CHAPTER 2

BACKGROUND

# 2.1 Introduction

Stratospheric aerosol plays an important role in the global radiative forcing balance by scattering solar irradiation and causing an overall cooling effect that depends on the particle size distribution and the concentration (*Kiehl and Briegleb*, 1993; *Stocker et al.*, 2013). These climate effects are an important and recent focus of research due to the potential contribution of stratospheric aerosol to the so-called global warming hiatus (*Solomon et al.*, 2011; *Haywood et al.*, 2014; *Fyfe et al.*, 2013) and efforts to quantify the variability and trends in the global stratospheric aerosol load are underway with various ground-based and satellite data sets (*Rieger et al.*, 2015; *Ridley et al.*, 2014).

Since its discovery with stratospheric balloon observations (*Junge et al.*, 1961), stratospheric aerosol has been measured with various techniques, although due to the variability of physical composition and particle size, the observations are always limited to some degree. In-situ balloon observations continue to be used and have provided highly valuable data sets, including most notably the long time series of optical particle counter measurements from Laramie, WY (*Deshler et al.*, 2003; 2008; *Kovilakam et al.*, 2015). Aircraft-borne nephelometers (*Beuttell and Brewer*, 1949; *Charlson et al.*, 1969) acquire detailed in-situ measurements, providing, for example, plume composition (*Murphy et al.*, 2014), but are spatially limited to the aircraft track. Ground based lidars have been used to do detailed studies of the extent of volcanic aerosol plumes (*Chazette et al.*, 1995; *Sawamura et al.*, 2012) and provide valuable insight into long term local variability and trends in the aerosol layer. For example, lidar observations were used by *Hofmann et al.* (2009) to first report the observed increase in stratospheric aerosol over approximately the last decade. However, the global distribution, which can only really be obtained with satellite observations, provides invaluable insight into aerosol processes and variability. A good example of this is the use of satellite observations by *Vernier et al.* (2011b) to determine that the increased stratospheric aerosol load reported by *Hofmann et al.* (2009) was in fact due to a series of relatively minor, mostly tropical, volcanic eruptions.

Several recent studies have highlighted the requirement for continued global stratospheric aerosol observations and especially the need to resolve, both vertically and horizontally, aerosol in the lowermost stratosphere and the upper troposphere. This is the case for tracking the evolution of aerosol from volcanic eruptions, which can have a substantial effect on the aerosol optical depth in the lowermost stratosphere (*Ridley et al.*, 2014; *Andersson et al.*, 2015). Furthering the understanding of the transport of aerosol near and across the tropopause would also benefit from higher spatial and temporal resolution observations. This is evident in the case of volcanic plumes, such as that from Nabro in 2011, the transport and origin of which has been studied extensively and somewhat controversially (*Bourassa et al.*, 2012c; 2013; *Vernier et al.*, 2013; *Fromm et al.*, 2013; 2014; *Fairlie et al.*, 2014; *Clarisse et al.*, 2014). However, this is also the case for the formation of background-level aerosol, particularly in the region of the Asian and North American monsoons, which have been identified as a source of substantial, seasonal and highly structured aerosol formation from precursor tropospheric source gases (*Vernier et al.*, 2011a; *Neely et al.*, 2014; *Thomason and Vernier*, 2013).

Continued stratospheric aerosol observations from space are drastically needed though few, if any, planned missions with such capability are underway. In this work, we present the design and test of a prototype instrument for potential future satellite-based stratospheric aerosol observation. The Aerosol Limb Imager (ALI) concept is a relatively small, low-cost, low-power, passive instrument, suitable for microsatellite deployment with the capability to provide high spatial resolution measurements, both vertically and horizontally, of the visible/NIR aerosol extinction coefficient. The basic idea is to leverage the clear advantages of the limb scatter technique as a passive, and therefore low mass and low power, means to obtain daily global coverage, with a two dimensional hyperspectral imager for filling cross-track observation.

The ALI instrument concept is built around the use of an Acousto-Optic Tunable Filter (AOTF), which is a novel filtering technology that provides the ability to rapidly select the central wavelength of an image with no moving parts. These filters, which have recently been developed as large aperture imaging quality devices, operate very efficiently in the red and near infrared spectral range, which is a well matched spectral range for limb scatter sensitivity to aerosol and cloud (*Rieger et al.*, 2014). Additionally, the spectral bandpass of the AOTF, which is typically between 3-6 nm at these wavelengths, is very suitable for the broadband scattering characteristics of the aerosol limb signal. The two dimensional imaging nature of the design provides the capability to achieve at least sub-kilometer resolution at the tangent point, which is on the order of the scale size of the upper troposphere and lower stratosphere (UTLS) aerosol features mentioned above.

It should be noted that the basic instrument design concept of ALI is very similar to that of the Atmospheric Limb Tracker for the Investigation of the Upcoming Stratosphere (ALTIUS) (*Dekemper et al.*, 2012), which is a Belgian instrument concept from the Belgian Institute for Space Aeronomy (BIRA). ALTIUS is designed to measure limb scattered sunlight; however, it also has solar, stellar, and planetary occultation modes and is scientifically focused on trace gas measurements, particularly for ozone, whereas ALI is optimized for aerosol observation.

