. Volcanos can inject such large amounts of sulfur that they in fact dominate the stratospheric aerosol layer perturbing the background levels. Examples of this perturbation of the aerosol layer were noted during the volcanic eruptions of El Chichon in 1982 (12-20 Tg of sulfur) (*McCormick and Swissler*, 1983; *Hofmann and Rosen*, 1983) and Mount Pinatubo in 1991 (20-30 Tg of sulfur) (*McCormick and Veiga*, 1992). However, after the Mount Pinatubo eruption a volcanically quiescent period occurred where aerosol layers returned to background. Following this period, a series of smaller volcanic eruptions have increased the background aerosol layer in the amount of 4-7% per year from 2000 to 2009 (*Vernier et al.*, 2011b) and has continued to the present day. Several examples of these small volcanic eruptions include Kasatochi (1.2-2.2 Tg) (*Prata et al.*, 2010), Nabro (1.0- 1.5 Tg) (*Clarisse et al.*, 2016), Calbuco (0.2 -0.5 Tg), and Kelut (0.1-0.3 Tg) (*Carn et al*., 2016).

## 2.2.2 Aerosol Microphysics

These sulfur sources undergo a series of reactions to be converted into H2SO­4 and have been found to form spherical, liquid droplets of aerosol that consist of approximately 25% H2O and 75% H2SO4 (*Rossen*, 1971; Wang *et al.*, 1989). These aerosol particles form droplets of various sizes on the order of 0.05 to 1.0 µm depending on the various contributions and stages of the processes of nucleation, evaporation and condensation (*Junge et al.* 1961; *Brock et al.*, 1995; *Bingen et al.*, 2004). A log-normal distribution is often used to approximate the particle size of aerosol in the form of

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| --- | --- |
|  | (2.1) |

where is the aerosol number density, is the mode radius, and is the mode width (*Jäger and Hofmann,* 1991; *Hamill et al.* 1997). In this case, the particle radii are distributed normally over the logarithm of the radius. For a typical non-volcanic background aerosol, with a mode radius and mode width of 0.08 µm and 1.6 respectively (*Thomason*, 1991), the log-normal distribution is shown in Figure 2-1.

Optical Particle Counters (OPC) have been used on board stratospheric balloon flights from Laramie, Wyoming over the past 40 years to measure particle sizes in bins between 0.15 to 2.0 µm (*Deshler et al*., 2003). These measurements provide a valuable if somewhat unique long term set of size-resolved measurements of sulfate aerosol. These particle size distributions can be considered to be primarily unimodal, known as a fine mode, over non-volcanic periods and can be used as an acceptable distribution to approximate background periods. But during volcanic episodes, a bimodal log-normal distribution of aerosol particles, which includes a coarse mode, is more representative of stratospheric aerosol (*Deshler et al.*, 2003; 2008; *Kovilakam et al.*, 2015). The coarse mode has larger particles than the fine mode and complicates the determination of aerosol microphysical parameters since the number of required parameters has increased to six: a number density for both the fine and coarse mode, two mode radii, and two mode widths. Figure 5 from *Deshler et al.* (2003), recreated in Figure 2-2, demonstrates two bimodal particle size distributions from balloon OPC. The first distribution is from a volcanic period in 1993 after the Mount Pinatubo eruption and another from a background period in 1999. It should be noted that even though a bimodal distribution is found for the background case in Figure 2-2, the number density of the coarse mode is very small and can generally be ignored in non-volcanic periods so a unimodal approximation is sufficient.



Figure 2-1: Sample log-normal distribution for typical non-volcanic stratospheric aerosol.



**Figure 2-2**: Bimodal particle size distributions fits from OPC. (a) Distributions from a volcanic period after the Mount Pinatubo eruption recorded in 1993. (b) Distributions from a background aerosol period recorded in 1999. Both of the aerosol distribution measurement are from 20 km altitude with the solid line being the fine mode and the dashed line is the coarse mode. Figure is recreated from Figure 5 of *Deshler et al.* (2003).

