

Master thesis

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2019-2020

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Part I

Questions

1 Questions for professor Sundqvist

- pcyg.f90: what is the parameter p ?
- pcyg.f90: what is meant by Eddington limb-darkening?
- Sundqvist+ 2009: what is the geometry of a *slice*?
- CMFAA course notes p.13 (the example) what is understood by plane-parallel geometry and is it 1D or 2D?
- CMFAA course notes p.15: why is this called diffusion $F = T^3 \frac{dT}{dx}$ (flux proportional to local gradient in temperature)?
- unclassified: what is the terminal velocity v_∞ ?
- unclassified: what is Sobo-distribution?

2 Questions for professor Samaey

- what is the difference between Monte Carlo and equation-free computing?

3 Solved questions

- Sundqvist+ 2009: what is thermal velocity (see Wikipedia)
- Sundqvist+ 2009: what is line force (see explanation Dylan)
- unclassified: what is a flux limiter? (see course notes)
- unclassified: what is cross section of scattering (see Wikipedia)
- Puls manual: p.26: how does the Milne equation appear? (see library book)

Part II

Monte Carlo Radiative Transfer

4 Glossary

- SED: spectral energy distribution
- (spectral) line-force: force on material in stellar atmosphere

5 Very broad introduction: Radiation Hydrodynamics

The material here originates from the master thesis of Nicolas Moens and the course notes *Introduction to numerical methods for radiation in astrophysics* from professor Sundqvist.

Heat flux diffusion equation $u_t = u_{xx}$. The flux

Specific intensity and its angular moments

specific intensity	$\Delta\epsilon = \boxed{I_\nu} A_1 A_2 / r^2 \Delta\nu \Delta t$
energy density	$E = \frac{1}{c} \iint I_\nu d\nu d\Omega$
flux vector	$F = \iint I_\nu n d\nu d\Omega$
pressure tensor	$P = \iint I_\nu n n d\nu d\Omega$
mean intensity	$J_\nu = \frac{c}{4\pi} E_\nu$
Eddington flux	$H_\nu = \frac{1}{4\pi} F_\nu$
Eddington's K	$K_\nu = \frac{c}{4\pi} P_\nu$

RHD equations The full RHD equations consist of

- five partial differential equations
- one HD closure equation, e.g. (i) variable Eddington tensor method or (ii) flux limited diffusion

Eddington factor In general, the Eddington factor is a tensor, for 1D systems it is reduced to a scalar.

$$f_\nu = \frac{K_\nu}{J_\nu} = \frac{P_\nu}{E_\nu} \quad (1)$$

- isotropic radiation field
- radiation field strongly peaked in radial (i.e. vertical in cartesian) direction

Radiation transport equations, diffusion, equilibrium

- black body radiation (Planck function $I_\nu = J_\nu = B_\nu$)
- in general, extinction (absorption, scattering) and emission

$$\frac{dI_\nu}{ds} = j_\nu - k_\nu I_\nu \quad (2)$$

– Cartesian coordinates:

$$\boxed{\frac{\partial I_{n,\nu}}{\partial t} \frac{1}{c} + n \nabla I_{n,\nu} = j_\nu - k_{n,\nu} I_{n,\nu}} \quad (3)$$

– spherical coordinates

– 1D-problem with only variation along z-axis $\mu \frac{dI}{dz} = j - kI$

– spherical symmetry $\mu \frac{\partial I}{\partial r} + \frac{1-\mu^2}{r} \frac{\partial I}{\partial \mu} = j - kI$

– plane-parallel approximation

$$\boxed{\mu \frac{dI}{dr} = j - kI} \quad (4)$$

The angle μ is constant throughout the computational domain. Dividing by k_ν , this yields

$$\mu \frac{dI}{k_\nu dr} = \mu \frac{dI}{k_\nu dz} = S - I \quad (5)$$

- 0th moment equation: integrate Equation (3) over ν and Ω , i.e. $\int d\nu d\Omega$. Conservation of energy
- first multiply Equation (3) with $\frac{n}{c}$ and then do integration

Radiative Diffusion Approximation

1. Black-body radiation in perfect equilibrium
2. Radiative transfer equation in the *near-surface* limit.

The approximation is the following: replace $\boxed{I = B}$ or $I_\nu = B_\nu$, once but not twice.

$$I_\nu = B_\nu - \mu \frac{dB_\nu}{k_\nu dz} \quad (6)$$

Derive this equation as a random walk of photons!