# 2.2 Stratospheric Aerosol

The atmosphere of earth is a complex and complicated system and is effected by human activities. In the late 18th century, it was known that atmospheric temperature decreased with altitude and a theory had been raised that at specific certain altitudes the temperature must eventually go to absolute zero. This lead to a series of balloon campaigns to discover this mysterious altitude in the atmosphere which were noisy and unreliable due to the technology available. However, in the late 19th century the technology used in these sounding balloon had improved to a point where the atmospheric temperature could be accurately measured and it was found that at approximately 12 km an inversion point was discovered where the temperature started to increase and thus the tropopause, which separates troposphere and the stratosphere, was discovered (*Hoinka*, 1997). The stratosphere is the region of the atmosphere above the temperature inversion of the troposphere, where atmospheric temperature increases and the lower bound of the stratosphere is in-between 10 and 16 km from the high latitude to the tropics (*Andrews*, 1987). The stratosphere, which extends up to approximately 50 km, is thermodynamically stable and fairly dry (*Boucher*, 2015). The characteristic stability of the region limits vertical transport of the stratosphere, leading to long lifetimes, spanning from months to years, for non-volatile species (*Volk et al,* 1997; *Brasseur and Solomon*, 2005)

The stratosphere undergoes exchange of air with the troposphere though a series of dynamical processes including tropical convection, the polar vortex, and tropopause folding (*Holton et al.,* 1995) and circulation within the stratosphere is dominated by the slow Brewer-Dobson circulation (*Plumb and Eluszkiewicz*, 1999). Some chemicals can cross the tropopause thermal barrier into the stratosphere which allows for chemicals reactions to occur. One such reaction forms stratospheric sulfate aerosol, discovered by *Junge et al* (1961) though stratospheric balloon sondes, which are droplets of hydrated sulfuric acid (H­2­SO­4) formed from the oxidation of sulfate compounds, primary OCS and SO2 (*Brock et al.*, 1995). This stable layer of aerosol exists in the stratosphere from the altitude of the tropopause to approximately 30 km.

## 2.2.1 Aerosol Sources and Microphysics

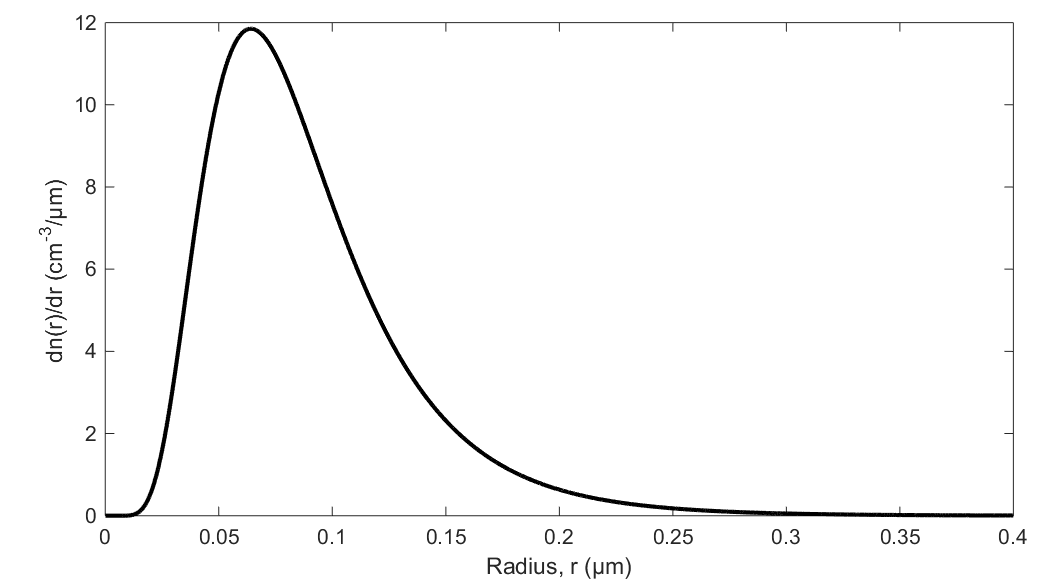
The sources for sulfate aerosol are primarily produced in the troposphere and formed though both natural and anthropogenic processes. One primary source of this sulfur is OCS which originates from marine processes, biomass burnings, and industry (*Kettle et al.*, 2002; *Notholt et al.*, 2003). OCS has a long lifetime in the troposphere and low solubility allowing for a significant portion to reach the stratosphere and some of it oxidizes and hydrates to form sulfate aerosol (H2SO4) and adds to the background aerosol layer (*Crutzen*, 1976).

