CHAPTER 5

STRATOSPHERIC BALLOON FLIGHT AND AEROSOL RETRIVALS

# 5.1 Stratospheric Balloon Flight

After the completion of the ALI system tests on August 18, 2014 the instrument was transported to Timmins, Ontario and preparation were underwent for the balloon launch from August 25, 2014 until September 19, 2014. During this balloon campaign there were seven balloon launches with ALI being a part of the seventh balloon. The flight of ALI took place on September 20, 2014.

## 5.1.1 Preflight Preparations

The Canadian Space Agency (CSA) balloon launch facility in Timmins, Ontario located at the Victor M. Power Airport (48.47oN 81.33oW). ALI arrived at the base on August 25, 2014 with a launch window from September 8 to 14, 2014. In between the arrival of ALI and the balloon launch, ALI had to be verified to have survived transportation, a seal within the CCD needed to be removed, thermal insulation needed to be added, and finally ALI needed to be integrated onto the Centre National d'Etudes Spatiales (CNES) CARMEN-2 gondola.

ALI was unpacked and set up on a test bench at the launch facility. A visual inspection occurred to verify that no obvious damage occurred to the instrument during transportation. Once completed, ALI was connected to its electronics and power boxes and ALI was powered on. A ground station computer was used to connect to ALI and preform a system check, including verification that of automated startup, telemetry connection was established, that the system powered on correctly with no error and that the science operation program functioned. With this test it was verified that no functional problems occurred to the device during transportation, all temperature and voltage sensors, GPS module, and CCD camera were reporting valid diagnostic values.

Once the ALI system was verified to be operational an imaging check was performed to check that no optical components suffered damage or slippage during transportation. An EIA 1956 resolution target was illuminated by a 250 W tungsten halogen light source and was imaged by ALI to verify the optical layout. The recorded images were very similar to the one taken in the laboratory before the leaving Saskatoon.

Following the successful test of ALI the final preparations were needed prior to beginning integration with CARMEN-2 were performed. First, the CCD used by ALI had a sealed chamber that was in a vacuum state designed to be at atmospheric pressure and would be required to be unsealed before the flight. The unsealing is done in order to not develop a strong pressure gradient between the CCD chamber and the low pressure of a 35 km environment causing permanent catastrophic damage to the CCD detector. At the launch facility ALI was taken to a semi-clean area to unseal the CCD chamber. A panel was removed on the side of the camera and the seal to be removed can be seen in Figure 5-1. The orange o-ring was removed with associated sealing components and the vacuum seal was broken. The chamber panel was replaced and ALI was moved back to the integration hall and another set of test resolution targets were taken to verify the correct operation of the ALI. All resolution target were similar with from the set before the chamber was unsealed expect there was approximately a 5% drop in counts which may have been caused by unsealing the chamber or a change in the lighting conditions of the resolution target.



**Figure 5-1:** Side of the QSI CCD with the panel that contains the vacuum seal opened. The orange o-ring seen in the cavity is removed from the chamber to open the vacuum seal to the camera's CCD chip.

The next step before ALI could be integrated was to add thermal in order to protect ALI from the thermal environment at approximately 35 km. The first thermal concern was the instrument falling to a temperature were the electronics were too cold to function. The instrument would have to be in complete darkness during the assent which would result in little to no solar heating. Furthermore, ALI will pass through the tropopause where temperatures can be as cold as -70oC. Insulation, in the form of foam, was added around the exterior of the instrument to give ALI thermal isolation from the cold environment. The second concern was once CARMEN-2 was at float altitude ALI would have to be able to survive the direct heating from the sun's radiation which could have overheating. The impact of the sun's energy was reduced on ALI by adding a thermal reflector to the outside of the thermal insulation which would reflect a portion of the incoming solar radiation away from ALI.

With the completion of the thermal management, ALI was ready to be mounted onto the CARMEN-2 gondola. ALI can be seen mounted on the CARMEN-2 gondola in Figure 5-2 and ALI used the power and communication subsystems of CARMEN-2. Testing was performed with collaboration from the CARMEN-2 team to check there were no issues between ALI and CARMEN-2’ systems. A problem was found in the communication module, named Siren, between ALI and the ground station computer. With as assistance from the CARMEN-2 team the correct Ethernet setting were determined and a correction to the ALI operation code was applied.

During the integration phase it should be noted that several instruments were also being verified with the CARMEN-2 systems for integration onto the gondola including four other Canadian instruments, including the OSRIS development model (*Kozun*, 2015; *Taylor*, 2015), and SHOW to measure water vapour.

The CNES gondola is an actively pointed gondola with azimuthal pointing precision better than 1’ with the use of an onboard star tracker. ALI was orientated so it would be maintained at 90o from the azimuthal direction of the sun, with an overall southern field of view during the mission.



**Figure 5-2:** The ALI instrument is mounted on board the CARMEN-2 gondola (top shelf on the right). ALI located next to SHOW, another Canadian instrument with collaboration between ABB, York University, and the University of Saskatchewan. ALI has its red tag cover over the optical entrance to protect the instrument from dust and other contaminates. Thermal insulation has been added to the instrument and during the flight sun side will be on the side of SHOW. Some of the reflective layer was blacked out to not cause additional stray light into SHOW optical path.

## 5.1.2 Balloon Flight

The flight plan for the CARMEN-2 gondola was once float altitude was reached and the sun had risen ALI, OSIRIS, and SHOW would perform their operational missions for the first four hours of the campaign. The operation objectives for ALI included a dark imaging suite for calibration purposes, and an aerosol imaging suite for aerosol measurements. A secondary goal was to test the sensitivity to aerosol of ALI with respect to SSA by recording images at various azimuth directions. After the end of the ALI mission, the instrument was to be powered off and other instruments on CARMEN-2 were to gather measurements.

The flight of CARMEN-2 was delayed past it launch window of September 8 to 14, 2014 due to poor weather conditions. On September 19, 2014 at 05:35 UTC (01:35 local time) ALI was launched as part of the Nimbus 7 mission from the CSA Timmins balloon launch facility. During the launch, the sky was clear with light winds allowing for a safe and uneventful launch. The ascent of the gondola occurred in darkness and reached its flight altitude of 36.5 km at 8:17 UTC. First light was observed by ALI at 9:39 UTC and spectral images were recorded until 14:42 UTC. ALI was powered off at 17:15 UTC.