5.1 Examples of radiation (diffusion equation)

1. Temperature structure in a static stellar atmosphere
- 2.

5.2 Applications and approximations for radiative forces

- definition of general radiative acceleration vector $g = \frac{1}{\rho c} \int \int n k_\nu I_\nu d\Omega d\nu$
 - continuum Thomson scattering
 - spectral line with extinction
 - * furthermore assume central continuum source
 - * then $g_{line} = \frac{F_\nu^0 k_L}{\rho c}$
- Sobolev approximation
- CAK theory

5.3 Recap

optical depth	optical depth along ray	line optical depth	Sobolev optical depth
$d\tau = k_\nu ds = \sigma_{nu} n ds = \kappa \rho ds$ $\tau_\nu = \int k_\nu ds = \int \sigma_\nu n ds$	$\tau_{\mu,\nu} = \int_z^{z_{max}} \frac{\alpha_{nu}(z')}{\mu} dz' = \frac{\tau_\nu(z)}{\mu}$	$\tau_\nu = \int k_L \phi_\nu dl = \int \kappa \rho ds$	

with

- σ cross-section
- n number density
- κ mass absorption density
- ρ mass density
- k_ν extinction coefficient

6 Overview of symmetry assumptions

plane-parallel	1D atmosphere bounded by horizontal surfaces	
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6.1 General equations - first year overview

6.1.1 Hydrodynamics

Euler equations, together with closing relation (e.g. ideal gas law).

primitive variables			
mass density	velocity	gas energy density	gas pressure
ρ	v	e	p

6.1.2 Radiation

Radiative transfer equation: intensity along a ray while interacting with medium. Photons are massless.

$$\left[\frac{1}{c} \partial_t + \vec{n} \cdot \vec{\nabla} \right] I_\nu = \eta_\nu - \chi_\nu I_\nu \quad (7)$$

frequency	intensity	emissivity	total absorption
ν	I_ν	η_ν	χ_ν

These deliver two equations

- the radiative energy equation (diffusion flux \vec{F})

$$\frac{\partial E}{\partial t} + \vec{\nabla} \cdot \vec{F} = \iint \dots d\nu d\Omega \quad (8)$$

- radiative momentum equation

$$\frac{d\vec{F}}{dt} = \iint \dots \vec{n} d\nu d\Omega \quad (9)$$

(after **integrating over all frequencies**). Depending on the geometry simplifications, one can e.g. integrate over all solid angles.

6.1.3 Radiation-Hydrodynamics

Combination delivers integral-diffusion equation

$$\begin{aligned} \frac{dI}{d\tau} &= S - I \\ &= \int I d\Omega - I \end{aligned} \quad (10)$$

6.1.4 Challenges

- combination with hydrodynamics
- current analysis: simplified geometries (symmetry). E.g. in 2D, an ADI method is used and now also a multigrid method.
- complex geometry difficult to show in ray-tracing scheme
- steady-state vs. time dependent
- focus on radiation equations

7 Introduction: course material (Sundqvist - CMPAA course)

7.1 EXERCISES: Introduction to numerical methods for radiation in astrophysics

1. introduction

2. radiation quantities

- exercise p.3:

- on one hand, we know that $\Delta\epsilon \sim C/r^2$
- on the other hand, from the definition we know that $\Delta\epsilon = I_\nu A_1 A_2 / r^2 \Delta\nu \Delta t$
- combining these equations shows that I_ν is independent from r

- exercise p.4:

–

- exercise 1:

- $F_x = \int_0^\pi \left[I_\nu(\theta) \sin^2(\theta) \int_0^{2\pi} \cos(\phi) \right] d\theta d\phi = 0$
- the same reasoning for $F_y = 0$

- exercise 2:

- the equation follows from $d\mu = d\cos(\theta) = \sin(\theta)d\theta$

- exercise 3:

- isotropic radiation field (i.e. $I(\mu) = I$) then we have $F_\nu = 2\pi \int_{-1}^1 I \mu d\mu = 2\pi I \left. \frac{x^2}{2} \right|_{-1}^1 = 0$

- exercise 4:

- $F_\nu = 2\pi \int_{-1}^1 I(\mu) \mu d\mu = 2\pi \int_{-1}^0 I_\nu^- \mu d\mu + 2\pi \int_0^1 I_\nu^+ \mu d\mu = 2\pi I_\nu^+$

- exercise p.7:

- isotropic radiation field:

- * although the radiation pressure is a tensor, we will denote it as a scalar $P_\nu = \frac{4\pi I_\nu}{c}$

- * the radiation energy density $E_\nu = \frac{12\pi I_\nu}{c}$

- * thus $f_\nu = \frac{1}{3}$

- very strongly peaked in radial direction (beam): $I_\nu = I_0 \delta(\mu - \mu_0)$ with $\mu_0 = 1$

- * pressure tensor $P_{nu} = \frac{1}{c} \int I_0 \delta(\mu - \mu_0) n n d\Omega$

- * energy density $E_\nu = \frac{1}{c} \int I_\nu d\Omega$

- * in this case $P_\nu = E_\nu$ thus $f_\nu = 1$

3. radiation transport vs. diffusion vs. equilibrium

- exercise p. 12: 1D, Cartesian geometry, plane-parallel, frequency-independent and isotropic emission/extinction