Another source of sulfur is SO2, one source originates in the troposphere though industry from the burning of fossil fuels. Sulfur dioxide (SO2), which has a short lifetime in the troposphere and its concentration varies regionally (*Thomason and Peter*, 2006), also hydrates into sulfate aerosol. A second source of SO2 is from volcanic eruptions which are highly variable in location and time and can insert a large amount of sulfur directly into the atmosphere. The sulfur amounts that can be injected by volcanos can be so large they in fact dominate the stratospheric aerosol layer perturbing from the background levels. Examples of this perturbation of the aerosol layer were noted during the volcanic eruptions of El Chichon in 1982 (12-20 Tg of sulfur) (*McCormick and Swissler*, 1983; *Hofmann and Rosen*, 1983) or Mount Pinatubo in 1991 (20-30 Tg of sulfur) (*McCormick and Veiga*, 1992). However, after the Mount Pinatubo eruption a volcanically quiescent period occurred where aerosol layers returned to background. Following this period, a series of smaller volcanic eruptions have increased the background aerosol layer in the amount of 4-7% per year from 2000 to 2009 (*Vernier et al.*, 2011).

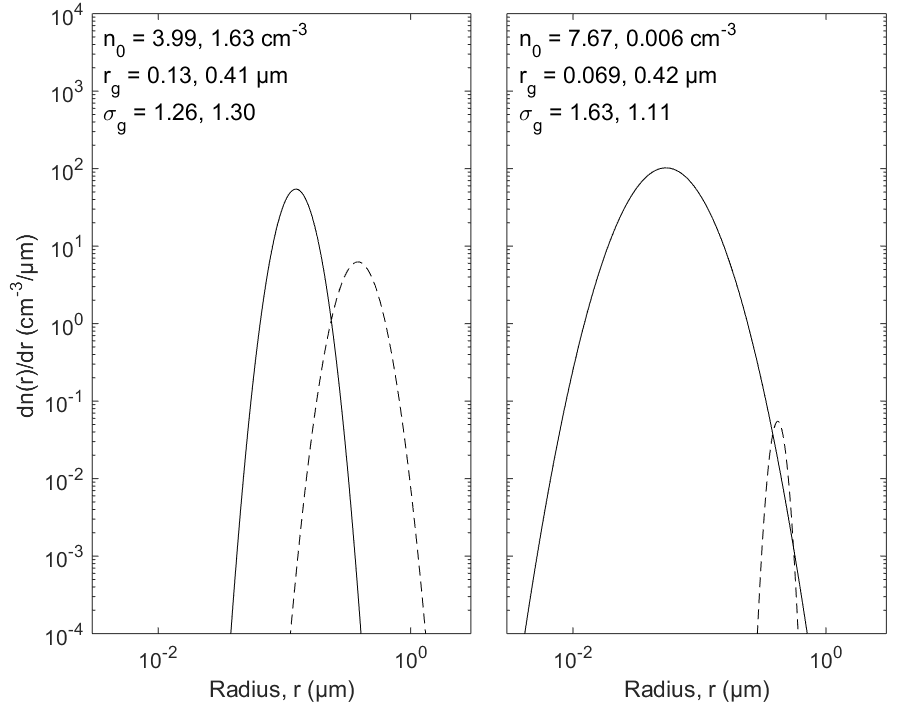
These sulfur sources undergo a series of reactions to be converted into H2SO­4 and have been found to form droplets of aerosol that consists of approximately 25% H2O and 75% H2SO4 (*Rossen*, 1971; Wang *et al.*, 1989). These aerosol particles form droplets of various sizes on the order of 0.05 to 1.0 µm depending on the processes of nucleation, evaporation and condensation that is underwent (*Junge et al.* 1961; *Brock et al.*, 1995; *Bingen et al.*, 2004). A log-normal distribution generally describes the particle size of aerosol in the form of

|  |  |
| --- | --- |
|  | (2.1) |

where is the aerosol number density, is the mode radius, and is the mode width (*Jäger and Hofmann,* 1991; *Hamill et al.* 1997). The particles are distributed normally over the logarithm of the radius. To fully describe aerosol for a single log-normal distribution the number density, mode radius and mode width is required and for a non-volcanic background aerosol layer an approximate mode radius and mode width of 0.08 µm and 1.6 respectively have been determined (*Thomason*, 1991) shown in Figure 2-1.



**Figure 2-1**: Log-normal distribution for non-volcanic background aerosol layer.



**Figure 2-2**: Bimodal particle size distributions fits from optical particle size. a) Distributions from a volcanic period after the Mount Pinatubo eruption recorded in 1993. b) Distributions from a background aerosol period eruption recorded in 1999. Both of the aerosol distribution measurement are from 20 km altitude with the solid line being the fine mode and the dashed line is the coarse mode. Figure is recreated from figure 5 of *Deshler et al.* (2003).

