

# Ryden Solutions

Rio Weil

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Solutions to problems from Ryden's "Introduction to Cosmology", 2nd ed.

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# 1 Introduction

## 2 Fundamental Observations

### 2.1 Human blackbody radiation

The energy density of blackbody radiation is given by:

$$\epsilon_\gamma = \alpha T^4 \quad (2.1)$$

Approximating a human being as a sphere with volume  $1\text{m}^3$ , and considering that photons travel at speed  $c$  we have that the rate of energy radiation is:

$$E = cA\epsilon_\gamma = cA\alpha T^4 \sim 2.10 \times 10^3 \text{W} \quad (2.2)$$

### 2.2 CMB photon rate

We can calculate the photon number density of the CMB to be:

$$n_\gamma = \beta T^3 = (2.029 \times 10^7 \text{m}^{-3} \text{K}^{-3})(2.8255 \text{K})^3 = 4.107 \times 10^8 \text{m}^{-3} \quad (2.3)$$

Approximating our body to be a perfect sphere with cross sectional area  $A \sim 1\text{m}^2$ , since photons travel at a rate  $\text{cm s}^{-1}$  the rate at which they pass through us is:

$$r \sim cAn_\gamma = (3.00 \times 10^8 \text{m s}^{-1})(1\text{m}^2)(4.107 \times 10^8 \text{m}^{-3}) = 1.23 \times 10^{17} \text{s}^{-1} \quad (2.4)$$

### 2.3 How long for the CMB to warm you up?

The energy per photon of the CMB is:

$$\frac{\epsilon_\gamma}{n_\gamma} = 1.02 \times 10^{-22} \text{J} \quad (2.5)$$

The energy required to raise my temperature by  $1\text{nK}$  is:

$$\Delta E = Cm\Delta T = (4200 \text{J kg}^{-1} \text{K}^{-1})(50 \text{kg})(10^{-9} \text{K}) = 2.1 \times 10^{-4} \text{J} \quad (2.6)$$

So solving for the time to heat up by 1 nanoKelvin:

$$t_{1\text{nK}} = \frac{\Delta E}{r \frac{\epsilon_\gamma}{n_\gamma}} \sim 16.8 \text{s} \quad (2.7)$$

### 2.4 Tired Light

We start with the energy loss per unit distance proposed by the “tired light hypothesis”:

$$\frac{dE}{dr} = -kE \quad (2.8)$$

This is the all-too-famous exponential decay ODE, which has solution:

$$E(r) = C \exp(-kr) \quad (2.9)$$

In principle we could use separation of variables to solve the ODE, but let's just verify that this is the correct solution:

$$\frac{d}{dr} (C \exp(-kr)) = -kC \exp(-kr) = -kE \quad \checkmark \quad (2.10)$$

Letting  $E(0) = E_0$  be the energy of the photon when emitted (before it has travelled and lost energy), we get:

$$E_0 = E(0) = C \exp(-k(0)) = C \implies C = E_0 \quad (2.11)$$

Therefore:

$$E(r) = E_0 \exp(-kr) \quad (2.12)$$

Now, using the energy-momentum relation and the DeBroglie wavelength relation, we have:

$$E = pc = \frac{hc}{\lambda} \quad (2.13)$$

So substituting this into (2.12) we get:

$$\frac{hc}{\lambda_r} = \frac{hc}{\lambda_0} \exp(-kr) \quad (2.14)$$

Where  $\lambda_r$  is the observed/measured photon wavelength at distance  $r$  from the source and  $\lambda_0$  is the photon wavelength measured at the source. Rearranging we obtain:

$$\frac{\lambda_r}{\lambda_0} = \exp(kr) \quad (2.15)$$

Now we recall the definition of redshift:

$$z = \frac{\lambda_r - \lambda_0}{\lambda_0} = \frac{\lambda_r}{\lambda_0} - 1 \quad (2.16)$$

Substituting this into (2.15) we get:

$$z = \exp(kr) - 1 \quad (2.17)$$

Which adding one and taking the logarithm of both sides we get:

$$\log(1 + z) = kr \quad (2.18)$$

If  $z \ll 1$ , then by Taylor expanding to first order we obtain:

$$\log(1 + z) \approx z \quad (2.19)$$

So in this limit we have:

$$\boxed{z = kr} \quad (2.20)$$

From which we see a linear distance-redshift relation. For the last part of the question, we recall Hubble's Law:

$$z = \frac{H_0}{c} r \quad (2.21)$$

So comparing this with (2.20), the value of  $k$  to yield a Hubble constant of  $H_0 = 68 \text{ km s}^{-1} \text{ Mpc}^{-1}$  must be:

$$\boxed{k = \frac{H_0}{c} = 2.3 \times 10^{-4} \text{ Mpc}^{-1}} \quad (2.22)$$

## 2.5 CMB Cutoff - The Cosmic Infrared background

We start with the number density for CMB photons:

$$n(f)df = \frac{\epsilon(f)df}{hf} = \frac{8\pi}{c^3} \frac{f^2 df}{\exp(hf/kT) - 1} \quad (2.23)$$

Since  $hf > E_0 \gg kT$ , we have that  $\exp(hf/kT) \gg 1$  and so  $\exp(hf/kT) - 1 \sim \exp(hf/kT)$ . This yields:

$$n(f)df \approx \frac{8\pi}{c^3} f^2 \exp\left(-\frac{hf}{kT}\right) df \quad (2.24)$$

Now, to get  $n(hf > E_0)$  we integrate this expression from  $E_0/h$  to  $\infty$ .

$$n(hf > E_0) = \int_{E_0/h}^{\infty} \frac{8\pi}{c^3} f^2 \exp\left(-\frac{hf}{kT}\right) df \quad (2.25)$$

Integrating by parts (with  $u = f^2, dv = \exp(-hf/kT)$ ), we have:

$$n(hf > E_0) = \frac{8\pi}{c^3} \left( f^2 \frac{-kT}{h} \exp\left(-\frac{hf}{kT}\right) \Big|_{E_0/h}^{\infty} - \int_{E_0/h}^{\infty} 2f \frac{-kT}{h} \exp\left(-\frac{hf}{kT}\right) df \right) \quad (2.26)$$

The term at infinity goes to zero, so:

$$n(hf > E_0) = \frac{8\pi}{c^3} \left( \frac{kTE_0^2}{h^3} \exp\left(-\frac{E_0}{kT}\right) + \frac{2kT}{h} \int_{E_0/h}^{\infty} f \exp\left(-\frac{hf}{kT}\right) df \right) \quad (2.27)$$

We now integrate by parts again (with  $u = f, dv = \exp(-hf/kT)$ ) to get:

$$n(hf > E_0) = \frac{8\pi}{c^3} \left( \frac{kTE_0^2}{h^3} \exp\left(-\frac{E_0}{kT}\right) + \frac{2kT}{h} \left( f \frac{-kT}{h} \exp\left(-\frac{hf}{kT}\right) \Big|_{E_0/h}^{\infty} - \int_{E_0/h}^{\infty} \frac{-kT}{h} \exp\left(-\frac{hf}{kT}\right) df \right) \right) \quad (2.28)$$

Again the term at infinity goes to zero and we get:

$$n(hf > E_0) = \frac{8\pi}{c^3} \left( \frac{kTE_0^2}{h^3} \exp\left(-\frac{E_0}{kT}\right) + \frac{2kT}{h} \left( \frac{kTE_0}{h^2} \exp\left(-\frac{E_0}{kT}\right) + \frac{kT}{h} \int_{E_0/h}^{\infty} \exp\left(-\frac{hf}{kT}\right) df \right) \right) \quad (2.29)$$

Finally the last integral is easy:

$$n(hf > E_0) = \frac{8\pi}{c^3} \left( \frac{kTE_0^2}{h^3} \exp\left(-\frac{E_0}{kT}\right) + \frac{2kT}{h} \left( \frac{kTE_0}{h^2} \exp\left(-\frac{E_0}{kT}\right) + \frac{kT}{h} \left( \frac{-kT}{h} \exp\left(-\frac{hf}{kT}\right) \Big|_{E_0/h}^{\infty} \right) \right) \right) \quad (2.30)$$

So after this tedious calculation, we have:

$$n(hf > E_0) = \frac{8\pi}{c^3} \exp\left(-\frac{E_0}{kT}\right) \left[ \frac{kTE_0^2}{h^3} + \frac{2k^2T^2E_0}{h^3} + \frac{2k^3T^3}{h^3} \right] \quad (2.31)$$

Using that  $E_0 \gg kT$  again, we can neglect all but the first term (so we could have really avoided the latter two integration steps, but alas):

$$n(hf > E_0) \approx \frac{8\pi kTE_0^2}{c^3 h^3} \exp\left(-\frac{E_0}{kT}\right) \quad (2.32)$$

Now taking the ratio of this with  $n_\gamma$ :

$$\frac{n(hf > E_0)}{n_\gamma} \approx \frac{\frac{8\pi kTE_0^2}{c^3 h^3} \exp\left(-\frac{E_0}{kT}\right)}{\frac{2.4041}{\pi^2} \frac{k^3}{h^3 c^3} T^3} = 0.42 \left(\frac{E_0}{kT}\right)^2 \exp\left(-\frac{E_0}{kT}\right) \quad (2.33)$$

which was the desired formula. Next, we calculate the fraction of “CMB” photons that are actually far-IR photons. We first consider the wavelength-frequency relation for light:

$$c = \lambda f \quad (2.34)$$

so  $\lambda < 1\text{mm}$  corresponds to  $f > 3 \times 10^{11}\text{Hz}$  and so:

$$E = hf > 2 \times 10^{-22}\text{J} \quad (2.35)$$

So using the above derived relation with  $T = 2.7255\text{K}$  we have:

$$\frac{n(hf > 2 \times 10^{-22}\text{J})}{n_\gamma} \approx 0.42 \left(\frac{2 \times 10^{-22}\text{J}}{k \cdot 2.7255\text{K}}\right)^2 \exp\left(-\frac{2 \times 10^{-22}\text{J}}{k \cdot 2.7255\text{K}}\right) = 0.058 \quad (2.36)$$

## 2.6 CMB Cutoff - The Other Direction

We again recall the number density of the CMB photons:

$$n(f)df = \frac{8\pi}{c^3} \frac{f^2 df}{\exp(hf/kT) - 1} \quad (2.37)$$

Since  $kT \gg E_0 > hf$ , we have that:

$$\exp\left(\frac{hf}{kT}\right) \approx 1 + \frac{hf}{kT} \quad (2.38)$$

by Taylor expanding to first order. The number density in this regime therefore becomes:

$$n(f)df \approx \frac{8\pi}{c^3} \frac{f^2 df}{1 + \frac{hf}{kT} - 1} = \frac{8\pi kT}{hc^3} f df \quad (2.39)$$

Integrating this from 0 to  $E_0/h$ , we obtain  $n(hf < E_0)$ :

$$n(hf < E_0) \approx \int_0^{E_0/h} \frac{8\pi kT}{hc^3} f df = \frac{4\pi kT}{hc^3} f^2 \Big|_0^{E_0/h} = \frac{4\pi kTE_0^2}{h^3 c^3} \quad (2.40)$$

So solving for the fraction of photons in the CMB with  $hf < E_0$  we have:

$$\frac{n(hf < E_0)}{n_\gamma} \approx \frac{\frac{4\pi kTE_0^2}{h^3 c^3}}{\frac{2.4041}{\pi^2} \frac{k^3}{h^3 c^3} T^3} = 0.21 \left(\frac{E_0}{kT}\right)^2 \quad (2.41)$$

Solving for the fraction of CMB photons with  $\lambda > 3\text{cm}$  (and hence capable of passing through the Earth’s atmosphere) we again use the wavelength-frequency relation for light of  $c = \lambda f$ .  $\lambda > 3\text{cm}$  corresponds to  $f < 10^{10}\text{Hz}$  so:

$$E = hf < 6.63 \times 10^{-24}\text{J} \quad (2.42)$$

So using our obtained relation with  $T = 2.7255\text{K}$  we have:

$$\frac{n(hf < 7 \times 10^{-24}\text{J})}{n_\gamma} \approx 0.21 \left(\frac{7 \times 10^{-24}\text{J}}{k \cdot 2.7255\text{K}}\right)^2 = 0.0065 \quad (2.43)$$

### 3 Newton versus Einstein

#### 3.1 Evidence for electrical neutrality

Through astronomical observations, we notice that gravitational forces dominate dynamics in the universe on large scales. However, electrostatic forces are much stronger than gravitational. A simple demonstration of this is given by comparing the magnitudes of the gravitational and electrostatic forces between an electron and proton. At a distance  $r$ , the gravitational force between an electron and proton (in the Newtonian picture) is given by:

$$|F_g| = \frac{Gm_p m_e}{r^2} = \frac{1}{r^2} (6.67 \times 10^{-11} \text{N kg}^{-2} \text{m}^2) (1.67 \times 10^{-27} \text{kg}) (9.11 \times 10^{-31} \text{kg}) = \frac{1.01 \times 10^{-66}}{r^2} \text{N} \quad (3.1)$$

And the electrostatic force is given by:

$$|F_e| = \frac{ke^2}{r^2} = \frac{1}{r^2} (8.99 \times 10^9 \text{N C}^{-2} \text{m}^2) (1.60 \times 10^{-19} \text{C})^2 = \frac{2.30 \times 10^{-28}}{r^2} \text{N} \quad (3.2)$$

Taking the ratio we have:

$$\left| \frac{F_e}{F_g} \right| = \left| \frac{\frac{2.30 \times 10^{-28}}{r^2} \text{N}}{\frac{1.01 \times 10^{-66}}{r^2} \text{N}} \right| = 2.28 \times 10^{38} \quad (3.3)$$

As we can see from this example, electrostatic forces are much stronger than gravitational (38 orders of magnitude stronger in this case!). Hence if there was a large charge imbalance in the universe, we would observe that electrostatic forces would dominate large-scale dynamics rather than gravitational. This however disagrees with our observations, and we conclude that the universe must be (mostly) electrically neutral.

