

Measuring the cosmological parameters with machine learning techniques.

Martín de los Ríos & Mariano Domínguez

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- Supervised learning.

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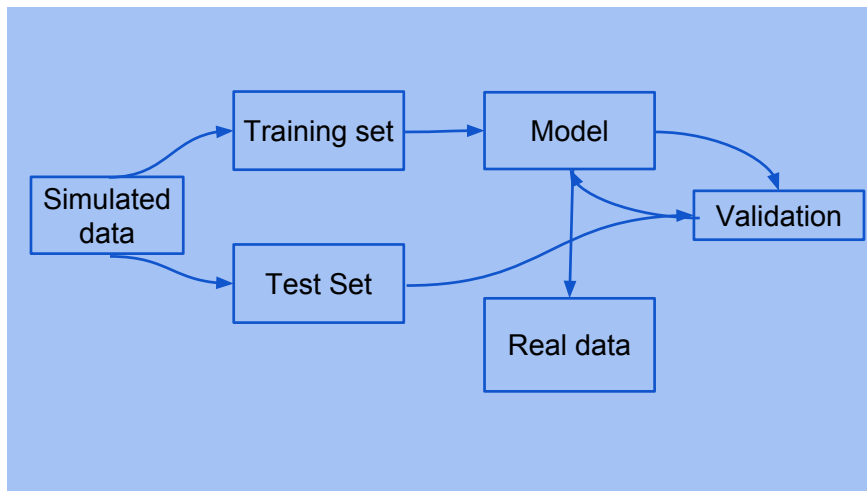
- The training sample.
- Applications.

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Final Remarks.

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Supervised Learning.



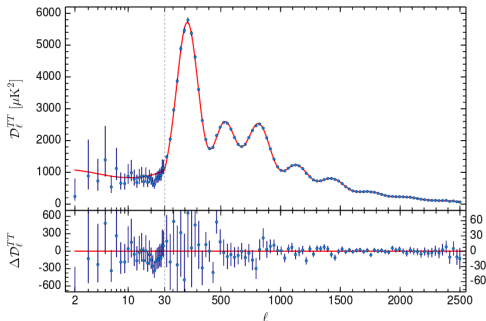
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The Standard model.

Homogeneous and isotropic Universe \rightarrow FRW metric

$$ds^2 = dt^2 - a^2(t) \left[\frac{dr^2}{1 - kr^2} + r^2(d\theta^2 + \sin^2\theta d\phi^2) \right]$$

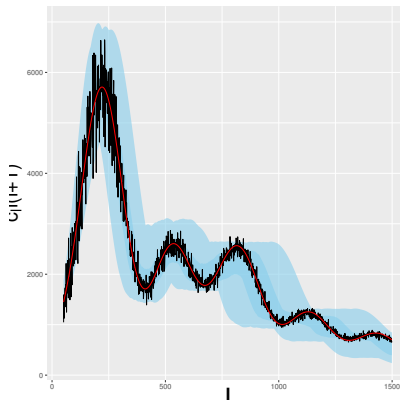
$$\left(\frac{H}{H_0}\right)^2 = \Omega_{rad} a^{-4} + \Omega_m a^{-3} + \Omega_\Lambda - Kc^2 a^{-2}$$



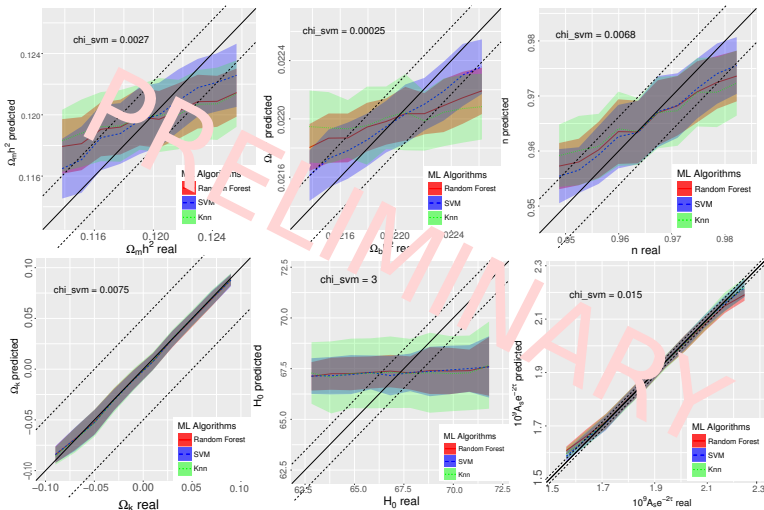
Planck Collaboration 2015 (1502.01589)

The training sample.

