

Monochromatic method for determining T_{\odot} .

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Measuring the Sun is important for understanding the ways that stars function, and allows us to make similar measurements for stars that are much further away. One important characteristic is the temperature of the sun T_{\odot} . Ensuring high-precision measurements of this value are paramount to the field of stellar astronomy. This study sought out to find the effectiveness of measuring T_{\odot} in a fashion different than is currently accepted within literature. We found a measurement of $T_{\odot} = 5050 \text{ K} \pm 12.5 \text{ K}$, suggesting that this method should not be adopted by the field for determining T_{\odot} .

I. INTRODUCTION

The Sun is of particular interest because it is the closest star. Investigation of the properties of the Sun allow for a deeper understanding of stars. One particularly important variable is the Sun's temperature. The temperature of the Sun is directly related to the luminosity of the sun.

A. Black Bodies

A black body is a hypothetical body that absorbs all incident light. At equilibrium, this means that any and all radiation that exits the black body are due to thermal radiation. Stars act as great approximations of black bodies. As such, the various equations that describe black bodies can be applied to stars including the Sun.

The radiant flux of a black body, j^* , can be modeled with the Stefan-Boltzmann Law [1],

$$j^* = \sigma T^4, \quad [\text{W m}^{-2}] \quad (1)$$

where σ is the Stefan-Boltzmann constant and T is the temperature of the black body.

Black bodies do not emit equally at all wavelengths of light as can be seen in figure 1. The energy emitted by a black body as a function of observed wavelength can be describe by Wien's Approximation [2],

$$E_{\lambda} = C_1 \lambda^{-5} \exp(-C_2/\lambda T), \quad (2)$$

where $C_1 = 2\pi h c^2 = 3.74 \times 10^{-16} \text{ W m}^{-2}$ and $C_2 = hc/k = 1.44 \times 10^{-2} \text{ m K}$.

Using equation 2, the temperature of two objects can be compared by observing emitted energies at one wavelength,

$$\left(\frac{1}{T_1} - \frac{1}{T_2} \right) = \frac{\lambda}{C_2} \ln \left(\frac{E_2}{E_1} \right). \quad (3)$$

This equation is the basis by which optical pyrometry is done.

B. Radiometry

Radiometry is the practice of measuring the intensity of light. In radiometry, the tool used for finding the temperature of an object is a pyrometer. Optical pyrometers function by comparing the temperature of a known source to the temperature of a source you are attempting to measure by utilizing equation 3 [3]. The specifics of the instrument used in this experiment will be discussed in more detail in section II.

C. Calculation of T_{\odot}

The accepted value for the temperature of the sun T_{\odot} , by the International Astronomical Union (IAU) is found using the Stefan-Boltzmann's law with the best estimates for the solar photospheric radius R_{\odot} , the astronomical unit, and the total solar irradiance S_{\odot} [5]. S_{\odot} is the mean total electromagnetic radiation from the sun, integrated across all wavelengths. S_{\odot} is measured using an absolute cavity radiometer [6]. One of the largest differences between the absolute cavity radiometer and a pyrometer is that the radiometer measures the intensity of light, whereas a pyrometer "interprets" the measured intensity and reads out a temperature. This study seeks to look at a monochromatic ground-based determination of T_{\odot} through the use of an optical pyrometer. The results of this study will inform us on the effectiveness of this method in comparison to the currently accepted method.

II. METHODS

The schematic for the apparatus used within this experiment can be seen in figure 2. The apparatus is housed within a hand held instrument. In order to collect data we first had to point the instrument at the sun. At this point, a filter was added that allowed for us to read higher temperatures.

In addition, a post hoc filter was added to the instrument that increased every temperature read out by about 3000 K. Without the second filter, it would not have been

