

# **Solar wind – Predicted magnetospheric influence and near-Sun environment**

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# **Solar wind – Predicted magnetospheric influence and near-Sun environment**

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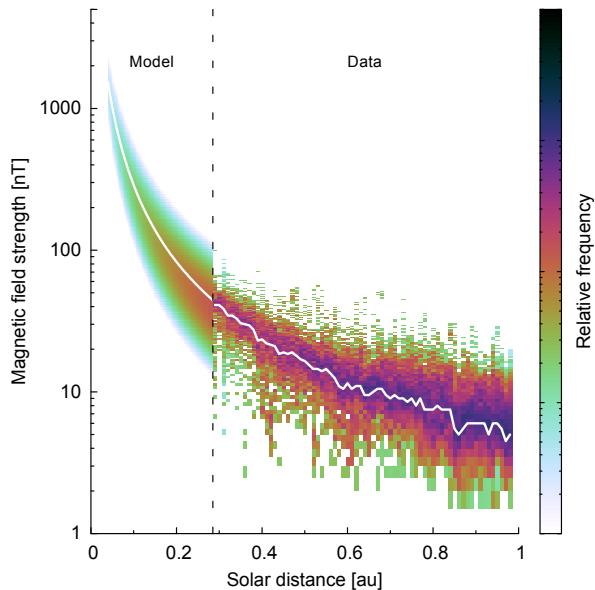
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*“Despite the ‘Dr.’ before his name, he had completed no course of study and received no degree. When people tried to pin him down about this, he would say that the letters were merely an abbreviation of his first name - Drummond - which he did not use. But it was as ‘Dr.’ Sam Laserowitz that he appeared in a number of science-fiction magazines; he was also known, in the circles of the fans of that genre, as a lecturer, and spoke on ‘cosmic’ themes at their many conferences and convention. Laserowitz’s speciality was earthshaking discoveries, which he happened upon two or three times a year. [...] We really have no idea what a multitude of con men and crackpots inhabit the domain that lies halfway between contemporary science and the insane asylum.”*

Excerpt from Stanisław Lem 1968, *His Master’s Voice* ([Lem & Kandel 1984](#), p. 38).



## **Abstract**

This thesis analyzes, and derives predictions of how strong solar wind and coronal mass ejections impact the terrestrial magnetosphere and how the solar wind environment scales in the near-Sun region. Near-Earth in-situ solar wind measurements, consisting of 35 years of minutely OMNI data, are analyzed together with the time series of the planetary geomagnetic disturbance indicator  $K_p$ . Correlations and functional dependencies are compiled with regard to nowcast the magnitude of the geomagnetic disturbances from in-situ solar wind measurements and to forecast them from remotely observed solar wind streams and CMEs. Further, 53 years of hourly OMNI data and sunspot number data are correlated for the purpose of deriving functional dependencies with the state of the solar cycle for the key solar wind parameters magnetic field strength, proton velocity, density, and temperature. Solar wind in-situ data from the Helios 1 and 2 missions, operational in the 1970s and orbiting the Sun within the distance range 0.3–1.0 au, are analyzed and empirical solar wind distance dependencies are derived. Additionally, in view of the planned near-Sun spacecraft mission Parker Solar Probe (PSP) with a planned launch in mid 2018, the solar wind environment is estimated down to PSP's closest perihelion at 9.86 solar radii.

# Contents

<b>1</b>	<b>Introduction</b>	<b>1</b>
<b>2</b>	<b>Background knowledge</b>	<b>3</b>
2.1	The Sun . . . . .	3
2.2	Solar dynamo . . . . .	6
2.3	Solar activity cycle . . . . .	7
2.4	Heliospheric magnetic field . . . . .	9
2.5	Solar wind . . . . .	11
2.5.1	Slow and fast streams . . . . .	12
2.5.2	Stream interaction regions . . . . .	14
2.5.3	Heliospheric current sheet . . . . .	15
2.5.4	Coronal mass ejections . . . . .	15
2.6	Space weather . . . . .	19
2.7	Magnetosphere . . . . .	22
2.7.1	Solar wind coupling mechanisms . . . . .	23
2.7.2	Dungey convection cycle . . . . .	24
2.7.3	Russell-McPherron effect . . . . .	25
2.7.4	Geomagnetic storms . . . . .	25
2.7.5	Geomagnetic indices . . . . .	26
2.7.6	<i>Kp</i> index . . . . .	26
2.8	Forecast methods . . . . .	26
2.8.1	<i>Kp</i> forecast . . . . .	26
2.8.2	CME forecast to Earth . . . . .	26
2.8.3	Stream forecast to Earth . . . . .	26
2.9	Magnetosphere . . . . .	26
2.9.1	Magnetopause reconnection . . . . .	26
2.9.2	Russell-McPherron effect . . . . .	27
2.9.3	Solar wind coupling . . . . .	28
2.9.4	Geomagnetic storms . . . . .	29
2.9.5	Geomagnetic indices . . . . .	30
2.9.6	<i>Kp</i> index . . . . .	30
2.10	Forecast methods . . . . .	31
2.10.1	<i>Kp</i> forecast . . . . .	31
2.10.2	CME forecast . . . . .	31
2.10.3	Stream forecast . . . . .	33
2.11	notes... . . . . .	33
<b>3</b>	<b>Instrumentation and data description</b>	<b>35</b>
3.1	Magnetometer . . . . .	35
3.2	Plasma detector . . . . .	36
3.3	<i>Kp</i> data . . . . .	36
3.4	Sunspot number . . . . .	38
3.5	OMNI data collection . . . . .	39
3.6	Helios probes . . . . .	39
<b>4</b>	<b>Solar wind and CME influence on the magnetosphere</b>	<b>43</b>
4.1	Introduction . . . . .	43
4.2	Long-term variations of the <i>Kp</i> index . . . . .	44
4.2.1	Solar activity influence . . . . .	44
4.2.2	Seasonal variations . . . . .	46
4.3	<i>Kp</i> nowcast from in-situ solar wind measurements . . . . .	47
4.3.1	Coupling function . . . . .	47
4.3.2	Data correlation . . . . .	48

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4.3.3	Functional dependency for solar wind . . . . .	48
4.4	<i>K<sub>p</sub></i> forecast from remote CME and stream observations . . . . .	50
4.4.1	Velocity forecast methods . . . . .	50
4.4.2	Solar Wind Structures list . . . . .	51
4.4.3	Data correlation . . . . .	52
4.4.4	Functional dependency for CME velocity . . . . .	53
4.4.5	Functional dependency for stream velocity . . . . .	55
4.5	Discussion and summary . . . . .	55
4.6	Applications and outlook . . . . .	57
4.7	notes... . . . .	57
<b>5</b>	<b>Empirical solar wind model for the inner heliosphere</b>	<b>59</b>
5.1	Introduction . . . . .	59
5.2	Parker Solar Probe mission . . . . .	59
5.3	Paper content . . . . .	59
5.4	Alternative magnetic field solar distance dependency . . . . .	60
5.5	Discussion . . . . .	64
5.6	Outlook . . . . .	65
5.6.1	flux conservation . . . . .	65
5.6.2	Seasonal solar wind variations . . . . .	66
5.6.3	Latitude dependency . . . . .	67
5.6.4	Radial evolution of solar wind structures . . . . .	68
5.7	notes... . . . .	68
<b>6</b>	<b>Paper: Solar-wind predictions for the Parker Solar Probe orbit</b>	<b>71</b>
	Abstract . . . . .	71
6.1	Introduction . . . . .	71
6.2	Frequency distributions of the solar-wind parameters . . . . .	72
6.3	Solar activity dependence of the solar-wind frequency distributions . . . . .	73
6.4	Solar distance dependency . . . . .	76
6.5	Empirical solar-wind model . . . . .	79
6.6	Model extrapolation to PSP orbit . . . . .	80
6.7	Discussion and summary . . . . .	82
	References . . . . .	83
<b>7</b>	<b>Summary and outlook</b>	<b>85</b>
7.1	Summary . . . . .	85
7.2	Outlook . . . . .	85
<b>A</b>	<b>Appendix</b>	<b>87</b>
A.1	Solar surface differential rotation . . . . .	87
A.2	Electromagnetism . . . . .	88
A.3	Plasma beta . . . . .	88
A.4	Alfvén velocity . . . . .	89
A.5	Sun-Earth orbit geometry . . . . .	90
A.6	GSE, GSM, and HGI coordinate systems . . . . .	91
A.7	Lognormal distribution . . . . .	92
A.8	Astronomical constants . . . . .	93
	<b>References</b>	<b>95</b>
	<b>Acknowledgements</b>	<b>107</b>
	<b>Curriculum vitae</b>	<b>109</b>



# 1 Introduction

Background -> problems -> motivation -> methods -> results

The abundance of modern technological systems that are sensitive to disturbances in the terrestrial magnetic field increases.

These include critical systems, such as ..., whose potential disruptions would not only have severe economical impacts but would affect human lives as well.

The disturbances in the terrestrial magnetic field (geomagnetic storms) are evoked by the field's interaction with streams of particles emitted by the Sun (solar wind).

Variations in some solar wind parameters lead to a direct response in geomagnetic activity.

This work aims to quantify this response for different solar wind forecast situations.

For when magnetic field and velocity information is available and for when only velocity information is available and it is known whether it is ambient solar wind or a CME event.

Open key questions are which mechanisms heat the corona and accelerate the solar wind.

The PSP mission is intended to clarify this.

As it is the first s/c to fly that near to the Sun, the solar wind environment is not known.

This is the motivation for the second part of this work.

This work derives estimates of that solar wind environment.

thesis goal: finding more precise relationships between parameters/quantities for being able to make better forecasts

catch from chapter abstracts...

This thesis presents quantitative analyses of the solar wind impact strength on the terrestrial magnetosphere and of the estimated near-Sun solar wind environment.

I structured this work as follows: The fundamentals about the Sun, its activity, solar wind, and space weather are laid in [Chapter 2](#). The instrumentation and data sources used in this work are described in [Chapter 3](#). The influences on the  $K_p$  index from solar activity, solar wind, CMEs, and solar wind streams are analyzed in [Chapter 4](#). In [Chapter 5](#), which is followed by the published paper on the same topic (integrated as [Chapter 6](#)), an empirical solar wind model for the inner heliosphere is developed and used to estimate the near-Sun solar wind environment for the planned PSP mission. [Chapter 7](#) summarizes the results and gives an outlook for further studies. Useful equations and information are located in the [Appendix A](#).



## 2 Background knowledge

This chapter summarizes the basic knowledge necessary for understanding the analyses performed in this work. First, the Sun's origin, inner structure, atmosphere and heliosphere are described. Then, the Sun's dynamics with its differential rotation and magnetic field generation are outlined. Further, the solar activity cycle is described, including the meridional flow circulation, appearance of active regions, the surface magnetic field change, and sunspot cycles. The heliospheric magnetic field is pictured from its photospheric emergence in magnetic bright points and coronal superradial expansion, through the formation of the heliospheric current sheet and the Parker spiral to the heliosheath. The solar wind and its properties, the origins of slow and fast streams, stream interaction regions, and coronal mass ejections are described. Furthermore, space weather, solar influence on Earth, the magnetosphere, solar wind-magnetosphere coupling, and geomagnetic indices are portrayed.

### 2.1 The Sun

13.8 billion years ago the Big Bang formed our universe. The energy density of our universe consists of 69.1 % dark energy, 25.9 % dark matter and 4.9 % baryonic matter, according to calculations using the inflationary  $\Lambda$ CDM<sup>1</sup> cosmology together with the latest cosmic microwave background temperature measurements (Planck Collaboration et al. 2016). After a few minutes the primordial nucleosynthesis left the universe in a state where the baryonic matter was composed of 75.33 %<sup>2</sup> hydrogen, 24.67 % helium and traces of deuterium, tritium and lithium (Planck Collaboration et al. 2016).

Over millions of years this gas cooled down and gravitationally accreted into molecular clouds and formed stars. The first generations of stars (Population III) fused this gas to heavier elements (metals) and supernovae distributed them into space as a foundation for the formation of new stars of low and high metallicity (Population II and I). Likewise, supernovae of these stars constantly enriched the interstellar medium with metals. Now, the interstellar medium in the Milky Way consists of about 32 % helium and traces of other metals (Daniger 1970).

Our Sun, a metal-rich Population I yellow dwarf star, emerged 4.6 billion years ago (Bahcall et al. 1995) from an accretion disk, formed by a collapsing rotating cloud. The compression within its center resulted in high temperatures, which initiated the fusion of hydrogen to helium (primarily pp chain reaction). The fusion reactions produce huge amounts of energy and heat the solar center to a temperature of 15.7 million kelvins (Christensen-Dalsgaard et al. 1996). The generated energy is transported through the solar body to its surface and eventually into space. The core region extends to about 0.25 solar radii ( $R_\odot$ ), where the declining temperature becomes insufficient for fusion reactions. The energy transport is dominated by thermal radiation until, because of declining ionization and density, at  $0.71 R_\odot$  up to the surface convective motion takes over (Christensen-Dalsgaard et al. 1991).

The temperature at this transition region, called tachocline, is about 2 million kelvins and decreases up to the solar surface to between 4400–6600 K<sup>3</sup>. Here at the photosphere, the energy is radiated away with an effective black body temperature of 5772 K (Mamajek et al. 2015), classifying the Sun as a spectral type G2V star. At this surface layer, granules, the tops of convection cells, and temporary sunspots are visible. Strong magnetic flux inhibits the convection at sunspots, leading to lower temperature and brightness (for more on sunspots see the following Sections 2.2 and 2.3). Figure 2.1 illustrates these photospheric features along with the inner solar structure.

Above the photosphere at the base of the chromosphere, the temperature declines to its solar minimum of 3800 K until it raises to 2–3 million kelvins in the corona (Billings 1959; Liebenberg et al. 1975). Up to now it is not fully understood why the corona is so much hotter than the underlying chromosphere – this question is known as the coronal heating problem (Klimchuk 2006; McComas et al. 2007; Fox et al. 2015). The generally considered energy transfer mechanisms are magnetic reconnections, wave heating and type II spicules or a combination of these (Cranmer et al. 2017).

<sup>1</sup> $\Lambda$ CDM: Lambda cold dark matter

<sup>2</sup>Percentages by mass.

<sup>3</sup>NASA Sun fact sheet: <https://nssdc.gsfc.nasa.gov/planetary/factsheet/sunfact.html>



**Figure 2.1** Image of the photosphere from 20 March 2016 together with a schema of the solar interior structure. The inset shows the granular surface with a sunspot. I created this figure based on a SDO/HMI continuum image, credit: NASA/SDO and the AIA, EVE and HMI science teams.

The chromosphere is a 2000 km thick region, whose features – numerous spicules, filaments, and prominences – can reach far into the corona. They consist of chromospheric material, channeled by the solar magnetic field, and are enveloped by a thin transition region where the temperature jumps up from about 30 000 K<sup>4</sup> to coronal temperatures. Reconnection of magnetic field lines can result in the eruption of filaments into the corona and beyond, termed coronal mass ejections (CMEs), see also [Subsection 2.5.4](#). Details of chromospheric features are shown in [Figure 2.2](#) – the images were taken on the same day as in [Figure 2.1](#).



**Figure 2.2** Composite image of the solar atmosphere from 20 March 2016 and some details of its features. Corona, chromosphere and photosphere are seen in wavelengths of 193 Å, 304 Å, and continuum. Chromospheric spicules are visible on the northern limb. The enlargements on the right show a prominence and a filament. The dark region at the south pole is a coronal hole. The left inset shows details of the active region belonging to the sunspots shown in [Figure 2.1](#). I created this figure based on SDO/AIA images, credit: NASA/SDO and the AIA, EVE and HMI science teams.

The Sun's atmosphere is dominated by the varying small- and large-scale solar magnetic field configuration. There are regions where the magnetic field lines arc back to the surface and regions with open field lines. In the latter areas the coronal plasma can – guided by the field – escape into space. Thus these coronal areas are less

<sup>4</sup>NASA Sun fact sheet: <https://nssdc.gsfc.nasa.gov/planetary/factsheet/sunfact.html>

dense, cooler and therefore appear darker in extreme ultraviolet (EUV) and are called coronal holes (CHs). In Figure 2.2 a coronal hole is visible at the solar south pole.

From Earth, the faint corona and chromosphere can only be observed during eclipses, because of the brightness of the solar disk. There are three effects contributing to the visibility of the corona: photons scattering off free electrons, producing a continuous spectrum; photons scattering off dust particles, their spectrum contains Fraunhofer absorption lines; and ion spectral emission lines – these contributions to the corona are termed K-, F- and E-corona<sup>5</sup>. Images from solar eclipses reveal the coronal plasma, shaped by the magnetic field, and red prominences from the chromosphere. The solar eclipse imaged in Figure 2.3 shows the magnetic field's dipole structure and the equatorial streamer belt, characteristic for a quiet Sun during cycle minimum.



**Figure 2.3** Total solar eclipse image of the inner corona up to a distance of five solar radii. The picture was taken in Mongolia, 1 August 2008 and is processed from multiple images. Credit: Miloslav Druckmüller, Peter Aniol, Jan Sládeček, 2008, reproduced with permission.

Due to the high coronal temperatures, plasma escapes the solar gravitational field (Parker 1958) with velocities of  $200\text{--}800\text{ km s}^{-1}$ . Its acceleration is linked to the coronal heating – however, the exact location and process remain an open question (Hollweg 1985; McComas et al. 2007; Fox et al. 2015; Cranmer et al. 2017). The magnetic field becomes too weak to guide the coronal plasma at a distance of a few solar radii. From this so-called source surface, the solar wind flows radially outward into space until it reaches the termination shock. Eventually it collides with the local interstellar medium, creating the boundary of the heliosphere, the heliopause. The heliopause is expected to be a bubble of either teardrop or croissant shape, caused by the Sun's relative velocity of  $23\text{ km s}^{-1}$  with respect to the local interstellar medium (Owens & Forsyth 2013; Opher et al. 2015). Thus, there may exist a leading bow shock. Measurements of the Voyager 1 and 2 spacecraft indicate their passage of the termination shock at about 94 au and 84 au, entering the heliosheath region (Owens & Forsyth 2013). Gurnett et al. (2013) report that in 2012 Voyager 1 actually crossed the heliopause into interstellar space at a solar distance of 121 au. The heliosphere and its surrounding flow structure is illustrated in Figure 2.4.



**Figure 2.4** Schema of the heliosphere and its surrounding flow structure, formed by the interaction of the solar wind (red) with the local interstellar medium (blue) at the heliopause. Credit: Owens & Forsyth (2013, Fig. 9), licensed under CC BY-NC 3.0 DE.

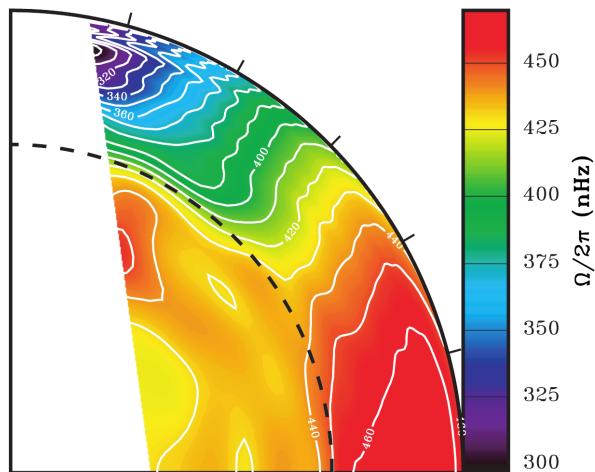
On its way outwards through the solar system, the solar wind – carrying the solar magnetic field – interacts

<sup>5</sup>K from kontinuierlich (continuous in german), F from Fraunhofer, and E from emission.

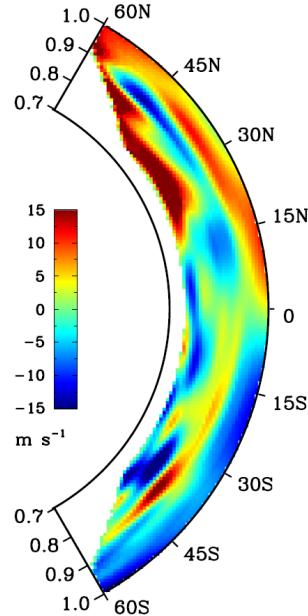
with the planets, their magnetic fields and other solar system bodies. This has various effects, for instance disturbances in planetary magnetic fields with appearance of aurorae and enhanced radiation, atmospheric losses and stripping of cometary tails. Some of these effects can have disruptive consequences for humans and their technology, creating a high interest in understanding space weather and forecasting its effects, the topic of space weather is further addressed in [Section 2.6](#). The magnitudes of these effects depend highly on spatial and temporal variations in the solar wind, which are rooted in the dynamics of the solar magnetic field, described in the following sections.

## 2.2 Solar dynamo

The conservation of the angular momentum in the contracting molecular cloud led to a rotation of the Sun. Although the Sun experiences loss of angular momentum due to solar wind ([Weber & Davis 1967](#)), its rotation still has a current average period of about 25 days. The radial convective motion within the solar interior above the tachocline leads to a transport of momentum away from the rotation axis and therefore to a slower polar and faster equatorial rotation in the convection zone ([Miesch 2005](#)). This differential rotation is visible on the surface and was first discovered from sunspot observations by [Scheiner \(1630\)](#). With a rotation period of about 34 days, the poles have a lag of almost 9 days (for further information on solar rotation see appendix [Section A.1](#)). The differential rotation in the solar interior can be inferred from helioseismological observations. Below the differential rotation of the convection zone, a nearly solid rotation with a period of about 26.6 days (this corresponds to a frequency of 435 nHz) exists in the radiation zone, as shown in [Figure 2.5](#).



**Figure 2.5** Rotation frequency profile of the solar interior. The location of the tachocline is indicated by the dashed line. The rotation frequency is inferred from helioseismology via observations from the Michelson Doppler Imager (MDI) at the Solar and Heliospheric Observatory (SOHO) spacecraft. Credit: [Thompson et al. \(2003, Fig. 3\)](#), © Annual Reviews, reproduced with permission.



**Figure 2.6** Meridional flow velocity profile in part of the convection zone. Positive values are directed towards north. The velocity is inferred from helioseismology via observations from the Helioseismic Magnetic Imager (HMI) at the Solar Dynamics Observatory (SDO) spacecraft. Credit: [Zhao et al. \(2013, Fig. 4, panel \(a\), I moved the colorbox\)](#), © AAS, reproduced with permission.

Turbulent plasma motions from convective flows in the convection zone generate and carry disorganized magnetic flux. The large rotational shear at the tachocline stretches and amplifies the magnetic fields to strong coherent toroidal flux ( $\omega$ -effect) with intensities of the order 1–10 T. These toroidal fields, generated near the bottom of the convection zone, can be stored in a deep magnetic layer located in the stably stratified region below the convection zone ([Ossendrijver 2003](#)). The stronger flux ropes become buoyant and raise to the surface. The Coriolis force twists them systematically on their way through the convection zone ( $\alpha$ -effect). The twist is stronger at higher latitudes (Joy's law). The flux ropes emerge then in the photosphere as bipolar active regions of opposite magnetic polarity – the stronger ones forming pairs of sunspots, as seen in [Figure 2.7](#). Turbulent convective diffusion of this surface flux contributes to the build-up of poloidal fields. Their resulting



**Figure 2.7** Continuum image of the two sunspots pictured in [Figure 2.1](#) (top left), magnetogram from the same region (bottom left), and magnetogram from the whole solar disk (right). The magnetogram shows the polarity of the line-of-sight magnetic field component at the photosphere (black/white: inward/outward polarity). The highly concentrated magnetic flux at the sunspots is visible as well as the extended bipolar magnetic field structure of the whole active region, which is divided by the so-called magnetic neutral line. The solar disk is scaled to the same size as in [Figure 2.1](#). I created the figure based on SDO/HMI continuum and magnetogram images from 20 March 2016, credit: NASA/SDO and the AIA, EVE and HMI science teams.

polarity is opposite to the prevailing global field due to the directional way the rotational shear at the tachocline and the Coriolis force in the convection zone act. Fluctuating motions further amplify the mean fields in these processes. This solar  $\alpha\omega$ -dynamo is thought to create the major part of the solar magnetic field. Still, with regard to the magnetic field's high variability, the long-term mean fields are governed by intermittent localized structures, that is, active regions, filaments and coronal loops ([Miesch 2005](#)).

## 2.3 Solar activity cycle

Helioseismic measurements reveal that the large-scale convective flow is aggregated into large convection cells with slow meridional flows of a few  $m s^{-1}$ , as can be seen in [Figure 2.6](#). A poleward subsurface flow and equatorward backflow beneath with a further poleward flow below are detected within each hemisphere, comprising a stacked double-cell profile ([Zhao et al. 2013](#)). The meridional circulation flow speed has a major influence on the average 22-year period of the emerging magnetic flux at the solar surface. This period varies and is influenced by the stochastic emergence rate and tilts of active regions and the diffusion from random convective motions ([Hathaway & Upton 2016](#)). The surface magnetic field configuration changes within one period from a dipole structure to a reversed dipole structure with opposite polarity and back, completing a so-called Babcock-Leighton dynamo cycle. Thus, the transition time from one dipole state to the next lasts about 11 years, this period is defined as one solar cycle.

In the transition phase, magnetic flux emerges in belts above and below the solar equator, manifesting as bipolar active regions with sunspots, resulting in a toroidal/multipolar structured magnetic field. Sunspots appear at about  $\pm 20^\circ$  latitude at the beginning of a cycle, this shifts towards lower latitudes at the end of a cycle. Thus the plot of sunspots over latitude and time reveals a butterfly pattern ([Maunder 1904](#)). This butterfly pattern appears in surface radial magnetic field observations as well, see [Figure 2.8](#). The leading polarity of bipolar regions is opposite in both hemispheres and the leading polarity changes with each solar cycle, this is called Hale's polarity law. The emerging flux is carried by the slow meridional surface flow poleward, canceling the current dominating polar field polarity and eventually resulting in the polar field switch ([Hathaway 2015](#)).

Since regions of strong magnetic flux are visible as sunspots on the photosphere, they were known well

## 2. Background knowledge

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**Figure 2.8** Magnetic butterfly diagram of the longitudinally averaged radial magnetic field on the solar surface. Yellow represents an outward directed magnetic field (positive), blue inward (negative). The data is obtained from instruments on Kitt Peak National Observatory and from the MDI at the SOHO spacecraft. Courtesy of David Hathaway, [Solar Cycle Science](#), 2018, updated version of [Hathaway \(2015, Fig. 17\)](#).

before the common era by greek and chinese scholars ([Clark & Stephenson 1978](#); [Vaquero 2007](#)). Systematic sunspot observations exist since 1610, shortly after the invention of the telescope. In 1843 Schwabe discovered the 11-year periodicity in the sunspot occurrence ([Schröder 2004](#), p. 124). In 1848 Wolf introduced the sunspot number (SSN) and the solar cycle number to record these cycles ([Hathaway 2015](#)). The SSN observations, see [Figure 2.9](#), show large variations in cycle length (9–14 years) and cycle amplitude, with peak SSNs in the range 0–300, ([Hathaway 2015](#)). There also exist long-term variations, such as secular cycles of different



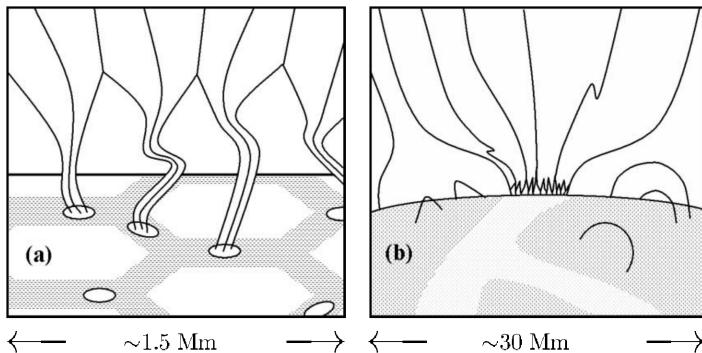
**Figure 2.9** Monthly mean sunspot number (blue) and 13-month smoothed monthly sunspot number (red) since 1954. Credit: [SILSO data/image](#), Royal Observatory of Belgium, Brussels, 2018.

periodicity or the 70-year Maunder Minimum, during which from 1645 on almost no sunspots were observed, ([Maunder 1890](#)). The source of the variations in the solar cycle periods and amplitudes are variations in the meridional circulation, because their fluctuations are larger than those found in the differential rotation and in the convective motions ([Hathaway 2015](#)).

As the SSN is commonly used as an indicator for solar activity, there exists interest in its prediction for the course of the actual and upcoming solar cycles. The continuing prediction of an already commenced activity cycle is reliable, but then the prediction of a cycle before it began is more difficult. Though, there are indications that the polar magnetic field strength during the preceding activity minimum is correlated to the strength of the next solar cycle ([Schatten & Sofia 1987](#)). However, [Hathaway & Upton \(2016\)](#) suggest that the predictability of solar cycles is generally limited – accumulated uncertainty produced by stochastic motions in the convection zone makes predictions further than the next solar cycle very unreliable.

## 2.4 Heliospheric magnetic field

In the quiet Sun during solar cycle minimum, open coronal regions – coronal holes – are the photospheric sources of the heliospheric magnetic field (HMF). Bright points between the granules on the photosphere are detected in G-band (430 nm) images. They are identified as magnetic flux tubes with field strengths of 100–200 mT (Cranmer & van Ballegooijen 2005). Together, these magnetic bright points (MBPs) cover around 1–2 % of the solar surface and carry many times the flux that active regions do (Sánchez Almeida et al. 2010). These thin flux tubes expand laterally in the low chromosphere and merge to homogeneous network fields, which expand and merge again to a large-scale canopy below the transition region (see Figure 2.10).



**Figure 2.10** Schemata of superradially expanding magnetic flux. (a) MBPs between granules on the photosphere are indicated by ellipses. The protruding lines are thin magnetic flux tubes that merge to a homogeneous network field. (b) Pictured is the network field which expands again to the large-scale canopy field of the lower corona. Credit: Cranmer & van Ballegooijen (2005, Fig. 1, panels (a) and (b)), © AAS, reproduced with permission.

The MBPs' convective appearance and stochastic motions on the photosphere result in wavelike fluctuations that propagate upward through the superradially expanding flux tubes. There exist three types of magnetohydrodynamic (MHD) waves within the plasma: compressional fast- and slow-mode waves, and an incompressible wave mode, which is the result of bending magnetic field lines (Alfvén 1942), called shear Alfvén wave. Alfvén waves propagate with a characteristic speed along magnetic field lines. As they transport energy from the photosphere outwards, it is assumed that they play a major role in the coronal heating process and that the solar wind is accelerated up to the so-called Alfvén critical surface at around  $17 R_\odot$ , where the local Alfvén speed equals the solar wind speed (Sittler & Guhathakurta 1999; Exarhos & Moussas 2000). Alfvén waves are dominant in coronal regions that have open magnetic field lines and thus they leak into the fast solar wind (Cranmer & van Ballegooijen 2005). Within solar wind at 1 au, their average velocity is about  $57 \text{ km s}^{-1}$  (Veselovsky et al. 2010) – for more details about the Alfvén velocity see appendix Section A.4.

The coronal plasma in CHs expands superradially, following the magnetic field lines. However, the field strength decreases with increasing solar distance and at a distance of about  $2.5 R_\odot$  the thermal pressure becomes larger than the magnetic pressure. Thereby the magnetic field gets frozen within the plasma and is carried by the solar wind radially outwards into the heliosphere. The distance from which the solar wind propagation gets released from the magnetic field lines is called the source surface (Schatten et al. 1969) and the thermal to magnetic pressure ratio is called plasma beta – for more details on plasma beta see appendix Section A.3. The magnetic field changes from superradial expansion below the source surface to a radial configuration above it, this field geometry is also visible in the total eclipse image in Figure 2.3.

Open field lines expand over adjacent closed field regions. Above the cusps of these regions' closed loops, the surrounding plasma flows encounter each other and stream outwards, forming so-called helmet streamers. Above the helmet streamers, magnetic boundaries are created by plasma flows, carrying opposite magnetic polarity. These boundaries constitute an extensive coronal neutral line around the Sun. Within the heliosphere, the two dominating magnetic polarity regions, originating from both solar magnetic poles, are separated by the extension of this neutral line: a large plasma boundary surface, called the heliospheric current sheet (HCS) (Smith 2001).

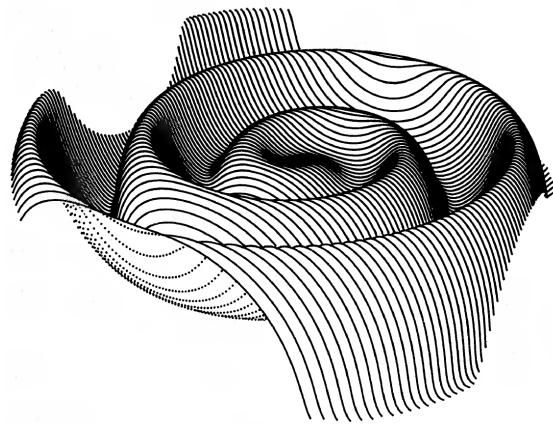
Under solar cycle minimum conditions, the magnetic dipole axis is near the rotation axis and thus the HCS is roughly located near the equatorial plane, dividing both hemispheres. The analytical solar magnetic field model for solar minimum conditions, constructed by Banaszkiewicz et al. (1998), shows this field geometry as seen in Figure 2.11. The quadrupole part of their dipole plus quadrupole plus current sheet (DQCS) model considers the closed equatorial fields and allows equatorial outflow along the current sheet. Commonly around solar minimum, the HCS's warped surface looks like a wavy ballerina skirt, due to the varying tilt angle between the dipole axis and the rotation axis, see Figure 2.12, and also due to local magnetic field variations (Jokipii & Thomas 1981).

During the field transition at solar maximum, the dipole axis rotates to lower latitudes, crosses the solar equator, and eventually the field ends up in a reversed dipole configuration (Jones et al. 2003). During this process, the HCS rotates almost rigidly together with the dipole axis and remains a single connected structure

## 2. Background knowledge



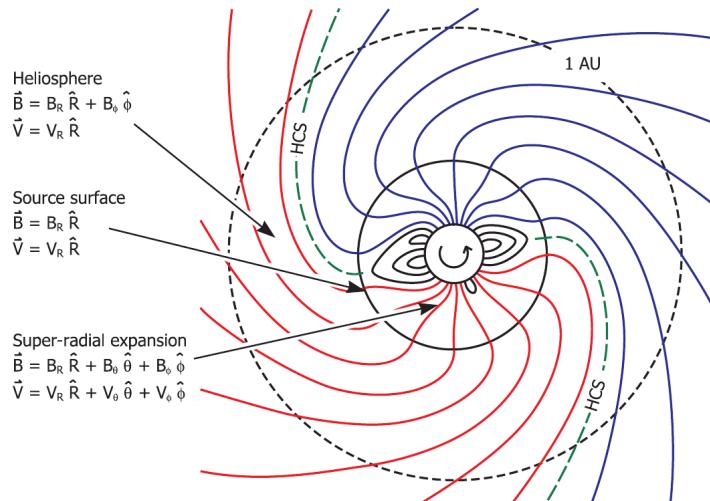
**Figure 2.11** Model of the solar magnetic field geometry in the polar plane for solar cycle minimum. Magnetic field lines (solid) and constant field strength surfaces (dashed) from the DQCS model are plotted. The field line spacing does not represent the field strength but provides better detail where needed. Credit: Banaszkiewicz et al. (1998, Fig. 3), © ESO, reproduced with permission.



**Figure 2.12** A simple HCS model, where its wavy surface shape solely stems from a solar dipole tilt of 15° to the rotation axis. The figure's extent is 25 au across. Credit: (Jokipii & Thomas 1981, Fig. 2), © AAS, reproduced with permission.

in the inner heliosphere (Jones et al. 2003). Hence during cycle maximum, the HCS has a very complex shape, is largely inclined to the solar equator, and reaches near-polar latitudes.

The solar wind source surface rotates with the Sun and thus shears the HMF into an Archimedean spiral pattern, adding an azimuthal component to the radial HMF. This geometry was anticipated by Parker (1958) and is today called Parker spiral. The Parker spiral, viewed in the ecliptic plane, is illustrated in Figure 2.13. The solar rotation axis tilt of up to 7.25° to the ecliptic leads to a slight diving into both hemispheres of opposite polarity. Thus, together with the ballerina topology of the HCS, the Parker spiral has typically a structure of either two or four sectors of alternating magnetic polarity (Ness & Wilcox 1965), which are separated by the HCS.



**Figure 2.13** Illustration of the Parker spiral formation in the ecliptic plane outside the source surface. The HCS (green) is located between solar wind flows of opposite magnetic field polarity (red/blue). Credit: Owens & Forsyth (2013, Fig. 1), adapted from Schatten et al. (1969, Fig. 1), licensed under CC BY-NC 3.0 DE.

The just described magnetic structures in the solar wind are overlaid by magnetic clouds (MCs) carried within CMEs. The MC's frequency and magnetic configuration vary with the solar activity cycle. Further, speed differences between solar wind streams, and between solar wind and CMEs cause enhanced field amplitudes and can result in shocks in the HMF.

That way, the HMF and its structures are carried out to the termination shock by the solar wind. MHD simulations, based on in-situ measurements of Voyager 1 and 2 within the heliosheath and based on IBEX observations of energetic neutral atoms, provide indications about the outer structure of the heliosheath. Behind the termination shock, the magnetic sector boundaries are compressed and they reconnect, forming magnetic

bubbles (Opher et al. 2011). These bubbles – unconnected to the HMF – flow away to the heliosheath tail region. Even beyond the termination shock, the solar wind plasma seems confined and collimated by the twisted solar magnetic field and driven into a northern and a southern jet (Opher et al. 2015). Hence, the Sun’s magnetosphere has likely a croissant-like shape with two turbulent tail-lobes, where eventually the solar wind and the HMF are being mixed into the interstellar medium.

## 2.5 Solar wind

It is observed that cometary tails lag a few degrees from the radial direction with respect to the Sun, sometimes they also show fluctuations and become kinked. As such behavior could not be explained by interaction with sunlight pressure, eventually Biermann (1951) concluded that cometary ion tails are influenced by a continuous flow of particles from the Sun. Parker (1958) considered the consequences of Biermann’s conclusions and built a solar wind model, adopting an expanding isothermal solar atmosphere. Parker also incorporated the implications for the solar magnetic field in his model and hence he laid the theoretical foundations for a continuous supersonic radial outflow of magnetized plasma. Thus, the existence of the solar wind was postulated before the first satellites measured it in situ in 1959 (Gringauz et al. 1960; Neugebauer & Snyder 1966). Since that time, spacecraft are able to measure the solar wind almost continuously with magnetometer and plasma instruments in situ (see Chapter 3). Pronounced solar wind structures, such as CMEs and streamers, become visible with the use of space-based coronagraph imagers. From Earth, the outflow of near-Sun solar wind can be observed during solar eclipses, see the eclipse photo in Figure 2.3.

The solar wind is a magnetized plasma consisting of electrons and ions. The ions are mainly composed of hydrogen, a small percentage of helium, and traces of oxygen, carbon, and other metals. The average abundance of helium is about 4.5 % and in slow wind at solar cycle minimum conditions less than 2 % (Feldman et al. 1978; Schwenn 1983; Kasper et al. 2012). The solar wind is commonly approximated by an ideal incompressible MHD plasma (viscosity  $\mu = 0$  and electrical conductivity  $\sigma = \infty$ ) and can be viewed as a neutral plasma. Also, its helium share is often viewed as being constant, in this case the proton density determines both the helium and electron densities. This work adopts these assumptions and treats the solar wind throughout as a proton plasma.

The properties of solar wind are highly variable in time and space. The key properties are determined by the values of the solar wind parameters magnetic field strength, proton velocity, density, and temperature. Their average magnitudes scale with solar activity, heliographic latitude, and solar distance. At a solar distance of Earth however, most of the time these parameter’s typical values lie in the ranges 3–8 nT, 300–500 km s<sup>-1</sup>, 2–8 cm<sup>-3</sup>, and 10<sup>4</sup>–10<sup>5</sup> K (Kivelson & Russell 1995, p. 92; Venzmer & Bothmer 2018). The low density of solar wind can be illustrated with a short comparison: 1 liter of air at standard pressure, expanded to a typical solar wind density of 6.5 cm<sup>-3</sup>, would occupy a volume of a cube with edge length of about 155 km. Solar wind quantities, such as particle flux densities, mass flux, pressures, and plasma beta, can be derived from the four listed parameters. Having the parameters in the aforementioned ranges, the solar wind is a plasma with a beta mostly greater than unity, that is, the average solar wind carries the HMF and its motions are not influenced by the field direction (for more on plasma beta see appendix Section A.3).