Optical particle counters have been used on board sondes from Laramie, Wyoming over the past 40 years to measure particle sizes in bins between 0.15 to 2.0 µm. The composition of the particle size distributions have been primarily unimodal, known as a fine mode, over non-volcanic periods and can be used as an acceptable distribution approximation during background periods. But during volcanic episodes, a bimodal log-normal distribution of aerosol particles is more representative of the stratospheric aerosols with the addition of a coarse mode (*Deshler et al.*, 2003; 2008; *Kovilakam et al.*, 2015). The coarse mode has larger particles than the fine mode and complicates determining full aerosol microphysical parameters since the number of required parameters has increased to six: a number density for both the fine and coarse mode, two mode radii, and two mode widths. Figure 5 from *Deshler et al.* (2003), recreated in Figure 2-2, demonstrated two bimodal particle size distributions from balloon optical particle counters. The first distribution is from a volcanic period in 1993 after the Mount Pinatubo eruption and another from a background period in 1999.

## 2.2.2 Climate Effects

Aerosol have several effects on the climate of the planet with a large amount of uncertainty to the overall effect (*Solomon et al.*, 2007). Through a direct effect, aerosol particles scatter incoming visible solar radiation away from earth increasing the albedo causing a cooling effect on the surface of the planet (*Lacis et al.*, 1992). A secondary effect from aerosols, that is dependent on aerosol particle size distributions, is a greenhouse effect that is caused by infrared radiation from the planet being scattered by aerosols (*Kiehl and Briegleb*, 1993). Aerosol also introduces an indirect effect to a radiative balance known as the cloud albedo effect. This is caused by condensation on existing aerosol particles forming cloud condensation nuclei which leads to an increase of the albedo for short wavelengths which also contributes to cooling the planet's surface. These types of cloud forming particles tend to increase the overall lifetime of the cloud increasing the cloud coverage of the planet thus increasing the albedo (*Charlson et al.*, 1992). Overall, the aerosol particles’ cooling effect dominates the warming effects and cools the surface of the planet (*Solomon et al.*, 2011).

During background aerosol periods there is little effect of the cooling from stratospheric aerosols but this can greatly change during periods of volcanic activity where the layer can be greatly perturbed. During the eruption of Mount Pinatubo in 1991 the sulfate aerosol load was increased by a large quantities causing a cooling of the lower atmosphere by 0.5oC (*McCormick et al.* 1995; *Soden et al.*, 2002) and 0.1 to 0.3oC on the surface (*Thompson et al.*, 2009; *Canty et al.*, 2013). And the surface temperatures did not return to pre-Pinatubo level until approximately 3 years after the eruption as the atmosphere filtered out the additional aerosol (*Hansen at al.*, 1996). More recently, a series of small to moderate volcanic eruptions have increased the background stratospheric aerosol layer (*Vernier et al.*, 2011b). The additional volcanic aerosol load may have been linked to a larger cooling effect, known as the global warming hiatus, when compared to background levels (*Solomon et al.*, 2011; *Haywood et al.*, 2014; *Fyfe et al.*, 2013). In order to fully quantify this effect, long term time series with microphysics information are requires to fully understand the aerosol forcing effect on climate modeling and change.

# 2.3 Aerosol Measurements

Two fundamental methods are used to measure aerosol concentrations within the atmosphere. The first of these methods are ground based and in-situ measurements which give good detail and information about a specific localized area. However, these measurements are limited in scope as they do not have global coverage that is inherent in satellite instrumentation. Both ground-based instruments and satellites have important roles in monitoring the planet’s aerosol content, however each of these methods have inherent advantages and disadvantages. An overview will be given on some of the common methods to determine aerosol concentration and why using different methods helps to increase the overall accuracy and precision of data sets.

## 2.3.1 In-Situ Measurements

In-situ measurement have occurred on balloon based platforms and aircrafts. Balloon instruments that use particle counters during the assent directly count the aerosol particles and can determine the particle size distributions. The optical particle counter is an active instrument that uses an incandescent light source internal to the device to optically count aerosol particles. This instrument has been launched from Laramie, Wyoming since 1971 to determine aerosol extinction and particle size (*Deshler et al.*, 2003; 2008; *Kovilakam et al.*, 2015). Another type of balloon instrument uses a passive light source, including the sun, moon, or stars, to determine aerosol extinctions. Instruments that use this type of technology are the Absorption par les Minoritaires Ozone et NOx (AMON) from 1992 to 2003 and Spectroscopie d’Absorption Lunaire pour l’Observation des Minoritaires Ozone et NOx - Nacelle 2 (SALOMON-N2) from 2007 onwards which use starlight and moon light respectively (*Berthet et al.*, 2002). Furthermore, aircrafts have been used to carry nephelometers to acquire in-situ measurements (*Beuttell and Brewer*, 1949; *Charlson et al.*, 1969).

In-situ measurements of aerosol extinction give direct measurement of scattered light from the altitude that the balloon is currently situated and allows for direct measurements of aerosol extinction and cross-section unlike remote sensing applications from satellites. However, these types of instruments do not achieve global coverage and only give aerosol extinction from a very localized region, like the Laramie, Wyoming optical particle counters, or have very few flights, for example AMON which had a total of six stratospheric balloon flights, three mid-latitude northern and three high-latitude northern flights. In order to achieve full global coverage satellite remote sensing instruments were created to fill the spatial gap.