## 2.2.3 Climate Effects

Stratospheric aerosol can have several effects on the climate of the planet, and particularly due to the volcanic contribution, there is a large amount of uncertainty in the overall effect (*Solomon et al.*, 2007). Through the so-called “direct effect”, aerosol particles scatter incoming visible solar radiation away from earth increasing the albedo causing a cooling of the surface of the planet (*Lacis et al.*, 1992). The albedo is the amount of incoming solar irradiance that is reflected back to space. A secondary direct effect from aerosols, which is highly dependent on aerosol particle size distribution, is a greenhouse-like effect that is caused by scattering of infrared radiation emitted from the earth’s surface (*Kiehl and Briegleb*, 1993). Aerosol also introduces a so-called “indirect effect” to the radiative balance. This is also known as the cloud albedo effect. This is caused by condensation of water onto existing aerosol particles. These become cloud condensation nuclei and stimulate cloud formation, which leads to an increase of the planetary albedo, which then also contributes to cooling the planet's surface. These types of cloud forming particles also tend to increase the overall lifetime of the cloud, increasing the overall cloud coverage and thus increasing the planetary albedo (*Charlson et al.*, 1992). Overall, the cooling effect of the aerosol particles dominates the warming effect and cools the surface of the planet (*Solomon et al.*, 2011).

Background aerosol periods result in relatively small cooling from stratospheric aerosols but this can greatly change during periods of volcanic activity where the layer can be greatly perturbed. After the eruption of Mount Pinatubo in 1991 the sulfate aerosol load was increased by 5 to 10 fold causing cooling of the lower atmosphere by 0.5oC (*McCormick et al.* 1995; *Soden et al.*, 2002) and 0.1 to 0.3oC on the surface (*Thompson et al.*, 2009; *Canty et al.*, 2013). And the surface temperatures did not return to pre-Pinatubo level until approximately three years after the eruption as the atmosphere filtered out the additional aerosol (*Hansen at al.*, 1996). More recently, a series of small to moderate volcanic eruptions have increased the background stratospheric aerosol layer (*Vernier et al.*, 2011b). This additional volcanic aerosol load has been proposed to be linked to a larger cooling effect, known as the global warming hiatus (*Solomon et al.*, 2011; *Haywood et al.*, 2014; *Fyfe et al.*, 2013). Even with the current knowledge of the cooling effect of aerosol on climate, there is a large uncertainty to the magnitude of the cooling effect due to the unknowns in aerosol microphysical parameters. To be able to fully resolve the effect of aerosol, new instrumentation is required that has the capabilities to not only determine aerosol concentration in higher resolutions, vertically, globally, and temporally, but also must be able to determine some form of particle size distribution information. This additional information will reduce the uncertainty of the effect of aerosol on climate.

# 2.3 Aerosol Measurements

Two fundamental methods are used to measure aerosol concentrations within the atmosphere. Two such methods are ground based and in-situ measurements which give good detail and information about a specific localized area. However, these measurements are limited in scope as they do not have global coverage that is inherent in satellite instrumentation. Both ground-based instruments and satellites have important roles in monitoring the planet’s aerosol content and each of these methods have inherent advantages and disadvantages. An overview is given on some of the common methods to determine aerosol and why using different methods helps to increase the overall accuracy and precision of merged data sets.

## 2.3.1 Balloon- and Aircraft-Based Measurements

In-situ measurement are typically performed using balloon- or aircraft-based platforms. In-situ balloon instruments directly measure aerosol particles during the assent and can determine the particle size distributions. The OPC is an active instrument that uses light source internal to the device to optically count aerosol particles. This type of instrument has been launched from Laramie, Wyoming since 1971, and has successfully measured aerosol mixing ratio and particle size distributions (*Deshler et al.*, 2003; 2008; *Kovilakam et al.*, 2015).

Another type of balloon instrument uses a passive light source, such as the sun, moon, or stars, to determine aerosol extinction. Some instruments that use this type of technology are the Absorption par les Minoritaires Ozone et NOx (AMON) from 1992 to 2003 and Spectroscopie d’Absorption Lunaire pour l’Observation des Minoritaires Ozone et NOx - Nacelle 2 (SALOMON-N2) from 2007 onwards which use starlight and moon light respectively (*Berthet et al.*, 2002).

In-situ measurements yield high quality direct measurements of aerosol from the altitude that the instrument is currently situated unlike remote sensing applications from satellites. Furthermore, passive instruments from balloon and aircraft platforms also acquire quality aerosol measurements. However, these types of instruments only give aerosol parameters from a very localized region, like the Laramie, Wyoming OPC, or have very few flights, for example AMON which had a total of six stratospheric balloon flights, three mid-latitude northern and three high-latitude northern flights. In order to achieve full global coverage satellite remote sensing instruments are needed to fill the spatial gap.