#### 3.2 Angular Width on a sphere

We recall that the metric for a sphere of radius  $R$  is given by:

$$dl^2 = dr^2 + R^2 \sin^2 \left( \frac{r}{R} \right) d\theta^2 \quad (3.4)$$

So we can write the width  $dl$  of the object using the above equation. Since  $dl \ll R$ , we can take the entire width of the object to be located at the same distance  $r$  away from us (the observer). In other words, in this limit we have  $dr = 0$  and hence the above reduces to:

$$dl^2 = R^2 \sin^2 \left( \frac{r}{R} \right) d\theta^2 \quad (3.5)$$

Isolating for the angular width of the object, we obtain:

$$d\theta = \frac{dl}{R \sin \left( \frac{r}{R} \right)} \quad (3.6)$$

As we take  $r \rightarrow \pi R$  (i.e. the object is located at the antipode) we have:

$$\lim_{r \rightarrow R\pi} d\theta = \lim_{r \rightarrow R\pi} \frac{dl}{R \sin \left( \frac{r}{R} \right)} = \infty \quad (3.7)$$

as the denominator goes to 0 ( $\lim_{r \rightarrow R\pi} \sin \left( \frac{r}{R} \right) = \sin \left( \frac{R\pi}{R} \right) = \sin(\pi) = 0$ ). The angular width is minimized when the object lies on the equator, where  $r = \frac{R\pi}{2}$  and  $d\theta = \frac{dl}{R}$ , and this angular width increases to infinity as the object approaches the antipode of the sphere. For an object at the antipode, all lines of sight from the observer lead to the object, and it therefore fills the horizon.

### 3.3 The Earth isn't flat!

We choose our coordinate system such that the origin is located in the center of the sphere, and we as the observer are located at the north pole. The circle drawn on the sphere is at a fixed distance  $r$  from us on the north pole, and therefore each point on the sphere is at a fixed polar angle  $\theta$ . The radius of this circle as measured as an outside observer (not located on the sphere) is  $R \sin \theta$ . The circumference of the circle

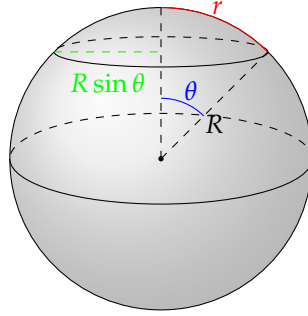


Figure 1: Illustration of setup of Problem 3.3

is therefore given by:

$$C = 2\pi R \sin \theta \quad (3.8)$$

Now, we recall that the arclength  $r$  is related to the polar angle and sphere radius by:

$$r = R\theta. \quad (3.9)$$

So substituting this into the circumference, we obtain:

$$C = 2\pi R \sin\left(\frac{r}{R}\right) \quad (3.10)$$

which was the claimed formula. For the second part of the question, we consider that for a flat space, the circumference would be measured to be:

$$C_f = 2\pi r. \quad (3.11)$$

In the limit  $r \ll R$ , we can Taylor expand the circumference as given in Eq. (3.10) to obtain that:

$$C_s \approx 2\pi R \left( \frac{r}{R} - \frac{r^3}{6R^3} \right) = 2\pi r - \frac{\pi r^3}{3R^2} \quad (3.12)$$

Hence the difference between the circumference measured in Euclidean space versus a sphere would be given by:

$$|C_f - C_s| \approx \left| 2\pi r - \left( 2\pi r - \frac{\pi r^3}{3R^2} \right) \right| = \frac{\pi r^3}{3R^2}. \quad (3.13)$$

We can measure distances within an error of  $\pm 1\text{m}$ , so to convince ourselves that Earth is spherical rather than flat, we require that the circumference difference be:

$$|C_f - C_s| > 1\text{m} \quad (3.14)$$

So approximately:

$$\frac{\pi r^3}{3R^2} > 1\text{m} \quad (3.15)$$



Rearranging, we have:

$$r > \sqrt[3]{\frac{3\pi R^2}{\pi}} \quad (3.16)$$

And substituting  $R = 6400\text{km}$  we get:

$$\boxed{r > 34\text{km}} \quad (3.17)$$

### 3.4 Area Bounds for Equilateral Triangles

**Case 1:**  $\kappa = +1$ . No, we cannot draw an equilateral triangle of arbitrarily large surface area  $A$  in this case. A simple counterargument is that a sphere has surface area  $4\pi R^2$ , so immediately no triangle with area larger than that is possible. For a calculation of the maximum area of a triangle, we require a thoughtful argument. First, recall the equation for the sum of three angles of a triangle of area  $A$  on a surface of a sphere of radius  $R$ :

$$\alpha + \beta + \gamma = \pi + \frac{A}{R^2} \quad (3.18)$$

For an equilateral triangle,  $\alpha = \beta = \gamma$ , so:

$$3\alpha = \pi + \frac{A}{R^2} \quad (3.19)$$

WLOG, we can choose our coordinate system that the first point of the triangle is on the north pole, the second point is a distance  $r$  from the north pole on the sphere with azimuthal angle  $\phi = 0$ , and the third point is also a distance  $r$  from the north pole on the sphere with azimuthal angle  $\phi = \alpha$ .

TODO

Figure 2: Illustration of the triangle drawn for the spherical case of Problem 3.4.

We now recall the area element on the sphere is given as:

$$dA = R d\theta \cdot R \sin \theta d\phi = R^2 d\theta d\phi \quad (3.20)$$

Where  $R d\theta$  is the polar distance and  $R \sin \theta d\phi$  is the azimuthal distance. With the setup as described, the area of the triangle may be calculated via integration. We integrate from  $0 \leq \theta \leq \frac{r}{R}$  (recalling that  $r = R\theta$ ) and from  $0 \leq \phi \leq \alpha$ :

$$A = \iint_{\text{triangle}} dA = R^2 \int_0^{r/R} \sin \theta d\theta \int_0^\alpha d\phi \quad (3.21)$$

Carrying out these two easy integrals, we obtain:

$$A = R^2 \left( 1 - \cos\left(\frac{r}{R}\right) \right) \alpha \quad (3.22)$$

Isolating for  $\alpha$ , we obtain:

$$\alpha = \frac{A}{R^2 \left( 1 - \cos\left(\frac{r}{R}\right) \right)} \quad (3.23)$$

Plugging this into (3.19), we get:

$$3 \frac{A}{R^2 \left( 1 - \cos\left(\frac{r}{R}\right) \right)} = \pi + \frac{A}{R^2} \quad (3.24)$$

Isolating for  $A$  we have:

$$A = \pi R^2 \frac{1 - \cos\left(\frac{r}{R}\right)}{2 + \cos\left(\frac{r}{R}\right)} \quad (3.25)$$

The largest that  $r$  can possibly be is when  $r = \pi R$  (the base of the triangle lies on the south pole), and so:

$$A_{\max} = \pi R^2 \frac{1 - \cos\left(\frac{\pi R}{R}\right)}{2 + \cos\left(\frac{\pi R}{R}\right)} = \pi R^2 \frac{1 - (-1)}{2 - 1} = 2\pi R^2 \quad (3.26)$$

In other words, the largest an equilateral triangle can be in this case is being half of the sphere!

**Case 2:**  $\kappa = 0$ . We can draw an equilateral triangle of arbitrarily large surface area  $A$  in this case. On a flat plane, there is no upper bound to how large you want to draw your shapes!

**Case 3:**  $\kappa = -1$ . No, we cannot draw an equilateral triangle of arbitrarily large surface area  $A$  in this case. Consider the equation that gives the sum of three angles of a triangle on a 2-D surface of uniform negative curvature (with radius of curvature  $R$ ):

$$\alpha + \beta + \gamma = \pi - \frac{A}{R^2} \quad (3.27)$$

For an equilateral triangle,  $\alpha = \beta = \gamma$  so:

$$3\alpha = \pi - \frac{A}{R^2} \quad (3.28)$$

$\alpha$  cannot be negative if we are to have a physical shape, so:

$$0 \leq \pi - \frac{A}{R^2} \quad (3.29)$$

Rearranging, we obtain an upper bound on the area of an equilateral triangle for a 2-D surface of uniform negative curvature:

$$A \leq R^2 \pi \quad (3.30)$$

Which gives us a maximum:

$$A_{\max} = R^2 \pi \quad (3.31)$$

in this limit, note that the angles of the triangle  $\alpha$  approach zero.

### 3.5 Equivalent Metrics

We wish to show the equivalence of the metrics:

$$dl^2 = dx^2 + dy^2 + dz^2 \quad (3.32)$$

and:

$$dl^2 = dr^2 + r^2 \left[ d\theta^2 + \sin^2 \theta d\phi^2 \right]. \quad (3.33)$$

Making the substitutions  $x = r \sin \theta \cos \phi$ ,  $y = r \sin \theta \sin \phi$ , and  $z = r \cos \theta$  into (3.32), we have:

$$dl^2 = (d(r \sin \theta \cos \phi))^2 + (d(r \sin \theta \sin \phi))^2 + (d(r \cos \theta))^2 \quad (3.34)$$

Liberal applying the product and chain rules, we obtain:

$$\begin{aligned} dl^2 = & (\sin \theta \cos \phi dr + r \cos \theta \cos \phi d\theta - r \sin \theta \sin \phi d\phi)^2 \\ & + (\sin \theta \sin \phi dr + r \cos \theta \sin \phi d\theta + r \sin \theta \cos \phi d\phi)^2 \\ & + (\cos \theta dr - r \sin \theta d\theta)^2 \end{aligned} \quad (3.35)$$

Now we can expand this expression:

$$\begin{aligned}
dl^2 = & \sin^2 \theta \cos^2 \phi dr^2 + r^2 \cos^2 \theta \cos^2 \phi d\theta^2 + r^2 \sin^2 \theta \sin^2 \phi d\phi^2 \\
& + 2r \sin \theta \cos \theta \cos^2 \phi dr d\theta - 2r \sin^2 \theta \sin \phi \cos \phi dr d\phi - 2r^2 \sin \theta \cos \theta \sin \phi \cos \phi d\theta d\phi \\
& + \sin^2 \theta \sin^2 \phi dr^2 + r^2 \cos^2 \theta \sin^2 \phi d\theta^2 + r^2 \sin^2 \theta \cos^2 \phi d\phi^2 \\
& + 2r \sin \theta \cos \theta \sin^2 \phi dr d\theta + 2r \sin^2 \theta \sin \phi \cos \phi dr d\phi + 2r^2 \sin \theta \cos \theta \sin \phi \cos \phi d\theta d\phi \\
& + \cos^2 \theta dr^2 - 2r \sin \theta \cos \theta dr d\theta + r^2 \sin^2 \theta d\theta^2 \quad (3.36)
\end{aligned}$$

Now grouping like terms, we have:

$$\begin{aligned}
dl^2 = & dr^2 \left[ \sin^2 \theta \cos^2 \phi + \sin^2 \theta \sin^2 \phi + \cos^2 \theta \right] + r^2 d\theta^2 \left[ \cos^2 \theta \cos^2 \phi + \cos^2 \theta \sin^2 \phi + \sin^2 \theta \right] \\
& + r^2 d\phi^2 \left[ \sin^2 \theta \sin^2 \phi + \sin^2 \theta \cos^2 \phi \right] + 2r dr d\theta \left[ \sin \theta \cos \theta \cos^2 \phi + \sin \theta \cos \theta \sin^2 \phi - \sin \theta \cos \theta \right] \\
& + 2r dr d\phi \left[ -\sin^2 \theta \sin \phi \cos \phi + \sin^2 \theta \sin \phi \cos \phi \right] + 2r^2 d\theta d\phi \left[ -\sin \theta \cos \theta \sin \phi \cos \phi + \sin \theta \cos \theta \sin \phi \cos \phi \right] \quad (3.37)
\end{aligned}$$

We can see the last line of terms all equals to zero (the brackets evaluate to zero, so):

$$\begin{aligned}
dl^2 = & dr^2 \left[ \sin^2 \theta \cos^2 \phi + \sin^2 \theta \sin^2 \phi + \cos^2 \theta \right] + r^2 d\theta^2 \left[ \cos^2 \theta \cos^2 \phi + \cos^2 \theta \sin^2 \phi + \sin^2 \theta \right] \\
& + r^2 d\phi^2 \left[ \sin^2 \theta \sin^2 \phi + \sin^2 \theta \cos^2 \phi \right] + 2r dr d\theta \left[ \sin \theta \cos \theta \cos^2 \phi + \sin \theta \cos \theta \sin^2 \phi - \sin \theta \cos \theta \right] \quad (3.38)
\end{aligned}$$

Now redrawing some brackets:

$$\begin{aligned}
dl^2 = & dr^2 \left[ \sin^2 \theta (\cos^2 \phi + \sin^2 \phi) + \cos^2 \theta \right] + r^2 d\theta^2 \left[ \cos^2 \theta (\cos^2 \phi + \sin^2 \phi) + \sin^2 \theta \right] \\
& + r^2 d\phi^2 \left[ \sin^2 \theta (\sin^2 \phi + \cos^2 \phi) \right] + 2r dr d\theta \left[ \sin \theta \cos \theta (\cos^2 \phi + \sin^2 \phi) - \sin \theta \cos \theta \right] \quad (3.39)
\end{aligned}$$

Now applying the  $\cos^2 \phi + \sin^2 \phi = 1$  identity to each term:

$$dl^2 = dr^2 \left[ \sin^2 \theta + \cos^2 \theta \right] + r^2 d\theta^2 \left[ \cos^2 \theta + \sin^2 \theta \right] + r^2 d\phi^2 \left[ \sin^2 \theta \right] + 2r dr d\theta \left[ \sin \theta \cos \theta - \sin \theta \cos \theta \right] \quad (3.40)$$

The last term vanishes, and to the first two terms we can apply the identity  $\cos^2 \theta + \sin^2 \theta = 1$ :

$$dl^2 = dr^2 + r^2 d\theta^2 + r^2 \sin^2 \theta d\phi^2 \quad (3.41)$$

And redrawing some brackets:

$$dl^2 = dr^2 + r^2 \left[ d\theta^2 + \sin^2 \theta d\phi^2 \right] \quad (3.42)$$

We identify this with (3.33) and hence we conclude.