- radiation energy equation

- * The equation follows by integrating Equation (4)

- * By definition, $E = \frac{1}{c} \iint I_\nu d\nu d\Omega$

- * thus $\frac{dE}{dr} = \int (j - kI) d\nu d\Omega$ thus $\boxed{\frac{dE}{dr} = \frac{(j - kI)4\pi(\nu_1 - \nu_0)}{c}}$

- * work out the integral taking into account frequency-independent and isotropic coefficients:
 - zeroth momentum equations
 - * One must also take into account the specific form of the flux vector

$$F = \iint I_\nu n d\nu d\Omega = 2\pi \int_{-1}^1 I_\nu(\mu) \mu d\mu$$
 - * thus $\frac{dF}{dr} = \frac{1}{c} \int (j - kI) n d\nu d\Omega$ thus $\frac{dF}{dr} = \frac{(j - kI)4\pi(\nu_1 - \nu_0)n}{c}$
 - first moment equation
 - * similar reasoning
 - * $\frac{dP}{dr} = \int (j - kI) n \cdot n d\nu d\Omega$ thus $\frac{dP}{dr} = \frac{(j - kI)4\pi(\nu_1 - \nu_0)n}{c}$
 - first exercise p. 15
 - $P = \frac{1}{c} \iint I_\nu \mu^2 d\Omega d\nu = \frac{2\pi}{c} \int_{-1}^1 \int_{-1}^1 I_\nu \mu^2 d\mu d\nu = \frac{4\pi}{3c} \int B_\nu d\nu = \frac{aT^4}{3} = \frac{E}{3}$
 - second exercise p.15
 - assuming the diffusion limit,
 - flux-weighted mean opacity $\kappa_F = \frac{\int F_\nu \kappa_\nu d\nu}{\int F_\nu d\nu}$
 - Rosseland mean opacity $\frac{1}{\kappa_R} = \frac{\int_0^\infty \frac{1}{\kappa_\nu} \frac{dB_\nu}{dT} d\nu}{\int_0^\infty \frac{dB_\nu}{dT} d\nu}$.
 - * in the diffusion limit, $F_\nu = -\frac{4\pi}{3} \frac{dB_\nu}{k_\nu dz}$ thus $\frac{dB_{nu}}{dT} =$
 - *
 - third exercise p.15
4. the equations of radiation-hydrodynamics
 5. numerical techniques for the radiative diffusion approximation
 6. applications and approximations for a dynamically important radiative force in supersonic flows
 - exercise p.27: $L_{SOB} = \Delta r = \frac{v_{th}}{dv/dr} = \frac{10[km/s]}{1000[km/s]/R_*} = 0.01 R_*$
 7. Appendix A: properties of equilibrium black-body radiation
 - exercise p. 29
 - this should be satisfied: $B_\nu d\nu = -B_\lambda d\lambda$ and also $\nu = \frac{c}{\lambda}$
 - this is equivalent to saying that $0 = \nu d\lambda + \lambda d\nu$ or $d\lambda = -\frac{\lambda}{\nu} d\nu$ thus $B_\lambda = \frac{\nu}{\lambda} B_\nu$
 - $B_\lambda(T) = \frac{\nu}{\lambda} \frac{2h\nu^3}{(\lambda\nu)^2} \frac{1}{e^{hc/\lambda kT} - 1} = \frac{2h\nu^2}{\lambda^3} \frac{1}{e^{hc/\lambda kT} - 1} = \frac{2hc^2}{\lambda^5} \frac{1}{e^{hc/\lambda kT} - 1}$
 - first exercise p.31
 - derive that $\lambda_{max} T = 2897.8[\mu m K]$
 - ...
 - second exercise p.31
 - this is about the spectra of (unknown) stars
 - first exercise p.32
 - see exercise 7
 - second exercise p.32

- BB radiation: $I_\nu = \frac{2h\nu^3}{c^2} \frac{1}{e^{h\nu/kT} - 1}$
- the radiative flux for isotropic BB radiation is zero. See also exercise 3. This also holds for BB radiation.
- exercise p. 33
 - HR-diagram

8. Appendix B: Simple examples to the radiative transfer equation

- first exercise p. 34
 - start from radiative transport equation $\mu \frac{dI}{ds} = \alpha - \eta I$ in which $\eta = 0$ thus $\boxed{\mu \frac{dI}{ds} = \alpha}$
 - solving the ODE in the general case that $\alpha(s)$ is not constant:
 - * integrate the equation $\mu I = \int_0^D \alpha ds$
 - * ...
 - second exercise p. 34
 - * case $\tau(D) \gg 1$: then $I(D) \approx S$
 - * case $\tau(D) \ll 1$: then $I(D) \approx I(0) + S(1 - 1) = I(0)$
 - first exercise p.35
 - * is the plane-parallel approximation valid for the solar photosphere?
 - second exercise p.35
 - * goal: find a solution to the equation $\mu \frac{dI_\nu}{d\tau_\nu} = I_\nu - S_\nu$ where $I(\tau, \mu)$
 - * solution
- second exercise p.35

