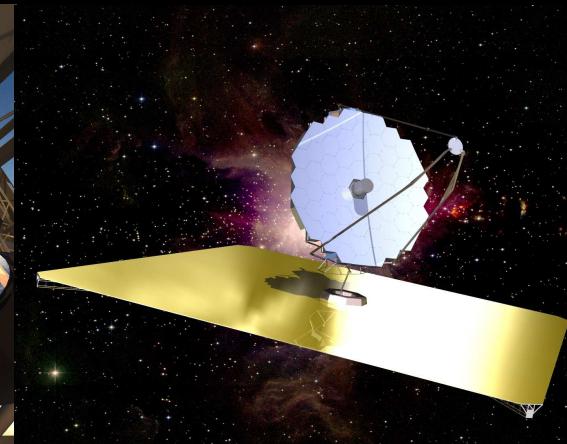
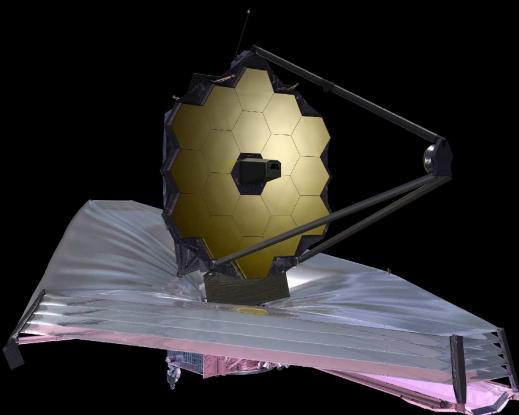


How will JWST measure First Light, Galaxy Assembly & Supermassive Blackhole Growth: New Frontiers after Hubble

Rogier Windhorst (ASU) — JWST Interdisciplinary Scientist

Collaborators: S. Cohen, L. Jiang, R. Jansen (ASU), C. Conselice (UK), S. Driver (OZ), & H. Yan (U-MO)

(Ex) ASU Grads: N. Hathi, H. Kim, M. Mechtley, R. Ryan, M. Rutkowski, B. Smith, & A. Straughn



1973~2018+;

1996~2029;

2000~2050+

2020~2050+?

ASU SESE Undergraduate Student Seminar, Camp SESE, Tontoazona, AZ

Saturday Sept. 12, 2015. All presented materials are ITAR-cleared.

JWST is like a hot bath. It feels good while you're in it; but the longer you stay, the more wrinkled you get.



WARNING: Both Hubble and James Webb are 30–40+ year projects:
You will feel wrinkled before you know it ... :)

Outline

- (1) Update on the James Webb Space Telescope (JWST), 2015.
- (2) What HST WFC3 has done: Measuring Galaxy Assembly and Supermassive Black-Hole Growth
- (3) How can JWST measure the Epochs of First Light & Galaxy Assembly, and Supermassive Black-Hole Growth?
- (4) The Future: Next generation 20-40 meter ground-based telescopes and ATLAST
- (5) Where do our students end-up? Possible NASA Careers
- (6) Summary and Conclusions.



Sponsored by NASA/HST & JWST



Edwin P. Hubble (1889–1953) — Carnegie astronomer

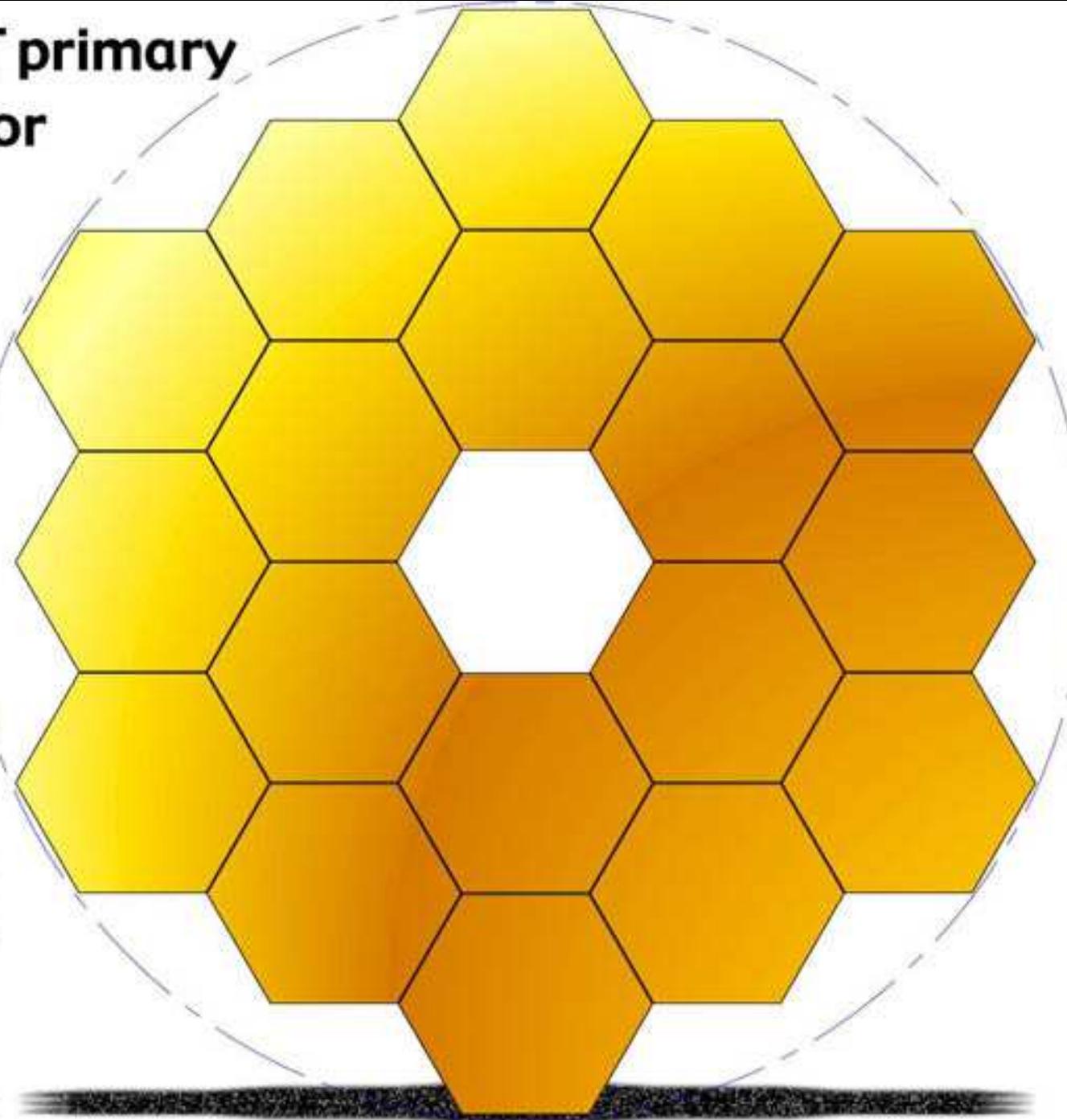


James E. Webb (1906–1992) — Second NASA Administrator

Hubble: Concept in 1970's; Made in 1980's; Operational 1990– \gtrsim 2015.

JWST: The infrared sequel to Hubble from 2018–2023 (–2029?).

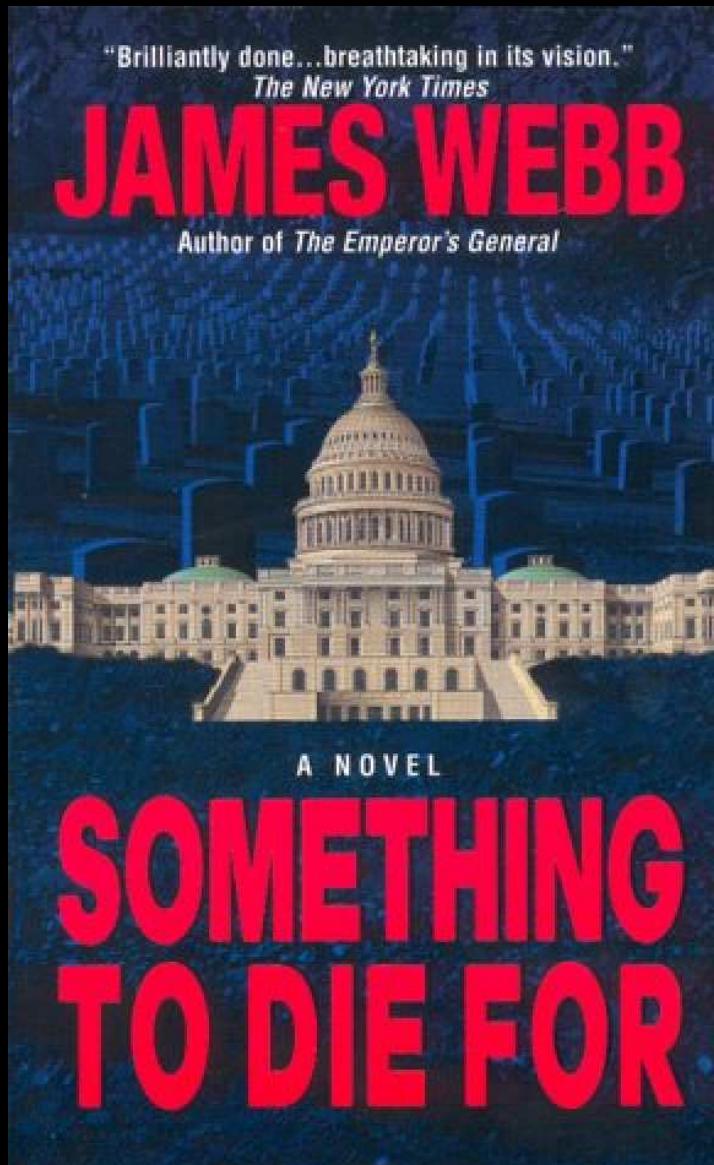
**JWST primary
mirror**



**Hubble primary
mirror**

JWST $\simeq 2.5 \times$ larger than Hubble, so at $\sim 2.5 \times$ larger wavelengths:
JWST has the same resolution in the near-IR as Hubble in the optical.

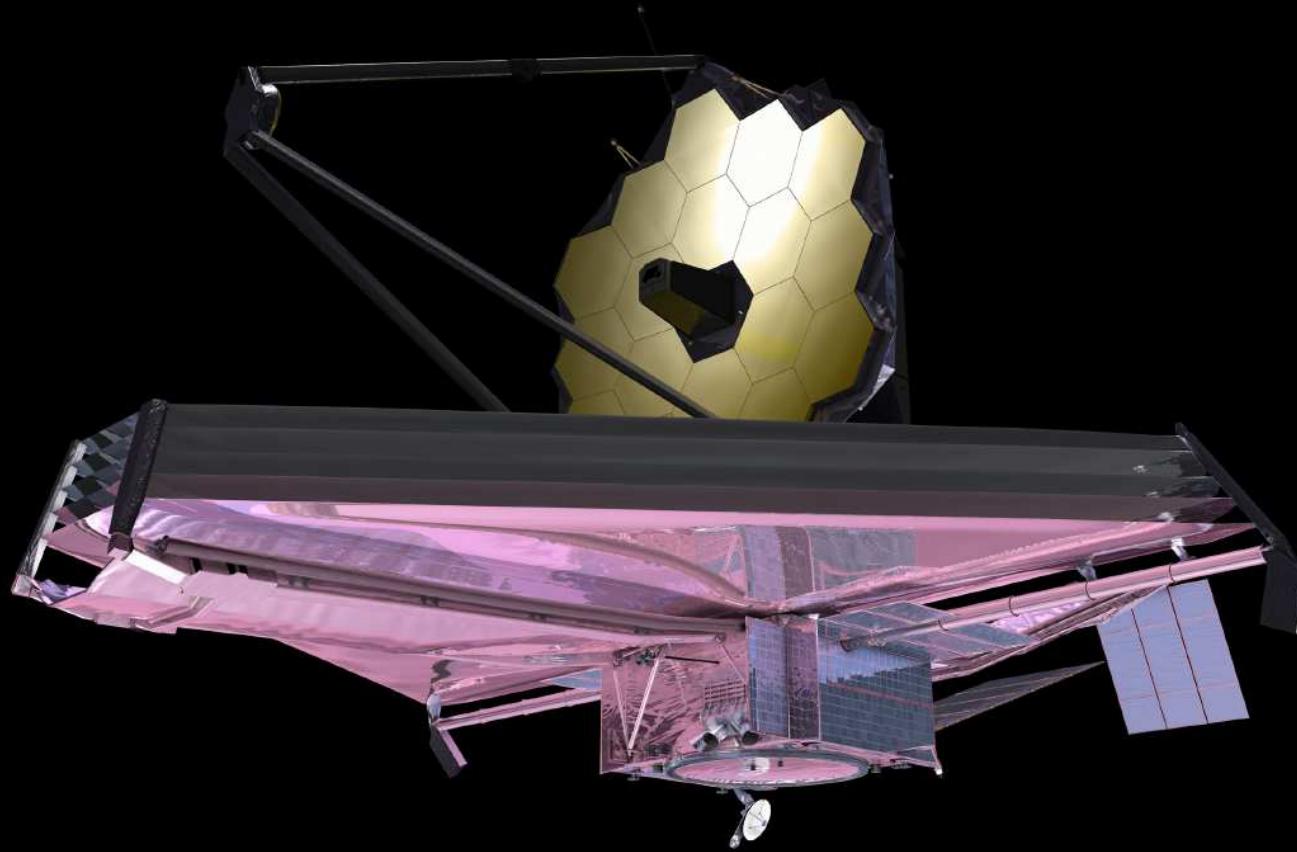
(1) Update of the James Webb Space Telescope (JWST), 2015



To be used by students & scientists after 2018 ... It'll be worth it.

(RIGHT) Life-size JWST prototype on the Capitol Mall, May 2007.

(1) Update of the James Webb Space Telescope as of 2015



- A fully deployable 6.5 meter (25 m^2) segmented IR telescope for imaging and spectroscopy at $0.6\text{--}28 \mu\text{m}$ wavelength, to be launched in Fall 2018.
- Nested array of sun-shields to keep its ambient temperature at 40 K, allowing faint imaging (31.5 mag \sim 1 FF from Moon), and spectroscopy.

THE JAMES WEBB SPACE TELESCOPE

JWST LAUNCH

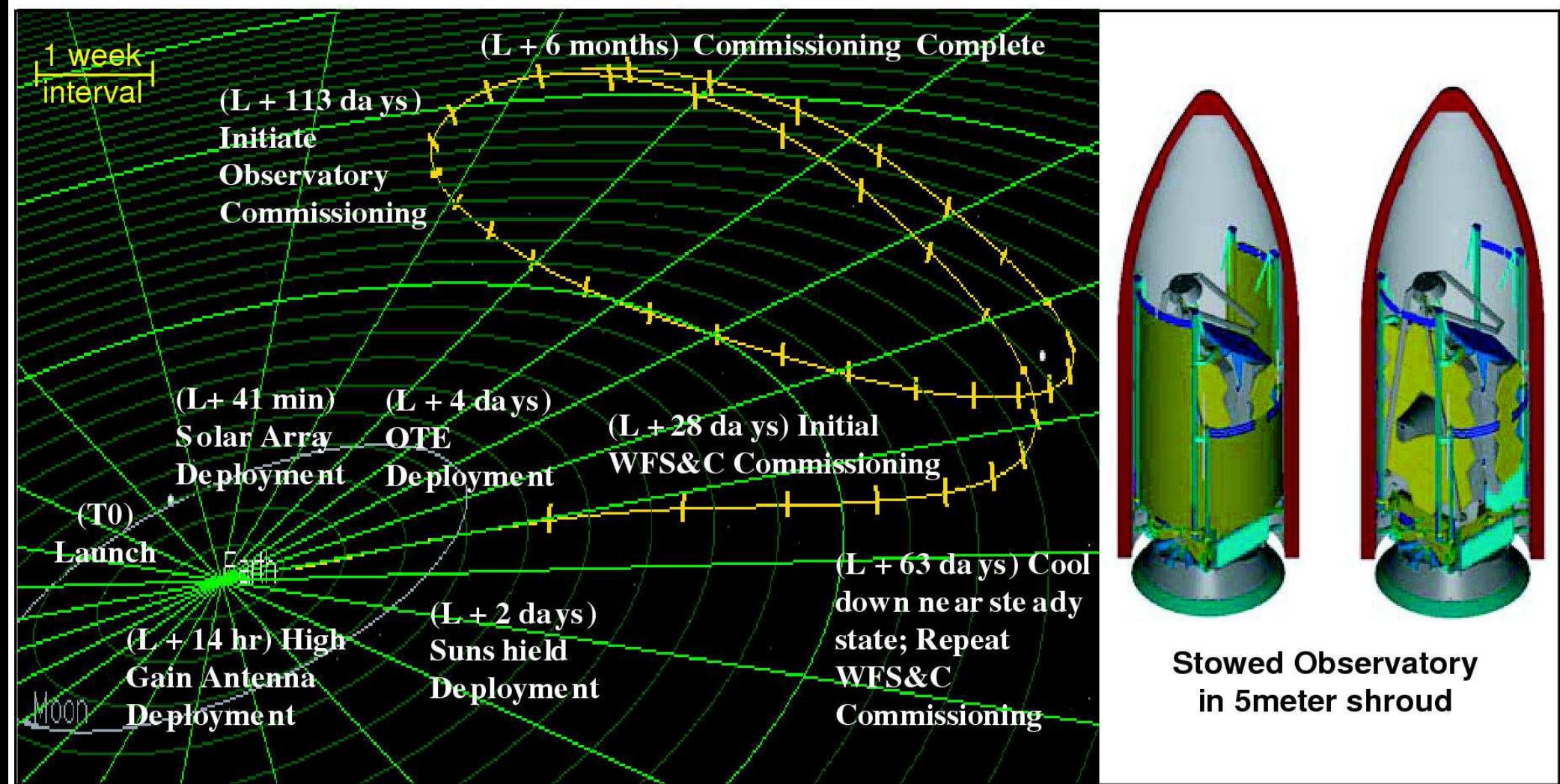
- LAUNCH VEHICLE IS AN ARIANE 5 ROCKET, SUPPLIED BY ESA
- SITE WILL BE THE ARIANESPACE'S ELA-3 LAUNCH COMPLEX NEAR KOUROU, FRENCH GUIANA



ARIANESPACE - ESA - NASA

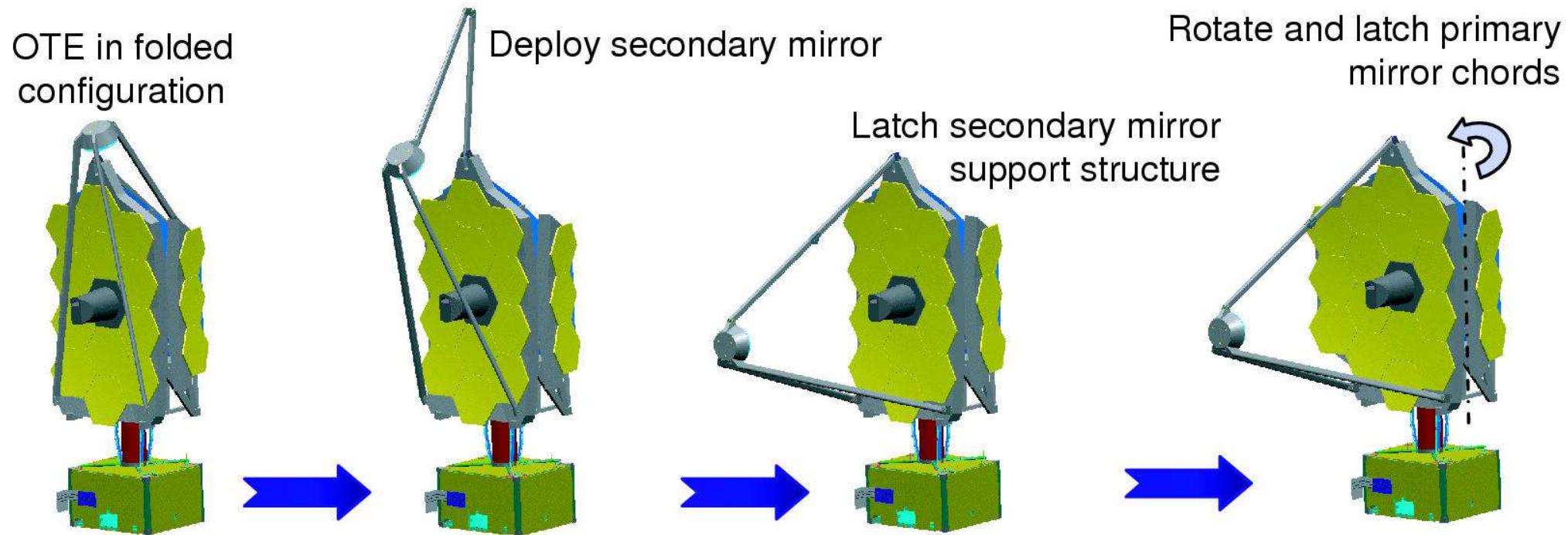
- The JWST launch weight will be $\lesssim 6500$ kg, and it will be launched to L2 with an ESA Ariane-V launch vehicle from Kourou in French Guiana.