## 4 Cosmic Dynamics

### 4.1 Does the Cosmological Constant Affect Planetary Motion?

In a sphere of radius 1AU, we have:

$$E_{\Lambda} = \epsilon_{\Lambda} V = \epsilon_{\Lambda} \frac{4}{3} \pi R^3 = 5200 \text{MeV m}^{-3} \frac{4}{3} \pi (1.5 \times 10^{11} \text{m})^3 = 1.18 \times 10^{25} \text{J} \quad (4.1)$$

Now calculating the rest energy of the sun, we have:

$$E_{\odot} = M_{\odot} c^2 = 1.99 \times 10^{30} \text{kg} (3.0 \times 10^8 \text{ms}^{-1})^2 = 1.79 \times 10^{47} \text{J} \quad (4.2)$$

We see a difference of 23 orders of magnitude; we conclude that the cosmological constant does not have a significant effect on the motion of planets within the Solar system.

### 4.2 Perturbing Einstein's Static Universe

If  $\Lambda = 4\pi G\rho$ , then the acceleration equation says:

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\epsilon + 3P) + \frac{\Lambda}{3} = -\frac{4\pi G\rho}{3} + \frac{4\pi G\rho}{4} = 0 \quad (4.3)$$

Where in the second equality we use that this hypothetical universe is filled solely with matter, so  $P \approx 0$ . Now, if some of this matter gets converted to radiation, we have that  $P_r = \frac{1}{3}\epsilon > 0$ , so  $P_{tot} > 0$  and hence:

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\epsilon + 3P) + \frac{\Lambda}{3} = -\frac{4\pi G P_{tot}}{3} < 0 \quad (4.4)$$

Note that while the pressure increases, the energy density  $\epsilon$  remains unchanged (the energy just converted form) so the energy density term and the cosmological constant term cancel out like before. From this, we conclude that  $\ddot{a} < 0$ , and so the universe contracts. The extra gravitational pressure from the radiation causes the universe to collapse; this shows that Einstein's static universe model isn't great, as even his universe with just one star would trigger a runaway collapse.

### 4.3 How Large is Einstein's Static Universe?

From Eq. 4.73 in Ryden, we have that in Einstein's static universe has radius of curvature:

$$R_0 = \frac{c}{2(\pi G\rho)^{1/2}} = 2 \times 10^26 \text{m} \sim 7 \text{Gpc} \quad (4.5)$$

If a photon were to circumnavigate this universe, it would take time:

$$T = \frac{2\pi R_0}{c} = 4 \times 10^{18} \text{s} \sim 132 \text{Gyr} \quad (4.6)$$

Which is longer than the age of the universe!

#### 4.4 Baseballs and Critical Density

The current critical density is given by Ryden Eq. 4.32 to be:

$$\rho_{c,0} = \frac{3}{8\pi G} H_0^2 = 8.7 \times 10^{-27} \text{kg m}^{-3} \quad (4.7)$$

We set the density of baseballs to be equal to the critical density:

$$\rho_{c,0} = \rho_{bb} = m_{bb} n_{bb} \quad (4.8)$$

Where  $n_{bb}$  is the number density of the baseballs. Rearranging, we get:

$$n_{bb} = \frac{\rho_{c,0}}{m_{bb}} = \frac{8.7 \times 10^{-27} \text{kg m}^{-3}}{0.145 \text{kg}} = 6.0 \times 10^{-26} \text{m}^{-3} \quad (4.9)$$

Given this density of baseballs, we can use Ryden Eq. 2.2 to solve for the average distance we could see before having our line of sight intersected by a baseball:

$$\lambda = \frac{1}{n_{bb} \pi r_{bb}^2} = \frac{1}{(6.0 \times 10^{-26} \text{m}^{-3}) \pi (0.0369 \text{m})^2} = 3.90 \times 10^{27} \text{m} \approx 126000 \text{Mpc} \quad (4.10)$$

The fact that we can see galaxies at a distance  $\sim c/H_0 \sim 4000 \text{Mpc}$  does not give us a useful upper bound on the density of intergalactic baseballs in this case (we see that the line of sight from the current calculation assuming critical density of baseballs is  $\sim 2$  orders of magnitude larger than what we can actually see already). However, for completeness we calculate what upper bound this does give on the density of intergalactic baseballs:

$$n_{bb} < \frac{1}{\lambda \pi r_{bb}^2} = \frac{1}{(4000 \text{Mpc}) (\pi (0.0369 \text{m})^2)} = 1.93 \times 10^{-24} \text{m}^{-3} \quad (4.11)$$

#### 4.5 Equation of State for Gases

The energy per-particle is given by:

$$E = (mc^2 + h^2 c^2 / \lambda^2)^{1/2} \quad (4.12)$$

And the total energy density of a gas of particles is given by:

$$\epsilon = nE \quad (4.13)$$

Combining the two, we have:

$$\epsilon = n(mc^2 + h^2 c^2 / \lambda^2)^{1/2} \quad (4.14)$$

Since  $n$  is the number density, we can write it as:

$$n = \frac{N}{V} = \frac{N}{k_1 a^3} \quad (4.15)$$

Where  $N$  is the number of particles in the gas (we assume this does not change, i.e. that no particles are created or destroyed), and  $V = k_1 a^3$  is the volume of the expanding universe (proportional to  $a^3$ ). Furthermore, we can write  $\lambda = k_2 a$  as the wavelength is linear in the scale factor. Putting this into (4.14) we have:

$$\epsilon = \frac{N}{k_1 a^3} (mc^2 + \frac{h^2 c^2}{k_2^2 a^2})^{1/2} \quad (4.16)$$

Now we recall the fluid equation:

$$\dot{\epsilon} + 3\frac{\dot{a}}{a}(\epsilon + P) = 0. \quad (4.17)$$

Substituting the equation of state:

$$P = w\epsilon \quad (4.18)$$

into the fluid equation, we have:

$$\dot{\epsilon} + 3\frac{\dot{a}}{a}\epsilon(1 + w) = 0 \quad (4.19)$$

Solving for  $w$ , we have:

$$w = \frac{-\dot{\epsilon}}{3\frac{\dot{a}}{a}\epsilon} - 1 \quad (4.20)$$

We will have to take the time derivative of (4.16) to substitute into (4.20). Noting that the only time-dependent parameter in  $\epsilon$  is  $a$ , we take the derivative (using the quotient rule and chain rule):

$$\dot{\epsilon} = \frac{\frac{N}{(mc^2 + \frac{h^2c^2}{k_2^2a^2})^{1/2}} \left( -2\frac{h^2c^2}{k_2^2a^3} \right) \dot{a}k_1a^3 - 3Nk_1a^2\dot{a}(mc^2 + \frac{h^2c^2}{k_2^2a^2})^{1/2}}{k_1^2a^6} \quad (4.21)$$

Simplifying slightly:

$$\dot{\epsilon} = \frac{-Nc^2\dot{a}(3k_2^2ma^2 + 4h^2)}{k_1k_2^2a^6(mc^2 + \frac{h^2c^2}{k_2^2a^2})^{1/2}} \quad (4.22)$$

Substituting (4.16) and (4.22) into (4.20) we have:

$$w = \frac{\frac{Nc^2\dot{a}(3k_2^2ma^2 + 4h^2)}{k_1k_2^2a^6(mc^2 + \frac{h^2c^2}{k_2^2a^2})^{1/2}}}{3\frac{\dot{a}}{a}\frac{N}{k_1a^3}(mc^2 + \frac{h^2c^2}{k_2^2a^2})^{1/2}} - 1 \quad (4.23)$$

Cancelling terms in the numerator and denominator, we have:

$$w = \frac{c^2(3k_2^2ma^2 + 4h^2)}{3k_2^2a^2(mc^2 + \frac{h^2c^2}{k_2^2a^2})} - 1 \quad (4.24)$$

Expanding out terms in the numerator and denominator:

$$w = \frac{3mc^2k_2^2a^2 + 4h^2c^2}{3mc^2k_2^2a^2 + 3h^2c^2} - 1 \quad (4.25)$$

In the highly relativistic limit, we have  $a \rightarrow 0$  and  $p \rightarrow 0$ . Note that while  $p$  does not appear explicitly in the above equation,  $a \rightarrow 0$  implies  $p \rightarrow \infty$  under the linear relationship of the scaling factor with  $\lambda$ , as  $k_2a = \lambda = h/p$ . In any case, taking  $a \rightarrow 0$  in the above expression, we have:

$$w_{\text{rel}} = \lim_{a \rightarrow 0} w = \lim_{a \rightarrow 0} \frac{3mc^2k_2^2a^2 + 4h^2c^2}{3mc^2k_2^2a^2 + 3h^2c^2} - 1 = \frac{4h^2c^2}{3h^2c^2} - 1 = \frac{4}{3} - 1 = \frac{1}{3} \quad (4.26)$$

This was precisely the claimed value. Now in the highly non-relativistic limit, we have  $a \rightarrow \infty$  and  $p \rightarrow 0$ . We again take the limit of  $a \rightarrow \infty$  in (4.25) to obtain:

$$w_{\text{nonrel}} = \lim_{a \rightarrow \infty} w = \lim_{a \rightarrow \infty} \frac{3mc^2k_2^2a^2 + 4h^2c^2}{3mc^2k_2^2a^2 + 3h^2c^2} - 1 = \frac{3mc^2k_2^2}{3mc^2k_2^2} - 1 = 1 - 1 = 0 \quad (4.27)$$

which is again the desired value.

## A Force-Based Derivation of the Newtonian Friedmann Equation

Now let's derive the "acceleration equation" (sometimes called "Friedmann's other equation"!) for the whole Universe from simple Newtonian physics. Imagine a sphere of constant density  $\rho(t)$  and radius  $r$ , with a test mass  $m$  at its edge. Write down the equation of motion for the test mass under the gravitational pull of the sphere. Now use the idea that the physical radius can be written as comoving radius times scale factor, i.e.  $r \equiv a(t)x$ . you should find that you can derive an equation for  $a$  which doesn't depend on  $x$  or on  $m$ ! In other words, the sphere that you used in the first place has disappeared and your equation of motion has ended up being for the scale factor itself. [Note that you're *not* being asked to solve this equation, just to derive it!]

The mass of the sphere is given by:

$$M(t) = \rho(t)V(t) = \rho(t)\frac{4}{3}\pi r(t)^3 \quad (4.28)$$

The distance from the center of the sphere to the test mass is just  $r(t)$  (the test mass is on the surface), so using Newton's second law and Newton's law of universal gravitation, we have:

$$m\ddot{r}(t) = F = -\frac{GM(t)m}{r(t)^2} \quad (4.29)$$

Substituting  $M(t)$  from (4.28), we have:

$$m\ddot{r}(t) = \frac{-G\rho(t)\frac{4}{3}\pi r(t)^3 m}{r(t)^2} = -G\rho(t)\frac{4}{3}\pi a(t)m \quad (4.30)$$

Cancelling out  $m$  from both sides and replacing  $r(t)$  with  $a(t)x$  we have:

$$x\ddot{a}(t) = -G\rho(t)\frac{4}{3}\pi a(t)x \quad (4.31)$$

The  $x$ s cancel on both sides, and dividing both sides by  $a(t)$  we get:

$$\boxed{\frac{\ddot{a}(t)}{a(t)} = -G\rho(t)\frac{4}{3}\pi} \quad (4.32)$$

## 5 Model Universes

### 5.1 Redshift in single-component universes

We can take Eq. 5.47 in Ryden as our starting point:

$$1 + z = \frac{a(t_0)}{a(t_e)} = \left( \frac{t_0}{t_e} \right)^{2/(3+3w)} \quad (5.1)$$

Taking the derivative w.r.t.  $t_0$  of both sides of this equation, we obtain:

$$\frac{dz}{dt_0} = \frac{\dot{a}(t_0) \frac{dt_0}{dt_0} a(t_e) - \dot{a}(t_e) \frac{dt_e}{dt_0} a(t_0)}{a(t_e)^2} \quad (5.2)$$

There's a variety of terms to process here. First, combining Ryden Eqs. 5.39 and 5.42 we get:

$$a(t) = \left( \frac{t}{t_0} \right)^{2/(3+3w)} \quad (5.3)$$

Where  $t_0$  is the age of the universe. Taking the time derivative, we have:

$$\dot{a}(t) = \frac{2}{3+3w} \left( \frac{t}{t_0} \right)^{2/(3+3w)} \frac{1}{t} = \frac{2}{3+3w} \frac{a(t)}{t} \quad (5.4)$$

We also have that:

$$\frac{dt_0}{dt_0} = 1 \quad (5.5)$$

The final quantity to determine is  $\frac{dt_e}{dt_0}$ . To solve for this, we recall Ryden 3.59:

$$\frac{1}{a(t_e)} \int_{t_e}^{t_e + \lambda_e/c} dt = \frac{1}{a(t_0)} \int_{t_0}^{t_0 + \lambda_0/c} dt \quad (5.6)$$

Since  $\lambda/c \ll 1$ , we can approximate  $\lambda_e/c \sim dt_e$ ,  $\lambda_0/c \sim dt_0$  to get:

$$\frac{1}{a(t_e)} \int_{t_e}^{t_e + dt_e} dt = \frac{1}{a(t_0)} \int_{t_0}^{t_0 + dt_0} dt \quad (5.7)$$

Which solving the integrals tells us that:

$$\frac{dt_e}{a(t_e)} = \frac{dt_0}{a(t_0)} \quad (5.8)$$

Which we can rearrange to obtain:

$$\frac{dt_e}{dt_0} = \frac{a(t_e)}{a(t_0)} = \frac{1}{1+z} \quad (5.9)$$

Combining the three obtained relations of (5.4), (5.5), (5.9) and substituting this into (5.2) we have that:

$$\frac{dz}{dt_0} = \frac{\frac{2}{3+3w} \frac{a(t_0)}{t_0} a(t_e) - \frac{2}{3+3w} \frac{a(t_e)}{t_e} \frac{1}{1+z} a(t_0)}{a(t_e)^2} \quad (5.10)$$

Factoring and using that  $\frac{a(t_0)}{a(t_e)} = 1 + z$  from (5.1) this becomes:

$$\frac{dz}{dt_0} = \frac{2}{3+3w} \left( (1+z) \frac{1}{t_0} - \frac{1}{t_e} \right) \quad (5.11)$$



