The Dark Matter Halo of the Milky Way, AD 2013

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Abstract. We derive the mass model of the Milky Way (MW), crucial for Dark Matter (DM) direct and indirect detection, using recent data and a cored dark matter (DM) halo profile, which is favoured by studies of external galaxies. The method used consists in fitting a spherically symmetric model of the Galaxy with a Burkert DM halo profile to available data: MW terminal velocities in the region inside the solar circle, circular velocity as recently estimated from maser star forming regions at intermediate radii, and velocity dispersions of stellar halo tracers for the outermost Galactic region. The latter are reproduced by integrating the Jeans equation for every modeled mass distribution, and by allowing for different velocity anisotropies for different tracer populations. For comparison we also consider a Navarro-Frenk-White profile. We find that the cored profile is the preferred one, with a shallow central density of $\rho_H \sim 4 \times 10^7 M_{\odot}/\mathrm{kpc}^3$ and a large core radius $R_H \sim 10\,\mathrm{kpc}$, as observed in external spirals and in agreement with the mass model underlying the Universal Rotation Curve of spirals. We describe also the derived model uncertainties, which are crucially driven by the poorly constrained velocity dispersion anisotropies of halo tracers. The emerging cored DM distribution has implications for the DM annihilation angular profile, which is much less boosted in the Galactic center direction with respect to the case of the standard ACDM, NFW profile. Using the derived uncertainties we discuss finally the limitations and prospects to discriminate between cored and cusped DM profile with a possible observed diffuse DM annihilation signal. The present mass model aims to characterize the present-day description of the distribution of matter in our Galaxy, which is needed to frame current crucial issues of Cosmology, Astrophysics and Elementary Particles.

Keywords: Rotation curves of galaxies, Dark Matter experiments, Galaxy dynamics

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1 Introduction

The evidence for the phenomenon explained in term of Dark Matter in Galaxies is outstanding. In particular, spiral rotation curves (e.g. [4, 38]) have revealed the presence in these objects of a dark "mass component". The *current* framework includes a central bulge, a stellar disk and an extended gaseous disk, all embedded in a spherical halo made by dark particles (of yet) unknown nature. In external galaxies, the dark-luminous mass distribution has been derived by means of a) hundreths of individual rotation curves (e.g. [39]) b) 1000 *coadded* RCs [41] and c) various measurements of galaxy's gravitational potential (see e.g. [40]). The result is that DM halos show a shallow central density profile, moreover, stellar disk radii, halo core and viral radii, gaseous and stellar masses, turn out to be all related and to lead to an Universal Rotation Curve [41], likely the final state of complex physical processes governing the formation of galaxies.

Let us now consider the Milky Way. It is logic to ask: what is its distribution of dark and luminous matter? Does it conform to that of external galaxies? Moreover, living "inside" the object does help us in better disentangling the different mass components? Also of interest, do we know well enough how the DM particles are distributed around the Galaxy, in order to predict the signals of their annihilations to be expected in current experimental surveys?

It is not easy to address these questions and obtain a proper mass model of the Galaxy. The MW is an intrinsically complex object: observations are difficult to interpret due to our location within it. We cannot measure the MW circular velocity, and we must resort to more indirect kinematical measurements to infer its gravitational field. In fact, inside the solar circle (see e.g. [25]) we measure the rotating HI disk terminal velocities V_T that relate to the circular velocity profile V(r) (with r the galactocentric radius) only once one assumes some values for R_{\odot} , the distance of the Sun from

the Galaxy center, and V_{\odot} , the value of the circular velocity at the Sun's position (see section 3.1 below). Both quantities are presently known within an uncertainty of 10% (e.g. [26] that triggers a not negligible uncertainty in the derived circular velocity V(r). Outside the solar circle, the circular velocities V(r) are inferred by applying the Jeans equation to the kinematics of populations of tracer stars. Remarkably, the result strongly depends on the (unknown) dynamical state of tracers and it is also affected by uncertainties in their photometric and kinematical measurements [2, 6, 49]. The mass modeling method features (see [8], and e.g. [46]): a) an ad hoc halo density profile b) a number of assumptions on the dynamical status of the tracers and c) on their density distributions. All this adds to the complications due to the fact that the MW rotation curve is quite flat and therefore it is intrinsically difficult to decompose in its dark and luminous components [48].

However, there are today good motivations to attempt deriving a robust and reliable MW mass model, i.e.: to test claims of Gamma-ray emissions from annihilating DM particles in high density regions of the Galaxy, to investigate the Nature of the dark particles from the galactic DM density distribution, to understand how the Galaxy and the Local Group formed. A novel approach is warranted since in the past few years, modeling techniques and measurements have progressed.

First, let us note that the mass distribution in the Galaxy is likely to be similar to that of any other galaxy of the same mass/luminosity/dimensions [41]. More in detail, the DM halo density profile in spirals is represented by the URC halo profile [15]. The different profiles adopted in several previous works, (NFW and/or pseudo-isothermal halo) are not supported by present day observations in external galaxies. In addition, in considering the NFW halo profile one should confront with the standard results emerging from simulations [21], a procedure not always followed in previous works.

Second, trigonometric parallaxes and proper motions of sources of maser emission associated with high-mass star forming regions in the disk of the Galaxy have been recently available [7]. These measurements allows one to determine source distances and proper motions in a direct and geometrical way, and to obtain their full 3-dimensional locations and velocity vectors leading, for particular line of sights, to the actual Galaxy circular velocity. This provides us with a (limited) number of true and direct determinations of the velocity speed on the Galactic disk.

Third, a very extended and large amount of radial motions of stars, from which to extract the dispersion velocity, has now become available out to 80 kpc and more [13, 19, 49].

Finally, with a profile independent method, we derived a most conservative value of the DM density at the Sun $\rho_{DM}(R_{\odot}) = 0.43 \pm 0.11 \pm 0.11$ [42]. This measure will be used as a cross-check of the Galaxy modeling.

The aim of this work is to obtain, despite various uncertainties and difficulties still present, a robust MW mass model of the Galaxy, at the best of the knowledge of 2013. In the next section we will present the adopted mass components, in section 3 we describe the observational data, in section 4 we report the fits to these data and their results; in sections 5, 6 we describe the consequences for direct and indirect DM detection, and in 7 we discuss our conclusions.