## 2.3.2 Occultation

Satellite instrumentation capable of remote sensing stratospheric aerosol has been in use since the 1970’s, beginning with limb sounding solar occultation measurements, and it operational geometry can be seen in Figure 2-3. Solar occultation measurements have provided a reliable, accurate and essentially continuous long term record of vertically resolved aerosol extinction coefficient measurements, mostly from the series of Stratospheric Aerosol and Gas Experiment (SAGE) instruments (*Russell and McCormick*, 1989; *Thomason and Taha*, 2003). These SAGE measurements, which have a vertical resolution of approximately 1 km, have generally compared well with ground-based and in-situ measurements, although there are challenges associated with comparing the retrieved extinction profiles to other microphysical parameters (*Russell and McCormick*, 1989; *Kovilakam et al.*, 2015). However, solar occultation is generally a robust and stable technique as it directly measures atmospheric optical depth, along with the exo-atmospheric solar spectrum with each scan, allowing for straight forward retrieval of aerosol extinction coefficient (*Damadeo et al*, 2013). The major drawback to occultation satellites is the number of measurements it can record in a single day since the instrument needs to be viewing a sunrise or sunset event which limits the geometry to only 16-48 measurements per day depending on the orbit. Although the currently operational MAESTRO and ACE-Imager instruments on SciSat (*McElroy et al.*, 2007; *Gilbert et al.*, 2007) have had some success producing stratospheric aerosol extinction products (*Vanhellemont et al.*, 2008; *Sioris et al.*, 2010), the era of solar occultation measurements essentially came to an end with SAGE III in 2006. However, a manifestation of SAGE III is planned for deployment on the International Space Station in 2016 (*Cisewski et al.*, 2014).



**Figure 2-3**: An occultation instrument monitoring the atmosphere by scanning the atmosphere by looking directly at the sun.

## 2.3.3 Lidar

Through the transmission of a laser pulses though the atmosphere, a method known as lidar can determine atmospheric parameters through the measuring of the intensity of the backscattered laser light at different wavelengths and polarizations. Lidar was been used at ground based facilities to measure aerosol layers dating back to the 1960s (*Fiocco and Grams*, 1964) and are still used today. More recently LIDAR instruments have been used on satellite missions including the Ice, Cloud, and land Elevation Satellite (ICESat) from 2002 to 2010 (*Schutz et al.*, 2005) and Cloud-Aerosol Lidar and Infrared Pathfinder Satellite Observations (CALIPSO) which launched in 2006 (*Winker et al.*, 2007). Traditional lidar instruments have looked in the nadir direction (either straight down or up) however some instruments have looked slightly off-nadir, both can be seen in Figure 2-4. Lidar measurements have been used to determine aerosol plumbs from volcanos (*Chazette et al.*, 1995; *Sawamura et al.*, 2012) and long term trends (*Hofmann et al.*, 2009).



Figure 2-4: Lidar instrument showing a measurements in both the nadir and off-nadir lines of sight.

CALIPSO is a joint mission developed between the National Aeronautics and Space Administration (NASA) and the Centre National d'Etudes Spatiales (CNES) of the United States and France respectively. It uses a two wavelength polarized lidar system to achieve high resolution aerosol and cloud retrievals along the satellite's orbital track global coverage from 82oS to 82oN (*Young and Vaughan*, 2009). CALIPSO nominally measures backscatter profiles approximately every 300 m along track with approximately 200 m vertical resolution. However, the stratospheric backscatter signal is weak and requires averaging of only the night time measurements over several days and typically 0.5 km vertically and 500 km horizontally (*Vernier et al*, 2011b). Additionally, the uncertainty in the calibration with respect to the molecular background that is on the order of the stratospheric aerosol signal leads to a potential bias in the stratospheric measurements (*Rogers et al.*, 2011). CALIPSO was launched in 2006 and although it is presently still operational, it is also operating beyond its design lifetime.

## 2.3.4 Limb Scatter

The limb scatter technique measures the light that is scattered into the line of sight of the instrument from atmospheric interactions. These scattering interactions can undergo either single scatter or multiple scatter events. Single scatter is when light from the sun interacts with a particle in the atmosphere and scatters it directly into the line of sight of the instrument. Multiple scatter is when the photon of light undergoes several scattering events before entering the line of sight including scattering off of multiple particles in the atmosphere or scattering off of the ground into a scattering event into the line of sight. These events can occur any number of times before entering the instrument. The geometry for the limb scatter technique can be seen in Figure 2-5 and demonstrates the defining angles for this method. All angles are defined from the tangent point, which is the point where the distance between the line of sight and the surface of the earth is minimized, represented by the black dot. The Solar Zenith Angle (SZA) is the angle between the local vertical and the direction of the sun; the Solar Scatter Angle (SSA) is the angle between the direction of the sun and the line of sight and the Solar Azimuth Angle (SAA) is the angle between the projection of the sun on the plane of the line of sight and the line of sight. These angles can also be seen on Figure 2-5.



Figure 2-5: Limb scattering geometry measurement for an instrument where single and multiple scattering events occur.

