

Mie Scattering and Absorption (Version 1)

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Abstract

A set of Mie functions has been developed in Python to compute the four Mie coefficients a_n , b_n , c_n and d_n , efficiencies of extinction, scattering, backscattering and absorption, the asymmetry parameter, and the two angular scattering functions S_1 and S_2 . In addition to the scattered field, also the absolute-square of the internal field is computed and used to get the absorption efficiency in a way independent from the scattered field. This allows to test the computational accuracy. This first version of MATLAB Mie Functions is limited to homogeneous dielectric spheres without change in the magnetic permeability between the inside and outside of the particle. Required

input parameters are the complex refractive index, $m = m' + im''$, of the sphere (relative to the ambient medium) and the size parameter, $x = ka$, where a is the sphere radius and k the wave number in the ambient medium.

1 Introduction

This report is a description of Mie-Scattering and Mie-Absorption programs written in the numeric computation and visualisation software, MATLAB (Math Works, 1992), for the improvement of radiative-transfer codes, especially to account for rain and hail in the microwave range and for aerosols and clouds in the submillimeter, infrared and visible range. Excellent descriptions of Mie Scattering were given by van de Hulst (1957) and by Bohren and Huffman (1983). The present programs are related to the formalism of Bohren and Huffman (1983). In addition an extension (Section 2.5) is given to describe the radial dependence of the internal electric field of the scattering sphere and the absorption resulting from this field. Except for Section 2.5, equation numbers refer to those in Bohren and Huffman (1983), in short BH, or in case of missing equation numbers, page numbers are given. For a description of computational problems in the Mie calculations, see the notes on p. 126-129 and in Appendix A of BH.

2 Formulas for a homogeneous sphere

2.1 Mie coefficients and Bessel functions

Python function: mie_abcd

The key parameters for Mie calculations are the Mie coefficients a_n and b_n to compute the amplitudes of the scattered field, and c_n and d_n for the internal field, respectively. The computation of these parameters has been the most challenging part in Mie computations due to the involvement of spherical Bessel functions up to high order. With MATLAB's built-in double-precision Bessel functions, the computation of the Mie coefficients has so far worked well up to size parameters exceeding 10,000; the coefficients are given in BH

on p.100:

$$a_n = \frac{m^2 j_n(mx)[xj_n(x)]' - \mu_1 j_n(x)[mxj_n(mx)]'}{m^2 j_n(mx)[xh_n^{(1)}(x)]' - \mu_1 h_n^{(1)}(x)[mxj_n(mx)]'} \quad (1)$$

$$b_n = \frac{\mu_1 j_n(mx)[xj_n(x)]' - j_n(x)[mxj_n(mx)]'}{\mu_1 j_n(mx)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[mxj_n(mx)]'} \quad (2)$$

$$c_n = \frac{\mu_1 j_n(x)[xh_n^{(1)}(x)]' - \mu_1 h_n^{(1)}(x)[xj_n(x)]'}{\mu_1 j_n(mx)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[mxj_n(mx)]'} \quad (3)$$

$$d_n = \frac{\mu_1 m j_n(x)[xh_n^{(1)}(x)]' - \mu_1 m h_n^{(1)}(x)[xj_n(x)]'}{m j_n(mx)[xh_n^{(1)}(x)]' - \mu_1 h_n^{(1)}(x)[mxj_n(mx)]'} \quad (4)$$

where m is the refractive index of the sphere relative to the ambient medium, $x = ka$ is the size parameter, a the radius of the sphere and $k = 2\pi/\lambda$ is the wave number and λ the wavelength in the ambient medium. In deviation from BH, μ_1 is the ratio of the magnetic permeability of the sphere to the magnetic permeability of the ambient medium (corresponding to μ_1/μ in BH). The functions $j_n(z)$ and $h_n^{(1)}(z) = j_n(z) + iy_n(z)$ are spherical Bessel functions of order n ($n = 1, 2, \dots$) and of the given arguments, $z = x$ or mx , respectively, and primes mean derivatives with respect to the argument. The derivatives follow from the spherical Bessel functions themselves, namely

$$\begin{aligned} [zj_n(z)]' &= zj_{n-1}(z) - nj_n(z) \\ [zh_n^{(1)}(z)]' &= zh_{n-1}^{(1)}(z) - nh_n^{(1)}(z) \end{aligned} \quad (5)$$

For completeness, the following relationships between Bessel and spherical Bessel functions are given:

$$j_n(z) = \sqrt{\frac{\pi}{2z}} J_{n+0.5}(z) \quad (6)$$

$$y_n(z) = \sqrt{\frac{\pi}{2z}} Y_{n+0.5}(z) \quad (7)$$

Here, J_ν and Y_ν are Bessel functions of the first and second kind. For $n = 0$ and 1 the spherical Bessel functions are given (BH, p. 87) by

$$\begin{aligned} j_0(z) &= \sin z/z \\ j_1(z) &= \sin z/z^2 - \cos z/z \\ y_0(z) &= -\cos z/z \\ y_1(z) &= -\cos z/z^2 - \sin z/z \end{aligned} \quad (8)$$

and the recurrence formula

$$f_{n-1}(z) + f_{n+1}(z) = \frac{2n+1}{z} f_n(z) \quad (9)$$

where f_n is any of the functions j_n and y_n . Taylor-series expansions for small arguments of j_n and y_n are given on p. 130 of BH. The spherical Hankel functions are linear combinations of j_n and y_n . Here, the first type is required

$$h_n^{(1)}(z) = j_n(z) + iy_n(z) \quad (10)$$

The following related functions are also used in Mie theory (although we try to avoid them here):

$$\begin{aligned} \psi_n(z) &= zj_n(z) \\ \chi_n(z) &= -zy_n(z) \\ \xi_n(z) &= zh_n^{(1)}(z) \end{aligned} \quad (11)$$