However, solar wind is structured by its different sources in the solar corona. It consists of fast continuous streams, slow variable flows, and transient CME events. These different flows have highly variable velocities, which result in compressed or rarefied regions at their interfaces. Additionally, the source region’s magnetic field configuration organizes the HMF, transported within the solar wind plasma. Regardless, pronounced magnetic structures embedded in the solar wind, such as field polarity changes or magnetic clouds, still influence the properties of the plasma.

This multitude of structures is apparent in the two months – beginning in May 2013 – of in-situ measured solar wind, which I plotted as an example period in Figure 2.14. The IMF and solar wind ion parameters were measured with the MAG and SWEPAM instruments onboard the Advanced Composition Explorer (ACE) spacecraft, located around the first Lagrange point (L1). The data have a time resolution of 64 seconds and are obtained from the ACE Science Center web interface<sup>6</sup>.

Some general solar wind tendencies can be seen from this plot: The temperature of the solar wind scales with its stream velocity; compressed plasma regions enhance the magnetic field and the density; HCSs, magnetic sector boundaries, and MCs come with high densities and low temperatures; MCs in CMEs have high magnetic fields and low temperatures. I indicated the periods of occurring solar wind structures, that is, HSSs, CIRs, HCSs, and CMEs, with colored bars – these types are further described in the following sections.

<sup>6</sup> ACE Science Center website: <http://www.srl.caltech.edu/ACE/ASC/>

## 2. Background knowledge



**Figure 2.14** Solar wind with several structures, measured at L1 during the time period 1 May to 5 July in 2013. The parameters are the magnetic field strength, its field angle in the ecliptic in GSE coordinates, the proton velocity, density, and temperature. I indicated periods of prominent solar wind structures with color bars: HSSs in blue, CIRs in green, HCSs in purple, and CMEs in red. In the velocity panel also the field polarity is color coded – assuming a Parker spiral angle of 135° at L1. Blank periods indicate bad or missing data. The data are 64 s measurements from the ACE spacecraft.

There exist a lot more aspects related to solar wind, i.e., MHD waves, the electron pitch angle, solar energetic particles, dust particles, and neutral energetic atoms. Yet, they are not covered here, because their relevance to the analyses in this work is considered minor.

### 2.5.1 Slow and fast streams

It is observed at 1 au that the continuous solar wind comes in streams roughly focused at two major velocity ranges (Neugebauer & Snyder 1966; Schwenn 1983), slow and fast streams with  $250\text{--}450 \text{ km s}^{-1}$  and  $450\text{--}800 \text{ km s}^{-1}$  respectively. Both types possess differences in their typical characteristics and ion compositions. Apart from its higher speeds, fast solar wind has most prominently lower proton densities ( $\sim 3 \text{ cm}^{-3}$ ) and higher temperatures ( $\sim 2 \times 10^5 \text{ K}$ ) than the slow solar wind, which has higher densities ( $\sim 10 \text{ cm}^{-3}$ ) and lower temperatures ( $\sim 4 \times 10^4 \text{ K}$ ) (Schwenn 1990). The fast solar wind has a nature of coming in steady high-speed streams (HSSs) with a unique magnetic field polarity, whereas slow solar wind is much more variable in all its properties except its velocity (Bame et al. 1977). HSSs are further overlaid with Alfvén waves, which modulate the stream velocity with typical periods of 15–60 min.

First soft X-ray observations of the corona, made during sounding rocket flights in the early 1970s, showed clearly that the fast solar wind emerges from extended areas of reduced X-ray emission, subsequently called coronal holes (CHs) (Krieger et al. 1973; Hundhausen 1977). The magnetic field polarities found in CHs are associated with the magnetic field directions observed in HSSs, as seen in Figure 2.15. In coronal regions with closed magnetic field lines, the plasma is trapped, though in CHs it can escape, following the open magnetic field lines outwards into space. Wave-particle interactions heat and accelerate the ions in CHs, leading to the emission of a fast solar wind (Hollweg & Isenberg 2002). Superradial expansion of the magnetic field lines in the corona has an influence on the wind speed – actually the expansion factor is anticorrelated with the final wind velocity (Wang & Sheeley 1990). As the field expansion is larger near the border of CHs, faster wind emerges from the mid regions of CHs, forming into HSSs. However, there are indications that the slow and fast solar wind are not only generated at different sources but from distinct mechanisms (McGregor et al. 2011a).

The high variability in the slow solar wind points to the existence of different types of slow wind flows,



**Figure 2.15** Solar wind velocity with respect to its estimated source longitude (top) and coronal brightness contour map at  $0.5 R_{\odot}$  above the photosphere (bottom) for the Carrington rotation 1616. The velocity is based on IMP spacecraft data, back-extrapolated to  $20 R_{\odot}$ . Brightness values below a fixed threshold are shaded corresponding to the magnetic field polarity (+/-) of the underlying photosphere. The map is based on observations from the K-coronameter at the Mauna Loa Observatory. Credit: Hundhausen (1977, Fig. 10), © Colorado Associated University Press, reproduced with permission.

originating from separate coronal locations and mechanisms (Schwenn 1983). It is still under debate if the variability is produced by the formation mechanism of the slow solar wind or if the variability is caused during the acceleration/propagation phase (Sanchez-Diaz et al. 2016). Still, at least a part of its variability can be attributed to the interactions between slow and fast solar wind, which result in a general reduction in velocity differences and thus let solar winds of different speeds (having different properties as well) converge to a common intermediate speed regime in the range  $400\text{--}500 \text{ km s}^{-1}$  (McGregor et al. 2011b; Sanchez-Diaz et al. 2016). Studies using remote white-light tracing of coronal material and in-situ measurements of solar wind suggest that multiple sources of slow solar wind flows exist (Wang et al. 2000; Kilpua et al. 2016). To the best of my knowledge, the generally considered sources are listed in the following:

- CH boundaries and small CHs, because their plasma outflow is slower due to the high superradial expansion of its open field lines (Wang & Sheeley 1990).
- CH boundaries, when trapped plasma is released by reconnection between open and closed field lines (Madjarska et al. 2004).
- Helmet/pseudo-streamers in active regions, where transient plasma blobs are released from the cusps of closed field loops (Wang et al. 1998, 2000). This slow and dense material is associated with the heliospheric plasma sheet belt.
- Edges of active regions, which have hot plasma outflows with a single magnetic polarity (Kojima et al. 1999).
- Jets originating from coronal bright points might contribute to the slow solar wind (Subramanian et al. 2010).
- Slow unidentified CMEs can be a source of slow wind as well.

It is found to be difficult to use in-situ measurements for tracing the slow solar wind flow types to different origins and distinguishing between them, because most properties are also highly variable in time (Kilpua et al. 2016). However, some indicators show tendencies to differentiate between the slow winds from different source regions. Some notable indicators are: elemental ion ratios, heavy ion charge states, and the specific entropy.

The elemental composition of the coronal plasma varies with height/location in the solar atmosphere, therefore the solar wind's elemental ion ratios (e.g., He/H, Fe/O) are used to determine its origin. The charge states of coronal heavy ions depend on the local temperature. However, the density of the outwards expanding plasma decreases fast, preventing further ionization/recombination. The charge states decouple from the local temperature and freeze in still close to the Sun. Thus, heavy ion charge ratios (e.g.,  $\text{C}^{+6}/\text{C}^{+4}$ ,  $\text{O}^{+7}/\text{O}^{+6}$ ) in the solar wind track the coronal source temperature and especially the  $\text{C}^{+6}/\text{C}^{+4}$  ratio is sensitive to the solar wind type (Landi et al. 2012). At solar minimum the specific proton entropy is found to correlate with the  $\text{O}^{+7}/\text{O}^{+6}$  ratio and thus able to trace slow solar wind sources as well (Pagel et al. 2004).

The solar wind stream pattern varies strongly with solar activity. The Sun's ordered dipole structure during solar cycle minima leads to polar regions with open magnetic fields, constituting large coronal holes, and to a

## 2. Background knowledge

large equatorial belt region with closed magnetic fields – this state is clearly visible in Figures 2.3 and 2.11. This state results in fast solar wind coming exclusively from the poles and higher latitudes, whereas active regions form an equatorial streamer belt around the Sun, emitting slow solar wind. This structure was confirmed from solar wind speed measurements done by the Ulysses spacecraft, which flew in an out-of-ecliptic solar orbit and whose mission covered a duration of more than one solar cycle (McComas et al. 2008), see Figure 2.16. The transition of the solar magnetic field during the solar cycle maxima induces the chaotic appearance of



**Figure 2.16** Solar wind velocity and magnetic field polarity (red/blue) with respect to heliographic latitude for the three orbits of the Ulysses spacecraft during low and high solar activity (upper panels). The data starts top left and runs counter-clockwise. The corresponding smoothed SSN (black) and HCS tilt angle (red) are plotted beneath. The background consists of solar images for solar cycle 22 minimum (1996-08-17), solar cycle 23 maximum (2000-07-12), and solar cycle 23 minimum (2006-03-28). The solar disk, inner corona, and outer corona images are from SOHO/EIT (Fe XII at 1950 nm), Mauna Loa K coronameter (700–950 nm), and SOHO/C2 white light coronagraph. Credit: McComas et al. (2008, Fig. 1), © American Geophysical Union, reproduced with permission.

closed magnetic fields at higher latitudes and even at the poles. Further, coronal holes begin to invade parts of the equatorial region, this leads to recurring phases of HSSs in the ecliptic. This can be seen from the solar wind period in Figure 2.14, where recurrent HSSs of the same field polarity but changing peak velocity exist – beginning on 6 May, 2 June, and 29 June 2013. The succeeding streams of different velocity result in interaction regions and alternating magnetic polarities result in magnetic sector boundaries.

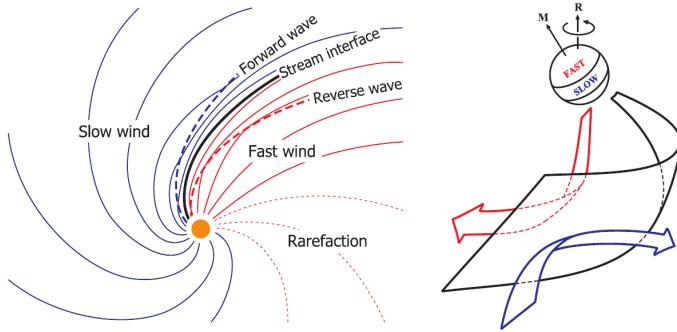
### 2.5.2 Stream interaction regions

The sources of the slow and fast solar wind rotate together with the Sun. Their uneven distribution on the solar surface – due to a significant inclination of the dipole axis or variations of the solar magnetic field – lead to the alternation of slow and fast wind streams that flow into the heliosphere (Owens & Forsyth 2013). In case of a fast stream being followed by a slow stream, a rarefaction region expands at their interface. If it is the other way around, the stream of fast solar wind catches up to that of slow wind ahead of it and a compression region forms at their interface, which is encompassed by two compressional waves (Balogh & Jokipii 2009). The latter are stream interaction regions (SIRs) and they take the shape of spiral fronts, as seen in the left panel of Figure 2.17.

When the solar dipole field is in a quasi-stable configuration, SIRs can stay for multiple solar rotations, recurrently sweeping over the heliosphere in 27-day periods (Gosling et al. 1972). Hence, they are referred to as co-rotating interaction regions (CIRs) (Smith & Wolfe 1976; Balogh et al. 1999). In the ecliptic at 1 au, CIRs occur commonly during the declining phase of the solar cycle when polar CHs form equatorial extensions (Balogh & Jokipii 2009).

The spiral shape of the stream interfaces and their inclination to the solar rotation axis lead to a deflection of both streams (Balogh & Jokipii 2009). Due to the fast and slow streams' collision, the fast wind is decelerated and the slow wind is accelerated and their flow directions are systematically deflected away from the interface, as shown in the right panel of Figure 2.17.

The plasma pressure inside SIRs is increasing with heliocentric distance and therefore, the leading and



**Figure 2.17** Schemata of the formation of a stream interface (left) and the deflection of streams along the interface (right). The stream interface (black) is located between regions of slow (blue) and fast (red) solar wind. Credit: [Owens & Forsyth \(2013, Fig. 7\)](#), right panel adapted from [Pizzo \(1991, Fig. 2\)](#), licensed under [CC BY-NC 3.0 DE](#).

trailing compressional waves form into forward and reverse shocks – typically at solar distances between 2–10 au ([Smith & Wolfe 1976](#); [Balogh & Jokipii 2009](#)). The solar wind speed increases abruptly at both shock fronts. With increasing solar distance, the leading and trailing shock fronts travel away from the stream interface. This widens the interaction regions and eventually they gain on the close-by interaction regions ([Burlaga et al. 1984](#)). Beyond 10 au they fuse to merged interaction regions, which are narrower and more compressed ([Burlaga et al. 1985](#)).

Due to the compression within SIRs, the magnetic field strength and the plasma density are higher than in the ambient streams. These signatures can be seen in the in-situ solar wind plot in [Figure 2.14](#). The CIRs on 5 May, 1 June, and 29 June 2013 are followed by recurrent HSSs and they contain sector boundaries as well. They are not yet accompanied by strong shocks owing to the measurement location at 1 au.

### 2.5.3 Heliospheric current sheet

The heliospheric current sheet (HCS), the boundary surface between solar wind streams of opposite magnetic polarity, is formed by open magnetic fluxes, expanding from both sides over closed field regions and coming into contact above. At the boundary between open and closed fields, local coronal plasma gets released and flows slowly between the fast wind streams along the current sheet into the heliosphere, creating a helmet streamer. This is why near the Sun, the HCS is typically located within slow solar wind stemming from the coronal streamer belt ([Owens & Forsyth 2013](#)). However, with increasing solar distance, the shock wave of an adjacent SIR can pass the HCS, so that sector boundaries are often found to be embedded within SIRs ([Gosling & Pizzo 1999](#)).

When sector boundaries in the solar wind are observed in situ, they come with significant depletions in  $\text{He}^{++}/\text{H}^+$  values ([Borrini et al. 1981](#)). The other solar wind parameters, such as velocity, density, and temperature, change with distance from the HCS as well ([Smith 2001](#)). The HCS region itself is quite narrow with a thickness of around 3000–10 000 km; it is embedded in a region of 20–30 times its thickness, the heliospheric plasma sheet (HPS) ([Winterhalter et al. 1994](#)). The HPS contains the magnetic polarity reversal and is of low magnetic field strength and high density, resulting in a significantly enhanced value of plasma beta ([Crooker et al. 2004](#)). In the slow solar wind, the HPS can also be identified from its particularly low specific entropy, arising from its low temperature and high density ([Kilpua et al. 2016](#)).

### 2.5.4 Coronal mass ejections

Coronal mass ejections (CMEs) are eruptions of coronal magnetized plasma, which expand within hours to bubbles with sizes of a few solar radii. They continuously expand further while moving farther away from the Sun into the heliosphere with velocities often surpassing even the high speed solar wind. When measured in-situ in interplanetary space, they are often called interplanetary CMEs (ICMEs). They are transient structures found to be embedded in the ambient slow and fast solar wind streams, having durations of a few days. CMEs make up about 5 % of the solar wind's flow share during solar cycle minima, but can represent up to about 50 % during solar cycle maxima ([Richardson & Cane 2012](#)). Indeed, their frequency correlates with the sunspot number ([Hildner et al. 1976](#)) and follows solar activity in amplitude and phase ([Webb 1991](#)).

Long before the origins of geomagnetic disturbances were actually attributed to CMEs, a solar influence was identified as their source by [Carrington \(1859\)](#). Indeed, CMEs are the major drivers for strong geomagnetic storms, because they carry the most extreme conditions found in the solar wind. Thus, they are of major importance to space weather – their impacts on the terrestrial magnetosphere are covered in [Section 2.6](#).

I will cover the CME formation processes only briefly, because this work is focused on in-situ measurements. Generally speaking, CMEs are products of the instability in the coronal magnetic field. The solar differential

## 2. Background knowledge

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rotation wraps the coronal magnetic field, which is rooted at the surface. Accumulated tension and emerging magnetic flux eventually initiate sudden field reconfigurations. These field line reconnections result in a lower state of potential energy, releasing a lot of energy in form of solar X-ray flares and CMEs. Indeed, CMEs are often associated with flares and eruptive prominences, coming from their source regions (Webb & Hundhausen 1987). The sources of CMEs are located near active regions and are frequently identified with eruptive prominences (Subramanian & Dere 2001).

CMEs emerge from coronal magnetic flux rope structures that hold plasma filaments embedded along their base (Webb & Hundhausen 1987; Cremades & Bothmer 2004). Reconfiguration of the ambient magnetic field can lead to a release and a subsequent ascension of the flux rope through the corona. During this process, the filament is lifted by the flux rope and forms a prominence eruption. The accompanying sudden magnetic reconnections often cause a multitude of other dynamic coronal phenomena: enhanced X-ray flaring; the formation of post-eruptive arcades where the filament disappeared; local decreases in soft X-ray intensity (coronal dimmings) due to depletion of the plasma density caused by magnetic field expansion; and large-scale coronal disturbances observed in extreme ultraviolet (EUV-waves) that can spread across the whole corona.

The onset of CMEs is often accompanied by solar energetic particles (SEPs) and by solar radio bursts. SEPs consist of protons, electrons, and ions with relativistic energies that can cover the distance to Earth within half an hour and pose a radiative threat to humans in space. The intensity of energetic protons correlates well with the CME speed (Kahler et al. 1978) and it is established that these coronal particles are accelerated by shocks formed in front of the rapidly expanding magnetic structures of CMEs in the corona (Cliver et al. 1982; Gosling 1993). Solar radio bursts are radio emissions from the corona that change with time in frequency space. There are different kinds of radio bursts, but especially type II radio bursts are seen as indicators for coronal shocks that accelerate electrons. They are emitted in the low corona and are drifting to lower frequencies with time. There is still a debate whether type II bursts originate from flares or CMEs, favoring the latter (Gosling 1993; Cliver et al. 2005; Cho et al. 2011).

CMEs were detected in the white-light observations of the K-corona made by the first space-based coronagraphs on board the OSO 7 satellite (Tousey 1973) and the Skylab space station (MacQueen et al. 1974). These observations show the steady outflow of solar wind, broken by intermittent ejections of coronal plasma. Subsequently, the kinematic properties of CMEs were identified from the white-light images (MacQueen 1980). Now, coronagraphs observe the corona and hence CMEs continuously from the first Lagrange point (L1) in front of Earth with the SOHO spacecraft and from changing equatorial perspectives with the STEREO Ahead (A) and Behind (B) spacecraft. The image of the corona from 23 September 2012, made by the SECCHI/COR2 coronagraph onboard STEREO A, shows a CME to the top right, see Figure 2.18.



**Figure 2.18** Image of the solar corona out to  $15 R_{\odot}$  from 23 September 2012 taken by the SECCHI/COR2 coronagraph onboard the STEREO A spacecraft. The solar disk is covered by the occulter disk and its position is indicated by the white circle. The bright blob to the top right is the CME; the smooth elongated radial lines are solar wind streamers. Courtesy of [STEREO/COR2 consortium \(NASA\)](#).



**Figure 2.19** Schema of a magnetic flux rope structure in a CME. The arrows depict passing through its flank (F) and its apex (A). Credit: Marubashi & Lepping (2007, Fig. 1, panel (a)), licensed under CC BY 3.0.

It was early determined by Gold (1962) that solar ejecta should drive shock waves ahead. In fact, shocks with trailing low proton temperatures caused by fast CMEs were then found in in-situ measurements (Gosling et al. 1973, 1974). Burlaga et al. (1981) analyzed magnetic field and plasma in-situ data from five spacecraft and identified a shock wave with a trailing turbulent sheath region followed by an organized helical magnetic structure that they called a magnetic cloud (MC). MCs have an enhanced magnetic field, a smooth rotation in the azimuthal magnetic field component and they show low densities and temperatures (Burlaga et al. 1981). Thus, MCs have a low thermal to magnetic pressure ratio (i.e., a small plasma beta) and the magnetic field dominates the plasma. Furthermore, the overall pressure in MCs is higher than in the ambient solar wind, resulting in the expansion of MCs on their way out. Shock-driving CMEs containing a helical MC are actually identified with magnetic flux ropes that expand self-similarly and that remain in connection with the solar surface (Chen et al. 1997). The surface connection is indicated by bi-directional streams of electrons that are found in MCs (Gosling et al. 1986). The shape and magnetic topology of such a magnetic flux rope is pictured in Figure 2.19. It is apparent that in-situ measurements should look significantly different depending on where the CME is pierced.

Solar wind in-situ measurements reveal the magnetic structure of CMEs. In particular, the orientation of magnetic flux ropes can be determined via applying a minimum variance analysis (MVA) to the MCs' magnetic field components (Sonnerup & Cahill 1967; Burlaga & Behannon 1982). An MVA determines the direction of minimum variance in the sequence of field vectors passing by during an MC encounter. This direction is interpreted as the principal axis of the flux rope.

Bothmer & Schwenn (1998) used the MVA on an extensive set of MCs which they found in the data of the Helios 1 and 2 probes. They related the results with the magnetic polarity structures of the MCs' apparent solar source regions. Connecting the derived flux rope directions with the orientation of disappearing filaments and magnetic neutral lines on the solar surface, they recognized a scheme that is able to infer the orientation and helicity of MCs found in CMEs. This Bothmer-Schwenn scheme (BSS) relates these magnetic flux rope properties to whether the solar cycle number is even or odd and depending on which solar hemisphere the CME originates from (northern/southern) – utilizing that the hemispheric polarity is alternating with each solar cycle, see Figure 2.20. As the probability is greater than 80 % that the magnetic topology of an active region



**Figure 2.20** Explaining sketch of the BSS. The two rows represent odd ( $n$ ) and even ( $n + 1$ ) solar cycle numbers. The magnetic polarity of sunspots, the structure of filaments, their helicity, and the corresponding flux rope type of magnetic clouds are shown. Credit: Bothmer & Schwenn (1998, Fig. 18), © EGS – Springer-Verlag, reproduced with permission.

conforms to the hemispheric rule (Wang 2013), an MC configuration predicted with the BSS is expected to have a reliability that is of the same order (Savani et al. 2015).

Three CMEs can be seen in the solar wind in-situ measurements showed previously in Figure 2.14, passing by the ACE spacecraft at L1 on 24 May, 6 June, and 27 June in 2013. The latter CME has a well-structured MC, which is presented in detail in Figure 2.21 – I use this event as an example throughout this section. In addition to showing the solar wind in-situ parameters, I indicated the shock, compression zone and the MC with dotted

## 2. Background knowledge



**Figure 2.21** Solar wind with a CME, measured at L1 during the time period 26 June to 2 July in 2013. The solar wind parameters are the magnetic field strength, its z-component and ecliptic field angle in GSE coordinates, the proton velocity, density, and temperature; in addition, the geomagnetic  $K_p$  index is plotted in the bottom panel. In the velocity panel also the field polarity is color coded – assuming a fixed Parker spiral angle of  $135^\circ$ . I indicated the shock, the compression zone, and the duration of the magnetic cloud with dotted lines. Blank periods indicate bad or missing data. The solar wind data was measured with the MAG and SWEPAM instruments onboard the ACE spacecraft and is obtained from the ACE Science Center. The  $K_p$  data is obtained from the GFZ Potsdam.

lines, and plotted the geomagnetic  $K_p$  index in order to visualize the CME's impact on the magnetosphere. This MC contains a sector boundary and is trailed by an interaction region caused by a following HSS.

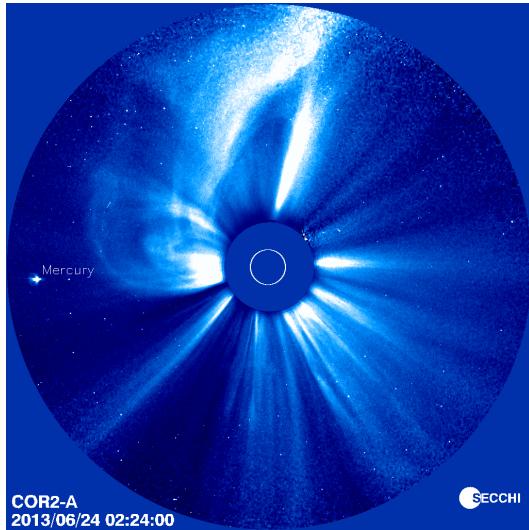
The CME occurred in solar cycle 24. The MC's IMF z-component changes from positive (northern) to negative (southern) values – during this process, the field angle in the ecliptic,  $\phi$ , stays pointing roughly towards  $270^\circ$  (west). Thus, it has a NWS configuration, which is expected from the BSS during even numbered solar cycles for CMEs with source regions located in the northern hemisphere.

The combination of white-light images with in-situ data enables relating the observed structures of CMEs. Even disturbances in front of fast CMEs can be identified as shock waves in the white-light images made by the SOHO coronagraph (Sheeley et al. 2000). The diffuse leading feature of a CME is the shock sheath, its brightness is caused by the density jump after the shock, which itself is not visible in coronagraphs. The trailing void is identical to the low density of the magnetic flux rope, which drives the whole structure.

The CME described before is indeed coming from the northern hemisphere, as can be seen from the STEREO A coronagraph image displayed in Figure 2.22 taken three days earlier. Apparently, the CME only grazes the ecliptic and its lower part passes Earth. Also in this image, the event originates from the backside of the solar disk, considering the observing spacecraft's position in relation to Earth at that time, see Figure 2.23.

Relating CME white-light images to observations of magnetic neutral lines on the solar surface, Cremades & Bothmer (2004) interpreted the white-light appearances of CMEs as projections of their 3D geometry and developed a scheme for their orientation depending on the orientation and location of the neutral lines on the solar surface. Neutral lines on the solar surface are the lines of polarity inversion between two areas of opposite magnetic polarity. The scheme is based on the topology of the flux rope model and the observation that its major axis stays roughly aligned to the underlying neutral line during the eruption process. Thus, Cremades & Bothmer (2004) concluded that CMEs look systematically different depending on their source region's position and magnetic configuration, as described in Figure 2.24. For the solar backside, the neutral lines are reversed

<sup>7</sup>STEREO Science Center website: <https://stereo-ssc.nascom.nasa.gov/>



**Figure 2.22** Image of the solar corona from 24 June 2013 made with the same coronagraph as the image in [Figure 2.18](#). The CME is the extended structure to the upper left. Courtesy of [STEREO/COR2 consortium \(NASA\)](#).



**Figure 2.23** Positions of the STEREO A and B space-craft for 24 June 2013 02:24 UT. This figure is made with the online [STEREO Orbit Tool](#) at the STEREO Science Center website<sup>7</sup>.

and according to the scheme so are the CME orientations – thus the broad backside CME seen in [Figure 2.18](#) matches this scheme. The scheme can effectively be applied for CMEs with source regions located less than  $50^\circ$  from the solar equator ([Webb & Howard 2012](#)).

This white-light CME projection scheme as well as the BSS for MCs are based on the notion that the orientation of flux ropes propagating outwards generally stays aligned with the neutral polarity inversion line on the solar surface where the filament erupted ([Marubashi 1997](#); [Bothmer & Schwenn 1998](#)). In fact, it is shown that flux ropes are not likely to rotate significantly after their initiation but maintain the orientation of their main axis parallel to the magnetic neutral line even in interplanetary space ([Marubashi et al. 2015](#)). However, during and shortly after their eruption, CMEs are frequently observed to be deflected to other directions by the surrounding coronal magnetic field structure ([Sterling et al. 2011](#)).

The revelation that all CMEs may be flux ropes ([Vourlidas et al. 2013](#); [Marubashi et al. 2015](#)) led to the expansion of the CME definition based on white-light images made by [Hundhausen et al. \(1984\)](#). [Vourlidas et al. \(2013\)](#); [Vourlidas \(2014\)](#) include magnetic flux ropes in their recent CME redefinition: “A CME is the eruption of a coherent magnetic, twist-carrying coronal structure with angular width of at least  $40^\circ$  and able to reach beyond  $10 R_\odot$  which occurs on a time scale of a few minutes to several hours.”

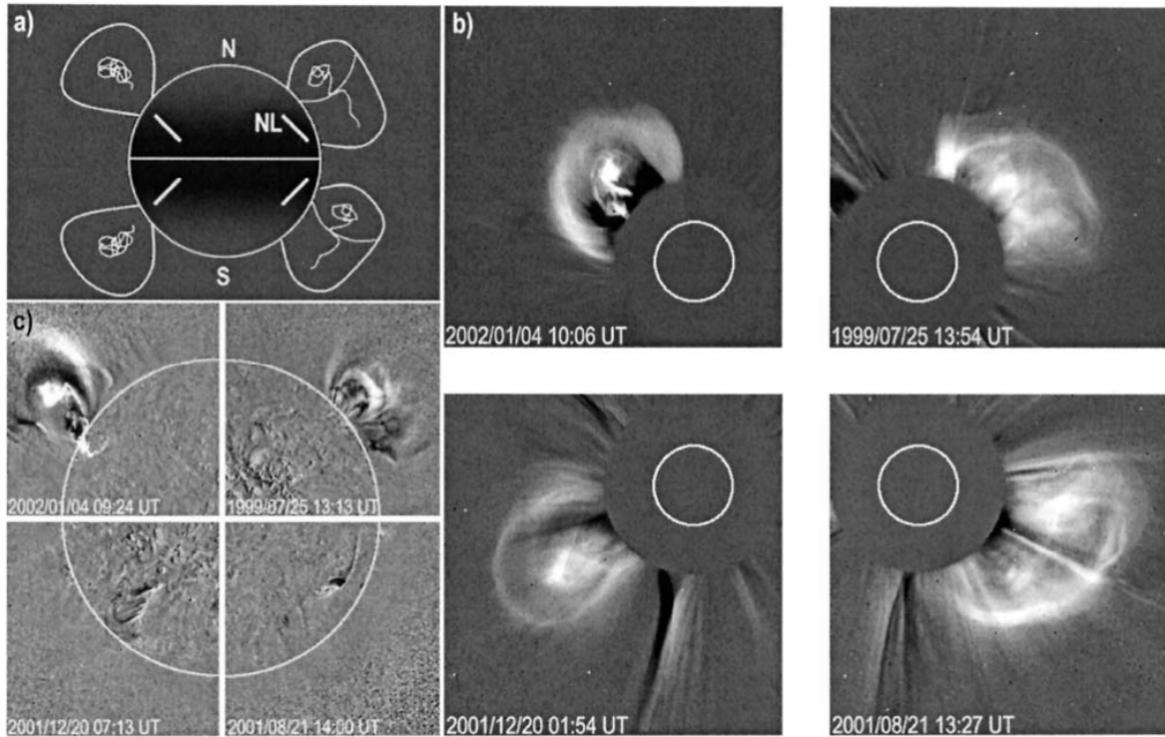
However, CMEs generally do not have a perfect flux rope geometry: Often their more complex structure can be seen in kinks of the corresponding filament before the eruption; and strong distortions can happen during the eruption process. Their cross section is only initially of circular shape, it distorts due to the pressure-driven self-expansion and the radial solar wind expansion ([Owens et al. 2006](#)). Though, the inner core of the flux rope is believed to stay of circular shape. For these reasons, the properties of CMEs vary widely and imaging projection effects contribute further to that ([Cremades & Bothmer 2004](#)).

There still exist a lot of unresolved questions about CMEs, their formation, and the effects observed along with them: Multiple mechanisms/processes are debated for the release of CMEs and their subsequent acceleration. It is also not known if there exist CMEs without magnetic flux ropes ([Vourlidas et al. 2013](#)). The relation between flux ropes observed in the near-Sun corona and measured in situ at interplanetary distances is still not entirely clear ([Vourlidas 2014](#)). Nobody knows yet when a CME will erupt on the Sun, at what exact time it will pass a given point in the heliosphere, and how its detailed magnetic configuration would look like in situ ([Gopalswamy 2016](#)).

## 2.6 Space weather

Space weather is the field of research that comprises all dynamic effects generated by the Sun in its domain of influence – the heliosphere. Naturally, it is especially focused on the solar effects observed on Earth and in its local environment – the ionosphere and the magnetosphere. There exist several different effects caused by

## 2. Background knowledge



**Figure 2.24** White-light CME projection scheme for frontside events, showing the case for each quadrant of the solar disk. (a) Schema showing the expected white-light appearance of CMEs and the orientation of the underlying magnetic neutral lines. (b) White-light images of CME example events, made with the LASCO C2 coronagraph on the SOHO spacecraft. (c) Details (prominences and post-eruptive arcades) of the source regions related to the CMEs shown in panel (b). Note that for events originating from the backside, panel (a) has to be mirrored vertically. Credit: Cremades & Bothmer (2004, Fig. 15), © ESO, reprinted with permission.

independent solar events occurring on different time scales that can affect humans directly or indirectly. Solar events affect sensitive technical systems and strong events are even able to disrupt them and pose a threat to humans (Bothmer & Daglis 2007). Space weather deals with effects on time scales similar to the terrestrial weather, that is, days and weeks.

The existence of a direct solar-terrestrial relation was known of early on. Carrington (1859) made the first connection between solar flares and disturbances in the terrestrial magnetic field when he observed a brightening on the solar disk on 1 September 1859 and suggested its connection to a strong geomagnetic storm that occurred about 17 hours later. Categories of solar-terrestrial correlations were already listed by Bartels (1962): Events initiated by irregular solar X-ray flares and CMEs, by the 11-year solar activity cycle, and by the periodic 27-day solar rotation. Other time related variations include seasonal effects caused by the Earth's orbital distance change over the year and the inclinations of the Sun's and Earth's rotation axes to the ecliptic normal. The terrestrial magnetic field's dipole axis tilt to the rotation axis contribute to the daily effects.

The major driving forces behind severe space weather events are CMEs, solar X-ray flares, and solar energetic particle (SEP) events. Most of the terrestrial space weather effects are introduced through their influence on the magnetosphere, ionosphere, and upper atmosphere. Space weather events are commonly classified into geomagnetic storms, solar radiation storms, and radio blackouts<sup>8</sup>.

The solar wind's internal structures, such as CMEs and CIRs, affect the magnetosphere and can cause disturbances in its magnetic field, that is, geomagnetic storms and substorms, both are described in more detail in the following sections. Notable consequences of these disturbances in the terrestrial magnetic field are listed in the following:

- Impact on the magnetic navigation and behavior of animals, such as migratory birds (Moore 1977) and whales (Vanselow et al. 2017).
- Stronger auroral activity and shifting of the auroral oval to lower latitudes<sup>9</sup>.
- The auroral oval shifts equatorward, which leads to direct risks for humans from enhanced radiation

<sup>8</sup>NOAA Space Weather Scales website: <https://www.swpc.noaa.gov/noaa-scales-explanation>

<sup>9</sup>NOAA/SWPC Tips on Viewing the Aurora: <https://www.swpc.noaa.gov/content/tips-viewing-aurora>

doses during high altitude flights through these regions<sup>10</sup>.

- The Van Allen radiation belt intensifies and can disturb satellite operations and lead to faulting electronics.
- Geomagnetic storms heat the upper atmosphere, causing it to expand, and generate an increased drag for satellites in low Earth orbits<sup>11</sup>.
- Geomagnetically induced currents, created in long conducting infrastructure, such as oil pipelines, rail tracks, and power lines, can lead to increased steel corrosion and power grid outages.
- The sunward magnetopause distance can shrink under extreme solar wind pressure below the geostationary Earth orbit (GEO) at  $6.6 R_E$ . GEO satellites crossing the magnetopause are directly exposed to SEPs.

Solar extreme ultraviolet (EUV) and X-ray radiation affect the ionosphere apart from their regular day/night influence. Solar flares can create sudden ionospheric disturbances (Gosling 1993), that is, abrupt enhancements in the ionospheric total electron content (TEC). These TEC variations can threaten the following technical systems (Kraaijkamp & Verbeeck 2015):

- Increased TEC leads to larger positioning errors in global navigation satellite systems (GNSS).
- Scintillation and absorption of high frequency signals on their way through the ionosphere lead to disturbance and loss of satellite communications.
- Induced changes in the stratification of the ionosphere (D-, E-, and F-layers) disturb radio communication and lead to outages in different frequencies.

Flares and shock fronts of CMEs accelerate coronal particles and produce SEPs. High fluxes of SEPs induce solar radiation storms with durations ranging from hours to days that have the following effects (NOAA/SWPC...):

- Radiation damage to electronic circuits in satellites.
- Radiation damage to DNA in humans located in space.
- The SEPs are guided by the terrestrial magnetic field down to the polar regions. Humans flying in aircraft near polar latitudes are exposed to higher radiation doses during strong radiation storms – especially humans in orbit outside the protective atmosphere.
- Free electrons, created by SEPs colliding with the atmosphere, form ionospheric layers that disturb/block high frequency radio communications.

Notable space weather effects on the rest of the solar system bodies:

- Inner solar wind structures can strip off cometary tails and enhance atmospheric losses in bodies without magnetosphere.
- Other planetary magnetospheres are impacted in similar ways as the terrestrial is.
- The average solar wind pressure and HMF scale with solar activity and therefore modulate the shape and size of the heliosphere. As cosmic rays are influenced by the HMF, the local cosmic ray radiation level scales inversely with solar activity.
- Even the passing of strong MCs lead to reductions in local cosmic ray radiation, which are called Forbush decreases.
- The solar system regions outside of protective magnetospheres suffer from episodes of enhanced solar radiation. This becomes important in the near future, when space weather directly impacts humans on space missions to Moon or Mars.

As can be seen, most space weather effects are connected with economic interests and some impacts on technical systems can even influence human lives. This creates a strong interest in forecasting the space weather. Space weather is a relatively new field of study, however, it becomes more and more important, not only because of the increasing use of sensitive technology at Earth, but also because of the anticipated expansion of interplanetary space activities within the near future. There exist lots of dedicated institutions and a multitude of forecasting techniques for the different aspects of space weather and more are being developed.

Predictions of magnetospheric disturbances are based on knowledge of the solar wind properties, which the magnetosphere will encounter. Thus in a first step, forecast methods have to predict the solar wind at Earth in advance and derive its expected impact on the magnetosphere in a second step. The first part of this present study contributes to the second step with predicting the amplitude of magnetospheric disturbances caused by solar wind. The basic structure of the magnetosphere, its coupling to the solar wind, and how disturbances are quantified/measured is described in the following Section 2.7. Existing forecast methods for geomagnetic

<sup>10</sup>NOAA/SWPC Galactic Cosmic Rays: <https://www.swpc.noaa.gov/phenomena/galactic-cosmic-rays>

<sup>11</sup>NOAA/SWPC Satellite Drag: <https://www.swpc.noaa.gov/impacts/satellite-drag>

disturbances and for their solar wind sources are described in [Section 2.8](#).

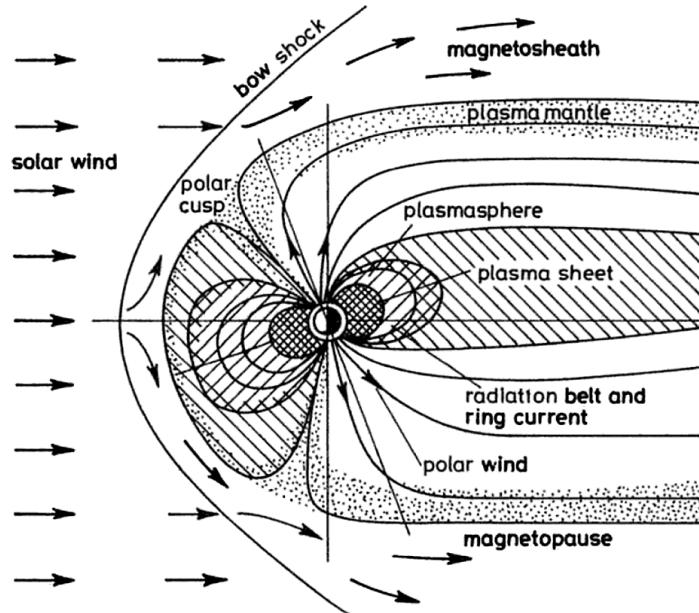
strong radio bursts...

## 2.7 Magnetosphere

In first-order approximation, the terrestrial magnetic field can be seen as a dipole aligned with the rotation axis of the Earth. Its magnetic south pole is located near the geographic North Pole and there is currently only a small tilt of about  $9.5^\circ$  between both axes that varies slightly over the years ([Thébault et al. 2015](#)). The magnetic field strength on the surface at the poles is around  $60 \mu\text{T}$ , whereas at the equator it is about half that magnitude. At the Earth's surface, this dipole approximation neglects major field irregularities, such as the prominent South Atlantic Anomaly whose field is lower than  $24 \mu\text{T}$  ([Thébault et al. 2015](#)).