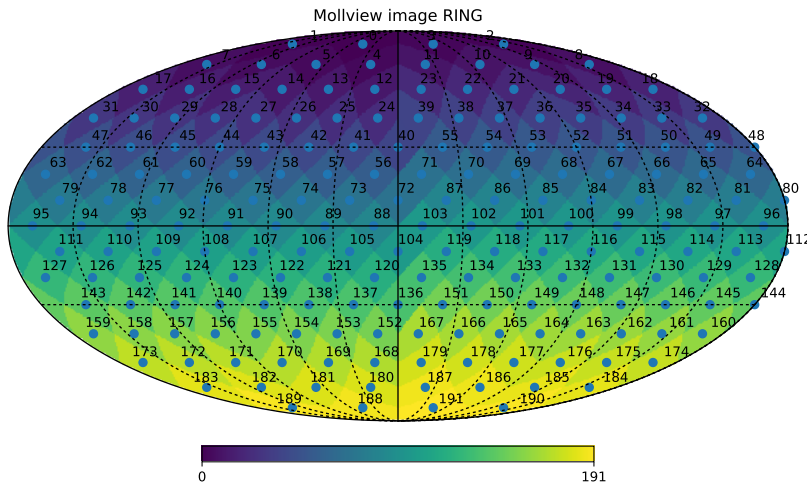
CAMB: Code for Anisotropies in the Cosmic Background



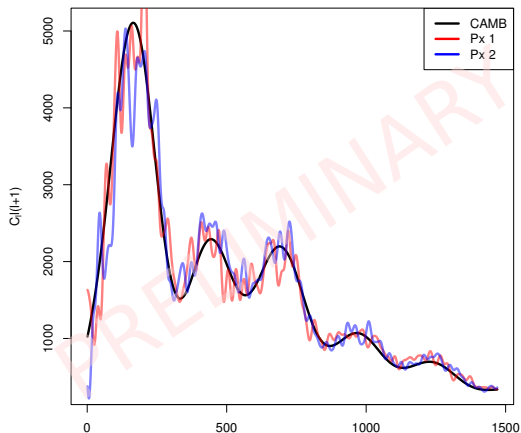
Studying different Machine Learning algorithms.



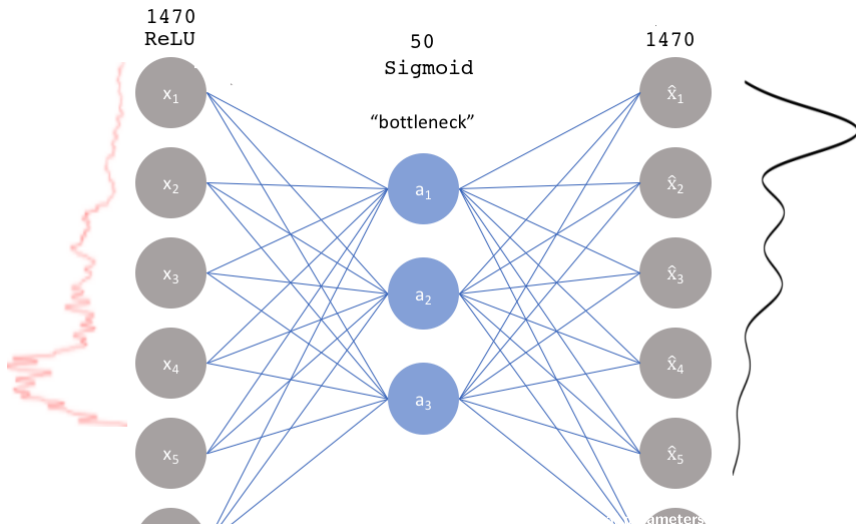
Measuring the cosmological parameters angular distributions.



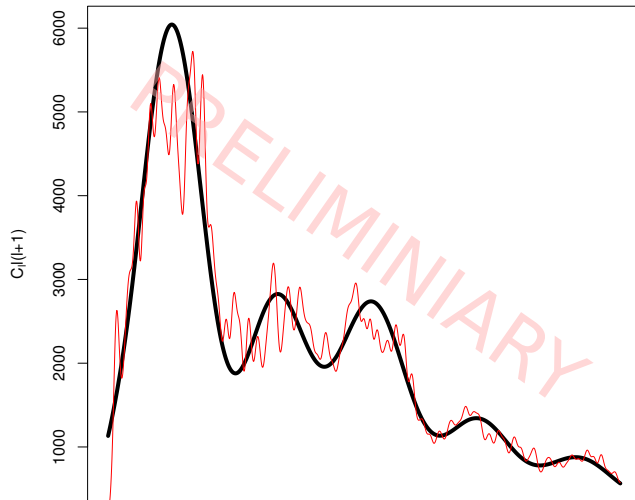
Measuring the cosmological parameters angular distributions.