Substituting  $t_e = t_0(1+z)^{-(3+3w)/2}$  from (5.1) we have:

$$\frac{dz}{dt_0} = \frac{2}{3+3w} \left( (1+z) \frac{1}{t_0} - (1+z)^{(3+3w)/2} \frac{1}{t_0} \right) \quad (5.12)$$

Now using Ryden Eq. 5.42:

$$t_0 = \frac{2}{3+3w} H_0^{-1} \quad (5.13)$$

We obtain:

$$\boxed{\frac{dz}{dt_0} = H_0(1+z) - H_0(1+z)^{(3+3w)/2}} \quad (5.14)$$

which was the desired relation. The observed redshift increases in time if  $\frac{dz}{dt_0} > 0$ , so rearranging the above expression we have:

$$H_0(1+z) - H_0(1+z)^{(3+3w)/2} > 0 \implies (1+z) > (1+z)^{(3+3w)/2} \quad (5.15)$$

And since  $1+z \geq 1$ , the LHS is greater than the RHS if  $(3+3w)/2 < 1$ , i.e. if:

$$\boxed{w < -\frac{1}{3}}. \quad (5.16)$$

## 5.2 Redshift Change Timescale for Flat Matter-Only Universe

In a flat matter-only universe, we can use the result from 5.1 with  $w = 0$  to obtain that:

$$\frac{dz}{dt_0} = H_0(1+z) - H_0(1+z)^{3/2} \quad (5.17)$$

We want to find the time  $dt_0$  for the galaxy to change by  $\frac{dz}{z} = -10^{-6}$ , and since  $z = 1$  to start  $dz = -10^{-6}$ . Rearranging the above expression to solve for  $dt_0$  we have:

$$dt_0 = \frac{dz}{H_0} \frac{1}{(1+z) - (1+z)^{3/2}} \quad (5.18)$$

So plugging in  $dz = -10^{-6}$ ,  $z = 1$ , and  $H_0 = 68 \text{ km s}^{-1} \text{ Mpc}^{-1}$  we numerically get:

$$\boxed{dt_0 \approx 5.49 \times 10^{11} \text{ s} \approx 17400 \text{ yrs}} \quad (5.19)$$

## 5.3 Present age of universe for positively curved matter-only universe

We recall the parametric solutions for the scale factor  $a$  and the time  $t$  for a positively curved universe filled only with matter:

$$a(\theta) = \frac{1}{2} \frac{\Omega_0}{\Omega_0 - 1} (1 - \cos \theta) \quad (5.20)$$

$$t(\theta) = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\theta - \sin \theta). \quad (5.21)$$

In the present day,  $a = 1$ , so solving (5.20) for  $\theta$  we have:

$$\theta = \cos^{-1} \left( 1 - \frac{2(\Omega_0 - 1)}{\Omega_0} \right) = \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \quad (5.22)$$

Substituting this into (5.21) we have:

$$t_0 = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \left( \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) - \sin \left( \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right) \right) \quad (5.23)$$

The last term looks complicated, but we can view  $\cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right)$  as the angle for a triangle with hypotenuse  $\Omega_0$  and adjacent side  $2 - \Omega_0$ . The sine of this will therefore be the ratio of the opposite side and the hypotenuse of this triangle. The length of the opposite side is given (by Pythagoras) as:

$$\sqrt{\Omega_0^2 - (2 - \Omega_0)^2} = 2\sqrt{\Omega_0 - 1} \quad (5.24)$$

So  $\sin \left( \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right)$  is given as:

$$\sin \left( \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right) = \frac{2\sqrt{\Omega_0 - 1}}{\Omega_0}. \quad (5.25)$$

So substituting this into (5.23) we get:

$$H_0 t_0 = \frac{1}{2} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \left( \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) - \frac{2\sqrt{\Omega_0 - 1}}{\Omega_0} \right) \quad (5.26)$$

Which distributing the product we get:

$$H_0 t_0 = \frac{1}{2} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \cos^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) - \frac{1}{\Omega_0 - 1} \quad (5.27)$$

which was the desired expression. A plot of  $t_0$  vs.  $\Omega_0$  for  $1 \leq \Omega_0 \leq 3$  is given below.

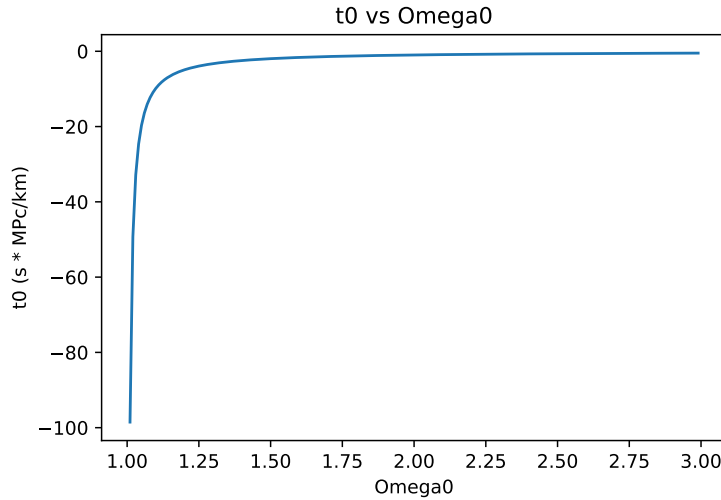


Figure 3: Plot of  $t_0$  vs.  $\Omega_0$  for  $1 \leq \Omega_0 \leq 3$ .

## 5.4 Present age of universe for negatively curved matter-only universe

We recall the parametric solutions for the scale factor  $a$  and the time  $t$  for a negatively curved universe filled only with matter:

$$a(\eta) = \frac{1}{2} \frac{\Omega_0}{1 - \Omega_0} (\cosh \eta - 1) \quad (5.28)$$

$$t(\eta) = \frac{1}{2H_0} \frac{\Omega_0}{(1 - \Omega_0)^{3/2}} (\sinh \eta - \eta). \quad (5.29)$$

In the present day,  $a = 1$ , so solving (5.28) for  $\eta$  we have:

$$\eta = \cosh^{-1} \left( \frac{2(1 - \Omega_0)}{\Omega_0} + 1 \right) = \cosh^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \quad (5.30)$$

Substituting this into (5.29) we get:

$$t_0 = \frac{1}{2H_0} \frac{\Omega_0}{(1 - \Omega_0)^{3/2}} \left( \sinh \left( \cosh^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right) - \cosh^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right) \quad (5.31)$$

Now, using that  $\sinh(\cosh^{-1} x) = \sqrt{x^2 - 1}$ , the first term of the above expression becomes:

$$\sinh \left( \cosh^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right) = \left( \frac{(2 - \Omega_0)^2}{\Omega_0^2} - 1 \right)^{1/2} = \left( \frac{4 - 4\Omega_0}{\Omega_0^2} \right)^{1/2} = \frac{2}{\Omega_0} (1 - \Omega_0)^{1/2} \quad (5.32)$$

So substituting this result we obtain:

$$H_0 t_0 = \frac{1}{2} \frac{\Omega_0}{(1 - \Omega_0)^{3/2}} \left( \frac{2}{\Omega_0} (1 - \Omega_0)^{1/2} - \cosh^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right) \right) \quad (5.33)$$

Distributing the terms, we obtain the desired result:

$$\boxed{H_0 t_0 = \frac{1}{1 - \Omega_0} - \frac{\Omega_0}{2(1 - \Omega_0)^{3/2}} \cosh^{-1} \left( \frac{2 - \Omega_0}{\Omega_0} \right)}. \quad (5.34)$$

Finally, a plot of  $t_0$  vs.  $\Omega_0$  for  $0 \leq \Omega_0 \leq 1$  is given below.

## 5.5 Phantom Energy and the Big Rip

The supposed “phantom energy” with equation of state parameter  $w_p < -1$  would have energy density:

$$\epsilon_p(a) = \epsilon_{p,0} a^{-3(1+w_p)}. \quad (5.35)$$

Comparitively, matter has energy density:

$$\epsilon_m(a) = \epsilon_{m,0} a^{-3}. \quad (5.36)$$

At equality, we have that:

$$1 = \frac{\epsilon_p(a)}{\epsilon_m(a)} = \frac{\epsilon_{p,0} a^{-3(1+w_p)}}{\epsilon_{m,0} a^{-3}} = \frac{\Omega_{p,0}}{\Omega_{m,0}} \frac{1}{a^{3w_p}} \quad (5.37)$$

Since  $\Omega_{p,0} = 1 - \Omega_{m,0}$ , we have:

$$1 = \frac{1 - \Omega_{m,0}}{\Omega_{m,0}} \frac{1}{a^{3w_p}} \quad (5.38)$$

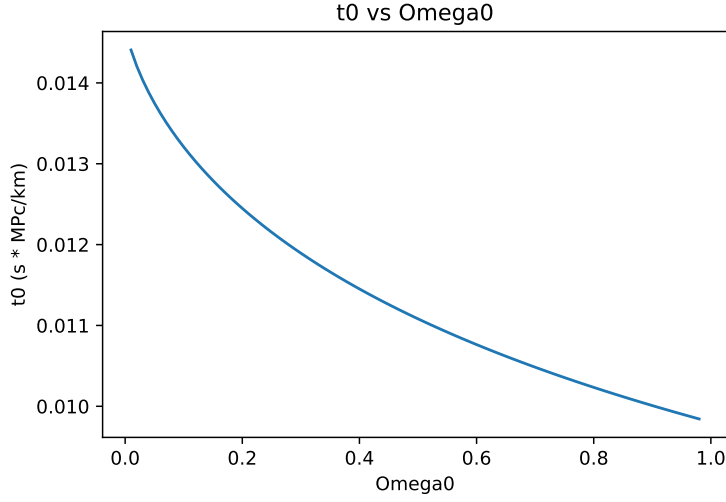


Figure 4: Plot of  $t_0$  vs.  $\Omega_0$  for  $0 \leq \Omega_0 \leq 1$ .

Which we can rearrange to solve for the scale factor  $a_{mp}$  where equality holds:

$$a_{mp} = \left( \frac{1}{\Omega_{m,0}} - 1 \right)^{1/3w_p} \quad (5.39)$$

The Friedmann equation in this universe reads:

$$\frac{H^2}{H_0^2} = \frac{\Omega_{m,0}}{a^3} + \frac{\Omega_{p,0}}{a^{3(1+w_p)}} = \frac{\Omega_{m,0}}{a^3} + \frac{1 - \Omega_{m,0}}{a^{3(1+w_p)}} \quad (5.40)$$

In the limit  $a \gg a_{mp}$ , the first term becomes negligible and hence:

$$\frac{H^2}{H_0^2} \approx \frac{1 - \Omega_{m,0}}{a^{3(1+w_p)}} \quad (5.41)$$

We now rearrange this equation to set up for the integration:

$$\frac{\dot{a}}{H_0} \approx \frac{(1 - \Omega_{m,0})^{1/2}}{a^{3(1+w_p)/2}} \implies H_0 dt \approx a^{3(1+w_p)/2} (1 - \Omega_{m,0})^{-1/2} da \quad (5.42)$$

Integrating from  $t_0$  to  $t_{rip}$ , corresponding to from  $a(t_0) = 1$  to  $a(t_{rip}) = \infty$  on the RHS, we have:

$$\int_{t_0}^{t_{rip}} H_0 dt \approx \int_1^\infty a^{3(1+w_p)/2} (1 - \Omega_{m,0})^{-1/2} da \quad (5.43)$$

Carrying out the integral, we get:

$$H_0(t_{rip} - t_0) \approx \frac{2}{3(1+w_p)} a^{3(1+w_p)/2} (1 - \Omega_{m,0})^{-1/2} \Big|_1^\infty \quad (5.44)$$

Since  $w_p < -1$ , the RHS goes to 0 at  $a = \infty$ , so:

$$H_0(t_{rip} - t_0) \approx -\frac{2}{3(1+w_p)} (1 - \Omega_{m,0})^{-1/2} \quad (5.45)$$

Which we can write as:

$$H_0(t_{\text{rip}} - t_0) \approx \frac{2}{3|1 + w_p|} (1 - \Omega_{m,0})^{-1/2} \quad (5.46)$$

which was the desired result. Finally, with  $H_0 = 68 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_{m,0} = 0.3$ , and  $w_p = -1.1$ , we can solve numerically for the time remaining until the “Big Rip” to be:

$$t_{\text{rip}} - t_0 \approx 2.53 \times 10^{18} \text{ s} = 80.4 \times 10^{10} \text{ yrs} \quad (5.47)$$

## 5.6 Pulling an Einstein

In this universe with matter and dark energy, the acceleration equation reads:

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3c^2}(\epsilon + 3P) = -\frac{4\pi G}{3c^2}(\epsilon_m + \epsilon_q + 3w_q\epsilon_q) \quad (5.48)$$

Where we have used that  $P = (0)\epsilon_m = 0$  for matter and  $P = w_q\epsilon_q$  for the dark energy. To have a static universe with  $\ddot{a} = 0$ , it must follow that:

$$\epsilon_m = -(1 + 3w_q)\epsilon_q \quad (5.49)$$

The Friedmann equation reads:

$$H(t)^2 = \frac{8\pi G}{3c^2}(\epsilon_m + \epsilon_q) - \frac{\kappa c^2}{R_0^2 a(t)^2} \quad (5.50)$$

Since we have a static universe,  $\dot{a} = 0$  and hence  $H(t) = 0$ . Furthermore, using that  $\epsilon_m = -(1 + 3w_q)\epsilon_q$  from earlier, the Friedmann equation becomes:

$$0 = \frac{8\pi G}{3c^2}(-3w_q)\epsilon_q - \frac{\kappa c^2}{R_0^2 a(t)^2} \quad (5.51)$$

Since  $-1 < w_q < -1/3$ ,  $w_q$  is negative, and hence the only way the above equality is satisfied is if  $\kappa = 1$  so the positive term can be cancelled by the second term. So in this scenario, the universe is positively curved.