2 Mass Components

In modeling the Galaxy, the main mass components are: the central bulge, the (thin) stellar disk and the DM halo. The observational uncertainties and the systematic biases on these components, which will be discussed below, make negligible (for our aims) other minor mass components present: the central bar, the thick and HI disks, and the stellar halo.

Bulge. The MW central bulge region has been subject to a number of probes which along the years have increased our knowledge of its mass distribution. Nevertheless, due to severe extinction

in the very central region in diverse wavelengths, there remain crucial uncertainties in the inner 1–2 kpc. In particular, the state of the art [3, 32, 37] knowledge seems to point quite clearly to a two-components structure: a more massive 'boxy' bulge extending up to ~ 1.5 kpc superimposed on a minor subleading 'bar', extending up to ~ 4 kpc. The estimates of the total mass of these components is however very difficult, because observations can only probe the mass density at a minimal angle of 2° from the galactic plane, i.e. is at the outer border of the boxy bulge distribution, thus requiring some sort of arbitrary modeling of the distribution in order to infer the total mass. Kinematical probes such as surveys of the rotational velocities are also limited by the large velocity anisotropies (see e.g. [47]). As a result, the total bulge mass has to be considered as an unknown parameter, subject to a lower bound which we take conservatively as ~ 0.8 –1 $\times 10^{10} M_{\odot}$.

Due to the above uncertainties and non circularities evident in the central region we will restrict our analysis of the rotation curve to radii larger than 2.5 kpc. At these distances, the shape of the mass distribution of the (dominant) bulge component is irrelevant, and the bulge can be safely be treated as a point-mass in the center. No results in this paper depend on the bulge density profile, as opposed to the value of its mass that plays a crucial role in mass modeling.

Disk. As in most of Spirals the MW surface density of the stellar disk is exponential, which for the thin disk we parametrize as [16]

$$\sigma_D = \frac{M_D}{2\pi R_D^2} e^{-r/R_D}.$$
 (2.1)

with $R_D = 2.6 \pm 0.5$ kpc [36].

The value of the thin disk mass M_D is rather unknown, although it can be constrained to lie in the range $M_D = 5-7 \times 10^{10} M_{\odot}$ [17, 32, 35, 36]. At the same time the estimate and uncertainty of the disk scale length R_D may be subject to significant biases (substructures, coverage, etc, see [36]). The disk thickness is constrained to be of the order of few hundreths of parsecs and in this work its presence and its variations have a negligible impact on the stellar contribution to the MW rotation curve (less than few %, see e.g. [42]). It is therefore justified to take the infinitesimal thin disk approximation. At the same time, the thick and HI disks, whose masses are estimated to be \sim one tenth of the thin disk one [28, 36] can be neglected in the present analysis and be considered to be part of the above range of M_D .

Dark Halo. We adopt, after [41] and [15], the Universal Rotation Curve (Burkert) profile. In addition, as a comparison, we consider also a NFW dark halo. These density profiles are parametrized through the density scale ρ_H and scale radius R_H

$$\rho_{NFW} = \rho_H \frac{1}{x(1+x)^2},\tag{2.2}$$

$$\rho_{Bur} = \rho_H \frac{1}{(1+x)(1+x^2)},\tag{2.3}$$

where $x = r/R_H$. Notice that R_H has a different physical significance in the two models: the radius at which $d \log \rho_{NFW}/d \log r = -2$ and the radius of the region of approximately constant density.

3 Observational Data

3.1 Terminal velocities

The circular velocities inside the solar circle R_{\odot} can be "reconstructed" from HI terminal velocities V_{Ti} as $(r_i, V_i) = (r_i, V_{Ti} + y_i V_{\odot})$, where $y_i = r_i/R_{\odot}$ and r_i is the galactocentric radius of the individual

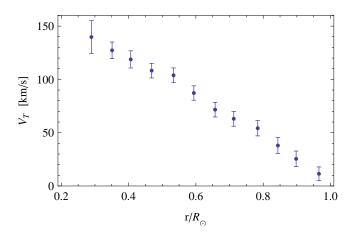


Figure 1. Terminal velocities, rebinned, see appendix A.

measured gas clouds. Notice that y_i is actually an angular measurement, with uncertainty which is negligible for the present work. The fit consists in comparing V_{Ti} with the terminal velocities as predicted by the model. In doing so, y_i , is translated into r_i , thus involving R_{\odot} in addition to V_{\odot} .

Data from a number of works [1, 12, 23, 25] are reported in appendix A and are binned in intervals of ~ 0.5 kpc, see figure 1. The binning procedure is delicate and we stress that some care has to be used in estimating the uncertainties later used for the fit. In fact, the errors on single measures are very small, even below 1 km/s, but data themselves are widely spread, due to differences between the values coming from different regions (e.g. at opposite longitudes). As a result, the uncertainties are mainly data-driven, and they have to be estimated accordingly. We believe that binning data separately for each measurements series provides a better description of the observative constraints with respect to the usual procedure of a global increase of the individual errors. We describe the binning in appendix A.

3.2 Maser velocities

Very precise measures of position and proper motions of maser sources, located at diverse galactocentric distances, e.g. [18, 34, 43] provide us with a number of circular velocities. While the galactocentric radii referred to these measurements are not crucial (in view of the almost flatness of the MW rotation profile) their amplitude is important in that it can resolve the degeneracy that V_{\odot} and R_{\odot} have. The outermost of such measurements reach the galactocentric radii of ~ 13 kpc and hint to a flat rotation curve, $V(r \simeq 13 \text{kpc}) \simeq V(R_{\odot})$. They also set the scale of the circular velocity, otherwise uncertain between 200 km/s and 250 km/s, to the upper end, $V_{\odot} \sim 250$ km/s

Thus, the two important outcomes of the maser observations [5, 7] are: i) the quantity $\omega = V_{\odot}/R_{\odot} = 30.3 \pm 0.9$, consistent with the previous determinations through the motion of Sagittarius A*, and ii) a direct estimate of the sun circular velocity of $V_{\odot} \simeq 239 \pm 7$ km/s, which is also in agreement with the knowledge of the Sun's galactocentric radius via ω (see e.g. [44]). See also [18] for more recent similar measurements.

These two results play an important role in the mass modeling: we will use i) as a constraint between V_{\odot} and R_{\odot} because its precision is much higher than the uncertainties in other quantities, and use ii) as the crucial measure of the Sun's LSR rotation velocity.