The geomagnetic field's strength decreases while it extends out into space. Its influence ends where its magnetic pressure balances the ram pressure of the solar wind. This outer boundary is the magnetopause – in average it has a sunward distance of about  $8.2 R_E$  ([cite?](#)). In response to the continuous changes in the solar wind pressure, the magnetopause usually is in a state of fast radial inward-outward motion and its displacement can reach several Earth radii ([de Keyser et al. 2005](#)). The solar wind flow shapes the magnetopause into a surface of teardrop form, named magnetosphere by [Gold \(1959\)](#), whose so-called magnetotail length varies around  $100 R_E$  and points away from the Sun.

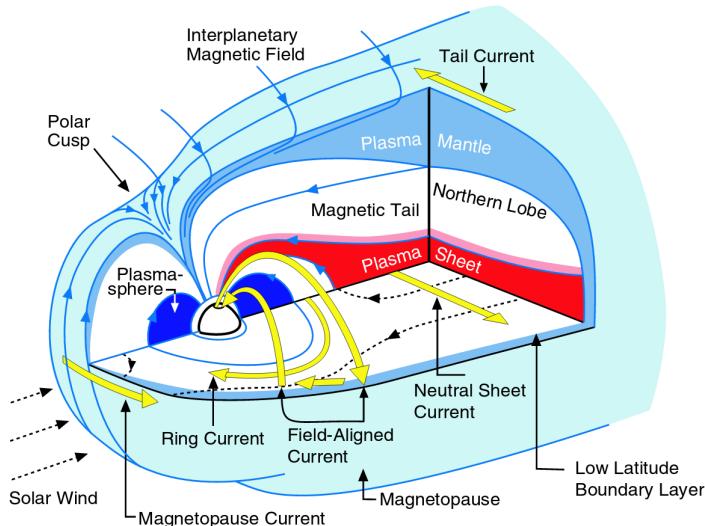
The solar wind compression at the front of the magnetopause creates a bow shock, whose standoff distance to the magnetopause is a few Earth radii but varies largely under extreme solar wind conditions. The shocked solar wind plasma flows within the magnetosheath around the magnetosphere as seen in [Figure 2.25](#). The IMF



**Figure 2.25** Schema of the magnetosphere's geometry in the plane spanned by the solar wind flow direction and the ecliptic normal (vertical line). The arrows show the flow of solar wind around the Earth's magnetic field. The diagonal line indicates the inclination of the rotation axis to the ecliptic. Credit: [Davies \(1990, Fig. 2.12\)](#), © IET, ... get permission...

in the magnetosheath is weaker but has a larger variability than the terrestrial magnetic field on the other side of the magnetopause ([de Keyser et al. 2005](#)).

At the front of the magnetopause, the incoming solar wind ions and electrons are being deflected in direction of dawn and dusk by the magnetospheric field, creating a current layer at the surface of the magnetopause. This current induces another magnetic field that cancels the geomagnetic field outside of the magnetopause and enhances it inside to about twice the strength of a pure dipole field at that distance ([de Keyser et al. 2005](#)). At the magnetotail side, the magnetopause current flows into the opposite direction – an overview, also of the inner magnetosphere's current systems is illustrated in [Figure 2.26](#).

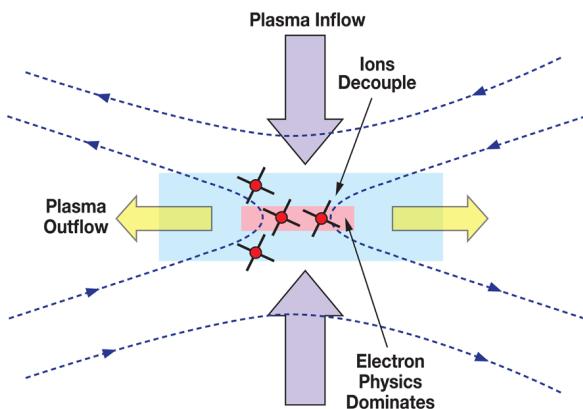


**Figure 2.26** Schema of the magnetosphere's inner 3D structure with focus on its current systems and plasma regions. The directions of the magnetic field lines and of the electric currents are indicated with blue and yellow arrows. Credit: de Keyser et al. (2005, Fig. 2.12), adapted from Kivelson & Russell (1995, Fig. 1.18), © Springer, reproduced with permission.

### 2.7.1 Solar wind coupling mechanisms

There exist several ways of how solar wind couples to the magnetosphere and deposits energy and plasma within. The contributions of the different mechanisms by which energy is transferred and solar wind plasma is able to penetrate the magnetopause are not yet established (Phan et al. 2005). The major interaction processes and underlying physical mechanisms are magnetic reconnection, compression, turbulence and are described in the following.

- Magnetic reconnection occurs where the IMF comes into antiparallel contact with the terrestrial magnetic field. At this region on the magnetopause, the magnetosphere opens up to the IMF and reconnection of the field lines occurs, resulting in a change of the magnetic topology (Phan et al. 2005). The oppositely directed magnetic field lines reconnect along a line with an X-geometry at the magnetopause. The reconnection process at this so-called X-line harbors a narrow diffusion region, where the plasma ions and electrons decouple to get accelerated in jets of particles by the reconnected field lines (Phan et al. 2005), see the [Figure 2.27](#). Thus, the magnetopause is left with a small magnetic field normal to it that has an



**Figure 2.27** Schema of an X-line reconnection region. The dashed arrowed lines represent magnetic field lines and their direction. The large arrows indicate the plasma flow direction and the shaded areas are the ion and electron diffusion regions. The red crosses represent the four Magnetospheric Multiscale (MMS) spacecraft that are built to analyze the magnetopause reconnection. Credit: NASA/GSFC MMS mission.

opposite polarity on each side of the X-line (de Keyser et al. 2005).

Although the microphysical processes leading to magnetic reconnection are yet little understood (Phan et al. 2005), there is evidence that magnetic reconnection is the dominant plasma transport mechanism into the magnetosphere (de Keyser et al. 2005).

The reconnection location on the magnetopause shifts depending on the direction of the IMF. During periods of southern IMF, the terrestrial field and IMF are antiparallel at the subsolar point of the magnetosphere, owing to the Earth's dipole orientation. The rate of reconnection becomes highest in this case, whereas, when the IMF is directed northward, reconnection tailward of the polar cusps has been observed (Phan et al. 2005). The magnetic flux opened at the front of the magnetosphere is in average balanced by reconnections in the magnetotail, closing the field again. This process is part of the Dungey convection cycle which is described in the next [Subsection 2.7.2](#).

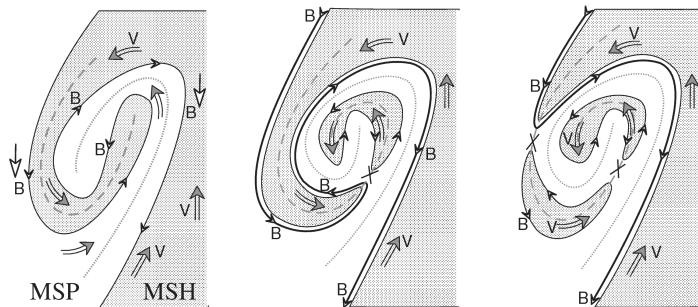
## 2. Background knowledge

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- High solar wind pressure during times when the IMF direction is parallel to the terrestrial magnetic field leads to compression of the sunward magnetosphere.
- Viscous interaction with the solar wind plasma is also able to insert energy into the magnetosphere ([Alfvén 1942](#)). Solar wind drag at the flanks of the magnetopause creates Kelvin-Helmholtz (KH) instabilities in form of turbulent eddies. It is observed that even during northern IMF, these turbulent vortices are able to channel solar wind plasma into the magnetosphere – either through forced magnetic reconnection or non-reconnection processes ([Otto & Nykyri 2003](#); [Phan et al. 2005](#)). MHD simulations of the velocity shear at the magnetopause during northern IMF suggest the presence of a double-vortex sheet structure ([Merkin et al. 2013](#)).

see [Figure 2.28](#)

**Figure 2.28** Schemata showing reconnection in a turbulent vortex, forming between the magnetosphere (MSP) and the magnetosheath (MSH, shaded area) during parallel magnetic fields. The boundary magnetic field line is indicated by the arrowed line, the plasma flow direction by the gray arrows, and the locations of reconnection by crosses. Credit: [Merkin et al. \(2013, Fig. 5\)](#), © American Geophysical Union, get permission...



- Induction of currents created by solar wind interaction.
- 

ring current systems

auroral ovals

“Reconnection erodes the magnetopause, causing the auroral oval to move lower in latitude. The By polarity shifts the auroral oval either toward dawn or toward dusk.”

radiation belts

### 2.7.2 Dungey convection cycle

half-wave rectifier

plasmoids in magnetotail

opening and closing flux rate

definition of geomagnetic storm versus substorm...

substorm cycle...

([Milan et al. 2007](#)): Magnetic flux transport in the Dungey cycle

Approximately half of the events can possibly be identified with a trigger, the other half appearing spontaneous.

We find no evidence for a low background rate of nightside reconnection between these events and conclude that substorms and other large reconnection bursts provide the major or only source of flux closure on the nightside.

In 1961, Dungey proposed that the dynamics of the magnetosphere are driven by the interaction between the magnetized solar wind and the terrestrial magnetic field, this interaction being realized by magnetic reconnection at the dayside magnetopause and in the magnetotail [[Dungey, 1961, 1963](#)]. The Dungey cycle begins when reconnection at the subsolar magnetopause opens previously closed terrestrial magnetic flux. These open field lines are swept antisunward by the flow of the solar wind to form an extended magnetotail [[Dungey, 1965](#); [Cowley, 1991](#); [Milan, 2004a](#)]. Eventually, reconnection in the central plane of the tail recloses open flux, this closed flux proceeding to the dayside to complete the cycle. This concept naturally explains several known features of the magnetosphere, such as the presence of an extended magnetotail [e.g., [Wolfe et al.,](#)

1967], erosion of the dayside magnetopause during periods of southward interplanetary magnetic field (IMF) [e.g., Aubrey et al., 1970; Holzer and Slavin, 1979], and the twin-cell ionospheric convection pattern [e.g., Cowley and Lockwood, 1992]. Support for magnetic reconnection being the mediator of the interaction can be found in the modulation of geomagnetic activity by the orientation of the interplanetary magnetic field, activity being a maximum when the IMF is directed southward such that reconnection at the subsolar magnetopause should be most efficient [e.g., Fairfield and Cahill, 1966; Fairfield, 1967], and numerous studies which have subsequently demonstrated the control of the ionospheric convection pattern by the orientation of the IMF.

[3] In Dungey's original proposal the rates of dayside and nightside reconnection were considered equal, and the Dungey cycle was envisaged as a steady-state process. However, there is no reason why the two reconnection rates should be instantaneously identical, the dayside rate being modulated by the orientation of the IMF and the nightside rate being governed by conditions in the magnetotail. Conditions in the tail will be influenced by the dayside reconnection rate, by for instance the loading of the magnetospheric lobes with open flux during southward IMF. However, the nightside reconnection rate cannot instantaneously readjust itself to changes in the dayside rate and there must be a lag between the two of at least the propagation time of the solar wind from the subsolar magnetopause to the tail, which can be several 10s of minutes. In this case the proportion of terrestrial magnetic flux that is open at any one time will vary and the ionospheric region threaded by this flux, the "polar cap," will expand or contract, as discussed by Russell [1972], Siscoe and Huang [1985], Cowley and Lockwood [1992], Lockwood and Cowley [1992], Taylor et al. [1996], Milan et al. [2003, 2006a], and Lockwood et al. [2005]. Siscoe and Huang [1985] encapsulated this expanding/contracting polar cap (ECPC) paradigm in the following formulation of Faraday's Law:

where FPC is the amount of open flux mapping into the northern and southern polar ionospheres, comprising the dim regions inside the auroral ovals, and FD and FN are the dayside and nightside reconnection rates. If FD exceeds FN, then the amount of open flux in the magnetosphere increases; conversely; if FN dominates, then the open flux decreases. Observations of changes in FPC can be used, with suitable assumptions, to quantify both FD and FN. Lockwood and Cowley [1992] discussed at length the relationship between the ECPC paradigm and the substorm cycle, specifically equating the growth phase with periods of open flux creation and the expansion phase with periods of open flux destruction and the combined action with the excitation of ionospheric convection; it is the aim of the present study to place this synthesis on a more quantitative basis.

Most of the tail reconnection events would be classed as substorms, though a few could be considered steady magnetospheric convection events (SMCs) and one occurred during a prolonged interval of northward IMF, during which no substorm signatures were observed. Up to 60% of the events could be identified with a possible trigger in the upstream conditions, either a northward or southward turning of the IMF or a step-like enhancement in the solar wind ram pressure. The remaining 40% occurred spontaneously during periods of ongoing southward interplanetary magnetic field.

[40] During periods of southward IMF the polar cap expands at a rate consistent with reconnection along an effective dayside magnetopause merging line of approximately 5 RE in length. During the 73 hours of observations the polar cap was not observed to exceed 0.9 GWb in flux or decrease below 0.2 GWb. Above polar cap fluxes of 0.5 GWb the occurrence of tail reconnection is greatly enhanced, which subsequently leads to a reduction of the open flux. We interpret our observations as providing strong supporting evidence for the expanding/contracting polar cap paradigm. In addition, we find no evidence for a trickle of nightside reconnection outside of the major events, making substorms the major nightside flux closure process.

[41] Tail reconnection events varied greatly in flux closure rate and duration, though with mean values of 85 kV and 70 min. On average 0.3 GWb of flux was closed in each event, with the mean open flux at onset being 0.65 GWb, with 0.45 GWb left remaining after; on average 0.1 GWb of flux was opened by ongoing dayside reconnection during each event. The probability of onset increases sharply once F PC exceeds about 0.7 GWb, which we interpret as a response to the increased stress in the tail.

### 2.7.3 Russell-McPherron effect

### 2.7.4 Geomagnetic storms

geomagnetic storms

geomagnetic storms (several days, caused by CMEs)

## 2. Background knowledge

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substorms  
substorm distance about 10 Re (VB)  
substorms (few hours, frequently, occur when plasmoids separate in the tail)

### 2.7.5 Geomagnetic indices

#### 2.7.6 $K_p$ index

## 2.8 Forecast methods

both Newell relations

### 2.8.1 $K_p$ forecast

$K_p$ -flux rate relation  
 $K_p$  prediction models/methods  
This paper adds more empirical relations.

### 2.8.2 CME forecast to Earth

CME forecast methods and models

### 2.8.3 Stream forecast to Earth

Stream forecast methods and models

## 2.9 Magnetosphere

### 2.9.1 Magnetopause reconnection

Phan2005, Magnetopause Processes:  
at the magnetopause, the large-scale spatial and temporal nature of reconnection is still not well understood.  
One of the unanswered questions is whether reconnection at the magnetopause is necessarily intermittent, or can be continuous, or even quasi-steady.

Phan2005, Magnetopause Processes:  
When the IMF changes, the reconnection site moves and flux transfer events (FTEs) (presumably due to intermittent reconnection) can be formed. The accumulated evidence suggests that the key controlling factor of the steadiness of the reconnection process is the steadiness of the solar wind and IMF. Thus intermittency does not appear to be a fundamental property of the reconnection process at the magnetopause.

Phan2005, Magnetopause Processes:  
The fundamental question is whether reconnection is intrinsically intermittent, even when the solar wind plasma and magnetic field orientation are steady. We describe two studies, one for southward and the other for northward IMF, which provide evidence that reconnection at the magnetopause is continuous when the solar wind conditions are steady. When the IMF changes direction, a joint IMAGE-Cluster study reveals that the reconnection site moves, but reconnection never stops. Thus reconnection is still continuous in a global sense but may appear intermittent (or transient) at the local magnetopause.

Phan2005, Magnetopause Processes:  
Both intermittent and continuous (and even quasi-steady) reconnection have been convincingly observed. The accumulated evidence suggests that the steadiness of the solar wind determines the steadiness of reconnection.

When the IMF and solar wind pressure are stable, reconnection at the local magnetopause can be continuous. If the IMF direction changes, the reconnection site moves and reconnection at the local magnetopause can become intermittent, but viewed globally, reconnection at the magnetopause never ceases.

Phan2005, Magnetopause Processes:

Tailward-of-the-cusp reconnection has been observed only when the IMF has a northward component. The occurrence rate of cusp reconnection is near 100% when the IMF is northward.

Phan2005, Magnetopause Processes:

Cluster findings include:

Convincing evidence for both intermittent and continuous reconnection at the magnetopause has been obtained from multi-point observations. Reconnection at the local magnetopause is continuous when the IMF is stable. When the IMF orientation changes, the reconnection site moves but reconnection never ceases. Thus reconnection may be intermittent at the local magnetopause but continuous on a global scale. Changes in the solar wind dynamic pressure can also change the reconnection rate and produce flux transfer events.

A strong ‘guide field’ detected at a reconnection X-line, i.e., a finite magnetic field along the X-line, has provided direct evidence for component merging.

Tailward-of-the-cusp reconnection has been found to occur only when the IMF has a northward component. The occurrence rate of cusp reconnection is nearly 100% when the IMF has a northward component, implying that cusp reconnection in the northern and southern hemispheres must be common. The high occurrence rate (in contrast to a rate of 50% at the subsolar magnetopause) is thought to be due to the presence of a plasma depletion layer. In this layer, the plasma beta is reduced, rendering the magnetosheath flow sub-Alfvenic and allowing the establishment of a stable X-line at the high-latitude magnetopause.

## 2.9.2 Russell-McPherron effect

annual Variation in Magnetic Disturbances

The semianual variation in geomagnetic activity has been recognized for a long period of time which shows the maximum appears around equinoxes while the minimum appears around solstices; by ([Cortie 1912](#))

A model, in which the interaction, ordered in solar magnetospheric (GSM) coordinates, is zero for northward components of the interplanetary field while the interaction is proportional to magnitude of the southward components predicts the correct phase and provides a yearly variation in the strength of the interaction sufficient to cause the observed effect. ([Russell & McPherron 1973](#))

This mechanism is the varying probability throughout the year of a southward component of the interplanetary magnetic field as seen by the magnetosphere. This arises from the changing orientation between the solar equatorial coordinate system (GSEQ), in which the interplanetary field is ordered, and the solar magnetospheric coordinate system (GSM), in which the interaction with the interplanetary field is ordered. ([Russell & McPherron 1973](#))

the R-M effect holds that the angle between Z axis in geocentric solar magnetospheric (GSM) coordinate system and Y axis in geocentric solar equatorial (GSEQ) coordinate system plays an important role. Figure 1a shows the semianual and diurnal variation of the angle  $q$  between the Z axis in GSM coordinate system and the Y axis in GSEQ coordinate system, that is, the controlling parameter of the R-M effect. According to the R-M effect, the probability of southward IMF increases when the angle  $q$ , which is smaller than 90 degrees, decreases, so that the dayside reconnection can be more efficient and more energy can be conveyed into the magnetosphere. ([Zhao & Zong 2012](#))

whereas Russell et al. [2003] demonstrate that the tilt of dipole axis controls the size of dayside reconnection region and thus the reconnection rate and geomagnetic activity. Nowadays, the R-M effect is one of the most prevailing hypotheses. ([Zhao & Zong 2012](#))

Thus, the R-M effect with positive/negative IMF polarity is more reasonable to explain seasonal and diurnal variation of geomagnetic activity under different IMF polarity and extreme solar wind conditions. ([Zhao & Zong 2012](#))

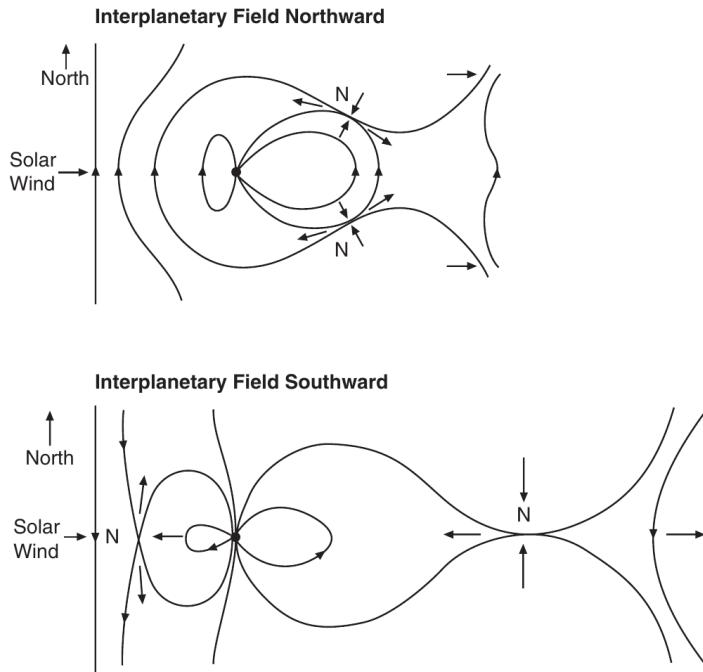
solar-terrestrial tilt angle and orbit geometry figure?

### 2.9.3 Solar wind coupling

(qualitative)

at Earth the solar wind total energy flux ( $1.45 \text{ mW/m}^2$ ) is only about one millionth of the solar radiation flux (see Schwenn (1990, p. 153))

reconnection behind the cusps or at the front of the magnetopause (see Figure 2.29)



**Figure 2.29** Sketch indicating where reconnection of magnetic field lines on the magnetopause occurs, for the cases of northward (upper panel) and southward (lower panel) IMF. The central dot indicates the position of Earth and the arrowed lines the direction of the magnetic field. The locations of X-line reconnection are denoted with N and the short arrows indicate the movement of the field lines. Credit: Bothmer & Daglis (2007, p. 116, Fig. 4.8), adapted from Dungey (1961, 1963). get permission...

for overview see chapter 4 in Bothmer2007book

E-field produced by solar wind acts on magnetospheric plasma

$$\mathbf{E}_{\text{IMF}} = -\mathbf{V} \times \mathbf{B}_{\text{IMF}} \quad (2.1)$$

from Lorentz force  $\mathbf{F} = -e \mathbf{v} \times \mathbf{B}$

$\mathbf{E} = -\mathbf{F}/e$

Because of high plasma conductivity the E-field is not existent.

Newell et al. (2007) and Newell et al. (2008): coupling consists of merging and viscous part (reconnection and turbulence)

merging part: universal sw-magnetosphere coupling function; rate magnetic flux is opened at the magnetopause ( $d\Phi_{\text{MP}}/dt$ )

viscous part: reconnection due to Kelvin-Helmholtz instabilities at the boundary ( $n^{1/2} v^2$ )

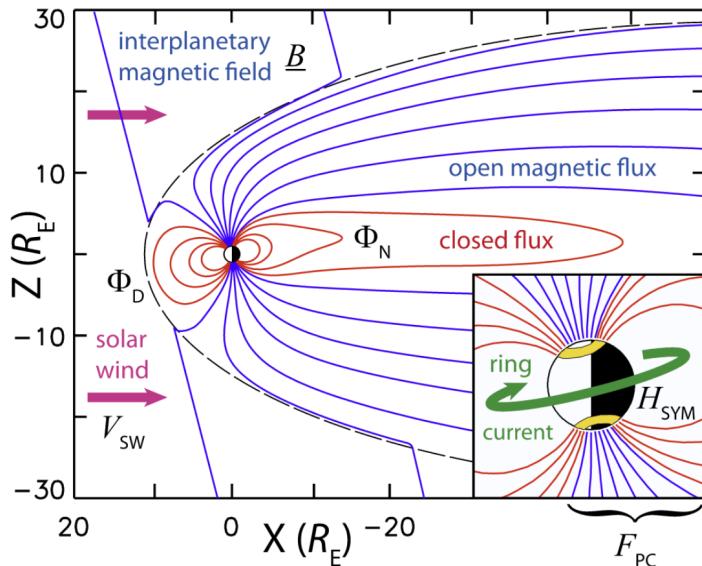
equation for the least variance linear prediction of Kp:  $Kp = 0.05 + 2.244 \times 10^{-4} d\Phi_{\text{MP}}/dt + 2.844 \times 10^{-6} n^{1/2} v^2$   
combination of both terms works best ( $r = 0.866$ )

Dungey convection cycle, see Milan2006  
magnetic flux cycle, see Figure 2.30

“the classical twin-cell convection streamlines. Pulsed nightside reconnection is occurring as part of the magnetospheric substorm cycle”

dayside reconnection:

“ $E_y$  is the rate at which southward magnetic flux is convected to the magnetosphere by the solar wind ( $-v_x \cdot B_z$ ) in GSM coordinates,” (Russell 2007)



**Figure 2.30** Schema of the magnetosphere for the case of southern IMF direction with emphasis on open (blue lines) and closed (red lines) magnetic flux during times of southern IMF. The magnetopause is indicated by the dashed line. The opening flux rate at the dayside is denoted with  $\Phi_D$  and the closing flux rate at the nightside with  $\Phi_N$ . The inset shows day- and nightside of the Earth and the relation between auroral ovals (yellow areas) and the open flux at the polar caps ( $F_{PC}$ ). The ring current intensity  $H_{SYM}$  and direction (green arrow) are plotted as well. Credit: Milan (2009, Fig. 1), © American Geophysical Union, reproduced with permission.

In this work I settle for  $vB_z$  as the coupling function – that is why I focus on this relation here.

The solar wind electric field  $E_y$  is the product of the proton velocity  $v$  and the magnetic field z-component  $B_z$ :

$$E_y = -v_x \times B_z. \quad (2.2)$$

Savani2017:

...a function is required that characterizes the coupling of the solar wind to the magnetosphere. Many functions that couple the solar wind to a wide variety of magnetospheric activity have been proposed in the past, often incorporating the magnetic field orientation [Lockwood et al., 2013, and references therein]. A recent study suggests that one parameter correlates best with 9 out of 10 indices of terrestrial activity [Newell et al., 2007]. This parameter,  $d\phi/dt$ , represents the rate at which magnetic flux is opened at the magnetopause and is defined as

$$\frac{d\phi}{dt} = v^{4/3} |B|^{2/3} \sin^{8/3}(\theta_c/2) \quad (2.3)$$

where  $v$  is the velocity of the solar wind;  $|B|$  is the magnetic field magnitude; and the interplanetary magnetic field (IMF) clock angle is defined by  $\theta_c = \tan^{-1}(By/Bz)$ . The predicted magnetic field time series in GSM coordinate system is used to calculate a theoretical magnetic flux rate.

Savani2017:

Extended periods of southward magnetic field, often found within and surrounding Earth-directed CMEs, result in periods of enhanced reconnection on Earth's dayside and energy input into the magnetosphere. In contrast, northward directed fields inhibit reconnection, resulting in a weaker magnetospheric response (Dungey 1961).

Milan2006: “We find no evidence for a low background rate of nightside reconnection between these events and conclude that substorms and other large reconnection bursts provide the major or only source of flux closure on the nightside.”

## 2.9.4 Geomagnetic storms

(causes)

Solar wind influences the terrestrial magnetosphere and can cause strong disturbances therein that are called geomagnetic storms.

Forbush1937 found the rapid reduction in cosmic ray intensity (Forbush decrease) during geomagnetic storms (caused by CMEs).

## 2. Background knowledge

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Gosling (1993) pointed out that the emphasis on flares is misplaced as far as geomagnetic storms and large SEP events are concerned and he referred to this as 'solar flare myth'. Gosling (1993)'s conclusions represented the culmination of research on CMEs in the 1970s and 1980s bringing them to the forefront as the main players in affecting Earth's space environment.

Coronal mass ejections have the central role in the chain of events leading from the Sun to near-Earth space (relationship between solar activity and interplanetary and geomagnetic events) (Gosling 1993).

Coronal mass ejections (CMEs) have been established as the primary source of major geomagnetic storms and large solar energetic particle (SEP) events (Gosling 1993)

88 major geomagnetic storms: single ICME, multiple ICME, and CIR events; numbers 53 (60%), 24 (27%), and 11 (13%) (Zhang et al. 2007).

The critical connection between CMEs and geomagnetic storms was made by Wilson (1987), who found that the southward component of the magnetic field in magnetic clouds is responsible for the storms. Wilson (1987) performed a superposed epoch analysis of 19 magnetic clouds and found that the Dst index simultaneously decreased to a large negative value at the onset of a large, sustained southward magnetic field in the magnetic cloud; the storm recovery started when the magnetic field turned northward.

Results from this study clearly show that magnetic clouds (bubbles) are associated with geomagnetic storms and that the association is statistically significant. Wilson (1987)

connection between geomagnetic disturbances and the southward component of the external interplanetary magnetic field ((Dungey 1961); Fairfield and Cahill Jr 1966).

(Zhang & Burlaga 1988): The magnitude of the change in the Dst index for the case when southward fields arrive first is comparable to that for the case when northward fields arrive first, and the phase is such that geomagnetic activity is associated with southward fields.

Savani2017:

As CMEs usually generate the most geoeffective events [e.g., Tsurutani and Gonzalez, 1997; Zhang et al., 2007; Zhang and Moldwin, 2014],

### 2.9.5 Geomagnetic indices

Geomagnetic observatories are distributed widely over the globe, measuring the local magnetic field at their position. Several sets of stations, covering specific regions, are defined to monitor the state of different parts of the magnetospheric system. Magnetic measurements from these sets define several geomagnetic indices. The International Association of Geomagnetism and Aeronomy (IAGA) supports the following global geomagnetic indices which are serviced by the International Service of Geomagnetic Indices (ISGI)<sup>12</sup>: The *aa* index is designed to represent the amplitude of the global geomagnetic activity, normalized to a geomagnetic latitude of  $\pm 50^\circ$ . The *am* index characterizes the global geomagnetic activity. The *Kp* index is designed to measure geomagnetic disturbances from solar particle radiation – it is described in more detail in the following section. The *Dst* index monitors the intensity of the magnetospheric ring current. The *PC* index monitors the polar cap magnetic activity – it approximates the amount of energy which entered the magnetosphere through solar wind coupling. The *AE* index and its relatives *AU*, *AL* and *AO* measure the magnetic effects of the northern auroral electrojet. The first three listed indices (*aa*, *am* and *Kp*) are calculated from different sets of local 3-hourly *K* indices, which measure the local magnetic disturbances at the observatories.

Lockwood et al. (2014) even used geomagnetic indices to reconstruct the near-Earth solar wind magnetic field strength and velocity back to 1845. which indices?

### 2.9.6 *Kp* index

The MC plotted in Figure 2.21 created a *Kp* of 6+.

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<sup>12</sup>ISGI website: <http://isgi.unistra.fr/>

## 2.10 Forecast methods

"The principal users affected by geomagnetic storms are the electrical power grid, spacecraft operations, users of radio signals that reflect off of or pass through the ionosphere, and observers of the aurora." NOAA cite

Methods:

- MC Kp impact forecast; Savani2015+Savani2017
- WSA-ENLIL sw forecast (CMEs without B-field)
- CH sw forecast (Uni Graz)
- NOAA/SWPC; see Savani2017
- NASA/SWRC; see Savani2017

### 2.10.1 *Kp* forecast

(quantitative)

Space weather effects can be directly attributed to CME sub-structures. While the outermost structure (shock) accelerates particles and causes sudden commencement, the interior structures (sheath and flux rope) cause geomagnetic storms if they possess a strong southward component of the magnetic field ([Gopalswamy 2016](#)).

see Wing Kp model description...

Southward interplanetary field can also occur in the shock sheaths (Gonzalez and Tsurutani 1987; Gosling and McComas 1987), which have been shown to be equally important in causing geomagnetic storms ([Tsurutani et al. 1988](#)).

The results of this study indicate the equal importance of both sheath fields or draped fields and driver gas fields for the generation of major geomagnetic storms. Because of the importance of the sheath fields the intensity and duration of geomagnetic storms cannot be predicted by solar observations of active regions alone ([Tsurutani et al. 1988](#)).

read Savani2015: Predicting the magnetic vectors within coronal mass ejections arriving at Earth: 2. Geomagnetic response

It makes sense to use the geocentric solar magnetospheric (GSM) coordinate system, which is aligned with the Earth's magnetic dipole axis.

Savani2015:

The compressed solar wind plasma in between supersonic magnetic flux rope obstacles and their driven shock fronts has not been addressed in this article, even though they have been shown to be significant drivers of magnetospheric storms [Huttunen and Koskinen, 2004].

Savani2017:

The correlation coefficient of  $d\phi/dt$  with the *Kp* index is  $r = 0.76$  [see Newell et al., 2007, Table 3]. A Kp prediction is then generated by the empirical correlation.

$$Kp(BSS) = 9.5 - e^{A-B(d\phi/dt)} \quad (2.4)$$

where  $A = 2.18$ ,  $B = 5.20 * 10^5$ , and with the velocity and magnetic field measured in km/s and nT, respectively [Emmons et al., 2013; Mays et al., 2015a].

### 2.10.2 CME forecast

Predicting the magnetic field variations in ICMEs from solar observations is a central subject for space weather forecasting because a long-lasting southward-directed field is a primary driver of major geomagnetic storms ([Zhang et al. 2007](#)).

## 2. Background knowledge

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Identification of a flux rope near the Sun makes it easier to predict the expected magnetic configuration at Earth, modified by the interaction with the solar wind on its way to Earth.

Bz prediction:

read Savani2015: Predicting the magnetic vectors within coronal mass ejections arriving at Earth: 1. Initial architecture

Savani2015:

Thus, models used for routine CME forecasts by various space weather services do not include magnetic structures within the simulated CMEs [e.g., Zheng et al., 2013; Shiota et al., 2014]. For example, ENLIL models the propagation of CMEs from 20 solar radii ( $R_s$ ) to beyond Earth at 215  $R_s$  and includes the background solar wind magnetic field. However, the CME is simplified to a high-pressure plasma pulse with a size and propagation direction estimated from solar imagery [Zheng et al., 2013]. CME arrival time predictions from these models provide lead times of 2–3 days, and their accuracy has been well investigated [Taktakishvili et al., 2010; Vršnak et al., 2014; Colaninno et al., 2013]. In contrast, the important magnetic vector information is only revealed when in situ measurements are made by spacecraft upstream of Earth at the first Lagrangian position (L1) 1 h prior to the CME arriving at Earth, thereby severely limiting the lead time available for reliable, magnetic field-based, storm warnings.

Savani2015:

In this paper, we highlight three key components of a proof of concept developed to improve the prediction of a storm's severity: (1) the use of the hemispheric helicity rule to provide a robust initial magnetic configuration at the Sun; (2) define a “volume of influence” of the CME, within the heliosphere, for which the Earth's trajectory can be estimated; and (3) incorporating magnetic vectors from a simplified magnetic flux rope model to create a time series upstream of Earth.

constant-alpha force-free (CAFF) flux ropes (Burlaga1988)

adjusted BSS scheme (Savani2015)

Savani2017:

inferring the magnetic field direction along the trajectory of Earth through the CME would be a major advance in geomagnetic activity prediction.

Savani2017:

This event highlights the need to forecast the field vectors not only in the flux rope but also in the sheath [e.g., Huttunen and Koskinen, 2004].

It is important to know in advance if and when CMEs arrive at Earth, because of their possible strong impact on the terrestrial magnetosphere. Knowing the geometry of a CME, it is possible to infer its direction and speed from white-light images. Through kinematic forward-modeling techniques that take advantage of the self-similar expansion observed in CMEs, it is possible to derive an estimated arrival time at Earth. As MCs are the drivers for most severe geomagnetic disturbances (Bothmer & Schwenn 1995; Cane & Richardson 2003), it is also important to forecast their magnetic field strength and configuration.

An early 3D model for a general CME geometry is the ice-cream cone model, which assumes a simple bubble-like CME structure (Fisher & Munro 1984). Though, building on the study by Cremades & Bothmer (2004), Thernisien et al. (2006) created the more complex graduated cylindrical shell (GCS) model for flux rope-like CMEs. This model assumes an empirically derived electron distribution in order to create synthetic coronagraph images. The GCS model is successfully applied to images of CMEs with well-developed white-light structure (Bosman et al. 2012) and is implemented as a part of current CME forecast procedures.

Depending on whether CMEs are faster or slower than the surrounding solar wind, they decelerate/accelerate on their way through the heliosphere due to the drag forces. Fast CMEs already decelerate substantially during their first few solar radii (Sachdeva et al. 2017). In case a CME interacts with a preceding CME, they both travel on with an intermediate speed (Manoharan et al. 2004; Temmer et al. 2012).

Extreme CME velocities above  $2000 \text{ km s}^{-1}$  near the Sun are rare, nevertheless, in some cases speeds around  $3000 \text{ km s}^{-1}$  were measured. From the free energy available in active regions, Gopalswamy et al. (2005) concluded that CME speeds cannot be far greater than  $3000 \text{ km s}^{-1}$ . This would implicate that CMEs need at least half a day to travel from their source region on the Sun to reach Earth. Gopalswamy et al. (2010) even esti-

mated a maximal limit of free energy active regions can store, which means that CMEs with  $4000 \text{ km s}^{-1}$  are not possible. Faster CMEs usually come with higher magnetic field strengths, for example the CME observed by STEREO A on 23 July 2012 had a shock speed of about  $2000 \text{ km s}^{-1}$  at 1 au and its magnetic field magnitude reached values of up to 100 nT ([Russell et al. 2013](#)).

observed coronal transient: 3d structure, directed at Earth, associated with shock wave; ([Howard et al. 1982](#))  
-> halo CME

### 2.10.3 Stream forecast

Cranmer2017:

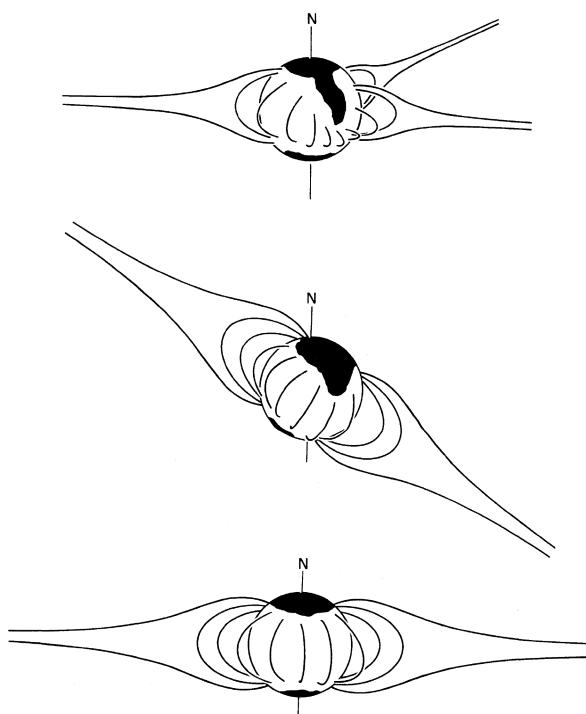
Complicating our understanding even more is the fact that processes such as turbulence, stream-stream interactions, and Coulomb collisions can make it difficult to unambiguously map a parcel measured at 1 AU back down to its coronal source.

## 2.11 notes...

These events (GS) are able to disturb sensitive technical systems. Understanding the properties of the solar wind helps with the prediction and protection of humanity's assets.

reference to [Bothmer & Daglis \(2007\)](#), maybe images

put fig. somewhere earlier (p. 10?): see [Figure 2.31](#)



**Figure 2.31** Schemata of different coronal configurations during the solar cycle. Visualized are the locations of closed coronal magnetic fields (lines) and coronal holes (black areas) for a post maximum distorted dipole (top panel), for a pre minimum stable tilted dipole (middle panel), and a post minimum axial dipole (bottom panel). Credit: ([Hundhausen 1977](#), Fig. 20, panels (b-d)), © Colorado Associated University Press, reproduced with permission.

The weaker solar activity cycle 24 seems to have important consequences for CMEs in the heliosphere: CMEs expand anomalously due to reduced heliospheric pressure leading to the increased observed rate of small CMEs, halos originating far from the disk center, and mild space weather ([Gopalswamy et al. 2014](#)),

## [2. Background knowledge](#)

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2015a; Petrie 2015).

The total pressure in the heliosphere (magnetic + plasma) is reduced by ca. 40%, which leads to the anomalous expansion of CMEs explaining the increased slope. The excess CME expansion contributes to the diminished effectiveness of CMEs in producing magnetic storms during cycle 24, both because the magnetic content of the CMEs is diluted and also because of the weaker ambient fields ([Gopalswamy et al. 2014](#)).

magnetosphere variations due to solar wind

magnetosphere protects from cosmic radiation (maybe from solar wind stripping atmosphere away?)