9. Appendix C: connecting random walk of photons with radiative diffusion model

- exercise p. 38. Computing the average photon mean-free path inside the Sun.

$$l = \frac{1}{\kappa \rho} = \frac{V_o}{\kappa M_o} [cm]$$
- exercise p.39. Computing the random-walk time (diffusion time) for photons

7.2 Implicit 1D solver (20-11-2018)

See computer code

7.3 ADI 2D Solver

See computer code

7.4 Area of a circle

See computer code

7.5 Limb Darkening

See Section 15.

8 The mathematics of Radiative Transfer

The material in this section is based on the book ...

8.1 Auxiliary mathematics

- $\cos(\Theta) = \cos(\theta) \cos(\theta') + \sin(\theta) \sin(\theta') \cos(\phi - \phi')$

- phase function
$$p(\mu, \phi, \mu', \phi', \tau) = \sum_{n=0}^N \omega_n P_n(\cos(\Theta))$$

- isotropic scattering $p(\tau) = \omega_0(\tau)$

- equation of transfer
$$\mu \frac{\partial I(\tau, \mu, \phi)}{\partial \tau} = I(\tau, \mu, \phi) - \mathcal{S}(\tau, \mu, \phi)$$

with $\mathcal{S}(\tau, \mu, \phi) = B_1(\tau) + \frac{1}{4\pi} \int_{-1}^1 d\mu' \int_0^{2\pi} I(\tau, \mu', \phi') p(\mu, \phi, \mu', \phi') d\phi'$

- axially symmetric with isotropic scattering

$$\mathcal{S}(\tau) = \frac{\omega_0(\tau)}{2} \int_{-1}^1 I(\tau, \mu') d\mu' = B_1(\tau) + \frac{\omega_0(\tau)}{2} \int_0^{\tau_1} \mathcal{S}(t) E_1(|t - \tau|) dt$$

- the Milne equation of the problem $(1 - \omega_0 \bar{\Lambda}) \{ \text{mahtcal} S(t) \} = B(\tau)$

- * solve for $\mathcal{S}(t)$

- * then find $I(\tau, \mu)$

8.2 The H-functions

- characteristic equation

9 Monte Carlo and Radiative Transfer (Puls)

9.1 basic definitions and facts

9.2 about random numbers

9.3 MC integration

9.4 MC simulation

Radiative transfer in stellar atmospheres

- GOAL: spatial radiation energy density $E(\tau)$ in an atmospheric layer
 - only photon-electron scattering
 - τ is the optical depth
- Milne's integral equation
$$E(\tau) = \frac{1}{2} \int_0^\infty E(t) E_1(|t - \tau|) dt$$
 - analytical solution $\frac{E(\tau)}{E(0)} = \sqrt{3}(\tau + q(\tau))$
 - MC simulation
 - * emission angle
 - * optical depth until next scattering event
 - * scattering angle
- HOW DOES THIS WORK?

Algorithm 1 Limb darkening: compute quantity of photons

create photons

probability distribution for emission angle $\mu = \cos(\theta)$: $p(\mu)d\mu = \mu d\mu$

optical depth until next scattering event: $p(\tau)dt \approx e^{-\tau} d\tau$

isotropic scattering angle at low energies: $p(\mu)d\mu \approx d\mu$

follow all photons until they leave the atmosphere or are scattered back into stellar interior

9.5 Exercise 1: RNG

9.6 Exercise 2: Planck-function

1. analytical method
2. MC method

9.7 limb darkening

See section 15.

10 Introduction to Monte Carlo Radiation Transfer (Wood+)

The material is taken from

- (Wood, Wittney, Bjorkman, Wolff - 2001)
- (Wood, Wittney, Bjorkman, Wolff - 2013)

10.1 Elementary principles

specific intensity	I_ν
radiant energy	dE_ν
surface area	dA
angle	θ
solid angle	$d\Omega$
frequency range	$d\nu$
time	dt
flux	F_ν
cross section	σ
scattering angle	χ $\mu = \cos(\chi)$
mean intensity	J
flux	H
radiation pressure	K

intensity	$I_\nu(l) = I_\nu(0)e^{n\sigma l}$
angular phase function of the scattering particle	$P(\cos(\chi))$

inverse method	$\xi = \int_0^{x_0} P(x)dx$ with $\xi \in \mathcal{U}(0, 1)$
rejection method	

