Determining the Dark Matter distribution in galaxies with Deep Learning (2111.08725)

As part of the darkmachines projects challenges: https://darkmachines.org/

Martín de los Rios *IFT/UAM (Madrid Spain) - ICTP/SAIFR (Sao Paulo Brasil)*
Mihael Petac *CAC (Slovenia) - LUPM (France)*
Bryan Zaldivar *IFIC (Valencia Spain)*
Nina Bonaventura *Cosmic Dawn Center (Copenhagen Denmark)*
Francesca Calore *CNRS LAPTh (Annecy France)*
Fabio Iocco *Dipartimento di Fisica INFN (Napoles Italy)*

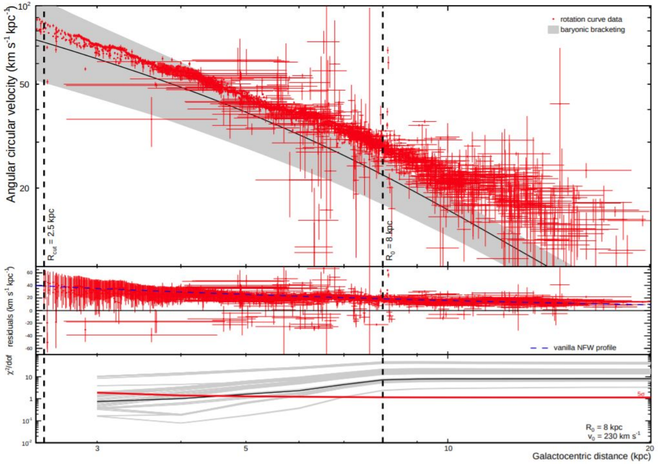
Outline

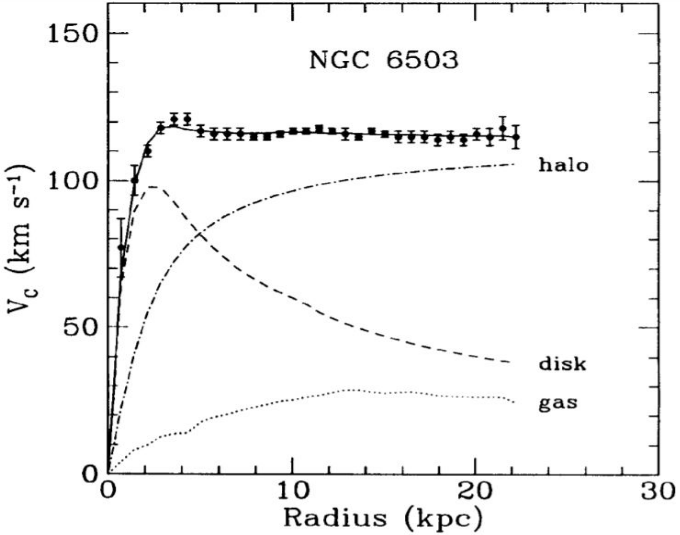
- Introduction
- Constructions of the dataset
  - TNG100 Simulations (1707.03401, 1707.03395, 1707.03395, ...)
  - SKIRT (2003.00721)
  - MARTINI (https://github.com/kyleaoman/martini)
- Results
  - Prediction of the Dark matter profile
  - Comparison between different architectures
  - Comparison between different inputs
  - Comparison with Rotation Curve method
- Conclusions and Future work

Introduction

[Evidence for dark matter in the inner Milky Way. 1502.03821](https://arxiv.org/abs/1502.03821)

Fabio Iocco, Miguel Pato & Gianfranco Bertone



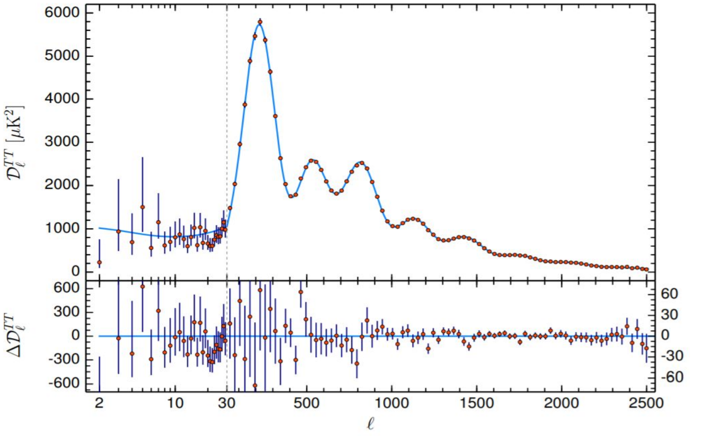


[NGC 6503 Rotation Curve. 0812.4005](https://arxiv.org/abs/0812.4005)

Katherine Freese

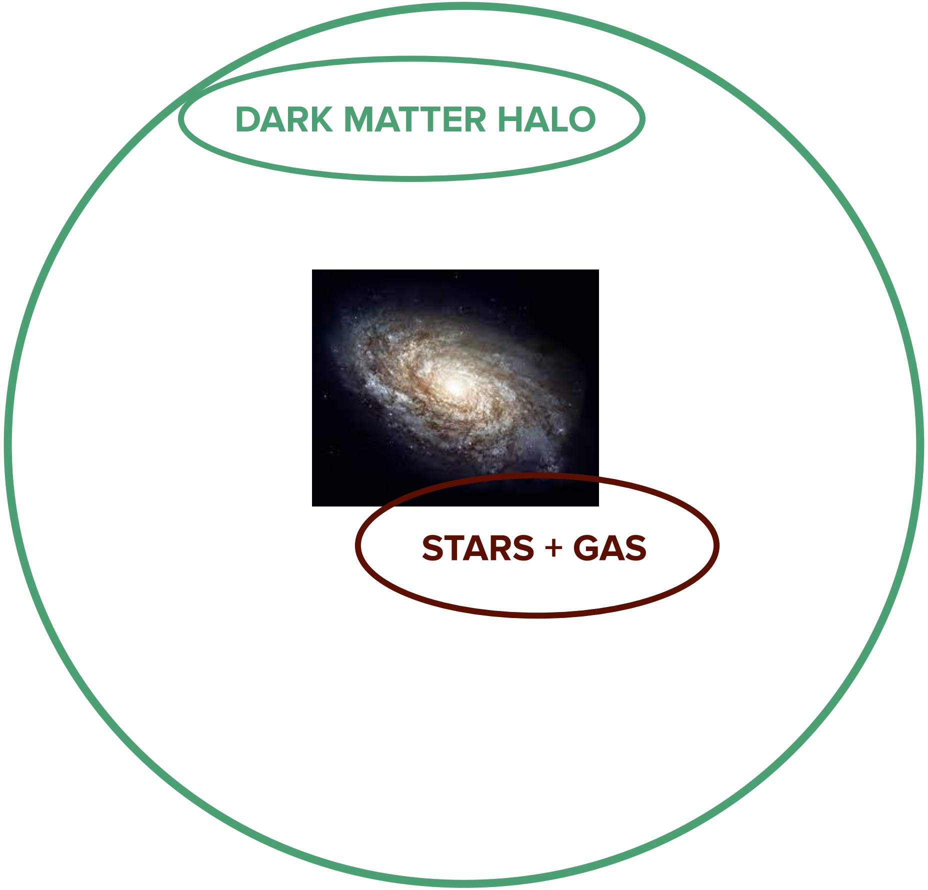
Markevitch et al.