## 2.3.2 Occultation

Satellite instrumentation capable of remote sensing stratospheric aerosol has been in use since the 1970s, beginning with limb sounding solar occultation measurements, and its operational geometry can be seen in Figure 2-3. Solar occultation measurements have provided a reliable, accurate and essentially continuous long term record of vertically resolved aerosol extinction coefficient measurements, mostly from the series of Stratospheric Aerosol and Gas Experiment (SAGE) instruments starting in with SAGE I in 1979 (*Russell and McCormick*, 1989; *Thomason and Taha*, 2003). These SAGE measurements, which have a vertical resolution of approximately 1 km, have generally compared well with ground-based and in-situ measurements, although there are challenges associated with determining microphysical parameters and comparison between instruments can be challenging (*Russell and McCormick*, 1989; *Kovilakam et al.*, 2015). However, solar occultation is generally a robust and stable technique as it directly measures atmospheric optical depth, along with the exo-atmospheric solar spectrum with each scan, allowing for straightforward retrieval of aerosol extinction coefficients (*Damadeo et al*, 2013). The major drawback to occultation instruments is that a sunrise or sunset event is required to perform a measurement limiting is the number of scans per day to 16-48 measurements depending on the orbit.



**Figure 2-3**: An occultation instrument monitoring the atmosphere by scanning the atmosphere by looking directly at the sun.

The SAGE III mission came to an end in 2006 and the occultation measurements have continued from the currently operational MAESTRO and ACE-Imager instruments on SciSat (*McElroy et al.*, 2007; *Gilbert et al.*, 2007) and have had some success producing stratospheric aerosol extinction products (*Vanhellemont et al.*, 2008; *Sioris et al.*, 2010). Furthermore, a manifestation of SAGE III is planned for deployment on the International Space Station in 2016 (*Cisewski et al.*, 2014).

## 2.3.3 Lidar

A method known as lidar can determine atmospheric parameters through the pulsing of a laser and measuring of the intensity of the backscattered laser light at different wavelengths and polarizations. Lidar has been used at ground based facilities to measure aerosol layers dating back to the 1960s (*Fiocco and Grams*, 1964) and are still used today. More recently lidar instruments have been used on satellite missions including the Ice, Cloud, and land Elevation Satellite (ICESat) from 2002 to 2010 (*Schutz et al.*, 2005) and Cloud-Aerosol Lidar and Infrared Pathfinder Satellite Observations (CALIPSO) which launched in 2006 (*Winker et al.*, 2007). More recently Cloud Aerosol Transport System (CATS) (*Chuang et al.*, 2013) has been mounted on the international space station in 2015 with a three year planned mission. Traditional lidar instruments have looked in the nadir direction (either straight down or up) however some instruments have looked slightly off-nadir, both can be seen in Figure 2-4. Lidar measurements have been used to determine aerosol plumbs from volcanos (*Chazette et al.*, 1995; *Sawamura et al.*, 2012) and long term trends (*Hofmann et al.*, 2009).

CALIPSO is a joint mission developed between the National Aeronautics and Space Administration (NASA) and the Centre National d'Etudes Spatiales (CNES) of the United States and France respectively. It uses a two wavelength polarized lidar system to achieve high resolution aerosol and cloud retrievals along the satellite's orbital track with global coverage from 82oS to 82oN (*Young and Vaughan*, 2009). CALIPSO nominally measures backscatter profiles approximately every 300 m along track with approximately 200 m vertical resolution. However, the stratospheric backscatter signal is weak and requires averaging of only the night time measurements over several days and typically yielding resolutions of 0.5 km vertically and 500 km horizontally (*Vernier et al*, 2011b). Additionally, the uncertainty in the calibration with respect to the molecular background is on the order of the stratospheric aerosol signal and leads to a potential bias in the stratospheric measurements (*Rogers et al.*, 2011). CALIPSO was launched in 2006 and although it is presently still operational, it is also operating beyond its design lifetime.



Figure 2-4: Lidar instrument showing a measurements in both the nadir and off-nadir lines of sight.