We rearrange the above equation to solve for the radius of curvature  $R_0$  when  $a(t) = 1$ :

$$R_0 = \sqrt{\frac{c^4}{8\pi G|w_q|\epsilon_q}} \quad (5.52)$$

## 5.7 Big Crunch

For a positively curved matter only universe, we recall the parametric solutions for  $a(\theta)$ ,  $t(\theta)$  (Ryden Eqs. 5.90/5.91):

$$a(\theta) = \frac{1}{2} \frac{\Omega_0}{\Omega_0 - 1} (1 - \cos \theta) \quad (5.53)$$

$$t(\theta) = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\theta - \sin \theta) \quad (5.54)$$

The big bang occurs at  $\theta = 0$ , and the Big crunch at  $\theta = 2\pi$ . Given the parametric solutions above, the time between these can be computed as (Ryden Eq. 5.92):

$$t_{\text{crunch}} = \frac{\pi}{H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \quad (5.55)$$

Furthermore, we can solve for the present  $t_0$  at which  $a(\theta_0) = 1$ :

$$a(\theta_0) = 1 = \frac{1}{2} \frac{\Omega_0}{\Omega_0 - 1} (1 - \cos \theta_0) \implies \theta_0 = \arccos \left( 1 - \frac{2(\Omega_0 - 1)}{\Omega_0} \right) = \arccos \left( \frac{2}{\Omega_0} - 1 \right) \quad (5.56)$$

Therefore the time between Dr. Niwde's observations at  $t = t_0 = t(\theta_0)$  and the final big crunch is given by:

$$\Delta t = t_{\text{crunch}} - t(\theta_0) = \frac{\pi}{H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} - \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\theta_0 - \sin \theta_0) \quad (5.57)$$

Or more concisely:

$$\boxed{\Delta t = \frac{1}{H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \left( \pi - \frac{1}{2} (\theta_0 - \sin \theta_0) \right)} \quad (5.58)$$

Where  $\theta_0$  is given in (5.56). To determine the highest amplitude blueshift, we recall the redshift-scale factor relation  $1 + z = \frac{1}{a}$ , so:

$$z = \frac{1}{a} - 1 \quad (5.59)$$

If we want to minimize  $z$  (i.e. have it be the most negative/highest magnitude blueshift), we want to maximize  $a$ . We can determine this from the Friedmann equation (a la Ryden Eq. 5.86) but it also can easily be read off from the parametric solution above to be:

$$a_{\text{max}} = \frac{\Omega_0}{\Omega_0 - 1} \quad (5.60)$$

which is attained at  $\theta = \pi$ . So, the highest amplitude blueshift that Dr. Niwde can observe is at:

$$\boxed{z_{\text{blue, max}} = \frac{\Omega_0 - 1}{\Omega_0} - 1 = -\frac{1}{\Omega_0}} \quad (5.61)$$

At this blueshift, as stated previously we have  $\theta = \pi$ . So this occurs at:

$$t_{\text{blue, max}} = t(\pi) = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\pi - \sin \pi) = \frac{\pi}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \quad (5.62)$$

So the lookback time is:

$$t_0 - t_{\text{blue, max}} = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\theta_0 - \sin \theta_0) - \frac{\pi}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} \quad (5.63)$$

Or with some factoring:

$$\boxed{t_0 - t_{\text{blue, max}} = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\theta_0 - \sin \theta_0 - \pi)} \quad (5.64)$$

## 5.8 Big Bounce

The Friedmann equation in such a universe would read:

$$\frac{H^2}{H_0^2} = \Omega_0 + \frac{1 - \Omega_0}{a^2} \quad (5.65)$$

At the point where  $a$  has an extrema,  $H(t) = 0$  and so:

$$0 = \Omega_0 + \frac{1 - \Omega_0}{a^2} \implies a_{\text{bounce}} = \left( \frac{\Omega_0 - 1}{\Omega_0} \right)^{1/2} \quad (5.66)$$

Rearranging the Friedmann equation, we have:

$$\frac{\dot{a}}{a} = H_0 \sqrt{\Omega_0} \sqrt{1 - \frac{a_{\text{bounce}}^2}{a^2}} \implies \frac{da}{dt} = \frac{H_0}{\sqrt{\Omega_0}} \sqrt{a^2 - a_{\text{bounce}}^2} \quad (5.67)$$

Integrating, we obtain:

$$\sqrt{\Omega_0} H_0 \int_{t_{\text{bounce}}}^{t_0} dt = \int_{a_{\text{bounce}}}^a \frac{1}{\sqrt{a'^2 - a_{\text{bounce}}^2}} da' \quad (5.68)$$

Integrating both sides (making use of an integral table for the RHS), we have:

$$\sqrt{\Omega_0} H_0 (t - t_{\text{bounce}}) = \text{arccosh} \left( \frac{a}{a_{\text{bounce}}} \right) \quad (5.69)$$

Which we can rearrange to obtain:

$$a(t) = a_{\text{bounce}} \cosh \left( \sqrt{\Omega_0} H_0 (t - t_{\text{bounce}}) \right) \quad (5.70)$$

At  $t_0$  we have that  $a(t_0) = 1$  by convention, so solving for  $t_0 - t_{\text{bounce}}$  using (5.69) we have:

$$t_0 - t_{\text{bounce}} = \frac{1}{H_0 \sqrt{\Omega_0}} \text{arccosh} \left( \left( \frac{\Omega_0}{\Omega_0 - 1} \right)^{1/2} \right) \quad (5.71)$$

## 5.9 $\Omega_{m,0}$ for $t = H_0^{-1}$

In a spatially flat and matter + cosmological constant filled universe, the Friedmann equation can be integrated to yield the analytic solution relating  $t$  and  $a$  (Ryden Eq. 5.101):

$$H_0 t = \frac{2}{3\sqrt{1 - \Omega_{m,0}}} \ln \left[ \left( \frac{a}{a_{m\Lambda}} \right)^{3/2} + \sqrt{1 + \left( \frac{a}{a_{m\Lambda}} \right)^3} \right] \quad (5.72)$$

Where  $a_{m\Lambda}$  is defined by:

$$a_{m\Lambda} = \left( \frac{\Omega_{m,0}}{1 - \Omega_{m,0}} \right)^{1/3} \quad (5.73)$$

At the present time, we have  $t = t_0$ , and  $a = a(t_0) = 1$  by convention. Substituting this, as well as (5.73) into (5.72) we get:

$$H_0 t_0 = \frac{2}{3\sqrt{1 - \Omega_{m,0}}} \ln \left[ \sqrt{\frac{1 - \Omega_{m,0}}{\Omega_{m,0}}} + \sqrt{1 + \frac{1 - \Omega_{m,0}}{\Omega_{m,0}}} \right] \quad (5.74)$$

In order to have  $t_0 = H_0^{-1}$  we require the RHS of the above equation to exactly equal one:

$$1 = \frac{2}{3\sqrt{1 - \Omega_{m,0}}} \ln \left[ \sqrt{\frac{1 - \Omega_{m,0}}{\Omega_{m,0}}} + \sqrt{1 + \frac{1 - \Omega_{m,0}}{\Omega_{m,0}}} \right] \quad (5.75)$$

Rewriting the last term:

$$1 = \frac{2}{3\sqrt{1-\Omega_{m,0}}} \ln \left[ \sqrt{\frac{1-\Omega_{m,0}}{\Omega_{m,0}}} + \sqrt{\frac{1}{\Omega_{m,0}}} \right] = \frac{2}{3\sqrt{1-\Omega_{m,0}}} \ln \left[ \frac{1 + \sqrt{1-\Omega_{m,0}}}{\sqrt{\Omega_{m,0}}} \right] \quad (5.76)$$

Using the hyperbolic secant identity of  $\operatorname{arcsech} x = \ln \left( \frac{1+\sqrt{1-x^2}}{x} \right)$ , we have:

$$1 = \frac{2}{3\sqrt{1-\Omega_{m,0}}} \operatorname{arcsech}(\sqrt{\Omega_{m,0}}) \quad (5.77)$$

Or:

$$\frac{3\sqrt{1-\Omega_{m,0}}}{2} = \operatorname{arcsech}(\sqrt{\Omega_{m,0}}) \quad (5.78)$$

This equation can now be solved graphically for  $\Omega_{m,0}$  (I don't think an analytical closed form exists):

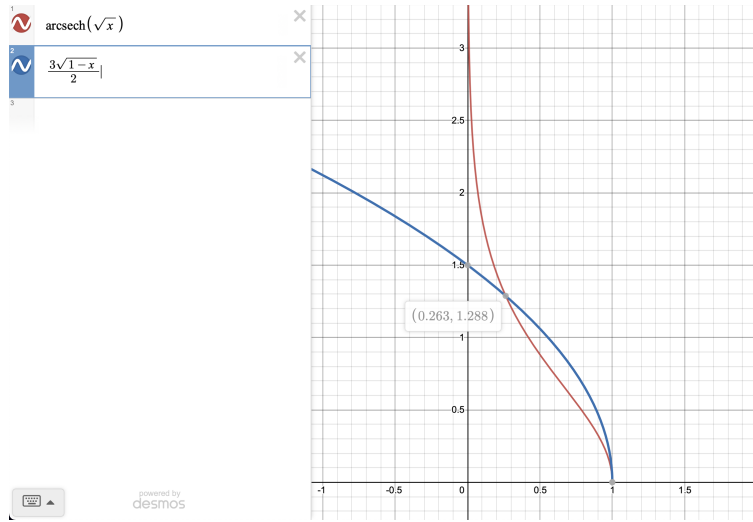


Figure 5: Numerical solution for (5.78).

From which we conclude:

$$\Omega_{m,0} \approx 0.263 \quad (5.79)$$

## 5.10 Amounts in the Benchmark model

In the Benchmark model, we have (from Ryden Table 5.2) that:

$$\Omega_{m,0} = 0.31, \quad \Omega_{\gamma,0} = 5.35 \times 10^{-5}, \quad \Omega_{\text{bary},0} = 0.048 \quad (5.80)$$

We that  $\Omega = \frac{\epsilon(t)}{\epsilon_c(t)}$  where  $\epsilon_c$  is the critical density, which is currently  $\epsilon_{c,0} 7.8 \times 10^{-10} \text{J m}^{-3}$  (up to uncertainty). Therefore obtaining the energy density for photons in the benchmark model, we have:

$$\epsilon_{\gamma,0} = \epsilon_{c,0} \Omega_{\gamma,0} = 4.17 \times 10^{-14} \text{J m}^{-3} \quad (5.81)$$

It will be more convenient to solve for the mass density for the matter and baryons:

$$\rho_{m,0} = \epsilon_{c,0} \Omega_{\gamma,0} / c^2 = 2.69 \times 10^{-27} \text{kg m}^{-3}, \quad \rho_{\text{bary},0} = \epsilon_{c,0} \Omega_{\text{bary},0} / c^2 = 4.16 \times 10^{-28} \text{J m}^{-3} \quad (5.82)$$



In addition, in the Benchmark model we have a finite horizon distance (Ryden Eq. 5.115):

$$d_{\text{hor}}(t_0) = 14000 \text{Mpc} = 4.33 \times 10^{26} \text{m}. \quad (5.83)$$

Therefore the volume of the universe within the horizon distance is:

$$V_{\text{hor}}(t_0) = \frac{4}{3} \pi d_{\text{hor}}(t_0)^3 = 3.39 \times 10^{80} \text{m}^3 \quad (5.84)$$

Now, solving for the total mass of all matter within our horizon, we have:

$$M_m = V_{\text{hor}}(t_0) \rho_{m,0} = 9.12 \times 10^{53} \text{kg} \quad (5.85)$$

Next, solving for the total amount of energy of photons within the horizon distance, we have:

$$E_{\text{photons}} = V_{\text{hor}}(t_0) \epsilon_{\gamma,0} = 1.41 \times 10^{67} \text{J} \quad (5.86)$$

Finally, we solve for the number of baryons in the universe. This will be the total mass of baryons within the universe divided by the mean mass per Baryon. Approximating baryonic matter as 75% hydrogen and 25% helium, we have the mean mass is given by:

$$m_{\text{avg}} = 0.75m_{\text{H}} + 0.25m_{\text{He}} \approx 0.75(m_p) + 0.25(4m_p) = 1.75m_p = 2.92 \times 10^{-27} \text{kg} \quad (5.87)$$

So therefore the total number of baryons within the horizon, we have:

$$N_{\text{bary}} = V_{\text{hor}}(t_0) \rho_{\text{bary},0} / m_{\text{avg}} = 4.83 \times 10^{79} \text{baryons} \quad (5.88)$$

## Time and Scale Factor for Matter-Only Universe

Let's make sure we can work through the mathematical steps for a closed matter-only (i.e. 'Matter + Curvature') universe. Try to do this without looking up the book for every step! Start by writing the Friedmann equation for this case, using  $a$  rather than  $1+z$ , and with the single parameter  $\Omega_0$  (where this implicitly means 'matter', here). Thus write an integral expression for  $t$ . It may not look trivial to solve this for  $t(a)$  or  $a(t)$ , but you should be able to *show* that the following parametric solution works [a 'parametric solution' means that you can write doesn't  $y(\phi)$  and  $x(\phi)$ , both in terms of a parameter  $\phi$ , even if an explicit expression for  $y(x)$  is hard or impossible]:

$$a(\theta) = \frac{1}{2} \frac{\Omega_0}{\Omega_0 - 1} (1 - \cos \theta) \quad (5.89)$$

$$t(\theta) = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\theta - \sin \theta). \quad (5.90)$$

To be clear: you are being asked to show that in the appropriate Friedmann equation, the LHS is equal to the RHS if you assume the above solution (*or* you could solve the integral, but that's harder!). [You might have to do a bit of rearranging of trig functions - but persevere, because it works!] Lastly, by letting  $\theta$  run from 0 to  $2\pi$ , sketch  $a$  vs.  $t$  for this model.