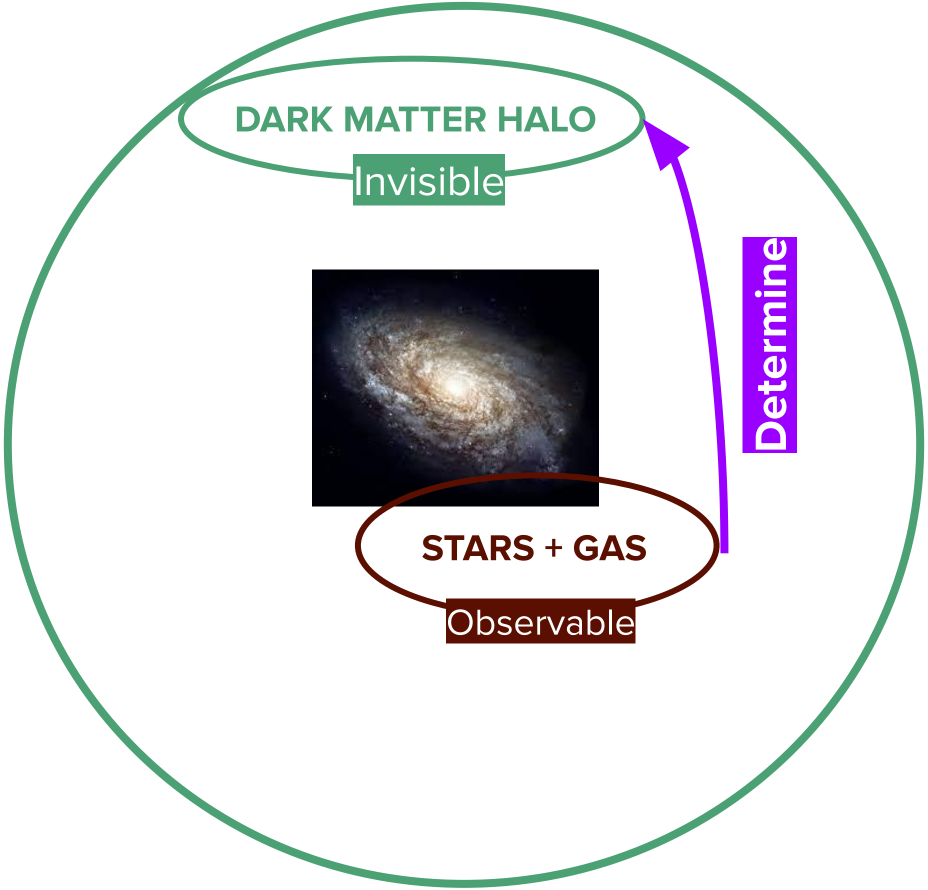
https://arxiv.org/abs/astro-ph/0309303



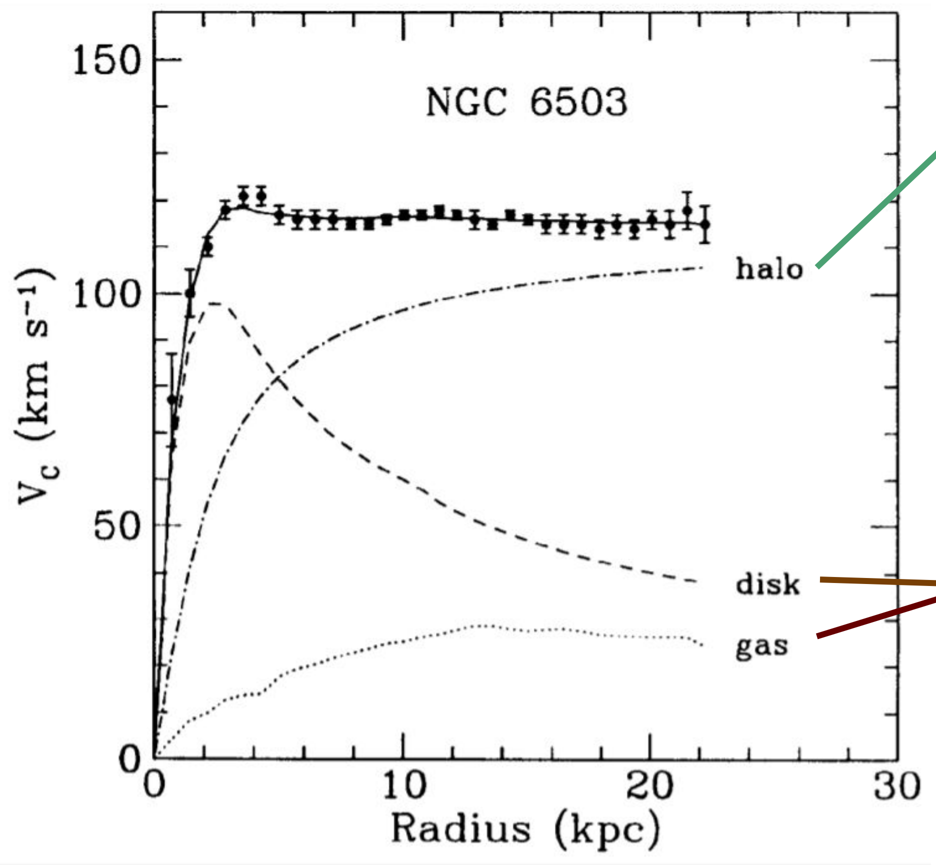
Planck Collaboration

https://arxiv.org/abs/1807.06209


DARK MATTER HALO
STARS + GAS


DARK MATTER HALO
Invisible
Determine
STARS + GAS
Observable

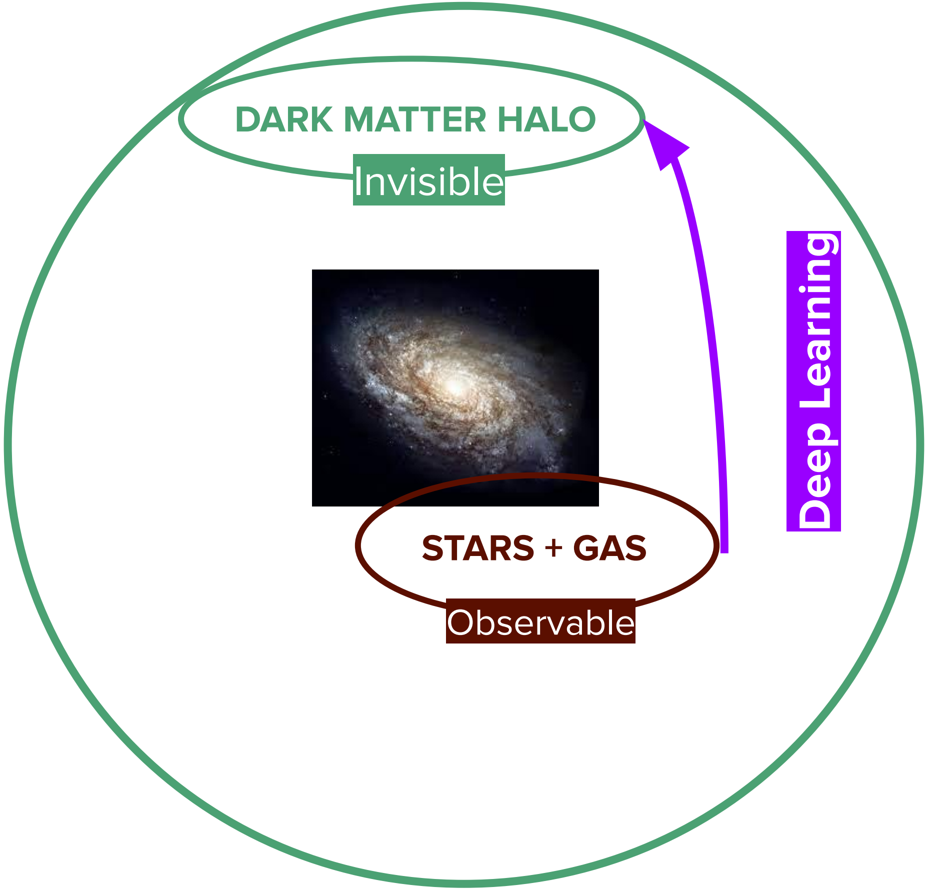
NGC 6503 Rotation Curve



Assume an NFW Profile

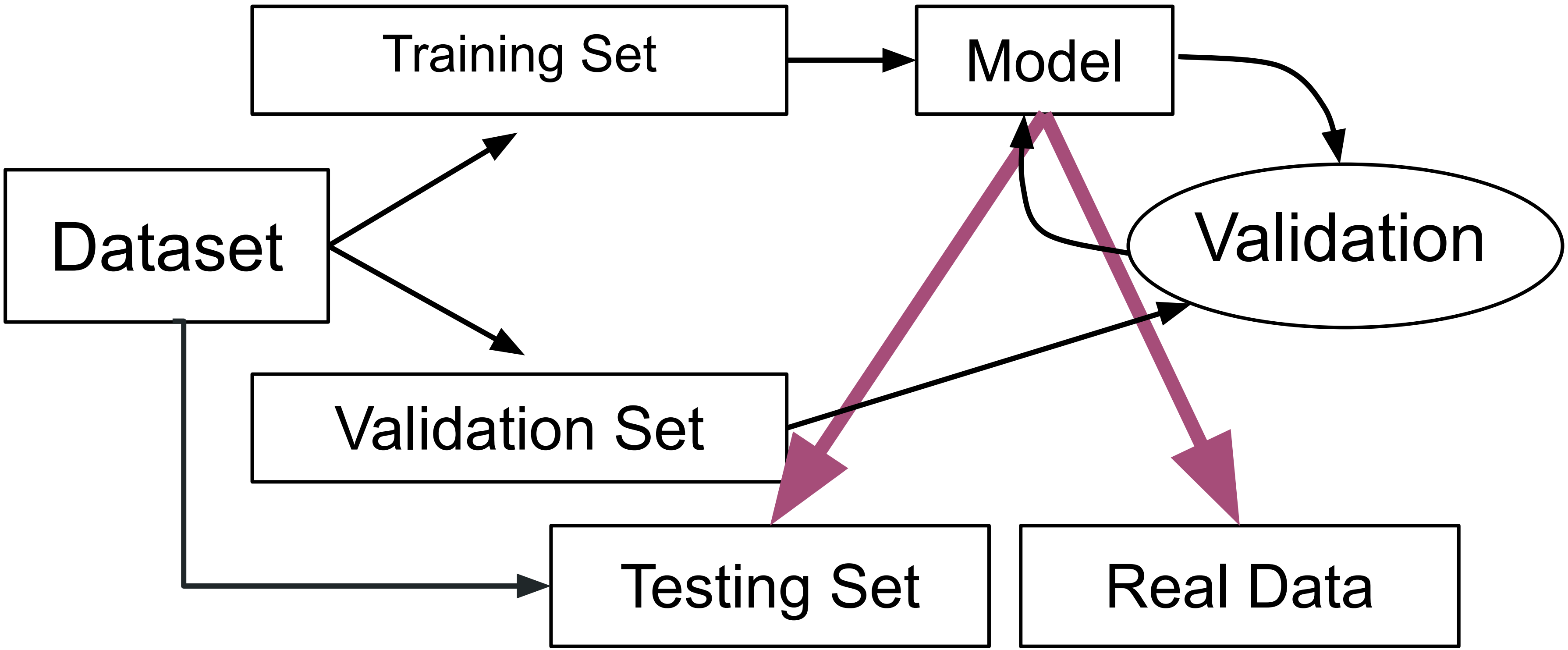
Estimated best-fit parameters

Inferred from observations

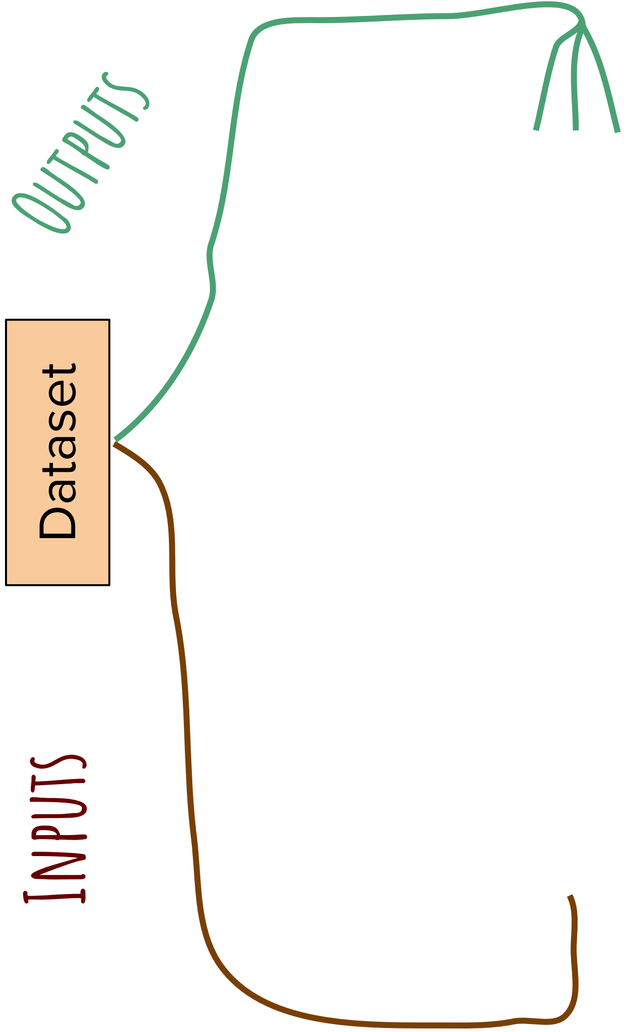

DARK MATTER HALO
Invisible
STARS + GAS
Observable
Deep Learning

Brief (1 slide) Introduction to Supervised Learning

Supervised Learning



Construction of the Dataset

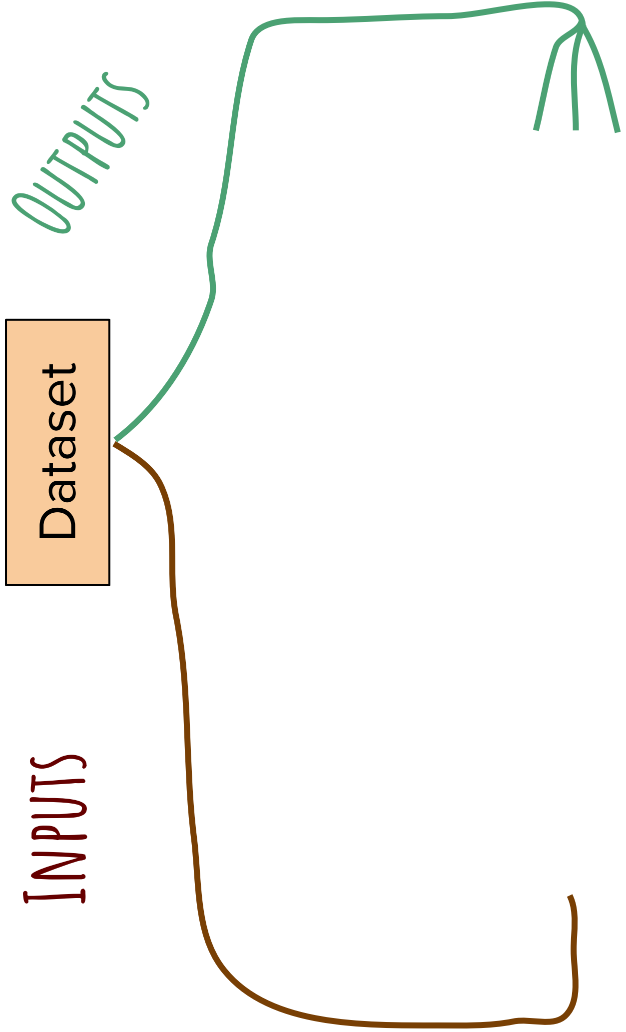

Outputs
Dataset
Inputs

TNG100 Cosmological Hydrodynamical Simulation
(https://www.tng-project.org/)