(1a) How will JWST travel to its L2 orbit?



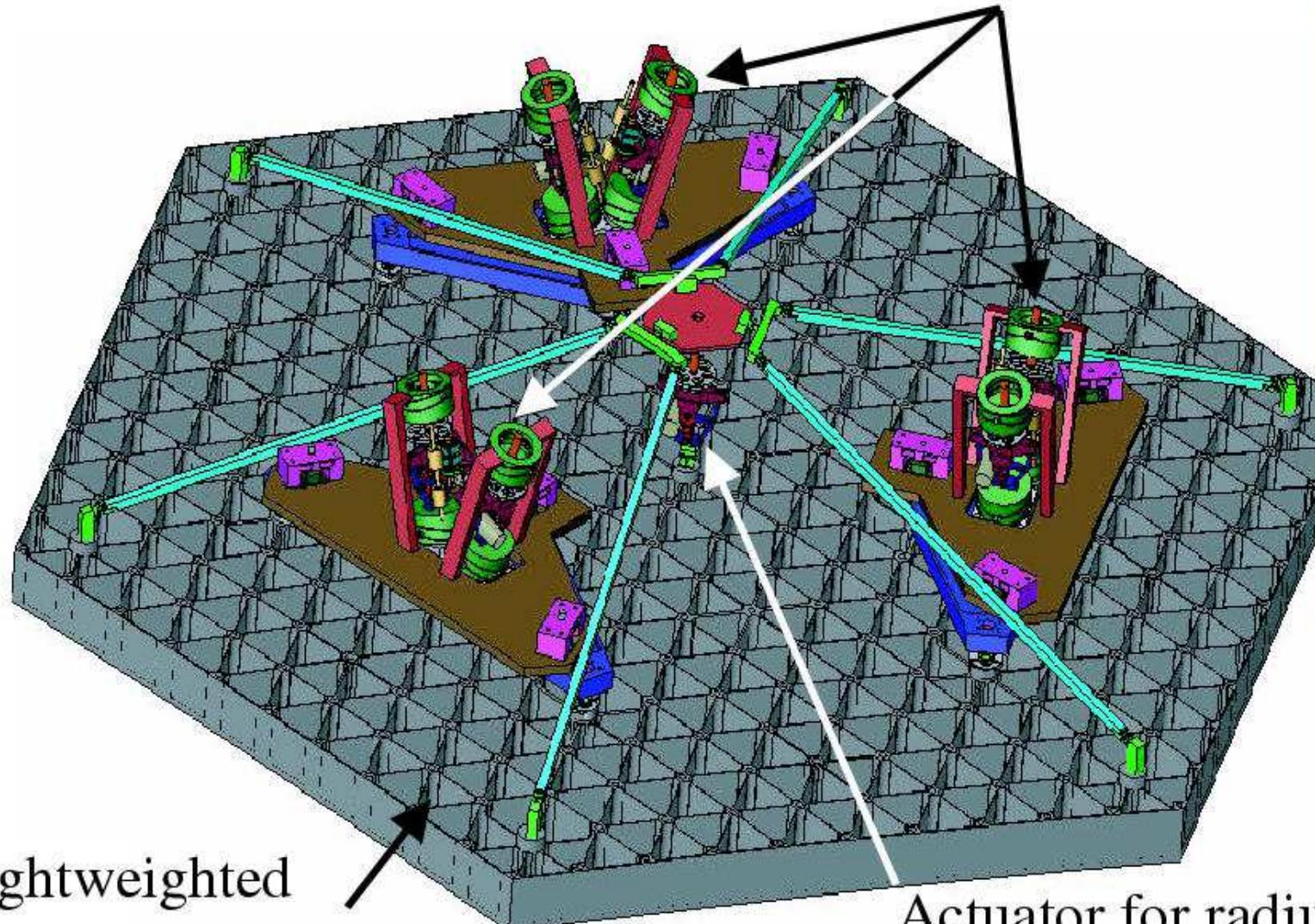
- After launch in (Oct.) 2018 with an ESA Ariane-V, JWST will orbit around the Earth–Sun Lagrange point L2, 1.5 million km from Earth.
- JWST can cover the whole sky in segments that move along with the Earth, observe $\gtrsim 70\%$ of the time, and send data back to Earth every day.

- (1b) How will JWST be automatically deployed?



- During its two month journey to L2, JWST will be automatically deployed, its instruments will be cooled, and be inserted into an L2 orbit.
- The entire JWST deployment sequence is being tested several times on the ground — but only in 1-G: component and system tests in 2014–2016 at GSFC (MD), Northrop (CA), and JSC (Houston).
- Component fabrication, testing, & system integration is on schedule: 18 out of 18 flight mirrors completely done, and meet the 40K specifications.

Actuators for 6 degrees of freedom rigid body motion



Actuator
development
unit

Lightweighted
Beryllium Mirror

Actuator for radius
of curvature adjustment

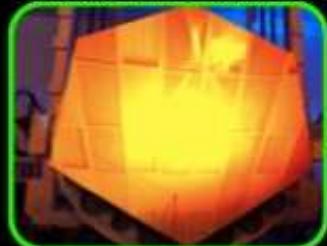
Active mirror segment support through “hexapods”, similar to Keck.

Redundant & doubly-redundant mechanisms, quite forgiving against failures.



JWST Hardware Status

Primary Mirror Segment



Aft Optics System



PM Flight Backplane



Tertiary Mirror

Secondary Mirror Pathfinder Strut



Secondary Mirror Hexapod



Secondary Mirror



Membrane Mgmt

Fine Steering Mirror



ISIM Flight Bench



Pathfinder Membrane



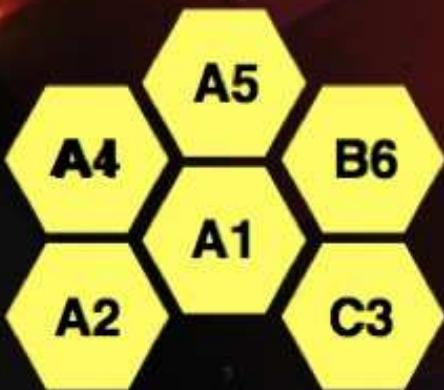
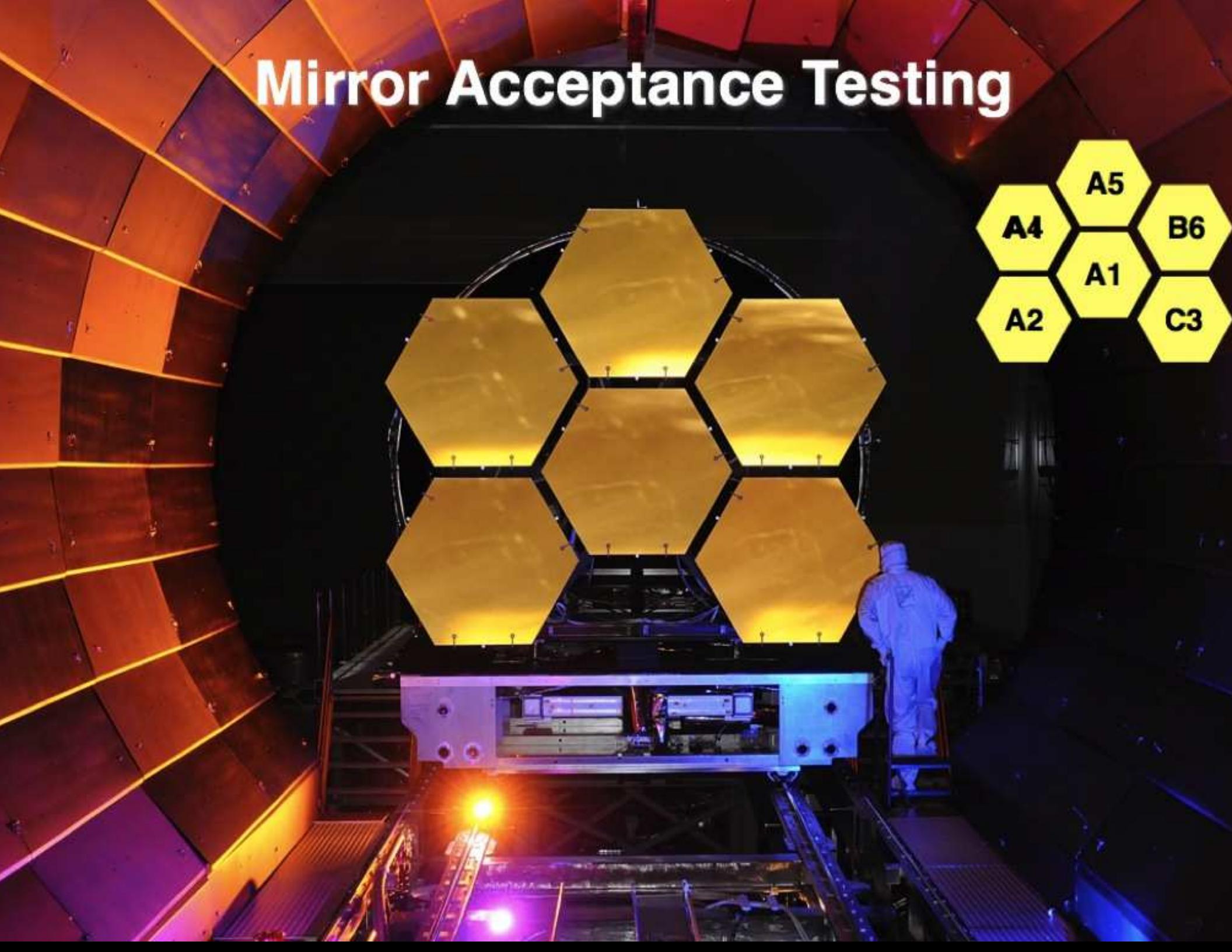
Spacecraft computer Test Unit



Mid-boom Test

Summer 2015: $\gtrsim 98\%$ of launch mass³ designed and built ($\gtrsim 65\%$ weighed).

Mirror Acceptance Testing



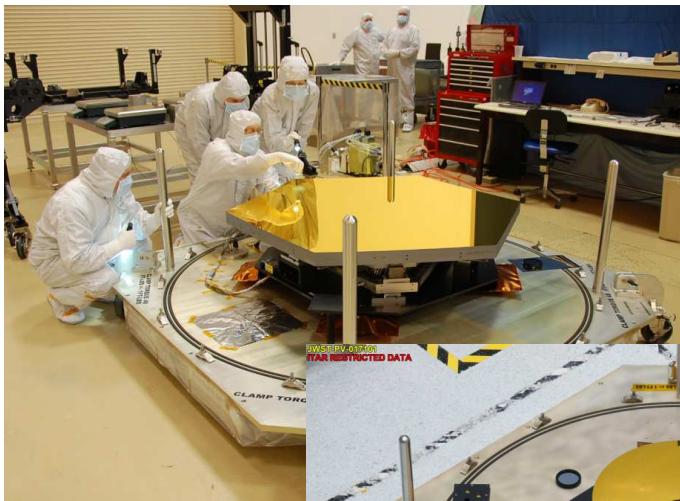




Mirror Status

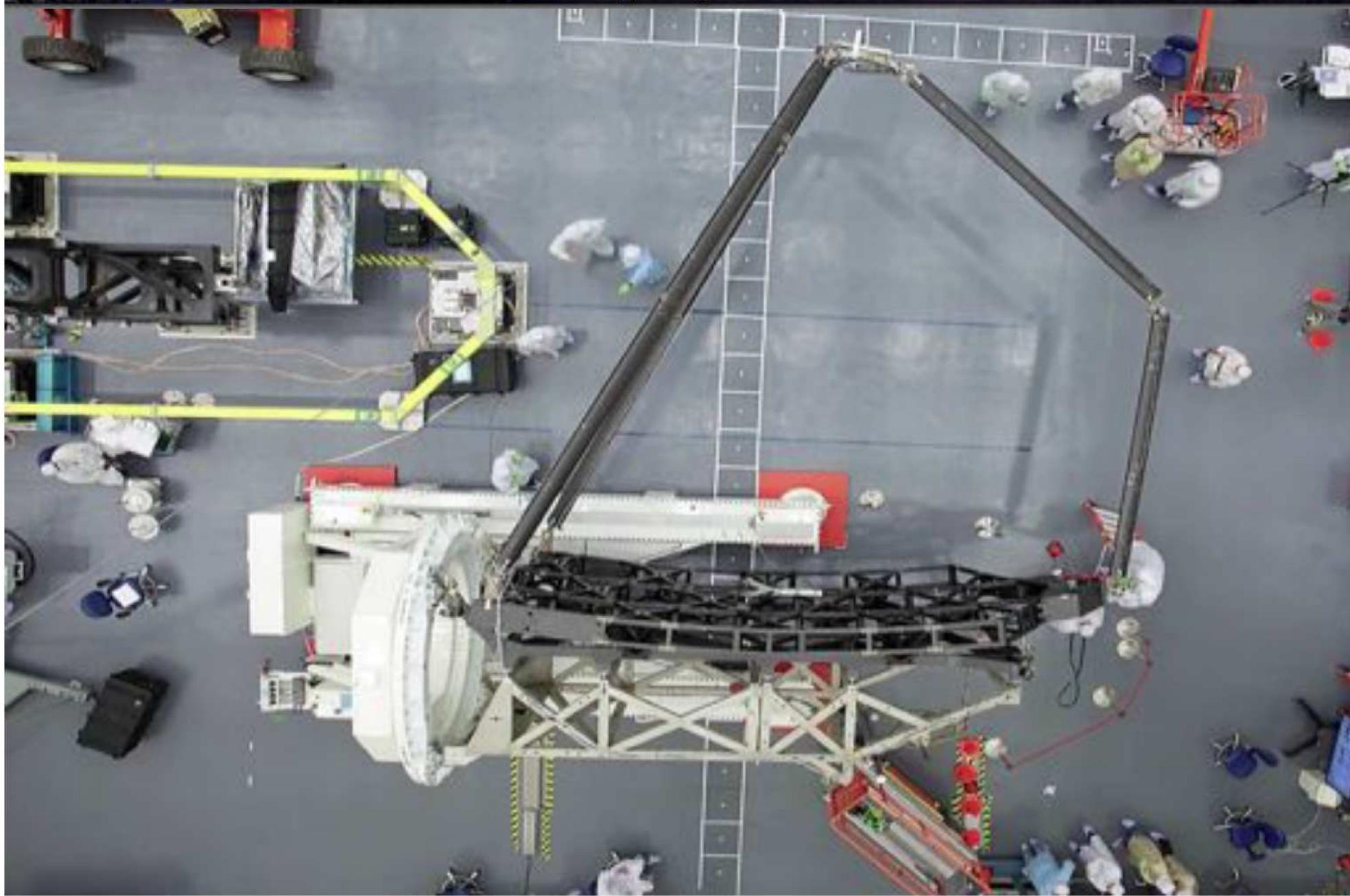


- **15 flight primary mirrors and the flight secondary mirror are at GSFC in storage**
 - All spares were at GSFC in storage (SM spares, 3 PMSA spares)
 - 2 EDU mirrors sent back to Ball for gear motor rework
 - All flight gear motor refurbishment is complete
 - All flight mirrors will be at GSFC by end of year, needed in 2015



Spring 2014: All 18 flight mirrors delivered to NASA GSFC (MD).

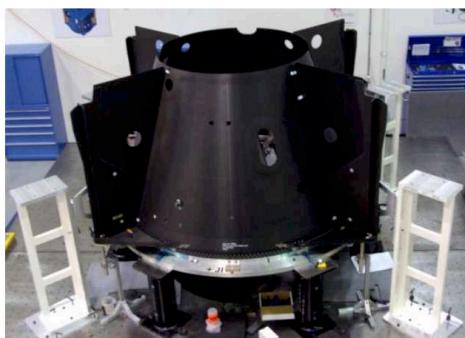
Pathfinder: Powered Deployment of SMSS



July 2014: Secondary Mirror Support deployment successfully tested.



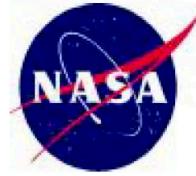
JWST Hardware Progress



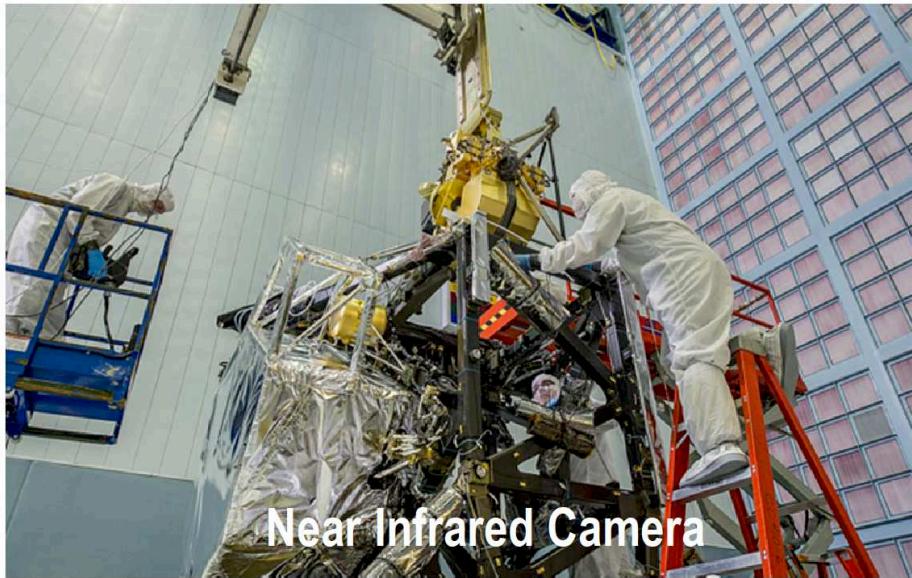
JWST remains on track for an October 2018 launch within its replan budget guidelines

29

July 2014: ● Secondary Mirror Support deployment successfully tested.
2015: ● Engineering sunshield successfully deployed at Northrop (CA).



All Instruments Integrated



(1c) JWST instrument update: US (UofA, JPL), ESA, & CSA.