The coupling between the solar wind and the magnetosphere is governed by reconnection and compression of the magnetic field lines.

the dayside reconnection is asymmetric

half-wave rectifier coupling

To describe this, coupling functions with different complexity were proposed (Newell, cites? and list). Newell's relation is being implemented into forecast procedures ([Savani2017](#))

causes (see [citetRangarajan1997](#) p. 1282 and mention [Bartels1963](#) too)  
read [Bothmer1998 Ch 3...](#)

[Sonnerup & Cahill \(1967\)](#): The rotational discontinuity seems to occur predominantly during magnetic storms and two of these cases, involving substantial normal-field components, provide compelling evidence that field reconnection takes place during the storm main phase.

read [Milan2009](#)

sources of southward Bz:

- MCs
- compressed material in front of CMEs (highest driving velocities)
- SIRs/CIRs front of HSSs
- sw streams

3-part structure figure ... find event with shock and flux rope?  
abbreviate magnetic flux rope to MFR?

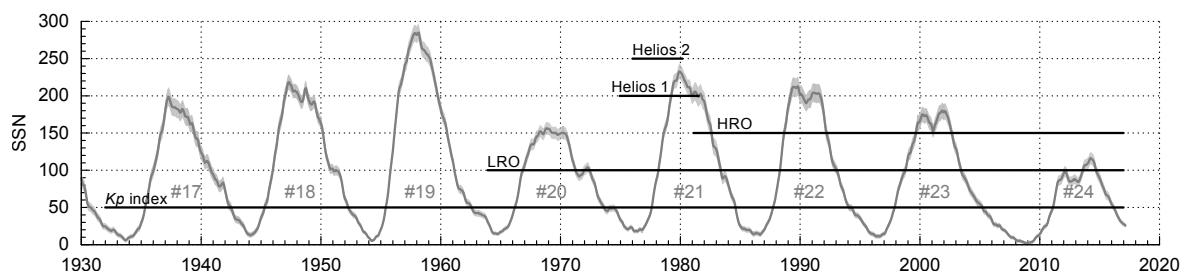
why HCS? what currents are there...

why is the CS not the divide between both polarities?

### 3 Instrumentation and data description

For analyzing the solar wind and related areas on the Sun, there exist remote instruments, such as solar imagers and coronagraphs, and in-situ instruments, such as magnetometers and plasma detectors. In this chapter the basic principles of the latter are described, because the analyses performed in this thesis are based entirely on in-situ measurements – except for the sunspot time series.

Different kind of data, collected from observatories at various locations, are used for the investigations performed in this thesis. In the following sections, I describe the data sets of the magnetospheric disturbance index  $K_p$  and of the solar activity indicator sunspot number (SSN). Further, the solar wind in-situ data sets are described: the near-Earth OMNI data collections consisting of minutely and hourly data, that is, low resolution OMNI (LRO) and high resolution OMNI (HRO), and the data obtained from the Helios 1 and Helios 2 space-craft in the 1970s. I plotted a time coverage overview of these data sets in relation to the solar activity cycles in [Figure 3.1](#).



**Figure 3.1** Time coverage of the data sets used in this work with the solar activity from 1930 until end of 2016. The individual data sets are the  $K_p$  index, the low and high resolution OMNI (LRO and HRO) collections, and the Helios 1 and Helios 2 data sets. I plotted the SIDC 13-month smoothed monthly SSN and the cycle number in the background for orientation.

#### 3.1 Magnetometer

Spacecraft nowadays carry two different magnetometer types, one for measuring the magnetic field direction and its strength and the other for observing the magnetic flux and detecting waves.

A flux-gate magnetometer consists of two coils around a core – one coil with alternating current, which is compared with the induced current signal from the other. Without external magnetic field both patterns match. The core is easier magnetized in direction of an existing external magnetic field, in which case the patterns differ. It measures...

In a search coil magnetometer one coil is placed around a core; measures plasma waves – where?

Because these magnetometer types are directional, they often are placed in two sets of triaxial configurations, attached on booms to minimize the influence of the spacecraft's own magnetic field.

L-> which is generated by surface charges?/electrons?/ionization?/the instruments?

figures and references...!

Ground based magnetometers measure the local magnetic field of the Earth. They already exist since...

examples:

ACE/MAG – flux-gate magnetometer

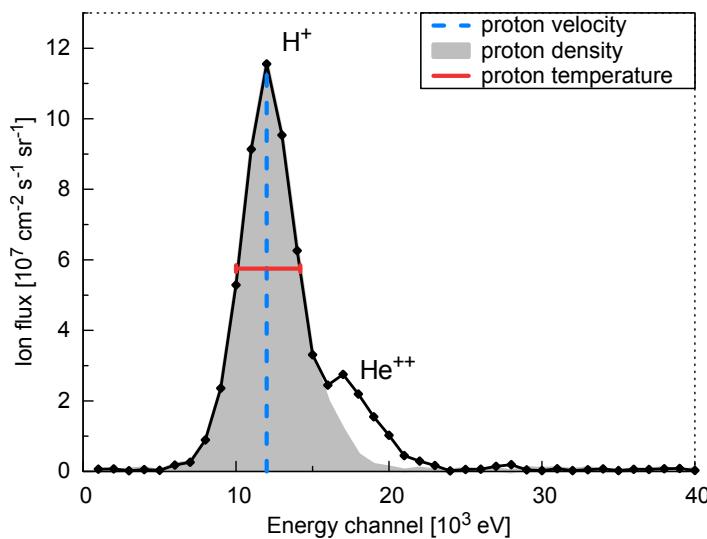
## 3.2 Plasma detector

several spectrometers with different energy ranges

- isotope spectrometer - isotopic abundances of SEPs
- ionic charge analyzer - charge state of SEPs
- solar wind ion mass spectrometer -
- solar wind ion composition spectrometer -
- radio burst tracker

A plasma detector measures the ion energy frequency distribution, which consists basically only of protons and alphas in solar wind. I prepared a synthetic ion energy spectrum for illustration in [Figure 3.2](#).

...read <http://www.goembel.biz/sun.html>



**Figure 3.2** Example of an ion energy spectrum for which I prepared some synthetic data points, including a small helium peak. Capitalize Proton...

The velocity, density, and temperature can be derived from the energy spectrum.  
more details...!

The bulk velocity is derived from the distribution's average energy (or position of Maxwell-Boltzmann distribution?).

The number density is the area of the distribution.

The temperature scales with the distribution's width.

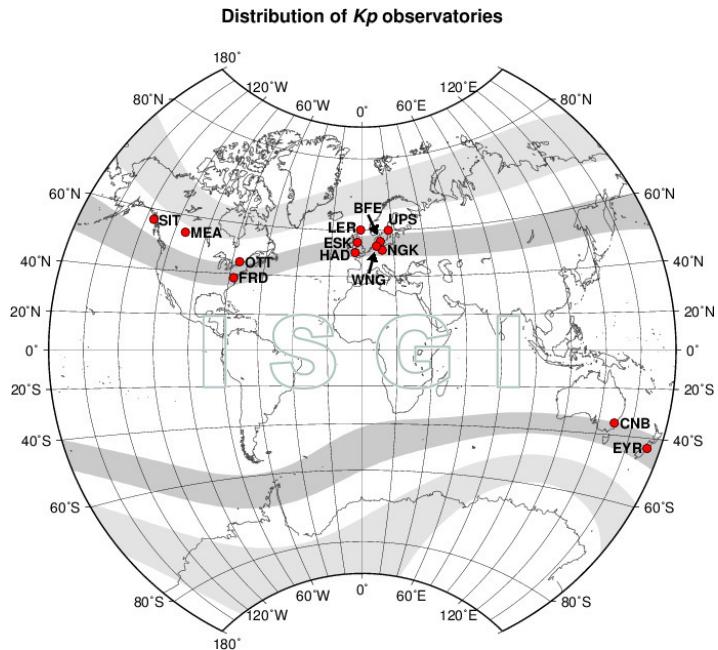
ACE/SWEPAM

## 3.3 K<sub>p</sub> data

Julius Bartels introduced the *K* index in 1938 and designed it to measure the intensity of geomagnetic disturbances ([Bartels et al. 1939](#)). Its name originates from 'Kennziffer' – the german word for characteristic digit. The *K* index is a measure for the maximal variation of the surface magnetic field, observed in a magnetogram within 3-hour intervals. Its scale in the range 0–9 is a quasi-logarithmic representation of the actual magnetic field strength's variations.

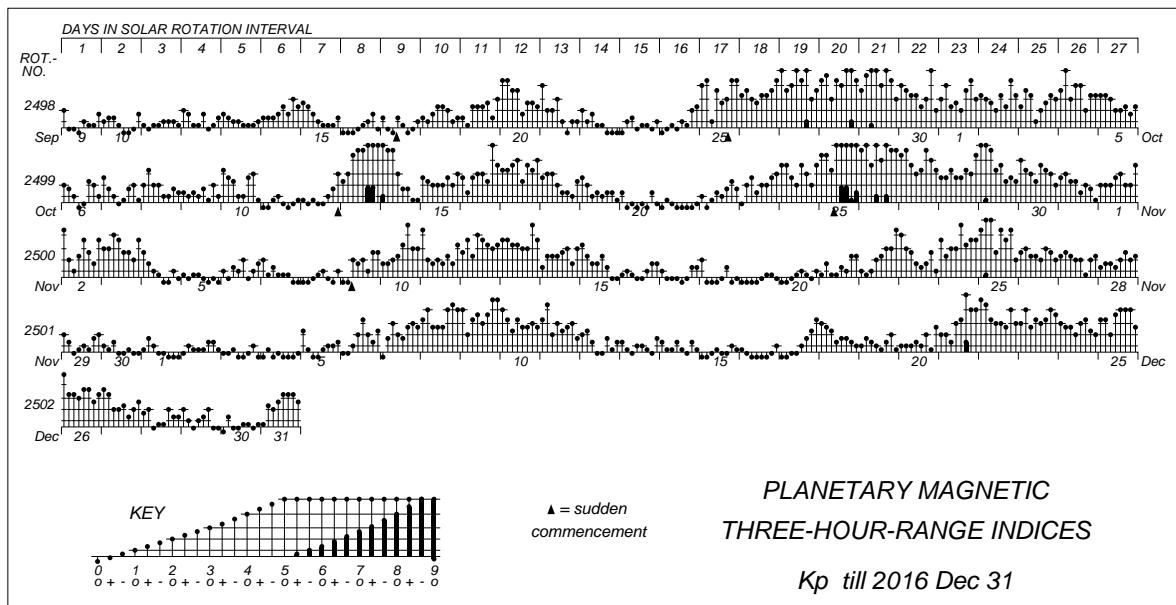
The Planetary *K* index (*K<sub>p</sub>*) is a planetary geomagnetic disturbance index, introduced by Bartels in 1949 at the Institute for Geophysics, University of Göttingen ([Bartels & Veldkamp 1949](#)). *K<sub>p</sub>* is the weighted average of 13 *K* indices, which are the standardized versions of the local *K* indices measured at 13 observatories. These contributing observatories are located around ±50° geomagnetic latitude and their distribution is biased towards Europe, see the map in [Figure 3.3](#).

To benefit from its higher precision, its scale, in the range 0–9 as well, is further divided into thirds, represented by the suffixes '+', 'o' and '-' (e.g., 3o, 3+, 4-, 4o). The *K<sub>p</sub>* indices are often visualized in musical



**Figure 3.3** Distribution of the 13  $K_p$  observatories. The shaded belts indicate regions of equal geomagnetic latitude. Courtesy of International Service of Geomagnetic Indices (ISGI), 2013.

diagrams, where they are stacked into periods of 27 days to enable the detection of recurrent activities, as seen in Figure 3.4.



**Figure 3.4** Bartels musical  $K_p$  diagram for the time period from September until end of December 2016. Two sudden commencements with following geomagnetic storms, having a maximal  $K_p$  of 6+, can be seen in October. Credit: GFZ Potsdam, 2017, licensed under CC BY 4.0.

The  $K_p$  index can be converted to the 3-hour equivalent  $ap$  index, which represents the magnetic field strength at a surface position of about 50° dipole latitude. The conversion is done via a table specified by Bartels, in which the value of the  $ap$  index is scaled in units of 2 nT, as seen in Table 3.1. There are further geomagnetic indices which are derived from the  $K_p$  index. They include  $Ap$ , the daily  $ap$  average,  $Cp$ , the daily  $ap$  sum mapped via a fixed table to the range 0–2.5, and  $C9$ , a mapping of  $Cp$  via a fixed table to the range 0–9. The definitions of Q-days (quiet days) and D-days (disturbed days) are also obtained from the  $K_p$  index.

The International Association of Geomagnetism and Aeronomy (IAGA) adopted the  $K_p$  index in 1954. The  $K_p$  index was maintained in Göttingen until January 1997 – now the German Research Centre for Geosciences (GFZ) in Potsdam supplies the  $K_p$  index and thereof derived indices. The GFZ provides historical

### 3. Instrumentation and data description

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**Table 3.1** Table for the fixed conversion from the  $K_p$  index to the equivalent  $ap$  index, which represents the magnetic field strength in units of 2 nT.

$K_p$	0o	0+	1-	1o	1+	2-	2o	2+	3-	3o	3+	4-	4o	4+
$ap$	0	2	3	4	5	6	7	9	12	15	18	22	27	32
$K_p$	5-	5o	5+	6-	6o	6+	7-	7o	7+	8-	8o	8+	9-	9o
$ap$	39	48	56	67	80	94	111	132	154	179	207	236	300	400

and quicklook data of the indices via their website<sup>1</sup>. The data series was extended backwards using existing measurements and is now available from 1932 onwards.

There exist several indicators/quantities that scale or are based on the  $K_p$  index:

- NOAA's Space Weather Prediction Center (SWPC) developed its NOAA G-Scale<sup>2</sup> for geomagnetic storms which relates the  $K_p$  index to five levels from G 1 to G 5.
- The equatorward auroral boundary position correlates with the  $K_p$  index (cite?).
- The variation of the total electron content (TEC) of the ionosphere correlates with the  $K_p$  index (cite?). The TEC has influence on global navigation satellite systems (GNSS). A part of their positional error scales directly with TEC (in extreme cases up to about 30 m).

'The  $K_p$  index is designed to measure solar particle radiation by its magnetic effects.'

Because of the geomagnetic disturbance impacts on sensitive technical systems, various methods for now-and forecasting the  $K_p$  index are developed:

- GFZ Potsdam Nowcast  $K_p$  index<sup>3</sup>.
- USAF Wing  $K_p$  model (Wing et al. 2005); 1–4 hours foercast from real-time solar wind measurements<sup>4</sup>
- ?Alexej  $K_p$  correlation forecast model; developed within AFFECTS (link)
- In Chapter 4 of this work, I derive relations to enable  $K_p$  nowcast from in-situ solar wind measurements and to enable  $K_p$  forecast from remote CME/solar wind stream observations.

Savani2017:

$K_p$  is the global magnetospheric index often used by forecasters to indicate the severity of a space weather event (Wing et al. 2005).

Savani2017:

The  $K_p$  difference of 1.5 is tested as there is evidence that a limitation in accuracy is present in the underlying empirical  $K_p$  formulation [Mays et al., 2015a]. [...] which is approximately consistent with recent results by Mays et al. [2015a], the  $K_p$  empirical formulation is accurate to about  $K_p = 1.5$ .

see also 2.7.6

## 3.4 Sunspot number

The number of sunspots occurring on the solar surface is commonly used as a long-term solar activity indicator, for more information on solar activity and sunspots see Section 2.3. The international sunspot number (SSN) is maintained by the World Data Center – Sunspot Index and Long-term Solar Observations (WDC-SILSO) at the Solar Influences Data Center (SIDC), Royal Observatory of Belgium (ROB). The SIDC provides an online catalogue<sup>5</sup>, where I obtained the SSN data. The SSN version 2.0, the current recalibrated version introduced in July 2015, is used in this work. Chapter 4 uses the SSN data from the time period 1932–2016, see Figure 3.1.

Short-term predictions of the SSN are provided by several institutions. The SIDC itself provides 12-month SSN forecasts derived via different methods. The Space Weather Prediction Center (SWPC) at the National

<sup>1</sup>GFZ website for geomagnetic indices: <http://www.gfz-potsdam.de/en/kp-index/>

<sup>2</sup>NOAA Space Weather Scales website: <http://www.swpc.noaa.gov/noaa-scales-explanation>

<sup>3</sup>GFZ Potsdam Nowcast  $K_p$  index website: [http://www-app3.gfz-potsdam.de/kp\\_index/ql\\_bar.gif](http://www-app3.gfz-potsdam.de/kp_index/ql_bar.gif)

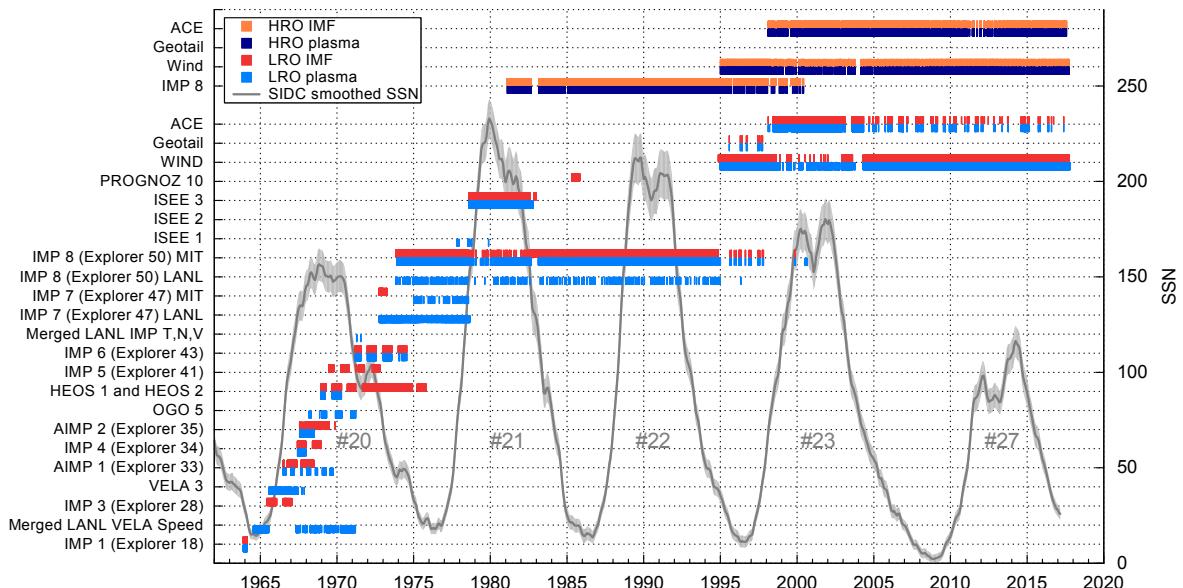
<sup>4</sup>USAF Wing  $K_p$  model website: <https://www.swpc.noaa.gov/products/wing-kp>

<sup>5</sup>WDC-SILSO website: <http://www.sidc.be/silso>

Oceanic and Atmospheric Administration (NOAA) supports the SSN prediction of the Solar Cycle 24 Prediction Panel<sup>6</sup>.

### 3.5 OMNI data collection

Solar wind was measured in situ for the first time by spacecraft in 1959 and since 1963 near-Earth measurements were done almost continuously. The OMNI 2 data collection (King & Papitashvili 2005) merges data from solar wind magnetic field and plasma, energetic proton fluxes, geomagnetic indices, and solar indices. The included solar wind data starts in 1963-11-27, the temperature data not before 1965-07-26, and is continuously maintained until today. As the data covers decades from multiple spacecraft at varying locations, the solar wind data is composed of intercalibrated data, which has been time-shifted to the nose of the magnetosphere's bow shock upstream of Earth. I created an overview of the various spacecraft contributing to the IMF and solar wind plasma data and their time coverages to the data set in Figure 3.5.



**Figure 3.5** IMF and solar wind plasma data sources (spacecraft) for the high and the low resolution OMNI (HRO and LRO) data sets until the end of 2016. I plotted this figure using the spacecraft identifiers noted in the OMNIWeb Data Documentation<sup>7</sup>. The SIDC 13-month smoothed monthly SSN and the cycle number are plotted in the background.

The OMNI data set is being maintained at NASA's Space Physics Data Facility (SPDF), Goddard Space Flight Center (GSFC). Their OMNIWeb interface<sup>8</sup> and their Coordinated Data Analysis Web<sup>9</sup> (CDAWeb) provide the data – the latter is where I obtained it.

Zhang2015 do with their data: “The data have been lagged by 5 min to allow for propagation from the nose of bow shock to the magnetopause.”

### 3.6 Helios probes

In order to observe solar wind in situ within the inner heliosphere, the nearly identical solar probes Helios 1 and Helios 2, one is pictured in Figure 3.6, were launched in December 1974 and January 1976 respectively.

Until today, these both probes were the only spacecraft that measured solar wind in situ over large solar distance ranges with perihelia as close as 0.31 au and 0.29 au respectively. Their highly elliptical orbits in the ecliptic covered a solar distance range up to 0.98 au. Launched during solar cycle minimum, the data of both probes cover the rise to the maximum of solar cycle 21, that amounts to about 6.5 years of data at varying

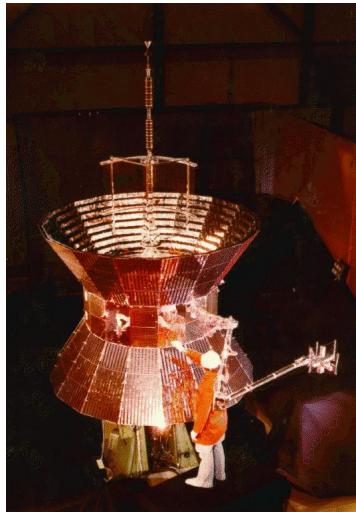
<sup>6</sup>Solar Cycle 24 Prediction Panel website: <http://www.swpc.noaa.gov/products/solar-cycle-progression>

<sup>7</sup>OMNIWeb Data Documentation: [https://omniweb.gsfc.nasa.gov/html/ow\\_data.html](https://omniweb.gsfc.nasa.gov/html/ow_data.html)

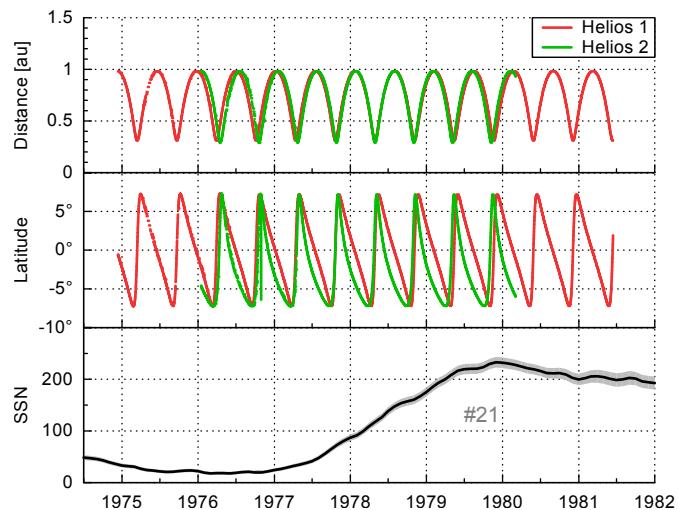
<sup>8</sup>GSFC OMNIWeb interface: <http://omniweb.gsfc.nasa.gov/>

<sup>9</sup>GSFC CDAWeb interface: <http://cdaweb.gsfc.nasa.gov/>

<sup>10</sup>I was not able to find out which Helios spacecraft this is.

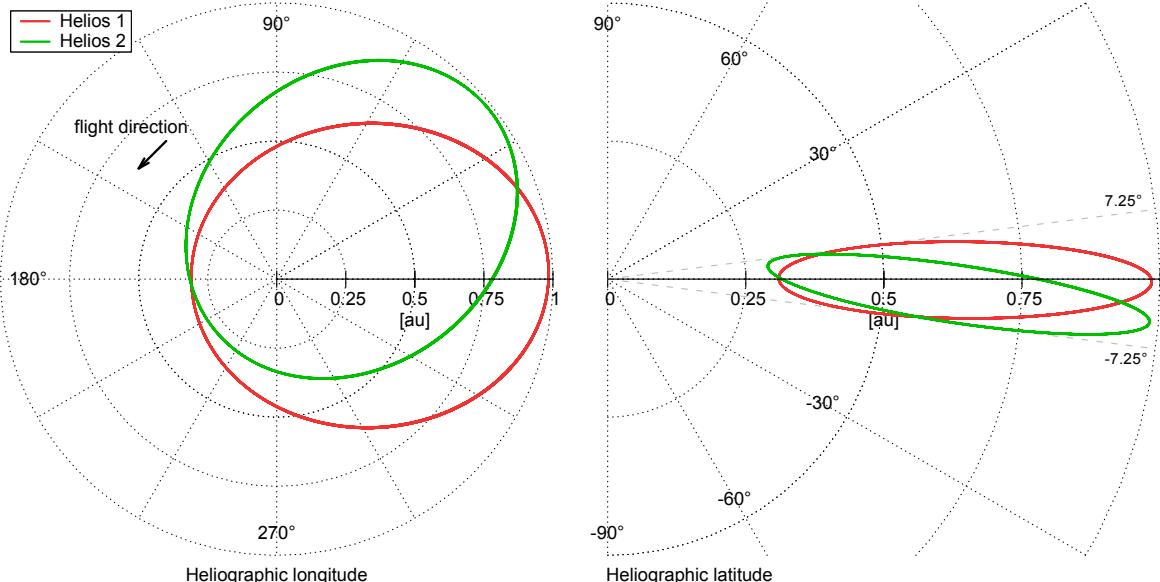


**Figure 3.6** One of the nearly identical twin Helios spacecraft<sup>10</sup>. Credit: NASA/Max Planck Institute for Solar System Research.



**Figure 3.7** Solar distance (top) and heliographic latitude (middle) of Helios 1 (red) and Helios 2 (green) with respect to their mission time. The trajectory data are from GSFC/SPDF and is plotted in HGI coordinates. I plotted the SIDC 13-month smoothed monthly SSN and its cycle number in the bottom panel for orientation.

solar distances. I plotted the probes' solar distance and heliographic latitude with respect to their mission time and sunspot number for illustration in Figure 3.7. The daily trajectory data are available from NASA's Space Physics Data Facility (SPDF) at the Goddard Space Flight Center (GSFC)<sup>11</sup> and are drawn in Heliographic Inertial (HGI) coordinates, see appendix Section A.6. I also plotted the orbits of Helios 1 and Helios 2 in a solar equatorial plane view and in a solar polar plane view, see Figure 3.8.



**Figure 3.8** Orbits of the Helios 1 (red) and Helios 2 (green) spacecraft in the solar equatorial plane (left) and in the solar polar plane (right). I obtained the trajectory data from GSFC/SPDF and plotted the orbits in HGI coordinates.

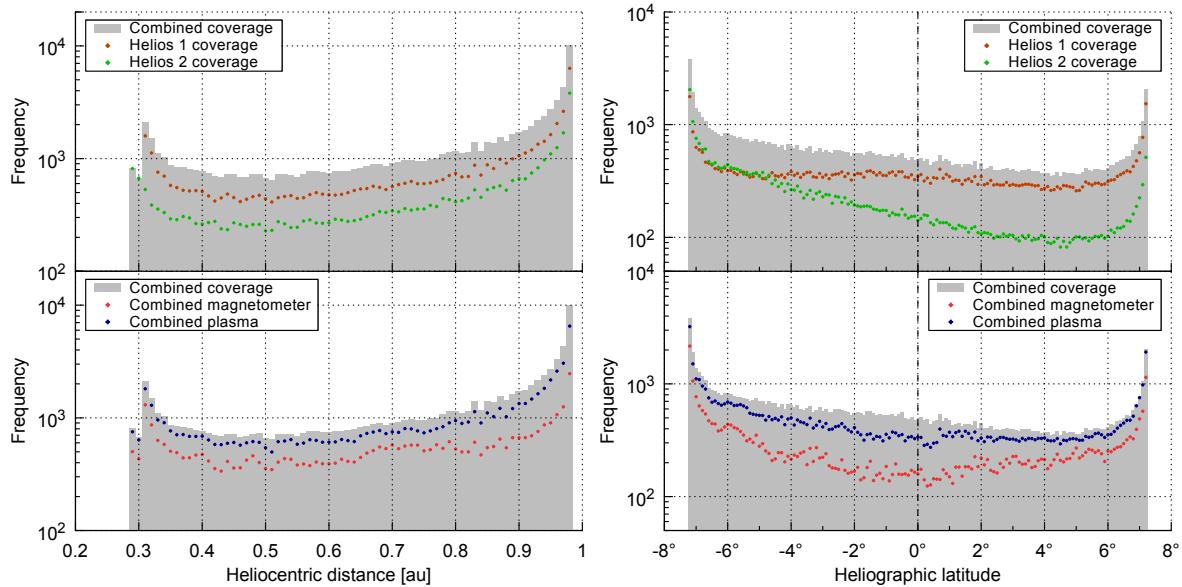
The scientific instruments carried by the spacecraft for measuring the magnetic field and solar wind plasma are two different Flux-gate Magnetometers and the Plasma Experiment Investigation. The data from the magnetometer and plasma instruments are merged into a data set with hourly resolution (Rosenbauer et al. 1977). In case of Helios 1 this data set includes about 12.5 orbits in the time range 1974-12-10 to 1981-06-14 and in case of Helios 2 about 8 orbits in the time range 1976-01-01 to 1980-03-04. The Helios data is available via

<sup>11</sup>SPDF Helios 1 trajectory data: <http://spdf.sci.gsfc.nasa.gov/pub/data/helios/helios1/traj/>

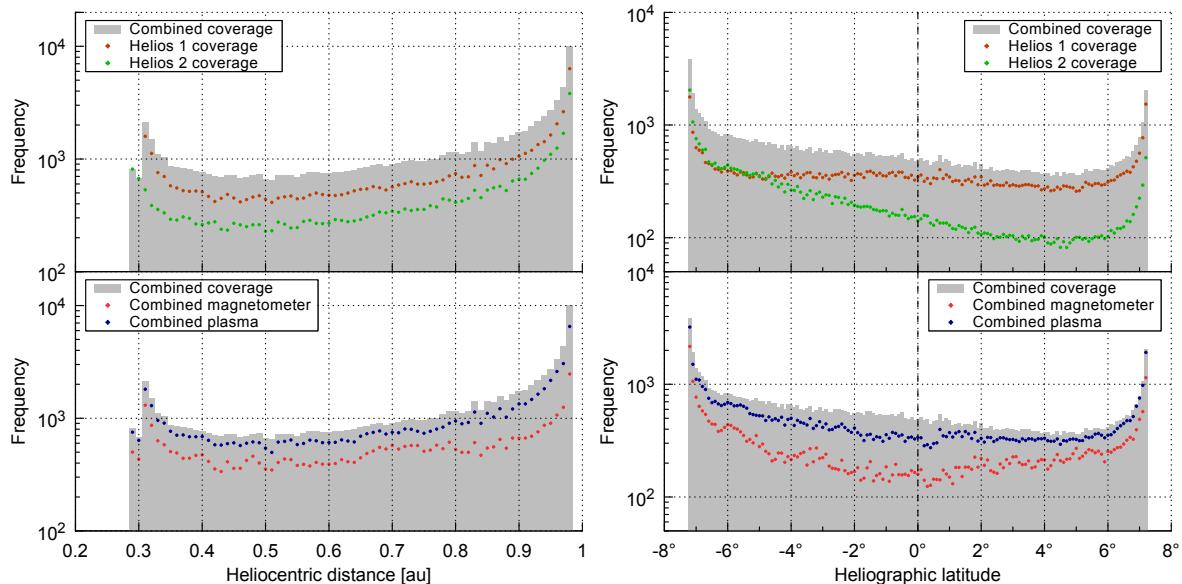
the GSFC/SPDF CDAWeb interface<sup>12</sup>.

The Helios 1 (Helios 2) magnetometer data coverage is about 43 % (54 %) and amounts to 2.8 years (2.3 years) in total. The plasma data coverage is 76 % (92 %) and amounts to 5.0 years (3.9 years) in total.

The Helios magnetic field and plasma data frequency over heliocentric distance and over heliographic latitude are plotted in Figure 3.9.



**Figure 3.9** Helios data frequency over heliocentric distance with bins of 0.01 au (left panels) and over heliographic latitude with bins of 0.1° (right panels). The frequency data is based on the hourly merged magnetometer and plasma data sets for Helios 1 and Helios 2. The top panels show the frequencies for Helios 1 and Helios 2 individually and the bottom panels those for the magnetometer and plasma data.



**Figure 3.10**

<sup>12</sup>GSFC CDAWeb interface: <http://cdaweb.gsfc.nasa.gov/>



## 4 Solar wind and CME influence on the magnetosphere

Impact estimations derived from empirical correlations between in-situ solar wind measurements and the geomagnetic  $K_p$  index

Variations in the Earth's magnetosphere are largely evoked by influence through the solar wind. These magnetospheric disturbances have diverse effects on the terrestrial environment. Especially the effects of severe geomagnetic storms created by coronal mass ejections (CMEs) pose various threats to sensitive technical systems and exposed humans. Thus, the development of quantitative forecasts for magnetospheric impacts caused by solar wind and CMEs is of major importance. The analyses in this chapter are based on my work done for the project Advanced Forecast For Ensuring Communications Through Space (AFFECTS).

This study's goals are to estimate the magnetospheric impact from solar activity in general, from solar wind and also to predict it for CMEs in particular. Empirical dependencies between specific solar wind parameters and the magnetospheric disturbance index  $K_p$  are presented. These dependencies allow to nowcast the  $K_p$  index from upstream (L1) solar wind in-situ measurements. Hence, also the magnetospheric impact of CMEs is estimated solely based on their arrival velocities, predicted from coronagraph observations. The prediction of solar wind stream velocities, e.g., from coronal hole observations, enables to estimate their impact as well.

First, I estimate the long-term variations of the yearly average  $K_p$  values, which are contributed by solar activity. This is achieved via logarithmic fitting of a yearly sunspot number (SSN) dependency. In order to nowcast the  $K_p$  index from general solar wind conditions, I apply a correlation with the product of the parameters velocity and magnetic field z-component in GSM coordinates ( $vB_z$ ). In order to forecast the  $K_p$  index from estimated CME and stream velocities, I furthermore filter the solar wind data, using flagged CME times from the solar wind structures (SWS) list provided by Richardson & Cane (2012). The solar wind data considered in these analyses consists of 35 years (1981–2016) of high-resolution minutely OMNI data, which is composed of multi-spacecraft intercalibrated in-situ measurements from 1 au. I analyze the  $K_p$  frequency distributions with respect to the depending parameters  $vB_z$  and velocity, derive their mean  $K_p$  per interval and further compile functional dependencies via logarithmic fitting.

The obtained functional relations enable empirical estimations of the mean  $K_p$  impact from measured solar activity, in-situ solar wind, and remotely determined CME and stream velocities.

### 4.1 Introduction

It is known since the early 19th century that variations in the solar wind evoke disturbances in the magnetosphere (Bartels 1962). Especially strong disturbances, called geomagnetic storms, can be provoked by coronal mass ejections (CMEs), which are embedded within the solar wind. The causes of the strongest geomagnetic storms are the compression of the solar wind magnetic field lines within the CME shock front and the enhanced field strengths of magnetic clouds, which are enclosed in CMEs (Bothmer 1993; Bothmer & Schwenn 1995). The consequences of strong geomagnetic disturbances are a threat to sensitive technical systems and exposed humans. Therefore it is important to know when magnetospheric disturbances will occur and how large they will become.

$K_p$  is designed for... (cite?). I use this magnetospheric disturbance index to correlate it with near-Earth solar wind measurements. More detailed information on the  $K_p$  index can be found in Subsection 2.7.6.

current comparable  $K_p$  forecast models take into account... (NOAA/SWPC, CCMC/SWRC, Wing forecasts)  
Elliott et al. (2013): The  $K_p$  index and solar wind speed relationship: Insights for improving space weather forecasts

differences to existing studies:

compare with Newell relation...

With the results presented in this chapter, I elaborate the step from the predicted CME and stream velocities to the forecast of the possible impact strength on the terrestrial magnetosphere. I derive an empirical correlation between the solar wind speed and the geomagnetic  $K_p$  index, in order to obtain the capability to forecast  $K_p$

## 4. Solar wind and CME influence on the magnetosphere

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values solely based on the predicted CME and stream velocities. The derived functional dependencies can be used to nowcast/forecast the  $K_p$  index.

In-situ measurements of solar wind are made almost continuously, e.g., at the first Lagrange point (L1), in front of the magnetosphere. Since 1963 several spacecraft collected more than 50 years of solar wind data. The latest spacecraft, e.g., Wind, ACE and DSCOVR (launched in early 2015), provide real-time solar wind data online<sup>1 2 3</sup>. These solar wind real-time data are used to nowcast various effects on the Earth's magnetosphere, such as the position of the magnetospheric bow shock in front of the Earth, the magnitude of geomagnetic disturbances (e.g.,  $K_p$  index), the positions of the polar auroral ovals, the variation of the total electron content (TEC) of the ionosphere, and the positional error of global navigation satellite systems (GNSS).

The velocity and the direction of CMEs can be determined in their early near-Sun stages via remote tracking with coronagraph white-light observations. Using these parameters as input for CME propagation models, their possible arrival time and arrival velocity at Earth can be derived.

There are efforts to predict the direction and strength of the magnetic field from the flux rope geometry and orientation ([Savani et al. 2015](#)). However, the compressed solar wind magnetic field in the sheath region is of at least similar importance for magnetospheric disturbances ([Huttunen & Koskinen 2004](#)). As the solar wind  $B_z$  is quite random with low autocorrelation time, its prediction is not yet implemented in the current forecast.

[Savani et al. \(2015\)](#): Predicting the magnetic vectors within coronal mass ejections arriving at Earth: 1. Initial architecture

[Savani et al. \(2017\)](#): Predicting the magnetic vectors within coronal mass ejections arriving at Earth: 2. Geomagnetic response

Coronal holes are the origin of the fast solar wind; their area on the solar disk, seen in EUV wavelengths, correlates with the measured velocity of solar wind streams ([Vršnak et al. 2007](#)). This relation is used to predict the Earth-arrival velocity of solar wind streams about 4 days in advance ([Rotter et al. 2012](#)), as is done in real-time within the empirical solar wind forecast (ESWF) at the University of Graz<sup>4</sup>.  
+27 day forecast...

mention ENLIL forecast simulation

The objectives of the analyses performed and found in this chapter are to estimate the magnetospheric impact of solar wind and to predict it for CMEs in particular. In [Section 4.2](#) I determine the magnitudes of the long-term  $K_p$  changes due to solar activity and measure the extent of seasonal variations stemming from the Earth's orbit. In order to nowcast the  $K_p$  index, I quantify the solar wind influence on  $K_p$  by deriving a functional relation with the product of solar wind velocity and magnetic field z-component ([Section 4.3](#)). Finally, for the purpose of enabling  $K_p$  forecasts from remote observations, I estimate the  $K_p$  impact coming from CMEs and solar wind streamers separately by deriving functional dependencies with their velocities ([Section 4.4](#)).

## 4.2 Long-term variations of the $K_p$ index

### 4.2.1 Solar activity influence

The  $K_p$  data is obtained from the GFZ Potsdam<sup>5</sup>, where the index is currently maintained. The data used in this analysis covers the time period 1932–2016, see also [3.3](#). Its frequency distribution shows that the highest frequencies occur around low  $K_p$  values of 1+. Going to higher  $K_p$  values, the frequencies decline asymptotically towards zero (see [Figure 4.1](#)) – a  $K_p$  value of 90 occurred only 29 times in this time interval.

Obviously the general  $K_p$  distribution (seen in [Figure 4.1](#)) averages over solar activity. Solar activity is generally tracked with the international sunspot number (SSN). In the present analyses SSN data from the time

<sup>1</sup>Wind real-time data website: <https://pwg.gsfc.nasa.gov/windnrt/>

<sup>2</sup>ACE real-time data website: <http://www.swpc.noaa.gov/products/ace-real-time-solarwind>

<sup>3</sup>DSCOVR real-time data website: <http://www.swpc.noaa.gov/products/real-time-solarwind>

<sup>4</sup>ESWF website: <http://swe.uni-graz.at/index.php/services/solarwind-forecast>

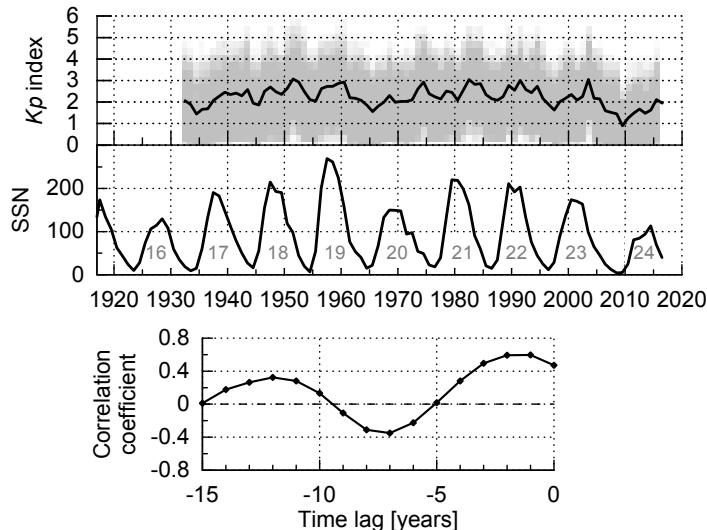
<sup>5</sup>GFZ Potsdam website: <http://www.gfz-potsdam.de/de/kp-index/>



**Figure 4.1**  $K_p$  frequency distribution for the time period 1932–2016. The inset shows a zoomed-in view of the high-value tail. The  $K_p$  data is obtained from the GFZ Potsdam.