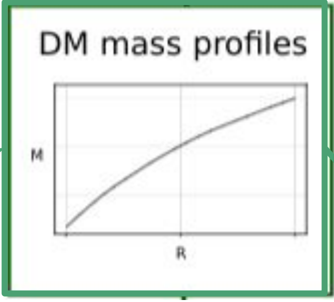
- Planck cosmology
- 106.5 Mpc by side
- 1820^3 DM particles
- 1820^3 hydrodynamic cells
- DM resolution 7.5 *10^6 M☉
- Baryon resolution 1.4*10^6 M☉
- 136 snapshots from z=127 to z=0

Property	Criterium
Simulation snapshot	99 ($z = 0$)
Stellar mass	$10^{10}\ M_\odot \leq M_\star \leq 10^{12}\ M_\odot$
Star formation rate	SFR $\geq 0.1\ M_\odot$/yr
Central galaxy	SubhaloParent = 0
Cosmological origin	SubhaloFlag = 1

TNG100 (z = 0)



TNG100 (z = 0)



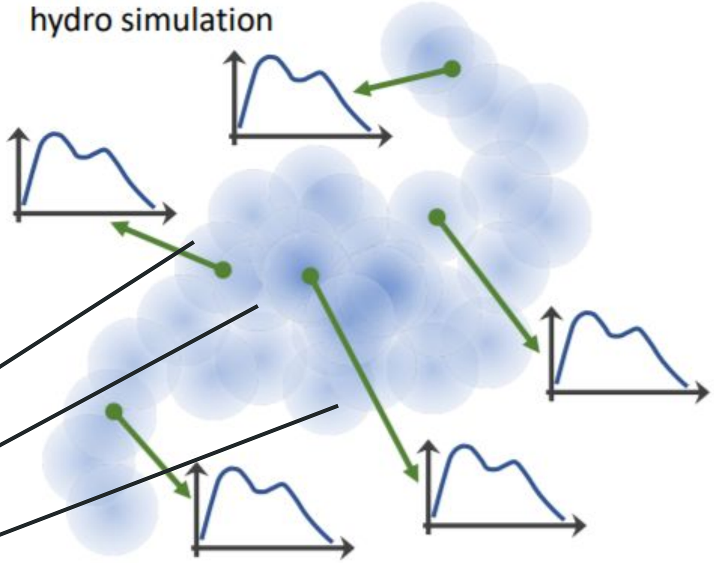
OUTPUTS

Dataset

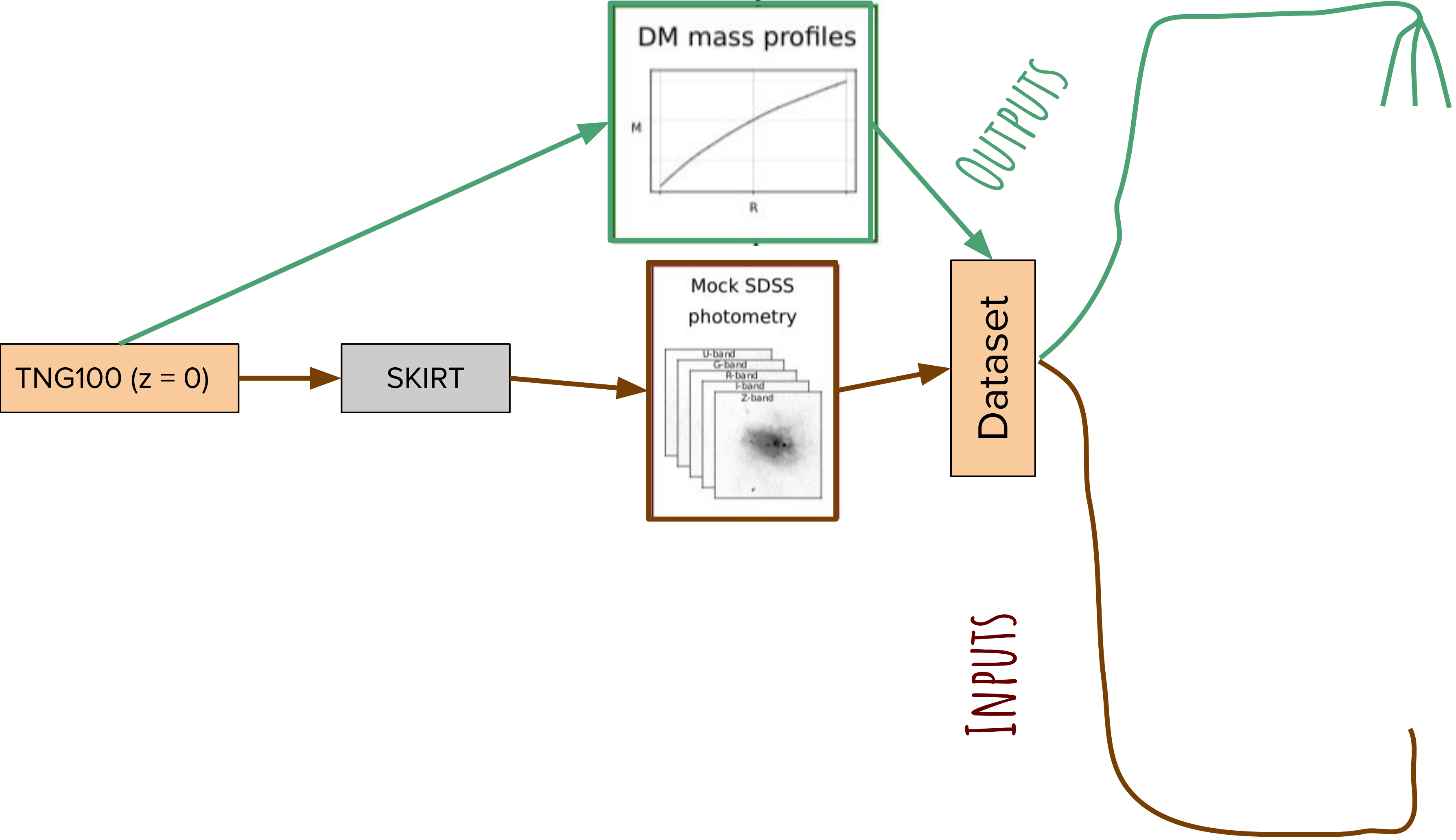
INPUTS

SKIRT* (2003.00721, skirt.ugent.be)

Radiative transfer code which emulates the stellar emissions and subsequent light-ray propagation to the observer, taking into account the absorption and re-emission by dust.