The limb scatter method yields vertical resolution and allows for measurements to be taken during any daylight period with good SNR, however it requires the use of a complex forward model of multiple scattering processes along with at least some a priori knowledge of the aerosol scattering cross section in order to retrieve the extinction coefficient profile. The model needs to accurately determine the effect of multiple scatter since it consists of 10-50% of the measured signal depending of the specific geometry (*Oikarinen et al.*, 1999). Furthermore, due to the complex nature of the problem; a large amount of computational time and memory is required for an accurate calculation.

The first use of limb scatter was on the Solar Mesosphere Explorer (SME) (*Barth et al.*, 1983) to measure mesospheric ozone profiles in 1981. Much later, other limb scatter instruments were launched into low earth orbit that had the capability to measure aerosol extinctions including the Optical Spectrograph and InfraRed Imaging System (OSIRIS) instrument launched on the Odin satellite in 2001 (*Llewellyn et al.*, 2004) and the SCanning Imaging Absorption spectroMeter for Atmospheric CHartographY (SCIAMACHY) instrument on Envisat launched in 2002 (*Bovensmann et al.*, 1999). Both of these instruments are scanning grating spectrometers which can gather a single tangent point and scans the atmosphere to complete a vertical profile.

The OSIRIS version 5.07 data product provides 750 nm extinction profiles at approximately 2 km vertical resolution (*Bourassa et al.*, 2007) and has been shown to agree relatively well with SAGE II and SAGE III occultation measurements (*Bourassa et al.*, 2012b; *Rieger et al.*, 2015). The SCIAMACHY instrument uses a retrieval technique essentially similar to OSIRIS to retrieve aerosol profiles at 750 nm with approximately 3 km vertical resolution (*Ernst et al.*, 2012; *von Savigny et al.*, 2015). However SCIAMACHY observations ceased with the demise of Envisat in 2012 and although OSIRIS continues to operate, it is now in the fourteenth year of a mission designed for two years.

The most recently launched limb scatter instrument is the Ozone Mapping Profiler Suite Limb Profiler (OMPS-LP) on the Suomi-NPP satellite. Although similar in spectral range and vertical resolution to OSIRIS, OMPS-LP is an imaging spectrometer that vertically images the limb in a single measurement. The imaging capability of OMPS-LP provides a decrease in the time required to obtain a limb profile and so increases the along track sampling. Recent work on the feasibility of aerosol retrieval from OMPS-LP measurements show promising results (*Rault and Loughman*, 2013).

An instrument that is currently under development is ALTIUS (*Dekemper et al.*, 2012), which is a Belgian instrument concept from at the BIRA. ALTIUS is designed to image limb scattered sunlight, both vertically and horizontally across the track through the use of the Acoustic-Optic Tunable Filter (AOTF) technology (see section 3.1) and additionally has solar, stellar, and planetary occultation modes. ALTIUS is scientifically focused on trace gas measurements, particularly for ozone and the instrument has three channels, each channel with a separate AOTF, measuring for wavelengths from 250-2000 nm which could eventually be used for aerosol extinction measurements.

# 2.4 Radiative Transfer

To be able to use the limb scatter technique to determine aerosol extinction and particle size information, an understanding of radiative transfer is required. However, modeling the complex scattering interaction of light within the atmosphere is difficult. In this section, an overview of scaler radiative transfer will be performed followed by the necessary modifications to form polarized radiative transfer equations. Next, a description of scattering interactions important to aerosols will be underwent. Finally, an overview of the SASKTRAN radiative transfer model used within this work will be introduced.

## 2.4.1 Scalar Radiative Transfer

The scattering and absorption processes in the atmosphere are non-trivial and an adequate method to model this interaction is needed. The following will present a derivation of radiative transfer equations for the atmosphere with scaler radiance. The scaler radiance is the radiance without accounting for polarization. In order to accurately discuss radiative transfer, a coordinate system must first be defined. If we assume that a ray of light, , is propagating in a given direction, , and starts at a location, , with the initial position of , then the position of the ray along the path direction can be completely defined by its path length, . The basis of path length will be used to define the radiative transfer equations.

The fundamental theory for radiative transfer is known as Beer-Lamberts law. The law describes the change in intensity or radiance of light, , as it interacts with a thin layer of space or atmosphere, . The thin layer has particles which affect the attenuation of the light which is dependent on the number of particles, n, and the particle cross section, . If there are several different particles, the attenuation is a summation of the number densities and cross sections. The Beer-Lambert Law gives the following form

|  |  |
| --- | --- |
|  | (2.2) |

The extinction of the partials is a measure of the loss of light over a given distance and is defined as

|  |  |
| --- | --- |
|  | (2.3) |

Integrating Equation 2.2 forms the following result

|  |  |
| --- | --- |
|  | (2.4) |

The optical depth, , is defined as the extinction over the path length simplifying Equation 2.4 to

|  |  |
| --- | --- |
|  | (2.5) |

The above gives the radiance at point after it has gone through attenuation from .