**Figure 4.2** Yearly  $K_p$  frequency distributions during the period 1932–2016, sorted and colored by yearly SSN. All distributions are normed to be of equal area. The  $K_p$  data is obtained from the GFZ Potsdam and the yearly SSN data from the SILSO World Data Center.



**Figure 4.3** Yearly  $K_p$  index distributions (shaded area) with their mean values for the time period 1932–2016 and yearly SSN with cycle number for the time period 1917–2016 (top panels). The Pearson correlation coefficients with the yearly SSN are calculated for time lags back to –15 years (bottom panel). The  $K_p$  data is obtained from the GFZ Potsdam and the yearly SSN data from the SILSO World Data Center.

period 1917–2016 is used. The data is obtained from the online catalog<sup>6</sup> provided by the World Data Center – Sunspot Index and Long-term Solar Observations (WDC-SILSO), Solar Influences Data Analysis Center (SIDC), Royal Observatory of Belgium (ROB).

The  $K_p$  frequency distributions' shape varies with different states of solar activity (cite?). This is visible in the yearly distributions, sorted and colored by yearly SSN in Figure 4.2.

The distribution's peak position scales with SSN, that is, a high yearly SSN results in a higher abundance of large  $K_p$  values as well (cite?).

The time series of yearly average  $K_p$  values of the years 1932–2016 shows an imprint of the solar cycles (see the top graphs in Figure 4.3). The  $K_p$  pattern follows the solar cycle minima and maxima as well as the changes in magnitude between solar cycles (cite?). The yearly mean  $K_p$  shifts about 1 unit for both variations separately. As expected, the  $K_p$  index correlation with solar activity shows an 11-year period (cite?) (see bottom graph in Figure 4.3). The highest correlation coefficient of 0.60 is found with a time lag of –1 year, that is, the yearly average  $K_p$  follows the SSN of the previous year.

The yearly mean  $K_p$  index with respect to the 1-year lagged SSN shows a raise with increasing SSN, this is seen in Figure 4.4. I perform a fit in order to obtain an analytical relation for this dependency.  $K_p$  itself is a quasi-logarithmic index, so it is apparent to use a logarithmic fit function:

$$f(x) = a \cdot \ln(x) + b. \quad (4.1)$$

<sup>6</sup>WDC-SIDC website: <http://www.sidc.be/silso/>



**Figure 4.4** Yearly mean  $K_p$  index with respect to 1-year lagged SSN (+) with the weighted logarithmic fit (dashed line). The error bars denote the SSN standard deviation and the relative weight from the yearly data coverage. The shaded area represents the fit error band derived from the estimated standard deviations of the fit parameters. The logarithmic function (Equation 4.1) is used for the weighted fit. The yearly  $K_p$  mean values are calculated from GFZ Potsdam data and the yearly SSN is obtained from the SILSO World Data Center.

The fit parameters are  $a = 0.281(43)$  and  $b = 1.05(19)$  and lead to the relation

$$Kp(ssn) = 0.28 \cdot \ln(ssn) + 1.1 , \quad (4.2)$$

which is plotted in Figure 4.4. This means that for a SSN of 1 the mean  $K_p$  is 1.05(20) and for a SSN of 300 it is 2.65(31). The numbers in parentheses are the errors on the corresponding last digits of the quoted value. They are calculated via error propagation from the estimated standard deviations of the fit parameters.

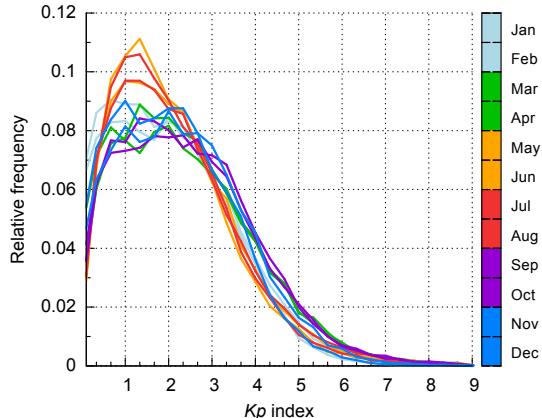
## 4.2.2 Seasonal variations

On top of the yearly variations, seasonal variations exist in the magnetospheric disturbances as well. In the months May–August the  $K_p$  peak frequency is higher than in the remaining months of the year, whereas in March/April and September/October the  $K_p$  values  $>3$  are more abundant (cite?). This is apparent from looking at the monthly  $K_p$  frequency distributions, plotted in Figure 4.5. These  $K_p$  changes arise from seasonal variations of the solar wind parameters at Earth, which stem from Earth’s yearly changes in orbital distance and heliographic latitude. Another seasonal effect stems from the Earth’s rotation axis tilt, which changes the direction of the Earth’s magnetic dipole axis to the Sun over the year (see top panel of Figure 4.6).

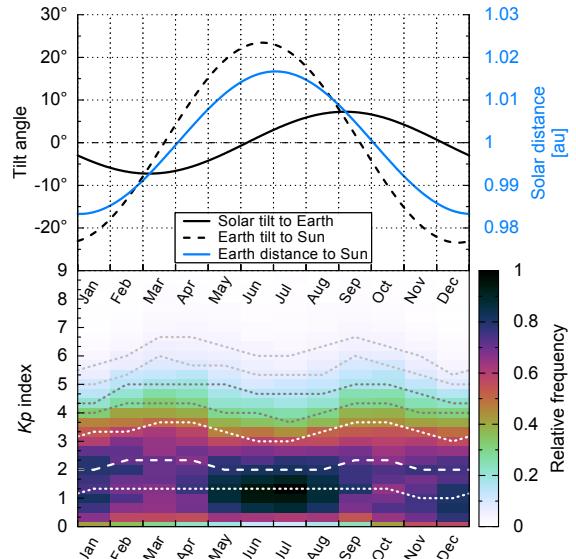
Earth’s distance to the Sun varies over the course of a year by  $\pm 1.67\%$  (see appendix Section A.5). The solar wind parameters scale via power-law dependencies with solar distance, as it is described in the following Chapter 5 and accordingly in Venzmer & Bothmer (2018). For example, the solar wind proton density scales with about  $r^{-2}$ , this leads to yearly variations in density of up to 3.4 %. These yearly solar wind changes have a direct influence on the  $K_p$  index.

The Sun’s rotation axis tilt angle to the ecliptic is  $\pm 7.25^\circ$  and that for Earth is  $\pm 23.44^\circ$ , see also appendix Section A.5. The solar wind magnetic field strength varies with heliographic latitude (cite?). The solar wind influence on the  $K_p$  index depends on its coupling efficiency with the magnetosphere. Furthermore, the rate of magnetic reconnection between solar wind and the Earth’s magnetosphere depends on both fields’ orientation to each other (parallel/antiparallel). Additionally, the tilt of the magnetic dipole axis to the rotation axis – a few degrees for the Sun during cycle minima (extreme during solar maxima...) and about  $10^\circ$  for the Earth – complicates this system even more.

So the  $K_p$  variation effects originate from the seasonal change in the solar tilt, the Earth’s tilt, and the Earth’s distance. Thorough analyses of the seasonal variations were already performed by (Cortie 1912) (more cites!). Thus, I just quantify the bulk magnitude of these effects in order to consider them as relative uncertainties when using the solar activity relation (4.2). Looking at the  $K_p$  frequency distributions by month – seen in the bottom panel of Figure 4.6 – it is apparent that for high  $K_p$  values ( $> 4$ ), there exist yearly frequency maxima at the equinoxes and frequency minima at the solstices, as described by (Cortie 1912). It can be seen from the plotted quantiles that this semiannual variation amounts to  $1/3$   $K_p$  unit at small  $K_p$  values and up to  $4/3$  units at higher  $K_p$  values.



**Figure 4.5** Average monthly  $K_p$  frequency distributions of the time period 1932–2016, colored by month of the year. The  $K_p$  data is obtained from the GFZ Potsdam.



**Figure 4.6** Seasonal changes in the solar tilt angle to Earth, Earth tilt angle to Sun, and Earth distance to Sun – the curves are approximated with trigonometric functions (top panel).  $K_p$  frequency distributions by month for the time period 1932–2016 (bottom panel) with median (white dashed) and quartile values (white dotted). The other dotted lines mark the upper eighth, 16th, 32nd, and 64th parts. The bin size is 1 month and 1/3  $K_p$  unit respectively.

## 4.3 $K_p$ nowcast from in-situ solar wind measurements

### 4.3.1 Coupling function

The coupling between the solar wind and the magnetosphere is governed by reconnection and compression of the magnetic field lines, as described in [Subsection 2.7.1](#).

In this work I settle for [Equation 2.2](#) ( $vB_z$ ) as the coupling function.  
For which reason do I choose this coupling function over others?

The solar wind electric field  $E_y$  is the product of the proton velocity  $v$  and the magnetic field z-component  $B_z$ :

$$?checkvectors E_y = -v_x \times B_z. \quad (4.3)$$

If not specified otherwise,  $B_z$  is always meant to be in GSM coordinates hereafter.

The solar wind velocity sticks mostly to its radial flow direction, that is, it rarely deviates up to  $0.0^\circ$  (correct and cite...). Thus, the absolute flow speed can be used instead of the vector component.

- argue for  $vB_z$ :
- 3min( $vB_z$ ) performs in rank correlation slightly better than the sophisticated Newell formula. really?
  - simple to calculate
  - ...

It is also known that the solar wind velocity itself already correlates strongly with the  $K_p$  index.

I use this velocity relationship for obtaining  $K_p$  proxies from CME and solar wind stream data.

### 4.3.2 Data correlation

The  $K_p$  time series started in 1932 when there existed no spacecraft to measure solar wind in situ. Thus, the maximal surveyed time range is defined by the available in-situ solar wind data.

Savani2017:

“Although  $K_p$  is defined for 3 h periods, predicted variations in the IMF are made on shorter time scales and can influence the level of geomagnetic activity (e.g., geomagnetically induced currents). Hence, averaging the IMF over intervals that match those of  $K_p$  may suppress features that are important drivers of geomagnetic activity. Nevertheless, for the purpose of this paper, we choose to average the  $K_p$  estimate from the field vectors predicted by the BSS model over the same 3 h intervals as the  $K_p$  values, as this currently represents the service provided by NOAA/SWPC.”

The  $K_p$  index represents maximal variations within 3-hour time intervals (see [Subsection 2.7.6](#)). Any solar wind parameter that will be correlated with  $K_p$  also has to have the same time resolution. In addition to adapting the time resolution, it has to be considered by which means this should be done. Simple 3-hour average values are expected to have a weaker correlation than the solar wind parameter’s 3-hourly maximal variation. The 3-hour maximal variations should obviously be stronger when using high resolution data. Thus, to be able to correlate  $K_p$  with solar wind data, high resolution data (that is, much shorter than 3 hours) are needed to determine the maximal solar wind variations within each 3-hour interval.

The OMNI data set collects the longest continuous solar wind measurements at 1 au. The longest time coverage has the hourly OMNI data set (since 1963), however, I choose to use the minutely OMNI data with the time range 1981–2016, to benefit from the stronger correlation (see [Figs?](#)).

The  $K_p-vB_z$  Pearson correlation coefficients for 3-hour mean and minimum data are plotted in [Figure 4.7](#). The largest correlation is found for the 3-hour minimum data without time shift. It is a negative correlation



**Figure 4.7**  $K_p-vB_z$  correlation coefficients for different time shifts up to  $\pm 24$  hours. The minutely OMNI data from 1981–2016 is processed with mean (black) and minimum (red) 3-hour averaging.

whose coefficient is  $-0.72$ . The  $K_p-vB_z$  frequency distribution is asymmetrically shifted to negative values when using 3-hour minimum values, this is visible in [Figure 4.8](#).



**Figure 4.8** Frequency distributions of  $vB_z$  for 3-hour mean (black) and minimum (red) minutely OMNI data from 1981–2016.

### 4.3.3 Functional dependency for solar wind

The frequency distribution in  $K_p-vB_z$  space is shaped like a candle flame, inclined to negative values by a light breeze, see top panel in [Figure 4.9](#). In order to determine a functional dependency, I focus on the relative frequencies per  $vB_z$ -interval and their mean  $K_p$  values, which are plotted in the bottom panel of [Figure 4.9](#). The mean absolute deviation (MAD) of the mean has a mean size of  $0.7 K_p$  units. This probability distribution is asymmetrically V-shaped around zero, having a larger and steeper negative arm than positive arm. The asymmetry also exists for 3-hour mean data, thus this effect is not a result of the data reducing method (3-hour



**Figure 4.9**  $K_p$  versus  $vB_z$  frequency distribution (top) and its relative distribution (bottom) with the mean  $K_p$  values (solid line) and their mean absolute deviation (dotted lines). It is 3-hour minimum data from the minutely OMNI data set (1981–2016). The bin size is  $500 \text{ km s}^{-1} \text{ nT}$  and  $1/3 K_p$  unit respectively.

minimum) (fig...?). Rather the steeper negative arm is a consequence of the half-wave rectifier coupling of the solar wind magnetic field direction to the magnetosphere, as described in Subsection 2.7.1.

stress that it is an empirical fit...

Since the  $K_p$  index has a quasi-logarithmic scaling (see Subsection 2.7.6), a logarithmic function is the obvious choice as a fit function. Furthermore, the depending argument consists of a product of two solar wind parameters which individually scale logarithmically with  $K_p$ . These reasons are why I use the logarithm of a parabola for the fitting approach:

$$f(x) = \ln(x^2) . \quad (4.4)$$

I also introduce a horizontal shifting parameter  $x'$  because the distribution's center is slightly offset. To be able to replicate the asymmetry in both arms, I further split the fit function into a negative and a positive part:

$$f(x) = \begin{cases} f_-(x) & \text{for } x < 0, \\ f_+(x) & \text{for } x \geq 0. \end{cases} \quad (4.5)$$

This way, both arms can be scaled individually with scaling factors for the negative and positive parts  $a_-$  and  $a_+$ . The resulting logarithmic fit function parts are

$$f_-(x) = a_- \cdot \ln((x + x')^2 + b) + y' , \quad (4.6)$$

$$f_+(x) = a_+ \cdot (f_-(x) - f_(-x')) + f_(-x') , \quad (4.7)$$

with the vertical shifting parameter  $y'$  and the depth parameter  $b$ . The resulting fit curve is plotted in Figure 4.10 with the fit coefficients  $a_- = 1.258(19)$ ,  $x' = 163(20)$ ,  $b = 1.416(68) \times 10^6$ ,  $y' = -17.04(33)$ , and  $a_+ = 0.467(20)$  for units of  $[\text{km s}^{-1} \text{ nT}]$ . The mean MAD is about 0.7  $K_p$  units. Thus, the solar wind dependency relation condenses to:

$$Kp_-(vB_z) = 1.258 \cdot \ln((vB_z + 163)^2 + 1.416 \times 10^6) - 17.04 , \quad (4.8)$$

$$Kp_+(vB_z) = 0.467 \cdot (Kp_-(vB_z) - Kp_-(-163)) + Kp_-(-163) . \quad (4.9)$$

This relation can be used together with real-time in-situ measurements from spacecraft located at L1 (see Section XX rt sources) to nowcast the actual  $K_p$  index.



**Figure 4.10** Mean  $K_p$  values (+) and MAD values (dotted lines) per  $vB_z$  interval. The error bars represent the relative data count. The logarithmic fit (dashed line) is plotted with a mean MAD band (shaded area). The splitted function (4.5) is used for the weighted fit.

In order to demonstrate this relation with actual in-situ data, I calculate the  $K_p$  estimate for the example CME from Figure 2.21 and include the 3-hour minimum value of  $vB_z$  in the plot, see Figure 4.11. It can be seen that in this period the  $K_p$  estimate traces the actual  $K_p$  index pretty well within the mean MAD band. However, deviations are found at the times of the initial shock, and the start and middle of the MC.

## 4.4 $K_p$ forecast from remote CME and stream observations

It makes sense to separately look at CMEs and solar wind streams -> see event comparison fig.  
same velocity -> different  $K_p$  effect

CMEs are already sighted raising from their source regions on the solar surface. Some CME properties can be estimated from remote coronagraph observations and modeled to Earth, such as its propagation direction, its arrival time, and arrival velocity (cites?). Thus, early observations enable a heads-up time depending on the CME's propagation speed to Earth. This travel duration can be more than 4 days for slow events with average solar wind speeds, about 40 hours for fast events with speeds of  $1000 \text{ km s}^{-1}$ , and down to 20 hours and even below for the rare extreme cases, e.g., about 21 hours for the event on 23 July 2012 (Russell et al. 2013; Temmer & Nitta 2015) and about 19 hours for the event observed by Carrington (1859) on 1 September 1859. The theoretical minimal heads-up time for CMEs is estimated to be half a day (Gopalswamy et al. 2005).

To make use of the heads-up time for CMEs, I neglect its magnetic field part, which is difficult to determine from remote observations and simplify the coupling relation (2.2) from before. Therefore, the only coupling parameter remaining is the solar wind velocity.

Savani2017:

*"Prior evidence has suggested the reason be due to southerly IMF strength being the most dominant solar wind parameter for driving geomagnetic activity at Earth. The strong correlation of solar wind speed to geomagnetic activity is evident during CME conditions but is only considered the most dominant parameter during high-speed streams [Holappa et al., 2014]."*

### 4.4.1 Velocity forecast methods

CME velocity estimation  
methods and modeling...?  
GCS, CAD modeling -> propagation direction and apex height-time profile -> acceleration and velocity kinematics...  
-> example event CME?  
see Basics...

Solar wind stream velocity estimation  
CH analysis (Graz)



**Figure 4.11** Solar wind measurements, official  $K_p$  index, and estimated  $K_p$  for the time period 26 June to 2 July 2013. The solar wind parameters are the magnetic field strength, its z-component in GSM coordinates, and the velocity. I also plotted the product  $vB_z$  with its 3-hour minimum for illustration. The solar wind data are from the minutely OMNI data set. The official  $K_p$  index is obtained from the GFZ Potsdam. The  $K_p$  estimate is derived via Equations 4.8 and 4.8. The orange band indicates the mean MAD.

L4 shift (Graz)  
 WSA-ENLIL MHD simulation  
 see Basics...

#### 4.4.2 Solar Wind Structures list

For the following analysis, I use the list of solar wind structures (SWS) created and updated by Richardson et al. (2000) and Richardson & Cane (2012), who characterized the near-Earth solar wind structures since 1963. All periods related to ICMEs in the OMNI solar wind data set were identified and flagged.

list identifying criteria... look into their paper!

The SWS list for 1963–2016 was kindly provided by Ian Richardson (private communication). into acknowledgments...

SWS list for 1963–2015 by (Richardson et al. 2000; Richardson & Cane 2012) is available via registration at CEDARweb<sup>7</sup>.

The CME fraction in the OMNI data is 15.8 % during the time period 1981–2016 (this accumulates to 5.53 years). This percentage is an average value. It varies heavily with the solar activity cycle and with the appearance of individual active regions on the solar surface. In the following, I refer to the remaining part without CMEs as “solar wind streams”, because it is composed entirely from a mixture of slow wind flows, fast wind streams, and their interaction regions.

<sup>7</sup>CEDARweb website for Solar Wind Structures: [http://cedarweb.vsp.ucar.edu/wiki/index.php/Tools\\_and\\_Models:Solar\\_Wind\\_Structures](http://cedarweb.vsp.ucar.edu/wiki/index.php/Tools_and_Models:Solar_Wind_Structures)

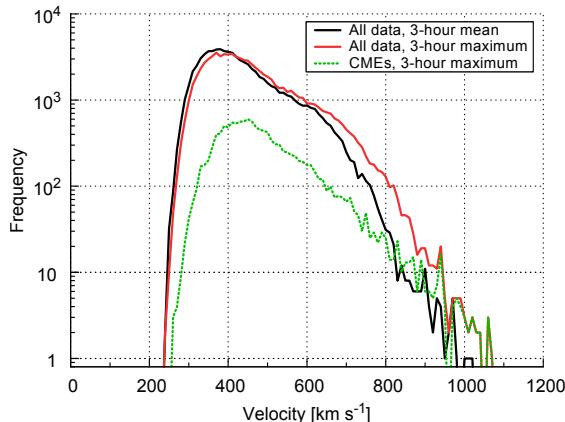
**Table 4.1** Time lags with the highest correlation coefficients for the  $K_p$ -velocity relation for the whole OMNI data, for stream data, and for CME data. The values are based on the 3-hour maximum of the minutely high resolution OMNI data from the time period 1981–2016.

Data	Time lag [hours]	Correlation coefficient
All data	6	0.622
Streams	9	0.661
CMEs	0	0.511

from OMNI data only?  
permission received.

#### 4.4.3 Data correlation

Again, I calculate 3-hour extreme values using the minutely OMNI data to profit from higher correlation coefficients, as done before for the data processing of the  $vB_z$  analysis in Subsection 4.3.2. The velocity only has positive values, thus its extreme values are 3-hour maximum values. The comparison between the 3-hour maximum and the 3-hour mean frequency distributions shows that their mean position shifts from a velocity of 405 to 425 km s $^{-1}$ , see Figure 4.12. The CME part and solar wind stream part of the data can be examined



**Figure 4.12** Solar wind velocity frequency distributions for 3-hour mean (black), maximum (red), and maximum of the CME part (green). The minutely OMNI data from the period 1981–2016 is used.



**Figure 4.13**  $K_p$ -velocity correlation coefficients for time shifts in the range –24–48 hours. The correlations are plotted for the whole solar wind data (solid), for solar wind streams without CMEs (dashed), and for CMEs only (dotted). The data used is the 3-hour maximum of the minutely high resolution OMNI data from the period 1981–2016.

separately, filtering the CME related periods using the SWS list. Their frequency distributions show that the CME share is rising in faster solar wind until eventually in the region above about 900 km s $^{-1}$  there only exist CMEs, see Figure 4.12.

The CME part of the data is correlated with the  $K_p$  index independently from the remaining solar wind streams, see Figure 4.13. The correlation for CME related data is lower than that for the regular solar wind (all data). Its maximal correlation coefficient with a value of 0.51 is without time shift, see Table 4.1. The solar wind streams show a higher correlation with  $K_p$  and the maximal coefficient of 0.66 has a positive time shift of 9 hours, that is, the  $K_p$  index forecasts the velocity of solar wind streams 9 hours in advance.

The positive time shift can be explained with the occurrence of interaction regions followed by high speed streams (HSS). When a slow solar wind stream is followed by a fast one, the compression at their interface leads to enhanced solar wind densities and magnetic field strengths. The peak velocity of a HSS naturally appears after the interaction region. Therefore the  $K_p$ -impact from the enhanced magnetic field is correlated with the higher velocity of the HSS, yielding the observed positive time shift.

In fact, Machol et al. (2013) even proposed a linear function of the  $K_p$  index as a best proxy for corrupted real-time velocity measurements made by the Advanced Composition Explorer (ACE) spacecraft.



**Figure 4.14** *Kp*-velocity distributions for all solar wind data, for solar wind streams and for CMEs. The data used is the 3-hour maximum of the minutely high resolution OMNI data from the time period 1981–2016. The SWS list from Richardson & Cane (2012) is used for the separation between CME and stream data. The bin size is 10 km s<sup>-1</sup> and 1/3 *Kp* unit respectively.

#### 4.4.4 Functional dependency for CME velocity

The general *Kp*-velocity dependency in the solar wind is apparent from the tilt of its distribution, see top panel of Figure 4.14. The distribution is inclined to positive values but very broad. The comparison with the filtered data shows that *Kp* values  $>6$  and velocities  $>850$  km s<sup>-1</sup> are almost always associated with CME related periods, see middle and bottom panel of Figure 4.14.

In order to find a functional relation for the mean *Kp* value, I look at the relative frequencies per velocity interval, which are plotted in the bottom panel of Figure 4.15. The mean *Kp* value seems to scale almost linear with the solar wind velocity. The MAD of the mean has a mean size of about 1.1 *Kp* units.

Again, as the *Kp* index has a quasi-logarithmic scaling, a logarithmic function is the obvious choice for the fitting process. Thus, the logarithmic function

$$f(x) = a \cdot \ln(x + x') + y' \quad (4.10)$$

is used for the fit, with the scaling factor  $a$ , the location parameter  $x'$ , and the vertical shifting parameter  $y'$ . The resulting fit is plotted in Figure 4.16 and its parameters are  $a = 10.6(34)$ ,  $x' = 8.1(43) \times 10^2$ , and  $y' = -73(28)$ , with the velocity in units of [km s<sup>-1</sup>]. The MAD is about 1.1 *Kp* units. This leads to the CME dependency function

$$Kp(v) = 10.6 \cdot \ln(v + 810) - 73, \quad (4.11)$$

which can be used to forecast the *Kp* index from the estimated CME arrival velocity.

add HSS example plot...

It becomes obvious how important the influence of the magnetic field z-component is, when looking at the same example CME from Figure 4.11. I plotted its in-situ magnetic field strength and velocity in Figure 4.17

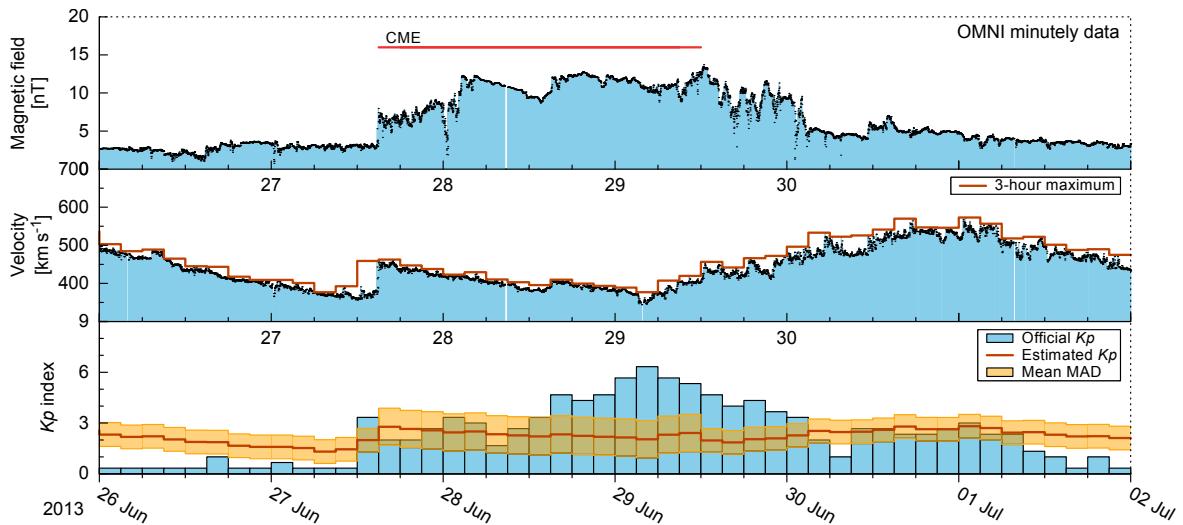


**Figure 4.15** CME part of the  $K_p$ -velocity distribution (same as third panel of Figure 4.14) and its relative distribution per velocity interval with the mean  $K_p$  values (solid line) and their mean absolute deviation (dotted lines). The bin size is  $10 \text{ km s}^{-1}$  and  $1/3 K_p$  unit respectively.



**Figure 4.16** Mean  $K_p$  values (+) and MAD values (dotted lines) per velocity interval for the CME part of the data. The error bars represent the relative data count. The logarithmic fit (dashed line) is plotted with a mean MAD band (shaded area). The function (4.10) is used for the weighted fit.

together with the official  $Kp$  index and the derived  $Kp$  estimate. The latter do not coincide well during this



**Figure 4.17** Solar wind measurements, official  $Kp$  index, and estimated  $Kp$  for the time period 26 June to 2 July 2013. The solar wind parameters are the magnetic field strength and the velocity. I also plotted the velocity's 3-hour maximum for illustration. The solar wind data are from the minutely OMNI data set. The official  $Kp$  index is obtained from the GFZ Potsdam. I plotted the  $Kp$  estimate depending on whether the period is flagged as a CME (red line in top panel) or not, that is, via Equation 4.11 or Equation 4.12. The orange band indicates the mean MAD.

period. The  $Kp$  estimate performs okay during the initial sheath region and around the peak of the trailing HSS. The rarefaction regions of declining velocity at the start and end of the considered interval are overestimated, whereas the MC and the trailing compression region are underestimated.

#### 4.4.5 Functional dependency for stream velocity

The procedure in this section is similar to that in the previous section. The correlation coefficient is higher for solar wind stream velocities when the data is shifted by 9 hours, see Figure 4.13. I use the shifted data and look at the relative frequencies per velocity interval in order to find a functional dependency for the mean  $Kp$  value, see bottom panel of Figure 4.18. Again, the mean  $Kp$  value scales almost linear with the velocity. The MAD of the mean has a mean size of about 0.7  $Kp$  units.

Again, I use the logarithmic function (4.10) for the fitting process. The resulting fit is plotted in Figure 4.19 and the fit parameters are  $a = 5.88(38)$ ,  $x' = 2.99(49) \times 10^2$ , and  $y' = -3.70(29) \times 10^1$ , with the velocity in units of [ $\text{km s}^{-1}$ ]. The MAD is about 0.7  $Kp$  units. This leads to the solar wind stream dependency function

$$Kp(v) = 5.88 \cdot \ln(v + 299) - 37.0, \quad (4.12)$$

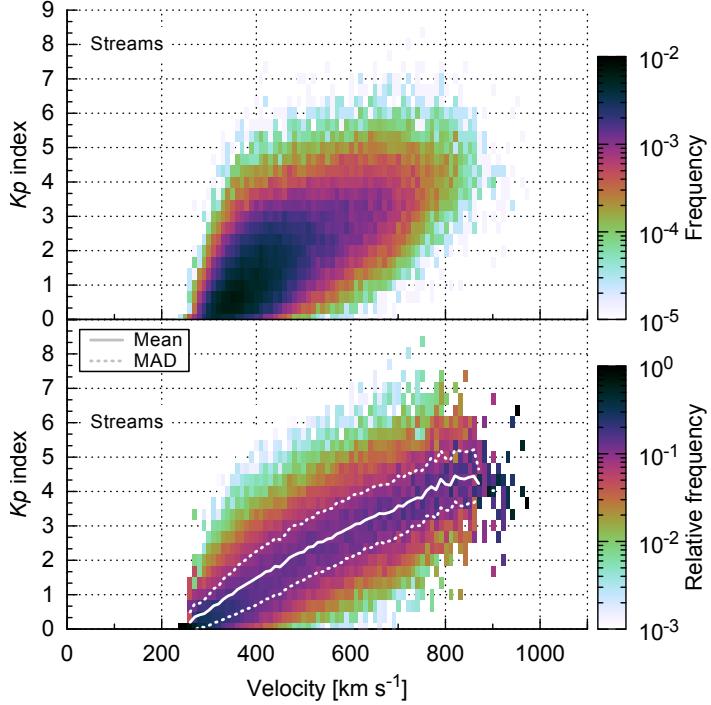
which can be used to forecast the  $Kp$  index from the estimated stream velocity, e.g., obtained from remote coronal hole analyses.

## 4.5 Discussion and summary

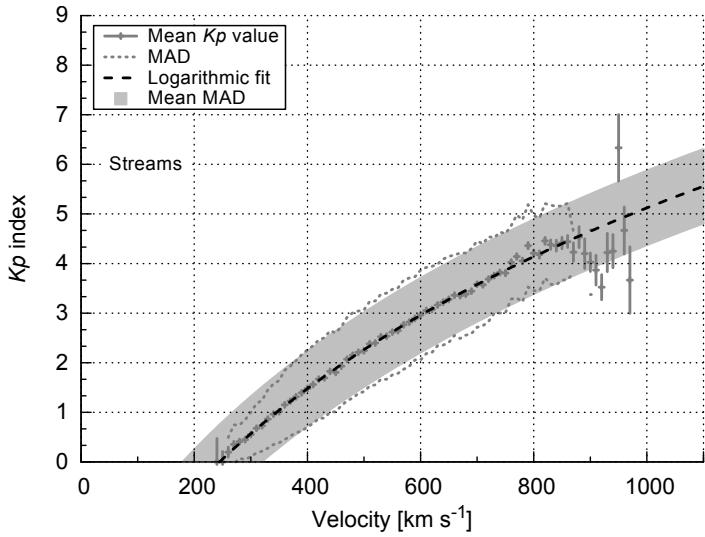
The following results/relations are obtained from the analyses:

- solar activity:  $Kp$ -SSN relation with an error of about 1/3  $Kp$  unit
- seasonal variations of up to 4/3  $Kp$  units
- solar wind nowcast:  $Kp$ - $vB_z$  relation (average and worst case)
- CME forecast:  $Kp$ -velocity relation (average and worst case)
- stream forecast:  $Kp$ -velocity relation (average and worst case)

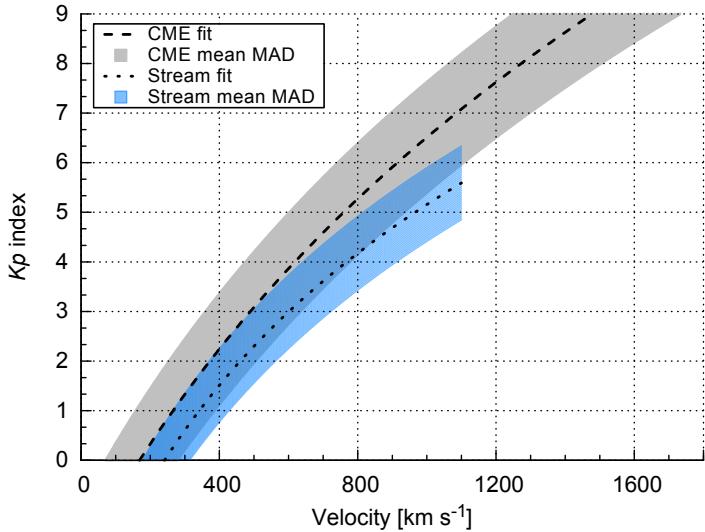
CMEs can be faster than the maximal velocities included in the OMNI data. Instrumental effects (specify...) lead to data gaps during these periods. Yet, CME speeds of up to XX km/s at 1 au were observed (cite?). According to Equation 4.11, a CME speed of 2000 km s<sup>-1</sup> at L1 would lead on average to a theoretical  $Kp$  of



**Figure 4.18** Stream part of the  $K_p$ -velocity distribution (similar to second panel of Figure 4.14, but with the data shifted by 9-hours) and its relative distribution per velocity interval with the mean  $K_p$  values (solid line) and their mean absolute deviation (dotted lines). The bin size is  $10 \text{ km s}^{-1}$  and  $1/3 K_p$  unit respectively.



**Figure 4.19** Mean  $K_p$  values (+) and MAD values (dotted lines) per velocity interval for the stream part of the data, shifted by 9 hours. The error bars represent the relative data count. The logarithmic fit (dashed line) is plotted with a mean MAD band (shaded area). The function (4.10) is used for the weighted fit.



**Figure 4.20** Logarithmic fit curves for CMEs (dashed line) and streams (dotted line) with their corresponding mean MAD bands (shaded areas).

12.2, however, the  $K_p$  scale is capped at 90. A  $K_p$  of 90 is reached already at a velocity of  $1489 \text{ km s}^{-1}$ , see Figure 4.20.

The maximal velocity of regular solar wind is limited by the coronal temperatures (Parker 1958) and is observed to be around  $900 \text{ km s}^{-1}$ . That is why streams on average provoke  $K_p$  values below 50.

comparison:  
 $K_p$ -velocity correlation  
similar to Elliott et al. (2013); different data time period, resolution and averaging method (3-hour maximum of 1 min data)

It is obvious that the derived relations are not suited to be used inversely, i.e., to deduce solar wind properties from existing  $K_p$  data.

## 4.6 Applications and outlook

'scientists develop an advanced prototype space weather warning system to ensure the operation of telecommunication and navigation systems on Earth to the threat of solar storms' from AFFECTS

Prototype/precursor relations are integrated into applications developed within the Advanced Forecast For Ensuring Communications Through Space (AFFECTS) project which ran from 2011 to 2013. The following services, accessible via the AFFECTS website<sup>8</sup>, contain results from this  $K_p$  study:

- Real-time plot: [Solar Wind and  \$K\_p\$  forecast plot](#) – DSCOVR real-time solar wind and  $K_p$  forecast plot. The  $K_p$  is estimated from the solar wind relation.
- RSS feed: [L1 Kp Alert](#) – threshold based RSS feed that gets triggered when a specified  $K_p$  is reached.
- RSS feed: [L1 GNSS Alert](#) – threshold based RSS feed that gets triggered when a specified GNSS error is reached... The value is derived from the  $K_p$  estimate.
- RSS feed: [L1 Aurora Alert](#) – threshold based RSS feed that gets triggered when a specified auroral location is reached... The value is derived from the  $K_p$  estimate.

note: repair RSS feeds before submitting!

Further applications of resulting  $K_p$ -relations:

- CME  $K_p$  impact (part of UGOE DDC)
- iPhone app L1 Alerts (Solar wind latest 2-hour extreme values and derived forecast values): <http://www.affects-fp7.eu/app-services/L1-Alerts/dataL1Alerts.txt>, <https://itunes.apple.com/au/app/affects/id893579846>
- Android app L1 Alerts... <https://play.google.com/store/apps/details?id=com.afects.forecasts>
- has SW-Display  $K_p$  forecast??

outlook:

Separate structures such as CIRs and HCSs for their  $K_p$ -impact separately...

All provided web links in this work were existent in [date].

CME  $K_p$  impact as part of DDC

- $K_p$  nowcast with L1 solar wind measurements (L1 alerts, disseminated as RSS feeds; integrated in smartphone app and space weather display)
- Forecast of the possible CME impact on the Earth's magnetosphere ( $K_p$  index) from the predicted CME arrival velocity (integrated in UGOE CME forecast chain (aka DDC))

## 4.7 notes...

What kind of solar wind structures create the individual regions in this distribution? (B-V- $K_p$  circle plot)  
What is their individual contribution to the  $K_p$  ranges (e.g. high  $K_p$ : CMEs 70% and CIRs 30%)?

<sup>8</sup>AFFECTS website: <http://www.affects-fp7.eu/services/>

#### 4. Solar wind and CME influence on the magnetosphere

ACE solar wind time series and event list (ACE OPTIMAP “Zeitreihe”-events)

How can the impact field strength of CMEs be forecasted (V->B correlation for CMEs)?

Internal solar wind correlations: B-V correlation

ACE MAGSWE 64 s data -> yearly overlay plot

rt data errors/gaps... vs science data (see paper Kp as V replacement)

DSCOVR as replacement was launched on 11 Februar 2015. It is NOAA's SWPC real-time solar wind prime source since 27 July 2016.<sup>9</sup>

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<sup>9</sup><http://www.swpc.noaa.gov/products/real-time-solarwind>

# 5 Empirical solar wind model for the inner heliosphere

## 5.1 Introduction

The analyses in the previous chapter are focused on the solar wind's influence on the magnetosphere – this chapter changes the main focus to the solar wind upstream of the magnetosphere, in a first step down to a solar distance of 0.3 au and then even further down to the region around  $10 R_{\odot}$ , close to the solar wind's origin near the Sun. The solar wind's evolution on its way to 1 au from the near-Sun region is modeled with the goal of predicting the solar wind environment for the upcoming PSP mission.

context of motivation in thesis and paper introduction and context of results in thesis

This chapter is constructed as follows, the PSP mission, its scientific goals, and the spacecraft are described, I introduce the analyses done in the publication [Venzmer & Bothmer \(2018\)](#), which is a major part of this chapter. An alternative to the magnetic field solar distance dependency is derived, which changes the power law from our article to support a Parker magnetic field geometry. An outlook is given, describing different effects that further could be analysed, quantified, and implemented into the solar wind model, such as flux conservation, seasonal and latitudinal effects, and the distance behavior of solar wind structures.