The Friedmann equation in this case reads:

$$\frac{H^2}{H_0^2} = \frac{\Omega_0}{a^3} + \frac{1 - \Omega_0}{a^2} \quad (5.91)$$

Now, since  $H = \frac{\dot{a}}{a}$ , we can rearrange this to obtain that:

$$\frac{\dot{a}^2}{H_0^2} = \left( \frac{\Omega_0}{a} + 1 - \Omega_0 \right) \quad (5.92)$$

Or in other words:

$$\frac{da}{dt} = H_0 \sqrt{\frac{\Omega_0}{a} + 1 - \Omega_0} \quad (5.93)$$

So solving for  $t$  we integrate:

$$H_0 t = \int_0^t dt' = \int_0^a \frac{da'}{\sqrt{\frac{\Omega_0}{a'} + 1 - \Omega_0}}. \quad (5.94)$$

Now, we verify that (5.92) holds for the parametric solution given by (5.89) and (5.90). First, we determine what  $\dot{a}$  is given this solution. By the chain rule, we have that:

$$\dot{a} = \frac{da}{dt} = \frac{da}{d\theta} \frac{d\theta}{dt} = \frac{da}{d\theta} \dot{\theta} \quad (5.95)$$

Solving for  $\frac{da}{d\theta}$  by differentiating (5.89) we have:

$$\frac{da}{d\theta} = \frac{1}{2} \frac{\Omega_0}{\Omega_0 - 1} \sin \theta. \quad (5.96)$$

And solving for  $\dot{\theta}$  by implicitly differentiating (5.90) we have:

$$1 = \frac{1}{2H_0} \frac{\Omega_0}{(\Omega_0 - 1)^{3/2}} (\dot{\theta} - \dot{\theta} \cos \theta) \implies \dot{\theta} = \frac{2H_0(\Omega_0 - 1)^{3/2}}{\Omega_0(1 - \cos \theta)} \quad (5.97)$$

Hence, we find that:

$$\dot{a} = \frac{H_0(\Omega_0 - 1)^{1/2}}{1 - \cos \theta} \sin \theta \quad (5.98)$$

Evaluating the LHS of (5.92), we have:

$$\frac{\dot{a}^2}{H_0^2} = \frac{\Omega_0 - 1}{(1 - \cos \theta)^2} \sin^2 \theta \quad (5.99)$$

Evaluating the RHS of (5.92), we have:

$$\frac{\Omega_0}{a} + 1 - \Omega_0 = \left( \frac{\Omega_0}{\frac{1}{2} \frac{\Omega_0}{\Omega_0 - 1} (1 - \cos \theta)} + 1 - \Omega_0 \right) \quad (5.100)$$

Multiplying the first term by  $1 = \frac{1 - \cos \theta}{1 - \cos \theta}$  and the second term by  $1 = \left( \frac{1 - \cos \theta}{1 - \cos \theta} \right)^2$  we obtain:

$$\frac{\Omega_0}{a} + 1 - \Omega_0 = \frac{2(\Omega_0 - 1)}{(1 - \cos \theta)^2} (1 - \cos \theta) - \frac{\Omega_0 - 1}{(1 - \cos \theta)^2} (1 - \cos \theta)^2 \quad (5.101)$$

Which after some expanding and cancellation, becomes:

$$\frac{\Omega_0}{a} + 1 - \Omega_0 = \frac{\Omega_0 - 1}{(1 - \cos \theta)^2} (1 - \cos^2 \theta) \quad (5.102)$$

Using the famous trig identity  $\sin^2 \theta + \cos^2 \theta = 1$ , we can identify (5.99) with (5.102) to conclude that this is indeed the correct solution.

For the plot, we observe that the given parametric equations are exactly of those for a cycloid, so the curve of  $a$  vs.  $t$  will be exactly that of a cycloid. It is given below.

TODO

Figure 6: ;caption;

## 6 Measuring Cosmological Parameters

### 6.1 Magnitudes and Polar Bear Feet

First solving for the bolometric absolute magnitude of the bear's foot, we have:

$$M_B = -2.5 \log_{10} \left( \frac{L}{L_\odot} \right) = -2.5 \log_{10} \left( \frac{10W}{78.7 L_\odot} \right) = -2.5 \log_{10} \left( \frac{10W}{78.7 \cdot 3.82 \times 10^{26} W} \right) \quad (6.1)$$

Numerically, this is:

$$M_B = 68.7 \quad (6.2)$$

Now for the apparent magnitude at luminosity distance of  $d_L = 0.5\text{km}$ . We first calculate the flux of the polar bear foot to be

$$f_B \frac{L}{4\pi d_L^2} = \frac{10W}{4\pi (500\text{m})^2} = 3.18 \times 10^{-6} \text{Wm}^{-2} \quad (6.3)$$

So therefore calculating the apparent magnitude, we have:

$$m_B = -2.5 \log_{10} \left( \frac{f_B}{f_x} \right) = -2.5 \log_{10} \left( \frac{3.18 \times 10^{-6} \text{Wm}^{-2}}{2.53 \times 10^{-8} \text{Wm}^{-2}} \right) \quad (6.4)$$

Which numerically we calculate to be:

$$m_B = -5.25 \quad (6.5)$$

If we have a bolometer that can detect the the bear's foot at a maximum luminosity distance of  $d_L = 0.5\text{km}$ , then  $f_B$  (the flux of the bear foot at this distance) defines the maximum flux  $f_{\text{max}}$  that is capable of being detected, so solving for the maximum luminosity distance it could detect the sun, we have:

$$d_{\text{max, sun}} = \sqrt{\frac{L_\odot}{4\pi f_{\text{max}}}} = 3.09 \times 10^{15} \text{m} \quad (6.6)$$

Repeating this calculation for the supernova, we have:

$$d_{\text{max, supernova}} = \sqrt{\frac{L_{\text{supernova}}}{4\pi f_{\text{max}}}} = \sqrt{\frac{4 \times 10^9 L_\odot}{4\pi f_{\text{max}}}} = 1.96 \times 10^{20} \text{m} \quad (6.7)$$

### 6.2 Angular Size and Polar Bear Feet

We know that:

$$d_A = \frac{l}{\delta\theta} \quad (6.8)$$

so given  $l = 0.16\text{m}$  and  $d_A = 0.5\text{km}$ , we can rearrange to solve for the angular size to be:

$$\delta\theta = \frac{l}{d_A} = \frac{0.16\text{m}}{500\text{m}} = 3.2 \times 10^{-4} \text{rad} \quad (6.9)$$

The critical redshift of the benchmark model is  $z_C = 1.6$ , where  $d_{A,\text{max}} = 5.31 \times 10^{25} \text{m}$ , so:

$$\delta\theta_{\text{max}} = \frac{l}{d_{A,\text{max}}} = 3 \times 10^{-27} \text{rad} \quad (6.10)$$

### 6.3 Maximizing $d_A$ in a flat, single-component universe

In a spatially flat, single-component universe, the scale factor is given as (Ryden Eq. 5.39):

$$a(t) = \left( \frac{t}{t_0} \right)^{2/(3+3w)} \quad (6.11)$$

So we can integrate to obtain the proper distance (Ryden Eq. 5.49):

$$d_p(t_0) = c \int_{t_e}^{t_0} \frac{dt}{a(t)} = c \int_{t_e}^{t_0} \left( \frac{t}{t_0} \right)^{-2/(3+3w)} dt = ct_0 \frac{3(1+w)}{1+3w} \left[ 1 - (t_e/t_0)^{(1+3w)/(3+3w)} \right] \quad (6.12)$$

Furthermore, using that  $1+z = \frac{a(t_0)}{a(t_e)} = (t_0/t_e)^{2/(3+3w)}$  (Ryden Eq. 5.47) to find  $t_e$ , we obtain (Ryden Eq. 5.48):

$$t_e = \frac{t_0}{(1+z)^{(3+3w)/2}} = \frac{2}{3(1+w)H_0} \frac{1}{(1+z)^{3(1+w)/2}} \quad (6.13)$$

Substituting this into (??) we find the current proper distance in terms of  $z$ :

$$d_p(t_0) = \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z)^{-(1+3w)/2} \right] \quad (6.14)$$

Further, in the case of a spatially flat universe, we can use Ryden Eq. 6.37 to obtain the current angular and luminosity distances:

$$d_A(t_0) = \frac{d_p(t_0)}{1+z} = \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z)^{-(1+3w)/2} \right] \frac{1}{1+z} \quad (6.15)$$

$$d_L(t_0) = d_p(t_0)(1+z) = \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z)^{-(1+3w)/2} \right] (1+z) \quad (6.16)$$

To solve for the critical redshift  $z_c$  where  $d_A$  has the maximum value, we take the derivative of (6.15) with respect to  $z$  and set it to 0:

$$\frac{dd_A}{dz} = \frac{c}{H_0} \frac{2}{1+3w} \left( -(1+z)^{-2} + \frac{3+3w}{2} (1+z)^{-(5+3w)/2} \right) = 0 \quad (6.17)$$

The terms in the brackets must vanish, so:

$$-(1+z)^{-2} + \frac{3+3w}{2} (1+z)^{-(5+3w)/2} = 0 \quad (6.18)$$

Cancelling out a factor of  $(1+z)^{-2}$ , we get:

$$\frac{3+3w}{2} (1+z)^{-(1+3w)/2} = 1 \quad (6.19)$$

Now rearranging to solve for  $z_c$ , we find:

$$z_c = \left( \frac{2}{3+3w} \right)^{(1+3w)/2} - 1 \quad (6.20)$$

We can now substitute this back into (6.15) to find what the maximum redshift is:

$$d_{A,\max} = \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z_c)^{-(1+3w)/2} \right] \frac{1}{1+z_c}$$

$$= \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - \left( \left( \frac{2}{3+3w} \right)^{(1+3w)/2} \right)^{-(1+3w)/2} \right] \left( \frac{2}{3+3w} \right)^{-(1+3w)/2} \quad (6.21)$$

So we conclude:

$$d_{A,\max} = \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - \left( \frac{2}{3+3w} \right)^{-(1+3w)^2/4} \right] \left( \frac{2}{3+3w} \right)^{-(1+3w)/2} \quad (6.22)$$

## 6.4 Difference between relative and absolute magnitude

At small redshift ( $z \ll 1$ ), the luminosity distance is approximately (Ryden Eq. 6.50, also seen in class):

$$d_L \approx \frac{c}{H_0} z \left( 1 + \frac{1-q_0}{2} z \right). \quad (6.23)$$

By Ryden Eq 6.49, the distance modulus is given by:

$$m - M = 5 \log_{10} \left( \frac{d_L}{1 \text{Mpc}} \right) + 25. \quad (6.24)$$

Substituting the first equation into the second, we get:

$$m - M \approx 5 \log_{10} \left( \frac{\frac{c}{H_0} z \left( 1 + \frac{1-q_0}{2} z \right)}{1 \text{Mpc}} \right) + 25 \quad (6.25)$$

From here on out, we will suppress the 1Mpc in the denominator for clarity:

$$m - M \approx 5 \log_{10} \left( \frac{c}{H_0} z \left( 1 + \frac{1-q_0}{2} z \right) \right) + 25 \quad (6.26)$$

Using that  $\log(ab) = \log(a) + \log(b)$  we get:

$$m - M \approx 5 \log_{10} \left( \frac{c}{H_0} z \right) + 5 \log_{10} \left( 1 + \frac{1-q_0}{2} z \right) + 25 \quad (6.27)$$

Using the approximation  $\log_{10}(1+x) \approx 0.4343 \ln(1+x) \approx 0.4343x$  for small  $x$  on the second term, we get:

$$m - M \approx 5 \log_{10} \left( \frac{c}{H_0} z \right) + 5(0.4343) \left( \frac{1-q_0}{2} z \right) + 25 \quad (6.28)$$

Simplifying the numerical terms in the second term, and multiplying by one in the argument of the first term, we get:

$$m - M \approx 5 \log_{10} \left( \frac{c}{H_0} z \frac{68 \text{km s}^{-1} \text{Mpc}^{-1}}{68 \text{km s}^{-1} \text{Mpc}^{-1}} \right) + 1.086(1-q_0)z + 25 \quad (6.29)$$

Further application of the  $\log(ab) = \log(a) + \log(b)$  rule yields:

$$m - M \approx 5 \log_{10} \left( \frac{c}{68 \text{ km s}^{-1} \text{ Mpc}^{-1}} \right) + 5 \log_{10} z + 5 \log_{10} \left( \frac{68 \text{ km s}^{-1} \text{ Mpc}^{-1}}{H_0} \right) + 1.086(1 - q_0)z + 25 \quad (6.30)$$

Applying the  $\log\left(\frac{1}{x}\right) = -\log(x)$  rule we get:

$$m - M \approx 5 \log_{10} \left( \frac{c}{68 \text{ km s}^{-1} \text{ Mpc}^{-1}} \right) + 5 \log_{10} z - 5 \log_{10} \left( \frac{H_0}{68 \text{ km s}^{-1} \text{ Mpc}^{-1}} \right) + 1.086(1 - q_0)z + 25 \quad (6.31)$$

Before we evaluate the first term numerically, we recall the suppressed Mpc, so:

$$m - M \approx 5 \log_{10} \left( \frac{300000 \text{ km s}^{-1}}{68 \text{ km s}^{-1}} \right) + 5 \log_{10} z - 5 \log_{10} \left( \frac{H_0}{68 \text{ km s}^{-1} \text{ Mpc}^{-1}} \right) + 1.086(1 - q_0)z + 25 \quad (6.32)$$

We use a calculator to check the value of the first term:

$$m - M \approx 5(3.6446) + 5 \log_{10} z - 5 \log_{10} \left( \frac{H_0}{68 \text{ km s}^{-1} \text{ Mpc}^{-1}} \right) + 1.086(1 - q_0)z + 25 \quad (6.33)$$

Grouping the numerical terms:

$$m - M \approx 43.23 - 5 \log_{10} \left( \frac{H_0}{68 \text{ km s}^{-1} \text{ Mpc}^{-1}} \right) + 5 \log_{10} z + 1.086(1 - q_0)z. \quad (6.34)$$

## 6.5 Surface Brightness

First, we recall the angular diameter distance  $d_A$  to a standard yardstick to be (Ryden Eq. 6.35):

$$d_A \equiv \frac{l}{\delta\theta} = \frac{S_\kappa(r)}{1+z} \quad (6.35)$$

Rearranging, we find:

$$\delta\theta = \frac{l(1+z)}{S_\kappa(r)} \quad (6.36)$$

The observed flux is related to the luminosity  $L$  and the observed flux  $f$  as (Ryden 6.27):

$$f = \frac{L}{4\pi S_\kappa(r)^2 (1+z)^2} \quad (6.37)$$

So,  $\Sigma$  as a function of redshift is:

$$\Sigma \propto \frac{f}{(\delta\theta)^2} = \frac{\frac{L}{4\pi S_\kappa(r)^2 (1+z)^2}}{\left(\frac{l(1+z)}{S_\kappa(r)}\right)^2} = \frac{L}{l^2 4\pi} \frac{1}{(1+z)^4} \propto \frac{1}{(1+z)^4} \quad (6.38)$$

So:

$$\Sigma = \frac{\Sigma_0}{(1+z)^4} \quad (6.39)$$

for some constant  $\Sigma_0$ . Since the surface brightness  $\Sigma$  only depends on the redshift and not cosmological parameters, we cannot use it to measure a cosmological parameter  $q_0$ .

