

PHSX 671: Homework #9

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Problem 1

Electromagnetic radiation (i.e., blackbody radiation) at temperature T_i fills a thermally insulated cavity of volume V . How much work is done by the radiation if the cavity expands adiabatically ($\Delta S = 0$) to volume rV ? Express your answer in terms of T_i , V , r , c , and the Stefan-Boltzmann constant σ only. Hint: use the equation for the entropy of this system to determine how the initial and final temperatures are related.

Solution:

The adiabatic condition implies that the entropy before and after remains the same, and that $Q = 0$

$$S_i = S_f, \quad \Delta E = Q - W$$

And the entropy initially is $S_i = \frac{4}{3}aVT^3$. We have already derived mean energy, which saves me a few steps:

$$E = aVT^4$$

So now the problem is reduced to arranging these things to suitably encompass the effects of thermal expansion. I can start by finding an expression coupling rV and the temperature:

$$\begin{aligned} S_i &= S_f \\ \frac{4}{3}aVT_i^3 &= \frac{4}{3}a(rV)T_f^3 \\ T_i &= (r)^{1/3}T_f \end{aligned}$$

Using the first law of thermodynamics, I can write the final expression for work:

$$\begin{aligned} -W &= \Delta E \\ W &= aVT_i^4 - aVT_f^4 \\ W &= aVT_i^4 - aV \frac{T_i^4}{r^{1/3}} \end{aligned}$$

$$W = aVT_i^4(1 - r^{-1/3}), \quad a = \frac{4\sigma}{c} = \frac{4\pi^2 k_B^4}{60(\hbar c)^3}$$

Problem 2

A Dyson shell is a hypothetical megastructure that was originally described by Freeman Dyson as a uniform solid shell of matter around a star meant to completely encompass the star and thus capture its entire energy output. Such a structure would also provide an immense surface which many envision being used for habitation. If engineers wanted an average temperature on the inner surface of the shell to be 300 K, what must the radius of the shell be? The temperature and radius of the sun are 5800 K and 7×10^8 m, respectively.

Solution:

It is feasible to directly write out mean energy of the blackbody radiation emitted by the sun, and simply say that it is conserved, which allows me to express:

$$\bar{E}_{\text{sun}} = \bar{E}_{\text{dy}}$$

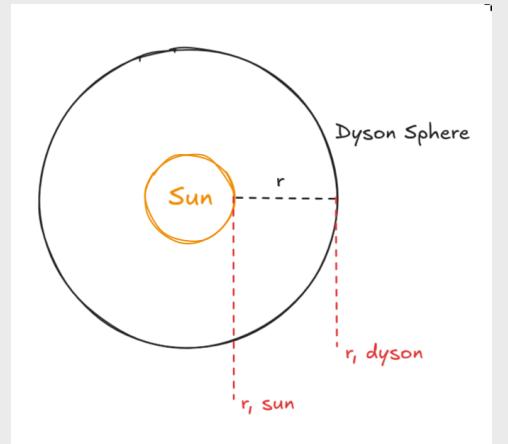
We have derived expressions for photon gasses in 3-D, and I get the sense that trying to hand-wavily take the sun to be a blackbody volume and work out average energy for the volume between the sun and dyson sphere would be quite dubious, but just to confirm:

$$aV_{\text{sun}}T_{\text{sun}}^4 = aV_{\text{space}}T_{\text{dy}}^4$$

$$r_{\text{sun}}^3 T_{\text{sun}}^4 = r_{\text{dy}}^3 T_{\text{dy}}^4$$

$$r_{\text{dy}} = \frac{r_{\text{sun}} T_{\text{sun}}^{4/3}}{T_{\text{dy}}^{4/3}}$$

$$r_{\text{dy}} \approx 3.63 \times 10^{10} \text{ m} = 0.243 \text{ au}$$



I would assume this to be horribly wrong, since the earth sits at 1 au distance, things would be far, far hotter if the earth was instead 4 times closer to the sun. Time for me to bite the bullet and derive an expression for surfaces instead of volumes:

$$\begin{aligned} \Gamma(k) &= 2 \left(\frac{1}{4} \frac{\pi k^2}{(\pi/L)^2} \right) dk = \frac{L^2 k^2}{2\pi} dk \\ D(k) dk &= \frac{L^2 k}{\pi} dk \\ k &= \frac{\omega}{c}, \quad dk = \frac{d\omega}{c} \\ D(\omega) d\omega &= \frac{A\omega}{\pi c^2} d\omega \end{aligned}$$