The major part of this chapter is published under the title “Solar-wind predictions for the Parker Solar Probe orbit” in [Venzmer & Bothmer \(2018\)](#), which is referred to as ‘the paper’ in this work. The article, published online on 20 March 2018 in *Astronomy and Astrophysics* (A&A), is included following this chapter, denoted as [Chapter 6](#), copyright to ESO and reproduced with permission.

The analyses described in the paper were entirely done by me, as well as the tables, figures and equations. My coauthor Volker Bothmer contributed to the text and the anonymous referee helped clarifying a few aspects. The text was further improved by the A&A language editor Joshua Neve.

## 5.2 Parker Solar Probe mission

motivation written in [McComas et al. \(2007\)](#)

mission goals (see paper + [Fox et al. \(2015\)](#))

A photo of PSP during testing is shown in [Figure 5.1](#) and an illustration of PSP flying near the Sun is shown in [Figure 5.2](#).

orbit figure? ref. to distance figure?

WISPR instrument

CGAUSS project – context of this work in project objectives

Sun angular diameter comparison, see presi S<sup>3</sup>

## 5.3 Paper content

We obtain lognormal representations of the frequency distributions' shapes of the four key solar wind parameters magnetic field strength, proton velocity, density and temperature. We derive analytical relations for the parameters' solar activity dependencies and for their solar distance scaling. An empirical solar wind model is built from the combination of the obtained frequency distributions, SSN dependence relations and solar distance dependence functions, representing the solar wind's solar activity and distance behavior. This empirical model is fed with SSN predictions and is extrapolated to the orbit of PSP. We estimate solar wind median values during PSP's first perihelion and model the values for PSP's closest perihelia.



**Figure 5.1** PSP in the Acoustic Test Chamber at NASA’s Goddard Space Flight Center in November 2017. Credit: [NASA/Johns Hopkins APL/Ed Whitman](#), 2017.



**Figure 5.2** Artist’s concept of PSP near the Sun. Credit: [NASA/Johns Hopkins APL](#), 2015.

In order to derive the solar wind environment for the PSP orbit, finally the general solar wind model derived in the previous sections will be extrapolated to the PSP orbit, taking into account predictions of the SSN.

## 5.4 Alternative magnetic field solar distance dependency

note: comparison plot between models in range 0–20 au...

In our article, we scaled the magnetic field strength via a power-law function and obtained a solar-distance dependency proportional to  $r^{-1.662}$ . We also noted there that the model’s near-Sun field magnitude, extrapolated to PSP’s closest perihelion, will be lower than the actual values to be found. However, it is a simplified approach – in the following I present a procedure, leading to an improved distance dependency, which is intended to account for a Parker IMF geometry, respecting the different power-law scalings of the contributions of the individual field vector components. This new distance dependency, extrapolated to the PSP orbit and combined with the solar activity dependency described in the paper, yields 8 % higher values at the first perihelion and 32 % higher values at the 22nd perihelion.

A possibly more accurate distance dependency for the model is derived in the following.

The coronal magnetic field near the Sun rotates rigidly, holding on to the coronal plasma. At the solar wind source surface (at around  $2.5 R_{\odot}$ ), where the thermal plasma pressure overcomes the magnetic pressure, the magnetic field gets radially transported outwards, maintaining only its radial vector component  $\mathbf{B}_r$ . From there on, the solar rotation begins to shear the magnetic field, building up a longitudinal component  $\mathbf{B}_{\phi}$  as well. The solar wind magnetic field model formulated by [Parker \(1958\)](#) has the following components in spherical coordinates ( $\theta$  is the colatitude):

$$\mathbf{B}_r(r) = B_0 \left( \frac{r_0}{r} \right)^2 \cdot \mathbf{e}_r, \quad (5.1)$$

$$\mathbf{B}_{\phi}(r) = -B_0 \left( \frac{r_0}{r} \right)^2 \cdot \frac{\omega r \sin \theta}{v_{sw}} \cdot \mathbf{e}_{\phi}, \quad (5.2)$$

$$\mathbf{B}_{\theta}(r) = 0 \cdot \mathbf{e}_{\theta}. \quad (5.3)$$

$B_0$  represents the radial field component at the solar distance  $r_0$ . The solar surface equatorial angular rotation rate  $\omega$  and the solar wind velocity  $v_{sw}$  are involved as well. From these equations it can be seen that  $\mathbf{B}_r$  scales

with increasing solar distance with  $r^{-2}$ , as it is expected for a spherical outflow, and  $\mathbf{B}_\phi$  scales with  $r^{-1}$  as expected for a two-dimensional point source. In-situ observations support the component's scaling according to these theoretical exponents, although in slow solar wind the scaling for  $\mathbf{B}_\phi$  deviates somewhat from theory (Mariani et al. 1978).

From the magnetic field's absolute value

$$B(r) = \sqrt{\mathbf{B}_r^2 + \mathbf{B}_\phi^2 + \mathbf{B}_\theta^2}, \quad (5.4)$$

$$B(r) = \sqrt{\left(B_0 \left(\frac{r_0}{r}\right)^2\right)^2 + \left(-B_0 \left(\frac{r_0}{r}\right)^2 \cdot \frac{\omega r \sin \theta}{v_{sw}}\right)^2}, \quad (5.5)$$

it is apparent that the magnetic field strength does not scale with a simple power law. If the solar dipole axis tilt to the solar rotation axis is neglected and the radial field component  $B_0$  is set to be in the equatorial plane at  $r_0 = 1$  au, then  $\theta = 90^\circ$  and therefore

$$B(r) = B_0 \cdot \sqrt{r^{-4} + \left(\frac{\omega r}{v_{sw}}\right)^2 r^{-4}}, \quad (5.6)$$

with the distance  $r$  in units of [au].

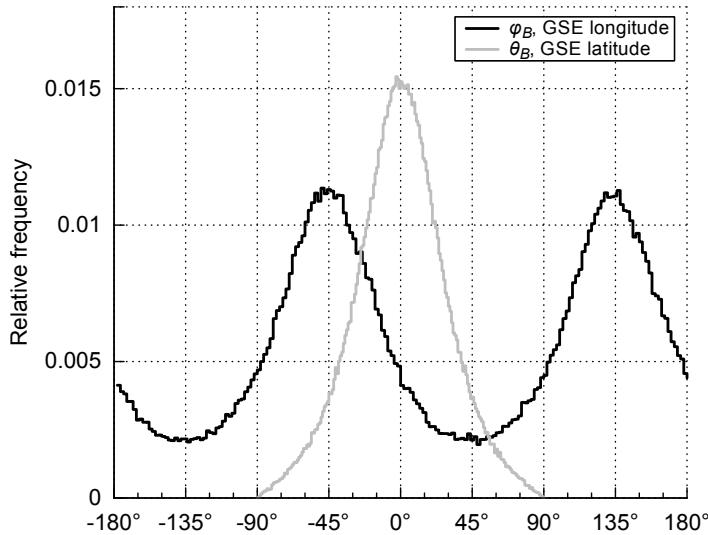
The magnetic field angle in the solar equatorial plane is distance dependent and is calculated via

$$\phi_B(r) = \arctan\left(-\frac{\omega \cdot r}{v_{sw}}\right). \quad (5.7)$$

The angle becomes  $\phi_B(1 \text{ au}) = -46.23^\circ$  when using the Sun's equatorial angular velocity  $\omega_{eq} = 14.37^\circ \text{ d}^{-1}$  (for more details on solar rotation see appendix Section A.1) and the median solar wind velocity  $v_{sw} = 416 \text{ km s}^{-1}$  from our paper.

During solar cycle minimum, the solar wind in the equatorial plane outside the source surface originates from polar regions at about  $\pm 60^\circ$  heliolatitude. This leads to a slightly underwound Parker spiral, because of the slower rotation rate at higher latitudes (Banaszkiewicz et al. 1998). Using the differential rotation rate  $\omega(\pm 60^\circ) = 13.69^\circ \text{ d}^{-1}$  (see appendix Equation A.1), the field angle becomes  $\phi_B(1 \text{ au}) = -44.84^\circ$ .

As the solar rotation axis tilt to the ecliptic normal only has a maximum angle of  $7.2^\circ$ , I consider its influence on  $\phi_B$  minor, and choose to neglect it in this calculation. Thus from theory, in the ecliptic at a solar distance of 1 au, the Parker spiral's observed magnetic field angle  $\phi_B$  is centered most of the time at  $-45^\circ$  or for the opposite field polarity at  $135^\circ$ . Indeed, this is confirmed by the angle's frequency distribution, apparent in measurements made at 1 au, such as the OMNI data of the time period 1963–2016 plotted in Figure 5.3.



**Figure 5.3** Frequency distributions of the magnetic field angles  $\phi_B$  and  $\theta_B$  in GSE coordinates. The frequencies are based on the hourly OMNI data during the period 1963–2016.

For  $\phi_B$  being in average centered around these two directions, both vector components  $\mathbf{B}_r$  and  $\mathbf{B}_\phi$  have to be of equal amplitude. It can be seen from Equations 5.4 and 5.6 that this leads to

$$B(1 \text{ au}) = B_0 \cdot \sqrt{2} \quad \text{with} \quad \left(\frac{\omega \cdot 1 \text{ au}}{v_{sw}}\right) = 1. \quad (5.8)$$

## 5. Empirical solar wind model for the inner heliosphere

**Table 5.1** Fit coefficients for the curve fits (Equation 5.9) and for the distribution fit (using the lognormal function of Equation (4) from the paper), derived from the combined Helios 1 and 2 data. The numbers in parentheses are the errors on the corresponding last digits of the quoted value. The crossing distance indicates where the median and mean fits intersect each other. This distance is calculated numerically and its corresponding error only represents an estimated value.

Fit	Factor $a$	Exponents		Crossing distance [au]
		$e_1$	$e_2$	
Curve	Median	4.04(13)	-1.852(25)	-0.97(23)
	Mean	4.47(12)	-1.740(18)	-0.99(23)
Distribution	Median	3.833(27)	-1.858(42)	-1.32(12)
	Mean	4.081(29)		-

However this way,  $\omega$  and  $v_{sw}$  are treated as time independent constants.

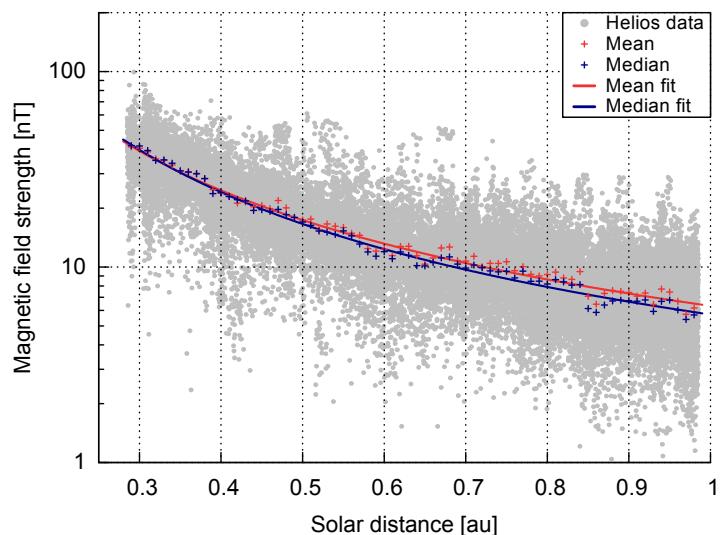
Considering the longitudinal field component, the magnetic field strength's distance dependency is derived in a similar way to that in the paper, see paper [Section 6.4](#). Also, the same Helios data set is used for the fit as is in the paper.

In order to get an analytical representation of the radial dependence of the magnetic field, I derive a suitable fit function, assuming a nonexistent  $\mathbf{B}_\theta$ ,

$$x(r) = a \cdot \sqrt{(r^{e_1})^2 + (r^{e_2})^2} \quad (5.9)$$

with the scaling parameter  $a$ , and the exponents  $e_1$  and  $e_2$ , is used in a least squares regression fit. From the theoretical considerations sketched above, the expected fit parameter values are:  $a \approx \frac{B(1\text{ au})}{\sqrt{2}}$ ,  $e_1 \approx -2$ , and  $e_2 \approx -1$ .

The resulting fit curves for the mean and median field strength are plotted with respect to solar distance in [Figure 5.4](#). In the Helios distance region 0.3–1.0 au, the fit curves are very similar to those in the paper – as they are expected to be. This is why this graph is visually almost indistinguishable to the corresponding one in the paper (paper [Section 6.4](#), Figure 7). The resulting fit parameters are presented in [Table 5.1](#)<sup>1</sup>.



**Figure 5.4** Magnetic field strength with respect to solar distance. The mean and median per 0.01 au data bin and their fit curves are plotted as well. The hourly Helios data has a native distance resolution of 0.01 au, thus, to make the distribution visible in this plot, I added a random distance value of up to  $\pm 0.005$  au.

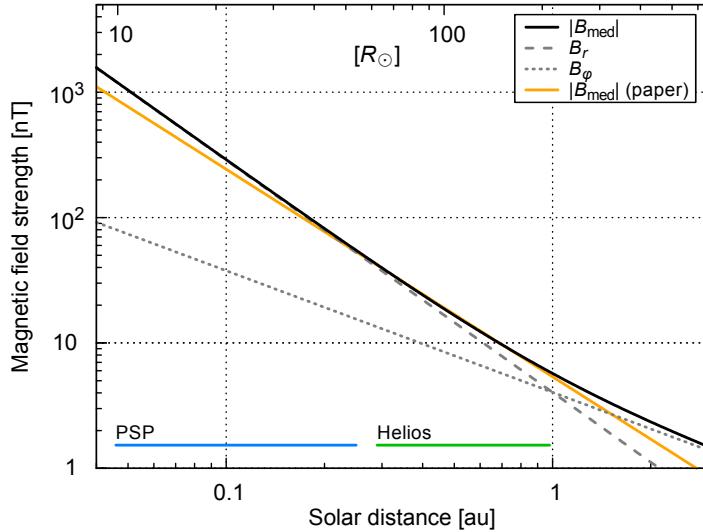
The weighted sum of the squared residuals (WSSR) of this fit slightly improved in comparison to that of the simple power-law function  $x(r) = d \cdot r^e$  used in [Venzmer & Bothmer \(2018\)](#) – 6 % for the median and 9 % for the mean. However, the mean and median fit curves are fairly similar and still cross each other at a similar value as in the paper.

The here obtained fit parameter values are closer to the theoretical values than those derived in the paper. The exponents  $e_{1,med} = -1.852$  and  $e_{1,avg} = -1.740$  are up to 13 % larger than the theoretical exponent 2 and

<sup>1</sup> Concerning the error notation, I adhere to the parentheses notation documented in the “Guide to the expression of uncertainty in measurement” (GUM) published by [Joint Committee for Guides in Metrology \(2008\)](#), where the numbers in parentheses are the errors on the corresponding last digits of the quoted value.

lead to a less steep slope near the Sun. However, they are about 12 % smaller than the power-law exponents derived in the paper, which have values of  $-1.655$  and  $-1.546$ .

It can be seen from [Figure 5.5](#) that although the new fit curve is slightly curved, both fits are congruent in the Helios data range.

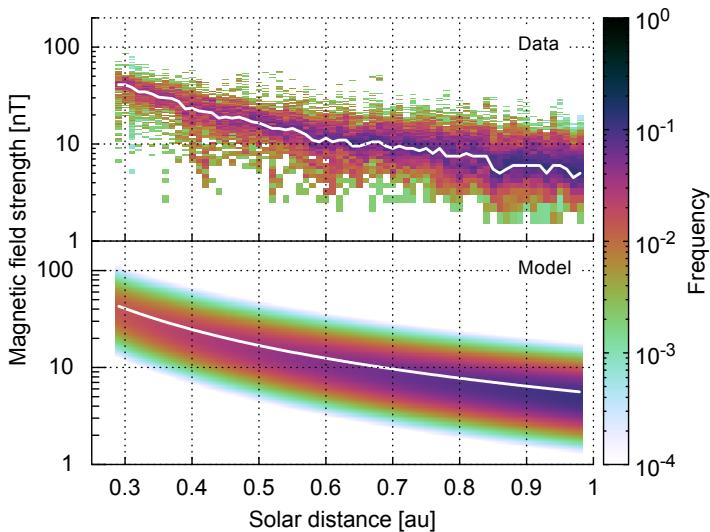


**Figure 5.5** Magnetic field strength fit curves with respect to solar distance. Plotted are the derived absolute median field strength and its radial and azimuthal components. I added the absolute median field strength from Equation (10) of the paper as well. The orbital distance ranges from PSP and the Helios probes are indicated by the blue and green lines.

In the following paragraph I use the same method and reasoning as explained in the paper [Section 6.4](#), thus I do not go into detail here.

...and it can be fitted via a lognormal function

see fit distribution in [Figure 5.6](#) and resulting fit parameters in [Table 5.1](#).



**Figure 5.6** Frequency distribution of the solar wind magnetic field strength with respect to solar distance. I plotted the binned Helios data (top panel) and the square-root power-law lognormal fit model (bottom panel) with their median values (white lines).

The result of this analysis can be used to scale the SSN-relation derived in the paper:

$$B(ssn, r) = \frac{B(ssn)}{\sqrt{2}} \cdot \sqrt{(r^{e_1})^2 + (r^{e_2})^2}. \quad (5.10)$$

The resulting equation, combined with the values obtained from the solar activity analysis in the paper [Sec-](#)

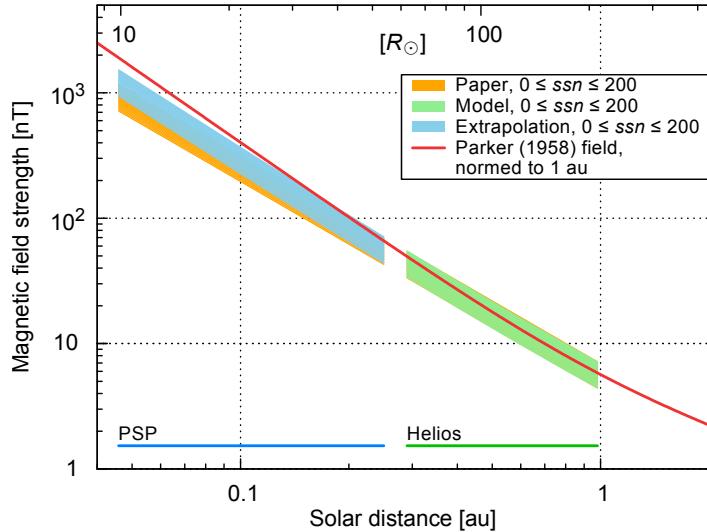
## 5. Empirical solar wind model for the inner heliosphere

tion 6.3, yields

$$B_{\text{med}}(ssn, r) = (0.0131 \text{ nT} \cdot ssn + 4.29 \text{ nT}) \cdot \sqrt{(r^{-1.86})^2 + (r^{-1.3})^2}, \quad (5.11)$$

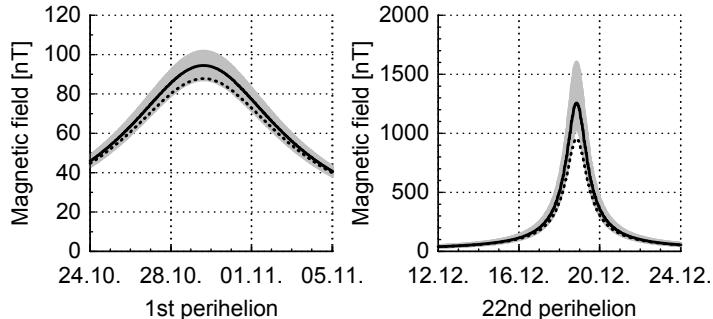
$$B_{\text{avg}}(ssn, r) = 1.0879 \cdot B_{\text{med}}(ssn, r). \quad (5.12)$$

This magnetic field strength model is extrapolated to PSP's orbital range in [Figure 5.7](#).



**Figure 5.7** Radial extrapolation of the solar wind magnetic field strength to the PSP orbit region. The median value from the model, obtained from Helios and OMNI measurements, is extrapolated to the PSP region for SSN values between solar minimum and maximum, that is,  $0 \leq ssn \leq 200$ . The lower edges of the shaded areas correspond to solar minimum, the upper edges to solar maximum. I also added the radial dependency of the field model by Parker (1958), normed to 1 au (red). The orbital distance ranges from PSP and the Helios probes are indicated by the blue and green lines.

Using the SSN prediction as described in the paper, the solar wind environment is forecasted to PSP's orbit and mission time, see [Figure 5.8](#).



**Figure 5.8** Estimated solar wind parameter medians (solid lines) and their error bands (gray area) during 12 days in 2018 and 2024 with PSP's first perihelion at about 0.16 au and PSP's 22nd (the first closest) perihelion at 0.046 au. I indicated the prediction from Venzmer & Bothmer (2018) as well (dotted lines).

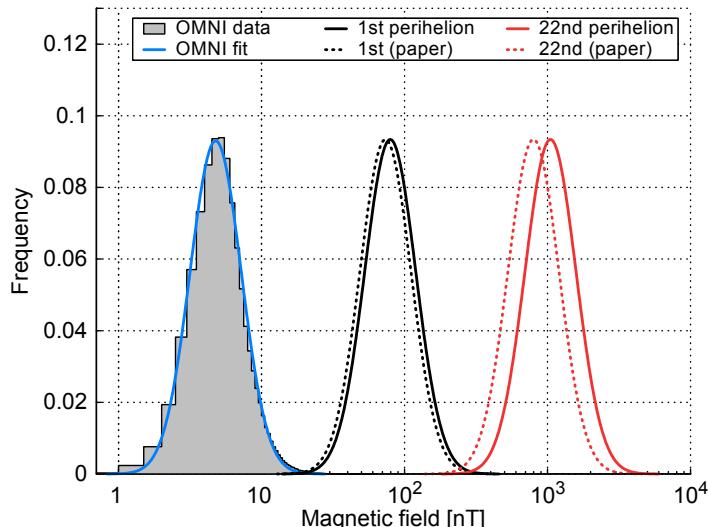
The estimated magnetic field strength median value of 94 nT for PSP's first perihelion is 8 % higher than that derived in the paper. However, the median value for the first closest perihelion is even 32 % higher with 1241 nT.

The estimated frequency distributions of the solar wind magnetic field strength at PSP's 1st and 22nd (first closest) perihelion are plotted in [Figure 5.9](#). During the first closest perihelion, high field strengths have multiple times the frequency of those derived in the paper.

## 5.5 Discussion

The newly derived dependency relation can be used to replace that found in the paper.

comparison with Burlaga1984: Large-scale IMF...; found for 1-5 au:  $B(r) = 4.75 \sqrt{1 + R^{**2}} / R^{**2}$



**Figure 5.9** Frequency distributions of the solar wind magnetic field strength (OMNI data) and those estimated with the solar wind model for PSP's 1st and 22nd (first closest) perihelion. In this plot the frequencies of both extrapolated curves are scaled for visibility to the same height as the 1 au distribution. I indicated the prediction from Venzmer & Bothmer (2018) is as well (dotted lines).

## 5.6 Outlook

Currently, the solar wind model described in this chapter and the paper is purely empirical and its four solar wind quantities are characterized independently from each other. In further steps, theoretical relations connecting the parameters could be introduced to make the model self-consistent. For example, flux conservation could be implemented into the radial distance dependencies, relating the two parameters density and velocity (see section XX). Or also, the velocity within the magnetic field's  $B_\phi$ -component (Equation 5.2) could be respected as a variable (not as a constant as it is done in the previous section).

inclusion of seasonal effects into the model, see Figure 5.10 and Figure 5.11; or error estimation...

latitudinal effects...

further investigations should be done into solar wind structure extrapolations  
solar wind structure investigation. solar wind structures differ from the average distance scaling laws, e.g., MCs show a different extrapolation behavior (cite?).

Data captured from spacecraft that studied Mercury (perihelion distance of 0.3 au), such as Mariner 10 and MESSENGER, could be included to refine the distance dependency of the solar wind model. However, PSP's near-Sun in-situ measurements, starting in fall 2018, will eventually reveal how accurate the predictions of the derived models are.

### 5.6.1 flux conservation

conserved quantities:

- momentum conservation... VBbookp112
- flux conservation...

With consideration of continuity, the mass flux per solid angle has to be constant:  $\dot{m} = \text{const}$   
conserved quantities:

- mass flux:  $\dot{m} = \rho v A$  (with mass density  $\rho$ , velocity  $v$  and solid angle  $A$ )
- particle fluxes (proton flux, electron flux, etc.)
- proton flux:  $j_p = n_p v_p A$  (with proton density  $n_p$  and proton velocity  $v_p$ )

(with proton mass density  $\rho = n_p m_p$  (with proton number density  $n_p$  and proton mass  $m_p$ )).

the individual radial dependencies for a spherical radial outflow are:  
 $A(r) \propto r^2 \rightarrow A/r^2 = \text{const}$   
and assuming exponential dependencies,

## 5. Empirical solar wind model for the inner heliosphere

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$$n_p(r) = n_0 r^{c_n}, \\ v(r) = v_0 r^{c_v}$$

$$j_p = \text{const} \quad (5.13)$$

$$n_p v_p A = \text{const} \quad (5.14)$$

$$n_0 r^{c_n} v_0 r^{c_v} r^2 = \text{const} \quad (5.15)$$

$$r^{c_n} r^{c_v} r^2 = \text{const} \quad (5.16)$$

$$\Rightarrow c_n + c_v + 2 = 0 \quad (5.17)$$

$$c_n + c_v = -2 \quad (5.18)$$

Thus, an increasing velocity results in a steeper density fall-off.

validity of mass flux continuity: within the heliosphere mass to energy conversion and vice versa is negligible, but there can be flux from and to higher latitudes as the Helios data is localized to a small latitude range in the ecliptic plane.

larger errors should be located near CMEs and CIRs (nonradial flows from deflections)

there is a proton flux difference between slow and fast solar wind streams (see book Schwenn 1990 p. 146)

estimate the possible size of error:

(constant mass flux only for mean)

$$c_n = -2.114$$

$$c_v = 0.0990$$

$$c_n + c_v = -2.024$$

difference to -2 is 0.024

simple view: For a spherical constant velocity mass outflow, a per distance squared law is expected for the density, because of the mass flux conservation per solid angle for different distances. Measurements up to the outer heliosphere confirm the  $r^{-2}$  dependency (1–38 au by Voyager 2, (Belcher et al. 1993) newer paper?)

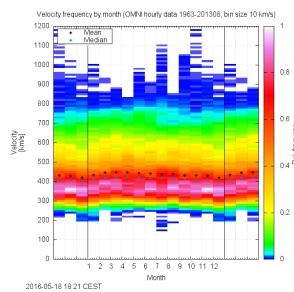
### 5.6.2 Seasonal solar wind variations

seasonal variation by month

quantify variation amplitudes

see [Figure 5.10](#) and [Figure 5.11](#)

make 4-panel figure...



**Figure 5.10** Diagram of the velocity frequency by month for the period 1963/01–2013/08. Mean and median values are shown as well.

derived exponent values from simple trigonometric fit on monthly values:

$$c_n = -2.234$$

maybe figure?



**Figure 5.11** Diagram of magnetic field frequency by month for the period 1963/01–2013/08. Mean and median values are shown as well as the expected course from the solar distance variation (obtained from Helios data).

expected influence from Earth's perihelion/aphelion (see Appendix...) distance vs observations we expect for the mean proton density (scaling law  $n(r) \propto r^{-2.114}$ ):

$$n(0.983 \text{ au}) = 7.098 \text{ cm}^{-3}$$

$$n(1 \text{ au}) = 6.845 \text{ cm}^{-3}$$

$$n(1.017 \text{ au}) = 6.605 \text{ cm}^{-3}$$

we expect for the magnetic field strength (paper scaling law  $n(r) \propto r^{-1.662}$ ):

$$B(0.983 \text{ au}) = 5.870 \text{ nT}$$

$$B(1 \text{ au}) = 5.705 \text{ nT}$$

$$B(1.017 \text{ au}) = 5.547 \text{ nT}$$

### 5.6.3 Latitude dependency

refer to Ulysses [Figure 2.16](#)

Ulysses swoops polar plots...

see Schwenn1990's Fig. 3.14

see also [Schwenn \(1990\)](#) p. 127

see also [Richardson et al. \(1995\)](#)

Balogh et al. (1999) p. 162 ff (origin and formation of CIRs in inner heliosphere with Helios data; latitude V dependence)

Helios latitude; see [Figure 5.12](#)



**Figure 5.12** The four solar wind parameter's HGI latitude dependency. Their mean and median values per  $0.1^\circ$  bin are plotted as well. I added a random distance value of up to  $\pm 0.005$  au in order to make the distribution visible in this plot.

with the exponential dependencies to 1 au projected solar wind parameters; there are only small changes with latitude in the range  $-7.25\text{--}7.25^\circ$   
have a look on distribution widths...

dependence from latitude in interval  $-7.25\text{--}7.25^\circ$  in Helios data negligible?, see Figure 5.13.  
estimate error ranges from latitude influence...



**Figure 5.13** Solar wind parameters projected to 1 au with respect to latitude. And their mean values, including weighted fit. I added a random distance value of up to  $\pm 0.005$  au in order to make the distribution visible in this plot. add projected median...

...plot Ulysses data into plot?

influence from latitude variation in data negligible? (see Ulysses Figure 2.16 in introduction). Helios probes within ecliptic  $\Rightarrow$  variation span equal to solar tilt:  $-7.25\text{--}7.25^\circ$ ; solar tilt/obliquity to ecliptic:  $i_\odot = 7.25^\circ$  (Sun fact sheet: <http://nssdc.gsfc.nasa.gov/planetary/factsheet/sunfact.html>)

big part of Helios data is from latitudes  $> \pm 5^\circ$ , see Figure XX (data count over latitude) and see Figure 3.8 (Helios orbit polar plane in data section)

## 5.6.4 Radial evolution of solar wind structures

see CGAUSS report2015\_2

Helios event lists HSSs, SLOWs, CIRs, CMEs...; event lists for all Helios data  
see Liu2004 for Helios ICME list and radial dependencies of  $B$ ,  $n$ ,  $T$  and  $v$ ...

200 km s<sup>-1</sup> slow solar wind at  $10 R_\odot$  is in agreement with blob measurements from Wang2000

very slow sw (VSSW) gets accelerated; see Sanchez-Diaz2016:

radial diameter of MCs increase between 0.3 au and 4.3 au proportional to the distance as  $r^{0.8}$  (Bothmer & Schwenn 1998)

MC central axial magnetic field strength radial density dependence  $B = 18.1 r^{-1.64}$  Leitner et al. (2007)

MC average diameter  $D = 0.23 r^{1.14}$  Leitner et al. (2007)

sw structure marked plots

## 5.7 notes...

Schwenn1983 intro -> sw-averaging comment (beer and wine) (cite him...)

see Hellinger2013 p.1353

see astro70/CGAUSS/dropbox\_presis/... (presi 1.07 Inside Helios-Origins and Evolution-Salem.ppt -> Helios reloaded; radial functions for all parameters)

see Balogh1999 from p. 162 (Helios CIR results)

see Marsch1999... (model constraints)

On solar wind acceleration and SPP proposition: McComas2007

Parker1963 book, p. 75 -> isothermal expansion figure

at larger distances (specify...) heating outbalances the adiabatic temperature part (adiabatic cooling vs. pickup proton and stream–interaction heating; 1–68 au by Voyager 2; [Richardson & Smith \(2003\)](#))  
solar wind ram pressure  $p_{\text{ram}} = \rho V^2$

replace paper with actual print published A&A version...

update paper reference...

remove sidc bib reference...



# Solar-wind predictions for the Parker Solar Probe orbit

## Near-Sun extrapolations derived from an empirical solar-wind model based on Helios and OMNI observations

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### ABSTRACT

**Context.** The Parker Solar Probe (PSP; formerly Solar Probe Plus) mission will be humanity’s first in situ exploration of the solar corona with closest perihelia at 9.86 solar radii ( $R_{\odot}$ ) distance to the Sun. It will help answer hitherto unresolved questions on the heating of the solar corona and the source and acceleration of the solar wind and solar energetic particles. The scope of this study is to model the solar-wind environment for PSP’s unprecedented distances in its prime mission phase during the years 2018–2025. The study is performed within the Coronagraphic German And US SolarProbePlus Survey (CGAUSS) which is the German contribution to the PSP mission as part of the Wide-field Imager for Solar Probe.

**Aims.** We present an empirical solar-wind model for the inner heliosphere which is derived from OMNI and Helios data. The German-US space probes Helios 1 and Helios 2 flew in the 1970s and observed solar wind in the ecliptic within heliocentric distances of 0.29–0.98 au. The OMNI database consists of multi-spacecraft intercalibrated in situ data obtained near 1 au over more than five solar cycles. The international sunspot number (SSN) and its predictions are used to derive dependencies of the major solar-wind parameters on solar activity and to forecast their properties for the PSP mission.

**Methods.** The frequency distributions for the solar-wind key parameters, magnetic field strength, proton velocity, density, and temperature, are represented by lognormal functions. In addition, we consider the velocity distribution’s bi-componental shape, consisting of a slower and a faster part. Functional relations to solar activity are compiled with use of the OMNI data by correlating and fitting the frequency distributions with the SSN. Further, based on the combined data set from both Helios probes, the parameters’ frequency distributions are fitted with respect to solar distance to obtain power law dependencies. Thus an empirical solar-wind model for the inner heliosphere confined to the ecliptic region is derived, accounting for solar activity and for solar distance through adequate shifts of the lognormal distributions. Finally, the inclusion of SSN predictions and the extrapolation down to PSP’s perihelion region enables us to estimate the solar-wind environment for PSP’s planned trajectory during its mission duration.

**Results.** The CGAUSS empirical solar-wind model for PSP yields dependencies on solar activity and solar distance for the solar-wind parameters’ frequency distributions. The estimated solar-wind median values for PSP’s first perihelion in 2018 at a solar distance of 0.16 au are 87 nT,  $340 \text{ km s}^{-1}$ ,  $214 \text{ cm}^{-3}$ , and 503 000 K. The estimates for PSP’s first closest perihelion, occurring in 2024 at 0.046 au ( $9.86 R_{\odot}$ ), are 943 nT,  $290 \text{ km s}^{-1}$ ,  $2951 \text{ cm}^{-3}$ , and 1 930 000 K. Since the modeled velocity and temperature values below approximately  $20 R_{\odot}$  appear overestimated in comparison with existing observations, this suggests that PSP will directly measure solar-wind acceleration and heating processes below  $20 R_{\odot}$  as planned.

**Key words.** solar wind – Sun: corona – Sun: heliosphere

### 1. Introduction

From observations of cometary tail fluctuations, Biermann (1951) inferred the presence of a continuous flow of particles from the Sun. With his theoretical solar-wind model, Parker (1958) formulated the existence of the solar wind even before the first satellites measured it in situ in 1959 (Gringauz et al. 1960; Neugebauer & Snyder 1966). The idea of a space mission flying through the solar corona dates back to the founding year of NASA in 1958 (McComas et al. 2008). Since then several space missions have measured the solar wind in situ at a wide range of heliocentric distances. In the case of Voyager 1, this was as far away as 140 au<sup>1</sup> in October 2017, having crossed the heliopause into interstellar space at a distance of 121 au (Gurnett et al. 2013). Various spacecraft have provided a wealth of solar-wind measurements near Earth’s orbit, with WIND (Lepping et al. 1995; Ogilvie et al. 1995), SOHO (Domingo et al. 1995), ACE

(Stone et al. 1998), and DSCOVR (Burt & Smith 2012) currently providing observations while orbiting around the L1 point 1.5 million km ahead of Earth in the sunward direction. Additional measurements at other solar distances were provided by planetary missions to Venus and Mercury, such as PVO (Colin 1980) or MESSENGER (Belcher et al. 1991). Ulysses was the first probe that orbited the Sun out of the ecliptic plane and thus could measure solar wind even at polar latitudes (McComas et al. 1998). The in situ solar-wind measurements closest to the Sun to date were made by the Helios missions. Helios 1, launched in 1974, reached distances of 0.31 au. Helios 2, launched two years later, approached the Sun as close as 0.29 au (Rosenbauer et al. 1977). The NASA Parker Solar Probe<sup>2</sup> (PSP), formerly Solar Probe Plus, six years after its planned launch date in mid 2018, will reach its closest perihelia at a distance of 9.86 solar radii ( $R_{\odot}$ ), that is, 0.0459 au

<sup>1</sup> <https://voyager.jpl.nasa.gov/>

<sup>2</sup> <http://parkersolarprobe.jhuapl.edu/>

(Fox et al. 2016). This distance will be achieved through seven Venus gravity assists with orbital periods of 88–168 days. In its prime mission time 2018–2025 PSP provides 24 orbits with perihelia inside 0.25 au (Fox et al. 2016). Even its first perihelion, 93 days after launch in 2018, will take PSP to an unprecedented distance of 0.16 au ( $35.7 R_{\odot}$ ). In comparison, the ESA Solar Orbiter mission with a planned launch in February 2019 will have its closest perihelia at 0.28 au (Müller et al. 2013).

The key PSP science objectives are to “trace the flow of energy that heats and accelerates the solar corona and solar wind, determine the structure and dynamics of the plasma and magnetic fields at the sources of the solar wind, and explore mechanisms that accelerate and transport energetic particles” as stated in Fox et al. (2016). To achieve these goals, PSP has four scientific instruments on board: FIELDS for the measurement of magnetic fields and AC/DC electric fields (Bale et al. 2016), SWEAP for the measurement of flux of electrons, protons, and alphas (Kasper et al. 2016), IS $\odot$ IS for the measurement of solar energetic particles (SEPs; McComas et al. 2016) and WISPR for the measurement of coronal and inner heliospheric structures (Vourlidas et al. 2016).

The study presented in this paper is undertaken in the Coronagraphic German And US SolarProbePlus Survey (CGAUSS) project, which is the German contribution to the PSP mission as part of the Wide-field Imager for Solar PRobe (WISPR). WISPR will contribute to the PSP science goals by deriving the three-dimensional structure of the solar corona through which the in situ measurements are made to determine the sources of the solar wind. It will provide density power spectra over a wide range of structures (e.g., streamers, pseudostreamers and equatorial coronal holes) for determining the roles of turbulence, waves, and pressure-balanced structures in the solar wind. It will also measure the physical properties, such as speed and density jumps of SEP-producing shocks and their coronal mass ejection (CME) drivers as they evolve in the corona and inner heliosphere (Vourlidas et al. 2016). In order to help optimize the WISPR and PSP preplanning of the science operations, knowledge of the expected solar-wind environment is needed. For this purpose the solar-wind environment is extrapolated down to the closest perihelion of  $9.86 R_{\odot}$  distance to the Sun using in situ solar-wind data from the Helios probes and near 1 au data from various satellites compiled in the OMNI solar-wind database.

Generally, two types of solar wind are observed in the heliosphere – slow and fast streams (Neugebauer & Snyder 1966; Schwenn 1983). Slow solar wind has typical speeds of  $<400 \text{ km s}^{-1}$  and fast solar wind has speeds  $>600 \text{ km s}^{-1}$  (Schwenn 1990, p. 144). Their different compositions and characteristics indicate different sources and generation processes (McGregor et al. 2011b). Fast streams are found to originate from coronal holes as confirmed by Ulysses’ out-of-ecliptic measurements (McComas et al. 1998). The source of slow wind, and its eventually different types (Schwenn 1983), is still a subject of controversial discussions because several scenarios are possible to explain its origin from closed magnetic structures in the solar corona, such as intermittent reconnection at the top of helmet streamers and from coronal hole boundaries (Kilpua et al. 2016). The occurrence frequency of these slow and fast streams varies strongly with solar activity and their interactions lead to phenomena such as stream interaction regions which may persist for many solar rotations (“co-rotating” interaction regions) if the coronal source regions are quasi-stationary (Balogh et al. 1999). Embedded in the slow and fast solar-wind streams are transient flows of CMEs – the faster ones driving shock waves ahead (Gosling et al. 1974). Their rate follows the

solar activity cycle and varies in near 1 au measurements between only one CME every couple of days during solar cycle minima up to multiple CMEs observed over several days at times of solar maxima, that is, the CME-associated flow share of the solar wind raises from about 5% up to about 50% (Richardson & Cane 2012).

It is not known which specific solar-wind type or structure PSP will encounter at a given time during its mission, therefore we extrapolate the probability distributions of the major solar-wind parameters from existing solar-wind measurements and take solar cycle dependencies into account. As a baseline we describe the solar-wind environment through the key quantities of a magnetized plasma: magnetic field strength, density and temperature. Furthermore, the bulk flow velocity is the defining parameter of the two types of solar wind. Solar-wind quantities, like flux densities, mass flux, and plasma beta, can be directly derived from these four parameters. In the analyses, we treat the solar wind as a proton plasma – the average helium abundance is about 4.5% and in slow wind at solar cycle minimum is even less than 2% (Feldman et al. 1978; Schwenn 1983; Kasper et al. 2012).