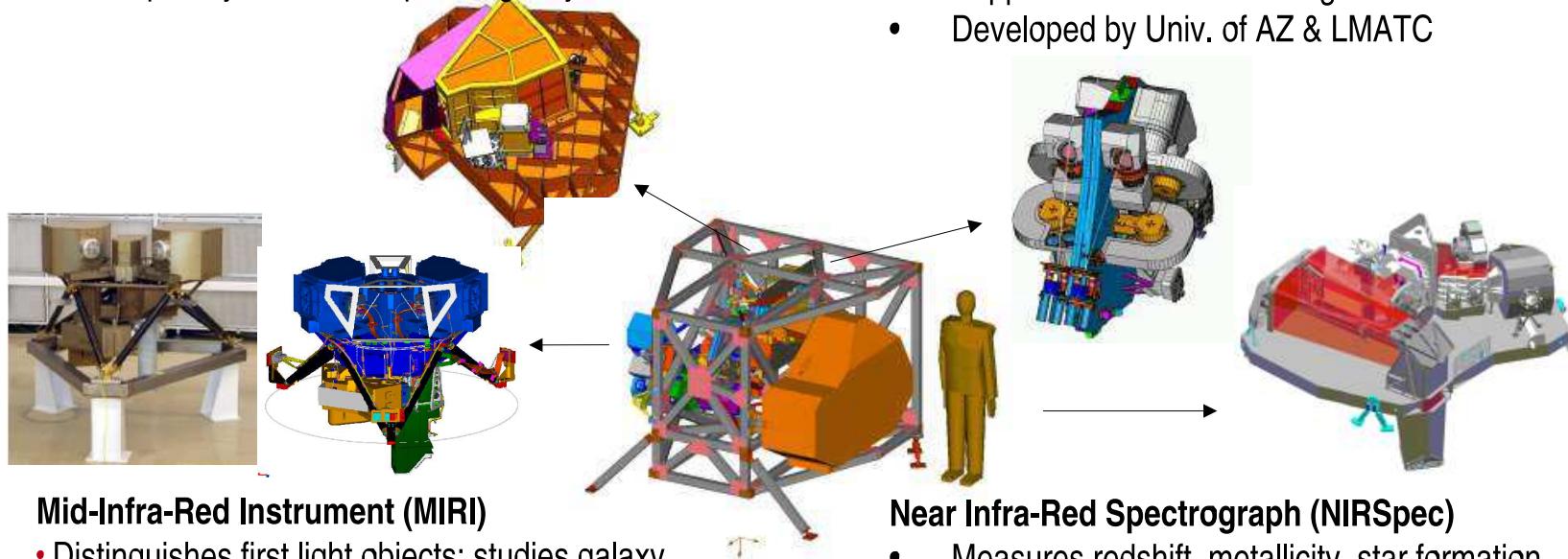


Instrument Overview



Fine Guidance Sensor (FGS)

- Ensures guide star availability with >95% probability at any point in the sky
- Includes Narrowband Imaging Tunable Filter
- Developed by Canadian Space Agency & COM DEV



Mid-Infra-Red Instrument (MIRI)

- Distinguishes first light objects; studies galaxy evolution; explores protostars & their environs
- Imaging and spectroscopy capability
- 5 to 27 microns
- Cooled to 7K by Cyro-cooler
- Combined European Consortium/JPL development

Near Infra-Red Camera (NIRCam)

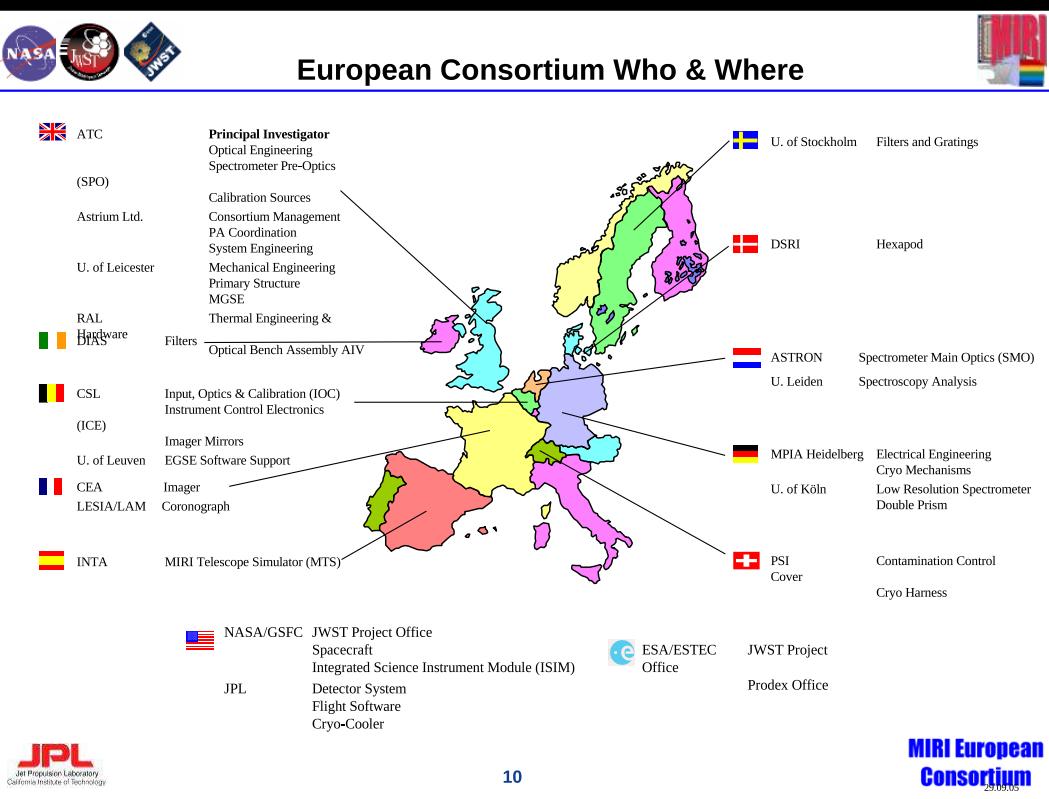
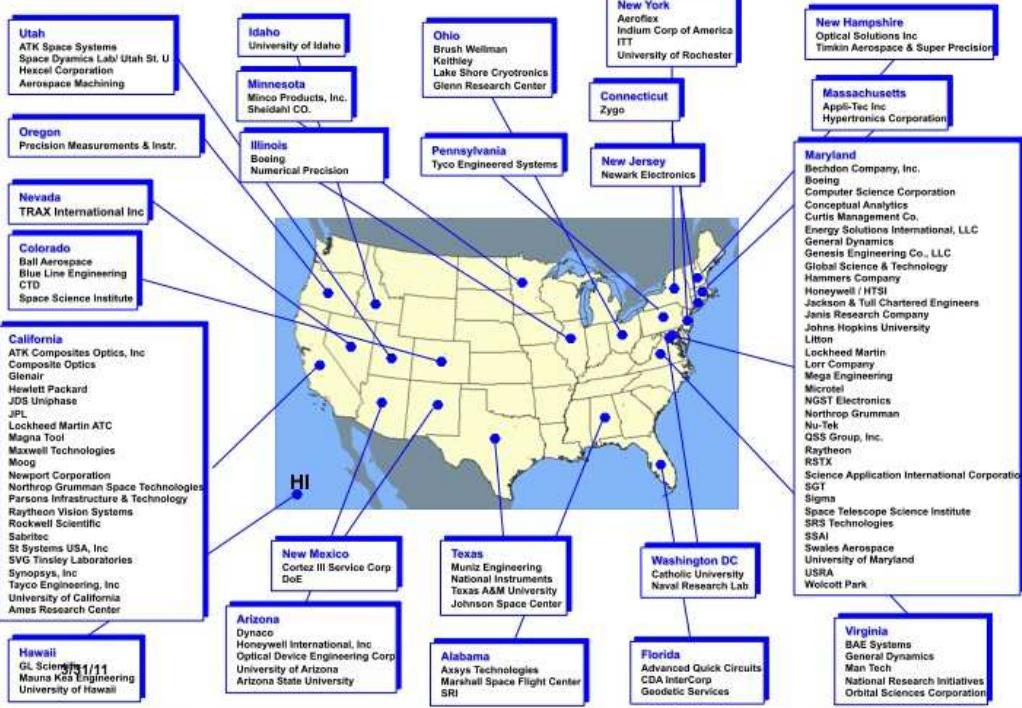
- Detects first light galaxies and observes galaxy assembly sequence
- 0.6 to 5 microns
- Supports Wavefront Sensing & Control
- Developed by Univ. of AZ & LMATC

Near Infra-Red Spectrograph (NIRSpec)

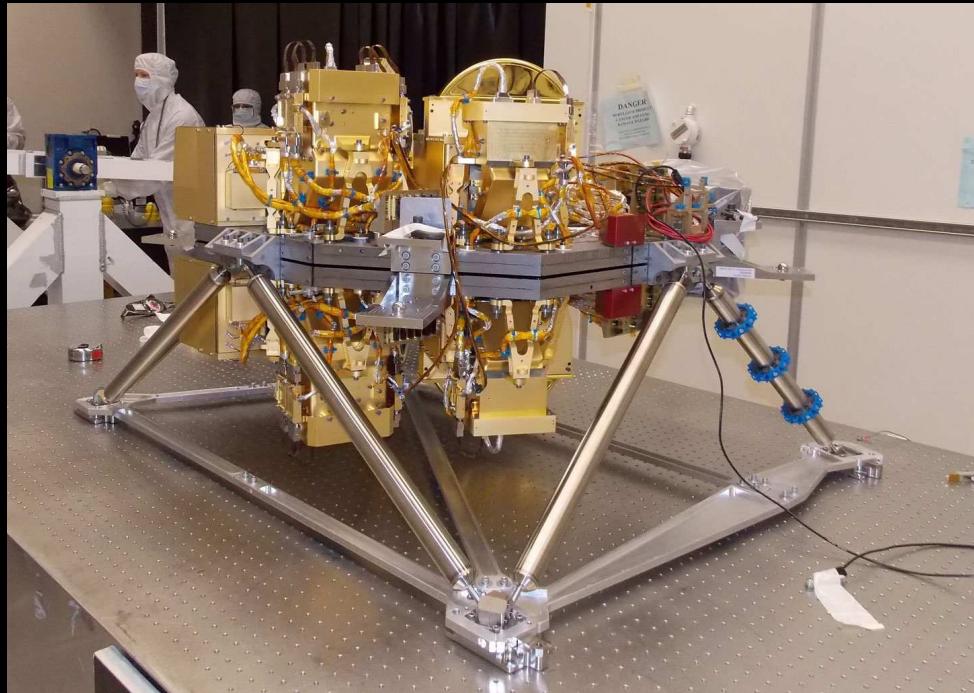
- Measures redshift, metallicity, star formation rate in first light galaxies
- 0.6 to 5 microns
- Simultaneous spectra of >100 objects
- Developed by ESA & EADS with NASA/GSFC Detector & Microshutter Subsystems

All delivered: MIRI 05/12; FGS 07/12; NIRCam 07/13, NIRSpec 9/13.

JWST: A Product of the Nation



- JWST hardware made in 27 US States: $\gtrsim 98\%$ of launch-mass finished.
- Ariane V Launch & NIRSpec provided by ESA; & MIRI by ESA & JPL.
- JWST Fine Guider Sensor + NIRISS provided by Canadian Space Agency.
- JWST NIRCam made by UofA and Lockheed.

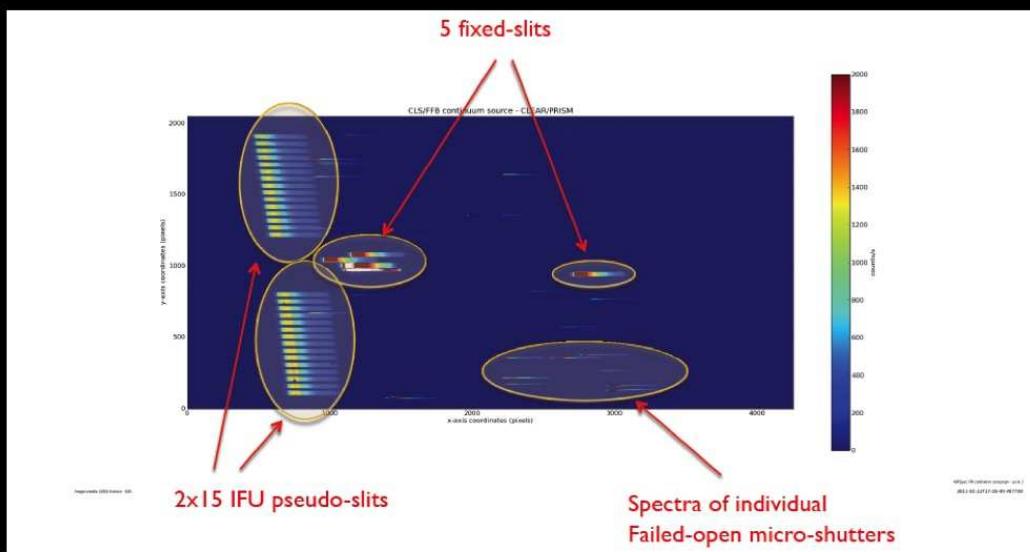


JWST's short-wavelength ($0.6\text{--}5.0\mu\text{m}$) imagers:

- NIRCam — built by UofA (AZ) and Lockheed (CA).
- Fine Guidance Sensor (& $1\text{--}5\mu\text{m}$ grisms) — built by CSA (Montreal).
- FGS includes very powerful low-res Near-IR grism spectrograph (NIRISS).
- FGS delivered to GSFC 07/12; NIRCam delivered 07/13.
- Detectors replaced in 2015 between CryoVacuum tests CV2 and CV3.



Flight NIRSpec First Light



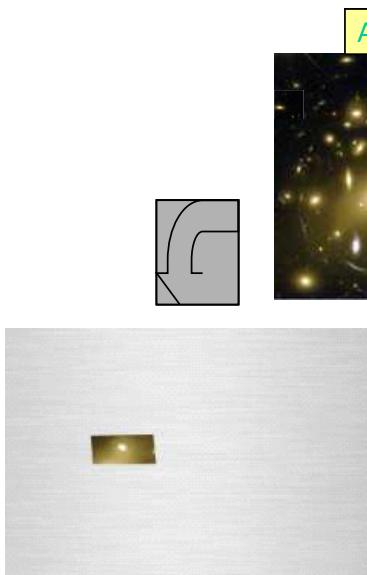
JWST's short-wavelength ($0.6\text{--}5.0\mu\text{m}$) spectrograph:

- NIRSpec — built by ESA/ESTEC and Astrium (Munich).
- Flight build completed and tested with First Light in Spring 2011.

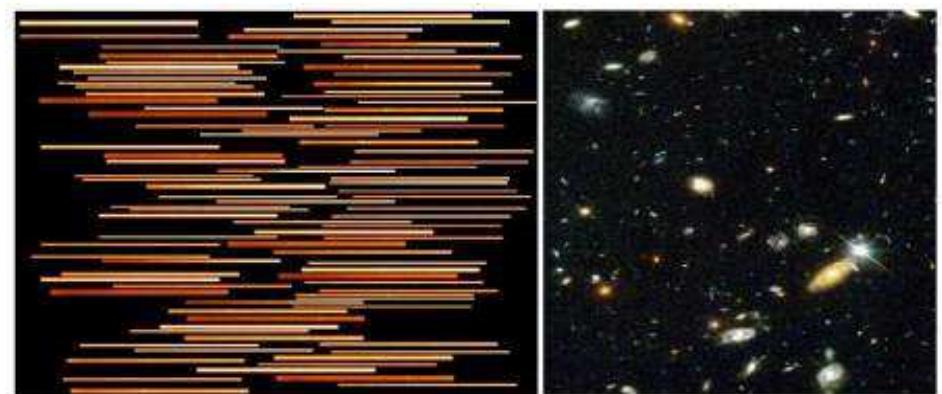
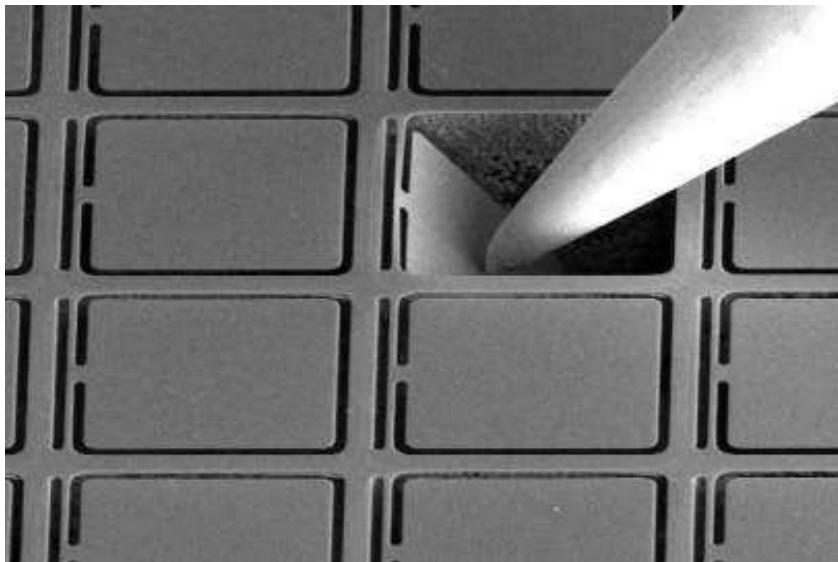
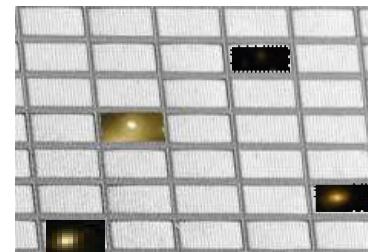
NIRSpec delivered to NASA/GSFC in 09/13.

- Detectors replaced in 2015 between CryoVacuum tests CV2 and CV3.

Micro Shutters

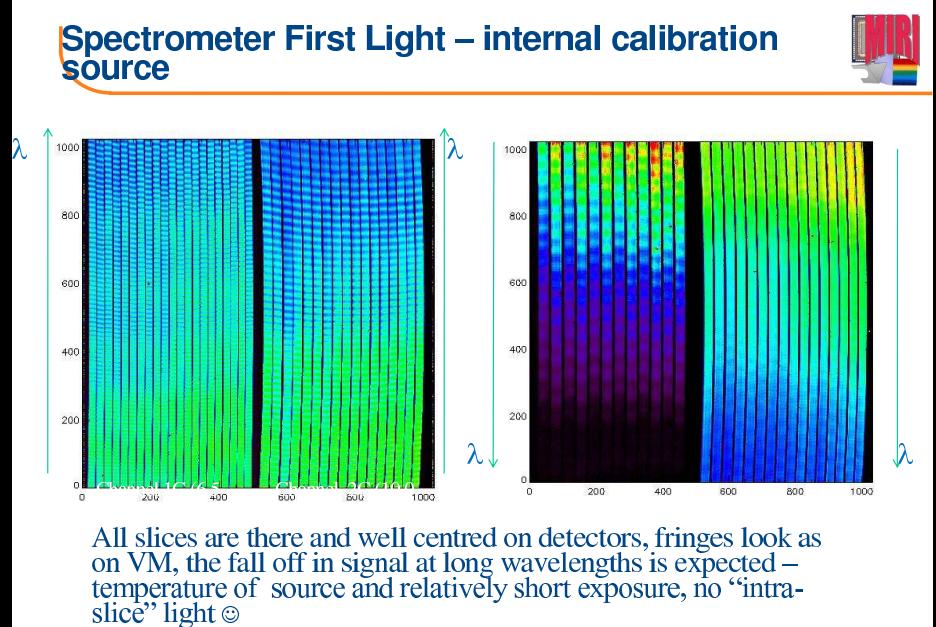


Metal Mask/Fixed Slit





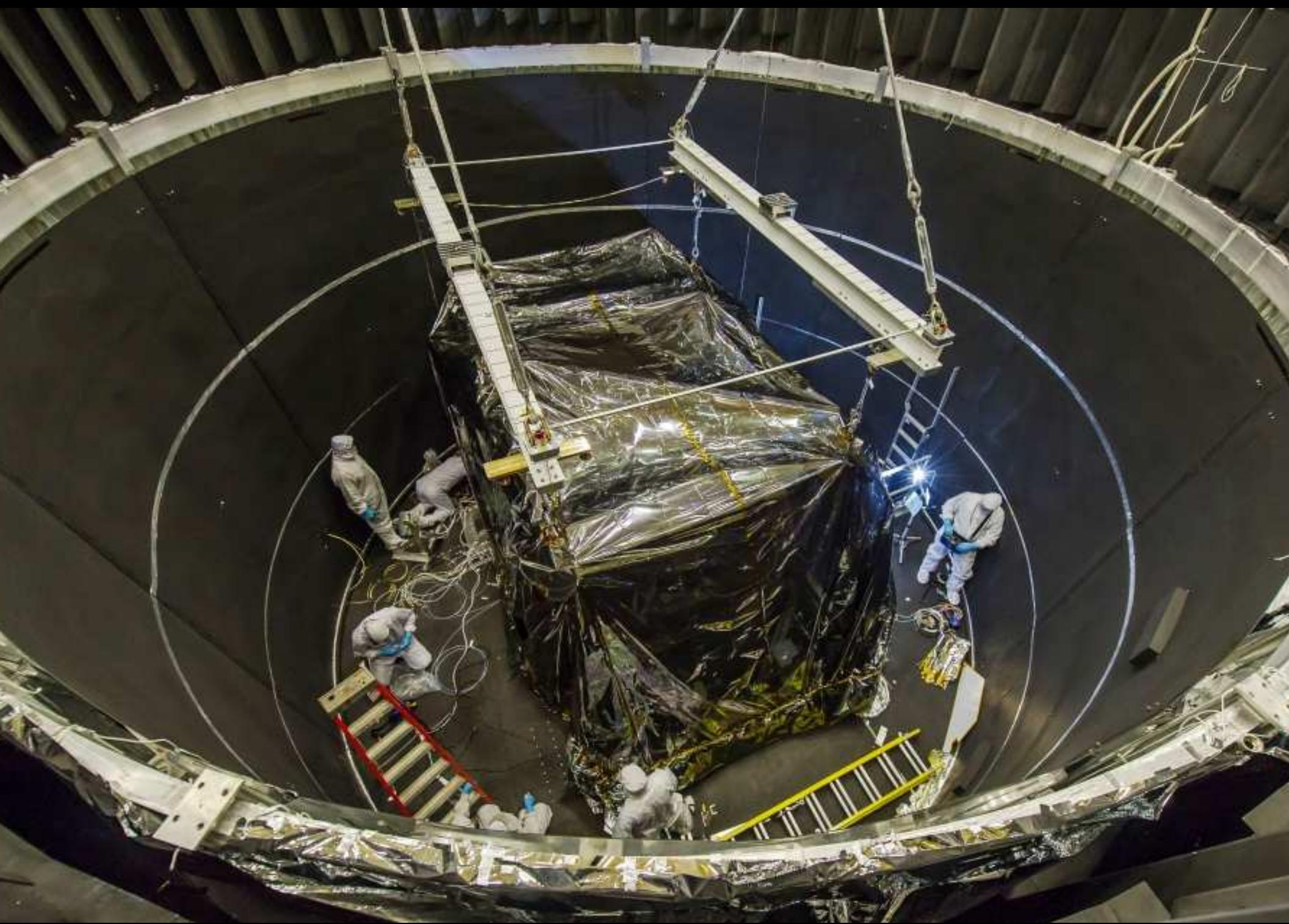
Flight MIRI



JWST's mid-infrared ($5\text{--}29\mu\text{m}$) camera and spectrograph:

- MIRI — built by ESA consortium of 10 ESA countries & NASA JPL.
- Flight build completed and tested with First Light in July 2011.