10.2 Eddington factors

10.3 Example: plane parallel atmosphere

1. emission of photons: select two angles (3D space). In isotropic scattering

- θ met $\mu = \cos(\theta)$
 - $\mu = 2\xi - 1$ (isotropic scattering)
 - $\mu = \sqrt{\xi}$ (A slab is heated from below. Then $P(\mu) = \mu$)
- $\phi = 2\pi\xi$

2. propagation of photons

- sample optical depth from $\tau = -\log(\xi)$
- distance travelled $L = \frac{\tau z_{max}}{\tau_{max}}$

3. conclusion of emission and propagation

$$\begin{aligned}
x &= x + L \sin(\theta) \cos(\phi) \\
y &= y + L \sin(\theta) \sin(\phi) \\
z &= z + L \cos(\theta)
\end{aligned}
\tag{11}$$

4. Binning: once the photon exists the slab. Produce histograms of the distribution function. Finally, we wish to compute the output flux or the intensity.

I have seen that a newer version of the paper is available, which was also used in these notes (which contains amongst other up-to-date references to code fragments).

A Plane Parallel, Isotropic Scattering Monte Carlo Code

10.4 Monte Carlo Radiative Transfer

From a macroscopic perspective, RT calculations rest on the transfer equation

- emissivity η (how much energy is added to radiation field due to emission)
- opacity χ (how much energy is removed due to absorption)
- the source function $S = \frac{\eta}{\chi}$
- optical depth τ captures the opaqueness of a medium

$$\left(\frac{1}{c} \frac{\partial}{\partial t} + \nabla \cdot \mathbf{n} \right) I = \eta - \chi I \quad (12)$$

$$d\epsilon = I d\nu dt d\Omega dA \cdot n \quad (13)$$

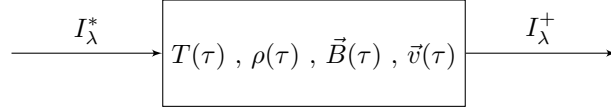
10.5 P Cygni profile for beta-velocity law and given opacity Monte Carlo simulation

10.5.1 Structure of the code

- module common
- module my_inter
- program pcyg
 - INPUT xk0, alpha, beta
 - OUTPUT
 - PROGRAM FLOW: loop over all photons
 - * get xstart and vstart
 - *
 - then do normalisation
- function func(r)
- function xmueout(xk0,alpha,r,v,sigma)
- function rtbis(func,x1,x2,xacc)

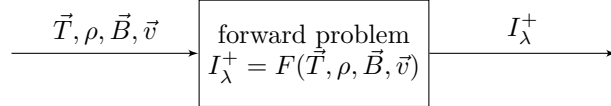
11 Challenges in Radiative Transfer (Ivan Milic)

11.1 Overview of the problem



Forward problem

The forward problem is schematically represented



In fact solve for intensity vector $\vec{I} = \begin{pmatrix} I \\ Q \\ \alpha \\ V \end{pmatrix}$ obeying the equation

$$\frac{d\vec{I}}{d\tau} = -X(\vec{T}, \rho, \vec{B}, \vec{v})\vec{I} - \vec{j}(\vec{T}, \rho, \vec{B}, \vec{v}) \quad (14)$$

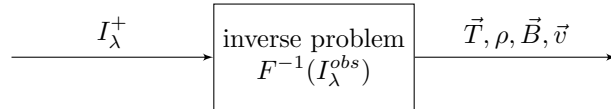
and the solution

$$I_{\lambda}^+ = I_0^+ e^{-\int} + \int \vec{j} e^{-\int} d\tau \quad (15)$$

Example Source function $S = a\tau + b$ then $\int_0^{\tau_{max}} (a\tau + b)e^{-\tau} d\tau = \dots$

Inverse problem

The inverse problem is schematically represented



Via least-squares approximation

$$\min_{\vec{T}, \rho, \vec{B}, \vec{v}} \sum \left(I_{\lambda}^{obs} - I_{\lambda}(\vec{T}, \rho, \vec{B}, \vec{v}) \right)^2 \quad (16)$$

11.2 Challenging domains of application

- Lyman alpha in Galaxy Halos
- Dusty torii (AGD)
- protoplanetary disks
- circumstellar disks
- atmospheres

12 Asymptotic Preserving Monte Carlo methods for transport equations in the diffusive limit (Dimarco+2018)

13 Splitting methods

From notes by professor Frank.