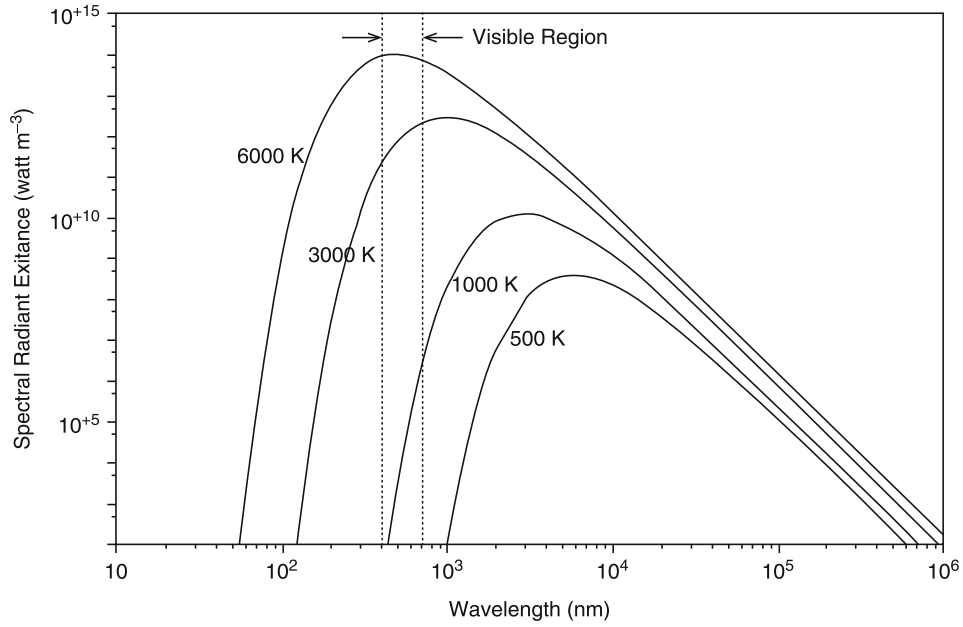


FIG. 1. Black body spectra. The shape of the black body spectra for bodies at varying temperatures. [4]

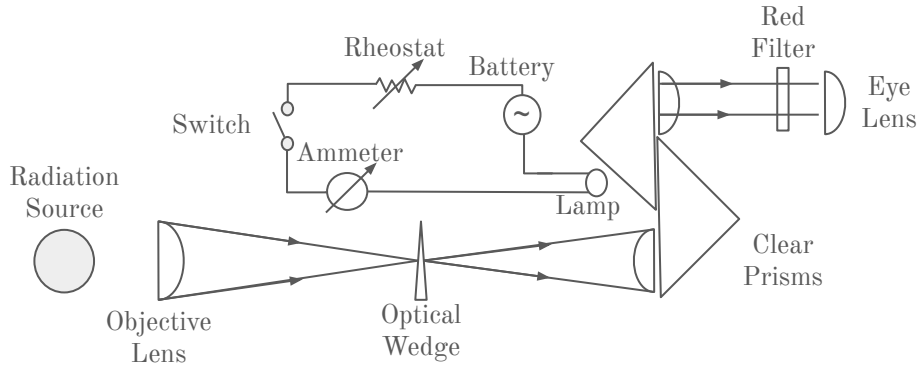


FIG. 2. Experimental Apparatus. Incoming radiation is focused by the objective lens. Then using an optical wedge the operator is able to lower the intensity of light. The operator manipulates the optical wedge until the incoming light matches intensity with an internal lamp with a known temperature. The instrument then reads out a measured temperature.

possible to collect a temperature for an object as hot as the Sun. 12.5 K.

III. RESULTS

Little scientifically viable data has been able to be collected as of right now because of the issues with being able to figure out how to operate the pyrometer used in this experiment. Two measurement were taken briefly that gave readings of T_{\odot} as 5000 ± 12.5 K and 5100 ± 12.5 K, giving $\bar{T}_{\odot} = 5050 \pm 12.5$ K. Since there were no calculations done other than averaging for this experiment, there is no need to propagate uncertainty. The uncertainty of our measurements were half the distance between each tick on the pyrometers readout, which cooresponds to

IV. CONCLUSIONS

The results from this study vary largely from widely accepted values of T_{\odot} . The IAU's nominal value for T_{\odot} is 5772 K. Several factor play into why the nominal value falls outside of uncertainty of the value that we calculated in this experiment.

One possible reason could be due to the instrument itself. Pyrometers can only calculate a valid temperature when the object being measured takes up the entire field of view. With the instrument used in the experiment, at the highest level of magnification, the Sun did not take up the entire field of view. The reason this matters is

because the Pyrometer itself does not actually account for distance. The way that it compensates has to do with the fact that both solid angle and radiant flux fall off as r^{-2} . Ensuring that the object takes up the whole field of view allows for these two values to cancel each other out.

Another possible source of error could be the atmospheric extinction. As light travels through the atmosphere some of it is absorbed. Without accounting for this, it will appear that the Sun is colder than it really

is.

A final source of error could be the monochromatic nature of the instrument. The pyrometer used in this experiment measures light at 655 nm, whereas the sun's spectra peaks at around 500 nm, this also suggests a lower temperature than what we've previously measured.

In all, the results of this experiment suggest that this method of calculating the temperature of the Sun is quite flawed. With the use of a radiometer that reads bolometric intensity, many of the issues found in this experiment could be solved.

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