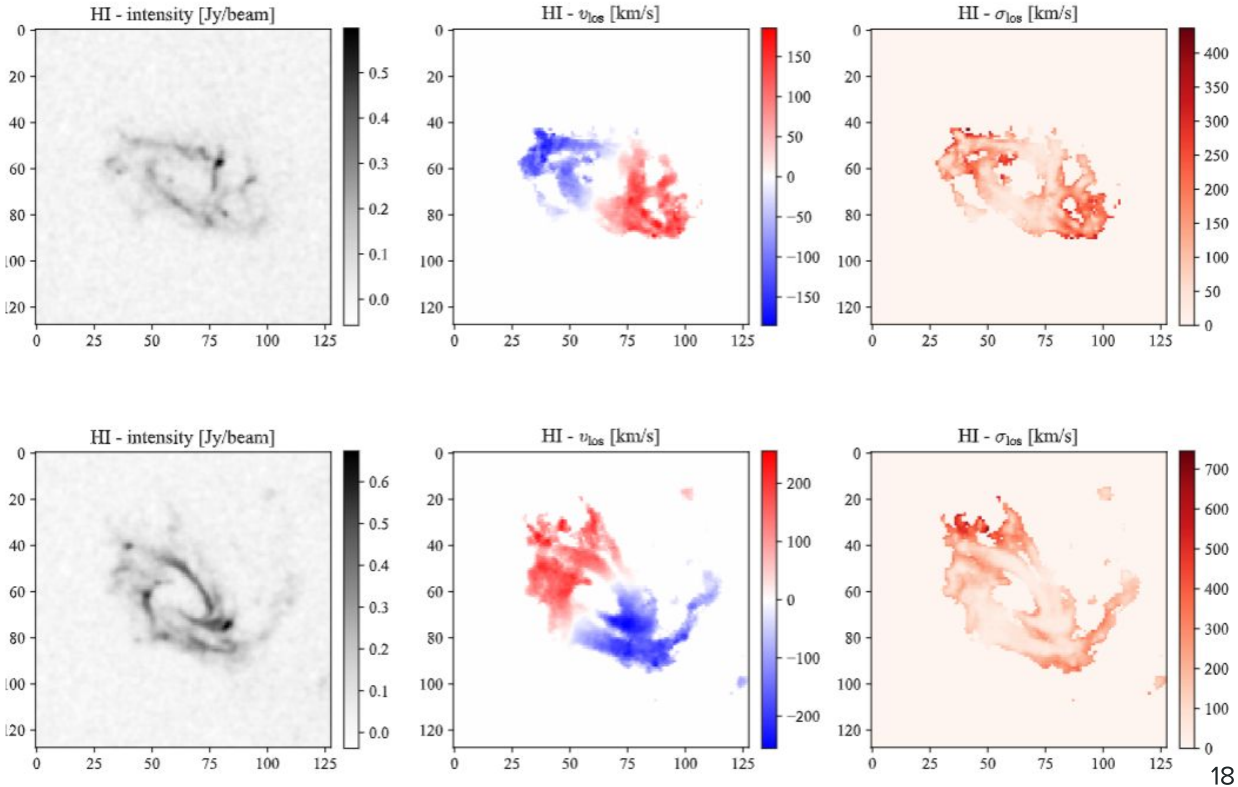
SED interpolated from template family for each particle or cell

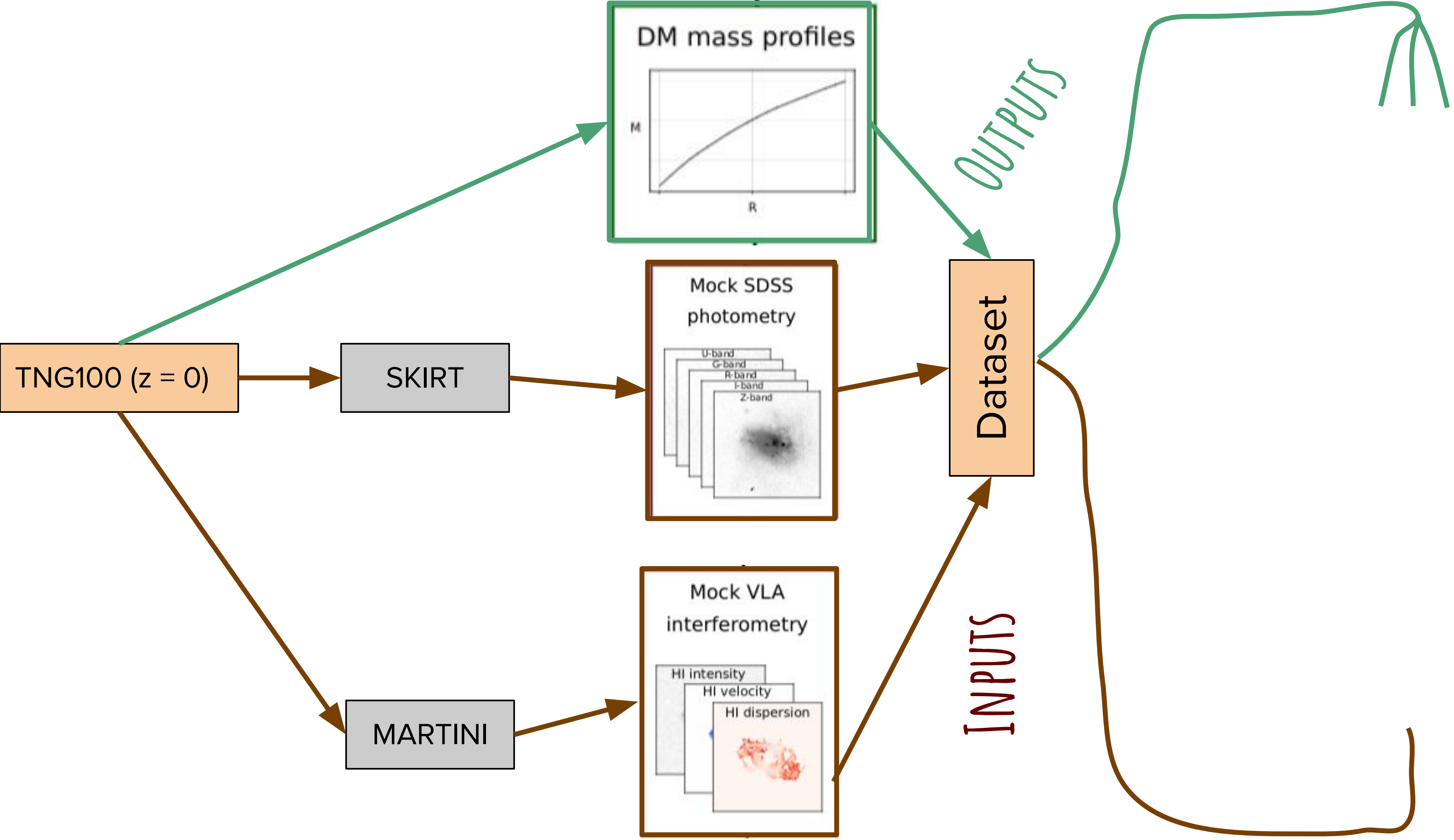


DM mass profiles
M
R
OUTPUTS
Mock SDSS photometry
U-band
G-band
R-band
I-band
Z-band
TNG100 (z = 0)
SKIRT
Dataset
INPUTS

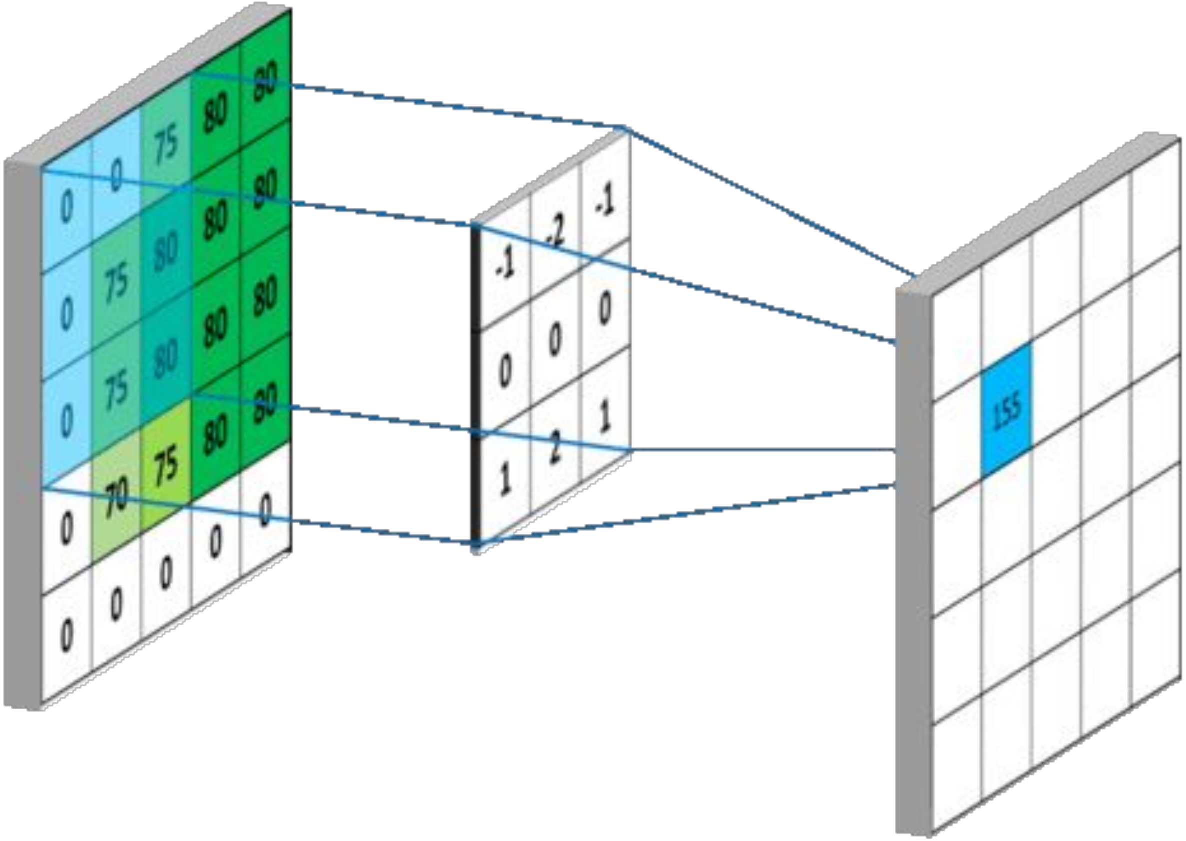
MARTINI (1706.07478 ; https://github.com/kyleaoman/martini)

Allows for the creation of synthetic resolved HI line observations (i.e. data cubes) directly from the snapshot of a hydrodynamic simulation, and its posterior analysis.