The partition function is then:

$$\begin{aligned} \ln Z &= - \int_0^\infty \ln(1 - e^{-\beta\hbar\omega}) \left(\frac{A\omega}{\pi c^2} \right) d\omega \\ &= - \frac{A}{\pi c^2 (\beta\hbar)^3} \int_0^\infty x \ln(1 - e^{-x}) dx, \quad \begin{aligned} x &= \beta\hbar\omega \\ d\omega &= \frac{dx}{\beta\hbar} \end{aligned} \\ &= \frac{\pi^4}{45} \frac{A}{\pi (c\beta\hbar)^2} = \frac{A\pi^3}{45(\beta\hbar c)^2} \end{aligned}$$

Mean energy is the derivative with respect to β , giving:

$$\bar{E} = \frac{2A\pi^3}{45(\hbar c)^2\beta^4}$$

and I'll assign the variable a to simplify:

$$\bar{E} = aAT^4, \quad a = \frac{2\pi^3 k_B^4}{45(\hbar c)^2}$$

Now it's pretty straightforward to relate energy emitted by the sun to energy at the surface of the dyson sphere.

$$r_{dy}^2 = \frac{r_{sun}^2 T_{sun}^4}{T_{dy}^4}$$

$$r_{dy} \approx 1.75 \text{ au}$$

Which seems much more reasonable to me.

Problem 3

The molar heat capacity at constant volume of an ideal gas of bosons of mass m at temperature T before the condensation temperature T_c is given by

$$C_V = 2R \left(\frac{T}{T_c} \right)^{3/2}$$

- (a) What is the internal energy per mole of this gas as a function of T for $T < T_c$?

Solution:

Probably the most straightforward way to get to internal energy is using $C_V = (\frac{\partial E}{\partial T})_V$:

$$\begin{aligned} E &= \int_0^T 2R \left(\frac{T}{T_c} \right)^{3/2} dT \\ &= \frac{4R}{5} \frac{T^{5/2}}{T_c^{3/2}} \end{aligned}$$

$$E = \frac{4R}{5} T \left(\frac{T}{T_c} \right)^{3/2}$$

- (b) What is the entropy per mole of this gas as a function of T for $T < T_c$?

Solution:

Using the Helmholtz Free Energy,

$$\int_0^S dS = \int_0^T \frac{C_V}{T} dT$$

So, entropy per mole s is:

$$\begin{aligned} s &= 2R \int_0^T \left(\frac{T}{T_c} \right)^{3/2} \frac{1}{T} dT \\ &= \frac{2R}{(T_c)^{3/2}} \int_0^T T^{1/2} dT \end{aligned}$$

$$s = \frac{4R}{3} \left(\frac{T}{T_c} \right)^{3/2}$$

Problem 4

What is the chemical potential of a three-dimensional photon gas?

Solution:

$$\mu = -k_B T \left(\frac{\partial \ln Z}{\partial N} \right)_{T,V}, \quad \ln Z = \frac{V\pi^2}{45(\beta\hbar c)^3}$$

$$\Rightarrow \mu = 0$$

The partial derivative of this is just zero, since there is no dependence on N , making μ equal zero. *In fact, nowhere in any of the equations surrounding photon gas is there an explicit dependence on N .* This is a bit strange, though, each photon has an energy, I would have expected it to take energy to add a photon to the system. This is definitely correct though, as a quick google search confirms. It seems that my notion about how chemical potential works was a bit incorrect, since yes, it is a measure of how energy changes when particles are added or removed from a system, but it is ONLY the act of adding a particle to the system that changes the energy. Particles can be given energy or momentum separately from adding them to the system. Yes we can introduce photons with kinetic energy, but the act of giving them momentum is separate from inserting them into the system. Relativity should hint that this is the case, since every particle has some mass-energy, but for the photon this is zero!