## 2.3.4 Limb Scatter

The limb scatter technique measures light that is scattered into the line of sight of the instrument from atmospheric interactions. These scattering interactions can undergo either single or multiple scattering events. Single scatter is when light from the sun interacts with a particle in the atmosphere and scatters it directly into the line of sight of the instrument. Multiple scatter is when the photon of light undergoes several scattering events before entering the line of sight including scattering off of multiple particles in the atmosphere or scattering off of the ground. These events can occur any number of times before entering the instrument. The geometry for the limb scatter technique can be seen in Figure 2-5 and demonstrates the defining angles for this method. All angles are defined from the tangent point, which is the point where the distance between the line of sight and the surface of the earth is minimized, represented by the black dot. The Solar Zenith Angle (SZA) is the angle between the local vertical and the direction of the sun; the Solar Scattering Angle (SSA) is the defined to be between the direction of the sun and the line of sight and the Solar Azimuth Angle (SAA) is the angle between the projection of the sun on the plane of the line of sight and the line of sight. These angles can also be seen on Figure 2-5.



Figure 2-5: Limb scattering geometry measurement for an instrument where single and multiple scattering events occur.

The limb scatter method yields relatively good vertical resolution and allows for measurements to be taken during any daylight period with good Signal to Noise Ratio (SNR), however it requires the use of a complex forward model to calculate the scattering events along with some a priori knowledge of the aerosol scattering cross section in order to retrieve the extinction coefficient profile. The model needs to accurately determine the effect of multiple scatter since it consists of 10-50% of the measured signal depending of the specific geometry and wavelength (*Oikarinen et al.*, 1999). Furthermore, due to the complex nature of the problem; a large amount of computational time and memory is required for an accurate calculation.

The first use of limb scatter was on the Solar Mesosphere Explorer (SME) (*Barth et al.*, 1983) to measure mesospheric ozone profiles in 1981. Much later, other limb scatter instruments were launched into low earth orbit that had the capability to determine aerosol extinction including the Optical Spectrograph and InfraRed Imaging System (OSIRIS) launched on the Odin satellite in 2001 (*Llewellyn et al.*, 2004) and the SCanning Imaging Absorption spectroMeter for Atmospheric CHartographY (SCIAMACHY) on Envisat launched in 2002 (*Bovensmann et al.*, 1999). Both of these instruments are scanning grating spectrometers which can gather a single tangent point and scans the atmosphere to complete a vertical profile.

The OSIRIS version 5.07 data product provides 750 nm extinction profiles at approximately 2 km vertical resolution (*Bourassa et al.*, 2007) and has been shown to agree relatively well, generally within 15%, with SAGE II and SAGE III occultation measurements (*Bourassa et al.*, 2012b; *Rieger et al.*, 2015). The SCIAMACHY instrument uses a retrieval technique essentially similar to OSIRIS to retrieve aerosol profiles at 750 nm with approximately 3 km vertical resolution (*Ernst et al.*, 2012; *von Savigny et al.*, 2015). However SCIAMACHY observations ceased with the demise of Envisat in 2012 and although OSIRIS continues to operate, it is now in the sixteenth year of a mission designed for two years.

The most recently launched limb scatter instrument is the Ozone Mapping Profiler Suite Limb Profiler (OMPS-LP) on the Suomi-NPP satellite in 2011. Although similar in spectral range and vertical resolution to OSIRIS, OMPS-LP is an imaging spectrometer that vertically images the limb in a single measurement. The imaging capability of OMPS-LP provides a decrease in the time required to obtain a limb profile and so increases the along track sampling. Recent work on the feasibility of aerosol retrieval from OMPS-LP measurements show promising results (*Rault and Loughman*, 2013).

An instrument that is currently under development is ALTIUS (*Dekemper et al.*, 2012), which is a Belgian instrument concept from the Belgium Institute for Space Aeronomy. ALTIUS is designed to image limb scattered sunlight, both vertically and horizontally across the track through the use of the Acousto-Optic Tunable Filter (AOTF) technology (see section 3.1) and additionally has solar, stellar, and planetary occultation modes. ALTIUS is scientifically focused on trace gas measurements, particularly for ozone and the instrument has three channels, each channel with a separate AOTF, measuring wavelengths from 250-2000 nm which could eventually be used for aerosol extinction measurements.

The limb scatter technique is selected for the ALI instrument. ALI also uses an AOTF to spectrally image the filtered scattered signal. This technique was selected due to the relatively lax measurement requirement of only requiring sunlit atmosphere to be able to record a high number of quality measurements. Additionally, the imaging nature of the AOTF will allow quick measurements that are used to retrieve vertical profiles of aerosol extinction.