A visualization of the flight path with major landmarks noted can be found in Figure 5-3a. Temperature profiles for the ambient atmosphere and instrument are shown in Figure 5-3b. The black curve is the ambient atmospheric temperature at the gondola altitude and location during the flight as obtained from ECMWF reanalysis (*Dee et al.*, 2011). The blue, green, and red are from temperature sensors onboard ALI located on the baffle, camera, and RF driver respectively. The baffle temperature sensor was attached just on the inside of the ALI right by the entrance aperture for the system and monitors the temperature at the front of the system. The camera sensor is attached to the back of the CCD camera and the RF driver sensors measures the surface temperature of the RF driver. ALI was thermally insulated to keep the system warm whereas the baffle temperature sensor is relatively uninsulated from the extreme cold of the tropopause. The effect of the cold tropopause can be seen on the gondola at approximately 6:00 UTC. The cooling effect can even be seen on the interiors CCD and RF driver sensors which are isolated from the exterior temperature. After, the internal temperature drop the system reaches an equilibrium temperature until the sun light rises and solar radiation comes into contact on the instrument at approximately 10:00 UTC at which point there is an increases in the systems temperature. The temperature of the system are kept within operating range with the aid of the reflective material during the flight.



**Figure 5-3:** (a) The GPS data from ALI during the Nimbus 7 mission generated via Google Earth. The colour of the line represents the absolute speed of the gondola during the mission. Important landmarks are noted on the image. The end of mission represents the end of the primary aerosol mission. No GPS data was collected from ALI after power down. The location of image 208 is the red label. (b) The temperature and altitude profiles from the Nimbus 7 flight. The time of image 208 is shown by the cyan vertical line and first light measured by ALI is occurs at the magenta vertical line.

During the mission, ALI operated in two primary acquisition modes, a calibration mode and an aerosol imaging mode. The first mode, the calibration mode, was primarily used during ascent when the gondola was in the darkness and intermittently between the aerosol mode during sunlit conditions. During this mode the filtering of the AOTF was not enabled and the system imaged essentially only dark current during the ascent in darkness and stray light during sunlit conditions. Eight exposures are taken in the calibration mode with 0.05, 0.1, 0.5, 1, 2, 3, 5, 10 second exposure times.

The second operational mode, the aerosol mode, recorded measurements in a cycle that contained 13 pairs of images across the spectral range (650-950 nm every 25 nm), the pairs being a calibration image with the “AOTF-off” and an image of the limb. Each cycle took approximately 10 minutes with each measurement set taking approximately 40 seconds to acquire with initial exposure times shown in Figure 5-4 in blue which were the exposure times determined during the roof testing of ALI (see section 3.6.1). However, during the flight it was determined that the calculated exposure times were not long enough and the images were underexposed. The underexposure is believe to be caused by the initial exposure time calibration curves being calculated with simulated scaler radiance since the SASKTRAN-HR polarization module had not yet been completed development. So the exposure time curve was recalibrated during the flight using the image statics that were sent down with the house keeping. A comparison of the two exposure time curves with the percent increase can be seen in Figure 5-4. The percent increase is given by

|  |  |
| --- | --- |
|  | (5.1) |

where is the exposure time for the original calibration and are the updated exposure times calculated from the flight.

After a successful flight that lasted for 16 hour 19 minute with the landing at 21:54 UTC. During the flight ALI successfully gathering 216 aerosol images. The gondola landed 70 km from Amos, Quebec or approximately 250 km from the launch facility. CARMEN-2 was recovered by the balloon recovery team and was returned to base on September 21, 2015. ALI was removed from the gondola, repacked and transported back to Saskatoon, Saskatchewan were the data could be processed.



**Figure 5-4:** During the flight the calibrated exposure times was updated. The blue curve represents the exposure times from the ground calibration and the red curve is the recalibration during the flight. The black curve is the percent change in between the pre-flight calibrated results and the during flight calibration.

# 5.2 Limb Measurements

After the successful post-flight recovery of ALI was unpacked and check for any damages in Saskatoon, SK on September 25, 2014. No obvious damage had occurred to ALI from the fight and the instrument was functioning correctly. 216 raw aerosol mode images were obtained from the flight and calibration was performed including pointing alignment and the calibrations as detailed in Section 3.3.

The images needed to have complete pointing information added to the images, the raw images only had position and altitude information from the onboard GPS. The azimuth and zenith directional information needed to be added to the images as they are needed to determine the line of sight of each pixel on the CCD. The CNES team supplied the SAA information from the flight and was correlated to the images using GPS time. However, no information was supplied about the zenith direction of the gondola, we were just notified that the gondola was zenithally stable throughout the entire flight. So some manual calibration of the zenith angle occurred. ALI was tilted at 3o so the zenith as assumed to be 93o for ALI. This starting guess was not accurate enough since features in the radiance profiles did not retain the same altitude over the course of a few images. To determine a more precise zenith angle, the zenith angle was varied from 92o to 94o in 0.1o intervals and the tangent altitude was calculated for each case. Then the radiance profiles for each zenith angle was compared to find zeniths where the features were align. The zenith angle with the optimal alignment was determined to be 92.55o with an uncertainty of 0.10o.

The calibration techniques discussed in section 3.3 were then applied to the raw images to find the final radiances. As an example image number 208 is used to demonstrate the steps in the calibration on a flight image. Image 208 is a 750 nm taken at 13:57 UTC with a solar zenith angle and solar scattering angle of 63o and 98o respectively. The dark current and DC offset have been removed form image 208 using the Equation 3.42. Next the stray light is removed by using the AOTF-off or calibration image and removing it from the AOTF-on or measurement image. The result of this procedure can be seen in Figure 5-5. In the first panel abnormal bright spots are noticed in the right side and the top right of the measurement. These same features are noticed in the stray light image. By subtracting the AOTF-off image from the measurement image a final smooth measurement image is seen. Finally, a flat fielding calibration is performed (see section 3.6.5) and a final calibrated image can be seen in Figure 5-6a. Remember that no absolute calibration was performed on ALI, so the radiance are relative radiance in arbitrary units to the 775 nm laboratory calibration. The error, , on a given pixel for the radiance measurements were given by

|  |  |
| --- | --- |
|  | (5.2) |

where is the readout uncertainty from the CCD, which is 15 counts, is the error in the DC offset calibration, is the error from the dark current in the CCD, is the error in the stray light calibration, and is the error in the flat field corrections.



**Figure 5-5:** Stray light removal technique is performed using image 208 which is a 750~nm measurement. The top panel is the image after the DC offset has been removed from the measurement. The middle panel is the associated AOTF-off image and stray light features are seen in the upper right of the image as well as light being registered in the entire right side of the image. The final panel is the first panel minus the second panel and the abnormal gradient has been removed from the final image, leaving a cleaner radiance profile.