Although this form of the Beer-Lambert’s Law is useful for describing the loss of light through scattering or absorbing from an initial source though a medium, the atmosphere also has incoming light that is scattered into the instruments line of sight from other directions or emitted from particles. To account for this additional source of light a source term, , is added to Equation 2.2 to yield

|  |  |
| --- | --- |
|  | (2.6) |

Using the fact that the change in optical depth is defined as

|  |  |
| --- | --- |
|  | (2.7) |

Equation 2.6 is rearrange to

|  |  |
| --- | --- |
|  | (2.8) |

Using the following derivative

|  |  |
| --- | --- |
|  | (2.9) |

and substituting it into Equation 2.8 yields

|  |  |
| --- | --- |
|  | (2.10) |

This form can now be integrated over the optical depth giving

|  |  |
| --- | --- |
|  | (2.11) |

Selecting the reference point at the observer to be and converting the equation back to path lengths yields

|  |  |
| --- | --- |
|  | (2.12) |

which gives the radiance as seen from an observer at a point, , along the line of sight.

With the full form of the radiative transfer equation, the source term must be determined. In the atmosphere there are three sources of additional radiation, blackbody emissions, photochemical reactions, and scattered light. For wavelengths from the visible to the near infrared (*i.e.* wavelengths less than 2 µm) there is little contributions from blackbody emissions, also known as thermal emission. Furthermore, as long as wavelengths where photochemical reactions occur are avoided this source term can also be ignored. This leaves the only significant source of light to be added into the line of sight to be from scattered sunlight. The source term for scattered sunlight is given by

|  |  |
| --- | --- |
|  | (2.13) |

The diffuse radiance is given by and is the radiation scattered into the line of sight from all directions. The phase function, , described the probability that a ray of light will be scattered from a direction, , into the line of sight propagation direction, . The scattering angle, , is defined as

|  |  |
| --- | --- |
|  | (2.14) |

Lastly, is the extinction only caused by scattering and not absorption. The term only allows the fraction of particles that scatter radiation, and not absorb it, to contribute to the source term.

As a final note, the calculation of the diffuse radiance is what makes this a computationally heavy problem. To completely solve for the diffuse radiance, the radiance at every point in the atmosphere must be determined. Furthermore, the light can be scattered multiple times in the atmosphere, requiring a diffuse radiance for each order of scatter. Each successive scattering adds smaller contributions to the final radiance at the observer. Through this iterative process the full multiple scatter solution to the radiative transfer equation is

|  |  |
| --- | --- |
|  | (2.15) |

The multiple scatter term is calculated until the contribution is sufficiently small to be negligible.

## 2.4.2 Vector Radiative Transfer

The scaler radiative transfer equation works well for systems that do not measure polarized light as the effect of polarization on the scaler radiance is small. However for instruments that measure polarized light, a vector radiative transfer equation is required. Before polarization can even be discussed, a method to quantify polarization must be defined which will be the Stokes vectors. The Stokes vectors are given as

|  |  |
| --- | --- |
|  | (2.16) |

where is the scaler or total radiance, is horizontally polarized light or light rays aligned with the local x-axis, is diagonally polarized light measured 45o counter clockwise from the local x axis, and is the counter clockwise circular polarized light (*Bickel and Bailey*, 1985). With the Stokes vectors the degree of polarization can be determined. If the equality holds true then the light is fully polarized, otherwise it is only partially polarized.

With the addition of polarization, the radiative transfer and source term equations (Equations 2.12 and 2.13) need to be rewritten with polarization included. The polarized versions are

|  |  |
| --- | --- |
|  | (2.17) |
|  | (2.18) |

which are the vector radiative transfer and source term equations respectively (*Mishchenko et al.*, 2002). The radiance and the source terms are now forms of the Stokes vectors in 4 by 1 matrices and the phase function is a 4 by 4 tensor that describes the probability of the incoming light to be scattered in the propagation direction with a specific polarization. The polarization equation adds extra computation and memory consumption since the polarization must be computed at each step in the interactive radiative transfer, which is not trivial, and stored in memory, which is four times the size of a standard scaler radiance calculations.

With the complete vector polarized radiative transfer expression the two scattering interactions that pertain to determining aerosol will be described. The first interaction is Rayleigh scattering which defines the scattering of the background atmosphere, and Mie scattering which determines how incoming light will scatter off of aerosol particles.

## 2.4.3 Rayleigh Scattering

Rayleigh scatter is the scattering performed on the molecular background of the atmosphere. The first calculation of molecular atmospheric scattering cross sections was by Lord Rayleigh where he assumed the molecules were dielectric spheres with radii much less than the wavelength of the light. Later, the King correction was added to the Rayleigh scattering cross section, , to yield the following expression

|  |  |
| --- | --- |
|  | (2.19) |

which is highly dependent on wavelength, , which is in cm. The parameters and are the volume polarizability, in cm-3, and the depolarization ratio, which is unitless TODO:CHECK UNITS (*Sneep and Ubachs*, 2005).