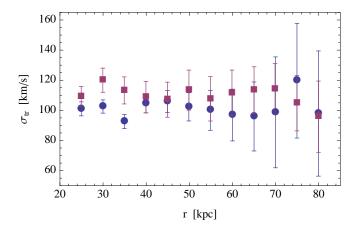


Figure 2. Tracers velocity dispersions: "1" Gnedin [19] (red squares), and "2" Xue [49] (blue circles).

3.3 Stellar halo velocity dispersions

From the measured radial velocities, the velocity dispersion $\sigma_i^2(r)$ of different halo populations of tracer (bright) stars have been estimated from the solar neighborhood out to $\sim 80 \,\mathrm{kpc}$. They provide unique constraints on the rotation curve at large radii and thus on the DM density profile. In particular, as we will see below, the halo scale radius R_H is significantly related to these measurements and to their uncertainties, that thus require a detailed discussion.

Assuming virialization, each population traces the gravitational potential, and we can use the spherical Jeans equation to link the measured velocity dispersion and the Galaxy gravitational potential. The density of each population is well represented by power law $\rho_i \propto r^{-\gamma_i}$, so that the Jeans equation can be written as

$$V^{2} = \sigma_{i}^{2} \left[\gamma_{i} - 2\beta_{i} - \frac{\partial \ln \sigma_{i}^{2}}{\partial \ln r} \right]. \tag{3.1}$$

Note that we must allow different β_i for each population of tracers, because although they share the same gravitational potential there is no argument for them to have the same kinematics, β_i and σ_i .

The Jeans equation (3.1) relates the circular velocity model $V^2(r)$, the velocity dispersions $\sigma_i(r)$ and the anisotropies β_i . Thus, for each given velocity model V(r) the equation can be integrated to find the predicted velocity dispersion as a function of radius $\sigma_i(r)$. For each population separately, the dispersion profile so computed from the model can then be directly compared with data. The integrated $\sigma_i(s)$ is determined up to a free constant multiplied by a known function, solution of the homogeneous equation, $\sigma_{hom,i}^2 \propto e^{\int (\gamma_i - 2\beta_i) d \ln r}$. Considering for instance a linear dependence of the anisotropy on the radius, $\beta_i(r) = \beta_i + \beta_i' r$, the function is simple, $\sigma_{hom,i}^2 = \bar{\sigma}_i^2 r^{\gamma_i - 2\beta_i} e^{-2\beta_i' r}$, where $\bar{\sigma}_i$ is the free constant, that we define so that $\sigma_i(80 \,\mathrm{kpc}) = \bar{\sigma}_i$. Although the constants $\bar{\sigma}_i$ enter the fit as additional free parameters of the mass model, their effect is actually very limited, because the dispersion kinematics is very uncertain at large radii (see figure 2). In practice, the effect of $\bar{\sigma}_i$ is just to reduce even more the constraining power of the data at large radii; So, for simplicity and without modifying the fit results, we set $\bar{\sigma}_{1,2} = 105 \,\mathrm{km/s}$. We checked directly that also the derivative β_i' of the tracers dispersion anisotropy has a negligible impact on the fits, due to the large uncertainties of the observed dispersion velocities and to the requirement that, at large radii, $\beta_i(80 \,\mathrm{kpc})$ should not be

¹We prefer to state explicitly this boundary condition at finite radius, rather than considering vanishing σ_i at infinity, which is not observationally motivated and may well turn out to be untrue.

unphysical. The dependence on β'_i can well be mimicked by a minor shift of the constant anisotropy β^0_i , and given the uncertain situation with the latter, we can safely set $\beta'_i = 0$.

The procedure of integrating the Jeans equation is of course equivalent to reconstructing *pseudo*-values of V^2 from the measures of σ^2 , which however would lead to increased and correlated errors due to the derivative in (3.1). For this reason we prefer to fit the actual kinematical data σ_i^2 .

We use two populations, the (mostly) BHB stars with velocities surveyed by HVS [19], hereafter "1", and the SDSS DR6 BHB stars [49] hereafter "2". Both series of binned data are taken as recalculated in [19] and are displayed in figure 2. For simplicity we consider spherical tracers density distributions, for which one has $\gamma_1 \approx 4$ and $\gamma_2 \approx 3.5$. These values provide a good spherical fit to BHB density and velocity dispersions as confirmed also by a recent detailed survey, see [13]. A very recent study using again BHB stars from the SDSS DR8 [20] confirms values for the radial velocity dispersions practically coincident with the older analysis of [49] that we use. This new analysis also provides an estimate of the dispersion anisotropies at various radii. For $r > 16\,\mathrm{kpc}$ it points to an average anisotropy of $\sim -0.6 \pm 0.5$. This is obviously an extremely useful information for determining the outer Galactic rotation curve.

We will not consider in this work the tracers velocity dispersions at small radii r < 25 kpc, because: i) data show a clear break both in the density and in the velocity dispersion [13]; ii) the dispersion anisotropy seems to vary rapidly in that region; iii) at such small radii the effect of the updated values of V_{\odot} and R_{\odot} may have to be reconsidered in the derivation of the dispersions from data, and iv) at these radii the masers information on V_{\odot} already provides a good constraint on the rotation curve.

From the two sets of data in figure 2 and the Jeans equation one can immediately infer some conclusion regarding the β_i . First, we note that both velocity anisotropies appear to have a fairly limited slope $|\partial \ln \sigma_i^2/\partial \ln r| \leq 0.1$, and their average values are $\sigma_i \sim 100$ –110. In the limit of flat rotation curve and considering the recent estimate of $V_{\odot} \simeq 240 \,\mathrm{km/s}$ (see previous section), one readily obtains from the Jeans equation (3.1) $\beta_i \simeq -0.5$ –0.8. So, the i.e. fairly tangential halo velocity dispersions. This agrees well with the evidence of a tangential outer halo [20]. Actually (see below) the preferred anisotropy would be even more tangential, i.e. $\beta_2 \lesssim -1$, due to the high V_{\odot} . In this respect, a decreasing rotation curve in place of a flat one at large radii can help relieving this tension.

As a result, the value of R_H resulting from the fit will be correlated with β_i , with higher β_i corresponding to smaller R_H . This will turn out to be the dominant source of uncertainty in the predicted values of R_H , as discussed below. The measure of β_2 in [20] is still preliminary; accordingly, we will use $\beta_2 \simeq -0.5 \pm 0.5$ as a fiducial range, keeping in mind possible updates in the future.