**Figure 5-6:** (a) Final calibrated 750 nm image, taken at 13:57 UTC located at 48.55oN, 80.00oW with a solar zenith angle and solar scattering angle of 63o and 98o respectively. (b) The same 750 nm image with the mean of the profile removed from the image leaving the residual signal that shows thin clouds in the troposphere.

From image 208 the horizontal structure across the image is nicely revealed by calculating the mean radiance profile across the image and then removing it from each profile. This is shown in Figure 5-6b, where thin clouds (2 km vertical extent or less) are clearly seen near and below the tropopause level, with substantial variation in tangent altitude across the horizontal field of view. These clouds were also observed from other instruments on board the gondola during the mission (B. Solheim, private communication). A brief check on the CALIPSO quick-look plots also shows clouds at a maximum height of approximately 13 km from measurements taken at 08:40 UTC at 47.24oN, 95.25oW, the nearest measurement point to the ALI location and time. Although these images only have a 35 km extent in the horizontal direction, there is also some indication of horizontal variation in radiance significantly above the cloud level, possibly due to real atmospheric variability in the aerosol layer. It should also be noted that some high altitude stray light is also visible in this mean residual image that was not observed in the laboratory tests. This may be due to contamination from scattering from a baffle vein or a nearby component of the gondola, although the true cause is unknown at this point.

For ease of further analysis, and to increase the precision of the measurements to a minimum of 0.6 MTF the images were averaged into cells of 25 pixels horizontally, and average vertically onto a 1 km tangent altitude grid. The errors for the averaged radiances, , is given by

|  |  |
| --- | --- |
|  | (5.3) |

where the errors for each pixel, , are summed in vertical, , and horizontal, , directions from the starting pixel, for the vertical and for the horizontal, to the final pixel in the average, for the vertical and for the horizontal. The radiance profiles from the center column of the images for all measurements obtained during the flight are shown in Figure 5-7. The first set of profiles, the dashed lines, which start near zero and move toward larger values, are the measurements that were recorded near and during sunrise so the gradual increase is therefore expected. Measurements obtained for solar zenith angles less than 90o are represented by the solid lines. These radiance profiles follow a similar, and expected exponential shape, with some variability at tangent altitudes below 12 km corresponding largely to changing cloud conditions.



**Figure 5-7:** Averaged ALI relative radiance vectors from 12 of the 13 wavelengths from the NIMBUS-7 flight. Each panel presents the radiance vectors from a different wavelength measured which is denoted in the top right corner. The dashed lines are radiance profiles where the solar zenith angle is greater than 90o and solid lines are profile where the solar zenith angle is less than 90o.

A full cycle of 13 spectral images (numbers 204-216) were used in Figure 5-8 to show the spectrum of relative calibrated radiances at selected tangent altitudes. The estimated uncertainty in the radiance is represented by the shading has calculated using Equations 5.2 and 5.3. The uncertainty is approximately five percent from 5 to 20 km and increases up to eight percent from 20 to 35 km. The spectra display the expected and relatively smooth fall off in intensity with increasing wavelength with Chappuis ozone absorption seen at the lower wavelengths; however, the reason for the peak in the spectra at 875 nm is not known and may be due to an inconsistency in the pre-flight calibration.



**Figure 5-8**: Relative radiances spectrally from 650 nm to 950 nm as measured from ALI at approximately 14:20 UTC consisting of images number 204 to 216 looking 90o in the azimuth from the sun facing southwards. These spectral profiles are presented at several tangent altitudes with a horizontal look direction of 0o. The shading represents the error on the radiances.

# 5.3 Aerosol Retrievals

Test

THESIS:

With the successful mission and relative radiance data from ALI the results need to be used to determine aerosol atmospheric parameters. The following sections will describe the MART retrieval method used, the results from using the MART method on the ALI data, and a comparison to the OSIRIS version 5.07 aerosol product and the from SALOMON a French gondola that flew before the CARMEN-2 gondola that measured aerosol extinction.

## 5.3.1 Aerosol Extinction Retrieval Methodology

Test

PAPER:

As a first application of the ALI measurements, we have applied a slightly modified version of the standard OSIRIS stratospheric aerosol extinction retrieval (Bourassa et al., 2012b) to the flight measurements. This inversion algorithm, which is applied from the tropopause to 30 km altitude, assumes log-normally distributed hydrated sulphuric acid droplets in order to calculate the aerosol scattering cross sections from the Mie scattering solution (Wiscombe, 1980). The modeled radiances for the nonlinear inversion were computed with the SASKTRAN High Resolution radiative transfer engine (SASKTRAN-HR) (Bourassa et al., 2008; Zawada et al., 2015) using the newly developed vector module for polarization (Dueck et al., 2015). The output of SASKTRAN-HR gives the Stokes vectors for the radiance on the model reference frame, which are then rotated into the instrument's coordinate system. Once rotated, the polarization signal required to match the ALI measurement is the vertical polarization given by

whereandare Stokes parameters defined by and . The variables and are the horizontal and vertical component of the electric field in the instrument reference frame.

The relative radiance measurements from ALI are used to create measurement vectors, , as specified in Bourassa et al. (2012) in the form,

where is the measured relative radiance from ALI and is the relative radiance at a high reference tangent altitude where there is little aerosol contribution. For the ALI measurements, the highest possible tangent altitude where the signal is above the noise threshold is approximately 30 km tangent height. The second term in Eq. 8 uses modeled radiances from SASKTRAN-HR with only the molecular atmosphere to approximately remove the Rayleigh signal. This is done to improve the speed of the convergence of the retrieval (Bourassa et al., 2012b). An initial guess state, , for the aerosol extinction and an assumed particle size distribution profile are set in the SASKTRAN-HR model. The forward model vector is then constructed similarly to the measurement vector, and used in combination with the measurement vector to update the aerosol extinction coefficient profile using Multiplicative Algebraic Reconstruction Technique (MART) algorithm,

where is the aerosol extinction at each model altitude, and denotes a tangent altitude from the measurements. is the weighting matrix that relates the importance of each element of the measurement vector to each shell altitude. This method described in detail by Bourassa et al. (2007).