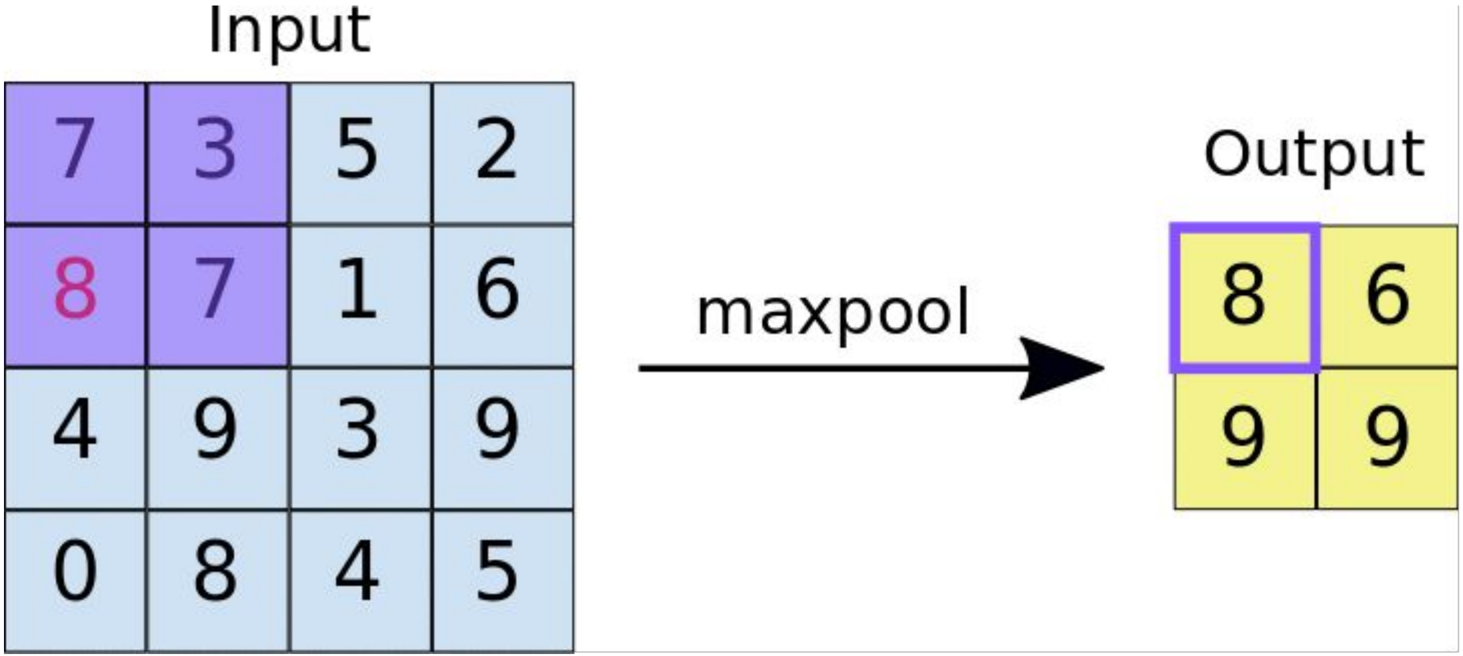



DM mass profiles
M
R
OUTPUTS
TNG100 (z = 0)
SKIRT
Mock SDSS
photometry
U-band
G-band
R-band
I-band
Z-band
Dataset
MARTINI
Mock VLA
interferometry
HI intensity
HI velocity
HI dispersion
INPUTS