## 6.6 Quasar Light Flux

The variation time scale at the time light was emitted is related to the variation timescale when it was observed via:

$$\delta t_0 = (1 + z)\delta t_e \quad (6.40)$$

So for redshift  $z = 5.0$  and  $\delta t_0$  of 3 days, the variation time scale when emitted is:

$$\boxed{\delta t_e = 0.5 \text{ days}} \quad (6.41)$$

$R_{\text{max}}$  for the observed quasar is:

$$\boxed{R_{\text{max}} = c(\delta t_e) = 6.48 \times 10^{12} \text{ m}} \quad (6.42)$$

From Ryden Figure 6.4, a standard yardstick with redshift  $z = 5.0$  has angular distance  $d_A \approx 0.3c/H_0$  in the Benchmark model, so the angular size is given by:

$$\boxed{\delta\theta = \frac{R_{\text{max}}}{d_A} = \frac{6.48 \times 10^{12} \text{ m}}{0.3c/H_0} = 1.58 \times 10^{-44} \text{ rad}} \quad (6.43)$$

## 6.7 Proper Area of a sphere

The FRW metric (Ryden Eq. 6.22) is:

$$ds^2 = -c^2 dt^2 + a(t)^2 [dr^2 + S_k(r)^2 d\Omega^2] \quad (6.44)$$

Expanding out  $d\Omega$ , this becomes:

$$ds^2 = -c^2 dt^2 + a(t)^2 [dr^2 + S_k(r)^2 d\theta^2 + S_k(r)^2 \sin^2 \theta d\phi^2] \quad (6.45)$$

For a space described by this metric, a surface element  $dA$  on a sphere of radius  $r$  will be given by:

$$dA = (S_k(r)d\theta)(S_k(r)\sin\theta d\phi) = S_k(r)^2 \sin\theta d\theta d\phi \quad (6.46)$$

Integrating this surface area to find the surface area of this sphere, we have:

$$A = \iint dA = \iint S_k(r)^2 \sin\theta d\theta d\phi = S_k(r)^2 \int_0^\pi \sin\theta d\theta \int_0^{2\pi} d\phi = S_k(r)^2 (2)(2\pi) = 4\pi S_k(r)^2 \quad (6.47)$$

Therefore setting  $r$  to be the proper radius  $r = d_p(t_0)$ , we obtain:

$$\boxed{A_p(t_0) = 4\pi S_k(r)^2} \quad (6.48)$$

## 6.8 Total Intensity of Standard Candles

The relation between the observed flux  $f$  and the Luminosity  $L$  of a distant light source is given by:

$$f = \frac{L}{4\pi S_k(r)^2 (1+z)^2} \quad (6.49)$$

Since we live in a flat universe,  $S_k(r) = r$  and so:

$$f = \frac{L}{4\pi r^2 (1+z)^2} \quad (6.50)$$

In a single-component universe, the proper distance  $r = d_p(t_0)$  for  $w \neq -1/3$  is given by (Ryden Eq 5.50):

$$r = d_p(t_0) = \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z)^{-(1+3w)/2} \right] \quad (6.51)$$

So substituting this into (6.50) we have:

$$f(z) = \frac{L}{4\pi \left( \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z)^{-(1+3w)/2} \right] \right)^2 (1+z)^2} \quad (6.52)$$

Which we can rewrite as:

$$f(z) = \frac{L(1+3w)^2}{16\pi(c/H_0)^2} \frac{1}{(1+z)^2} \left[ 1 - (1+z)^{-(1+3w)/2} \right]^{-2} \quad (6.53)$$

which was the desired expression. When  $w = -1/3$ , the scale factor in a spatially flat, single component universe is given by:

$$a(t) = \left( \frac{t}{t_0} \right)^{2/(3+3w)} = \left( \frac{t}{t_0} \right)^{2/(3-1)} = \frac{t}{t_0} \quad (6.54)$$

Therefore, the Hubble constant is given by:

$$H_0 = \left. \frac{\dot{a}}{a} \right|_{t=t_0} = \frac{\frac{1}{t_0}}{\frac{t_0}{t_0}} = \frac{1}{t_0} \quad (6.55)$$

The redshift is given by:

$$1+z = \left( \frac{t_0}{t_e} \right)^{2/(3+3w)} = \frac{t_0}{t_e} \quad (6.56)$$

Solving for the proper distance  $r = d_p(t_0)$  in this universe, we have:

$$r = d_p(t_0) = c \int_{t_e}^{t_0} \frac{dt}{a(t)} = ct_0 \int_{t_e}^{t_0} \frac{dt}{t} = ct_0 \ln \left( \frac{t_0}{t_e} \right) = \frac{c}{H_0} \ln(1+z) \quad (6.57)$$

Therefore the observed flux in this universe is given by:

$$f(z) = \frac{L}{4\pi(c/H_0)^2 \ln^2(1+z)} \frac{1}{(1+z)^2} \quad (6.58)$$

The number of stars located in the range  $r$  to  $r + dr$  in the sky per steradian will be given by:

$$N(r) = n_0 r^2 dr \quad (6.59)$$

So the number of stars in the range  $z$  to  $z + dz$  per steradian is given by:

$$N(z) = n_0 r^2 \frac{d}{dz} \left( \frac{c}{H_0} \frac{2}{1+3w} \left[ 1 - (1+z)^{-(1+3w)/2} \right] \right) \quad (6.60)$$

Taking the derivative:

$$N(z) = n_0 r^2 \left( \frac{c}{H_0} (1+z)^{-(3+3w)/2} \right) dz \quad (6.61)$$



So, finding the intensity from standard candles in the range  $z$  to  $z + dz$ , we multiply the earlier result for  $f(z)$  with the above result for  $N(z)$ :

$$dJ(z) = f(z)N(z) = \frac{L}{4\pi r^2(1+z)^2} n_0 r^2 \left( \frac{c}{H_0} (1+z)^{-(3+3w)/2} \right) dz \quad (6.62)$$

Which after cancelling terms:

$$\boxed{dJ(z) = \frac{n_0 L(c/H_0)}{4\pi} (1+z)^{-(7+3w)/2} dz} \quad (6.63)$$

which is the desired result. To find the total intensity  $J$ , we integrate over all redshifts:

$$J = \int_0^\infty dJ(z) = \int_0^\infty \frac{n_0 L(c/H_0)}{4\pi} (1+z)^{-(7+3w)/2} dz = \frac{n_0 L(c/H_0)}{4\pi} \frac{-(1+3w)}{2} (1+z)^{-(1+3w)/2} \Big|_0^\infty \quad (6.64)$$

This gives us the result:

$$J = \begin{cases} \frac{n_0 L(c/H_0)}{4\pi} \frac{1+3w}{2} & w > -\frac{1}{3} \\ \infty & w < -\frac{1}{3} \end{cases} \quad (6.65)$$

To obtain the  $w = -\frac{1}{3}$  result, we would have to repeat the analysis using (6.58), but we leave this as an exercise.

The result above tells us that the total intensity we have in a universe with  $w \leq -\frac{1}{3}$  is infinite! On some level this makes sense, as the horizon distance is infinite. However, in this universe, we claim the apparently paradoxical result that the brightness of the night sky is still finite. Why? Because the above calculation assumes that we see light flux from every single standard candle in the universe; this is simply NOT the case. There will be standard candles that block the sight of other standard candles to ours, so we simply do not see the light from every light source in the universe (and hence the above result is actually misleading, as it does not take into account the fact that light sources block other light sources). The maximum possible brightness we could have is if stars paved the sky (i.e. every sightline was blocked eventually by a star). In this scenario, we can repeat the calculation as done in Chapter 2 of Ryden. If a standard candle of radius  $R_*$  is at a distance  $r \gg R_*$ , its angular area in steradians is given by:

$$\Omega = \frac{\pi R_*^2}{4\pi r^2(1+z)^2} = \frac{R_*^2}{4r^2(1+z)^2} \quad (6.66)$$

and its measured flux will be:

$$f = \frac{L_*}{4\pi r^2(1+z)^2} \quad (6.67)$$

so the surface brightness of the star, in watts per square meter is:

$$\Sigma_* = \frac{f}{\Omega} = \frac{L_*}{\pi R_*^2} \quad (6.68)$$

which also gives the surface brightness of the paved sky, which while large, is most certainly finite!

## 6.9 Expansion Switch

The acceleration equation can be written as:

$$-\frac{\ddot{a}}{aH^2} = \frac{1}{2} \sum_{i=1}^N \Omega_i (1+3w_i) \quad (6.69)$$

where the sum is taken over the different components of the universe. In the Benchmark model, we have matter ( $w = 0$ ), radiation ( $w = \frac{1}{3}$ ), and the cosmological constant ( $w = -1$ ), and so:

$$-\frac{\ddot{a}}{aH^2} = \frac{1}{2}\Omega_m + \Omega_r - \Omega_\Lambda \quad (6.70)$$

We have that  $\Omega_m = \frac{\Omega_{m,0}}{a^3}$ ,  $\Omega_r = \frac{\Omega_{r,0}}{a^4}$ , and  $\Omega_\Lambda = \Omega_{\Lambda,0}$ , so:

$$-\frac{\ddot{a}}{aH^2} = \frac{1}{2}\frac{\Omega_{m,0}}{a^3} + \frac{\Omega_{r,0}}{a^4} - \Omega_{\Lambda,0} \quad (6.71)$$

Setting  $\ddot{a}$  to find the scale factor  $a$  for which the expansion of the universe switched from slowing down to speeding up, we have:

$$0 = \frac{1}{2}\frac{\Omega_{m,0}}{a^3} + \frac{\Omega_{r,0}}{a^4} - \Omega_{\Lambda,0} \quad (6.72)$$

Multiplying by  $a^4$  we get:

$$0 = \frac{1}{2}\Omega_{m,0}a + \Omega_{r,0} - \Omega_{\Lambda,0}a^4 \quad (6.73)$$

In the Benchmark model,  $\Omega_{m,0} = 0.31$ ,  $\Omega_{r,0} = 9.0 \times 10^{-5}$ , and  $\Omega_{\Lambda,0} \approx 0.69$ . The radiation term can be neglected to good approximation, yielding:

$$0 \approx a\left(\frac{1}{2}\Omega_{m,0} - \Omega_{\Lambda,0}a^3\right) \quad (6.74)$$

So solving for the positive root of this equation, we get:

$$a = \sqrt[3]{\frac{\Omega_{m,0}}{2\Omega_{\Lambda,0}}} \approx 0.608 \quad (6.75)$$

Which we note is less than the scale factor at matter-lambda equality of  $a_{m\Lambda} = 0.77$ .

## 7 Dark Matter

### 7.1 Dark Matter Candidates

THIS IS RIGHT, JUST FIX THE NUMBERS LATER

Also, you can use:

$$\frac{4}{3}\pi d^3 = \frac{4}{3}\pi R^3 \frac{m_{\text{BH}}}{m_{\text{gal}}} \quad (7.1)$$

Taking the radius of the halo to be  $R_{\text{halo}} \approx 75\text{kpc}$  and the mass of our galaxy to be  $M_{\text{gal}} \approx 9.6 \times 10^{11} M_{\odot}$ , we approximate that roughly all of the mass comes from dark matter, hence giving us  $N = M_{\text{gal}}/10^{-8} M_{\odot} = 9.6 \times 10^{19}$  black holes. The volume of our galaxy is given by:

$$V = \frac{4}{3}\pi R_{\text{halo}}^3 = 5.2 \times 10^{64} \text{m}^3 \quad (7.2)$$

The number density of these black holes in our galaxy is therefore given by:

$$n_{\text{BH}} = \frac{N}{V} = \frac{9.6 \times 10^{19}}{5.2 \times 10^{61} \text{m}^3} \approx 1.85 \times 10^{-41} \text{m}^{-3} \quad (7.3)$$

In other words, we can find one black hole per:

$$V_{\text{BH}} = \frac{1}{n_{\text{BH}}} = 5.42 \times 10^{40} \text{m}^3 \quad (7.4)$$

So the nearest black hole would (approximately) be a distance of:

$$d_{\text{BH}} \approx \sqrt[3]{V_{\text{BH}}} = 3.78 \times 10^{13} \text{m} \quad (7.5)$$

away. The mean free path before a black hole comes into a distance 1AU with our sun is given by:

$$\lambda_{\text{BH}} = \frac{1}{n_{\text{BH}}\sigma} = \frac{1}{n_{\text{BH}}\pi(1\text{AU})^2} = 7.67 \times 10^{17} \text{m} \quad (7.6)$$

So combining this with the solar galactic orbital speed of  $235\text{km s}^{-1}$ , we find that the frequency of such black hole pass-bys are given by:

$$f_{\text{BH}} \approx \frac{v}{\lambda_{\text{BH}}} = 3.1 \times 10^{-13} \text{Hz} \quad (7.7)$$

Repeating the exact same calculation for MACHOs with mass  $10^{-3} M_{\odot}$ , we find:

$$d_{\text{MACHO}} \approx \quad (7.8)$$

$$f_{\text{MACHO}} \approx \quad (7.9)$$

### 7.2 Draco Galaxy

If we assume that the velocity dispersion is isotropic, then the 3D RMS velocity is equal to the three times the 1D mean square velocity  $\sigma_r$ , so:

$$\langle v^2 \rangle = 3(10.5\text{km s}^{-1})^2 = 3.31 \times 10^8 \text{m s}^{-1} \quad (7.10)$$

Using the steady state virial theorem, we know that:

$$\frac{1}{2}M\langle v^2 \rangle = \frac{\alpha}{2} \frac{GM^2}{r_h} \quad (7.11)$$

Which we can rearrange for  $M$ :

$$M = \frac{\langle v^2 \rangle r_h}{\alpha G} \quad (7.12)$$

Assuming that  $\alpha = 0.45$ , we can calculate the mass of the Draco galaxy to be:

$$M = \frac{3.31 \times 10^8 \text{m s}^{-1} \cdot 120 \text{pc}}{0.45 * 6.67 \times 10^{-11} \text{m}^3 \text{kg}^{-1} \text{s}^{-2}} = 3.89 \times 10^{33} \text{kg} = 2.04 \times 10^7 M_\odot \quad (7.13)$$

The mass to light ratio is then:

$$\frac{M}{L} = \frac{2.04 \times 10^7 M_\odot}{1.8 \times 10^5 L_\odot} \approx 113 \frac{M_\odot}{L_\odot} \quad (7.14)$$

Some possible errors: The isotropy assumption of the velocity dispersion (on a local scale, the galaxy is certainly not isotropic). Assuming that the galaxy was in a steady state. Assuming  $\alpha = 0.45$ .