Once a retrieval has been completed for a measured radiance profile, the result is then used to estimate the error in the retrieved extinction. For each altitude, a gain matrix, , is calculated through successive numerical perturbation of the measurement vector and re-retrieval (Rodgers, 2000). A much faster method to use the Jacobian to determine the error has been performed (Bourassa et al., 2012a) but makes an assumption that the gain matrix is equal to the inverse of the Jacobian, as typically the averaging kernel is close to the identity matrix. However, this method adds additional uncertainty to the error estimate and with a limited set of balloon data, it is possible to calculate the gain matrix directly. The error at each retrieved altitude is then given by

where is the covariance matrix of the measurement vector and is the covariance of the retrieved aerosol profile (Rodgers, 2000). The reported precision for ALI aerosol extinction retrievals is the square root of the diagonal of .

Using the retrieved extinction profiles for the complete spectral range, we have attempted a determination of the Angström exponent using a method similar to that outlined by Rault and Loughman (2013) for the OMPS-LP analysis. In this method, the independently retrieved extinction profiles at each wavelength and altitude are fit with a straight line in log-wavelength, log-extinction space. The slope of this line corresponds to the Angstrom exponent. This is then used to find the best match to the spectral dependence of the Mie scattering cross section in order to update the particle size distribution. With only one piece of information, the mode-width of the log-normal distribution is fixed to 1.6 and the mode radius is updated. The extinction retrievals are then performed again at each wavelength and the process is iterated until the Angstrom exponent, corresponding to the determined mode radius, converges.

Ideally, the ALI measurements would be used independently to also retrieve ozone in the Chappuis band. However, due to the spectral range of the prototype, only a small fraction of the long wavelength side of the absorption band was captured. For this analysis, we have not retrieved the ozone profile but have set the ozone profile in SASKTRAN-HR to an average of the five closest coincident ozone profiles measured by OSIRIS at the ALI location and time. The surface albedo used is also from the OSIRIS scans since the two instruments share a similar measurement method and should determine a similar albedo for the cloudy conditions. Preferably albedo would be determined from the ALI following the method of Bourassa et al., 2012b, however due to the lack of an absolute calibration this was not possible.

THESIS:

In order to be able use the the measured radiances from ALI a forward model is needed to be able to determine the atmospheric state for a specific geometry given an atmospheric state. The SASKTRAN-HR polarized radiative transfer model \citep{Dueck2015} which was discussed in detail in \autoref{TODO:SASKTRAN}. This section focuses on the method used to solve the nonlinear inverse problem known as the Multiplicative Algebraic Reconstruction Technique (MART.)

The Multiplicative Algebraic Reconstruction Technique (MART) is a measurement inversion technique where a measured value can be converted into a wanted usable physical quantity via an iterative method. A measurement vector, $\mathbf{y}$, that has sensitivity to the desired physical state is constructed and is computed with a forward model, $F$ is used, in this case SASKTRAN, with an input state, $\mathbf{b}$, and a wanted parameter $\mathbf{x}$. The measurement vector is then calculated with the model and is yielded by

\begin{equation}

\mathbf{y} = \mathbf{F}(\mathbf{x},\mathbf{b}) + \boldsymbol{\epsilon}

\end{equation}

where $\boldsymbol{\epsilon}$ is the noise on the measurement vector.

The MART was developed by \citep{Bourassa2012a} used with limb scatter instruments. An initial guess of the atmospheric state, $\mathbf{x}$ or a priori is assumed which is iteratively updated based on the ratio of the measured state, $\mathbf{y}$, to the forward model, $\mathbf{F}(\mathbf{x},\mathbf{b})$ and is computationally efficient as it allows for the updating of the atmospheric state without calculating the jacobian \cite{Degenstein2009}. This method is similar to the Chahine relaxation method \citep{Chahine1972} with the addition of a weighing matrix, $\mathbf{W\_{ij}}$ which allows for relates the importance of each measurement vector to each shell altitude of the model, $i$ and the measured tangent altitude, $j$. The iterative method is given by

\begin{equation}

x\_{i}^{n+1} = x\_{i}^{n}\sum\_{j}\frac{y\_{j}}{F\_{i}(\mathbf{x},\mathbf{b})}W\_{ij}.

\end{equation}

Now the formation of the measurement vectors is required in order to fins the atmospheric state but before the measurement vector can be discussed a description of the radiance is required.

Now the formation of the measurement vectors is required in order to finds the atmospheric state but before the measurement vector can be discussed a description of the radiance is required. The polarized SASKTRAN-HR model gives the Stokes vectors for the radiance on its internal coordinate grid which can be retrieved from the model and then rotated to match ALI's coordinate system. Once rotated the polarization required to match ALI's measured results is the vertical polarization given by

\begin{equation}

I = \frac{1}{2}\left(\mathrm{S\_{0}}-\mathrm{S\_{1}}\right)

\end{equation}

where $\mathrm{S\_{0}}$ and $\mathrm{S\_{1}}$ are Stokes parameters representing total radiance and horizontally polarized radiance respectively. From here on any reference to the radiance, $I$, in the rest of the chapter will refer to the vertical polarization measured by ALI.

The relative radiance level 1 data from ALI are used to create measurement vectors, $y$, in the following form

\begin{equation}

\mathbf{y} = \log\left(\frac{\mathbf{I}(\mathbf{z},\lambda)}{I(z\_{ref},\lambda)}\right)-\log\left(\frac{\mathbf{I}\_{rayleigh}(\mathbf{z},\lambda)}{I\_{rayleigh}(z\_{ref},\lambda)}\right)

\label{eqn:5.3:measurementVector}

\end{equation}

where $\mathbf{I}(\mathbf{z},\lambda)$ is the measured radiance from ALI and $I(z\_{ref},\lambda)$ is the radiance at a reference altitude used to normalize the signal from a high altitude where there is little aerosol contribution, for ALI the highest possible altitude where the signal is above the noise threshold is around 25-30~km tangent height. The second term is modeled values from SASKTRAN-HR with only background neutral atmosphere to remove the rayleigh signal from the measured radiances which is done to improve the speed of the convergence of the retrieval. The left panel of \autoref{fig:5.3:measurementVectors} shows the measurement vector from a 750~nm image, number 208, from the 0\si{\degree} horizontal line of sight. The final measurement vector, $\mathbf{y}$ is shown in the black, with the first term of \autoref{eqn:5.3:measurementVector} is in blue and the second term is in red. The reason why removing the rayleigh component of the signal from the data increases the speed of the convergence of the solution is that most of the measured signal from ALI is comprised of the rayleigh component, in red, by removing the rayleigh component of the signal the resulting in a measurement vector just primarily sensitive to aerosol. The ALI measurement vector is similar to the measurement vector used for the OSIRIS retrieval \citep{Bourassa2007,Bourassa2011} and all of the the measurement vectors for the 750~nm 0\si{\degree} horizontal line of sight form the mission can be seen in the right of \autoref{fig:5.3:measurementVectors}.