# 2.4 Radiative Transfer

To use the limb scatter technique to determine aerosol extinction and particle size information, an understanding of radiative transfer is required. However, modeling the complex scattering interactions of light within the atmosphere is difficult. In this section, an overview of scalar radiative transfer is performed, followed by the necessary modifications to form polarized radiative transfer equations. A description of scattering interactions important to aerosols is developed. Finally, an overview of the SASKTRAN radiative transfer model used within this work is provided.

## 2.4.1 Scalar Radiative Transfer

The scattering and absorption processes in the atmosphere are non-trivial and an adequate method to model this interaction is needed. The following presents a derivation of radiative transfer equations for the atmosphere with scalar radiance which does not account for polarization. In order to accurately discuss radiative transfer, a coordinate system must first be defined. If we assume that a ray of light, , is propagating in a given direction, , and starts at a location, , with the initial position of , then the position of the ray along the path direction can be completely defined by its path length, . The basis of path length is used to define the radiative transfer equations.

The fundamental theory for radiative transfer is known as Beer-Lamberts law. The law describes the change in intensity or radiance of light, , as it interacts with a thin layer of space or atmosphere, . The thin layer has particles which affect the attenuation of the light which is dependent on the number of particles, n, and the particle cross section, . If there are several different particles, the attenuation is a summation of the number densities and cross sections. The Beer-Lambert Law gives the following form

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| --- | --- |
|  | (2.2) |

The extinction of the particles is a measure of the loss of light over a given distance and is defined as

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| --- | --- |
|  | (2.3) |

Integrating Equation 2.2 forms the following result

|  |  |
| --- | --- |
|  | (2.4) |

The optical depth, , is defined as the extinction over the path length simplifying Equation 2.4 to

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|  | (2.5) |

The above gives the radiance at point after it has gone through attenuation from .

Although this form of the Beer-Lambert’s Law is useful for describing the loss of light through scattering or absorbing from an initial source though a medium, the atmosphere also has incoming light that is scattered into the line of sight from other directions or emitted from particles. To account for this additional source of light a source term, , is added to Equation 2.2 to yield

|  |  |
| --- | --- |
|  | (2.6) |

Using the fact that the change in optical depth is defined as

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|  | (2.7) |

Equation 2.6 is rearrange to

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|  | (2.8) |

Using the following derivative

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|  | (2.9) |

and substituting it into Equation 2.8 yields

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|  | (2.10) |

This form can now be integrated over the optical depth giving

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| --- | --- |
|  | (2.11) |

Selecting the reference point at the observer to be and converting the equation back to path lengths yields

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| --- | --- |
|  | (2.12) |

which gives the radiance as seen from an observer at a point, , along the line of sight.

With the full form of the radiative transfer equation, the source term must be determined. In the atmosphere there are three sources of additional radiation, blackbody emissions, photochemical reactions, and scattered light. For wavelengths from the visible to the near infrared (*i.e.* wavelengths less than 2 µm) there is little contributions from thermal emissions. Furthermore, as long as wavelengths where photochemical reactions occur are avoided this source term can also be ignored. This leaves the only significant source of light to be added into the line of sight to be from scattered sunlight. The source term for scattered sunlight is given by

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|  | (2.13) |

The diffuse radiance is given by and is the radiation scattered into the line of sight from all directions. The phase function, , described the probability that a ray of light is scattered from a direction, , into the line of sight propagation direction, . The scattering angle, , is defined as

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| --- | --- |
|  | (2.14) |

Lastly, is the extinction only caused by scattering and not absorption. The term only allows the fraction of particles that scatter radiation, and not absorb it, to contribute to the source term.

As a final note, the calculation of the diffuse radiance is what makes this a computationally heavy problem. To completely solve for the diffuse radiance, the radiance at every point in the atmosphere must be determined. Furthermore, the light can be scattered multiple times in the atmosphere, requiring a diffuse radiance for each order of scatter. Each successive scattering adds smaller contributions to the final radiance at the observer. Through this iterative process the full multiple scatter solution to the radiative transfer equation is

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| --- | --- |
|  | (2.15) |

The multiple scatter term is calculated until the contribution is sufficiently small to be negligible.