Moreover, we note that the dispersion velocities of population 1 (red squares) lie on average above those of 2 (blue circles). Because at the same time $\gamma_1 \sim 4$ is larger than $\gamma_2 \sim 3.5$, to trace the same gravitational potential i.e. to satisfy eq. (3.1) with a common $V^2(r)$, the only possibility is to have $\beta_1 - \beta_2 \sim 0.5$ or more. This is formally possible, but since the two samples are not relative to radically different populations, it is unlikely that their dispersion anisotropies differ so much. Actually, this difference may be a systematic bias, probably linked to the different ways the dispersions are estimated from observative data.³ Therefore, we prefer to leave some tension in the fit than stretching the anisotropy values, and adopt $\beta_1 = \beta_2 + 0.5$.

²In [13], the velocity dispersion anisotropy at large radii was estimated to be $β_2 \sim +0.5$, by using an incorrect lower rotation curve as a prior knowledge, in particular with $V_0 ≈ 220$ km/s.

³Considering for instance the 10-20% overall change in the velocity dispersions resulting from different estimation methods, as reported by [6].

4 Fit

Each choice of model parameters determines V(r), $V_T(r)$ and by means of the integration procedure described above, the velocity dispersions $\sigma_i^2(r)$. We fit

- the binned terminal velocities of figure 1 against those calculated from the modeled rotation curve;
- the tracers velocity dispersions of figure 2 against the dispersion profiles calculated from V(r);
- the circular velocities of the Maser regions;
- V_{\odot} as determined above from maser observations.

Denoting with x_i all these different quantities (evaluated at their radii) and with $x_{i,exp}$ their observational constraint, a standard $\chi^2 = \sum_{i=1}^{N} [x_i - x_{i,exp}]^2 / \delta x_{i,exp}^2$ is employed, and with our data, N = 40.

The complete set of parameters is $\{\rho_H, R_H, M_B, M_D, R_D, R_O\}$ which specify the modeled galaxy rotation curve, plus the anisotropies $\{\beta_1, \beta_2\}$ of the two populations which determine the dispersion velocity profile. In practice however, the number of relevant parameters can be reduced. In fact,

- As discussed in section 3.3, we use $\beta_1 \beta_2 = 0.5$, and study β_2 in the range $\beta_2 = -0.5 \pm 0.5$;
- R_D is constrained by observations, as discussed in section 2; moreover the fits strongly prefer large values of R_D (especially in the NFW case) forcing it to be at the boundary of the allowed values;
- M_B and M_D are free parameters that lie between a minimum and a maximum value. As we will see below the fits prefer often the lowest values for them so the minimum is on the boundary, and they play effectively a minor role in the minimization;
- In view of the quite precise knowledge of the sun's angular velocity, $V_{\odot}/R_{\odot} \simeq 30.3\pm0.3$ km/s kpc as recalled above, we express R_{\odot} in terms of V_{\odot} by means of the above relation.

Finally, through the mass modeling there is a one-to-one correspondence between ρ_H and the circular rotation at the sun $V_{\odot} = V(r_{\odot})$ (all the other parameters kept fixed), and it is technically convenient to trade ρ_H for V_{\odot} , which is used in its place as independent variable.

4.1 Galaxy mass modeling results

The fits have been performed with values of β_2 in a range of $\beta_2 \in [-1, 0]$, and we report in figure 3 the resulting best fit halo parameters and confidence intervals for each case of $\beta_2 = 0$, -0.5, -1 (left to right in each frame). In figure 4 we show the central (preferred) best fit Galactic rotation curves, and in figure 5 we show the sections of χ^2 , describing also the effect of the other astrophysical parameters and their correlations.

As discussed above, given the estimate of $\beta_2 \simeq -0.5$ in [20], we consider this as our preferred fit point, and the other two as extremal cases. With these three choices, the values of the best reduced $\chi^2_{DOF=40}$ are respectively 0.59, 0.41, 0.35 for the Burkert profile, and 0.9, 0.46, 0.35 for the NFW profile. The fit is further improved with even more negative values of β , a fact which hints to possible tension between data and the model. Thus a careful reconsideration of the uncertainties in the tracers velocity dispersions could lead to a softening of this problem. Nevertheless, due to the low values of $\chi^2_{DOF} < 1$, we can say that both models are able to provide reasonable fits to the data.

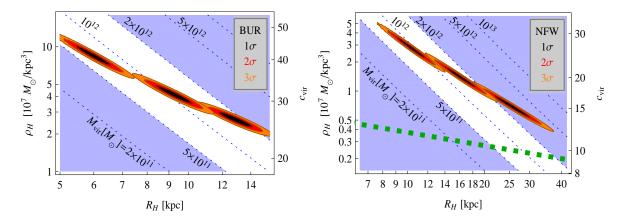


Figure 3. Best fits and confidence regions in the ρ_H – R_H plane at 1, 2, 3 σ C.L. (black, red, orange, yellow, two-parameter contours) for the Burkert (left) and NFW (right) profiles. In each frame the three regions correspond to tracers dispersion anisotropy $\beta_0 = 0, -0.5, -1$ (left to right). The superimposed blue (solid and dotted) contours mark $M_{vir}[10^{12} M_{\odot}] = 0.1-10$ and the shaded blue region is disfavored by estimates of the total MW mass. The thick green dotted line shows the c_{vir} - M_{vir} relation as resulting from cosmological simulations. Parameters other than R_H and ρ_H are fixed at their best fit preferred values (see text).

From figure 3 we note that, as anticipated, the predicted values of R_H are strongly dependent on the values of the dispersion anisotropy β_2 . The possible variation of the latter is thus the main source of uncertainty, driving the variation in R_H and the ranges presented in table 1 for the halo parameters and derived quantities. The observational (or theoretical) likelihood profile of β_2 is not known at present and thus one can not give a definite evaluation of confidence intervals by marginalizing on this dominant uncertainty. Nevertheless, the adopted range in β_2 allows us to estimate the expected ranges of parameters, in a consistent picture fitting all observational constraints, by extending the 95% C.L. fit uncertainties in a flat range of β_2 .

In figure 4 we display the best fit rotation curves, where the fitted terminal velocities, velocity dispersions, and maser circular velocities are compared with data. In figure 5 we display the cross correlations between the fit parameters M_D , M_B , R_D , R_H , V_{\odot} , by plotting the sections of χ^2 through the best fit point, as well as the constraints on astrophysical parameters. As a result of these constraints the best fits are sometimes constrained to lie just at the boundary of these regions. This is true in particular for R_D , and for M_B , M_D in the NFW case. In table 1 we report the DM halo parameters according to the preferred fit and also quote parameter uncertainties.