The error for a measurement vector is required to be able to determine the error on the final retrieved aerosol profile. To yield this term \autoref{eqn:5.3:measurementVector} is differentiated giving the following result

\begin{equation}

\delta\mathbf{y}^{2} = \left(\frac{\delta\mathbf{I}(\mathbf{z},\lambda)}{\mathbf{I}(\mathbf{z},\lambda)}\right)^{2} + \left(\frac{\delta I(z\_{ref},\lambda)}{I(z\_{ref},\lambda)}\right)^{2} + \left(\frac{\delta\mathbf{I}\_{rayleigh}(\mathbf{z},\lambda)}{\mathbf{I}\_{rayleigh}(\mathbf{z},\lambda)}\right)^{2} + \left(\frac{\delta I\_{rayleigh}(z\_{ref},\lambda)}{I\_{rayleigh}(z\_{ref},\lambda)}\right)^{2}.

\end{equation}

However, the only error considered is in the analysis is error due to the instrument, since the rayleigh component are modeled they are dropped from the error, simplifying the above result to

\begin{equation}

\delta\mathbf{y}^{2} = \left(\frac{\delta\mathbf{I}(\mathbf{z},\lambda)}{\mathbf{I}(\mathbf{z},\lambda)}\right)^{2} + \left(\frac{\delta I(z\_{ref},\lambda)}{I(z\_{ref},\lambda)}\right)^{2}.

\end{equation}

An example of this error on the measurement vector on image 208 is located on the middle panel of \autoref{fig:5.3:measurementVectors}.

A similar measurement vector is used for the forward model except base aerosol state or a priori, $\mathbf{x}$, is assumed for aerosol extinction profile and the both terms of the expression is calculated by SASKTRAN yielding the following

\begin{equation}

\mathbf{F}(\mathbf{x},\mathbf{b}) = \log\left(\frac{\mathbf{I}\_{mod}(\mathbf{z},\lambda)}{I\_{mod}(z\_{ref},\lambda)}\right)-\log\left(\frac{\mathbf{I}\_{rayleigh}(\mathbf{z},\lambda)}{I\_{rayleigh}(z\_{ref},\lambda)}\right)

\label{eqn:5.3:forwardModel}

\end{equation}

where $\mathbf{I}\_{mod}(\mathbf{z},\lambda)$ is the modeled radiance for the measurement and $I\_{mod}(z\_{ref},\lambda)$ is the measurement at the same reference altitude for normalization.

\begin{figure}

\includegraphics[width=1.0\textwidth]{./Images/5-3-MeasurementVectorsComparisons.pdf}

\caption[ALI Measurement Vectors]{Left: The measurement vector and two corresponding terms for \autoref{eqn:5.3:measurementVector} of the 0\si{\degree} line of sight of image 208. The black, blue, red curves represent the measurement vector, first term of \autoref{eqn:5.3:measurementVector}, and second term of \autoref{eqn:5.3:measurementVector}. Middle: Image 208 0\si{\degree} line of sight measurement vector with associated error represented by the shading. Right: A collection of all of the 0\si{\degree} line of sight measurement vectors at 750~nm during the mission.}

\label{fig:5.3:measurementVectors}

\end{figure}

## 5.3.2 Aerosol Extinction Retrievals

Test

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The above retrieval method was applied to a complete cycle of ALI spectral images (number 204-216 of the balloon mission). The retrieved aerosol extinction profiles can be seen in the left panel of Figure 12. Note the log scale. After the retrieval, the difference between the measurement and forward model vectors were less than 2% for the majority of the retrieval region, approximately 13 km to 28 km, across all wavelengths. Note the behavior of decreasing extinction with increasing wavelength as expected due to the dependence of the cross section with respect to particle size.

The ALI 750 nm aerosol extinction profile is shown in the right panel of Figure 12 in blue with the shading representing the precision of the retrieval. The error is strictly based on measurement error and neglects any model and atmospheric state errors. The green curve is the average 750 nm aerosol extinction profiles of the same five coincident OSIRIS scans used for the ozone profile. The retrieved extinction profiles from ALI and OSIRIS are within with the total retrieval uncertainty below 20 km. It is encouraging, however, that the instruments follow the same overall profile shape including the stratospheric layer and the steep increase below 15 km. However, the OSIRIS and ALI extinctions do not agree within error between 20 to 25 km. Aerosol is notoriously difficult to validate in remote sensing with various technique and instrument geometries, and yet the SAGE II, SAGE III and OSIRIS differences are generally below 20-30% up to 30 km (Bourassa et al., 2012b; Rieger et al., 2015) so the disagreement between OSIRIS and ALI from 20 to 25 km found here is somewhat puzzling. However, given the retrieved uncertainty, the OSIRIS profile is only outside the upper error bound of ALI by less than 10%. There are also several possible systematic errors not accounted for in the inversion including the choice of retrieval altitude ranges, particle size composition and distributions, stray light, and the high altitude aerosol load. This is also the first polarized limb scatter retrieval to our knowledge and so there may be further issues to explore with the polarized measurement and forward model. Regardless, the results are encouraging.

THESIS:

In order to be able to use the ALI measurement vectors in the MART method certain quantities were needed for the model; albedo, ozone concentration and cross sections, and aerosol cross sections. The albedo is required since an absolute radiance calibration was never preformed with ALI and the albedo cannot be determined directly from ALI's measurements, which is important since the amount of ground scatter in SASKTRAN-HR must correspond to the ground scatter during ALI's flight. Second, the ozone absorbtion features from the Chappuis band appear in the ALI measurements from 650 to 700~nm. The absorbtion must be accounted for to not artificially change the determined aerosol profiles. The ozone profiles were acquired from OSIRIS instrument using the five scans that were within 48 hours of the balloon flight and within 500~km of the launch facility were averaged together to be the ozone profile used in the SASKTRAN-HR model, with cross sections from \cite{Burrows1999}. The albedo is from the ADAM database which has monthly values for albedo over the surface on earth at a resolution of 0.1\si{\degree}~x~0.1\si{\degree} grid at 1~nm spectral resolution \citep{Muller2013}. The aerosol cross sections come from the Mie scattering derivation that was originally proposed by Mie and was implemented efficiently by \cite{Wiscombe1980}. For the purpose of the retrieval an a priori for particle was used with a mode radius of 0.08~$\si{\micro\metre}$ and a mode width of 1.6 in a log-normal distribution which is considered a standard size distribution for aerosol \citep{Deshler2003}.