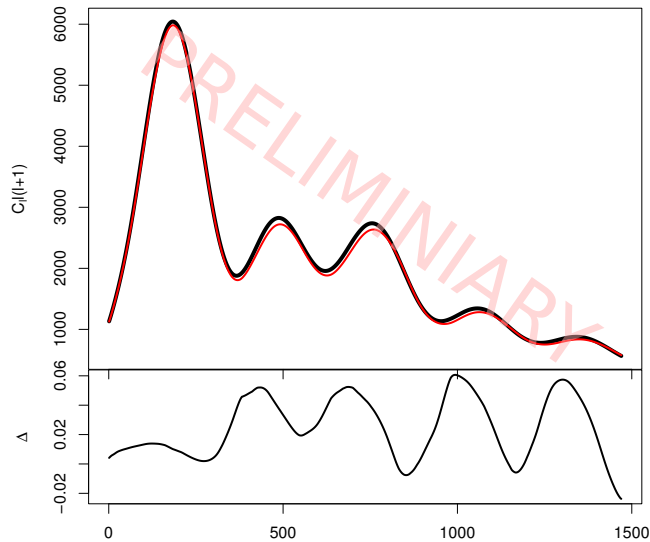
Denoising Autoencoders



ps reconstruction

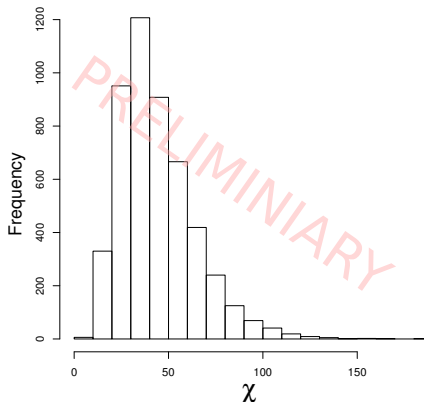


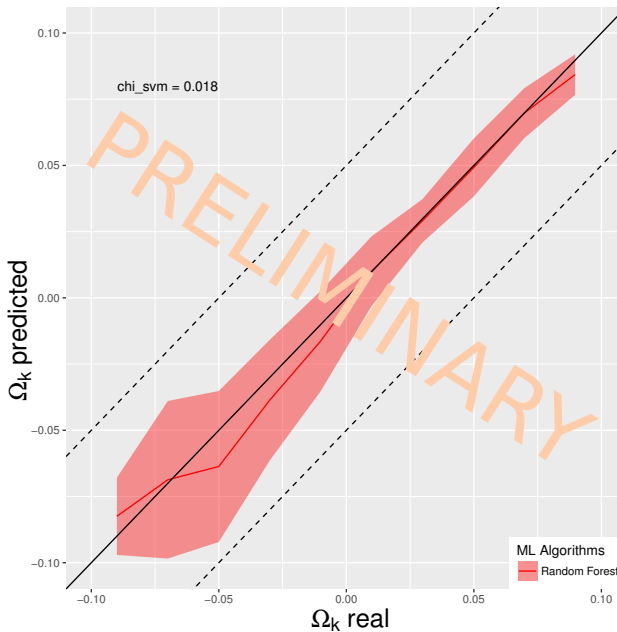
ps reconstruction

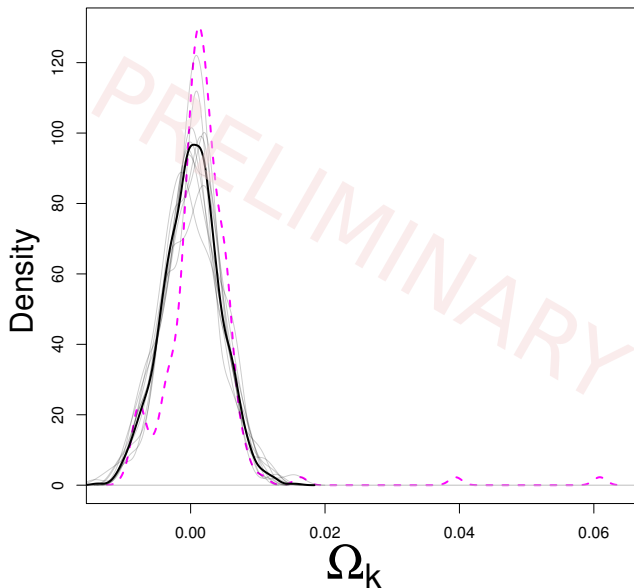


Measuring the cosmological parameters angular distributions.

$$\chi = \frac{\sum_{i=1}^{npix} |C_{l,real} - C_{l,rec}|}{npix}$$







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Final Remarks

- We developed a machine learning technique that estimate the cosmological parameters using CMB information.
- This technique can be easily extended to use more cosmological information as features (BAO, correlation function, SZ emission, etc.).
- As a first application we are studying the angular distribution of the cosmological parameters.
- We developed a machine learning technique that reconstructed the power spectra from pixels.
- We do not found any significant curvature departure from what is expected in an homogeneous and isotropic universe, with the exception of some pixels that are in the galactic plane.
- We will extend the parameters space and add polarization information.



Changing the minimum mutipole.