## 2.4.2 Vector Radiative Transfer

The scalar radiative transfer equation works well for systems that do not measure polarized light as the effect of polarization on the scalar radiance is small. However, for instruments that measure polarized light, a vector radiative transfer equation is required. Before polarization can even be discussed, a method to quantify polarization must be defined which is the Stokes vectors. The Stokes vectors are given as

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| --- | --- |
|  | (2.16) |

where is the scalar or total radiance, is the difference between horizontal polarization to vertical polarization, is the difference between +45o diagonal polarization to -45o polarization, and is the difference between the counter clockwise circular polarization to clockwise polarization (*Bickel and Bailey*, 1985). Using a reference frame where the local x-axis is defined to be the horizontal polarization leads to the following definition for the Stokes vector

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|  | (2.17) |

The electric field aligned with the x and y-axis are and respectively, the star is the complex conjugate, and and are the “real part of” and “imaginary part of” respectively. The degree of polarization can be determined with the Stokes vectors. If the equality holds true then the light is fully polarized, otherwise it is only partially polarized if .

With the addition of polarization, the radiative transfer and source term equations (Equations 2.12 and 2.13) need to be rewritten with polarization state included. The polarized radiative transfer equation are

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|  | (2.18) |
|  | (2.19) |

which are the vector radiative transfer and source term equations respectively (*Mishchenko et al.*, 2002). With polarization a scattering reference frame is defined and incoming radiance is rotated into the scattering frame multiplied by the scattering matrix then returned to the original propagation frame. The rotation matrix is defined as

|  |  |
| --- | --- |
|  | (2.20) |

where is the angle between the propagation and scattering reference frame. The radiance and the source terms are now Stokes vectors in 4 by 1 matrices and the scattering matrix, , is a 4 by 4 tensor that is related the probability of the incoming light to be scattered in the propagation direction with a specific polarization. As a note, the operation of is commonly referred to as the phase matrix. The polarization equation adds extra computation and memory consumption since the polarization must be computed at each scattering in the radiative transfer equation, which is nontrivial, and stored in memory, which is four times the size of a standard scalar radiance calculations.

With the complete vector polarized radiative transfer expression the two scattering interactions that pertain to determining aerosol will be described. The first interaction is Rayleigh scattering which defines the scattering of the background atmosphere, and Mie scattering which determines how incoming light scatters from aerosol particles.

## 2.4.3 Rayleigh Scattering

Rayleigh scatter is the scattering performed by the molecular background atmosphere, *i.e*. by molecules of the air. The first calculation of molecular atmospheric scattering cross sections was by Lord Rayleigh where he assumed the molecules were dielectric spheres with radii much less than the wavelength of the light. Later, the King correction was added to the Rayleigh scattering cross section, , to yield the following expression

|  |  |
| --- | --- |
|  | (2.21) |

which is highly dependent on wavelength, , which is in cm. The parameters and are the volume polarizability, in cm-3, and the depolarization ratio, which is unitless (*Sneep and Ubachs*, 2005).

The other important quantity for scattering is the scattering matrix. For Rayleigh scattering, the vector model scattering matrix is given by the Rayleigh-Gans approximation (*Mishchenko et al.*, 2002)

|  |  |
| --- | --- |
|  | (2.22) |

Each component of the scattering matrix itself is smooth which allows for easy and accurate calculation for Rayleigh scattering.

## 2.4.4 Mie Scattering

For larger particles, like sulfate aerosol, Rayleigh scattering no longer holds since the wavelengths are on the order of the size of the particles and Mie scattering must be used. *Mie* (1908) solved Maxwell’s equations in a general sense with a solution using spherical Bessel and Henkel functions. Only the fundamental Mie scatter equation is presented here but a full derivation of Mie scatter can be found in *van de Hulst* (1957). The scattering cross section from Mie theory is given by

|  |  |
| --- | --- |
|  | (2.23) |

where is the wavenumber, r is the particle radius and the coefficients and are given by

|  |  |
| --- | --- |
|  | (2.24) |
|  | (2.25) |

The index of refraction of the particle is given by , and and are the normalized half-integer order Bessel functions of the first kind and Henkel functions of the second kind respectively. The scattering matrix for Mie scatter for a vector solution has the following form (*Hansen and Travis*, 1974)

|  |  |
| --- | --- |
|  | (2.26) |

The terms in the scattering matrix, and , are known as the amplitude functions and are given by

|  |  |
| --- | --- |
|  | (2.27) |
|  | (2.28) |

where are the Legendre polynomials.