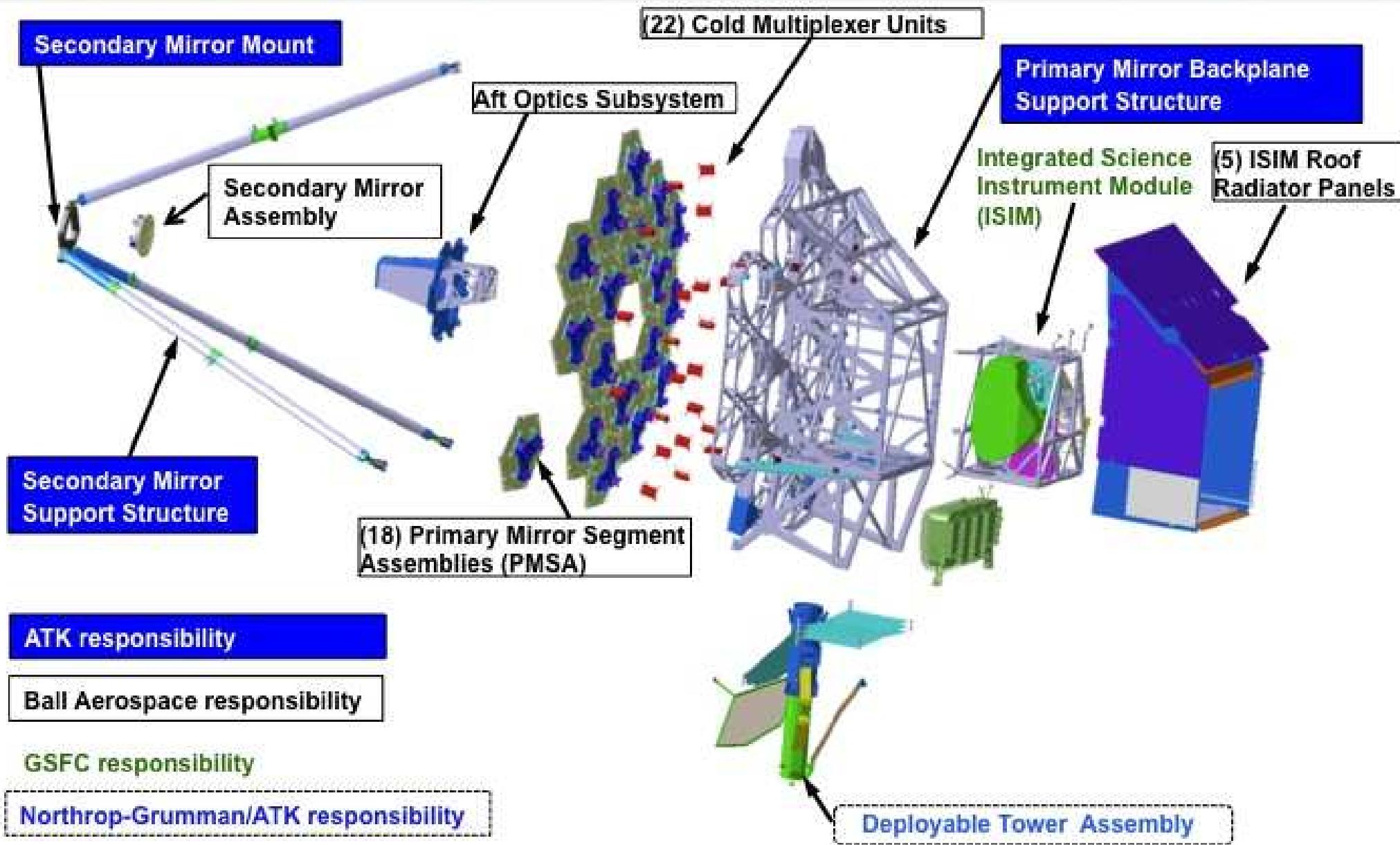
MIRI delivered to NASA/GSFC in May 2012.



June 2014: Flight ISIM (with all 4 instruments) in OSIM; Oct. 2015: CryoVac3.



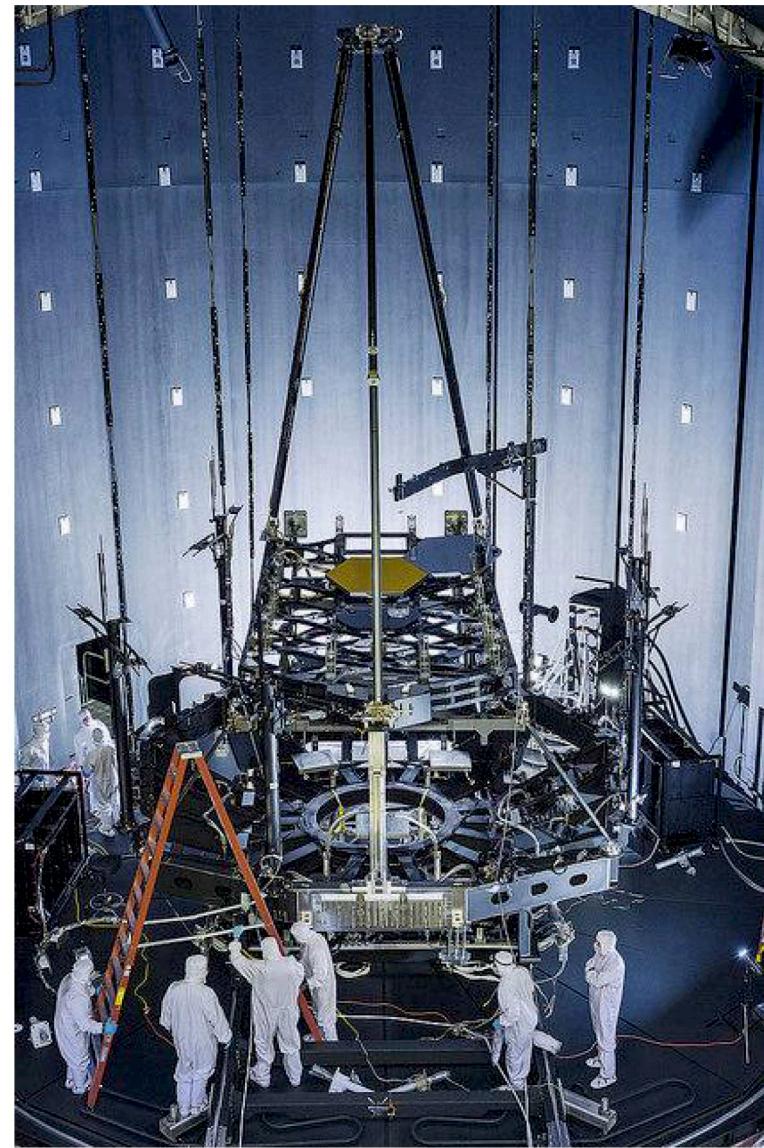
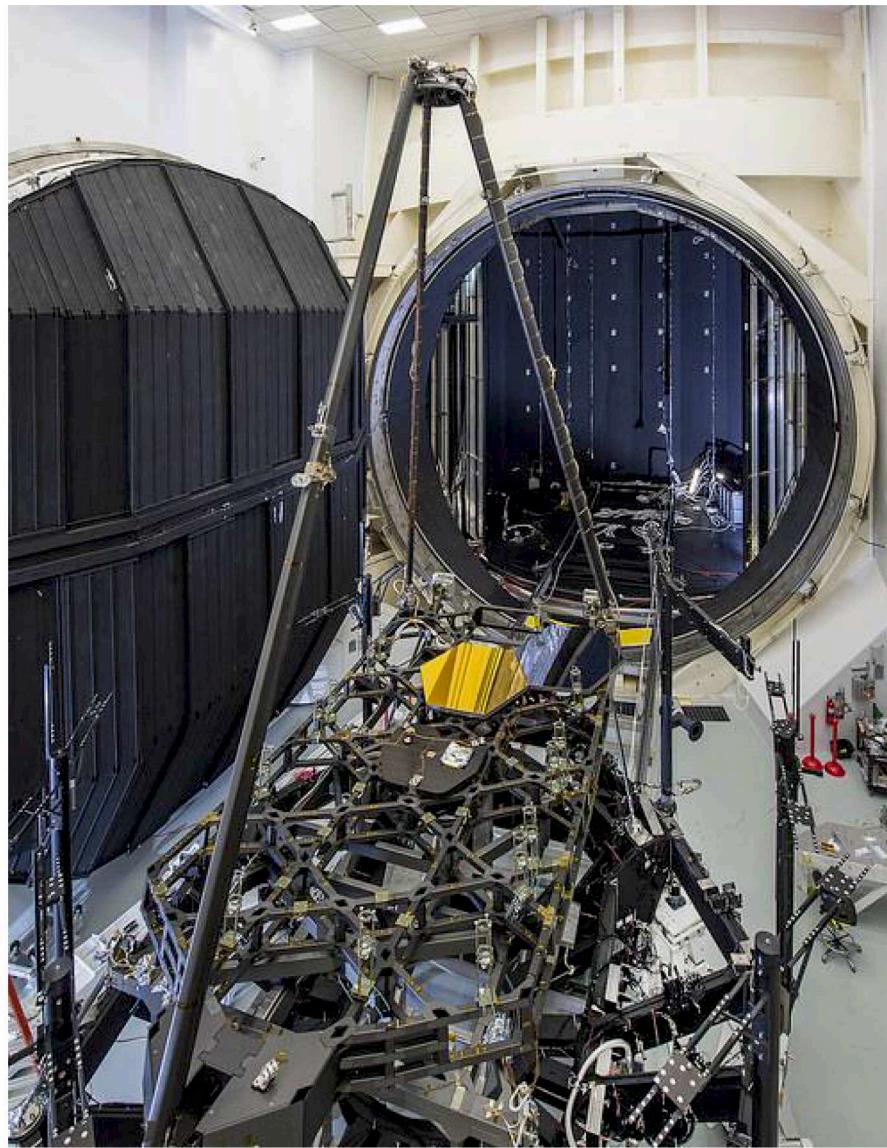
TELESCOPE ARCHITECTURE



3/31/11

2014–2016: Complete system integration at GSFC and Northrop.

Pathfinder & JSC Chamber A: getting ready for OGSE1 (and eventually OGSE2 & Thermal Pathfinder)



April 2015: Testing OTIS chamber with the JWST Engineering model.



OTIS Test GSE Architecture and Subsystems



Chamber Isolator Units

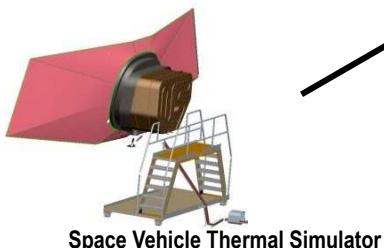
Dynamically isolates OTIS Optical Test
– Integration 6 units complete



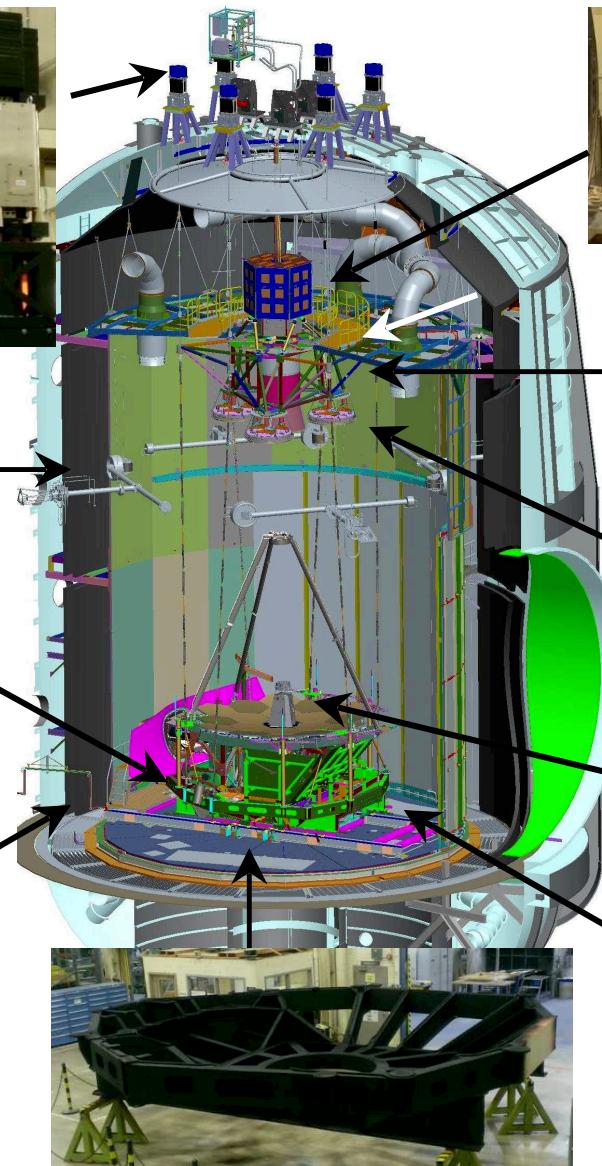
Cryo Position Metrology (CPM)
Photogrammetry System
Integration Complete



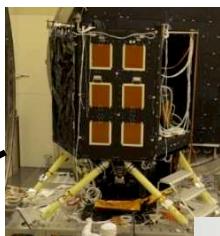
ADM - new Leica
delivered and under
test



Passed design review and started
Procurements and fab subcontracts



HOSS – OTIS support structure
HOSS – will be in the chamber for Bake out in June



Center of Curvature Optical Assembly (COCOA)

- Multiwavelength interferometer (MWIF), null, calibration equipment, coarse/fine PM phasing tools, Displacement Measuring Interferometer – COCOA was exercised at MSFC in December



USF Structural Frame – supports Metrology
ready for chamber integration and Cryo Load tests



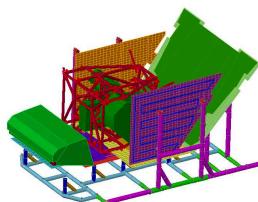
3 Auto collimating Flat Mirrors (ACFs)
1.5 M Plano for Pass and Half Testing
Cryo testing underway, ACF 1 complete, ACF 4 in
Cryo test complete , ACF 5 ready for Cryo.



AOS Source Plate
Sources for Pass and Half Test
72 optical fiber support cont.



Mag Damper Cryo
Test
Article
Fabrication started



Deep Space Edge Radiation Sink (DSERS)

Thermal modeling of payload and DSERS
started

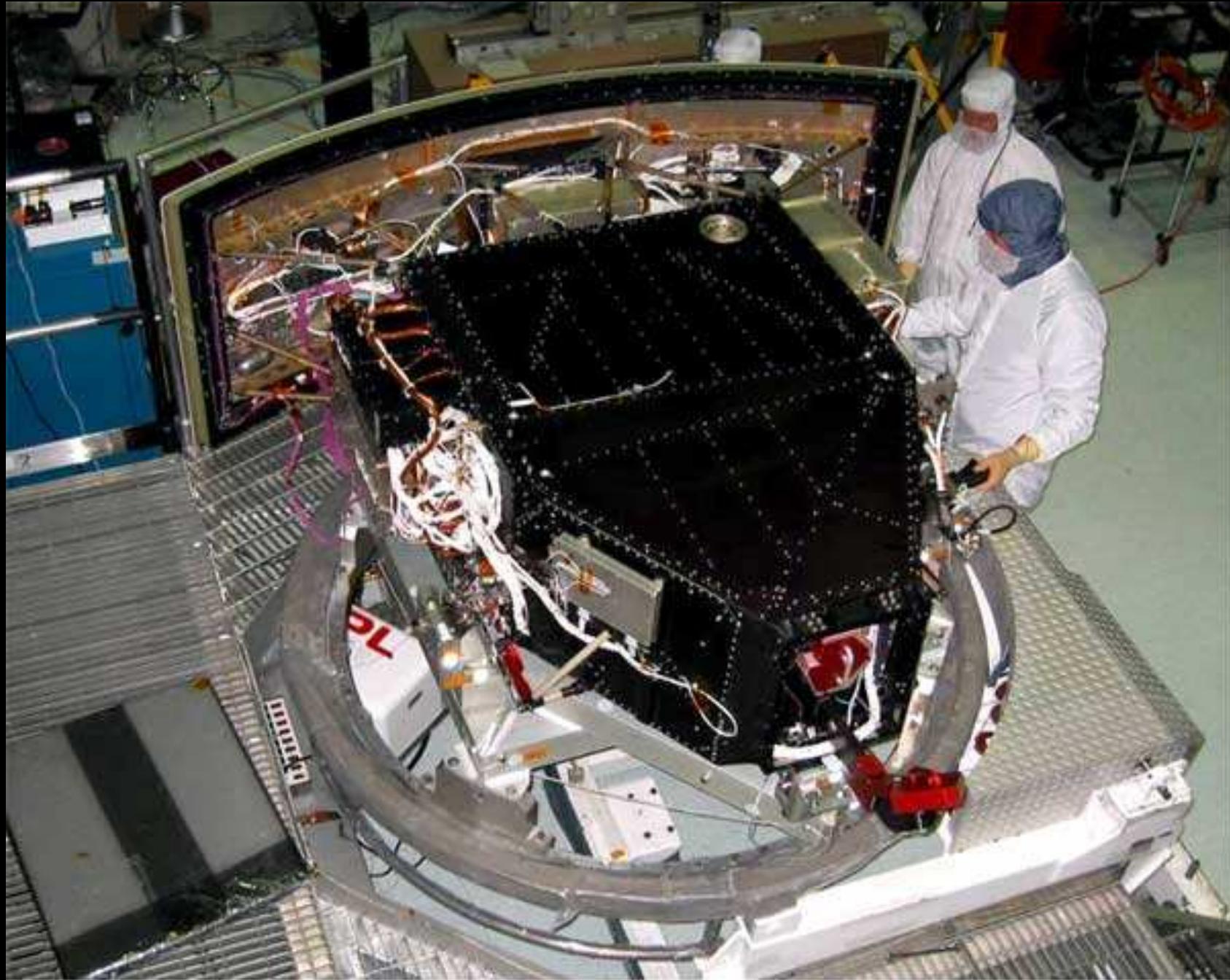
World's largest TV chamber OTIS: will test whole JWST in 2016–2017.

(2) HST WFC3: Measuring Galaxy Assembly and SMBH/AGN Growth?