Our approach is to obtain analytical representations of the shapes of the solar-wind parameter’s frequency distributions in Sect. 2, of their solar activity dependence in Sect. 3 and of their solar distance scaling in Sect. 4. The solar-wind parameters’ frequency distributions and solar activity dependence is derived from near-Earth solar wind and sunspot number (SSN) time series with a duration of almost five solar cycles. Their distance dependency is derived from Helios solar-wind measurements covering more than two thirds of the distance to the Sun and more than half a solar cycle. From a combination of the obtained frequency distributions, SSN dependence functions, and solar distance dependence functions, a general solar-wind model is built in Sect. 5, representing the solar activity and distance behavior. Finally, this empirical model is fed with SSN predictions and extrapolated to PSP’s planned orbital positions in Sect. 6.

## 2. Frequency distributions of the solar-wind parameters

The solar-wind parameters are highly variable due to short-term variations from structures such as slow and fast wind streams, interaction regions, and CMEs, whose rate and properties depend on the phase of the solar activity cycle. Hence, for deriving characteristic frequency distributions for the solar-wind parameters, measurements over long-term time spans are needed. The abundance of the near-Earth hourly OMNI data set is ideally suited for this purpose, because to date it spans almost five solar cycles.

The OMNI 2 data set (King & Papitashvili 2005) combines solar-wind magnetic field and plasma data collected by various satellites since 1963, currently by WIND and by ACE. This intercalibrated multi-spacecraft data is time-shifted to the nose of the Earth’s bow shock. The data is obtained from the OMNIWeb interface<sup>3</sup> at NASA’s Space Physics Data Facility (SPDF), Goddard Space Flight Center (GSFC). In this study the whole hourly data until 31 December 2016 is used, starting from 27 November 1963 (for the temperature from 26 July 1965). The data coverage of the different parameters is in the range 67–74%, corresponding to a total duration of 36–40 yr. We note that a test-comparison of hourly averaged data with higher-time-resolution data for the available shorter time span 1981–2016 did not show significant differences in our results. According to the OMNI

<sup>3</sup> <http://omniweb.gsfc.nasa.gov/>



**Fig. 1.** Frequency distributions of the four solar-wind parameters and their lognormal fits derived from the hourly OMNI data set. The histograms have bins of 0.5 nT, 10 km s<sup>-1</sup>, 1 cm<sup>-3</sup>, and 10 000 K. The fits' median and mean values are indicated as well. The *insets* show zoomed-in views of the high-value tails of the distributions.

data precision and maximal parameter ranges we specify bin sizes of 0.5 nT for the magnetic field strength, 10 km s<sup>-1</sup> for the velocity, 1 cm<sup>-3</sup> for the density, and 10 000 K for the temperature. The frequency distributions of the solar-wind magnetic field strength, proton velocity, density and temperature are shown in Fig. 1. The solar-wind magnetic field strength is in the range 0.4–62 nT, the velocity in the range 156–1189 km s<sup>-1</sup>, the density in the range 0–117 cm<sup>-3</sup>, the temperature in the range 3450–6.63 × 10<sup>6</sup> K, and the mean data values are at 6.28 nT, 436 km s<sup>-1</sup>, 6.8 cm<sup>-3</sup>, and 1.05 × 10<sup>5</sup> K. These ranges and mean values are as statistically expected from previous analyses of near 1 au solar-wind data (e.g., Table 3.3 in Bothmer & Daglis 2007, p. 39). Much higher or lower peak values at 1 au have been observed in extraordinary events, such as the 23 July 2012 CME with a speed of over 2000 km s<sup>-1</sup> and a peak field strength of about 100 nT that was observed by STEREO A (Russell et al. 2013), or the solar-wind disappearance event observed in May 1999 with density values even down to 0.2 cm<sup>-3</sup> (Lazarus 2000).

The frequency distributions of the solar-wind parameters, magnetic field strength, proton density, and temperature, can be well approximated by lognormal distributions, whereas the proton velocity's frequency has a differing shape, as shown in Veselovsky et al. (2010). We investigate how well all four solar-wind parameters' frequency distributions can be represented by lognormal functions, which we use in the process of a least squares regression fitting. The lognormal function,

$$W(x) = \frac{1}{\sigma \sqrt{2\pi} x} \exp\left(-\frac{(\ln x - \mu)^2}{2\sigma^2}\right), \quad (1)$$

depends on the location  $\mu$  and the shape parameter  $\sigma$ . Changes in  $\mu$  affect both the horizontal and vertical scaling of the function whereas  $\sigma$  influences its shape. The distribution's median  $x_{\text{med}}$  and mean  $x_{\text{avg}}$  (average) positions are easily interpreted and are directly calculated from  $\mu$  and  $\sigma$ :

$$x_{\text{med}} = \exp(\mu) \quad \Longleftrightarrow \quad \mu = \ln(x_{\text{med}}), \quad (2)$$

$$x_{\text{avg}} = \exp\left(\mu + \frac{\sigma^2}{2}\right) \quad \Longleftrightarrow \quad \sigma = \sqrt{2 \ln\left(\frac{x_{\text{avg}}}{x_{\text{med}}}\right)}. \quad (3)$$

It is apparent that the mean is always larger than the median. Replacing the variables  $\mu$  and  $\sigma$  with these relations, the lognormal function (1) becomes

$$W(x) = \frac{1}{2 \sqrt{\pi \ln\left(\frac{x_{\text{avg}}}{x_{\text{med}}}\right)} x} \exp\left(-\frac{\ln^2\left(\frac{x}{x_{\text{med}}}\right)}{4 \ln\left(\frac{x_{\text{avg}}}{x_{\text{med}}}\right)}\right). \quad (4)$$

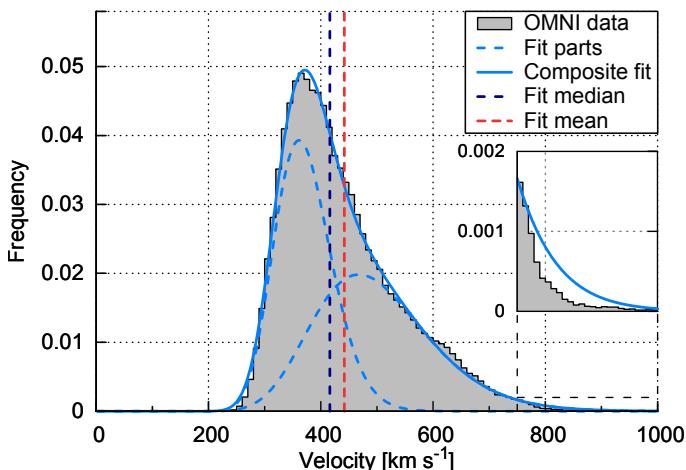
The values of  $x_{\text{med}}$  and  $x_{\text{avg}}$  obtained from fitting the individual solar-wind frequency distributions are listed in Table 1.

From visual inspection, the resulting fit curves describe the shape of the magnetic field strength, the density and the temperature distributions well, as can be seen in Fig. 1. However, for the velocity, the fit function appears not to be as good in describing the measured distribution's more complex shape around its peak and in the higher velocity range. This also can be inferred from the sum of absolute residuals (SAR) between data and fit, listed in Table 1 as a percentage of the distribution area, being almost three times larger than those from the other parameters. In order to find a better fit result for the velocity distribution, we

**Table 1.** Resulting fit coefficients from the fitting of the lognormal function (4) to the shape of the solar-wind parameters' frequency distributions from near 1 au OMNI hourly data.

Parameter	Median $x_{\text{med}}$	Mean $x_{\text{avg}}$	Balance $c$	SAR [%]
Magnetic field [nT]	5.661(16)	6.164(18)	–	6.83
Velocity [ $10^2 \text{ km s}^{-1}$ ]	4.085(19)	4.183(20)	–	18.69
Density [ $\text{cm}^{-3}$ ]	5.276(24)	6.484(34)	–	6.48
Temperature [ $10^4 \text{ K}$ ]	7.470(17)	11.301(32)	–	5.78
Velocity [ $10^2 \text{ km s}^{-1}$ ]	$W_1$ 3.68(20) $W_2$	4.89(14) 3.72(20) $W_{\text{II}}$	5.00(14) 4.42(14) <sup>a</sup>	0.504(62) – –
			$c$	4.20

**Notes.** For the velocity, the fit parameters of the double lognormal function (5) are also listed, as well as the median and mean values of the resulting velocity fit. The numbers in parentheses are the errors on the corresponding last digits of the quoted value. They are calculated from the estimated standard deviations of the fit parameters. For each parameter, the sum of absolute residuals between data and fit (in percentage of the distribution area) is also listed. <sup>(a)</sup> Error estimates derived from the individual fit part errors.



**Fig. 2.** Velocity frequency distribution (same as in Fig. 1) and its compositional lognormal fit. The fit's median and mean values and its two fit parts are indicated as well. The *inset* is a zoomed-in view of the high value tail of the distribution.

assume that the velocity distribution can be made up of at least two overlapping branches (McGregor et al. 2011a). Therefore a compositional approach is chosen by combining two lognormal functions (4), involving more fit variables:

$$W_{\text{II}}(x) = c \cdot W_1(x) + (1 - c) \cdot W_2(x). \quad (5)$$

The balancing parameter  $c$  ensures that the resulting function remains normalized as it represents a probability distribution. The fitting of  $W_{\text{II}}(x)$  to the velocity's frequency distribution yields the values of the now five fit parameters ( $c$ ,  $x_{\text{med},1}$ ,  $x_{\text{avg},1}$ ,  $x_{\text{med},2}$  and  $x_{\text{avg},2}$ ) as listed in Table 1 together with the median and mean values of the composed distribution, which can be derived by solving

$$\int W_{\text{II}}(x) dx = 0 \quad \text{and} \quad \int x W_{\text{II}}(x) dx = 0. \quad (6)$$

This more complex fit function is more accurate in describing the velocity's frequency distribution as shown in Fig. 2. Thus in the following Sections we keep the double lognormal ansatz for all velocity frequency fits.

For the bulk of the solar wind these static lognormal functions describe the parameters' distributions well. The abnormally high parameter values in the distribution functions can be

attributed to shock/CME events in agreement with the results of the OMNI solar-wind investigations by Richardson & Cane (2012). The simple lognormal fit functions underestimate the frequencies in their high-value tails, except for the temperature's tail which is overestimated, as seen in the insets of Fig. 1. This appears to be because CMEs do not come with abnormally high temperatures, but rather with temperatures lower than those of the average solar wind (Forsyth et al. 2006). The velocity's compositional lognormal fit only slightly overestimates its tail as seen in the inset of Fig. 2. The slow and fast part contribute almost equally ( $c \approx 0.5$ ) to the long-term velocity distribution function.

### 3. Solar activity dependence of the solar-wind frequency distributions

In the next step we investigate how the long-term solar-wind distribution functions presented in the previous section depend on general solar activity. Therefore we examine their correlation with the SSN, being a commonly used long-term solar activity index, and determine the time lags with the highest correlation coefficients.

For the correlations we fit lognormal functions to the frequency distributions as in Sect. 2, but implement linear relations to the yearly SSN, allowing shifting of the distribution functions with SSN. For the velocity the approach is different insofar as its two components are kept fixed and instead their balance is modified with the changing SSN. Thus we obtain solar-activity-dependent models for the frequency distributions of all four solar-wind parameters.

The international sunspot number (1963–2016) is provided by the online catalog<sup>4</sup> at the World Data Center – Sunspot Index and Long-term Solar Observations (WDC-SILSO), Solar Influences Data Analysis Center (SIDC), Royal Observatory of Belgium (ROB).

Yearly medians of the solar-wind parameters and the yearly SSN together with the solar cycle number are shown in the upper part of Fig. 3. The reason for correlating the SSN to the solar-wind median values is because the position of a lognormal function is defined by its median. The data are averaged to yearly values to avoid seasonal effects during the Earth's orbit around the Sun caused by its variations in solar latitude and distance. The solar-wind velocity, density, and temperature depend

<sup>4</sup> <http://www.sidc.be/silso/>

**Table 2.** Resulting fit coefficients from the OMNI data, based on the linear SSN dependencies (7) and (8).

Parameter	Median		Mean scaling factor $a_{\text{avg}}$	Balance		SSN lag [years]
	SSN factor $a_{\text{med}}$	baseline $b_{\text{med}}$		SSN factor $c_a$	baseline $c_b$	
Magnetic field [nT]	$1.309(19) \times 10^{-2}$	4.285(17)	$8.786(78) \times 10^{-2}$	—	—	0
Density [ $\text{cm}^{-3}$ ]	$3.81(25) \times 10^{-3}$	4.495(26)	$3.050(27) \times 10^{-1}$	—	—	6
Temperature [ $10^4 \text{ K}$ ]	$1.974(26) \times 10^{-2}$	5.729(19)	$6.541(28) \times 10^{-1}$	—	—	3
Velocity [ $10^2 \text{ km s}^{-1}$ ]	$W'_1$ $W'_2$	—	3.633(12) 4.831(81)	$1.008(37) \times 10^{-2}$ $2.31(20) \times 10^{-2}$	-1.799(95)e-3 0.638(32)	3

**Notes.** For the velocity the fit parameters from the double lognormal fit (5) and their balancing function (9) are given. The numbers in parentheses are the errors on the corresponding last digits of the quoted value. They are calculated from the estimated standard deviations of the fit parameters. The listed SSN time lags are used for the fits.



**Fig. 3.** Solar-wind parameter yearly frequencies (gray shading) with yearly medians (lines) derived from OMNI data and the yearly SSN from the SILSO World Data Center (1963–2016) with solar cycle number (top). Their correlation coefficients with the yearly SSN are calculated for time lags back to -15 yr (bottom).

on the state of the solar cycle (Schwenn 1983). For instance the fast solar wind occurs at times when polar coronal holes extend to lower latitudes, a typical feature of the declining phase of the solar cycle as pointed out by Bothmer & Daglis (2007, p. 75, Fig. 3.52). Therefore the solar-wind velocity, density, and temperature maxima exhibit time lags relative to the SSN maxima.

The correlation coefficients of the solar wind parameters with the yearly SSN shown in the bottom part of Fig. 3 are calculated for time lags back to -15 yr to cover a time span longer than a solar cycle. As expected, the amplitudes of the variations in the correlations of all parameters decline with increasing time lag and show a period of about 11 yr. The highest correlation coefficient of 0.728 to the SSN is found for the magnetic field.

strength; it has no time lag. This finding is anticipated because the SSN is found to be directly proportional to the evolution of the photospheric magnetic flux (Smith & Balogh 2003). Velocity and temperature show time lags of 3 yr with peak correlation coefficients of 0.453 and 0.540. The density with a correlation coefficient of 0.468 has a time lag of 6 yr, which is in agreement with the density anticorrelation to the SSN reported by Bougeret et al. (1984).

Next we create solar-activity-dependent analytical representations of the solar wind frequency distributions. This is achieved by shifting the median positions of the lognormal distributions as a linear function of the SSN. To enable these shifts, we add a linear SSN dependency to the median,

$$x_{\text{med}}(\text{ssn}) = a_{\text{med}} \cdot \text{ssn} + b_{\text{med}}, \quad (7)$$

using a factor to the SSN  $a_{\text{med}}$  with a baseline  $b_{\text{med}}$ . We relate the mean with a scaling factor to the median to transfer its SSN dependency:

$$x_{\text{avg}}(\text{ssn}) = (1 + a_{\text{avg}}) \cdot x_{\text{med}}(\text{ssn}). \quad (8)$$

These relations, substituted into the lognormal function (4), lead to a new SSN-dependent function  $W'(x, \text{ssn})$ . This function is then fitted to the yearly data, using the yearly SSN as input parameter. The SSN is offset with the individual time lags determined before for each parameter, to benefit from the higher correlation. The values of the three resulting fit coefficients ( $a_{\text{med}}$ ,  $b_{\text{med}}$ , and  $a_{\text{avg}}$ ) are presented in Table 2.

Naturally, the fit models match with the general data trends, as can be seen from Fig. 4, though single year variations are not replicated by the model (e.g., the high velocity and temperature values in 1974, 1994, and 2003). The comparison of this model with the yearly data median values with respect to the lagged SSN shows that the medians obtained from the modeling have a similar slope, as shown in Fig. 5.

Again, the solar-wind velocity needs a special treatment because of the application of the double lognormal distribution (5). Since it is well known that slow and fast solar-wind stream occurrence rates follow the solar cycle, we keep the two velocity components' positions SSN-independent ( $x_{\text{med}} = b_{\text{med}}$ ) and vary instead their balance with the SSN:

$$c(\text{ssn}) = c_a \cdot \text{ssn} + c_b. \quad (9)$$

The fit result (see Table 2) yields a model in which three years after solar cycle minimum (SSN of zero) the contribution of slow solar wind to the overall solar wind distribution reaches a maximum value (about 64%) and decreases with increasing SSN as shown in Fig. 6.



**Fig. 4.** Solar wind parameter yearly data frequencies and lognormal fit models, both with their median values (white lines) over the OMNI time period 1963–2016. The corresponding yearly SSN and the shifted SSN for the models are indicated by gray and black lines. The velocity median is derived from the SSN-weighted constant lognormal parts (dotted lines).

To investigate the amount of slow and fast wind contributions depending on solar activity, we apply the commonly used constant velocity threshold of  $v_{\text{th}} = 400 \text{ km s}^{-1}$  (Schwenn 1990, p. 144). The linear fit to the yearly data ratio and the derived model ratio show a good agreement (see Fig. 6). The to-some-degree steeper balance parameter of the double fit function used in this model cannot be compared directly with specific velocity thresholds between slow and fast solar wind. However, it appears to be a more realistic approach than just taking a specific velocity threshold for the slow and fast wind, in agreement with the overlapping nature of the velocity flows reported by McGregor et al. (2011a).

#### 4. Solar distance dependency

In order to derive heliocentric distance relationships of the bulk solar wind distribution functions, we apply and fit power law dependencies to the Helios data. We then examine how the fits may be extrapolated towards the Sun and in particular in to the PSP orbit. We use the fitting methods of Sect. 2 for the distance-binned combined data from both Helios probes. Helios' highly elliptical orbits in the ecliptic covered a solar distance range of 0.31–0.98 au in case of Helios 1 and 0.29–0.98 au in case of Helios 2. Launched during solar cycle minimum, the data of both probes cover the rise to the maximum of cycle 21, covering ~6.5 yr at varying distances to the Sun.

We investigate hourly averages of the Helios data in the same way as with the OMNI data. The Helios 1 merged

hourly data from the magnetometer and plasma instruments (Rosenbauer et al. 1977) include ~12.5 orbits for the time range 10 December 1974 to 14 June 1981, and those for Helios 2 include ~8 orbits for the time span 1 January 1976 to 4 March 1980. The data are retrieved from the Coordinated Data Analysis Web (CDAWeb) interface at NASA's GSFC/SPDF<sup>5</sup>.

The Helios 1 magnetometer data coverage for this data set is about 43% (i.e., 2.8 yr), and that of Helios 2 amounts to 54% (i.e., 2.3 yr). The plasma data coverage is 76% (i.e., 5.0 yr) in case of Helios 1 and 92% (i.e., 3.9 yr) in case of Helios 2. Thus, using this data, we point out that its time coverage is unequally distributed over the solar cycle. Considering the data gap distributions, the amount of data during solar cycle minimum up to mid 1977, that is, the transition from minimum to maximum, covers about 68% of this period whereas during maximum of cycle 21 data are available only 38% of the time. This Helios data bias towards solar minimum is one reason why in this study the Helios solar wind data are not used to derive long-term frequency distributions and solar-cycle dependencies for the key solar wind parameters.

The radial dependencies of the key solar-wind parameters over the distance range 0.29–0.98 au measured by both Helios probes are plotted in Fig. 7, together with their median and mean values for different solar distances, calculated for the minimal distance resolution 0.01 au of the data set. Assuming a radial solar-wind outflow, it is expected that the distance

<sup>5</sup> <http://spdf.gsfc.nasa.gov/>

**Table 3.** Fit coefficients for the median and mean solar distance dependencies (10) of the four solar wind parameters derived from the combined Helios 1 and 2 data.

Parameter	Median		Mean		Crossing distance [au]	Yearly variation $\Delta e$
	$d_{\text{med}}$	$e_{\text{med}}$	$d_{\text{avg}}$	$e_{\text{avg}}$		
Magnetic field [nT]	5.377(92)	-1.655(17)	6.05(10)	-1.546(18)	0.339(11)	0.11
Velocity [ $10^2 \text{ km s}^{-1}$ ]	4.107(28)	0.058(13)	4.356(24)	0.049(10)	$0.7(83) \times 10^3$	0.012
Density [ $\text{cm}^{-3}$ ]	5.61(27)	-2.093(46)	7.57(30)	-2.010(38)	0.027(73)	0.072
Temperature [ $10^4 \text{ K}$ ]	7.14(23)	-0.913(39)	9.67(21)	-0.792(28)	0.082(85)	0.050

**Notes.** The numbers in parentheses are the errors on the corresponding last digits of the quoted value. They are calculated from the estimated standard deviations of the fit parameters. The crossing distances indicate where the median and mean fits intersect each other. The yearly variation is the weighted standard deviation derived from the yearly fit exponents seen in Fig. 9.



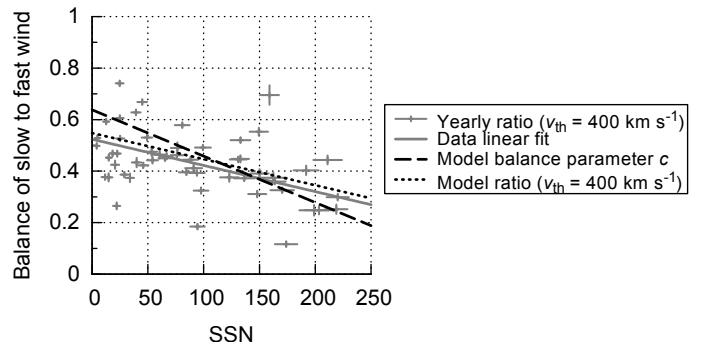
**Fig. 5.** Solar-wind parameter medians with respect to the lagged SSN. The yearly data medians (+) with their weighted linear fit (solid lines) are obtained from OMNI data. The error bars denote the SSN standard deviation and the relative weight from the yearly data coverage. The SSN-dependent median (dashed lines) is derived from the lognormal model fit. For the velocity the median is derived from the SSN-weighting (9) of the slow and fast model parts (dotted lines), whose magnitudes are SSN independent.

dependence of the solar-wind parameters over the Helios data range 0.29–0.98 au can be described through power law scaling. Therefore we use the power law function,

$$x(r) = d \cdot r^e, \quad (10)$$

for the regression fit of the median and mean, with  $r$  being the solar distance in astronomical units,  $d$  the magnitude at 1 au, and  $e$  the exponent. The fits are weighted through the different data counts per bin. The obtained coefficients for the median and mean power law fits ( $d_{\text{med}}$ ,  $e_{\text{med}}$ ,  $d_{\text{avg}}$ , and  $e_{\text{avg}}$ ) are listed in Table 3 and their corresponding curves are shown in Fig. 7.

Our derived exponents agree with those found in existing studies from the Helios observations: Mariani et al. (1978) derived the exponents for the magnetic field strength separately for the fast and slow solar wind as  $B_{\text{fast}} \propto r^{-1.54}$  and  $B_{\text{slow}} \propto r^{-1.61}$ , ours is  $B_{\text{avg}} \propto r^{-1.55}$ . The velocity exponent  $v_{\text{avg}} \propto r^{0.049}$  matches with the values found by Schwenn (1983, 1990), who derived the distance dependencies for both Helios spacecraft separately as  $v_{\text{H1}} \propto r^{0.083}$  and  $v_{\text{H2}} \propto r^{0.036}$ . The calculated density exponent  $n_{\text{avg}} \propto r^{-2.01}$  agrees well with the Helios plasma density model derived by Bougeret et al. (1984) yielding  $n \propto r^{-2.10}$ . The temperature exponent  $T_{\text{avg}} \propto r^{-0.79}$  is similar to those in the studies



**Fig. 6.** Ratio of slow to fast solar wind for a SSN lagged by 3 yr. The yearly ratios (+) and their weighted linear fit (solid line) are obtained from OMNI data with a threshold velocity of  $v_{\text{th}} = 400 \text{ km s}^{-1}$ . The error bars denote the SSN standard deviation and the relative weight from the yearly data coverage. The model's balance parameter (9) and derived ratio (same threshold) are plotted as dashed and dotted lines.

by Hellinger et al. (2011, 2013), who also derived the exponents separately for the fast and the slow solar wind:  $T_{\text{fast}} \propto r^{-0.74}$  and  $T_{\text{slow}} \propto r^{-0.58}$ .

The mean and median velocity fit exponents acquired from the Helios data are very similar, which indicates that they can be kept identical so that the basic shape of the frequency distribution does not change with distance. Conversely, the mean and median fits for the magnetic field strength cross each other at 0.339 au (see Table 3) and the mean becomes slightly lower than the median at smaller distances. Thus, below that distance the frequency distribution can no longer be well described by a lognormal function, because the mean of a lognormal function has to be larger than its median (as pointed out in Sect. 2), that is, the location of the crossing indicates that the parameter's distribution is no longer of a lognormal shape thereafter. The fits for the proton temperature show a similar behavior, having an extrapolated intersection at 0.082 au. Therefore the extrapolation of the magnetic field and temperature distribution frequencies to the PSP orbit by applying lognormal functions is limited. The crossing points limit the regions where the distribution's shapes can still be considered lognormal.

In order to still fit and extrapolate lognormal functions with the data, we assume that the shapes can be considered lognormal at all distances. For the frequency distribution fit function to be discussed in the following paragraph, we reduce the fit exponents  $e_{\text{med}}$  and  $e_{\text{avg}}$  to only one. We note that this simplification leads to slightly larger modeling errors, especially in case of the magnetic field strength.

Next we retrieve the frequency distributions of the four solar wind parameters in solar distance bins of 0.01 au, choosing the



**Fig. 7.** Helios hourly data plots of the four solar wind parameters over solar distance. The mean and median per 0.01 au data bin and their fit curves are plotted as well. The Helios data has a native distance resolution of 0.01 au, thus, to make the distribution visible in these plots, we added a random distance value of up to  $\pm 0.005$  au. The high velocity data points above  $800 \text{ km s}^{-1}$  (circled red) are identified as CME events (e.g., Sheeley et al. 1985; Bothmer & Schwenn 1996, 1998).

**Table 4.** Fit coefficients for the distance-dependent single lognormal function, based on Eq. (4) combined with (10) from the combined Helios data.

Parameter	Median $d'_{\text{med}}$	Mean $d'_{\text{avg}}$	Exponent $e'$	Balance $c'$	Seasonal variation $\Delta d [\%]$
Magnetic field [nT]	5.358(25)	5.705(28)	-1.662(11)	—	2.8
Density [ $\text{cm}^{-3}$ ]	5.424(33)	6.845(47)	-2.114(20)	—	3.6
Temperature [ $10^4 \text{ K}$ ]	6.357(64)	10.72(14)	-1.100(20)	—	1.9
Velocity [ $10^2 \text{ km s}^{-1}$ ]	$W''_1$ 3.707(13)	3.748(16)	0.0990(51)	0.557(45)	0.17
	$W''_2$ 5.26(13)	5.42(11)			
	$W''_{\text{II}}$ 4.13(13) <sup>a</sup>	4.47(11) <sup>a</sup>	—	—	—

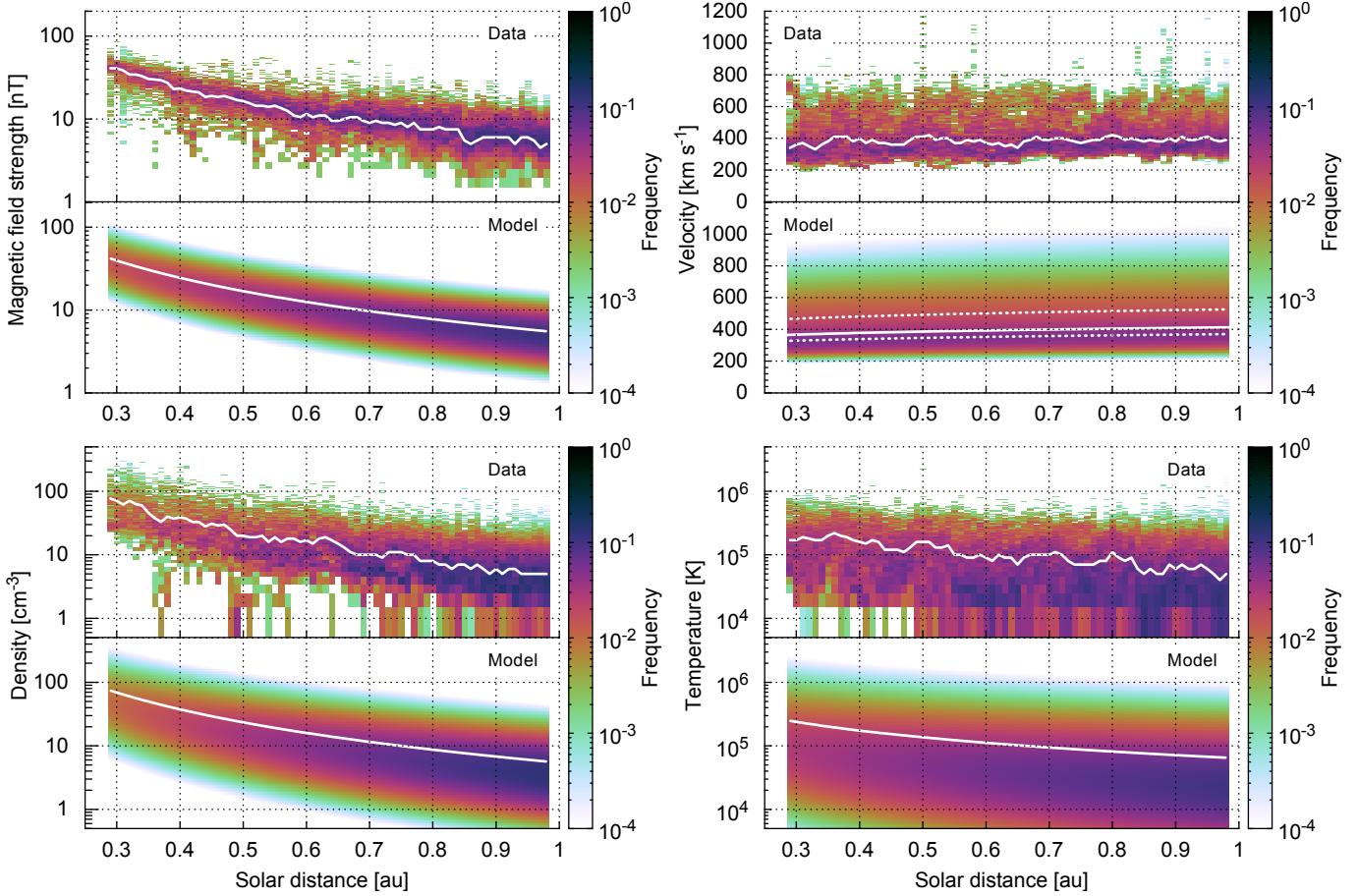
**Notes.** Regarding the velocity, the double lognormal function (5) is used instead. The numbers in parentheses are the errors on the corresponding last digits of the quoted value. They are calculated from the estimated standard deviations of the fit parameters. The seasonal variations are calculated from Earth's orbital solar distance variation and the derived exponents. <sup>(a)</sup> Velocity median and mean 1 au values for the resulting function. Error estimates derived from the individual fit part errors.

same resolution as for the OMNI data analyzed in Sect. 2 – the distributions and their median values are plotted in Fig. 8. For simplification, as mentioned before, we treat the exponents of the median and mean fit functions as being identical, using one fit parameter for both. Implementing the power law distance dependency (10) into the lognormal function (4), we get the fit parameters  $d'_{\text{med}}$ ,  $d'_{\text{avg}}$  and the common exponent  $e'$ . Again, we use the double lognormal function (5) for the velocity distribution fit – resulting in  $W''_{\text{II}}(x, r)$ . The additional fit parameters are the balancing parameter  $c'$  and for the second lognormal part  $d'_{\text{med},2}$ ,

and  $d'_{\text{avg},2}$ . The resulting fit coefficients for the four solar wind parameters are presented in Table 4.

The velocity balancing parameter  $c' = 0.557$  is in good agreement with the results for the SSN dependency (9), because with a mean SSN of 59 during the Helios time period,  $c(59) = 0.53$ , as can be seen from Fig. 6.

The power law lognormal models and the power law double lognormal model for the velocity, which result from the fitting, are plotted in Fig. 8 together with their median values. The model's magnetic field strength is broader around values



**Fig. 8.** Frequency distributions of the four solar-wind parameters with respect to solar distance. Plotted are the binned Helios data and the power law lognormal fit models with their median values (white lines). The double lognormal model is used for the velocity, its slow and fast parts are indicated by dotted lines.

of 40 nT at the lower distance boundary than the data's frequency distribution implies. This behavior is expected because of the applied distance-independent shape approximation. The velocity and temperature models' upper values generally show a higher abundance than the actual data; see also zoom boxes in Figs. 1 and 2. The high-velocity tail that increases with distance arises from using the same exponent for both slow and fast components. This effect is not seen in the data; more specifically, not only the slowest wind but also the fastest wind is expected to converge to more average speeds (Sanchez-Diaz et al. 2016).

## 5. Empirical solar-wind model

In order to estimate the solar-wind environment for the PSP orbit, we combine the results from the solar-wind frequency distributions' solar-activity relationships and their distance dependencies derived from the OMNI and Helios data. The result is an empirical solar-wind model for the inner heliosphere which is then extrapolated to the PSP orbit in Sect. 6.

This solar-wind model for the radial distance dependence is representative for the time of the Helios observations around the rise of solar cycle 21. The variations of the yearly power law fit exponents from fitting the solar-distance dependency (10) are shown in Fig. 9 together with the yearly SSN for the time period 1974–1982. It can be seen that during the Helios time period there might be some systematic variation of the exponents with solar activity – at least for the velocity and temperature



**Fig. 9.** Helios yearly variation of the solar-wind parameter power exponents for the dependence on radial distance together with the SIDC 13-month smoothed monthly SSN. The weighted standard deviations and average values for all years are indicated by the shaded areas. In this plot, the 21 days since Helios launch in the year 1974 are omitted because a distance range of merely 0.95–0.98 au was covered that year.

exponents. However, for simplicity we assume that the distance scaling laws can be treated as time independent and include

the calculated exponents' yearly variations  $\Delta e$ , summarized in Table 3, as relative uncertainties.

Since we neglect possible variations of the distance scaling laws, we combine the frequency distribution's median solar activity dependency (7) derived for 1 au from the OMNI data with the power law exponents (10) derived from the Helios data:

$$x_{\text{med}}(ssn, r) = (a_{\text{med}} \cdot ssn + b_{\text{med}}) \cdot r^{e'}. \quad (11)$$

Thus, implementing the median and mean relations into the lognormal function (4), we obtain the combined model function  $W'''(x, ssn, r)$  and for the velocity  $W''_I(x, ssn, r)$  with the double lognormal function (5). The corresponding median and mean relations for each solar-wind parameter, based on the values resulting from our analyses, are listed below. Their numerical values are the fit parameters from Table 2 and the exponents from Table 4.

- The magnetic field strength relations, depending on solar activity and solar distance, are:

$$B_{\text{med}}(ssn, r) = (0.0131 \text{ nT} \cdot ssn + 4.29 \text{ nT}) \cdot r^{-1.66}, \quad (12)$$

$$B_{\text{avg}}(ssn, r) = 1.0879 \cdot B_{\text{med}}(ssn, r). \quad (13)$$

- The proton velocity relations for the slow and fast components, depending on solar distance, are:

$$v_{\text{med}}^{\text{slow}}(r) = 363 \text{ km s}^{-1} \cdot r^{0.099}, \quad v_{\text{med}}^{\text{fast}}(r) = 483 \text{ km s}^{-1} \cdot r^{0.099}, \quad (14)$$

$$v_{\text{avg}}^{\text{slow}}(r) = 1.0101 \cdot v_{\text{med}}^{\text{slow}}(r), \quad v_{\text{avg}}^{\text{fast}}(r) = 1.023 \cdot v_{\text{med}}^{\text{fast}}(r). \quad (15)$$

The share of both components balanced with solar activity is found to be:

$$c(ssn) = -0.00180 \cdot ssn + 0.64. \quad (16)$$

- The derived relations of the proton density are:

$$n_{\text{med}}(ssn, r) = (0.0038 \text{ cm}^{-3} \cdot ssn + 4.50 \text{ cm}^{-3}) \cdot r^{-2.11}, \quad (17)$$

$$n_{\text{avg}}(ssn, r) = 1.305 \cdot n_{\text{med}}(ssn, r). \quad (18)$$

- The derived proton temperature relations are:

$$T_{\text{med}}(ssn, r) = (197 \text{ K} \cdot ssn + 57\,300 \text{ K}) \cdot r^{-1.10}, \quad (19)$$

$$T_{\text{avg}}(ssn, r) = 1.654 \cdot T_{\text{med}}(ssn, r). \quad (20)$$

These relations average over seasonal variations because they are based on yearly data. The OMNI data are time-shifted to the nose of the Earth's bow shock; this leads to yearly solar distance variations of  $\pm 1.67\%$  as it orbits the Sun. The resulting maximal solar-wind parameter variation amplitudes over the year can thus be derived from the derived power law exponents. They are estimated to be smaller than 4% as seen in Table 4. Bruno et al. (1986) and Balogh et al. (1999) have pointed out that the solar-wind parameters vary with latitudinal separation from the heliospheric current sheet. Its position in heliographic latitude is highly variable around the solar equator (Schwenn 1990) and, furthermore, the Earth's orbit varies over the course of the year by  $\pm 7.2^\circ$  in latitude. Since this latitudinal separation is highly variable and requires significant effort to calculate for an extended time series, we have ignored this aspect in this analysis.



**Fig. 10.** PSP's solar distance during its mission time (top). Consecutive Venus flybys bring its perihelia nearer to the Sun. Actual and predicted SSN (bottom), that is, SIDC 13-month smoothed monthly actual SSN, SIDC Standard Curves Kalman filter prediction and SWPC prediction with their corresponding expected ranges (shaded areas). The SSN from previous cycle 24, shifted by 11 yr, is plotted together with two alternative trends of half and twice its amplitude.

## 6. Model extrapolation to PSP orbit

To estimate PSP's solar-wind environment during its mission time for its orbital positions, predictions of the SSN during the mission are incorporated into the empirical solar-wind model, derived in the previous sections, and extrapolations down to the PSP perihelion region are performed.

PSP is planned to launch in mid 2018. With its first Venus flyby it will swing into Venus' orbital plane, reaching a first perihelion with a distance of 0.16 au just 93 days after launch, in November 2018. Seven additional Venus flybys allow the perihelion distance to be reduced to a minimum of  $9.86 R_\odot$ . This distance will be reached with the 22nd perihelion in December 2024 (Fox et al. 2016), as plotted in the top panel of Fig. 10.