Often $\mu_1 = 1$; then, (4) simplify to

$$a_n = \frac{m^2 j_n(mx)[xj_n(x)]' - j_n(x)[mxj_n(mx)]'}{m^2 j_n(mx)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[mxj_n(mx)]'} \quad (12)$$

$$b_n = \frac{j_n(mx)[xj_n(x)]' - j_n(x)[mxj_n(mx)]'}{j_n(mx)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[mxj_n(mx)]'} \quad (13)$$

$$c_n = \frac{j_n(x)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[xj_n(x)]'}{j_n(mx)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[mxj_n(mx)]'} \quad (14)$$

$$d_n = \frac{mj_n(x)[xh_n^{(1)}(x)]' - mh_n^{(1)}(x)[xj_n(x)]'}{m^2 j_n(mx)[xh_n^{(1)}(x)]' - h_n^{(1)}(x)[mxj_n(mx)]'} \quad (15)$$

The parameters used in radiative transfer depend on a_n and b_n , but not on c_n and d_n . The latter coefficients are needed when the electric field inside the sphere is of interest, e.g. to test the field penetration in the sphere, to study the distribution of heat sources or to compute absorption. The absorption efficiency Q_{abs} , however, can also be computed from the scattered radiation, Equation 18 and Equation 20 to be shown below.

2.2 Mie efficiencies and cross sections

Python functions: mie, mie_xscan The efficiencies Q_i for the interaction of radiation with a scattering sphere of radius a are cross sections σ_i (called C_i

in BH) normalised to the particle cross section, πa^2 , where i stands for extinction ($i = \text{ext}$), absorption ($i = \text{abs}$), scattering ($i = \text{sca}$), backscattering ($i = b$), and radiation pressure ($i = \text{pr}$), thus

$$Q_i = \frac{\sigma_i}{\pi a^2} \quad (16)$$

Energy conservation requires that

$$Q_{\text{ext}} = Q_{\text{sca}} + Q_{\text{abs}} \quad (17)$$

$$\sigma_{\text{ext}} = \sigma_{\text{sca}} + \sigma_{\text{abs}} \quad (18)$$

The scattering efficiency Q_{sca} follows from the integration of the scattered power over all directions, and the extinction efficiency Q_{ext} follows from the Extinction Theorem (Ishimaru, 1978, p. 14, van de Hulst, 1957, p. 31), also called Forward-Scattering Theorem, leading to:

$$Q_{\text{sca}} = \frac{2}{x^2} \sum_{n=1}^{\infty} (2n+1)(|a_n|^2 + |b_n|^2) \quad (19)$$

$$Q_{\text{ext}} = \frac{2}{x^2} \sum_{n=1}^{\infty} (2n+1)\text{Re}(a_n + b_n) \quad (20)$$

and Q_{abs} follows from (18). All infinite series can be truncated after n_{max} terms. For this number Bohren and Huffman (1983) proposed the value

$$n_{\text{max}} = x + 4x^{1/3} + 2 \quad (21)$$

and this value is used here as well. Furthermore, the asymmetry parameter $g = \langle \cos \theta \rangle$ indicates the average cosine of the scattering angle θ with respect to power; it is used in Two-Stream Models (Meador and Weaver, 1980), and it is related to the efficiency Q_{pr} of radiation pressure:

$$Q_{\text{pr}} = Q_{\text{ext}} - Q_{\text{sca}} \langle \cos \theta \rangle \quad (22)$$

$$Q_{\text{sca}} \langle \cos \theta \rangle = \frac{4}{x^2} \left\{ \sum_{n=1}^{\infty} \frac{n(n+2)}{n+1} \text{Re}(a_n a_{n+1}^* + b_n b_{n+1}^*) + \sum_{n=1}^{\infty} \frac{2n+1}{n(n+1)} \text{Re}(a_n b_n^*) \right\} \quad (23)$$

Finally, the backscattering efficiency Q_b , applicable to monostatic radar, is given by

$$Q_b = \frac{1}{x^2} \left| \sum_{n=1}^{\infty} (2n+1)(-1)^n (a_n - b_n) \right|^2 \quad (24)$$

2.3 The scattered far field

Python functions: `mie_s12`, `mie_pt`, `mie_thetascan` If the detailed shape of the angular scattering pattern is required, e.g. to get the phase matrix or phase function for radiative-transfer calculations (Chandrasekhar, 1960), the scattering functions S1 and S2 are required. These functions describe the scattered field \mathbf{E}_s . The scattered far field in spherical coordinates ($E_{s\theta}$, $E_{s\phi}$) for a unit-amplitude incident field (where the time variation $\exp(-i\omega t)$ has been omitted) is given by