

1. Introduction

2. Automation Program

The rc3 class creates an rc3 object containing its “basic info”¹ for each galaxy in the RC3 catalog. The practical purpose of the class (other than the use of OOP that actually makes sense) is to keep track of the number of iterations in the recursive step using the instance variable `num_iterations`.

Some parts of the program needs to be adjusted for survey-specific, but the core concept (and bulk of the code) should stay the same.

2.1. Technical Details

The program was written in Python 2.7.6 to survey-independent and interacts It uses IPAC’s Montage [1] using the AstroPy Montage wrapper². The final g,r,i image is created using Astromatic STIFF [2]. Our choice of program reflects best feature from both program : STIFF provides the flexibility of changing many variables for the final g,r,i; Montage creates scientifically-calibrated images by retaining astrometric accuracy during image reprojection. The use of two different program in the mosaicing step is due to Montage’s ability to create scientifically-calibrated mosaic FITS, but STIFF provided more flexible parameters for combining all bands into color images, so we get the best of both worlds. The STIFF setting needs to be adjusted for each survey depending on specs on the telescope’s CCD dependent parameters such as imaging bands. (?) The source extraction is done using SExtractor. If the mosaicing field is chosen correctly, then SExtractor’s skylevel estimation is fairly accurate.

2.2. Getting the Data

2.2.1. Search

Figure out how to convert positional values (ra,dec) to record-keeping parameters dependent on the survey’s telescope. (tiles, frames...etc)³ Usually this step can be done using the SQL search. Most surveys have an API that enables you to access data using SQL queries so that the mosaicing program can interact with so that it doesn’t have to click buttons or type in textboxes in the web GUI.

3. Algorithms

4. Pipeline

4.1. Class Hierarchy

5. Result from SDSS run

5.1. Known Errors

Even though a series of exception handling and error prevention mechanisms were put in place,

- mProjectExec is the montage Mosaic procedure that creates the reprojected image from the raw FITS files. Sometimes reprojected images are not created even when Montages’ debug statement clearly shows that the reprojection was successful and table and header files are corrupted. This results in an error in later mosaicing steps. We have implemented error prevention mechanism to ensure that mosaic procedures terminates correctly in such cases and wrote the problematic galaxy into failed_projection, which can be examined later.
- Sometimes in the step “apply background model to images” takes a long time and produce huge files in coordir ()

5.2. Performance

accelerate the process by “testing” only on the a single band fit file. For SDSS, that is the r band, best quality (longest wavelength =i, most photons). Most of the time is spent on downloading the raw FITS files from the survey’s specific server. This process can be dramatically sped up if the investigator has imaging data stored locally on a hardisk. We anticipate that this consist of simply write a subclass of the Server class and fill in some path-dependent details (where images are stored) to enable this feature. Computation time significantly increases The Montage mBgExec takes a long time .

6. Algorithm

7. Conclusion

8. Acknowledgements

This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

This research made use of Astropy, a community-developed core Python package for Astronomy (Astropy Collaboration, 2013).

Funding for SDSS-III has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, and the U.S. Department of Energy Office of Science. The SDSS-III web site is <http://www.sdss3.org/>.

SDSS-III is managed by the Astrophysical Research Consortium for the Participating Institutions of the SDSS-III Collaboration including the University of Arizona, the Brazilian Participation Group, Brookhaven National Laboratory, University of Cambridge, Carnegie Mellon University, University of Florida, the French Participation Group, the German Participation Group, Harvard University, the Instituto de Astrofísica de Canarias, the Michigan State/Notre Dame/JINA Participation Group, Johns Hopkins University, Lawrence Berkeley National Laboratory, Max Planck Institute for Astrophysics, Max Planck Institute for Extraterrestrial Physics, New Mexico State

¹ra,dec,radius, pgc

²<http://www.astropy.org/montage-wrapper/>

³Since SDSS images are taken in drift-scan mode, we need the run,camcol,field values to access each image frame

University, New York University, Ohio State University, Pennsylvania State University, University of Portsmouth, Princeton University, the Spanish Participation Group, University of Tokyo, University of Utah, Vanderbilt University, University of Virginia, University of Washington, and Yale University

- [1] G. B. Berriman, E. Deelman, J. C. Good, J. C. Jacob, D. S. Katz, C. Kesselman, A. C. Laity, T. A. Prince, G. Singh, and M.-H. Su. Montage: a grid-enabled engine for delivering custom science-grade mosaics on demand. In P. J. Quinn and A. Bridger, editors, *Optimizing Scientific Return for Astronomy through Information Technologies*, volume 5493 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pages 221–232, Sept. 2004. doi: 10.1117/12.550551.
- [2] E. Bertin. Displaying Digital Deep Sky Images. In P. Ballester, D. Egret, and N. P. F. Lorente, editors, *Astronomical Data Analysis Software and Systems XXI*, volume 461 of *Astronomical Society of the Pacific Conference Series*, page 263, Sept. 2012.