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HOMEWORK 3: INTERSTELLAR'S RING

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Contents

1	Setup	2
2	Warm up	2
3	Null geodesics in a Schwarzschild's spacetime	3
4	The observer's frame	3
5	Drawing the observed disk	3
6	Acknowledgement	3

1 Setup

We are interested in describing a Schwarzschild black hole of mass M described by the Schwarzschild coordinates (t', r', θ', ϕ') . In this coordinate system, the singularity of the black hole is approached as $r' \rightarrow 0$ and the event horizon is the null hypersurface with $r' = 2GM$. The rotational symmetry of the system is broken by adding an accretion disk of external radius a orbiting in the $\theta = \pi/2$ equatorial plane of the Schwarzschild coordinates. We are interested in the image of the light created by the accretion disk as seen by an observer at $r' \rightarrow \infty$. This observer will carry a local Minkowski coordinate system (t, x, y, z) . We take $t = t'$ forcing the observer to be static with respect to the black hole (using an infinitesimal acceleration to cancel its geodesic flow towards the black hole). The remaining coordinates x, y, z are related to the asymptotic Schwarzschild coordinates x', y', z' (respectively given by $r' \sin \theta' \cos \phi', r' \sin \theta' \sin \phi'$, and $r' \cos \theta'$) by a rotation. This rotation is fixed by taking $x = x'$ as the rotation axis and orienting the z axis in the direction of the observer (note that this implies that the radius coordinate value $r = \sqrt{x^2 + y^2 + z^2} = r'$ is shared between the two systems, but refer to very different metrics at small r). The angle between z and z' is denoted θ_0 .

The image at $r \rightarrow \infty$ is generated by the collection of light geodesic crossing the xy plane parallel to the z axis. For a given crossing point on the xy -plane we can time reverse (parameter reverse is more appropriate since the geodesic is null) the geodesic motion to see if this point is associated with an accretion disk initial condition and light the crossing point accordingly. The location of the crossing point on the xy plane is identified with an angle $\phi = \arctan_2(y, x)$ and a radius b (the impact parameter of the geodesic at $r \rightarrow \infty$). An important feature of Schwarzschild geodesics is that they are planar: if the geodesic starts at an angle ϕ' on the accretion disk it will stay at the same angle. For a given impact parameter b at angle ϕ , there is a point of maximal approach to the black hole (periastron) with radius $r = r_0(b, \phi)$.

2 Warm up

To better study geodesics around a Schwarzschild black hole, we use a different set of Schwarzschild coordinates (t, r, χ, ψ) (with the same time coordinate as the initial coordinate system) to simplify the calculations in the plane of motion. More precisely the plane of motion constitutes the equatorial plane $\chi = \pi/2$ of the coordinate system. Projecting all points of the geodesic on a constant t slice, expressed in the polar coordinates r, ψ of the equatorial plane. Finally, the $\psi = 0$ direction is aligned with the rays reaching the observer.

- (a) Figure 1 shows sketches of three geodesics in an equatorial plane orthogonal to the accretion disk. The red geodesic has impact parameter b such that the point of the accretion disc plane reached by moving backward in parameter is out of the accretion disc. Geodesics 1 and 2 correspond to light reaching the observer. They have different periastra r_0^1 and r_0^2 and different impact parameters.

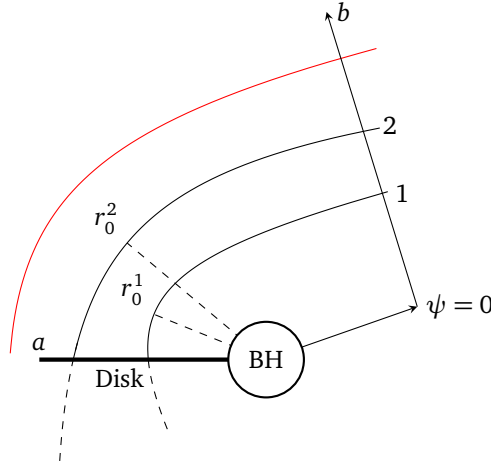


Figure 1: Equatorial plane sketch of geodesic motion with three different impact parameters and periastra . The periastra gets further away from the accretion disk point as we get closer to the outer ring of the accretion disk. If the disk was large, the geodesic starting on the outer ring would approximatively be a straight line starting at the disk point and directed towards the observer. In this limit, the periastron is located exactly in the middle of the line.

(b) Figure 2 compares the total deflexion angle for the two geodesics presented in 1

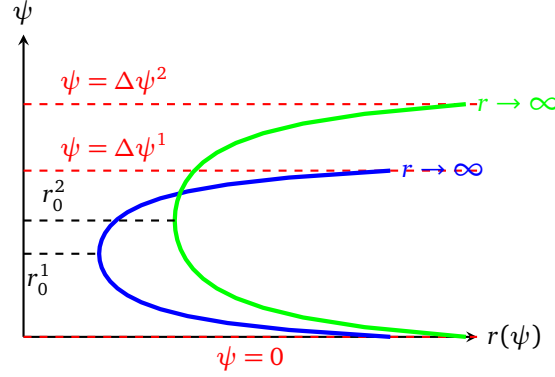


Figure 2: ψ coordinate describing geodesics 1 (blue) and 2 (green) from 1 as a function of the respective radial coordinate r^1 and r^2 . The full geodesics (extended beyond the accretion disc) start at $r \rightarrow \infty$ and end again at $r \rightarrow \infty$. We choose $\psi = 0$ to represent the observer angle for both geodesic forcing them to coincide on the $\psi = 0$ axis. The other asymptote represents the extended initial direction of motion of light consistent with the initial accretion disc point. The minimal value of r (the periastron) is realized in the middle of the geodesic motion. From the dashed lines in 1, we see that the deflexion angle $\Delta\psi^{1,2}$ between $\phi = 0$ and the extended motion asymptote is smaller for geodesic 1 compared to geodesic 2 (we have $\Delta\psi^1 < \Delta\psi^2$). This is consistent with the increasing gravitational pull with decreasing periastron.

(c)

3 Null geodesics in a Schwarzschild's spacetime

4 The observer's frame

5 Drawing the observed disk

6 Acknowledgement

Chat GPT was used to produce the initial code for the figures in the Warm up section.