

CHAPTER 1

Introduction

1.1 The Sun

Its significance, structure, and magnetic phenomena.

1.2 Magnetohydrodynamics

To build up a mathematical description of the Sun's plasma dynamics, let's motivate some assumptions.

The Sun's plasma, just like all matter in the Universe, is made up of particles¹, but the phenomena such as MHD waves that we are concerned with in this Thesis operate on a *macroscopic* level. By this we mean on length-scales much larger than the *mean free path*². This means that the *Knudsen number*, the dimensionless parameter defined by the ratio of the mean free path to a characteristic length scale, in the Sun is much less than unity. This motivates the **continuum assumption**, where the fluid is considered to “fill up” the space in which it is contained, so that small-scale inhomogeneities caused by particle dynamics are negligible. This gives us a coherent notion of fluid density, ρ and pressure, p .

We are gratefully gifted fundamental laws that are universally obeyed by classical mechanics systems upon which we can build our framework. These are the **conservation of mass**, *conservation of momentum*, and *conservation of energy*. The conservation of mass tells us that the change in density in a fixed infinitesimal volume is due only to mass entering or leaving the volume. Mathematically, we can equate the rate of change of density, $\partial\rho/\partial t$ to the incoming density flux $-\nabla(\rho\mathbf{v})$, *i.e.*

$$\frac{\partial\rho}{\partial t} + \nabla(\rho\mathbf{v}) = 0, \quad (1.1)$$

¹atoms or subatomic particles, depending on the temperature of its location in the Sun

²Approximately 1 cm - 1 km in the Sun. [REFERENCE](#)

where $\mathbf{v} = (v_1, v_2, v_3)$ is the velocity of the fluid.

1.3 Waves in the solar atmosphere

1.3.1 Magnetohydrodynamic waves in homogeneous plasma

1.3.2 Magnetohydrodynamic waves in homogeneous plasma

1.3.2.1 Interface

1.3.2.2 Symmetric slab

1.3.2.3 Cylinder

1.4 Thesis outline

Bibliography