10 filters with HST/WFC3 & ACS reaching AB=26.5–27.0 mag (10- σ) over 40 arcmin² at 0.07–0.15" FWHM from 0.2–1.7 μ m (UVUBVizYJH). JWST adds 0.05–0.2" FWHM imaging to AB \simeq 31.5 mag (1 nJy) at 1–5 μ m, and 0.2–1.2" FWHM at 5–29 μ m, tracing young+old SEDs & dust.

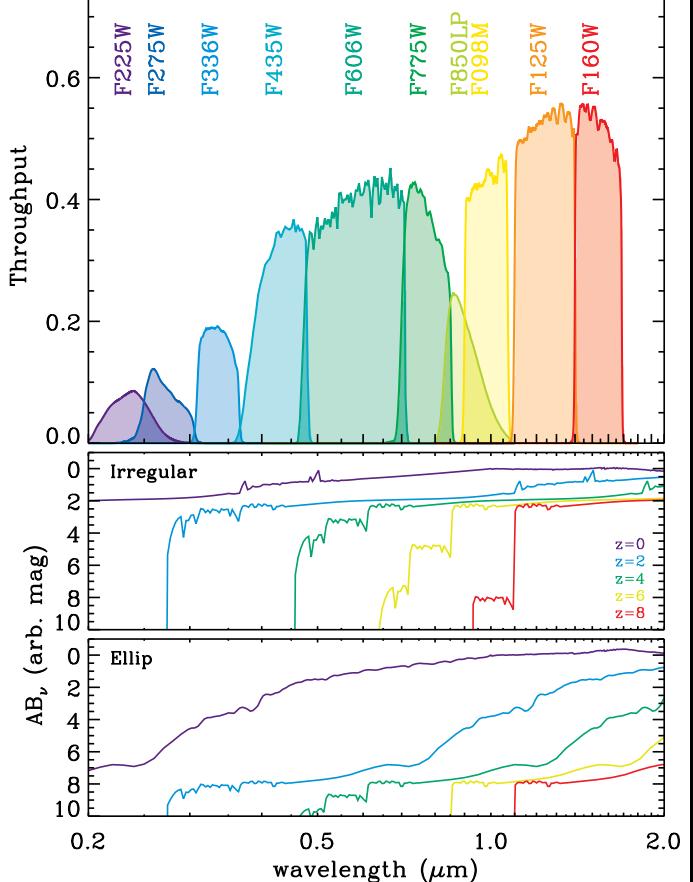
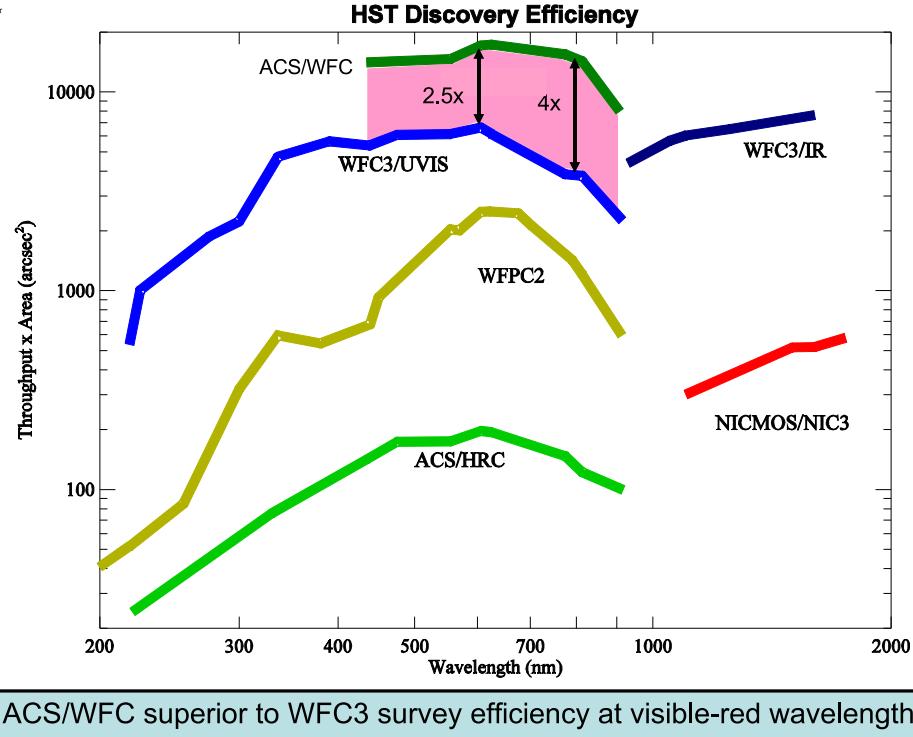
(2a) WFC3: Hubble's new Panchromatic High-Throughput Camera



HST WFC3 and its IR channel: a critical pathfinder for JWST science.



Role of ACS in HST Post-SM4 Imaging Capability



WFC3/UVIS channel unprecedented UV–blue throughput & areal coverage:

- QE $\gtrsim 70\%$, $4\text{k} \times 4\text{k}$ array of $0\farcs04$ pixel, FOV $\simeq 2\farcs67 \times 2\farcs67$.

WFC3/IR channel unprecedented near-IR throughput & areal coverage:

- QE $\gtrsim 70\%$, $1\text{k} \times 1\text{k}$ array of $0\farcs13$ pixel, FOV $\simeq 2\farcs25 \times 2\farcs25$.

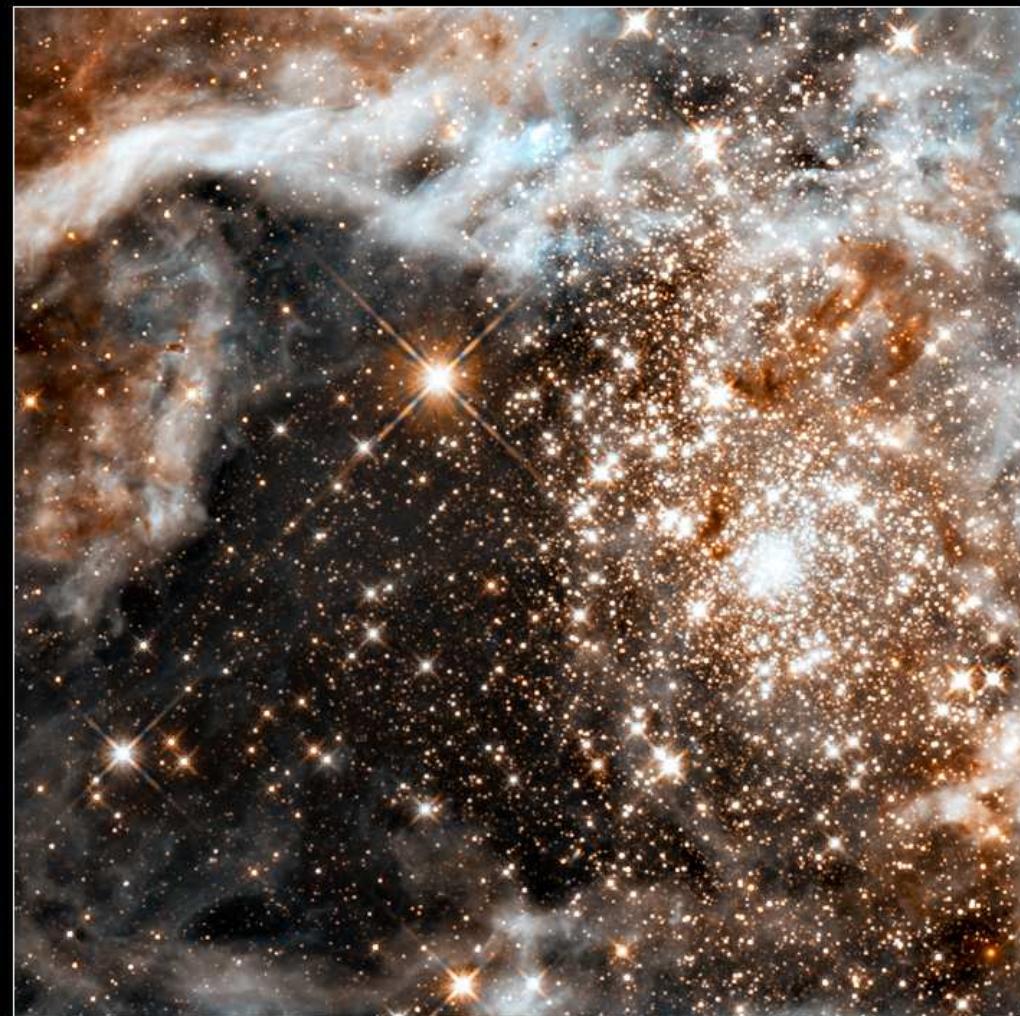
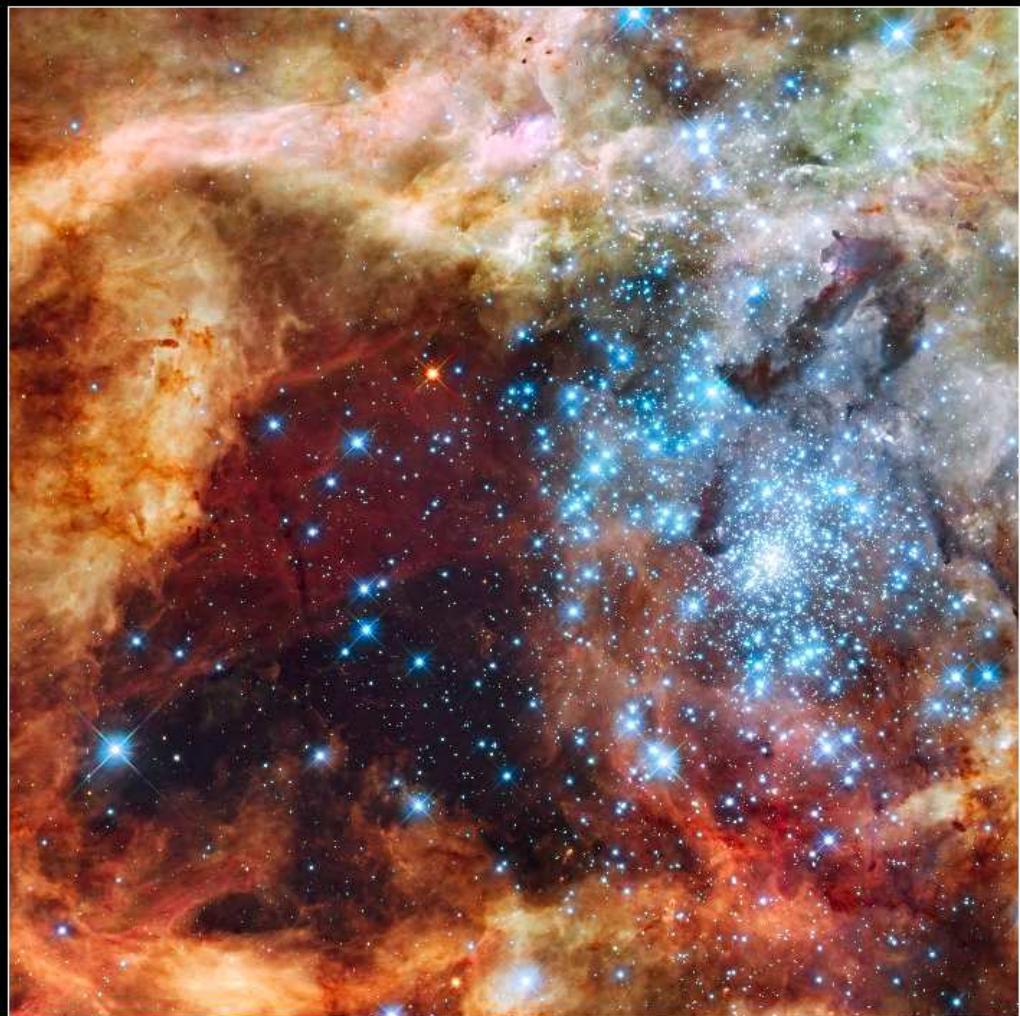
\Rightarrow WFC3 opened major new parameter space for astrophysics in 2009:

WFC3 filters designed for star-formation and galaxy assembly at $z \simeq 1\text{--}8$.

- HST WFC3 and its IR channel a critical pathfinder for JWST science.

Visible

Infrared



30 Doradus Nebula and Star Cluster

Hubble Space Telescope • WFC3/UVIS/IR

NASA, ESA, F. Paresce (INAF-IASF, Italy), and the WFC3 Science Oversight Committee

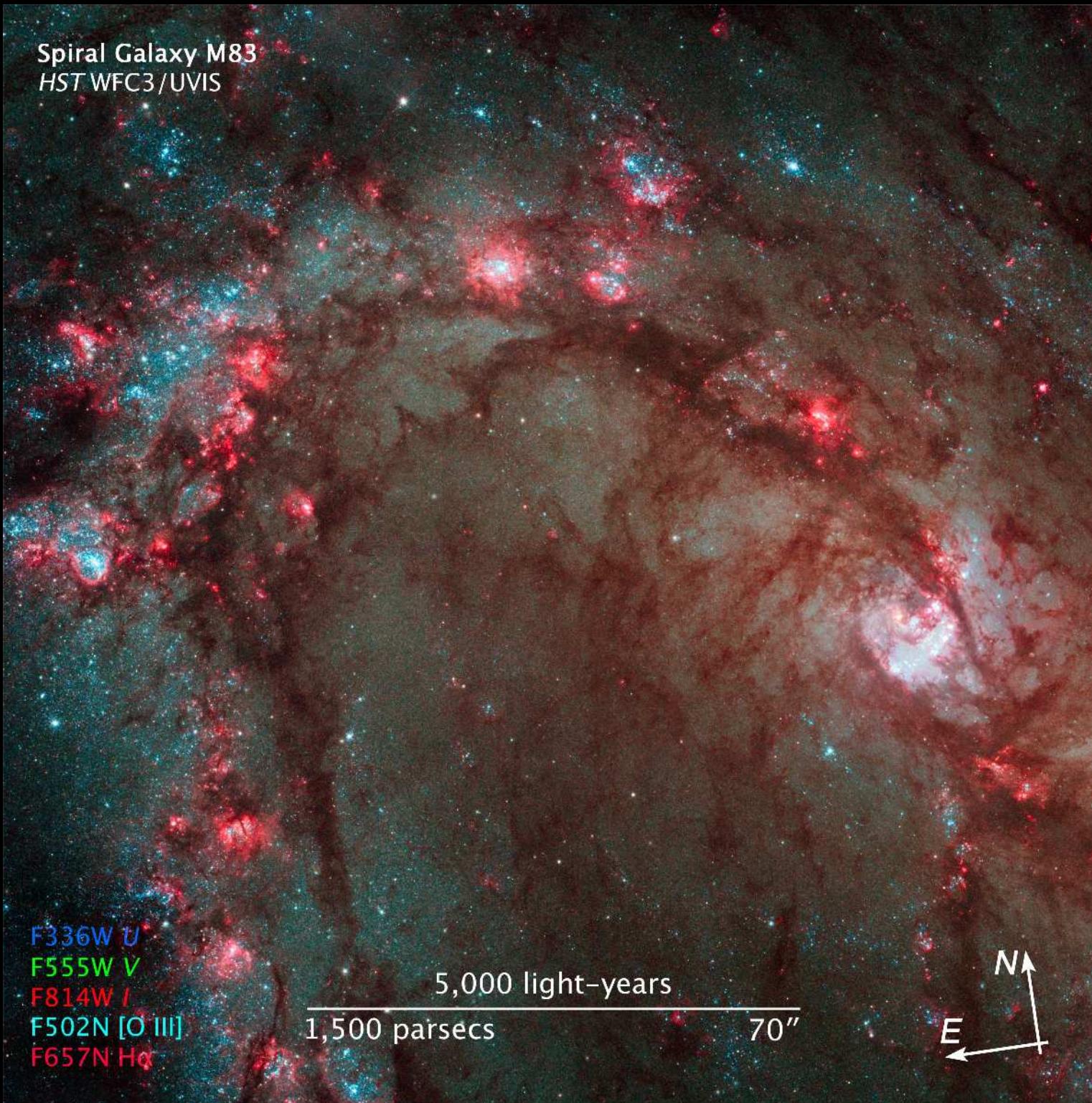
STScI-PRC09-32b

30 Doradus: Giant young star-cluster in Large Magellanic Cloud (150,000 ly), triggering birth of Sun-like stars (and surrounding debris disks).