The complete mission consisted of 216 images that were recorded in illuminated conditions. The MART retrieval method was run on a select set for the purpose of the analysis, specifically the last complete set of images from 650 to 950~nm consisting of images 204-216. They were chosen due to being the last in the mission and the sun was the highest in the sky giving the brightest atmosphere leading to the best signal to noise during the mission. A sample of the retrievals can be observed in \autoref{fig:5.3:AliRetreivals} which highlights the 725, 825, and 925~nm retrievals for images 206. The left panels shows the measurement vector from ALI in black with the forward model radiance profile from SASKTRAN-HR in blue. For each of the wavelengths, the algorithm determines altitudes where the value of the measurement vector is less than the known noise and does not allow aerosol to be retrieved in those regions. Instead the scaling factor, given by $\alpha = yF^{-1}$ is scaled to the aerosol profile above and below the last retrieved point to keep the aerosol profile smooth, as discontinuities are nonphysical and can lead to a nonphysical results in the MART algorithm. The middle panel shows the convergence between the measurement vector and the forward model result. For the full range of the wavelengths a difference of less than 1\% is seen from 12 to 24~km with a few outliers.

\begin{figure}

\includegraphics[width=1.0\textwidth]{./Images/5-3-AliRetreivals.pdf}

\caption[ALI Aerosol Retrievals]{An example of three aerosol retrievals from images 206, 208, and 214, with center wavelengths of 700, 750, and 900~nm respectively vertically displayed in the figure from top to bottom. The left column shows the measurement vector, $y$, in black with the retrieved forward model, $F$, in blue. The center column shows the ratio of the $y$ over $F$ known as $\alpha$ and is the convergence factor between the ALI measurement and the forward model. For both the first two columns the black line is barely viable due to the very good agreement of the forward model. The final column is ALI aerosol extinction in blue with the associated error represented by the light blue shading.}

\label{fig:5.3:AliRetreivals}

\end{figure}

The complete set of aerosol extinction profiles for the last cycle can be seen in the left panel of \autoref{fig:5.3:AliAerosolCycle} fromt he 0\si{\degree} horizontal line of sight. From the retrieval altitudes the aerosol extinctions do not show then expected decrease as the wavelength increases. The extinction is the number density multiplied by the scattering cross section, theoretically the number density should not change with respect to wavelength but the scattering cross section, according to Mie theory, should decrease with wavelength. However since a particle size distribution is assumed, and not determined, there will be a error associated with these extinctions and a method to determined particle size will be discussed in \autoref{sec:5.4:ParticleSizeDetermineation}.

With the successful aerosol retrieval from the mission the question that remains is how do these results compare to other aerosol extinction measuring instrument. Two other instruments are used to compare with the ALI results, the satellite based OSIRIS and stratospheric balloon based SALOMON. The ALI aerosol profile for the 750~nm aerosol extinction is shown in blue with the shading representing the error for the retrieval on the right of \autoref{fig:5.3:AliAerosolCycleNoPartSize}. The error is strictly based on measurement error and neglecting any model and atmospheric state errors. The green curve is the average 750~nm aerosol extinction profiles of the same five OSIRIS scans used for the ozone profile and OSIRIS uses a retrieval method for aerosol that is very similar to the method used for ALI \citep{Bourassa2007,Bourassa2011}. The red is the 750~nm aerosol extinction from SALOMON \citep{Berthet2002} which was balloon launched from the Timmins balloon base as the Nimbus 5 mission on September 12, 2014. The aerosol extinction for ALI and OSIRIS are within agreement of each other within the error for a majority of the 750~nm profile with some disagreement between 18 to 24~km. It also should be noted that the three instruments follow the same overall profile shape. First, a bend in the profile occurs at approximately 25~km, then increases approximately linearly until 15~km where aerosol extinction leaves the linear trend and forms the peak of the measurement. The agreement in shape is an excellent result for the ALI mission since it shows good sensitivity to aerosol.

\begin{figure}

\includegraphics[width=1.0\textwidth]{./Images/5-3-FullAerosolCycleComparisonNoPartSize.pdf}

\caption[Aerosol Retrievals for All Wavelengths and 750~nm Comparison with OSIRIS and SALOMON]{Left is the retrieved aerosol extinction profiles from the last complete cycle consisting of images 205 to 216 from the 0.0\si{\degree} horizontal line of sight. Right is the 750~nm ALI aerosol extinction in blue with its error represented by the shading compared to the 750~nm extinction measured by OSIRIS and SALOMON in green and red respectively.}

\label{fig:5.3:AliAerosolCycleNoPartSize}

\end{figure}

As noted in \autoref{fig:5.3:AliAerosolCycleNoPartSize} an error was determined for the ALI retrievals and was determined in the following method. Start by preforming the retrieval on the measurement vector to yield the determined aerosol state. Next since iterative states always slightly change with additional iteration and need to be removed to not be considered part of the noise error on the analysis the aerosol profile is is reretrieved, $\mathbf{k}$. Then perturb the measurement vector at the first retrieved altitude and retretrieve the aerosol profile yielding $\mathbf{k}\_{1}$, and create the first column of the error matrix, $\mathbf{E}$, by takeing $\delta\mathbf{k}\_{1} = \mathbf{k}\_{1}-\mathbf{k}$ which removed the initial profile plus the precision error from reretieveing the profile. Repeat for the rest of the retrieved altitudes, $n$, yielding the error matrix

\begin{equation}

\mathbf{E} =

\left[ \begin{array}{ccccc}

\vdots & \vdots & \vdots & \vdots \\

\delta\mathbf{k}\_{1} & \delta\mathbf{k}\_{2} & \dots & \delta\mathbf{k}\_{n} \\

\vdots & \vdots & \vdots & \vdots

\end{array} \right].

\end{equation}

Each row is the error matrix is caused by a specific height of the retrieval, $j$, the total error on each profile height is given is given by

\begin{equation}

\epsilon\_{j} = \left( \sum^{m}\_{i} E\_{ij}^{2} \right)^{0.5}

\end{equation}

where $m$ is the total number of retrievals altitudes. Since the ALI measurements are normalized to a 1~km grid for ALI $\mathbf{E}$ is a square matrix.