### 7.3 Gravitational Lensing of Earth

The local curvature of spacetime causes the photon to be deflected by angle:

$$\alpha_{\text{Earth}} = \frac{4GM}{c^2 R} = 2.8 \times 10^{-9} \text{rad} \quad (7.15)$$

For a white dwarf and a neutron star, we get:

$$\alpha_{\text{dwarf}} = 4.0 \times 10^{-4} \text{rad} \quad (7.16)$$

$$\alpha_{\text{neut}} = 0.74 \text{rad} \quad (7.17)$$

### 7.4 Halo Mass Density

For a spherically symmetric mass distribution, we can model the density as:

$$\rho(r) = \frac{1}{4\pi r^2} \frac{dM(r)}{dr} \quad (7.18)$$

Since  $M(r) = \frac{v^2 r}{G}$ ,  $\frac{dM(r)}{dr} = \frac{v^2}{G}$  and so the above becomes:

$$\rho(r) = \frac{v^2}{4G\pi r^2} \quad (7.19)$$

Note that we have assumed here that  $v$  is approximately constant with  $r$ . Indeed, we find that  $v(r) \approx 230 \text{km s}^{-1}$  out to  $r = 35 \text{kpc}$  for our galaxy, so we are justified as treating it as a constant in our derivation. Putting in this value for  $v$  and  $G$ , we get:

$$\rho(r) = \frac{1}{r^2} 6.3 \times 10^{19} \text{kg m}^{-1} = \frac{1}{r^2} 9.78 \times 10^{11} M_\odot \text{Mpc} \quad (7.20)$$

If we look at the mass density of the cosmological constant (assuming it to be uniform), we have:

$$\rho_\lambda = \Omega_\Lambda \rho_{\text{crit}} = 0.7 \cdot 1.28 \times 10^{11} M_\odot \text{Mpc}^{-3} \approx 10^{11} M_\odot \text{Mpc}^{-3} \quad (7.21)$$

So within our galactic halo (which only extends to  $\sim 75 \text{kpc}$ ), the mass density from the dark matter halo is evidently is much larger than the cosmological constant; it therefore shouldn't significantly affect the dynamics of our galaxy's halo.

## 7.5 Cluster Collisions

The number density of galaxies in this half-mass radius is:

$$n = \frac{N}{V} = \frac{N}{\frac{4}{3}\pi r_h^3} = 70.7 \text{Mpc}^{-3} \quad (7.22)$$

If the typical cross sectional area is  $\Sigma \approx 10^{-3} \text{Mpc}^2$ , then the mean free path of the Coma cluster before it hits another galaxy is:

$$\lambda = \frac{1}{\Sigma n} = \frac{1}{(10^{-3} \text{Mpc}^2)(70.7 \text{Mpc}^{-3})} = 14.1 \text{Mpc} \quad (7.23)$$

If the velocity dispersion of the Coma cluster is  $\sigma \approx 880 \text{km s}^{-1}$ , assuming isotropy we can obtain the 3D RMS velocity to be:

$$\langle v^2 \rangle = 3(880 \text{km s}^{-1})^2 = 2.32 \times 10^{12} \text{m}^2 \text{s}^{-2} \quad (7.24)$$

Then approximating  $\langle v \rangle \approx \sqrt{\langle v^2 \rangle}$ , we get:

$$\langle v \rangle \approx 1.52 \times 10^6 \text{m s}^{-1} = 4.9 \times 10^{-17} \text{Mpc s}^{-1} \quad (7.25)$$

Hence the average time between collisions is:

$$t = \frac{\lambda}{\langle v \rangle} = \frac{14.1 \text{Mpc}}{4.9 \times 10^{-17} \text{Mpc s}^{-1}} = 2.9 \times 10^{17} \text{s} = 9.2 \text{Gyr} \quad (7.26)$$

The Hubble time is  $H_0^{-1} \approx 14 \text{Gyr}$ , so  $t$  is less than that, but on the same order of magnitude.

## 7.6 Solar vs. CMB neutrinos

Assuming isotropic emission, neutrinos from the sun (by the time they reach Earth) are spread out over area:

$$A_{\text{shell}} = 4\pi R^2 \quad (7.27)$$

where  $R = 1 \text{AU}$ . Letting  $A_{\text{human}}$  be the cross-sectional area of a human being, the approximate number of solar photons that hit us per second is:

$$r_\nu = r_{\text{sun}} \frac{A_{\text{human}}}{A_{\text{shell}}} \quad (7.28)$$

Modelling the human body as a tube of length  $L_{\text{human}}$ , the time that a given photon will spend inside the body is:

$$t_\nu = \frac{L_{\text{human}}}{c} \quad (7.29)$$

So the number of neutrinos inside of us at any given moment will be:

$$N = r_\nu t_\nu = r_{\text{sun}} \frac{A_{\text{human}}}{A_{\text{shell}}} \frac{L_{\text{human}}}{c} = \frac{r_{\text{sun}} V_{\text{human}}}{4\pi R^2 c} \quad (7.30)$$

Approximating  $V_{\text{human}} \sim 0.1\text{m}^3$ , and using  $R = 1\text{AU} = 1.5 \times 10^{11}\text{m}$ ,  $c = 3.0 \times 10^8\text{m s}^{-1}$  and  $r_{\text{sun}} = 2 \times 10^{38}\text{neutrinos s}^{-1}$  (given in the question) we find:

$$\boxed{N_{\text{sun}} = 2.4 \times 10^5 \text{neutrinos}} \quad (7.31)$$

Which gives us our result for the number of solar neutrinos in our body at any given moment. The number density of neutrinos from the cosmic neutrino background is 3/11 times the number density of CMB photons (per neutrino flavour), so accounting for the 3 flavours, we get:

$$n_\nu = 3 \left( \frac{3}{11} \right) n_\gamma = \frac{9}{11} (4.108 \times 10^8 \text{m}^{-3}) = 3.36 \times 10^8 \text{m}^{-3} \quad (7.32)$$

So taking our volume to be roughly  $V \sim 0.1\text{m}^3$ , we have approximately:

$$\boxed{N_{\text{CNB}} = 3.36 \times 10^7 \text{neutrinos}} \quad (7.33)$$

cosmic neutrino background neutrinos inside of us at any moment. Hence there are around  $\boxed{100 \text{ times}}$  the amount of cosmic neutrino background neutrinos inside of us as there are solar.

## 8 The Cosmic Microwave Background

### 8.1 Baryon-to-Photon Ratio and Recombination Temperature

We recall the fractional ionization as solved for using the Saha equation:

$$X = \frac{-1 + \sqrt{1 + 4S}}{2S} \quad (8.1)$$

where  $S$  is given by:

$$S(T, \eta) = 3.84\eta \left( \frac{kT}{m_e c^2} \right)^{3/2} \exp\left(\frac{Q}{kT}\right) \quad (8.2)$$

where  $T$  is the temperature,  $\eta$  is the baryon-to-photon ratio, and  $Q$  is the ionization energy. We take  $k = 8.62 \times 10^{-5} \text{ eV K}^{-1}$ ,  $Q = 13.6 \text{ eV}$ ,  $m_e c^2 = 511000 \text{ eV}$ . For  $\eta = 4 \times 10^{-10}$  and for  $\eta = 8 \times 10^{-10}$ , we get:

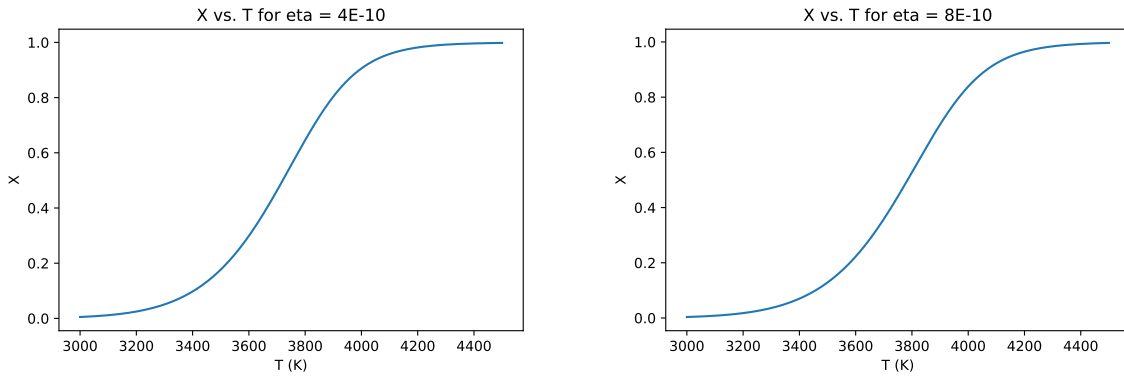


Figure 7: Plots of fractional ionization  $X$  as a function of temperature  $T$  (in Kelvin) for baryon-to-photon ratios  $\eta = 4 \times 10^{-10}$  and  $\eta = 8 \times 10^{-10}$ .

Taking  $T_{\text{rec}}$  to be when  $X = 1/2$ , for  $\eta = 4 \times 10^{-10}$ , we have  $T_{\text{rec}} = 3720\text{K}$ , and for  $\eta = 8 \times 10^{-10}$ , we have  $T_{\text{rec}} = 3784\text{K}$ . Doubling the photon-to-baryon ratio has a small effect (only a relative change of about 1.7%).

8.2

8.3

8.4

## 9 Nucleosynthesis and the Early Universe

9.1

9.2

9.3

9.4

### 9.5 Neutrino Detection

The cross-section for the interaction of a neutrino with a proton or neutron is:

$$\sigma_w \sim 10^{-47} \text{m}^2 \left( \frac{kT}{1\text{MeV}} \right)^2. \quad (9.1)$$

So for a typical CNB neutrino with  $E_\nu \sim kT_\nu \sim 5 \times 10^{-4} \text{eV}$ , we have:

$$\boxed{\sigma_w \sim 2.5 \times 10^{-66} \text{m}^2} \quad (9.2)$$

Fe-56 has 26 protons, 26 electrons, and 30 neutrons per atom. It has a per-atom weight of:

$$M = 26m_p + 26m_e + 30m_n \approx 56m_p = 56 \times 1.67 \times 10^{-27} \text{kg} = 9.34 \times 10^{-26} \text{kg} \quad (9.3)$$

So the atomic number density is:

$$n_a = \frac{\rho}{M} = \frac{7900 \text{kg m}^{-3}}{9.34 \times 10^{-26} \text{kg}} = 8.5 \times 10^{28} \text{m}^{-3} \quad (9.4)$$

The number density of protons/electrons is therefore:

$$\boxed{n_p = n_e = 26n_a = 2.2 \times 10^{30} \text{m}^{-3}} \quad (9.5)$$

and the number density of neutrons is therefore:

$$\boxed{n_n = 30n_a = 2.5 \times 10^{30} \text{m}^{-3}} \quad (9.6)$$

The total number density of sub-atomic particles is therefore:

$$n \sim n_p + n_e + n_n = 6.9 \times 10^{30} \text{m}^{-3} \quad (9.7)$$

though of course this is approximate; really all the subatomic particles are clustered in their respective atoms. The mean free path of a CNB neutrino in this medium is given by:

$$\boxed{\lambda = \frac{1}{n\sigma_w} = 5.8 \times 10^{34} \text{m}} \quad (9.8)$$

*Extremely* large; neutrinos are very hard to detect!



## 10 Inflation and the Very Early Universe

### 10.1

### 10.2

### 10.3

#### Gamow's CMB Prediction

A fascinating bit of cosmological history is that of George Gamow's prediction of the Cosmic Microwave Background in 1948. (Unfortunately, his prediction was premature; by the time the CMB was actually discovered in the 1960's, his prediction had fallen into obscurity.) Let's see if you can reproduce Gamow's line of argument. Gamow knew that nucleosynthesis must have taken place at a temperature  $T_{\text{nuc}} \simeq 10^9$  K, and that the age of the Universe is currently  $t_0 \simeq 10$  Gyr. Assume that the Universe is flat and contains only radiation. With these assumptions, what was the energy density  $\epsilon$  at the time of nucleosynthesis? What was the Hubble parameter  $H$  at the time of nucleosynthesis? What was the time  $t_{\text{nuc}}$  at which nucleosynthesis took place? What is the current temperature  $T_0$  of the radiation filling the Universe today? If the Universe switched from being radiation-dominated to being matter-dominated at a redshift  $z_{rm} > 0$ , will this increase or decrease  $T_0$  for fixed values of  $T_{\text{nuc}}$  and  $t_0$ ? Explain your answer.

## **11 Structure Formation: Gravitational Instability**

**11.1**

**11.2**

**11.3**

**11.4**

**11.5**

**11.6**

## **12 Structure Formation: Baryons and Photons**

**12.1**

**12.2**

**12.3**

**12.4**