In the atmosphere, various particle sizes occur and a log-normal distribution (Equation 2.1) is assumed for aerosols. In order to determine effective scattering cross-section, a weighted average over the particle radius is performed

|  |  |
| --- | --- |
|  | (2.29) |

The weighted average is similarly performed to determine the effective phase matrix for a particle size distribution.

It should be noted that although the theory is well founded to calculate the Mie scattering cross sections and phase matrices, in practice it is computationally intensive since the terms consist of infinite sums of Bessel functions. Work done by *Wiscombe* (1980) has allowed for effective computation of the Mie scattering coefficients, which has been implemented into the SASKTRAN radiative transfer engine.

## 2.4.5 SASKTRAN Radiative Transfer Model

The SASKTRAN radiative transfer was first developed to solve the scalar radiative transfer equation in a fully spherical atmosphere for both single and multiple scatter with a one dimensional atmosphere, *i.e*. considering the variation in altitude only (*Bourassa et al*., 2008). The first source term, , is the sunlight from the sun attenuated and scattered into the instrument line of sight, and it is assumed the incoming solar irradiance encounters the earth in parallel randomly polarized rays. To include higher order terms, a successive orders method is used to simulate second, third and higher orders of scattering within the atmospheric model. Another important assumption in the SASKTRAN model is that the ground reflection is assumed to have a depolarizing Lambertian distribution, which will evenly distribute the incoming radiance evenly in all outgoing directions with the efficiency of the planetary albedo.

Recent upgrades have been performed on SASKTRAN and have led to a new engine known as SASKTRAN High Resolution or SASKTRAN-HR (*Zawada et al.*, 2015) which has expanded the model to be able to perform radiative transfer calculations with a fully three dimensional atmosphere. This update allows the model to vary the atmospheric concentrations in not just the vertical direction, like the original SASKTRAN, but in both of the horizontal geometries (*i.e.* latitude and longitude), allowing for true variances observed in the atmosphere.

The most important update to the SASKTRAN-HR model for this work is the addition of the ability to calculate the vector or polarized radiances (*Dueck et al.*, 2016). Using the vector model allows for SASKTRAN to compute the Stokes vectors in the reference frame of the model, which can be rotated into any desired frame of reference through the use of a provided basis by SASKTRAN-HR. The polarization output from SASKTRAN-HR preforms polarized calculations up to an arbitrary order scattering interaction, all scattering past this arbitrary scattering are assumed to be scalar.

# 2.5 ALI Prototype and Stratospheric Balloon Flight

The work presented here is focused on the design of the Aerosol Limb Imager (ALI), a prototype polarized limb scatter instrument to spectrally image the atmosphere. The central feature of ALI is the use of the novel technology known as an AOTF which has the ability to rapidly select the central filtered wavelength with no moving parts. AOTFs have recently been developed with large apertures and high quality crystals allowing for imaging capabilities permits AOTF use in new applications, such as the ALI instrument with the ability to achieve sub-kilometer resolution, both horizontally and vertically. The AOTF technology operates efficiently in the red to near infrared which is well-matched for limb scatter sensitivity to aerosol and clouds (*Rieger et al.*, 2014) and has a typical spectral bandpass of 3‑6 nm which is ideal for the broadband scattering characteristics of aerosol.

However, the use of an AOTF has specific optical requirements which complicates the optical design on the system. A large focus of this work is the designing and testing of the optical system to verify excellent performance with the AOTF. Once the optical design is finalized, calibration and opto-mechanical design is undergone to prepare ALI for its stratospheric balloon flight from Timmins, Ontario in 2014 and is the focus of Chapter 3. Additionally, since ALI inherently measures polarized radiance due to the nature of the AOTF, special consideration had to be given to the orientation of the AOTF within the system. Measuring a linear polarized light instead of measuring the total radiance, like current generation instruments, changes the sensitivity of the instrument to measure aerosol. To verify good sensitivity to aerosol during the campaign a study was performed, contained in Chapter 4, to determine the advantages and disadvantages to measuring linear polarized light instead of the total radiance. Finally, Chapter 5 is focused on the stratospheric balloon launch and post-flight analysis including raw image conversion and the retrieving of aerosol. The retrieved aerosol is compared to OSIRIS measurements to verify the quality of the measurements recorded by ALI.