## 5.3.3 Particle Size Retrieval Methodology

Test

PAPER:

Using the retrieved extinction profiles for the complete spectral range, we have attempted a determination of the Angström exponent using a method similar to that outlined by Rault and Loughman (2013) for the OMPS-LP analysis. In this method, the independently retrieved extinction profiles at each wavelength and altitude are fit with a straight line in log-wavelength, log-extinction space. The slope of this line corresponds to the Angstrom exponent. This is then used to find the best match to the spectral dependence of the Mie scattering cross section in order to update the particle size distribution. With only one piece of information, the mode-width of the log-normal distribution is fixed to 1.6 and the mode radius is updated. The extinction retrievals are then performed again at each wavelength and the process is iterated until the Angstrom exponent, corresponding to the determined mode radius, converges.

THESIS:

To rectify the problem with the dependance between the extinction and wavelength a method will be preformed on the data in order to try to extract some particle size information But first a look into the information of particle size from limb measurements must be explored. From work done by \citep{Rieger2014} shows that different particle size distributions can effect the aerosol measurement vector. He uses an OSIRIS scan geometry and calculates the respective measurement vectors for a series of particle sizes which can be seen in reproduction of \autoref{fig:5.4:AerosolMeasurementParticleSizeInfo}. In panel A three different log-normal distributions are used which give a near identical measurement vector at 750~nm with simulation of the atmosphere with SASKTRAN. The three profiles are in blue is a single fine mode particle size distribution with a mode radius and width of 0.08~\si{\micro\meter} and 1.6 respectively, black is a bimodal particle size distribution the simulates volcanic conditions with the mode radius and width of 0.08~\si{\micro\meter} and 1.6 for the fine mode and 0.4~\si{\micro\meter} and 1.2 for the coarse mode, lastly red is a representative size distribution with mode radius and width of 77~\si{\micro\meter} and 1.75. Panel B shows the measurement vectors calculated with the three distributions across a series of wavelengths. The third panel, panel C, shows the difference of the measurement vectors compared to the bimodal distribution. Sensitivity to particle size is only seen past 800~nm measurement but great sensitivity does not occur until measurement recorded out to 1500~nm. Furthermore an error of 1\% in the radiance yields a relative error in the bimodal distribution measurement vector shown by the gray shading.

\begin{figure}

\centering

\includegraphics[width=1.0\textwidth]{./Images/5-4-AerosolMeasurementParticleSizeInfo.pdf}

\caption[Aerosol Measurement Vectors Sensitivity to Particle Size]{Reproduced from \citep{Rieger2014}. For OSIRIS scan 6432001 aerosol measurement vectors were calculated at 22.5~km. (A) The three size distributions used in the study. (B) The measurement vectors calculated via the SASKTRAN simulation (C) The relative percent difference of the fine and representative distributions with respect to the bimodal distribution. A 1\% error is the radiance yields an uncertainty in the bimodal measurement vector shown by the grey shading.}

\label{fig:5.4:AerosolMeasurementParticleSizeInfo}

\end{figure}

For ALI measurement were only gathered between 650 and 950~nm in wavelength, due to the CCD camera not having the required efficiency sample at wavelengths above this point. As such ALI only has some sensitivity to particle size form the longer wavelengths and only a small amount. As such it is not possible to determine both the mode radius and mode width due to information poor wavelengths. Instead the data from ALI will be used to determine the Angstr\"{o}m exponent. The Angstr\"{o}m exponent is an approximation to Mie scattering since the value of the Angstr\"{o}m exponent, $\alpha$, is an analogy to particle size. The scattering cross section for aerosol is dependant the particle size distribution assuming since the number density must be a constant across wavelength, as such the extinction value changes with a change in cross section. From \autoref{eqn:5.4:agstromCoefficient} the particle size profile from a series wavelength can be determined from the Angst\:{o}m exponent.Lower Angstr\"{o}m exponents correspond to larger particle sizes and vice versa for small particle sizes. Since the measurements observe relatively the same atmosphere over the time of one complete aerosol cycle; the differences between extinction ratios at the different wavelength can be used to gather a understanding of aerosol particle size in the form

\begin{equation}

\frac{n\sigma}{n\_{0}\sigma\_{0}} = \left(\frac{\lambda}{\lambda\_{0}}\right)^{-\alpha}

\label{eqn:5.4:agstromCoefficient}

\end{equation}

where $n$ is the aerosol concentration, and $\sigma$ is the scattering cross section. For the 750~nm wavelength for a variety of mode radius and widths a change in the cross section can be observed in \autoref{fig:5.4:scatteringCrossSections}. The Angstr\"{o} exponent can also be determine fitting a line through a series of points by rearangeing \autoref{eqn:5.4:agstromCoefficient} into the following

\begin{equation}

\alpha = -\frac{\log(n\sigma)-\log(n\_{o}\sigma\_{o})}{\log(\lambda)-\log(\lambda\_{o})}

\label{eqn:5.4:agstromCoefficientSlope}

\end{equation}

which show the Anstr\"{o}m exponent in simple the slope of the log of the extinction over the log of the wavelengths.

\begin{figure}

\centering

\includegraphics[width=1.0\textwidth]{./Images/5-4-CrossSectionDependance.pdf}

\caption[Log of Scatting Cross Section for Various Mode Radii and Widths]{Computed with the optical properties of the SASKTRAN engine. This varation of the corss section with respect to the mode radius and width allows for some determination of the particle size distribution through the Angstr\:{o} exponent.}

\label{fig:5.4:scatteringCrossSections}

\end{figure}

Once the retrieval has been preformed for a complete series of wavelengths; determination of the Angstr\"{o}m exponent occurs which is preformed in a similar manor as outlined by \cite{Rault2013} for the OMPS aerosol particle size retrieval. For the retrieval described here a single mode log-normal distribution is assumed with a mode radius, $r\_{g}$, and the mode width, $\sigma\_{g}$, which is the same as the OSIRIS version 5.07 aerosol product of 0.08\si{\micro\meter} and 1.6 respectively. The process for the retrieval is as follows:

\begin{enumerate}

\item Set up a MART retrievals for all wavelengths measured by ALI was a mode radius and width of 0.08\si{\micro\meter} and 1.6 respectively.

\item Preform the MART retrieval for each wavelength.

\item For all wavelengths determine the maximum minimum retrial altitude and minimum maximum retrieval altitude. These will be the bound for for the particle size retrievals.