Problem 5

Electromagnetic radiation at temperature T has a volume V . How much work is done by the radiation if it expands isothermally to a final volume rV ?

Solution:

(Isothermal just refers to the system being held at a fixed temperature). The first law of thermodynamics lets me say

$$0 = Q - W$$

Or, equivalently,

$$0 = TdS - PdV$$

We've already found that the pressure of a photon gas is $P = \frac{1}{3} \frac{\bar{E}}{V}$, so it's pretty straightforward to just integrate that:

$$W = \frac{1}{3} \int_V^{rV} aT^4 dV$$

$$= \frac{1}{3} aVT^4(r - 1)$$

Problem 6

What is the Fermi energy of a two-dimensional electron gas in terms of the charge density σ of the gas?

Solution:

Quickly reviewing how we got to fermi energy in 3D– we took $f(\epsilon)$ to be 1 when $T = 0$, since all states are occupied up to the fermi energy. No states above the fermi energy are occupied, either.

$$\begin{aligned} N &= \int_0^\infty f(\epsilon) D(\epsilon) d\epsilon \\ &= \int_0^{\epsilon_F} [1] \left[\left(\frac{V}{2\pi^2} \right) \left(\frac{2m}{\hbar^2} \right)^{3/2} \epsilon^{1/2} \right] d\epsilon \end{aligned}$$

Integrating and rearranging for ϵ_F gave:

$$\epsilon_F = \frac{\hbar^2}{2m} \left(\frac{3N\pi^2}{V} \right)^{2/3}$$

To take this to 2D, $1/4$ of phase space is positive, and the states are a collection of rings accessible, so

$$\begin{aligned} \Gamma(k) &= 2 \left(\frac{1}{4} \right) \left(\frac{\pi k^2}{\left(\frac{2\pi}{L} \right)^2} \right) = \frac{Ak^2}{4\pi} \\ D(k) dk &= \frac{Ak}{2\pi} dk D(\epsilon) d\epsilon = \frac{A}{2\pi} \sqrt{\frac{2m\epsilon}{\hbar^2}} \cdot \frac{1}{2\sqrt{\epsilon}} \sqrt{\frac{2m}{\hbar^2}} d\epsilon \\ D(\epsilon) d\epsilon &= \frac{Am}{2\pi\hbar^2} d\epsilon \end{aligned}$$

Now I can do the same computation with $D(\epsilon)$ in 2D:

$$\begin{aligned} N &= \int_0^{\epsilon_F} \frac{Am}{2\pi\hbar^2} d\epsilon \\ N &= \frac{Am}{2\pi\hbar^2} \epsilon_F \\ \epsilon_F &= \frac{2N\pi\hbar^2}{Am} \end{aligned}$$

If some total charge Q is just the sum of total charges Nq , charge per area σ is given by:

$$Q = Nq \implies \frac{Q}{A} = \sigma = \frac{Nq}{A}$$

In this case, $q = e$, so

$$N = \frac{\sigma A}{e}$$

Therefore,

$$\epsilon_F = \frac{2\sigma\pi\hbar^2}{me}, \quad e = -1.6 \times 10^{-19} \text{ coulombs}$$

Problem 7

The universe has a radius of about 14 billion light-years and contains 2.7 K background radiation left over from the Big Bang. Estimate the temperature of the universe when it had a volume of 1 m³. You may assume that all the energy of the present background radiation was in the universe when it had the smaller volume and ignore any coupling to matter.

Solution:

If we consider the temperature of the universe to mostly depend on the energy from radiation, and assume that it expands equally in every direction, then:

$$\begin{aligned}E_i &= E_f \\aV_i T_i^4 &= aV_f T_f^4 \\T_i^4 &= \frac{4\pi r_f^3 T_f^4}{3V_i}\end{aligned}$$

Taking $r_f = 14 \times 10^9$ light years, $T_f = 2.7K$, and $V_i = 1$, temperature when the universe had a volume of 1 cubic meter was something on the order of

$$T_i \approx 1.572 \times 10^8 \text{ kelvin}$$