From the parameter variations in table 1 and from figures 3 and 5, we see that the Milky Way mass modeling is still quite uncertain, and this is mainly due to our ignorance of the dominant baryonic mass components (bulge and disk) and even more of the kinematics in the outer galaxy. The latter may be expected to improve in the near future due to more extensive stellar surveys and maser probes; on the other hand, the knowledge of the mass enclosed in the central region appears to be essentially hampered by the severe extinction at the relevant wavelengths. Still, some messages emerge from the fit

The best fit Burkert model features a total halo mass of $M_{vir} \sim 10^{12} M_{\odot}$, a central DM density of $\sim 4 \times 10^7 M_{\odot}/\text{kpc}^3$ and a core radius of ~ 9 kpc. The emerging mass model results very similar to that of the URC relative at spirals like the Galaxy, although in the two cases the gravitational potential is obtained in completely different ways. The variations of β_2 can significantly affect the mass modeling, and drive the total mass also to values larger than $2 \times 10^{12} M_{\odot}$. The Burkert best fit

		BUR	NFW
$ ho_H$	$[10^7 M_{\odot}/\mathrm{kpc}^3]$	$4.13^{+6.2}_{-1.6}$	$1.40^{+2.9}_{-0.93}$
R_H	[kpc]	$9.26^{+5.6}_{-4.2}$	16.1 ^{+17.} -7.8
V_{\odot}	[km/s]	$241.^{+11.}_{-7.4}$	244.+6.3
R_{\odot}	[kpc]	$7.94^{+0.36}_{-0.24}$	8.08 ^{+0.2} _{-0.2}
$ ho_{\odot}$	[GeV/cm ³]	$0.487^{+0.075}_{-0.088}$	$0.471^{+0.048}_{-0.061}$
$M_{r<50\mathrm{kpc}}$	$[10^{12}M_{\odot}]$	$0.45^{+0.35}_{-0.20}$	$0.48^{+0.20}_{-0.15}$
$M_{r<100\mathrm{kpc}}$	$[10^{12}M_{\odot}]$	$0.67^{+0.67}_{-0.33}$	$0.81^{+0.60}_{-0.32}$
$M_{ m vir}$	$[10^{12}M_{\odot}]$	$1.11^{+1.6}_{-0.61}$	$1.53^{+2.3}_{-0.77}$
$c_{ m vir}$	[Δ = 200]	$31.4^{+13.}_{-5.3}$	$20.1^{+11.}_{-7.1}$

Table 1. Best fit halo parameters and uncertainties for the Burkert and NFW profiles. The best fit values are relative to $\beta_1 = -0.5$, while the reported ranges correspond to 2σ intervals (95.45% C.L.) calculated in the *two-parameter* space V_{\odot} - R_H (see figure 5) by considering also the range $\beta_1 = [0, -1]$. As a result, quoted uncertainties should not be considered as statistical gaussian intervals, but only as estimates of expected values.

values in table 1 give $\text{Log}(\rho_H R_H/M_{\odot} \text{pc}^2) \simeq 2$, in perfect accordance with the trend obtained for other galaxies in a wide span of magnitudes [15]. This is also the correlation between ρ_H and R_H which can be inferred from the left figure 3, so that this relation is stable under variations of β_2 .

We also analysed the model featuring a NFW DM profile plus baryons, shown in the right panel of figure 3, as well as in the lower panels of figure 4 and 5. In figure 3 we also show the contours of the total MW virial mass M_{vir} and the concentration parameter c_{vir} on the right axis (using an overdensity parameter $\Delta_{vir} = 200$).

As one can see the allowed range of halo scale radii turns out to be quite wide, $R_H = 10\text{-}30\,\mathrm{kpc}$. In addition for NFW, unlike the Burkert case, the correlation between ρ_H and R_H in the right panel of figure 3, is of the form $\rho_H R_H^2 \sim \mathrm{const.}$, which reflects just the velocity constraints along the curve (both maser velocities and halo dispersions), in practice just a byproduct of the parametrization of the DM density through ρ_H , devoid of deep significance. Similarly for the V_{\odot} - R_H correlation in the central frames of lower figure 5.

In addition, two more remarks can be made. First, the NFW plus baryons best fit models require minimal bulge and disk masses as well as the largest disk scale radius, all at the boundary of the constraining intervals. This can be clearly appreciated in figure 5, where these parameters are varied and ther correlation shown. If one fixes the more reasonable value of $R_D \simeq 2.5$ kpc, the NFW fit worsens to $\chi^2_{DOF} \simeq 2.2$, corresponding to a probability of less than 10^{-4} (the Burkert fit instead raises to $\chi^2_{DOF} = 0.7$ which is still very good). Thus, the acceptable fits of the NFW plus baryons model are only achieved by stretching the astrophysical parameters to their allowed values.

Second, a fairly high value for $c_{vir} \simeq 20 \pm 10$ is obtained for the NFW profile. Current Λ CDM cosmological N-Body simulations imply the relation $c_{vir} = 9.6 M_{vir}/(h^{-1}10^{12} M_{\odot})$ [33] (with $\Delta_{vir} = 200$), which is also depicted on figure 3 for comparison. The Galaxy reasonably has $M_{vir} \gtrsim 10^{12} M_{\odot}$,

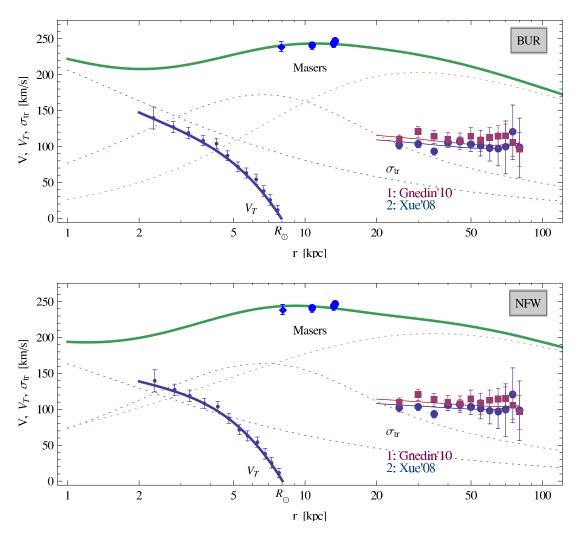


Figure 4. Best fit rotation curve, terminal velocities and velocity dispersions for the URC Burkert profile (upper panel) and for the NFW profile (lower panel), compared with the adopted observational data on terminal velocities, maser data and halo tracers velocity dispersions (left to right). Dotted curves show the bulge, disk and DM halo separate contributions (left to right)