We extrapolate the derived empirical solar-wind models to PSP's orbital distance range and compare the results with those from the existing models shown in Fig. 11. The model and its extrapolation are visualized for the SSN range between solar minimum and maximum ( $0 \leq ssn \leq 200$ ), indicated by the shaded regions in the figure. The magnetic field strength is found to increase from median values of about 43 nT at 0.25 au to 715 nT at 0.046 au for a SSN of zero. Taking a SSN of 200 increases the values to 69 nT and 1152 nT. Our extrapolation results are slightly flatter than those derived from the analytical magnetic field model by Banaszkiewicz et al. (1998), who constructed an analytic dipole plus quadrupole plus current sheet (DQCS) model for solar minimum. We note that one cannot easily compare the absolute values of our study with the values obtained by Banaszkiewicz et al. (1998) because the DQCS model assumes solar wind originating from coronal holes at higher heliographic latitudes only, neglecting the slow solar-wind belt. We suggest that the difference in slope is due to the previously mentioned (Sect. 4) changing shape of the frequency distribution with heliocentric distance, which for smaller distances deviates more from the model's lognormal distribution. The average velocity is found to decrease from  $340 \text{ km s}^{-1}$  at 0.25 au to about  $290 \text{ km s}^{-1}$  at 0.046 au 3 yr after a SSN of zero occurred, whereas using a SSN of 200 it decreases from  $390 \text{ km s}^{-1}$



**Fig. 11.** Radial extrapolation of the solar-wind parameters to the PSP orbit region. The median values from the models, obtained from Helios and OMNI measurements, are extrapolated to the PSP region for SSN values between solar minimum and maximum, that is,  $0 \leq \text{ssn} \leq 200$ . The lower edges of the shaded areas correspond to solar minimum, the upper edges to solar maximum. Also plotted are the radial dependencies of the analytic DQCS magnetic field model for solar minimum from Banaszkiewicz et al. (1998), the slow wind velocity models from Sheeley et al. (1997) and Wang et al. (2000), the density model from Leblanc et al. (1998) and the range of temperature measurements from Billings (1959) and Liebenberg et al. (1975).

to  $330 \text{ km s}^{-1}$ . Comparing the results with the measurements by Sheeley et al. (1997) and Wang et al. (2000) shows an overestimation in our extrapolated slow solar-wind velocity values for distances below approximately  $20 R_\odot$ . They used LASCO coronagraph observations to track moving coronal features (blobs) in the distance range  $2\text{--}30 R_\odot$  to determine speed profiles and sources of the slow solar wind and they derived temperature and sonic point values for slow solar wind with the isothermal expansion model from Parker (1958). Therefore, it generally can be expected that PSP will encounter a slower solar-wind environment close to the Sun than our model estimates. Thus PSP will measure solar-wind acceleration processes (McComas et al. 2008), maybe even still at  $30 R_\odot$  as the study by Sheeley et al. (1997) suggests. The proton density increases from about  $84 \text{ cm}^{-3}$  at  $0.25 \text{ au}$  to about  $3018 \text{ cm}^{-3}$  at  $0.046 \text{ au}$  6 yr after a SSN of zero occurred. Being almost independent of the SSN, the values for a SSN of 200 are only 17% larger. The results are in good agreement with those of Leblanc et al. (1998), who derived an electron density model from type III radio burst observations. Their model shows that the density distance dependency scales with  $r^{-2}$  and steepens just below  $10 R_\odot$  with  $r^{-6}$ . We assumed a solar-wind helium abundance of 5% to convert these electron densities to proton densities. The extrapolated proton temperature increases from about  $260\,000 \text{ K}$  at  $0.25 \text{ au}$  to about  $1\,690\,000 \text{ K}$  at  $0.046 \text{ au}$  3 yr after a SSN of zero occurred and from  $440\,000 \text{ K}$  to  $2\,860\,000 \text{ K}$  for a SSN of 200. Knowing

that near-Sun coronal temperatures are in the range of 2–3 MK (Billings 1959; Liebenberg et al. 1975), the model overestimates the extrapolated temperatures at the PSP perihelion distance.

Aside from the solar distance, the derived solar-wind parameter models depend on the SSN. Short-term predictions of the SSN can be used for the solar-wind predictions of PSP's early perihelia and also for refining the solar-wind predictions during PSP's mission. Several sources are available for SSN short-term predictions. The SIDC provides 12-month SSN forecasts<sup>6</sup> obtained from different methods (e.g., Kalman filter Standard Curve method). The SSN prediction of NOAA's Space Weather Prediction Center (SWPC) for the time period until the end of 2019 follows a consensus of the Solar Cycle 24 Prediction Panel<sup>7</sup>. The SSN for PSP's first perihelion will be small – certainly below 20 – whereas the SSN during the closest perihelia, which will commence at the end of 2024 at the likely maximum phase of cycle 25, cannot be predicted at this time. However, Hathaway & Upton (2016) found indications that the next solar cycle will be similar in size to the current cycle 24. Therefore we simply assume a pattern similar to the last cycle for the prediction of the next solar cycle and thus shift the last cycle by 11 yr. Additionally, we consider as possible alternatives

<sup>6</sup> <http://sidc.be/silso/forecasts>

<sup>7</sup> <http://www.swpc.noaa.gov/products/solar-cycle-progression>



**Fig. 12.** Estimated solar-wind parameter medians (black lines) and their error bands (gray) during 12 days in 2018 with PSP's first perihelion at about 0.16 au. For the velocity the combined median is calculated and also the SSN-independent slow and fast parts are plotted (dotted lines).



**Fig. 13.** Estimated solar-wind parameter medians (black lines) and their error bands (gray) during 12 days in 2024 with PSP's 22nd (the first closest) perihelion at 0.0459 au. For the velocity the combined median is calculated and also the SSN-independent slow and fast parts are plotted (dotted lines).

SSN patterns of half and twice its amplitude as shown in the bottom panel of Fig. 10.

Implementing the SSN predictions for the PSP mission time and the orbital trajectory data, we can infer which solar-wind parameter magnitudes can be expected. Figures 12 and 13 show the median values (7) of the considered different solar-wind parameters for 12-day periods, comprising the first perihelion in November 2018 and the first closest perihelion in December 2024. In the beginning of the mission median values of about 87 nT, 340 km s<sup>-1</sup>, 214 cm<sup>-3</sup>, and 503 000 K are estimated to be measured at 0.16 au, increasing to about 943 nT, 290 km s<sup>-1</sup>, 2951 cm<sup>-3</sup>, and 1 930 000 K during the first closest approach at 0.046 au. Monthly SSNs – shifted by the time lags specific to the solar-wind parameters – are used in the calculation of the solar-wind predictions. These SSNs are either actual smoothed values from the SIDC with their reported standard deviations, short-term predictions from the SWPC with



**Fig. 14.** Frequency distributions of the four solar-wind parameters (same as in Figs. 1 and 2) and those estimated with the solar-wind model for PSP's 1st and 22nd (first closest) perihelion. In these figures the frequencies of both extrapolated curves are scaled for visibility to the same height as the 1 au distribution.

their expected ranges, or actual smoothed values from the SIDC shifted by 11 yr with half/twice their values as uncertainties. The error bands given in both figures, calculated from error propagation, include these SSN ranges and the derived fit parameter errors.

Finally the estimated solar-wind environment can be derived from the function  $W'''(x, ssn, r)$ . The estimated frequency distributions of the four solar-wind parameters at PSP's 1st and 22nd (first closest) perihelion are plotted in Fig. 14. Again, we point out that the velocity and temperature distributions for the 22nd perihelion are only upper limits and the actual values to be encountered by PSP are expected to be smaller.

## 7. Discussion and summary

The scientific objective of this study, being part of the CGAUSS project – the German contribution to the WISPR instrument – is to model the solar-wind environment for the PSP mission to be launched mid 2018. For this purpose we started the development of the empirical solar-wind environment model for the near-ecliptic PSP orbit. We derived lognormal representations of the in situ near-Earth solar-wind data collected in the OMNI database, using the frequency distributions of the key solar-wind parameters, magnetic field strength, proton velocity, density, and temperature. Throughout the different analyses in our study, the velocity's frequency distribution is treated as a composition of a slow and a fast wind distribution. Each velocity part is fitted with a lognormal function, which allows for the overlap of both velocity ranges. The OMNI multi-spacecraft solar-wind data is intercalibrated and covers almost five solar cycles. It thus represents solar wind gathered at different phases of solar activity in the ecliptic plane. In the next step we investigated the yearly variation of the solar-wind distribution functions along with the SSN over 53 yr and derived linear dependencies of the solar-wind parameters with the SSN. The radial dependencies of the solar-wind distribution functions were then analyzed, using Helios 1 and 2 data for the distance range 0.29–0.98 au in bins of 0.01 au, deriving power law fit functions that were used to scale the previously calculated SSN-dependent 1 au distribution fit functions to the PSP orbit,

taking into account SSN predictions for the years 2018–2025, encompassing the prime mission up to the closest approach of  $9.86 R_{\odot}$ . The reason for performing the analysis this way is based on the fact that the OMNI solar-wind database is much larger than the Helios database.

For determining solar-activity- and solar-distance-dependent relations for the median and mean solar-wind values, we could have used the simpler approach of combining the radial dependence of averaged Helios data with averaged 1 au OMNI data scaled with the SSN. It is expected that the results of a simpler analysis would have similar distance scaling results, as can be inferred from the exponents in Tables 3 and 4. However, in our study we are not only interested in averages but rather in bulk distributions, that is, the whole range of values that might occur. For the determination of the frequency distributions the use of the more complex fit model is important, because the distance between median and mean values determines the width of the lognormal distributions.

It is clear that the calculated distribution functions only represent first-order estimates of the real solar wind to be encountered by PSP. The solar-wind environment to be encountered will depend at times of PSP on the structure of the solar corona and underlying photospheric magnetic field and on the evolution and interaction of individual solar-wind streams and superimposed CMEs and shocks. However, the derived results are in good agreement with existing studies about near-Sun solar-wind magnetic field strengths and densities as shown in Sect. 6. The extrapolation results of the velocity and the temperature differ from the direct measurements seen in existing studies. This suggests that below about  $20 R_{\odot}$  PSP may dive into the region where the acceleration and heating of the solar wind is expected to occur (see Fig. 11). The near-Sun solar-wind velocity at PSP perihelion is also expected to be slower than our model estimates, because the solar wind is assumed to be accelerated up to the height of the Alfvénic critical surface, which is predicted to lie on average around  $17 R_{\odot}$  (e.g., Sittler & Guhathakurta 1999; Exarhos & Moussas 2000), scaling with solar activity within a range of between  $15 R_{\odot}$  at solar minimum and  $30 R_{\odot}$  at solar maximum (Katsikas et al. 2010; Goelzer et al. 2014).

We have not specifically investigated the occurrences of extreme solar-wind parameters caused by CMEs or enhanced values due to stream interaction or co-rotating interaction regions. The Helios solar-wind measurements plotted over radial distance in Fig. 7 show several extreme values far above the usual solar-wind velocities, which are associated with individual CMEs. The results by Sachdeva et al. (2017) indicate that due to solar-wind drag, the speeds of fast CMEs will commonly slow down substantially from early distances of a few solar radii. Therefore, it is expected that PSP will encounter CMEs with much higher speeds than those observed during the Helios mission. Also, the magnetic field, density and temperature values are expected to be much larger than in the average solar wind in individual fast-shock-associated CME events. PSP will thus also substantially improve our understanding of the near-Sun evolution of CMEs and their expansion with radial distance.

With the resulting CGAUSS empirical solar-wind model for PSP, the following main results for the bulk solar-wind parameters and estimations for their median values at PSP's first perihelion in 2018 at a solar distance of 0.16 au and at PSP's closest perihelia beginning in 2024 at 0.046 au ( $9.86 R_{\odot}$ ) are obtained:

- The dependency of the magnetic field strength on solar activity and radial distance appears to be valid above  $20 R_{\odot}$ ,

however near PSP's closest perihelia, the actual values might be found to be slightly higher.

- The estimated magnetic field strength median values obtained from relation (12) for PSP's 1st and 22nd perihelion are 87 nT and 943 nT.
- The radial dependencies of the proton velocity median values for slow and fast solar wind (14) appear to be valid above about  $20 R_{\odot}$  solar distance; below they overestimate the actual solar-wind velocities obtained from remote measurements. The share of their frequency distributions to the overall solar-wind velocity distribution (5) depends on solar activity with their balance relation (16). Thus, at solar minimum, with a SSN of around zero, the slow-wind component contributes about 64% and drops to 28% during solar maximum conditions with a SSN around 200.
- The calculated median velocity values for PSP's 1st and 22nd perihelion are  $340 \text{ km s}^{-1}$  and  $290 \text{ km s}^{-1}$ .
- The proton density relation appears to be valid throughout the full PSP orbital distance range, even down to about  $8 R_{\odot}$ .
- The estimated density median values obtained from relation (17) for PSP's 1st and 22nd perihelion are  $214 \text{ cm}^{-3}$  and  $2951 \text{ cm}^{-3}$ .
- The derived correlation function for the proton temperature appears to provide overly high temperature values around PSP's closest perihelion in comparison to coronal measurements.
- The estimated temperature median values obtained from relation (19) for PSP's 1st and 22nd perihelion are  $503\,000 \text{ K}$  and  $1\,930\,000 \text{ K}$ .

The results of the modeled solar-wind environment will be useful to help optimize the WISPR and in situ instrument science plannings and PSP mission operations. This also applies for the Heliospheric Imager (SoloHI; Howard et al. 2013) and the in situ instruments on board the Solar Orbiter spacecraft.

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## 7 Summary and outlook

### 7.1 Summary

In order to derive the link from near-Earth solar wind properties to its  $K_p$  impact, 35 years of solar wind high-resolution measurements were correlated with the  $K_p$  index and empirical relations were derived that enable the prediction of  $K_p$  from the solar wind velocity and the IMF z-component. Early predictions of CME and ambient stream magnetic fields obtained from near-Sun observations currently either are not existent or come with high uncertainties. Thus, to provide  $K_p$  predictions, I derived relations exclusively from their arrival velocity.

I built an empirical solar wind model for the inner heliosphere in the ecliptic which accounts for the variations of the solar activity cycle and for solar distance. In order to obtain empirical estimates of the solar wind environment PSP is to encounter, this solar wind model is extrapolated down to PSP's planned near-Sun perihelia and is built to consider the expected solar activity during PSP's mission.

### 7.2 Outlook

Using the SSN prediction, the derived solar wind model allows the forecast and extrapolation of the solar wind environment occurring during the PSP mission's near-Sun encounters from end of 2018 on. The anticipated Parker Solar Probe data (near-Sun measurements) will reveal how far the solar wind estimates really are from reality and thus is able to locate the outer boundary of the solar wind acceleration region.

DSCOVR data (advantages over ACE? gain?)

other possible spacecraft missions that would be beneficial for space weather forecasting: sub-L1 for earlier in-situ CME magnitude warning (cite?) and L5 for early CME velocity and arrival warning ([Vourlidas 2015](#))



# A Appendix

## A.1 Solar surface differential rotation

The solar differential rotation is visible on the surface and was first discovered from sunspot observations by Scheiner (1630). (double)

Bartels (1934) set the synodic solar rotation period to 27 days for the definition of his solar rotation number. The Bartels' Rotation Number counts the solar rotations starting with 8 February 1832.

Carrington solar rotation period of 27.2753 days (Where Carrington Rotation Number is based upon, starting with November 9, 1853; Wikipedia...)

Solar surface rotation period at 16° latitude:  
sidereal: 25.38 d (of 609.12 h Sun Fact Sheet...), synodic: 27.2753 d (derived)

rotation axis tilt (see next section)

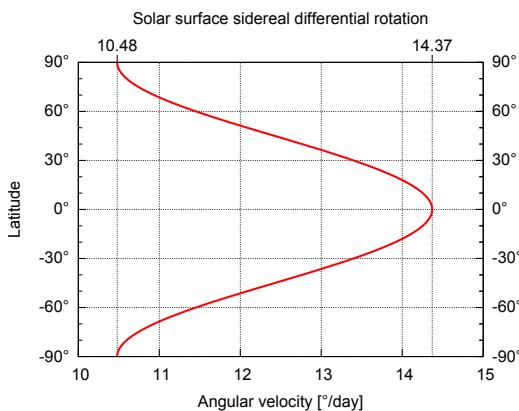
The Sun's inner thermal convective circulation results in a differential rotation caused by transport of angular momentum away from the rotation axis.

The Sun's sidereal differential angular velocity best-fitting function with values as stated in (Sun Fact Sheet...)<sup>1</sup> is

$$\omega_{\odot}(\theta) = \omega_{\text{eq}} + B \sin^2(\theta) + C \sin^4(\theta) \quad (\text{A.1})$$

with the heliolatitude  $\theta$ , the equatorial angular velocity  $\omega_{\text{eq}} = 14.37^\circ \text{d}^{-1}$ , the coefficients  $B = -2.33^\circ \text{d}^{-1}$ , and  $C = -1.56^\circ \text{d}^{-1}$  (see Figure A.1).

see Figure A.1



**Figure A.1** Diagram of the sidereal solar surface differential rotation. It shows the angular velocity for different latitudes. remove sides...

Thus, the solar equatorial rotation period (sidereal) is

$$\begin{aligned} T_{\odot}^{\text{eq}} &= 360^\circ / A \\ &= 25.05 \text{ d} \end{aligned} \quad (\text{A.2})$$

and the synodic period is

$$\begin{aligned} T_{\odot}^{\text{eq,syn}} &= 1 / (1/T_{\odot}^{\text{eq}} - 1/T_{\text{Earth}}) \\ &= 26.90 \text{ d} \end{aligned} \quad (\text{A.3})$$

with the Earth's orbital rotation period  $T_{\text{Earth}} = 365.25 \text{ d}$  (1/100 Julian century).

<sup>1</sup>NASA's Sun Fact Sheet (<http://nssdc.gsfc.nasa.gov/planetary/factsheet/sunfact.html>, accessed 2016-08-19).

Solar surface rotation period at equator  
sidereal: 25.05 d (Sun Fact Sheet...), synodic: 26.90 d (derived)  
Solar surface rotation period at poles:  
sidereal: 34.35 d (diff. rot. formula), synodic: 37.92 d (derived)  
are listed in [Table A.1](#).

**Table A.1** Solar surface rotation periods for equator,  $\pm 16^\circ$  latitude and poles (sidereal and synodic).

	Equator [d]	$\pm 16^\circ$ latitude [d]	Poles [d]
Sidereal	25.05	25.38	34.35
Synodic	26.90	27.2753 <sup>a</sup>	37.92

<sup>a</sup>Carrington solar rotation period

The meridional circulation is the proposed equatorial updrift and polar downdrift - a result of Reynolds stress and convective transport ([cite?](#)).

## A.2 Electromagnetism

Electromagnetism is one of the four fundamental forces at the common level of energy. In all situations examined in this thesis (solar wind plasma, magnetosphere) it is by far the strongest force and the others can be neglected.

The Lorentz force law defines the electric and magnetic field vectors  $\mathbf{E}$  and  $\mathbf{B}$ :

$$\mathbf{F} = q(\mathbf{E} + \mathbf{v} \times \mathbf{B}), \quad (\text{A.4})$$

with the charge  $q$  the force is acting on and its velocity  $v$ .

Full Ohm's law:

$$\mathbf{j} = \sigma(\mathbf{E} + \mathbf{v} \times \mathbf{B}) \iff \mathbf{E} = -\mathbf{v} \times \mathbf{B} + \frac{\mathbf{j}}{\sigma}. \quad (\text{A.5})$$

Solar wind is often approximated as an ideal MHD plasma having an infinite conductivity ( $\sigma = \infty$ ). Thus electrical fields do not generate electrical currents  $j$ :

$$\mathbf{E} = -\mathbf{v} \times \mathbf{B}. \quad (\text{A.6})$$

The Maxwell equations in differential notation:

$$\operatorname{div} \mathbf{B} = 0 \quad (\text{A.7})$$

$$\operatorname{div} \mathbf{D} = \rho \quad (\text{A.8})$$

$$\operatorname{rot} \mathbf{H} = \mathbf{j} + \dot{\mathbf{D}} \quad (\text{A.9})$$

$$\operatorname{rot} \mathbf{E} = -\dot{\mathbf{B}} \quad (\text{A.10})$$

With the magnetic flux density  $\mathbf{B}$  (aka magnetic field), the electric displacement field  $\mathbf{D}$ , the charge density  $\rho$ , the magnetic field  $\mathbf{H}$ , the electric field  $\mathbf{E}$  and the current density  $\mathbf{j}$ .

## A.3 Plasma beta

The thermal pressure of a plasma is defined as  $p = nk_B T$ , with the number density  $n$  and the Boltzmann constant  $k_B$ . However, according to magnetohydrodynamics (MHD), the magnetic energy density  $w_{\text{mag}} = \frac{B^2}{2\mu_0}$ , with the magnetic field strength  $B$  and the permeability constant  $\mu_0$ , behaves like an additional pressure that adds to the thermal pressure of a plasma ([Kivelson & Russell 1995](#), p. 50). The ratio of the thermal pressure to

the magnetic pressure determines the behavior of the plasma. If the thermal pressure dominates the magnetic pressure (warm plasma), the plasma movements transport the magnetic field, else the plasma movements are guided by the magnetic field lines (cold plasma). This ratio is called plasma beta:

$$\beta = \frac{p}{p_{\text{mag}}}, \quad (\text{A.11})$$

$$= \frac{2\mu_0 n k_B T}{B^2}. \quad (\text{A.12})$$

The plasma at the photosphere has typical beta values around 14 and that of the low corona around 0.2 (Gary 2001). Further up, beta raises again and the region where it equals 1 is defined as the source surface for the solar wind (Schatten et al. 1969). This surface is typically located at about 1.2–2.5  $R_\odot$  (Gary 2001), more cites...

The solar wind usually has plasma beta values higher than 1 – it carries the solar magnetic field away into the heliosphere. Together with the solar rotation, this effect creates the spiral form of the interplanetary magnetic field (Parker spiral). Yet in some solar wind structures, such as magnetic clouds,  $\beta \ll 1$  and thus the magnetic field can still contain the plasma.

Savani2017:

*“By using statistical analysis, Riley and Richardson [2013] suggest that a distinct magnetic flux rope object may be contained within an overall propagating CME structure and that the plasma beta is a good predictor variable. Savani et al. [2013b] used simulations to confirm that such a scenario of a distinct coherent core obstacle can occur within the CME structure and that a transition in plasma beta can be observed.”*

## A.4 Alfvén velocity

The incompressible wave mode within MHD plasmas, the shear Alfvén wave, consists of periodic disturbances in the magnetic field orthogonal to its direction (Alfvén 1942). Alfvén waves are prevalent in open coronal regions and therefore occur in fast solar wind (Cranmer & van Ballegooijen 2005). Their propagation velocity is an important parameter to characterize a plasma. In an ideal incompressible MHD plasma (viscosity  $\mu = 0$  and electrical conductivity  $\sigma = \infty$ ) the kinetic and magnetic energy density are of equal value (Kivelson & Russell 1995, p. 51):

$$w_{\text{kin}} = w_{\text{mag}} \quad (\text{A.13})$$

$$\frac{\rho v^2}{2} = \frac{B^2}{2\mu_0}$$

with the permeability constant  $\mu_0$  and the total mass density  $\rho$  of the charged plasma particles. Thus, the Alfvén velocity can be calculated from

$$v_A = \frac{|B|}{\sqrt{\mu_0 \rho}}. \quad (\text{A.14})$$

The wave's phase velocity is

$$v_{\text{ph}} = v_A \cos(\theta) \quad (\text{A.15})$$

with  $\theta$  as the angle between wave propagation direction and magnetic field line, that is, Alfvén waves travel along magnetic field lines. They consist of periodic disturbances in the magnetic field, the electric field, the plasma velocity, and the current density. Plasma density, pressure and magnetic field magnitude are not affected by them. Additionally, there exist two types of compressional wave modes within MHD plasmas, the fast-mode wave and the slow-mode wave. The phase speeds of the three MHD waves meet  $v_{\text{fast}} \geq v_A \geq v_{\text{slow}}$  (Kivelson & Russell 1995, p. 52). Within solar wind at 1 au, the typical frequency of Alfvén waves is 1–4 per hour and their average velocity is  $v_A = 56.8 \text{ km s}^{-1}$  (Veselovsky et al. 2010).

Alfvén critical surface...  
sonic critical surfaces...

shocks & MHD waves (Alfvén waves)  
 Alfvén velocity ca. 40 km/s (Kivelson1995)  
 Alfvén velocity 56.8 km/s (Veselovsky2009)

sonic and Alfvénic critical point positions (see Sittler & Guhathakurta (1999))  
 sonic point and slow solar wind origin (Sheeley et al. 1997)

## A.5 Sun-Earth orbit geometry

orbit defines ecliptic

Earth orbit parameters (cite?):  
 semimajor axis:  $a = 1.000001018$  au  
 eccentricity:  $e = 0.0167086$  au  
 distance at perihelion: (formula cite?, accuracy?)

$$\begin{aligned} r_p &= a(1 - e) \\ &= 0.98329 \text{ au} \end{aligned} \tag{A.16}$$

distance at aphelion:

$$\begin{aligned} r_p &= a(1 + e) \\ &= 1.0167 \text{ au} \end{aligned} \tag{A.17}$$

for calculation of heliospheric distance see HORIZONS Web-Interface at <http://ssd.jpl.nasa.gov/horizons.cgi>  
 perihelion/aphelion times...

### Solar distance

Sun-Earth distance over the course of the year.

In the year 2017 Earth's perihelion was on 5 January with a distance of  $-1.67\%$  from 1 au (Horizons On-Line Ephemeris System<sup>2</sup>, Solar System Dynamics Group, Jet Propulsion Laboratory).  
 The cosine approximation

$$r_E(t) = 1 - 0.0167 \cdot \cos\left(2\pi\left(t - 2017 - \frac{5}{365}\right)\right), \tag{A.18}$$

with  $t$  in years, suffices for our accuracy requirements.

seasonal variation function:  
 $X_{\text{avg}}(t) = a r_E(t)^b$

### Solar rotation axis tilt to the ecliptic

The inclination of the solar equator to the ecliptic (tilt/obliquity) is  $i_\odot = 7.25^\circ$  (U.S. Nautical Almanac Office 2015).

the rotation axis is tilted from the ecliptic normal

Viewed from Earth the projected solar rotation axis tilt angle varies as the Earth is moving on its orbit.

At the time XX the angle is zero.

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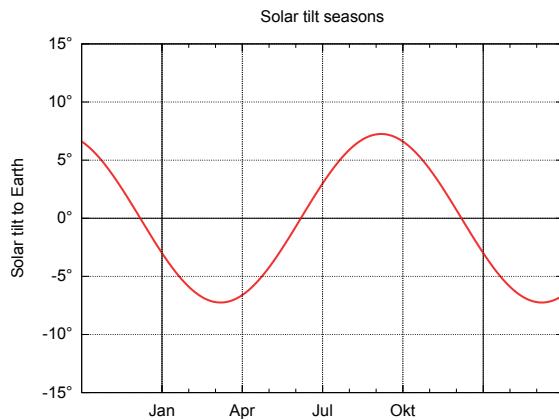
<sup>2</sup><http://ssd.jpl.nasa.gov/horizons.cgi>

The projected tilt angle to Earth over the year is

Hapgood (1992):

$$\omega = 73.67 + 0.013\,958 * (\text{today} - 1850.0) \quad (\text{A.19})$$

solar tilt over the year, see Figure A.2



**Figure A.2** Projected solar tilt angle over the year as viewed from Earth. remove sides...

### Earth rotation axis tilt to the ecliptic

## A.6 GSE, GSM, and HGI coordinate systems

Coordinate systems used in this thesis:

GSE - Geocentric Solar Ecliptic

GSM - Geocentric Solar Magnetospheric

HGI - Heliographic Inertial

refer to Hapgood (1992) for GSE and GSM

figures for GSE and GSM

### Geocentric Solar Ecliptic

The Geocentric Solar Ecliptic (GSE) coordinates are

GSE - Geocentric Solar Ecliptic

X = Earth-Sun Line

Z = Ecliptic North Pole

GSE coordinates are used in ACE solar wind data, etc.

see Jursa1985, p. 4-3

"the polar axis is the axis inclined 11.5° to the axis of rotation, intersecting the earth surface at the point 78.5°N, 291.0°E which defines the geomagnetic north pole. This was once at one time the axis of the best centered-dipole approximation to the field"

### Geocentric Solar Magnetospheric

GSM - Geocentric Solar Magnetospheric

X = Earth-Sun Line

Z = Projection of dipole axis on GSE YZ plane

GSM is defined with a time dependent dipole axis.

the dipole axis orientation changes over time; at 1995 the northern pole was at  $l = 288.59^\circ$  and  $b = 79.30^\circ$ ; more recent year (2015)?... cite?

## Heliographic Inertial

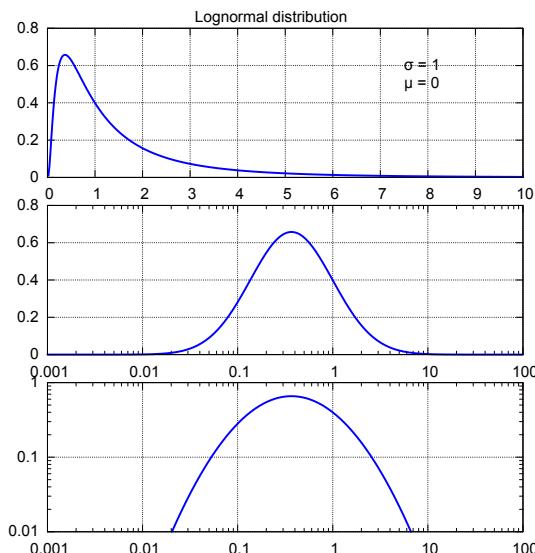
Heliographic Inertial (HGI) coordinates

HGI coordinates are Sun-centered with the z-axis directed along the solar rotation axis and directed northward of the solar equator. The solar equator plane is inclined  $7.25^\circ$  from the ecliptic.

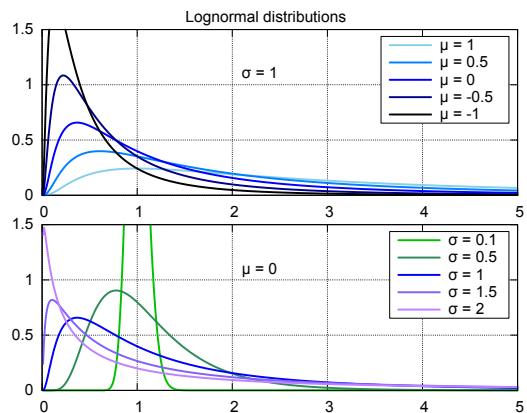
HGI coordinates; latitude range  $-7.25^\circ$  to  $-7.25^\circ$   
latitude variation (see Schwenn1990, p. 127)

## A.7 Lognormal distribution

This is a small summary about the lognormal probability distribution (Bronstein et al. 2000, p. 780). The lognormal distribution is the distribution of a random variable  $X$  if the logarithm of  $X$  conforms to a normal distribution. Its shape is highly asymmetric, however in a semi-log plot the Gaussian bell curve is recognizable (see the second panel of Figure A.3). Its probability density function is



**Figure A.3** The lognormal probability density function ( $\sigma = 1, \mu = 0$ ) plotted in a linear, semi-log and log-log way. remove borders...



**Figure A.4** Five lognormal distributions plotted with fixed  $\sigma$  (top) and fixed  $\mu$  (bottom). remove borders...

$$f(x) = \frac{1}{\sigma \sqrt{2\pi} x} e^{-\frac{(\ln x - \mu)^2}{2\sigma^2}} \quad (\text{A.20})$$

with the location ( $\mu$ ) and the shape parameter ( $\sigma$ ). Changes in  $\mu$  affect both the horizontal and vertical scaling of the function, whereas  $\sigma$  has an influence on its shape (see Figure A.4).

Because it is a probability distribution, its area is normalized

$$\int_0^\infty f(x) dx = 1. \quad (\text{A.21})$$

For a lognormally distributed random variable the geometric moments mean, standard deviation and variance are:

$$\begin{aligned} \mu_g &= e^\mu, \\ \sigma_g &= e^\sigma, \\ var_g &= e^{\sigma^2} \quad (!). \end{aligned}$$

Its arithmetic moments are:

$$\begin{aligned}\mu_a &= e^{\mu + \frac{\sigma^2}{2}}, \\ \sigma_a &= e^{\mu + \frac{\sigma^2}{2}} (e^{\sigma^2} - 1), \\ var_a &= \sigma_a^2.\end{aligned}$$

Other useful characteristics are the median and the mode

$$\begin{aligned}x_{\text{median}} &= e^\mu, \\ x_{\text{mode}} &= e^{\mu - \sigma^2}.\end{aligned}$$

Note that for the lognormal distribution its median is equal to its geometric mean.

Applications of lognormal distributions...

Most natural quantities which can only be positive are lognormally distributed. e.g. animal body sizes?, animal life expectancies, financial stock prices...; income distributions.

## A.8 Astronomical constants

Astronomical unit: 1 au = 149 597 870 700 m ([U.S. Nautical Almanac Office 2015](#))

Solar mass:  $M_\odot = 1.9884(2) \times 10^{30}$  kg ([U.S. Nautical Almanac Office 2015](#))

Nominal solar radius (photosphere):  $R_\odot = 695 700$  km ([Mamajek et al. 2015](#))

Solar rotation axis tilt:  $i_\odot = 7.25^\circ$  ([U.S. Nautical Almanac Office 2015](#))

Solar surface rotation period at equator, sidereal: 25.05 d (Sun Fact Sheet...)

Nominal solar effective temperature (photosphere):  $T_{\text{eff}\odot} = 5772$  K ([Mamajek et al. 2015](#))

Sun escape velocity:  $v_{\text{esc}} = 617.6$  km/s (Sun Fact Sheet...)

NASA maintains the Planetary Fact Sheets<sup>3</sup> online.

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<sup>3</sup>NASA Planetary Fact Sheets website: <https://nssdc.gsfc.nasa.gov/planetary/planetfact.html>



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- Kp data  
- SSN data

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CGAUSS: Solar wind model and extrapolation  
HELCATS: Minimum variance analyses of magnetic clouds  
OPTIMAP: Solar wind ACE time series

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## Curriculum vitae

Malte Venzmer, born 5 February 1984 in Bremerhaven, Lower Saxony, Germany, finished secondary school (Abitur) at the Gymnasium Ganderkesee in 2003. Subsequently he performed his civilian service working in a care home for disabled persons in Westerland, Sylt.

In 2004 he began to study physics at the University of Konstanz. During his studies, in summer 2008, he did an internship at the Ångström Laboratory, Uppsala University, Sweden, working in the field of materials physics. In 2010/2011 he wrote his diploma thesis in the *Radio- and X-ray astronomy* group of Dr. Jürgen Kerp at the Argelander-Institute for Astronomy, University of Bonn. There, following his diploma in physics, he continued working under a scholarship until end of 2011.

In 2012 he started as a PhD candidate in the *Solar, heliospheric and space weather research* group of Dr. Volker Bothmer at the Institute for Astrophysics, University of Göttingen. During his doctoral studies he worked as a research assistant for several national and international projects. Most of his main results are described in this very thesis.

## notes

Update figures before printing!

Check if footnotes are on same page!

Fragen:

- references vol # fett drucken?
- references journal kursiv drucken?
- maybe style references different. like in space science reviews?
- page # oben lassen?
- A5 format?

nice phrases:

...this leads to the question wh... . The answer to this question is developed in the course of the next section.

write basics without third persons view “we”

define if using astronomical symbols...

...and use style consistently

Letzte Änderungen:

Genehmigungen für die Abb. besorgen!!

check for topic sentences!

Am Anfang der Arbeit den folgenden logischen Aufbau der Kapitel erlaeutern.

Am Anfang jedes Kapitels den folgenden logischen Aufbau der Abschnitte erlaeutern.

Beides auf Aenderungen ueberpruefen.

print figures to check look of colors

adjust figure width to textwidth...

complete pdfinfo text...

Abkuerzungen CME, usw. konsequent nutzen und beim ersten Auftauchen ausschreiben.

englische Kommasetzung beachten!

schauen, ob Kommata oder Punkte durch Semikolon ersetzt werden koennen

Pruefen, ob Ueberschriften den Inhalt des Kapitels gut beschreiben/gut zum Kapitel passen.

check if AE spelling consistently, search and replace: analyse, etc.

check for thin spaces in numbers with 5 digits and more

Header und footer fuer notes, seitenzahl, Thema und Kapitel einfuegen

check gnuplot plot text sizes and finish figures

check for changed links and update access date

sinnlosen Satz irgendwo einfügen

76 %, ACME irgendwo einfügen (ACME is good...)

comic strip aus thesis extrahieren

put movie in header or footer

create hyperref version including bibtex links

adjust link colors and remove boxes... (hyperref package options)

remove google-books links!

maybe adjust References like in Hathaway2010

LaTeX commands:

block commenting: strg+d

strg+shift+d

$$N_{mean_0} = \text{\texttt{e\&e\ e\ e\ e\ e\ e\ e}} \quad (A.1)$$

$$N_{median_0} = \text{\texttt{||| | | | | | |}} \quad (A.2)$$

Provides generic commands  $\circ$ ,  $^{\circ}\text{C}$ ,  $\%$ ,  $\mu$  and  $\Omega$  which work both in text and maths mode.

[Lem & Kandel \(1984\)](#), ([Lem & Kandel 1984](#)), “`citep[see][page 45]Lem1984`” => (see [Lem & Kandel 1984](#), page 45)

included persons via citations: Volker Bothmer, Rainer Schwenn, Eckhard Bosman, Russel Howard, Manuela Temmer, Angelos Vourlidas, Craig DeForest, Tim Howard, Rodney Viereck

Style rules:

the Sun’s/solar rotation period

Earth’s/terrestrial

the Sun, sunspots; near-Sun measurements, near-relativistic electrons, in situ measurements without hyphen?

64 s data; but 64-second data

Therefore, a comma is set after certain adverbs at the beginning of a sentence.

lognormal or log-normal, not log normal

A&A style:

Italics should never be used for units

italics should be avoided for the following: mathematical signs such as “d” (total differential), “e” (base of natural logarithm), “i” (imaginary unit), “pi”

Physical constants such as the speed of light, the Boltzmann constant, the Hubble constant and the solar mass are also set in regular italics.

$$\text{e}^x = 10 \text{ km s}^{-1} \quad (A.3)$$

$$\text{e}^x = 10 \text{ km}\cdot\text{s}^{-1} \quad (A.4)$$

$$\text{e}^x = 10 \text{ km/s} \quad (A.5)$$

$$k_{\text{B}} \quad (A.6)$$

Sample input: 20 000 km, 1 000 000 s, HD 174 638 1950–1985, p. 11–21, this – written on a computer – is now printed, this—written on a computer—is now printed, signal-to-noise ratio, early-type, metal-poor, non-relativistic  $-30 \text{ K}$ ,  $-5 \text{ }^{\circ}\text{C}$  Dr. h.c. Rockefeller-Smith and Prof. Dr. Mallory

Nützliche Beispiele für LaTeX-Kommandos:

Autobahn

X-Ray

X-Ray

Dr. Huber

100 000 V

100 000 V

The food—which was delicious—reminded me of home.

Red, white, and blue—these are the colors of the flag.

$U = 20\,000 \text{ V}$

solar radius:  $R_{\odot} = 696\,000 \text{ km}$

$10 \text{ }^{\circ}\text{C}$

$10 \text{ }^{\circ}\text{C}$

$10 \text{ }^{\circ}\text{C}$

Ein Winkel von  $10^{\circ}$ .

$$\frac{T_D}{[K]} = 21,8 \cdot \left( \frac{\Delta v_{FWHM}^{LMC}}{[\text{km s}^{-1}]} \right)^2 \quad (\text{A.7})$$

$$\frac{T_D}{[K]} = 21,8 \cdot \left( \frac{\Delta v_{FWHM}^{LMC}}{[\text{km s}^{-1}]} \right)^2 \quad (\text{A.8})$$

$$k = 3.56 \text{ e-}6 \quad (\text{A.9})$$

$$k = 3.56 \cdot 10^{-6} \quad (\text{A.10})$$

$$k = 3.56 \times 10^{-6} \quad (\text{A.11})$$

### 1. CM LMC Longitude

drei CM LMC Latitude

Volume in  $(^{\circ})^2 \text{m s}^{-1}$ . Summe aller Voxel des Objektes.

**Volume** in  $(^{\circ})^2 \text{m s}^{-1}$ . Summe aller Voxel des Objektes.

- Volume

sdgffs CM GSR Velocity

label

asdf Volume

drei CM LMC Latitude

<sup>1</sup> oder <sup>2</sup>

$$a^2 + b^2 = c^2 \quad (\text{A.12})$$

Formel zitieren: (Siehe Formel A.12, Formel (A.12), Seite 112)

package siunitx:  
 $20.5 \times 10^{-3} \text{ kg}$   
 $20.5 \times 10^{-3} \text{ kg}$   
 $20.5 \times 10^{-3} \text{ kg}$

width defines the width of the resulting box as seen from the outside (This means it can be **smaller** than the material inside the box. You can even set the width to 0pt so that the text inside the box will be typeset without influencing boxes). Besides the length expressions, you can also use semi-log plots:

log-lin: logarithmic scale on the y-axis, and a linear scale on the x-axis

lin-log: logarithmic scale on the x-axis, and a linear scale on the y-axis

the naming is output-input (y-x), the opposite order from (x, y)

my common language errors:  
hyphen for adjectives  
solar wind  
solar-wind parameters  
high-value tail

serial comma (it is an american thing and optional)  
a, b, c and d.  
a, b, c, and d.

<sup>1</sup>All listed websites were existent on 2008-09-15.

<sup>2</sup>All listed websites were existent on 2008-09-15.

comma if object is at beginning of sentence  
In the case of A, B is C.  
In order to A, B can C.

BE vs AE  
catalogue  
catalog

grey  
gray