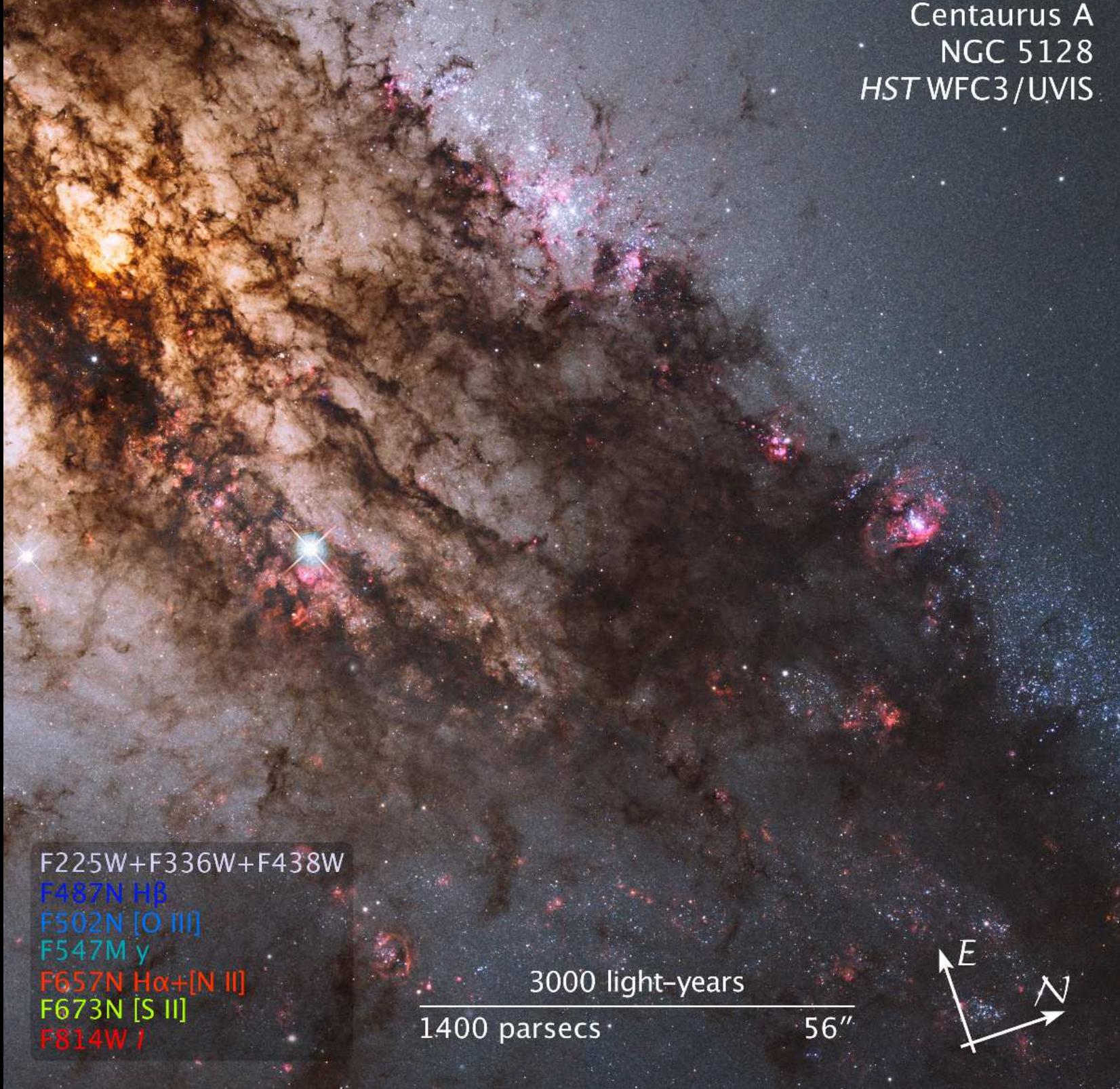


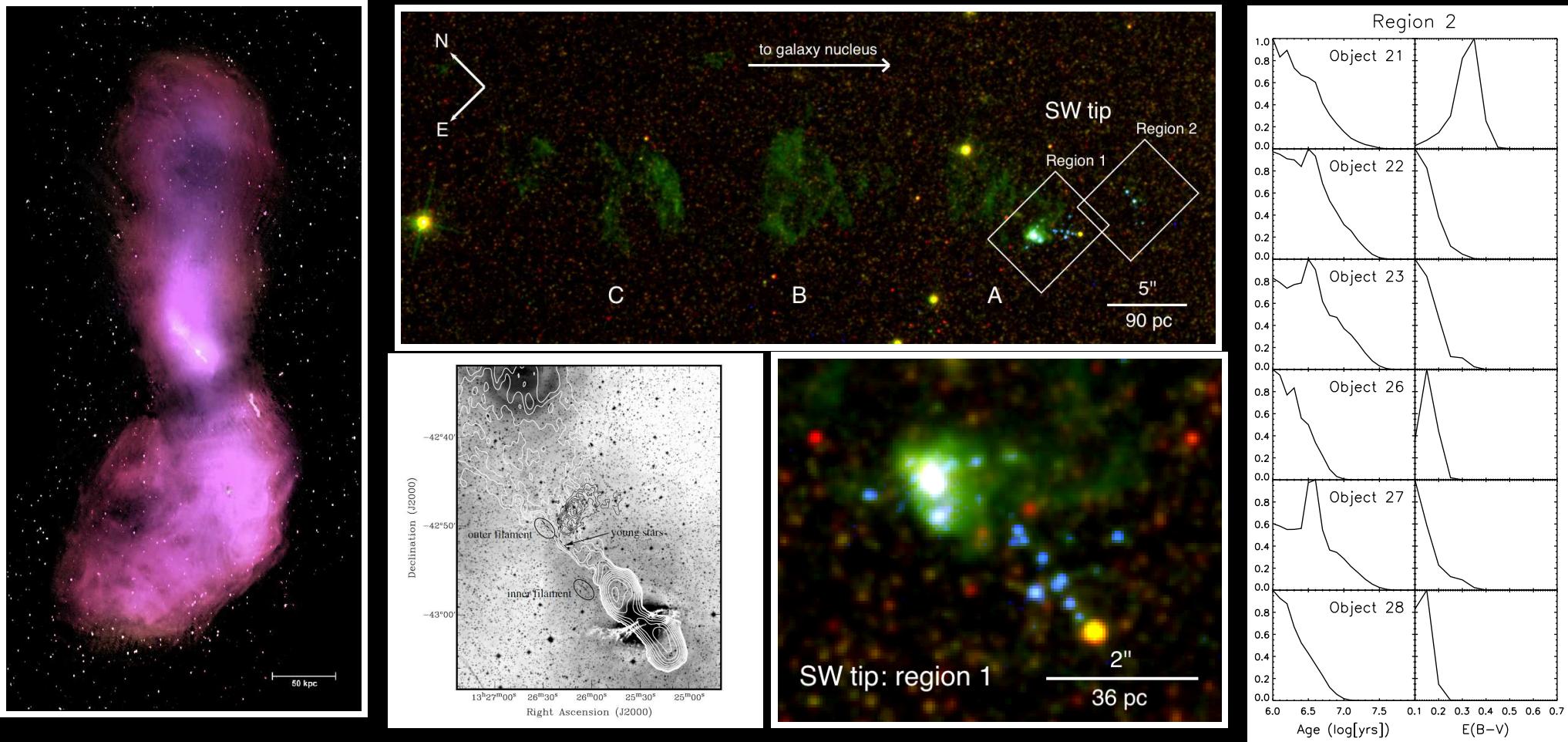


(2a) Measuring (Nearby) Galaxy Assembly and Supermassive Black-Hole Growth.



Centaurus A
NGC 5128
HST WFC3/UVIS

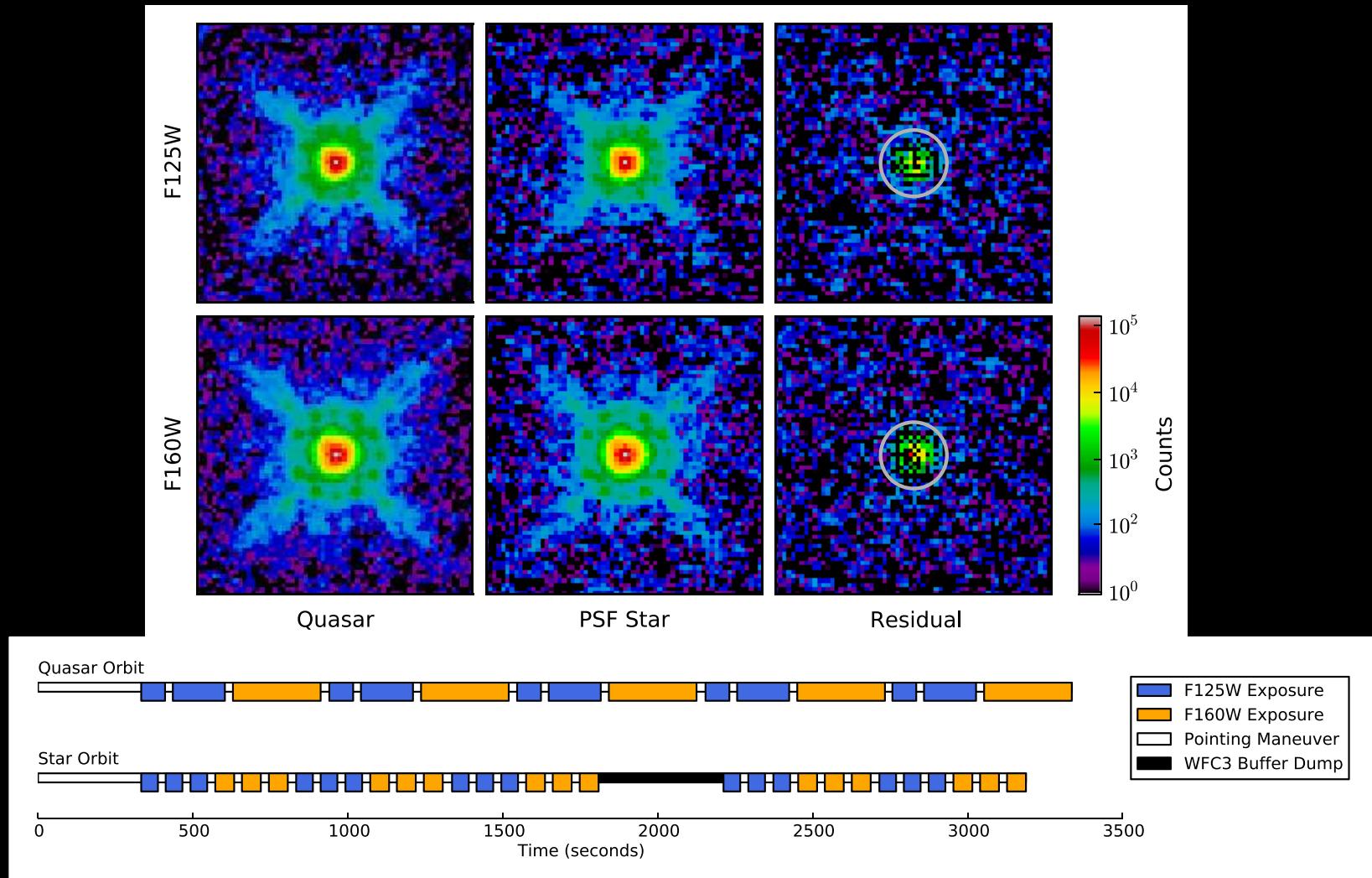




[Left] CSIRO/ATNF 1.4 GHz image of Centaurus A (Feain⁺ 2009). Fermi GeV source (Yang⁺ 12); & Auger UHE Cosmic Rays (Abreu⁺ 2010). [Middle] SF in Cent A jet's wake (Crockett⁺ 2012, MNRAS, 421, 1602). [Right] Well determined ages for young (~ 2 Myr) stars near Cen A's jet.

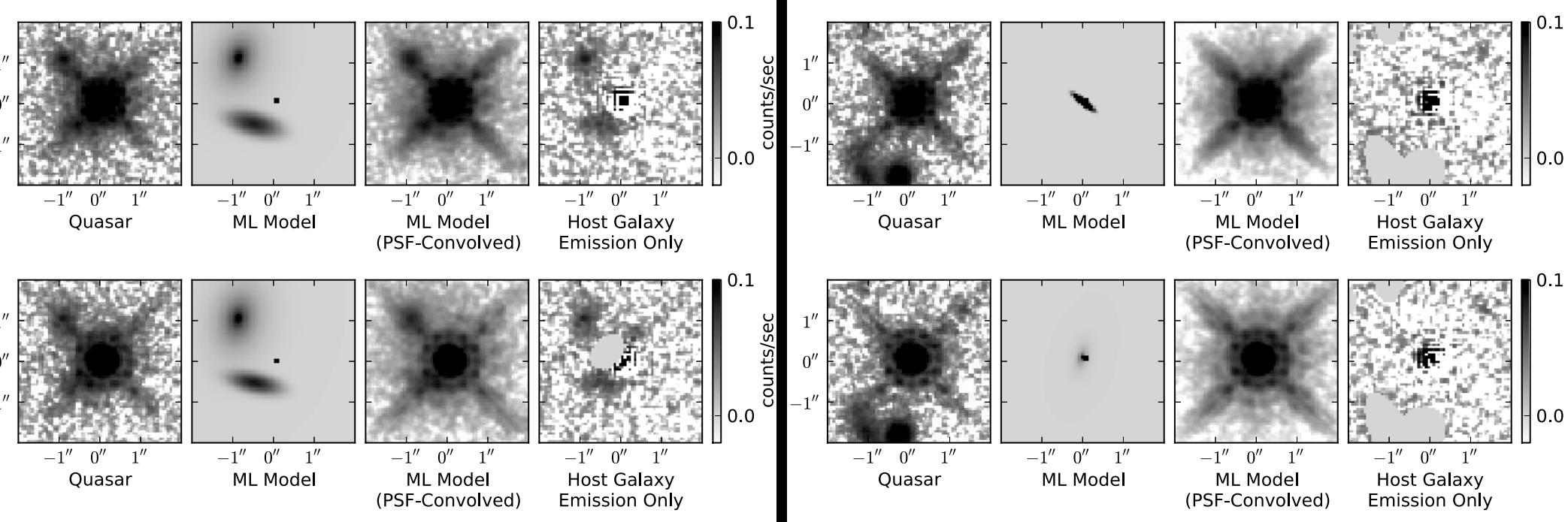
- JWST will trace older stellar pops and SF in much dustier environments.
- We must do all we can with HST in the UV–blue before JWST flies.

(2b) HST WFC3 observations of QSO host systems at $z \simeq 6$ (age $\lesssim 1$ Gyr)



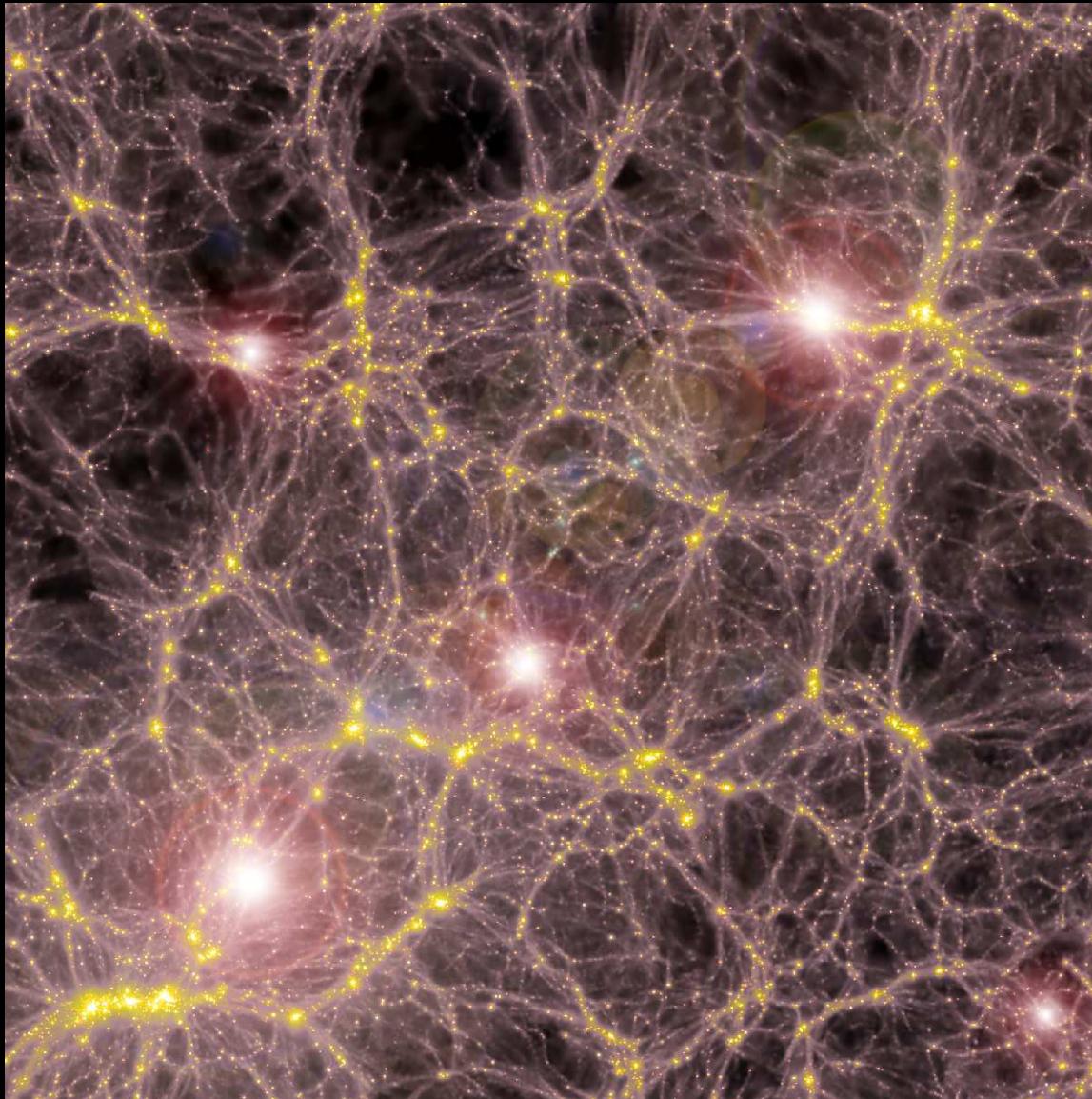
- Careful contemporaneous orbital PSF-star subtraction: Removes most of “OTA spacecraft breathing” effects (Mechtley ea 2012, ApJL, 756, L38).
- PSF-star ($AB \simeq 15$ mag) subtracts $z=6.42$ QSO ($AB \simeq 18.5$) nearly to the noise limit: NO host galaxy detected $100\times$ fainter ($AB \gtrsim 23.5$ at $r \gtrsim 0\farcs3$).

(2b) WFC3: Detection of one QSO Host System at $z \simeq 6$ (Giant merger?)



- Monte Carlo Markov-Chain of observed PSF-star + Sersic ML light-profile. Gemini AO images to pre-select PSF stars (Mechtley⁺ 2014).
 - First detection out of four $z \simeq 6$ QSOs [2 more to be observed].
 - One $z \simeq 6$ QSO host galaxy: Giant merger morphology + tidal structure??
 - Same J+H structure! Blue UV-SED colors: $(J-H) \simeq 0.19$, constrains dust.
 - $M_{AB}^{host}(z \simeq 6) \lesssim -23.0$ mag, i.e., ~ 2 mag brighter than $L^*(z \simeq 6)$!
- $\Rightarrow z \simeq 6$ QSO duty cycle $\lesssim 10^{-2}$ ($\lesssim 10$ Myrs); 1/4 QSO's close to Magorrian.
- JWST Coronagraphs can do this $10\text{--}100 \times$ fainter (& for $z \lesssim 20$, $\lambda \lesssim 28\mu\text{m}$).

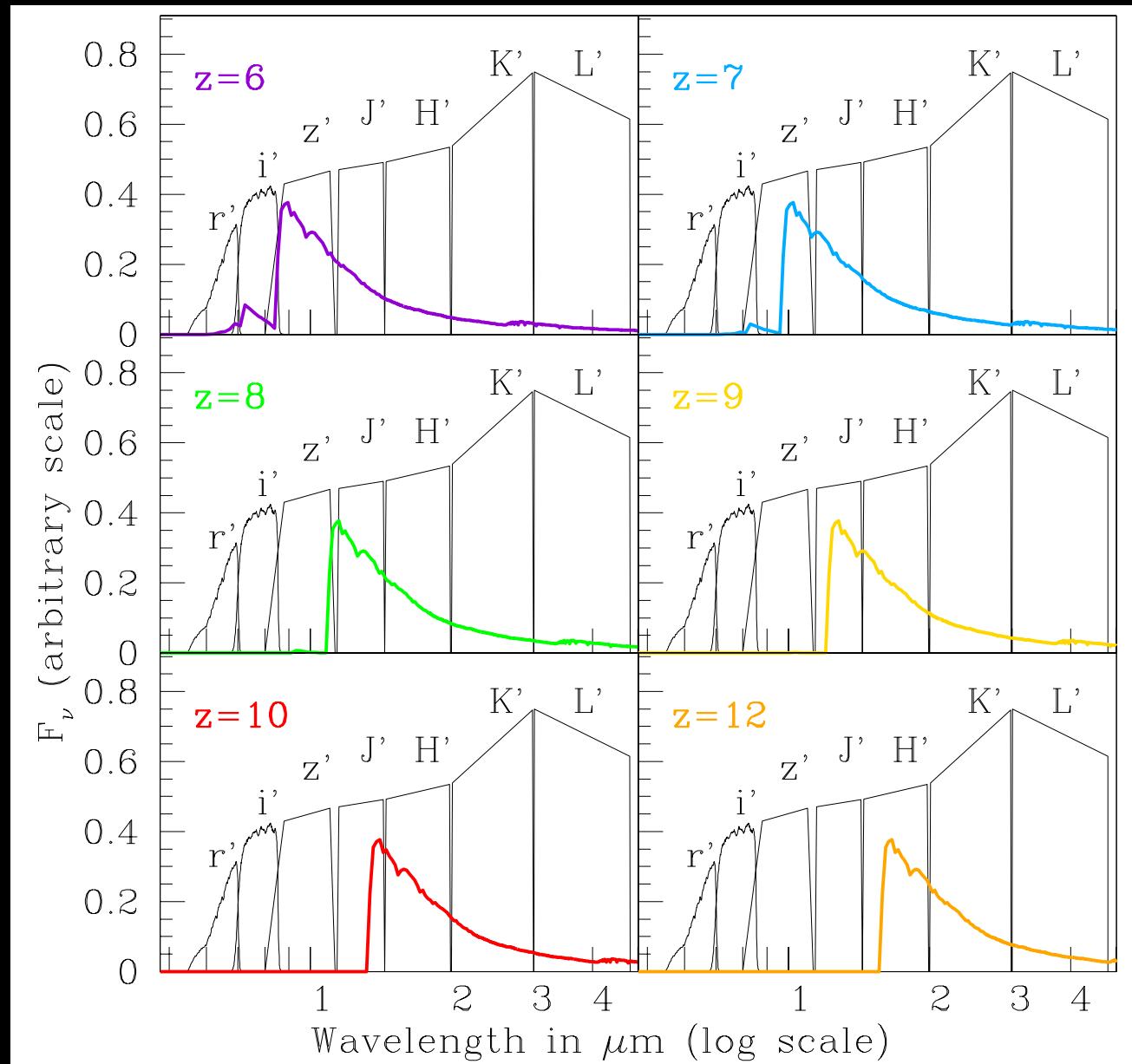
(3) How will JWST Observe First Light and Reionization?



- Detailed cosmological models (V. Bromm) suggest that massive “Pop III” stars ($\gtrsim 100 M_{sun}$) started to reionize the universe at $z \lesssim 10-30$ (First Light).
- This should be visible to JWST as the first Pop III stars or surrounding (Pop II.5) star clusters, and perhaps their extremely luminous supernovae at $z \simeq 10 \rightarrow 30$.

We must make sure that we theoretically understand the likely Pop III mass-range, their IMF, their duplicity and clustering properties, their SN-rates, etc., before JWST flies, so we know what to look for.