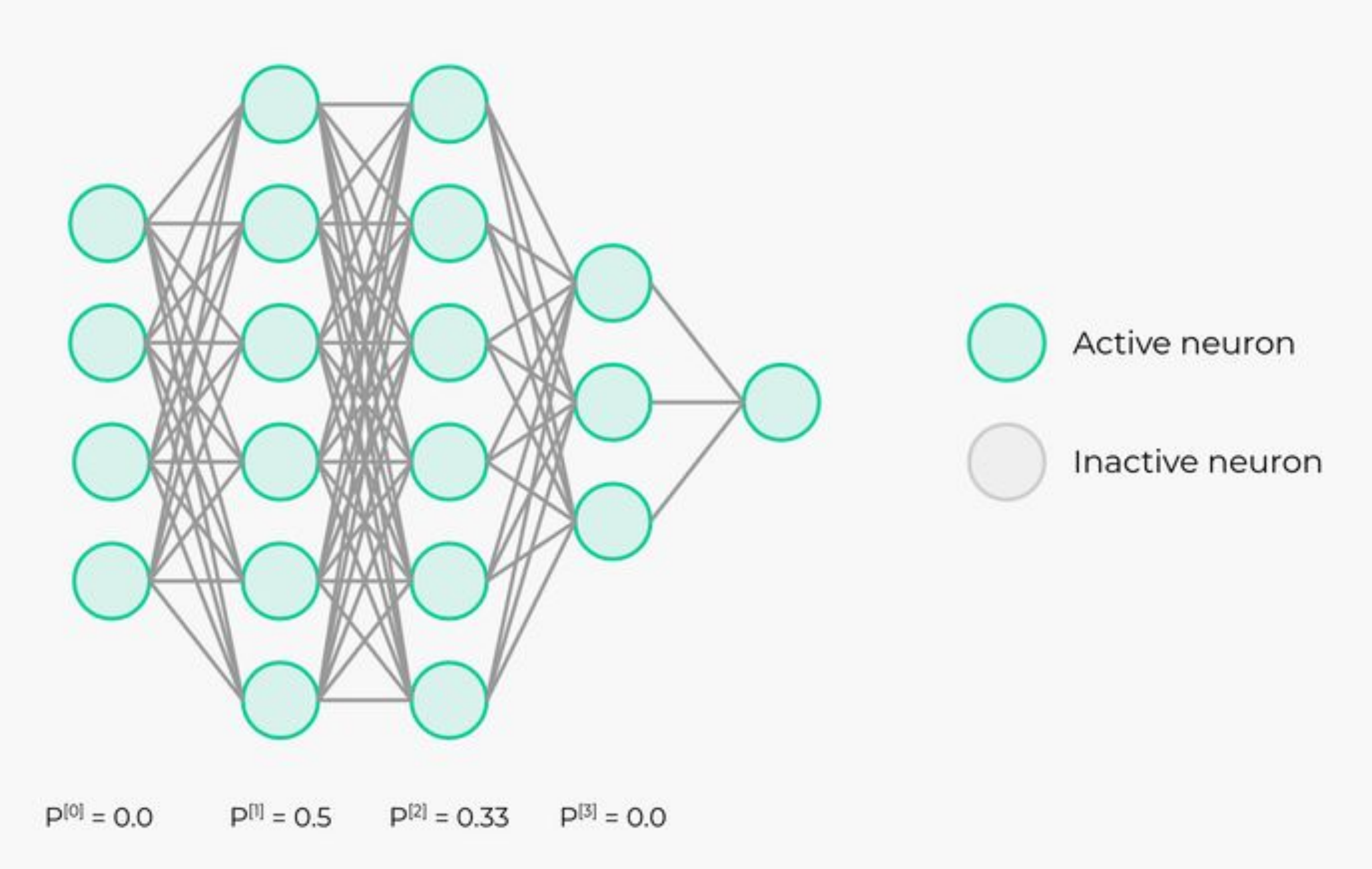
Convolutional layers



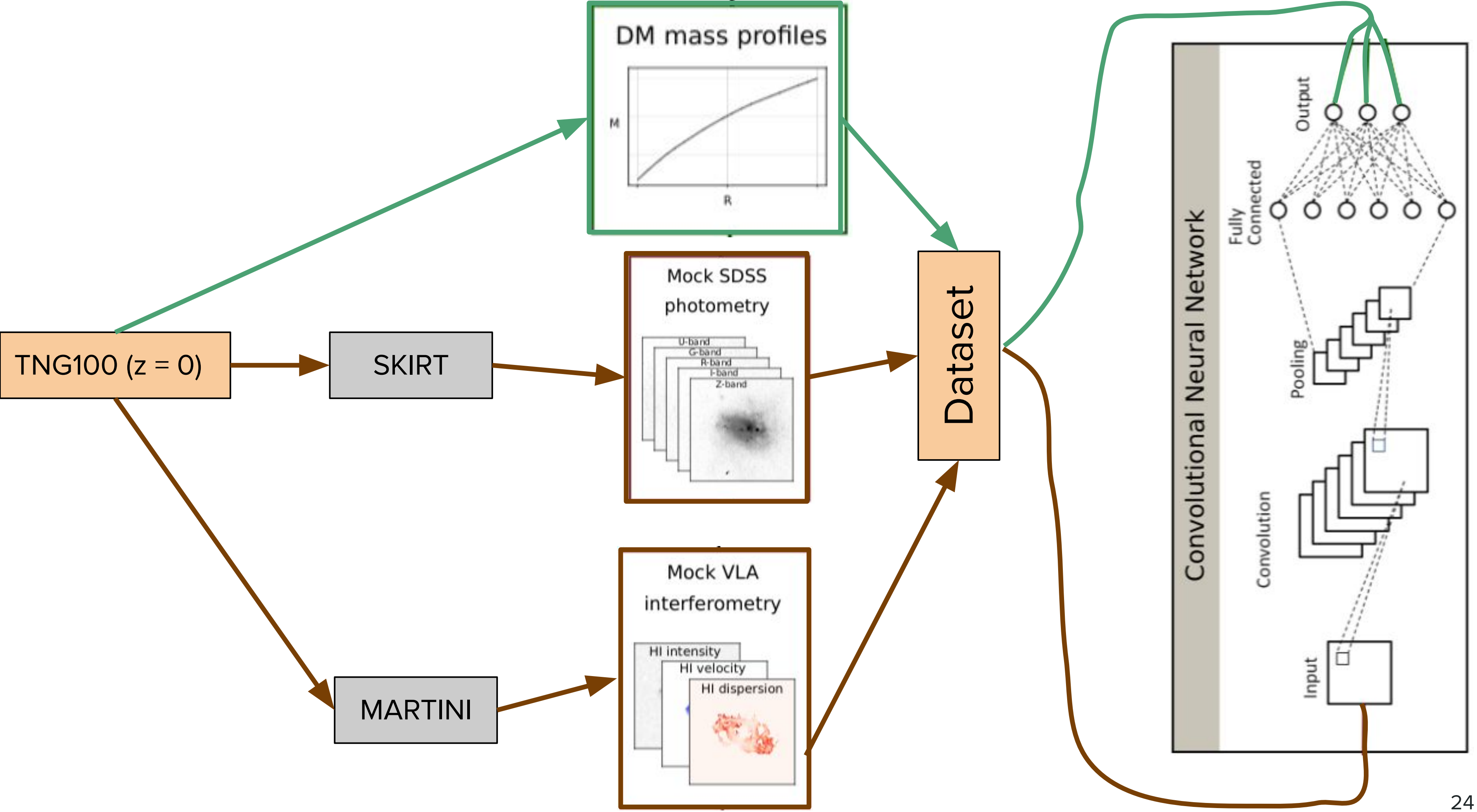
Pooling layers



Dropout



Layer	Details
2D convolution	64 kernels, 5×5 px kernel size, 2 px stride, ReLU activation
2D max pooling	2 px pooling
Dropout	50% dropout fraction
Batch normalization	
2D convolution	128 kernels, 5×5 px kernel size, 2 px stride, ReLU activation
2D max pooling	2 px pooling
Dropout	50% dropout fraction
Batch normalization	
2D convolution	256 kernels, 5×5 px kernel size, 2 px stride, ReLU activation
Batch normalization	
Dense	256 units, ReLU activation
Dropout	50% dropout fraction
Batch normalization	
Dense	128 units, ReLU activation
Dropout	50% dropout fraction
Batch normalization	
Dense	64 units, ReLU activation
Dropout	50% dropout fraction
Batch normalization	
Dense (output)	20 units, linear activation



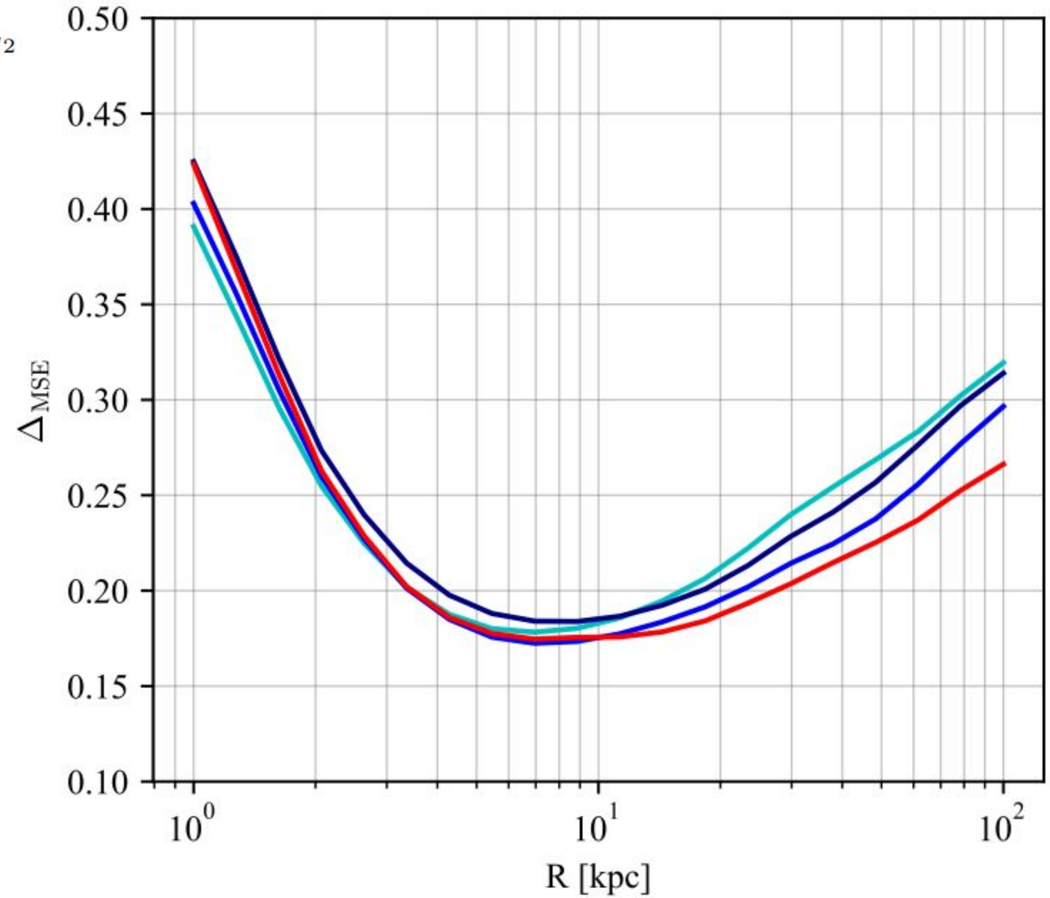
DM mass profiles
M
R
TNG100 (z = 0)
SKIRT
Mock SDSS photometry
U-band
G-band
R-band
I-band
Z-band
MARTINI
Mock VLA interferometry
HI intensity
HI velocity
HI dispersion
Dataset
Convolutional Neural Network
Output
Fully Connected
Pooling
Convolution
Input

Results

Comparison between different architectures

$$\Delta_{\mathrm{MSE}}(R_i) = \left[\frac{1}{N}\sum_{j=1}^{N}\left(\frac{\mu_j(R_i) - \hat{\mu}_j(R_i)}{\hat{\mu}_j(R_i)}\right)^2\right]^{1/2}$$