\item For each valid tangent height fit a least squares line through the extinctions and wavelengths via \autoref{eqn:5.4:agstromCoefficientSlope} to determine the Angstr\"{o}m exponent. A wavelength at a altitude was rejected if the forward model at that shell altitude was not within 3\% of the measurement vector.

\item For each tangent height convert the Angstr\"{o}m exponent into a mode radius assuming a mode width of 1.6.

\item Take the median of all the determined mode radii and use this as the new values for the mode radius. If this value is the same as the previous iteration stop the retrieval as the solution has converged, If it is different change the mode radius in the model to the new mode radius and repeat the sequence from step 2.

\end{enumerate}

This method generally achieved convergence within 4-5 iterations.

## 5.3.4 A Sample Particle Size Retrieval

Test

PAPER:

The particle size method outlined above was also applied to this measurement set. The retrieved extinction at a given altitude was rejected from the straight line fit if the converged forward model radiance at that altitude was not within 2% of the measurement vector. In the case shown in Figure 13, at the 14.5 km altitude point, only 10 of the 13 possible wavelengths contributed to the determination of the Angström exponent. The first panel of Figure 13 shows the median Angström exponent that was determined after each iteration and convergence can be seen after a couple iterations. The results are shown in the second panel of Figure 13, where the Angström exponent is between 2 and 3 throughout the altitude range from 13 to 22 km. Assuming a mode width of 1.6 yields a median mode radius of 0.077 µm. In comparison to typical levels of background aerosol from the Laramie, Wyoming OPC data (Deshler et al., 2003) the retrieved particle size parameters are certainly within an expected range, although there is a relatively large error bar on the retrieved value, limiting the usefulness of the retrieved particle size information for background aerosol. However, with these error bars, even this limited spectral range would have the sensitivity to detected particle size changes as seen by OSIRIS and SAGE II over recent decades due to small volcanic perturbations (Rieger et al., 2014).

THESIS:

A sample particle size retrieval is shown in \autoref{fig:5.4:ParticleSize} for the cycle consisting of images 204-216. The first panel of \autoref{fig:ParticleSize} shows the median Angstr\"{o}m exponent that was determined after each iteration and convergence can be seen after a couple iterations. The initial guess for the mode radius was changed to several different values and each run of the model resulted in the same final retrieved mode radius. The particle size determined for ALI in the last complete set of aerosol images can be seen in the second panel of \autoref{fig:ParticleSize} which yields a final Angstr\"{o}m exponent of between 2 and 3 throughout the altitude range from 13 to 22~km. Assuming a mode width of 1.6 yields a median mode radius of 0.105~\si{\micro\metre}. The final panel shows the least squares for the 20.5~km tangent altitude for this case only 11 of the 13 possible wavelengths contributed to the determination of the Angstr\"{o}m exponent.

\begin{figure}

\centering

\includegraphics[width=0.5\textwidth]{./Images/5-4-ParticalSize.pdf}

\caption[Particle Size Retrieval]{The top panel shows the convergence of two sample particle size retrievals over the iterations, blue and red represent a priori of 0.08 and 0.12~\si{\micro\metre} respectively. Both initial state converge to the same value over approximately 4 iteration in the particle size retrieval method. The second panel is the final Angstr\"{o}m exponents determined for images 204-217 during the Timmins 2014 campaign and are in black as the final profile was almost identical after each particle size retrieval, the shading represents the error associated with the least squares fit. The last panel demonstrate a least squares fit to determine the Angstr\"{o}m exponent at 20.5~km shell altitude.}

\label{fig:5.4:ParticleSize}

\end{figure}

The error in the Angstr\"{o} exponent was determined looking at the error in the slope possible from the least squares method. Assuming no error in the extinction points the error in the slope is given by

\begin{equation}

\delta\alpha = \frac{s}{SS\_{xx}}

\end{equation}

where $s$ is

\begin{equation}

s = \sqrt{\frac{SS\_{yy}-\alpha SS\_{xy}}{n-2}},

\end{equation}

and $SS\_{xx}$, $SS\_{yy}$, and $SS\_{xy}$ are the sum of squares of the wavelength, aerosol extinction, and cross term between the wavelength and aerosol extinction and $n$ is the number of points. However there is and associated error with each aerosol extinction profile. To determine the error on the Angstr\"{o}m exponent accounting for the precision in the aerosol extinction the error on the slope was calculated 10,000,000 times with each iteration a random amount of error from he known range was added to each extinction point. Then the mean from the all the errors of the least squares fit was used as the fine precision estimate on the Angstr\"{o}m coefficient which can be seen in the second panel of \autoref{fig:5.4:ParticleSize}.

The extinction cycle similar to \autoref{fig:5.3:AliAerosolCycleNoPartSize} is shown in \autoref{fig:5.4:AliAerosolCycle} except with updated particle size parameters form the retrieval. the aerosol extinction profiles now shows the characteristic decrease in aerosol extinction that is expected as wavelength increase. this extinction is self consistent with Mie theory and give some validity to the retrieved particle sizes determined from ALI.

\begin{figure}

\includegraphics[width=1.0\textwidth]{./Images/5-4-FullAerosolCycleComparison.pdf}

\caption[Aerosol Retrievals with Corrected Particle Size for All Wavelengths and 750~nm Comparison with OSIRIS and SALOMON]{Left is the retrieved aerosol extinction profiles from the last complete cycle consisting of images 205 to 216 from the 0.0\si{\degree} horizontal line of sight. Right is the 750~nm ALI aerosol extinction in blue with its error represented by the shading compared to the 750~nm extinction measured by OSIRIS and SALOMON in green and red respectively.}

\label{fig:5.4:AliAerosolCycle}

\end{figure}

# 5.4 Results

Test

PAPER:

The ALI prototype, which is telescopic acousto-optic imager, has been used to successfully measure two dimensional spectral images of the atmospheric limb from stratospheric balloon. The observed radiances appear to be of high quality and show both vertical and horizontal features of the cloud and aerosol layers. Aerosol extinction coefficient profiles were retrieved from the ALI data that show reasonable agreement with OSIRIS satellite measurements.

No large scale issues were found with the instrument performance; however, some future changes would be recommended. First, an absolute calibration of the instrument would allow ALI to determine the effect albedo directly, as is done with OSIRIS. This would remove some the uncertainly in the model inputs and likely yield higher quality results. This is simply a matter of having access to the calibration equipment. Also, even with the baffle and the robust method of removing stray light with the cycling of the AOTF, some stray light was still observed in the obtained images. Impact and mitigation of this should be tacked in future iterations of the instrument.