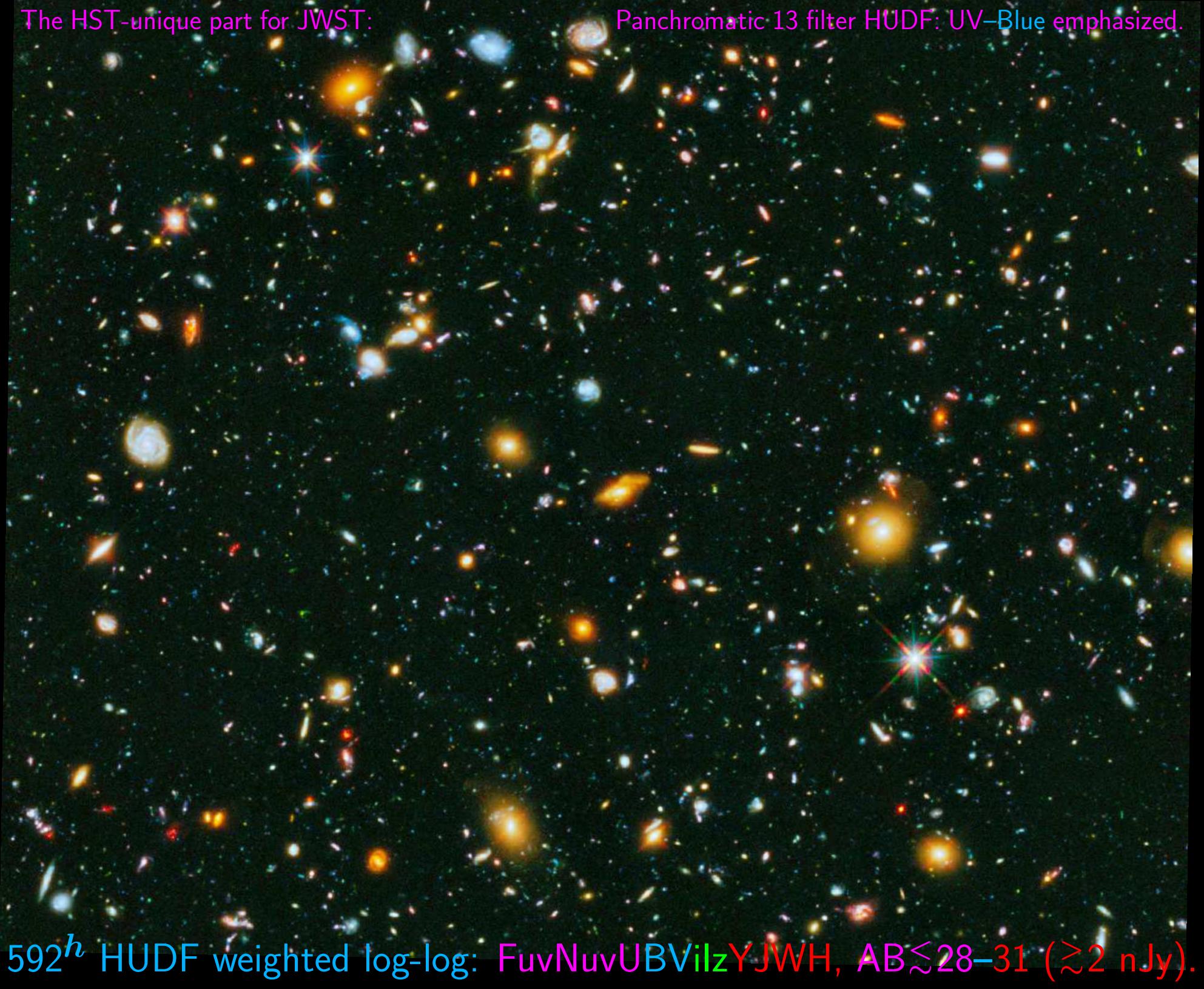
3) How will Webb measure First Light: What to expect in (Ultra)Deep Fields?



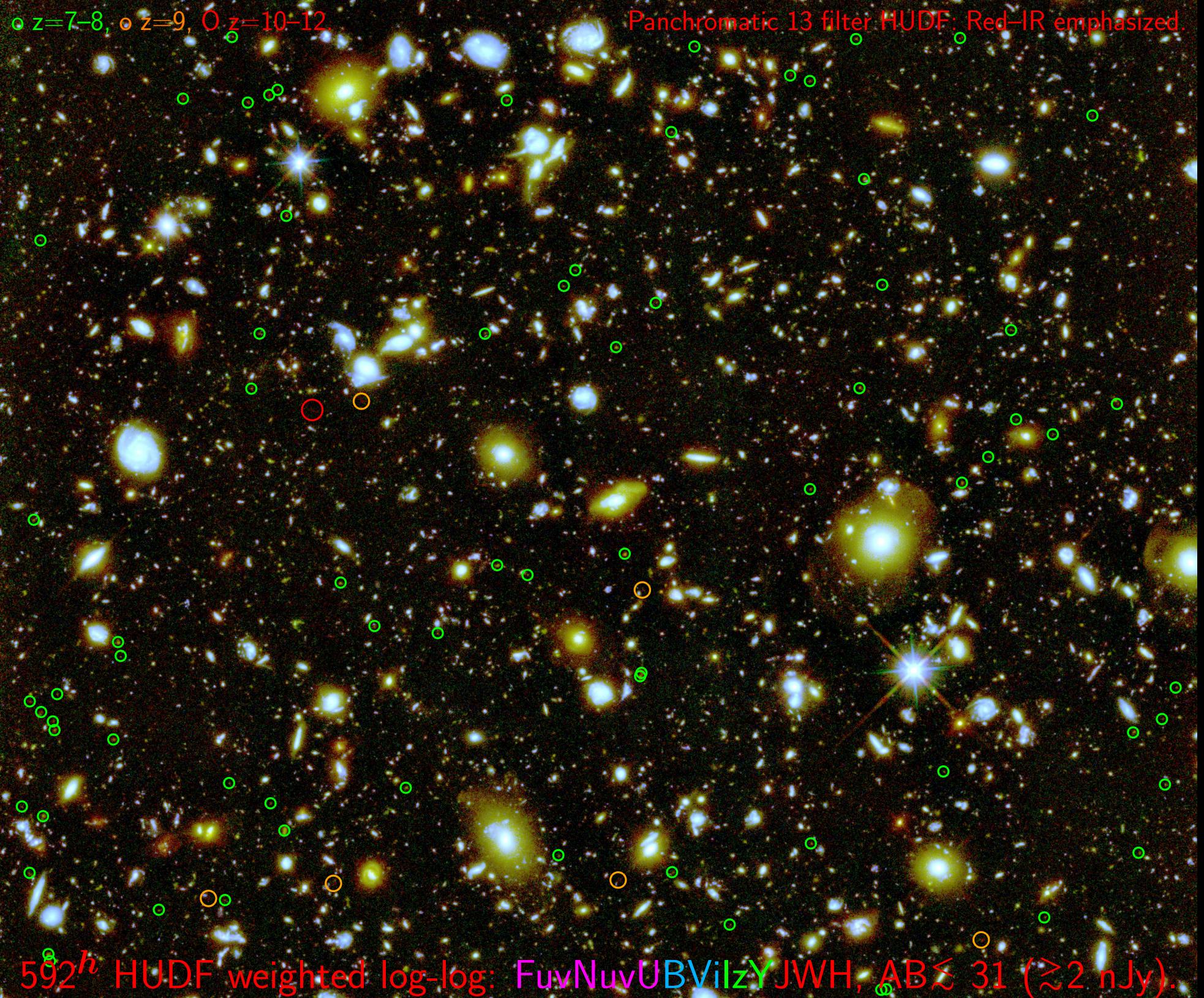
- Can't beat redshift: to see First Light, must observe near-mid IR.
⇒ This is why JWST needs NIRCam at 0.8–5 μm and MIRI at 5–28 μm .

The HST-unique part for JWST:

Panchromatic 13 filter HUDF: UV–Blue emphasized.

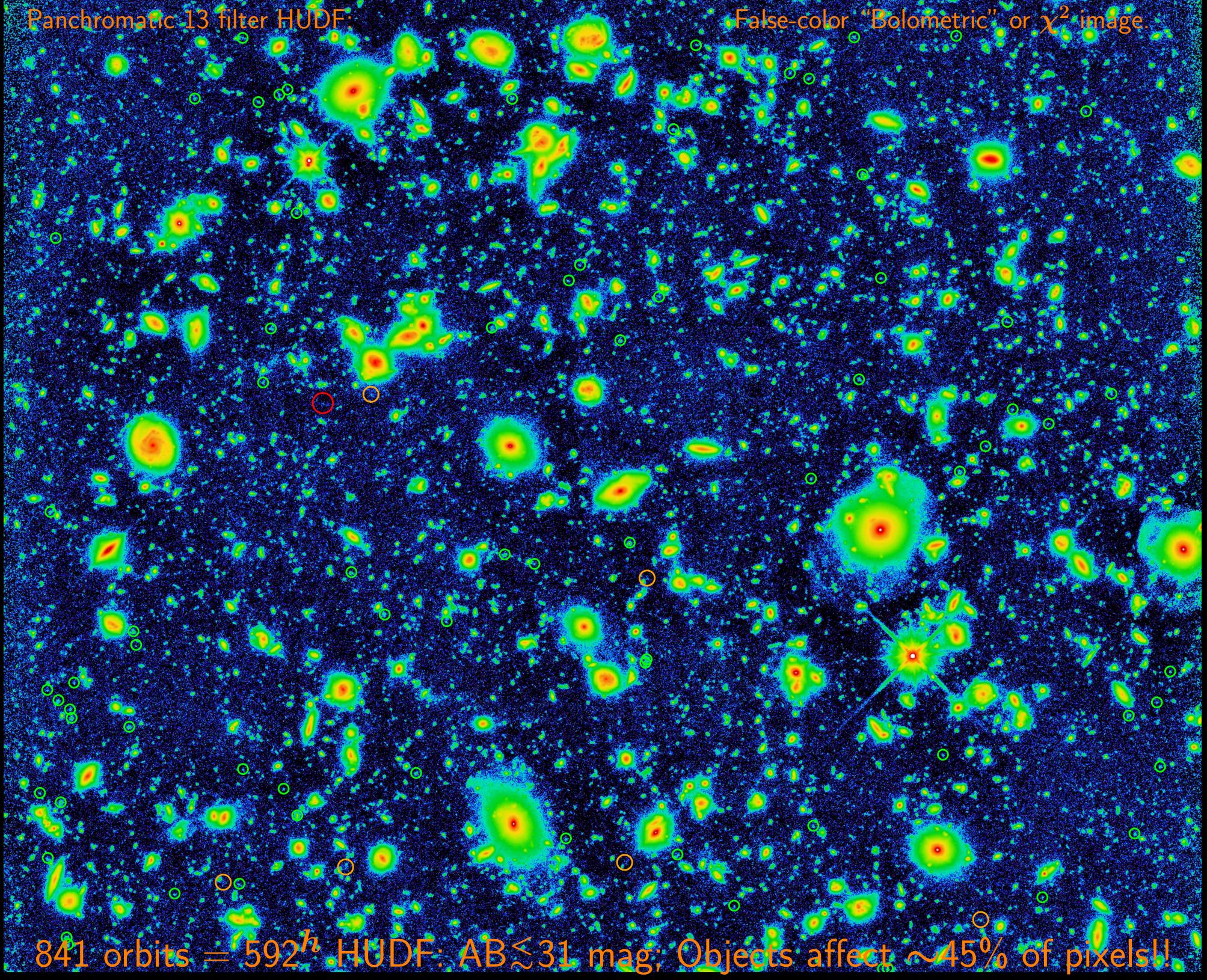


592^h HUDF weighted log-log: FuvNuvUBViIzYJWH, AB \lesssim 28–31 (\gtrsim 2 nJy).

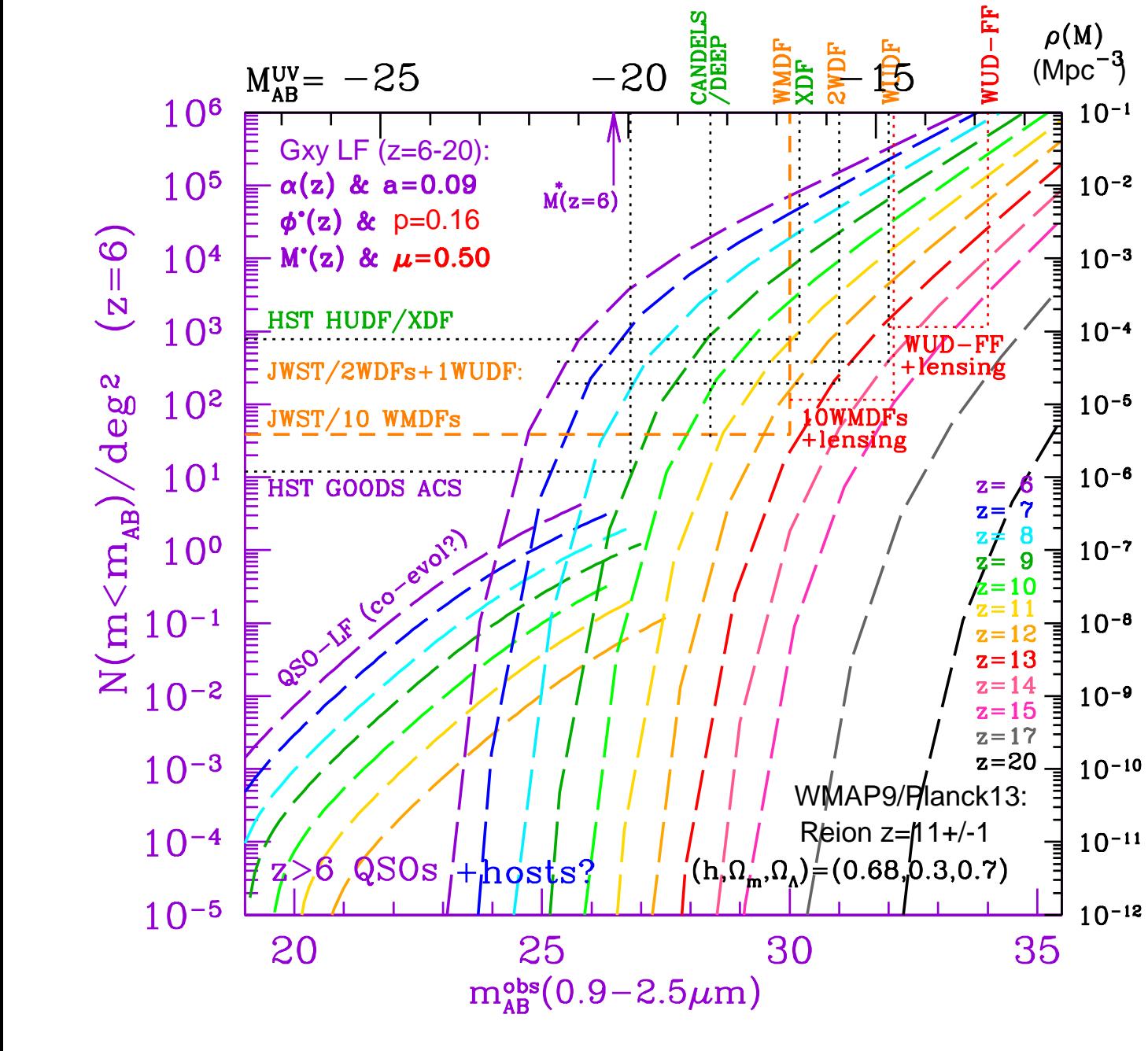


Panchromatic 13 filter HUDF

False-color "Bolometric" or χ^2 image.

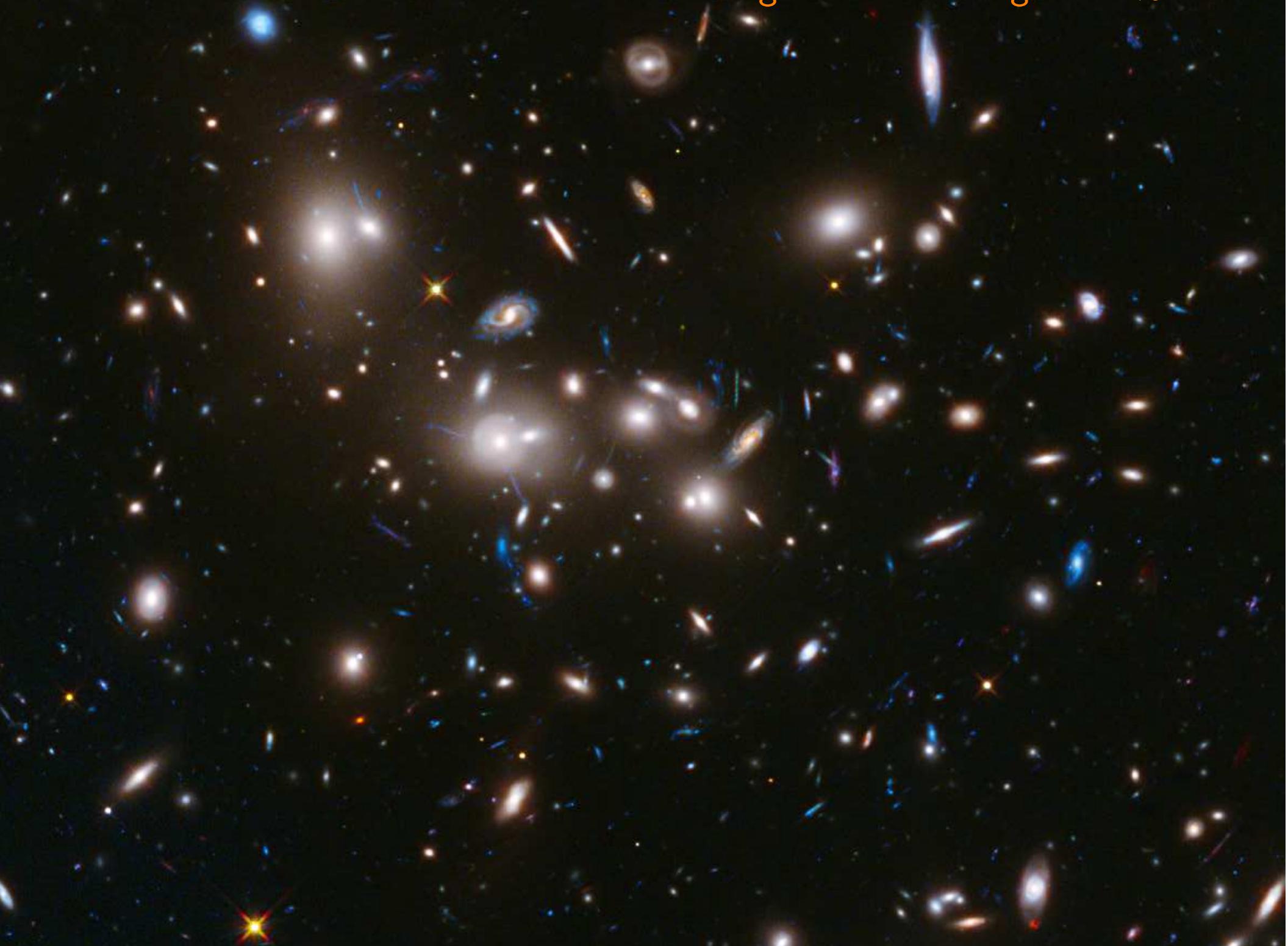


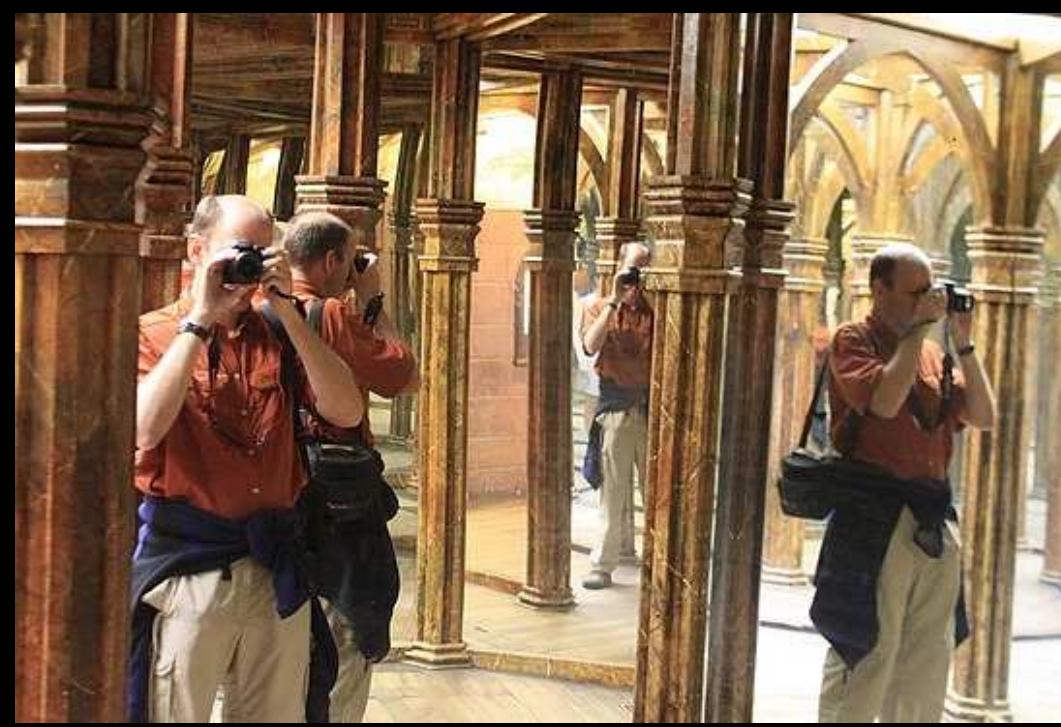
841 orbits = 592^h HUDF: AB \lesssim 31 mag; Objects affect $\sim 45\%$ of pixels!!