- Architecture A
- Architecture B
- Architecture C
- ResNet50



Comparison between different inputs

SDSS I

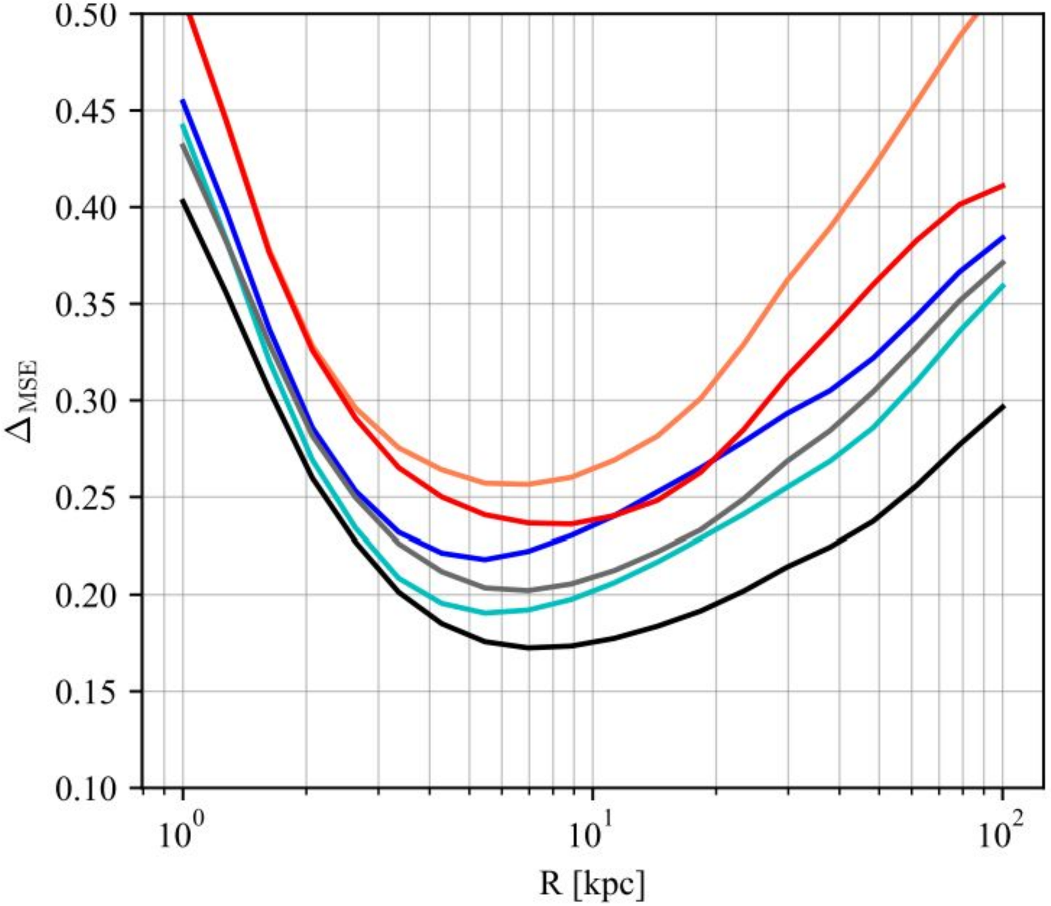
SDSS URZ

V_{HI}

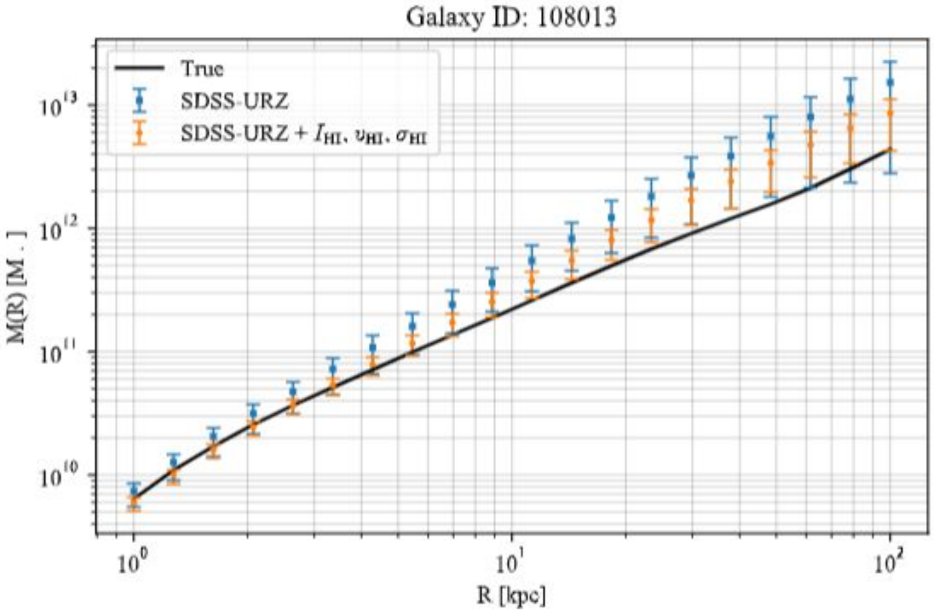
$I_{HI} + V_{HI} + \sigma_{HI}$

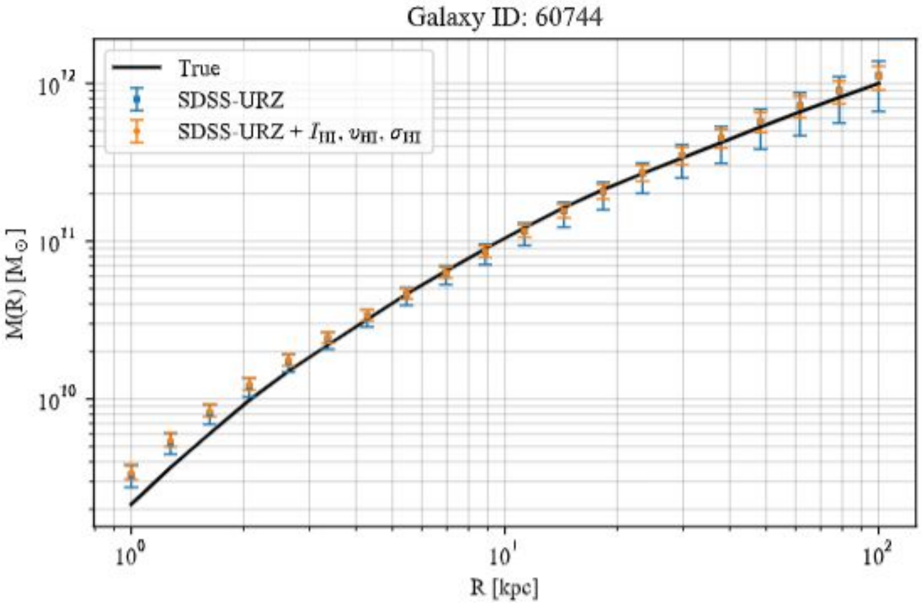
SDSS URZ + V_{HI}

SDSS URZ + $I_{HI} + V_{HI} + \sigma_{HI}$



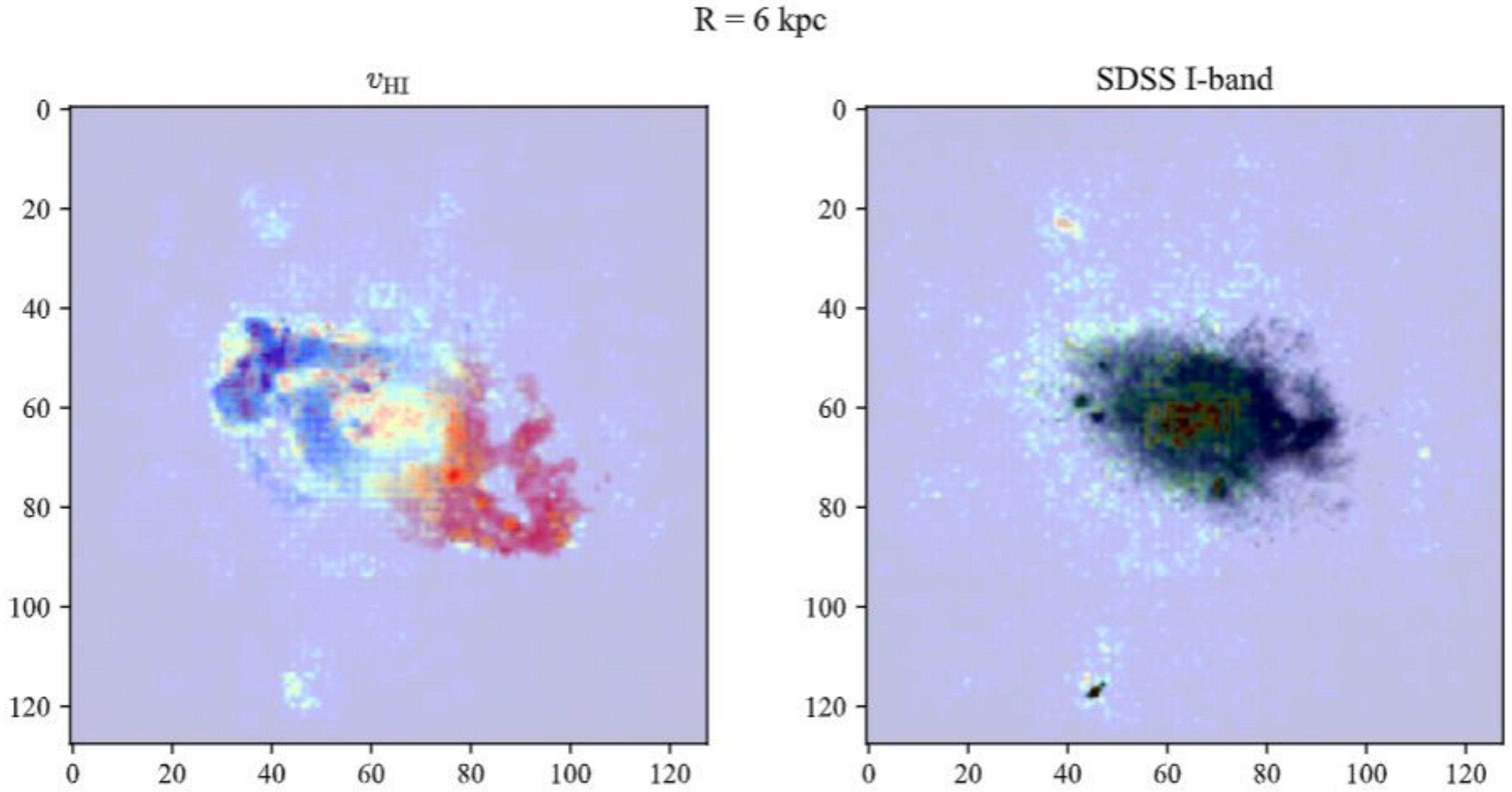
Prediction of the dark matter profile



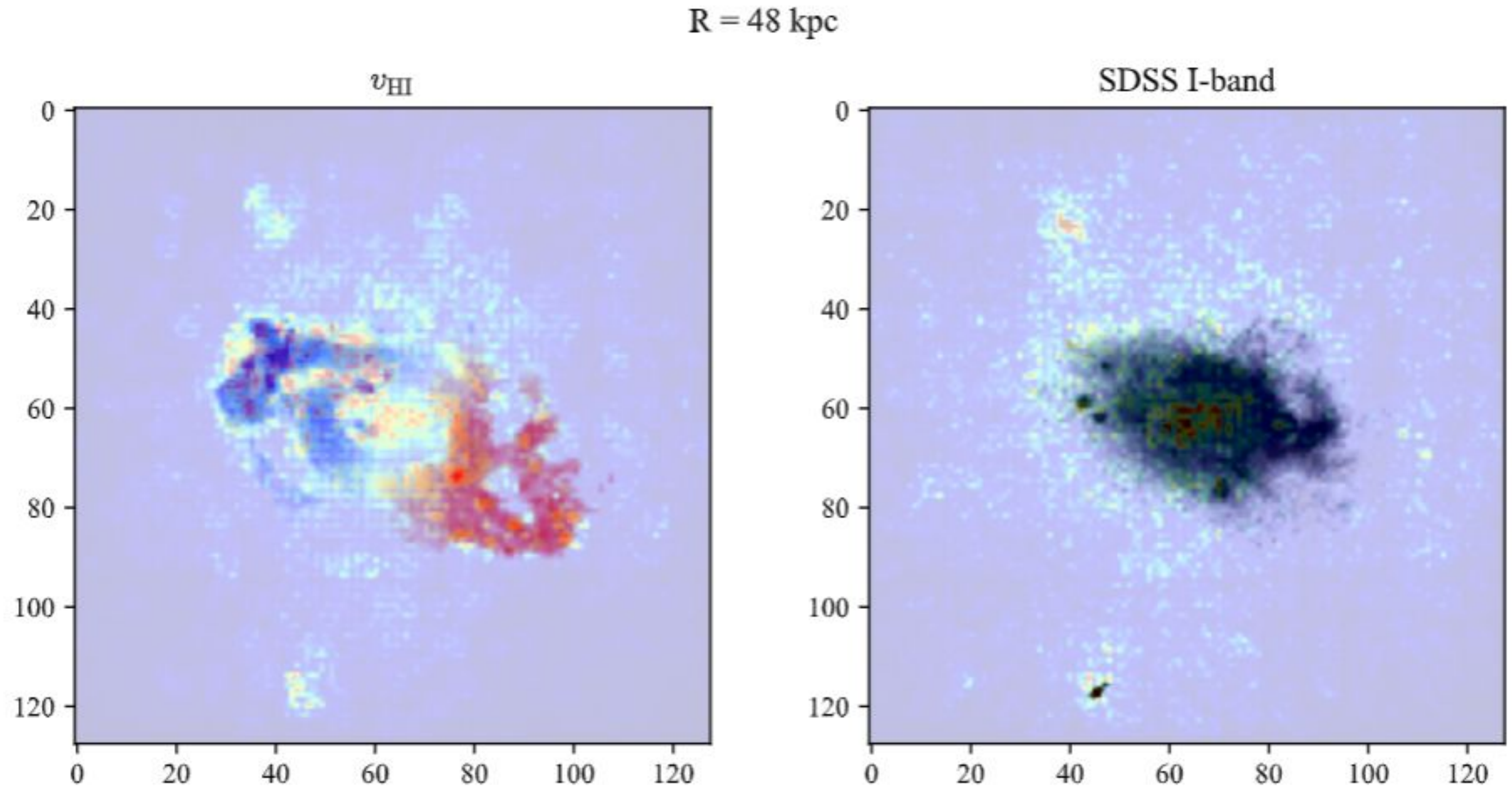


Understanding the results

$$S_{ij} \equiv \frac{\partial y}{\partial x_{ij}}$$

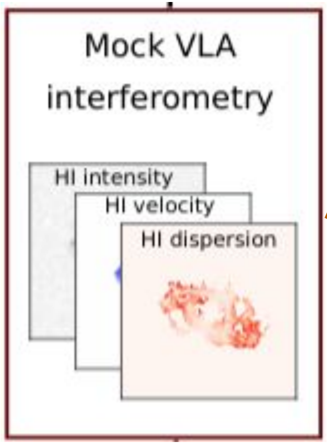


Understanding the results



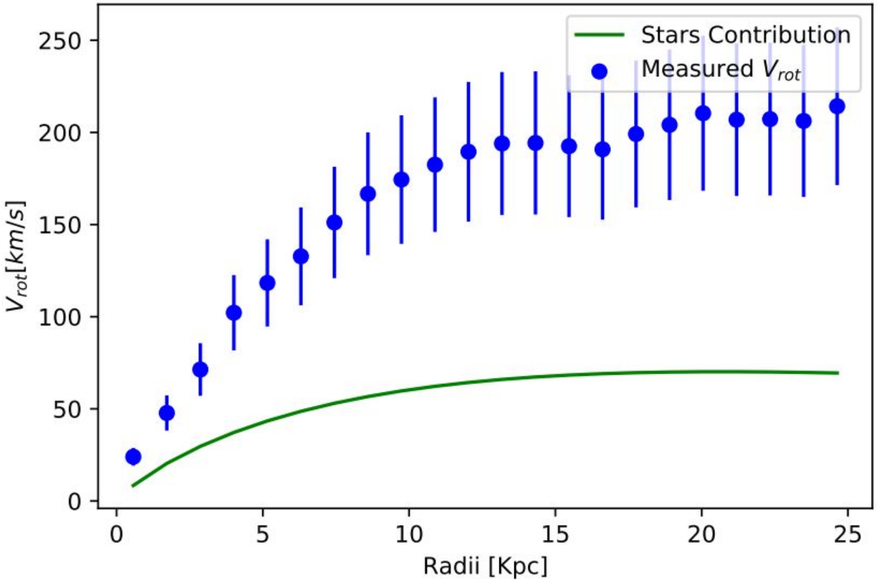
Comparison with RC analysis

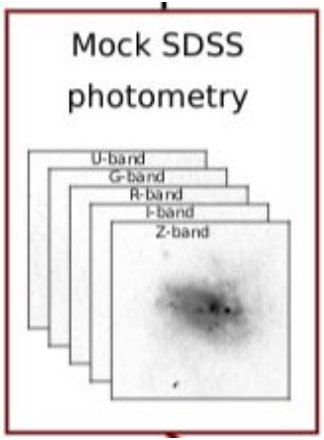
(preliminary)



^{3D}BAROLO

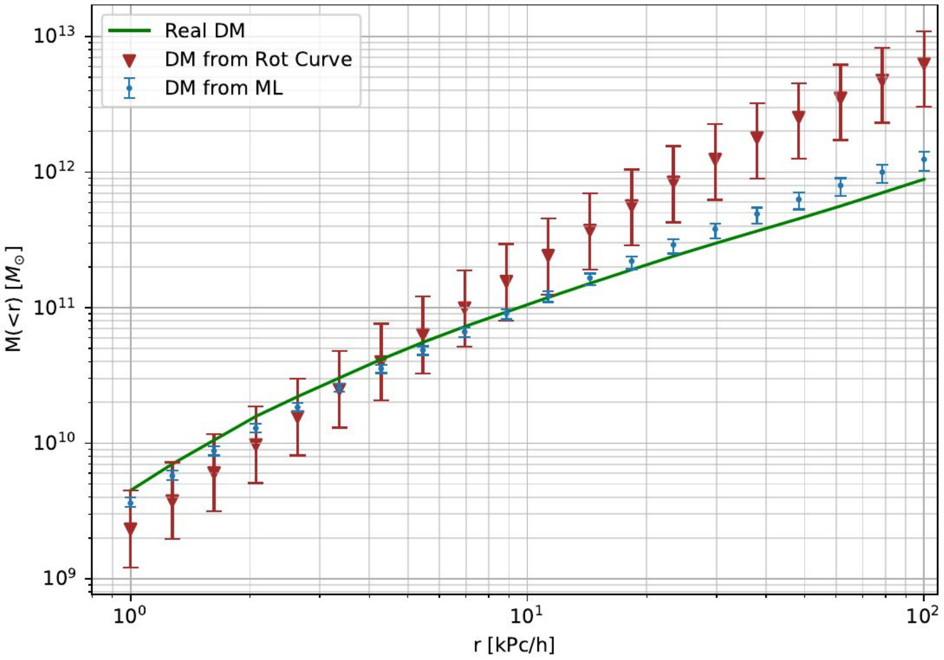
https://arxiv.org/abs/1505.07834

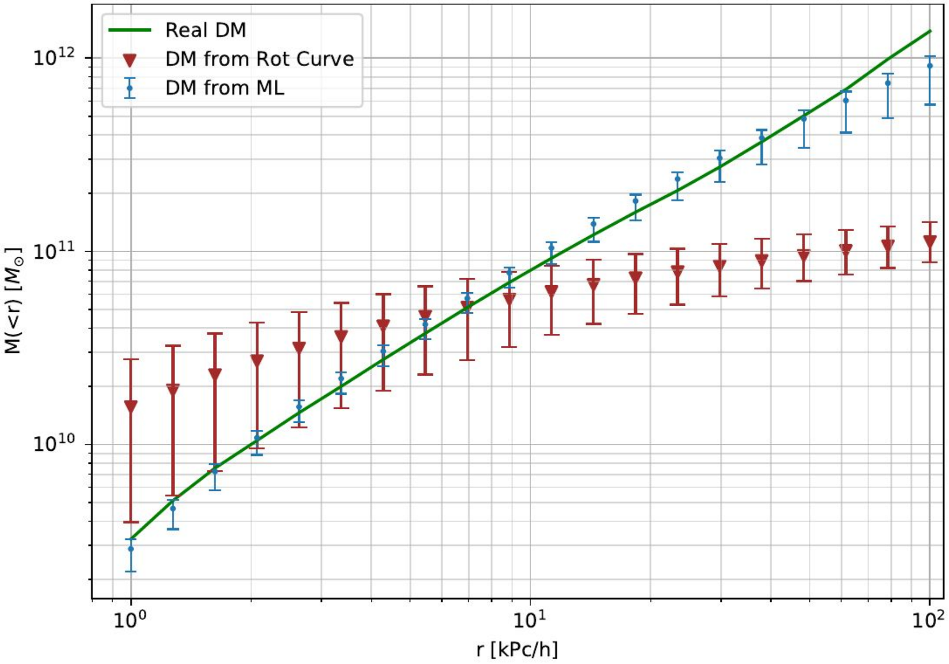




Autoprof

https://arxiv.org/abs/2106.13809


Real DM
DM from Rot Curve
DM from ML
M(<r) [$M_\odot$]
r [kPc/h]
10^{9}
10^{10}
10^{11}
10^{12}
10^{13}
10^{0}
10^{1}
10^{2}


Real DM
DM from Rot Curve
DM from ML
M(<r) [$M_\odot$]
r [kPc/h]
10^{10}
10^{11}
10^{12}
10^{0}
10^{1}
10^{2}

Conclusion and Future Work

- Our algorithm is able to reconstruct the DM distribution profile with high performance throughout the extension of the galaxy.

- The highest performance is achieved in the intermediate regions with a mean square error below 0.2 using all the photometric and spectroscopic information.

- Even in the absence of spectroscopic information, our method is able to recover the dark matter profile with a mean square error below 0.3 in the intermediate regions.

- Our reconstruction of the DM distribution is completely data-driven, and does not need any assumption on the shape nor the functional form of the DM profile.

- The method developed here is applicable to different types of galaxies since it does not rely on explicit physical assumptions regarding the dynamical state of the system.

- The results achieved have been obtained for galaxies with masses in the range ~10^{10}-10^{12} M⊙ but the methodology can be extended to a broader mass range.

- We will make a comparison with the dark matter profile obtained through the traditional rotation curve analysis for the simulated galaxies.

- Study the robustness of our results to the hydrodynamical cosmological simulation.

- Apply our method to real galaxies and compare the results with other estimations.

THANK YOU