13.1 Exercises

13.1.1 Exercise 1

Part III**Practical work and Exercises****14 Overview of exercises**

1. limb-darkening scattering exercise we did during the course. — You can look into your notes from that, and I attach here also a sample program which you can use as a base. After you have familiarised yourself with this, you can start to think about how you would go about to extend this to a 3D setting (assuming isotropic scattering).
2. (As prep for Monte-Carlo school) here is a script computing a UV resonance P-Cygni line in spherically symmetric wind with v beta-law. At top of routine, a few exercises are given, where you can modify and play around with code. Monte-Carlo program which computes a UV resonance spectral line from a fast outflowing spherically symmetric stellar wind (if you were not cc'd on that email, let me know so that I can send you the files as well). At the top of that little script, there are a few suggestions for exercises (additions) you could do to that program, in order to learn a bit more about the general workings of Monte-Carlo radiative transfer in this context. — So that might be a good idea for you to do as well ! (And you can also ask the others in the group for some tips etc. then.)
3. Some background reading:
 - Attached mc manual by Puls.
 - Paper by Sundqvist+ 2010 (Appendix, I think).

15 Limb darkening

15.0.1 2D Case

We again have $\mu = \cos(\theta)$. The solution of the radiative transfer equation in plane-parallel symmetry with frequency-independent absorption and emission, is

$$I(\mu) = I_1(0.4 + 0.6\mu) \quad (17)$$

In the Monte Carlo code, the photons are sorted according to the direction that they leave the atmosphere.

Goal Calculates the angular dependence of photon's emitted from a plane-parallel, grey atmosphere of radial optical depth **taumax**. The value of **tau** determines the position of the photon

Variables and Algorithm

- **muarray** contains emergent photons
- **na** number of channels
- **dmu** = 1/**na** width of channels
- **nphot** number of photons
- **taumax** maximum optical depth

Algorithm 2 Limb darkening: compute quantitiy of photons

```

initialization
  radial optical depth  $\tau$ 
  direction  $\mu$ 
for all photons do
   $\tau = \tau_{max}$ 
  while  $\tau \geq 0$  do
    compute scattering angle  $\mu$ 
    if  $\tau \geq \text{taumax}$  then  $\mu = \text{sqrt}(x)$  (initial distribution)
    else  $\mu = 2 * x - 1$  (isotropic scattering)
     $\tau_i = -\log(x_2)$ 
     $\tau = \tau - \tau_i * \mu$ 
  end while
  now we know that the photon has left the photosphere
  compute the distribution of all angles  $\mu$  at which the photon left the photosphere
end for
visualisation:
  • plot photon numbers from  $\mu d\mu$  against  $\mu$ 
  • plot specific intensity from  $d\mu$  against  $\mu$  against

```

Figure 1 is according to what is expected $I = I_0(0.4 + 0.6\mu)$

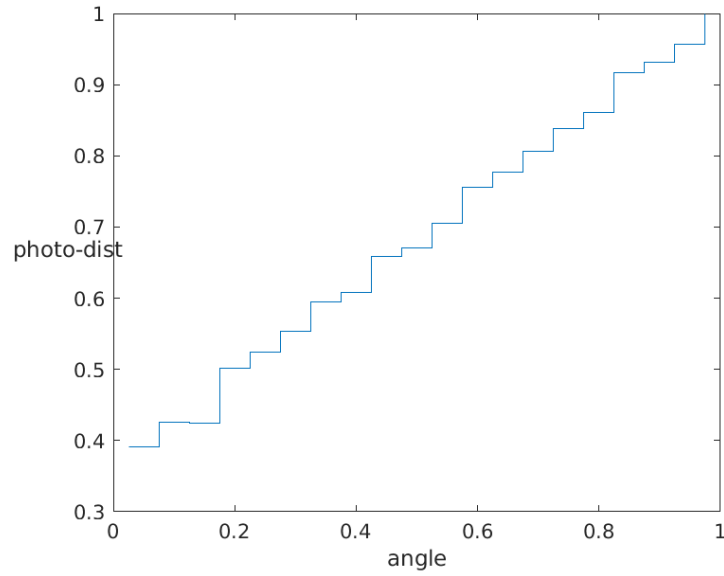


Figure 1: histogram for mu

15.0.2 3D Code

What changes is this:

- introduction of a new angle ϕ
- the optical depth has to be updated according to ϕ also