so according to simulations we should have found $c_{vir} < 10$, a value significantly lower than the best fit one. If we impose the above c_{vir} -mass relationship, the resulting NFW Milky-Way mass model becomes an unacceptable representation of data. In table 1 we also report the predicted value of $c_{vir} \sim 31^{+13}_{-5.3}$ for the Burkert fit; we stress that, because in the URC scenario c_{vir} is just a structural parameter, so far unconstrained by galaxy formation scenarios, this value should not be compared with Λ CDM simulations which only produce cusped DM profiles, and instead should be taken as the outcome (prediction) of our fitted MW model.

It is important to stress that the uncertainty in the tracers anisotropy discussed above can not be invoked to release this tension of Λ CDM simulations, since a) it would require very negative values such as $\beta_2 \sim -3$, i.e. the existence of very tangential motions at odds with results of cosmological simulations, and b) the total MW mass would end up being unacceptably high, $\sim 10^{13} \, M_{\odot}$. Considering that the NFW best fits are already obtained by stretching astrophysical parameters to their extrema, models that solve the high concentration issue would also be too contrived.

For these reasons, we can conclude that present data tend to prefer a cored profile with respect to a cusped one. Part of this preference can be tracked down to the terminal velocities, which require the lowest possible mass enclosed in the inner region (r < 8 kpc), given that the disk and bulge already provide enough gravitational potential (see figure 5).⁴ Acceptable fits can still be achieved, but at the price of finding a solution in tension with the ones arising within the Λ CDM paradigm.

The value of the DM density at the Sun's radius ρ_{\odot} is also an outcome of the fit, and it turns out to be consistent with the profile independent determination derived in [42] (see discussion below).

Recalling that the main source of uncertainty in our fit comes from the poor knowledge of the outer tracers anisotropies β_i and the new recent knowledge on the latter, it is useful to compare our findings with the results of some other works.

In [9], MW mass models were built in the process of estimating of DM local density. They have considered both NFW and URC Burkert dark matter halos and tested them with a number of dynamical observables for the Galaxy (similar but not coincidental with those in the present paper, in particular with a different likelihood function). Both models have been found to agree with their observational data, with values of parameters somewhat similar but coincidental with ours. Among those an higher value of c_{vir} . In [10], with a similar global fit, they study also the DM phase space distribution. In view of the uncertainties in the halo tracers dynamics and also in the other available observative constraints, we consider quite premature such an analysis.

By assuming state-of-the-art models for the distribution of baryons in the Galaxy, [22] used microlensing and dynamical observations of the Galaxy to show that these data are in good agreement with the predictions of Λ CDM Dark Matter profiles.

In [14] 2000 distant Blue Horizontal Branch stars at Galactocentric distances of 15 kpc < r < 50 kpc were used as kinematic tracers of the MilkyWay mass distribution. Their density was modeled as an oblate, power-law halo embedded within the spherical power-law DM potential. Be means of distribution function method they obtained the power-law potential exponent and the velocity anisotropy of the halo tracers. The resulting outer circular velocity profile for the Milky Way when reproduced by a NFW halo, led to a high concentration value ($c_{vir} = 20$) in agreement with present results.

In [27] a simple method for fitting a MW mass model to observational constraints was presented (not too different from those we used in this paper). By means of a Bayesian approach they took input from observational data and expectations from the self consistent NFW plus disc model. They also forced the value of c_{vir} to lie in a range compatible with N-Body simulations, i.e. between 7 and 12. For the values of the disc scale lengths of 3.00 ± 0.22 kpc Solar galactocentric distance of 8.29 ± 0.16 kpc and a circular speed at the Sun of 239 ± 5 km/s; MW stellar mass of $6.43 \pm 0.63 \times 10^{10} M_{\odot}$; a virial mass of $1.26 \pm 0.24 \times 10^{12} M_{\odot}$ and a local dark matter density of 0.40 ± 0.04 GeV/cm³ their model fits observations.

It is also worth to stress that the Deason, McMillan or Iocco approaches are not much different from ours, but they did not consider cored DM profiles. These works, with the present one, support the view that in the MW a NFW plus baryon model can reproduce data, but with somewhat stretched values of the astrophysical parameters.

⁴In fact, we performed a similar analysis for the Einasto profile, finding even worse fits in the inner region.

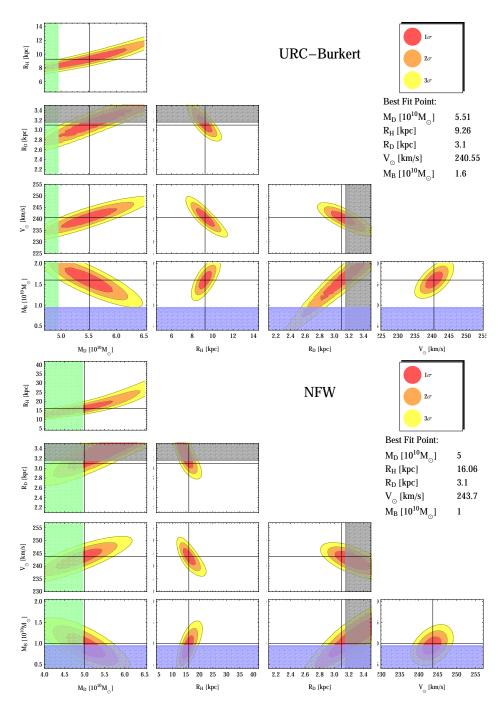


Figure 5. Two dimensional sections of χ^2 for the URC Burkert (upper) and NFW (lower) profiles, passing through the best fits, with fixed $\beta_2 = -0.5$. The shadowed regions mark the values of M_B (green left vertical band), M_D (blue bottom horizontal band) and R_D (the other grey bands), disfavored by bulge and disk studies. The contours mark $\Delta \chi^2 = (2.3, 6.18, 11.8)$, i.e. the *two-parameter* confidence levels of (68.27%, 95.45%, 99.73%).