The other important quantity for scattering is the phase function. For Rayleigh scattering, the vector model phase function is given by the Rayleigh-Gans approximation (*Mishchenko et al.*, 2002)

|  |  |
| --- | --- |
|  | (2.20) |

Each component of the phase matrix itself is smooth which allows for easy and accurate calculation for Rayleigh scattering.

## 2.4.4 Mie Scattering

For larger particles, like sulfate aerosol, Rayleigh scattering no longer holds since the wavelengths are on the order of the size of the particles and Mie scattering must be used. *Mie* (1908) solved Maxwell’s equations in a general sense with a solution using spherical Bessel and Henkel functions. Only the fundamental Mie scatter equation will be presented here but information on the Mie scatter can be found in *van de Hulst* (1957). The scattering cross section from Mie theory is given by

|  |  |
| --- | --- |
|  | (2.21) |

where is the wavenumber and the coefficients and are given by

|  |  |
| --- | --- |
|  | (2.22) |
|  | (2.23) |

With the radius of the particle given by , n is the index of refraction of air, and and are the normalized half-integer order Bessel functions of the first kind and Henkel functions of the second kind respectively. The phase function for Mie scatter for a vector solution has the following form (*Hansen and Travis*, 1974)

|  |  |
| --- | --- |
|  | (2.24) |

The terms in the phase matrix, and , are known as the amplitude function and given by

|  |  |
| --- | --- |
|  | (2.25) |
|  | (2.26) |

where are the Legendre polynomials.

In the atmosphere, various particle sizes occur and a log-normal distribution (Equation 2.1) is assumed for aerosols. In order to determine effective scattering cross-section, a weighted average over the particle radius is performed

|  |  |
| --- | --- |
|  | (2.27) |

The weighted average is similarly performed to determine the effective phase function for a particle size distribution.

It should be noted that although the theory is well founded to calculate the Mie scattering cross sections and phase functions, in practice it is computationally intensive since the terms consist of infinite sums. Work done by *Wiscombe* (1980) has allowed for effective computation of the Mie scattering coefficients, which has been implemented into the SASKTRAN radiative transfer engine.

## 2.4.5 SASKTRAN Radiative Transfer Model

The SASKTRAN radiative transfer was first developed to solve the scaler radiative transfer equation in a fully spherical atmosphere for both single and multiple scatter with a one dimensional atmosphere (*Bourassa et al*., 2008). The first source term, , is from the radiance from the sun which is assumed to encounter the earth in parallel unpolarized rays. Higher order source terms are based off of the scattering within the atmosphere to produce the diffuse radiance, however it is impractical to calculate the radiance at every point in the atmosphere. Instead, a series of diffuse radiance profiles are used to simulate the diffuse radiance term. Another important assumption in the SASKTRAN model is that the ground reflection is assumed to have a Lambertian unpolarizing distribution when any radiance encountering the surface of the earth is reflected at an efficiency, known as the albedo, and is evenly distributed in all directions.

Recent upgrades have been performed on SASKTRAN and have led to a new engine known as SASKTRAN High Resolution model or SASKTRAN-HR (*Zawada et al.*, 2015) which has expanded the model to be able to perform radiative transfer calculations with a fully three dimensional atmosphere. This update allows the model to vary the atmospheric concentration in the model not just in the vertical direction, like the original SASKTRAN, but in both of the horizontal geometries, allowing for a variance seen in the atmosphere.

The most important update to the SASKTRAN-HR model for this work is the addition to calculate the vector or polarized radiances (*Dueck et al.*, 2015). Using the vector model allows for SASKTRAN to compute the Stokes vectors in the reference frame of the model, which can be rotated into any desired frame of reference through the use of a provided basis by SASKTRAN-HR. The polarization output from SASKTRAN-HR only preforms polarized calculations up to the third order scattering interaction and all subsequent orders are assumed to be scaler.

# 2.5 ALI Prototype and Stratospheric Balloon Flight

New satellite instrumentation will be necessary to continue the stratospheric aerosol data set, especially since the current generation of remote sensing instruments are aging and are operating well beyond their designed lifetimes. One such instrument proposed to fulfill this role is the Aerosol Limb Imager (ALI) which will image the atmosphere in a limb scatter geometry and with the correct orbit able to achieve daily global aerosol extinctions. In order to achieve the imaging of the atmosphere an AOTF will be used which will complicate the optical design and has not been well tested in a space environment. However, before a satellite mission can be considered, a proof of concept or prototype mission needs to be undergone to verify feasibility of the technology. This work will entail the designing, fabricating, building, and calibrating of a prototype ALI instrument which will be discussed in chapter 3. Another important aspect of ALI is that the measurements are inherently polarized due to the nature of the AOTF and chapter 4 will perform a study to determine the best geometry and polarization for a future satellite or balloon missions. When ALI was completed, it was packed and shipped to Timmins, Ontario and was flown on a stratospheric balloon in the fall of 2014, performing an engineering test of the technology. After the mission, the radiance measurements were used to verify that this technology can retrieve aerosol profiles accurately which is the focus of chapter 5.