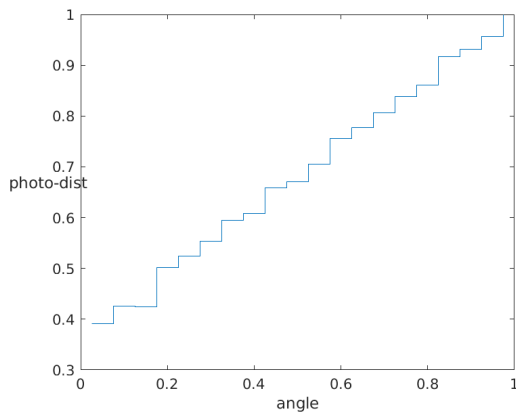


Figure 2: histogram for mu

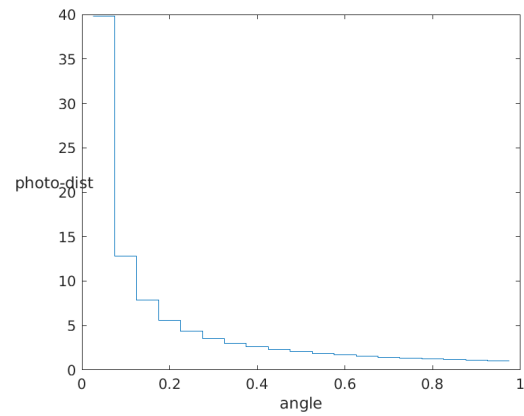


Figure 3: histogram for phi

Figure 2 and Figure 3 are according to what is expected, namely $I = I_0(0.4 + 0.6\mu)$ and a uniform distribution for ϕ , which corresponds to a $I \sim \frac{1}{\phi}$

16 Investigation of program: pcyg.f90

16.1 Overview of variables

name	explanation	scope
paramaters		
xk0		
alpha		
beta		
start frequency of the photon		
xstart	start frequency	
vmin		
vmax		
angle of the photon		
xmuestart	start angle	
xmuein	incident angle	
xmueou	outward angle	
pstart	impact parameter	
xnew	new photon frequency	
optical depth		
tau	optical depth	
number of photons admin		
nphot	number of photons	
nin	photons scattered back into core	
nout	photons escaped	
functions		
func	velocity profile r	distance from center of star
xmueout	sign of outwards angle xk0 alpha r v sigma	

16.2 Exercises (at top of program)

16.2.1 Investigation of original code

In original version of the code, all photons are released radially from photosphere, thus `xmuestart = 1`

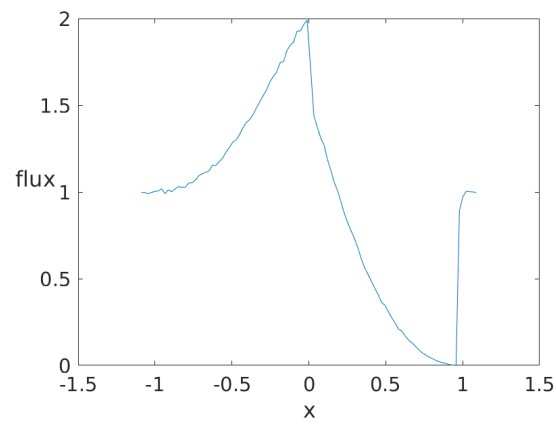


Figure 4: Original version of the code

16.2.2 First adaptation: what if all photons are released radially from photosphere?

What would happen with line-profile, if you assumed all photons were released radially from photosphere?

- test case number 1
- The code sets $mu = 1$. Results in Figure 5.

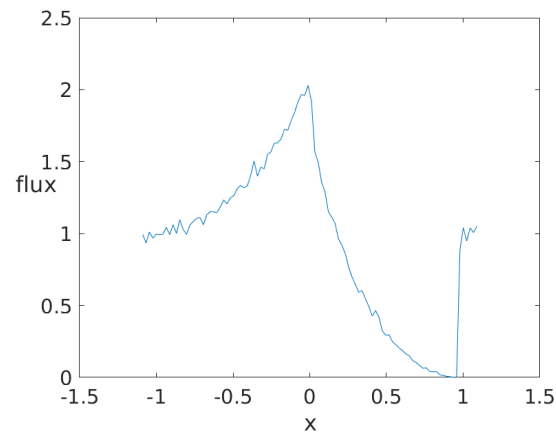


Figure 5: First adaptation

16.2.3 Second adaptation: isotropic scattering

What would happen to line-profile, if you assumed scattering was isotropic (i.e., NOT following Sobolev distribution)

- test case number 2

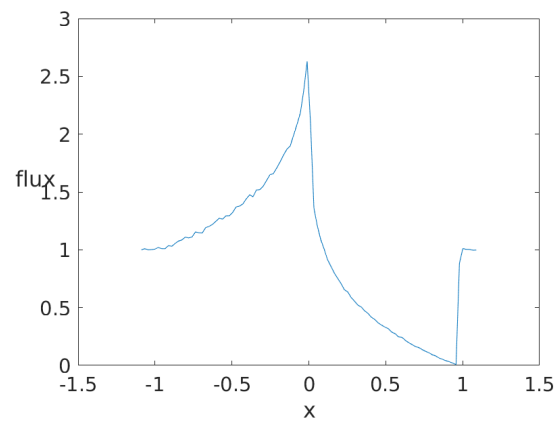


Figure 6: Second adaptation

The *grosso modo* form has not changed, although the scaling has changed.

16.2.4 Third adaptation: introduction of Eddington limb-darkening

Put Eddington limb-darkening in. What happens?

General discussion: Eddington limb darkening The data are taken from Christensen, 2015.