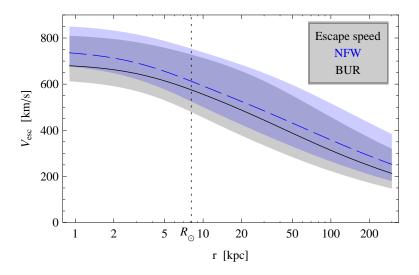


Figure 6. Galactic escape velocity as a function of galactocentric radius r for the best fit URC Burkert (black continuous) and NFW (blue dashed) models, with their 95% uncertainties in the halo parameters, including the uncertainty in the anisotropy β_2 . The dotted line marks the Sun's radius.

5 Direct DM search: local density and escape velocity

The derived mass models provide us with the local value of DM density. For the Burkert profile one finds $\rho_{\odot} \simeq 0.49^{+0.08}_{-0.09}\,\mathrm{GeV/cm^3}$, for the NFW values of $\simeq 0.47^{+0.05}_{-0.06}\,\mathrm{GeV/cm^3}$. These values are clearly consistent with the profile independent estimate of [42], as it must be since the gauss law and formula (11) there are satisfied. Nevertheless, the present global MW mass modeling has helped in reducing the uncertainties, from 0.2 to roughly 0.1 $\,\mathrm{GeV/cm^3}$, because the global fit constrains the values of M_D , R_{\odot}/R_D , V_{\odot} and rotation-curve slope more strictly than the set of the upper-lower limits taken in [42]. The central value is also slightly higher, due to the increased rotation speed at the Sun V_{\odot} , which is now better known.

From the mass models one can derive the Galactic escape velocity, whose value at the Sun's location is important for direct searches through DM nuclear recoil, because it affects the DM velocity distribution. We plot in figure 6 the escape velocity as a function of radius, out to 300 kpc. One can observe that at the Sun's location the Burkert and NFW best fits lead to quite similar estimates of the escape velocity, we have $V_{esc}^{BUR}(R_{\odot}) \simeq 576 \pm 124$ km/s, and $V_{esc}^{NFW}(R_{\odot}) \simeq 613 \pm 114$ km/s. These values are quite large and reflect the uncertainty due to the tracers anisotropy β_2 . They are however in agreement with recent estimates of the local escape speed using RAVE data [45]. Considering the overall variation of escape speed among the two models, one finds a large uncertainty of ~ 250 km/s, i.e. roughly 50%. This should be taken into account in direct search experiments, especially in the regime of light "WIMP" DM mass of the order of ~ 10 GeV or less, where one relies on the upper tail of the velocity distribution to estimate the number of expected recoil events (see e.g. the analysis in [24], [10]).

The uncertainty in the escape velocity at large radii r > 25 kpc is also considerable, (i.e. from 300 to $600 \, \text{km/s}$ at $50 \, \text{kpc}$) and this has an impact when estimating the halo dispersion velocities, through the selection of outliers versus acceptable halo tracers. Clearly in the future a modeling with this feedback would give a more self-consistent picture (see e.g. [6] for such an attempt).

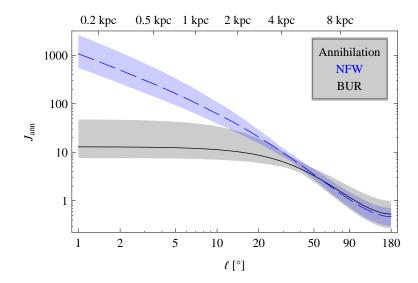


Figure 7. The "prompt" emission factor from DM annihilation as a function of the Galactic longitude, for the best fit Burkert (black continuous) and NFW (blue dashed) models, with their 2σ regions (95.45% C.L.). Note that for gamma rays in for instance the Fermi detector, below $< 0.1-1^{\circ}$ the point-spread function would smear the observed profile, making it effectively cored for any DM profile.

6 Indirect DM search: annihilation

The flux from DM annihilation is conveniently expressed in terms of the "prompt" emission factor

$$J_{ann}(\ell) = \frac{1}{\bar{\rho}_{\odot}^2 \bar{R}_{\odot}} \int_{l.o.s.} \rho_H^2(\mathbf{x}) \, d\mathbf{x}, \qquad (6.1)$$

which we normalize by using $\bar{\rho}_{\odot} = 0.4 \,\text{GeV}$ and $\bar{R}_{\odot} = 8.3 \,\text{kpc}$. The factor J_{ann} , as a function of the longitude ℓ from the galactic center, traces directly the angular profile of the dominant observed flux from annihilation into gamma rays. For annihilation into other (charged) particles, which are then subject to bremsstrahlung, scattering with ISR, and nontrivial galactic diffusion, see [11].

In fig 7 we plot J_{ann} for the URC Burkert + baryons and for the NFW + baryons models. We see that for each mass model the uncertainties in the galactic parameters lead to variations of the expected flux of a factor of ~ 5 , in the innermost region. On the other hand, in direction of about 40–60° from the galactic center the flux is predicted within a factor of 2 only, and independently of the profile chosen.

In fact, the differences between the two different mass models emerge only at $\ell < 15^\circ$, and with a clear discriminating power only below $\ell < 10^\circ$, corresponding to r < 1.5kpc, i.e. inside the bulge region. One should thus bear in mind that this plot extrapolates the DM density profile in the very central region where observations can not constrain it. Indeed, the DM density in the bulge region or at shorter scales may well deviate from purely cored or NFW profile, still without modifying the present global fits.

For instance, if one is willing to consider the scenario in which a DM core results from baryon feedback (mainly supernovae explosions) which erases the density cusp during galaxy formation, the same mechanism may well leave a 'mini-cusp' in the central region, which would give a small contribution to the total mass inside the solar circle, but depending on the very inner density slope, may contribute to an evident annihilation signal from the inner zone, with a very localized source region of at most few degrees in angular size.

7 Conclusions

In this paper we have derived a model for the DM in the Milky Way, by adopting state of the art inner terminal velocities, probes of maser star forming regions, and outer stellar halo velocity dispersions. We performed our analysis for both a cored (Burkert) and cusped (NFW) profiles.

Our fit confirms that the presence of a dark component in the Galaxy is beyond any doubt. We derived the preferred ranges of DM halo parameters ρ_H , R_H and their correlation in figures 3, and reported the values and derived parameters in table 1. The results show that the DM halo is still subject to a considerable degree of uncertainty. This is due to the limited constraining power of the available observations, especially in the outer stellar halo region, and in the Galactic center, bulge region. The main source of uncertainty has been traced to the unknown velocity dispersion anisotropy of the tracers in the stellar halo, which is still very poorly constrained by observations, and which drives the uncertainty in the halo scale radius R_H . Nevertheless, thanks to recent more precise measurements of the rotation speed at the Sun's position V_{\odot} via maser emission in high star-forming regions, the fitted halo parameters allow to draw some interesting conclusions.