- the source function $S = \langle I \rangle = a + b\tau_\nu$ with $a = \frac{\sigma}{2\pi}T_{eff}^4$ and $b = \frac{3\sigma}{4\pi}T_{eff}^4$
- solve the equation
- this yields $\frac{I(\theta)}{I(0)} = \frac{a + b \cos(\theta)}{a + b} = \frac{2}{5} + \frac{3}{5} \cos(\theta)$

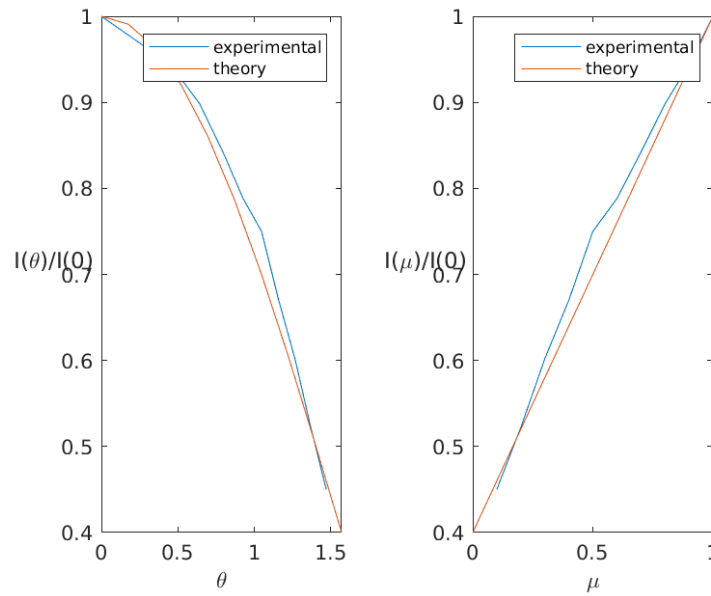


Figure 7: Eddington limb darkening (two times the same plot with $\mu = \cos(\theta)$)

for which star are the experimental data and what assumptions are used in the theory?

Application to exercise

- test case number 3 (not yet implemented)

16.2.5 Fourth adaptaion: photospheric line-profile

Challenging: Put photospheric line-profile (simple Gaussian) in !What happens? Test on $x_{k0}=0$ (opacity =0) case.

- test case number 4 (not yet implemented)

17 Mass loss from inhomogeneous hot star winds (Sundqvist)

- GOAL: synthesis of UV resonance lines from inhomogeneous 2D winds
 - clumped in density
 - clumped in velocity
 - effects of non-void inter-clump medium
- WIND MODELS
 - symmetry assumptions
 - * 1D: spherical symmetry
 - * 2D: symmetry in Φ
 - models
 1. time-dependent radiation-hydrodynamic from Puls and Owocki (POF)
 - * 1D
 - * isothermal flow
 - * perturbations triggered by photospheric sound waves
 2. time-dependent radiation-hydrodynamic from Feldmeier (FPP)
 - * 1D
 - * treatment of energy equation
 - * perturbations triggered by photospheric sound waves or Langevin perturbations (photospheric turbulence)
 3. stochastic model, clumped in density
 - * smooth winds with $v_\beta = (1 - b/r)^\beta$ with $\beta = 1$
 - * clumping factor f_{cl}
 4. stochastic model, clumped in density and in velocity (non-monotonic velocity field)
 - * smooth winds with $v_\beta = (1 - b/r)^\beta$ with $\beta = 1$
 - * clumping factor f_{cl}
- RADIATIVE TRANSFER (MC-2D)

18 Asymptotic preserving Monte Carlo methods for radiative transfer equation in diffusion limit (Dimarco+ 2018)

18.1 Goldstein-Taylor

18.2 Radiative transfer

Part IV

Equation meetings

19 10 April 2019

20 17 April 2019

21 14 August 2019

Literature Study

- General guidelines for good practices in scientific computing are found in [Wil+14].
- I went to the 2019 Petnica Summer school in Petnica, Serbia. This was a good general introduction and overview to astrophysics.
- A very interesting article about Monte Carlo methods for radiative transfer problems, from a mathematical point of view, is [DPS18]. I am currently trying to reproduce the numerical experiments that are reported in the article.

References

- [DPS18] G. Dimarco, L. Pareschi, and G. Samaey. “Asymptotic-Preserving Monte Carlo Methods for Transport Equations in the Diffusive Limit”. eng. In: *SIAM Journal on Scientific Computing* 40.1 (2018), pp. 504–528. issn: 1064-8275.
- [Wil+14] Greg Wilson et al. “Best Practices for Scientific Computing”. In: *PLoS Biology* 12.1 (2014), pp. 1–18. issn: 15449173. doi: 10.1371/journal.pbio.1001745. arXiv: arXiv:1210.0530v4.