First, if one considers a cored profile like the adopted URC-Burkert one, very good fits are obtained, which can accommodate all available observations. The uncertainty in the tracers velocity anisotropy is still driving the DM halo parameters uncertainty, resulting in the ranges reported in table 1, with $R_H = 9.26^{+5.6}_{-4.2}$ kpc and $\rho_H = 4.13^{+6.2}_{-1.6}$ with quite a strict correlation of $\rho_H R_H \sim const$, even varying the tracers anisotropy β_2 . The value of $\rho_H R_H \simeq 10^2$ is in excellent agreement with Dark-Matter fits in many other similar external galaxies, and constitute a reference point on which one may base further analyses of the Dark Matter in our Galaxy.

We studied also the case of NFW profile and showed that, while it also can be used to fit the observations, this was only possible by stretching other astrophysical parameters to their allowed limits; e.g., the disk scale length had to be set to 3.1 kpc which is at its marginal upper limit, and the bulge and disk masses had to be as small as possible. This is due to the profile of inner terminal velocities, which require as low mass as possible in the central region, where instead the NFW cusp is. Indeed, a similar analysis for the Einasto profile shows a further worsening of this tension. In other words, the present data rather prefer a cored DM distribution, though the presence of standard NFW halo cannot be excluded and, if one decides to adopt it, a great care must be taken in choosing the values of the model parameters.

Also, an important evidence is that especially due to the high V_{\odot} value, the NFW halo turns out to have a high $c_{vir} \sim 20$, at odds with results from numerical simulations in the Λ CDM setup. Of course this might be due to the effect of dirty baryonic physics. Recent simulations have started to account for the effect of baryons which have a twofold effect, namely, to: i) adiabatically contract the halo, increasing the inner halo slope even beyond 1–2, and ii) expel the inner dark matter due to supernovae feedback. It is not clear whether the net final effect is to reduce the concentration c_{vir} even further with respect to the baryon-less case. In this scenario a DM profile much different from the original NFW one is likely to emerge and, in our opinion, to converge with the URC one.

In the light of the fitted halo parameters, we analyzed the predicted Galactic escape velocity profile and the flux from DM prompt annihilation. While consistent with estimates with different methods, the escape speed is found to be very uncertain, independently from the profile chosen: at the Sun's radius it ranges in the interval 500–750 km/s, which has an impact on direct dark matter searches, especially for low (WIMP) DM masses of the order of 10 GeV. The escape speed turns out to be quite uncertain also at large radii in the outer galaxy, with an impact on the removal of unbounded outliers from the samples of stellar halo tracers.

Regarding the flux from DM annihilation, we assessed the discriminating power on the cored versus cusped DM halo hypotheses of indirect DM searches toward the galactic center. In figure 7 we displayed the expected fluxes from prompt DM annihilation, showing that a discrimination is possible only in the region inside 10° (or 1.5 kpc) from the Galactic center. This is right in the bulge region, and may be subject to a considerable astrophysical background uncertainty. Let us also stress in this respect that the observational data are able to constrain the DM profile only at radii larger than about 2 kpc; but because the physics leading to the inner DM density profile is still not understood, it may well be the case that the DM profile deviates from the simple core or cusp.

The DM density at the Sun's location is also an outcome of the mass modeling, and the values found are $\rho_H(R_\odot) \simeq 0.49^{+0.08}_{0.09}$ for the Burkert profile (and $0.471^{+0.05}_{0.06}$ for NFW). They are clearly in agreement with the profile-independent determination [42], but its uncertainties are reduced by a factor of two from ~ 0.2 down to roughly $\lesssim 0.1 \, \text{GeV/cm}^3$. This is due to the fact that in the present global modeling the rotation curve is more constrained by inner and outer galactic observations and the sun's rotation speed is known better.

In the present study, we tried to assess realistic ranges for the DM halo parameters, but still being far from claiming to give precise statistical significance intervals. The reason is due to the lack of knowledge of realistic likelihood profiles for many important quantities (the stellar halo anisotropy parameters but also the disk and bulge masses) so that any arbitrary choice of ad-hoc priors would translate in unrealistic or unstable results. We try to avoid such unjustified predictions in parameters, both as a matter of principles and also to avoid complicating the development of the field of Galactic modeling as well as the one of dark matter detection.

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A Terminal velocities

We describe here the adopted data relative to terminal velocities in the inner Galactic region. We use available data from [1, 12, 23, 25] versus the absolute value of the longitude. Since the data show systematic variations due to local gas motion, it is important not to trust either the small individual errors (often smaller than 1 km/s) or the high number of data points. Indeed, even an infinite increase in the number of measures would only amount to a more precise map of the gas motion, and not imply a more precise determination of the circular velocity. As a result, the adopted uncertainty in the circular velocities should follow the dispersion of the data, which directly traces the local variations and does not become precise with an increase of data points. We perform a binning of ~ 0.5 kpc and consider the weighted average and the dispersion inside each bin as data for the global fit. The binned average is performed by first binning each dataset separately (to avoid dominance of more populated datasets) using original errors enlarged by 7 km/s (to avoid the dominance of sets with tiny experimental error). Then, inside each bin, we average the results from different datasets, and consider the weighted average and the dispersion inside each bin as data for the global fit. The resulting binned terminal velocities are shown in figure 1. In figure 8 we show also the original data, together with the binned velocities, where for convenience we have reconstructed the circular velocities for $V_{\odot} = 242 \text{ km/s}$ and $R_{\odot} = 8 \text{ kpc}$.

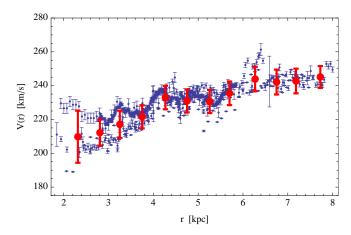


Figure 8. Circular velocities reconstructed from terminal velocities (blue, thin) and values obtained after their binning (red, thick). For $V_{\odot} = 242 \, \text{km/s}$ and $R_{\odot} = 8 \, \text{kpc}$.

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