Schechter LF ($z \lesssim 6 \lesssim 20$) with best-fit $\alpha(z)$, $\Phi^*(z)$, $M^*(z)$ & $\mu=0.50$.
 Area/Sensitivity for: Hubble Ultra Deep Field (HUDF)/X-ray Deep Field (XDF), 10 Wide-Field Multi-Object Databases (WMDFs), 2 Wide-Field Databases (WDFs), & 1 Wide-Field Ultraviolet Database (WUDF).
 ● May need lensing targets for WMDF–WUDFF to see $z \simeq 14$ –16 objects!

HST Frontier Field A2744: JWST needs lensing to see First Light at $z \gtrsim 11-15$.



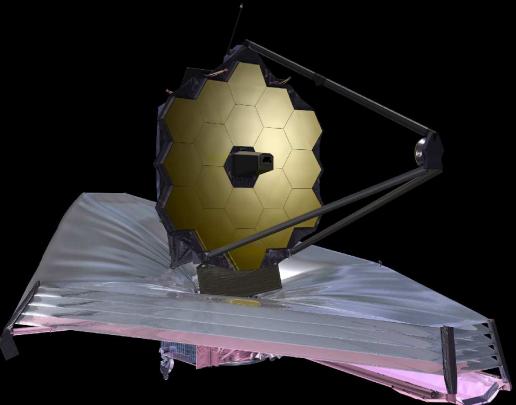


Two fundamental limitations may determine ultimate JWST image depth:

- (1) Cannot-see-the-forest-for-the-trees effect [Natural Confusion limit]:
Background objects blend into foreground because of their own diameter
⇒ Need multi- λ deblending algorithms.
- (2) House-of-mirrors effect [“Gravitational Confusion”]: Most First Light objects at $z \gtrsim 12-14$ may need to be found by cluster or group lensing.
⇒ Need multi- λ object-finder that works on sloped backgrounds.
⇒ If $M^*(z \gtrsim 10) \gtrsim -18$, need to use & model gravitational foreground.

(4) Future: Next generation 20-39 m ground-based telescopes and ATLAST

True relative size: Hubble, James Webb, & Giant Magellan Telescope

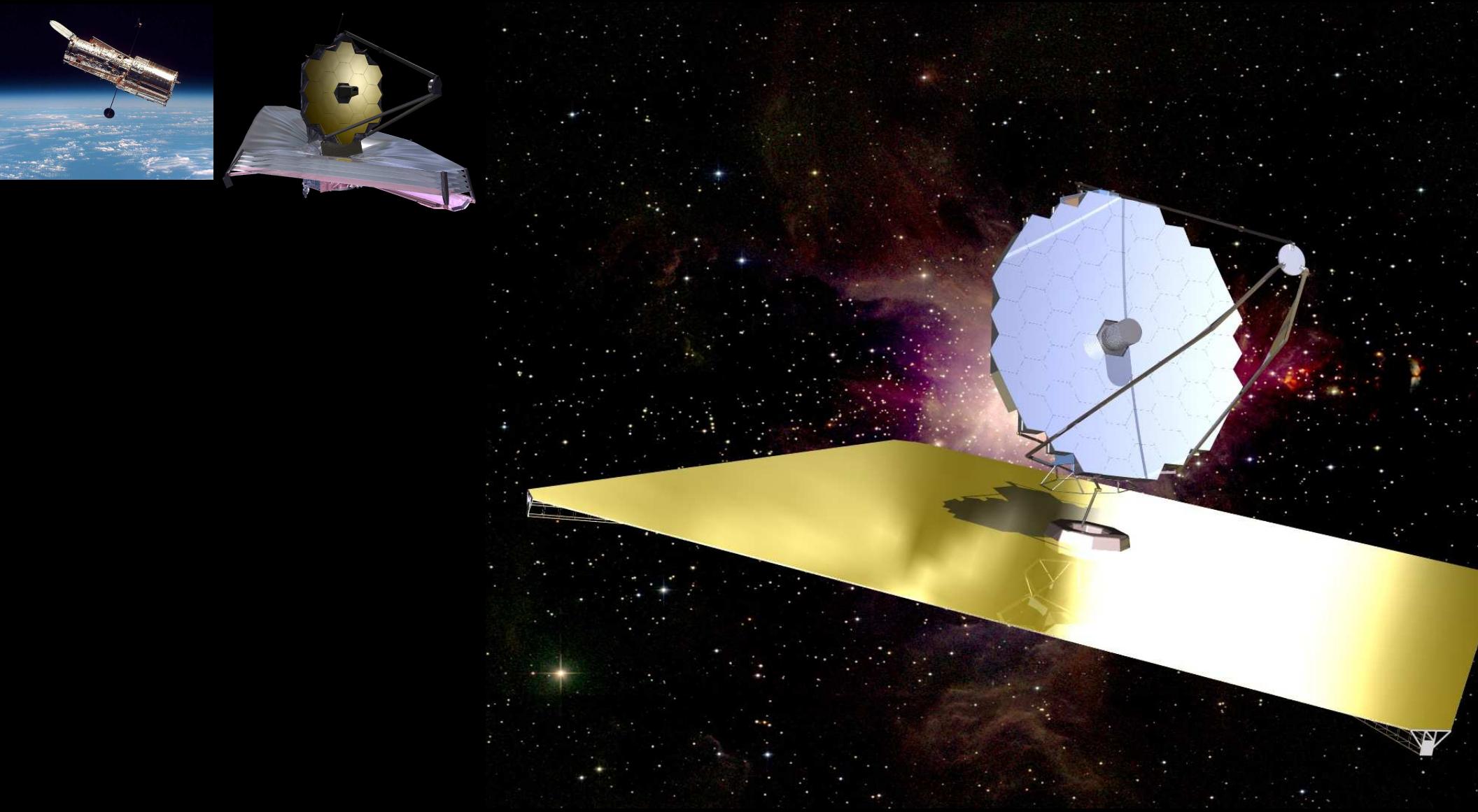


18 B\$ (1973~2018); 9 B\$ (1996~2029);

\sim 1 B\$ (2000~2050 $^+$).

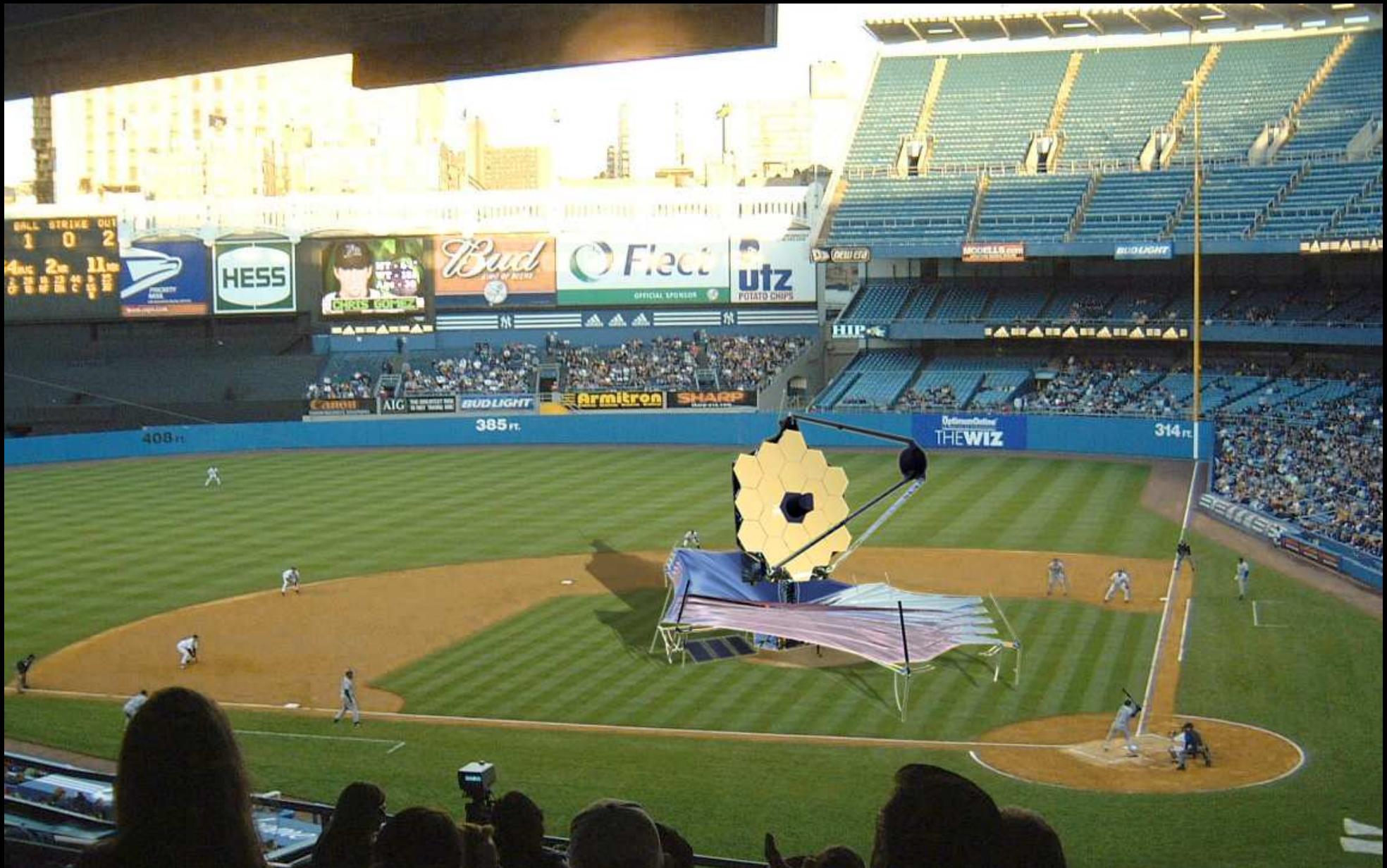
(4) Future: Next generation 20-39 m ground-based telescopes and ATLAST

True relative size: Hubble, James Webb, and ATLAST ...



18 B\$ (1973~2018); 9 B\$ (1996~2029); 15–20 B\$ (2020~2050⁺?).

(4) Future: How can we knock it out of the ball-park in the next 30 years?



Each of GMT and ATLAST facility nearly fills the whole Yankee ballpark ...

- New paradigm: They are too large for an individual university to take on.
- Universities need to collaborate nation-wide to make this happen.

(5) What do our Astrophysics College Graduates do?

Future Careers at NASA:

- Over the last 25 years, (ASU) Astrophysics College Graduates typically:
- (0) Have very low unemployment (\lesssim few %).
- (1) About 30% are faculty at Universities or 4-year colleges.
- (2) About 30% are researchers at NASA or other government centers.
- (3) About 20% work in Aerospace or related industries.
- (4) About 20% are faculty at Community Colleges or Highschools.

See also: <http://aas.org/learn/careers-astronomy>

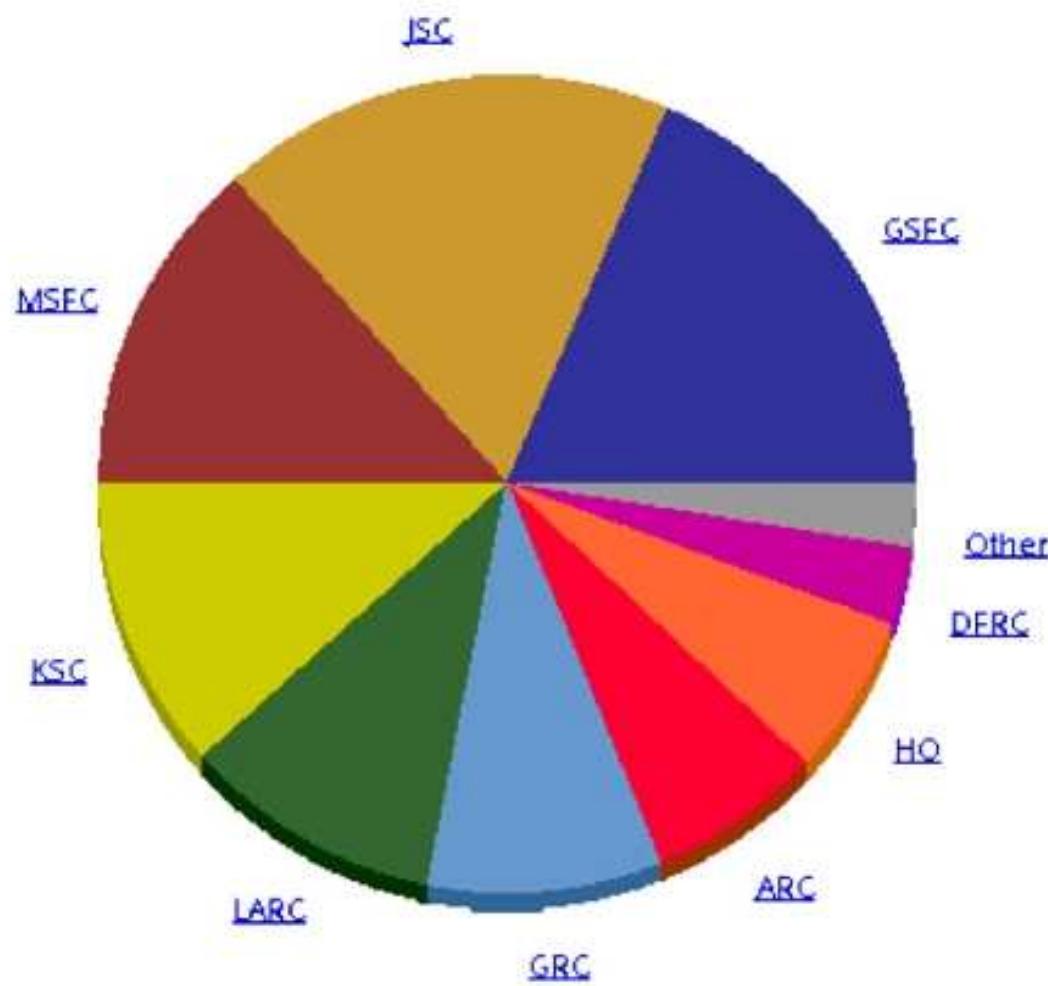
<http://www.aip.org/statistics/astronomy/>

<https://webapp4.asu.edu/programs/t5/careerdetails/19-2011.00?init=false&nopassive=true>

<http://scitation.aip.org/content/aip/magazine/physicstoday/article/68/6/10.1063/PT.3.2815>

CS Head Count

as values



NASA workforce as pie-chart and in numbers — 2013 total: about 18,000).

Nation-wide NASA contractors (Northrup, Lockheed, Boeing, etc): 150,000.

See also: <https://wicn.nssc.nasa.gov/generic.html>

Some of our ASU grad students do important outreach events:



Annual Girl Scout Stargazing at the White House South lawn (July 2015).

Our own Amber Straughn (right; now at NASA GSFC working for Nobel Laureate Dr. John Mather) informs the Obama's about NASA.

(6) Summary and Conclusions

(1) HST set stage to measure galaxy assembly in the last 12.7-13.0 Gyrs.

(2) JWST passed Preliminary & Critical Design Reviews in 2008 & 2010.

Management replan in 2010-2011. No technical showstoppers thus far:

- More than 98% of JWST H/W built or in fab, & meets/exceeds specs.

(3) JWST is designed to map the epochs of First Light, Reionization, and Galaxy Assembly & SMBH-growth in detail. JWST will determine:

- Formation and evolution of the first star-clusters after 0.2 Gyr.
- How dwarf galaxies formed and reionized the Universe after 1 Gyr.

(4) JWST will have a major impact on astrophysics this decade:

- IR sequel to HST after 2018: Training the next generation researchers.
- JWST will define the next frontier to explore: the Dark Ages at $z \gtrsim 20$.

SPARE CHARTS

- References and other sources of material shown:

<http://www.asu.edu/clas/hst/www/jwst/> [Talk, Movie, Java-tool]

<http://www.asu.edu/clas/hst/www/ahah/> [Hubble at Hyperspeed Java–tool]

<http://www.asu.edu/clas/hst/www/jwst/clickonHUDF/> [Clickable HUDF map]

<http://www.jwst.nasa.gov/> & <http://www.stsci.edu/jwst/>

<http://ircamera.as.arizona.edu/nircam/>

<http://ircamera.as.arizona.edu/MIRI/>

<http://www.stsci.edu/jwst/instruments/nirspec/>

<http://www.stsci.edu/jwst/instruments/fgs>

Gardner, J. P., et al. 2006, Space Science Reviews, 123, 485–606

Mather, J., & Stockman, H. 2000, Proc. SPIE Vol. 4013, 2

Windhorst, R., et al. 2008, Advances in Space Research, 41, 1965

Windhorst, R., et al., 2011, ApJS, 193, 27 (astro-ph/1005.2776).

What the Scientists See:

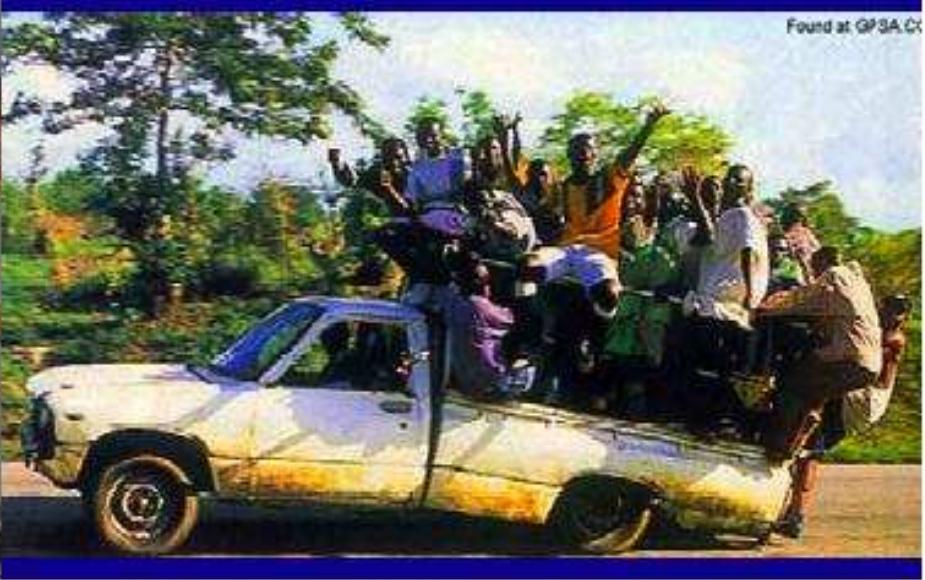


What the Project Manager Sees:



The Happy Balance

Found at GP3A.CX



Any (space) mission is a balance between what science demands, what technology can do, and what budget & schedule allows ... (courtesy Prof. R. Ellis).

Northrop Grumman Expertise in Space Deployable Systems

- Over 45 years experience in the design, manufacture, integration, verification and flight operation of spacecraft deployables
- 100% mission success rate, comprising over 640 deployable systems with over 2000 elements



Baseline "Cup Down" Tower Configuration at JSC (Before)



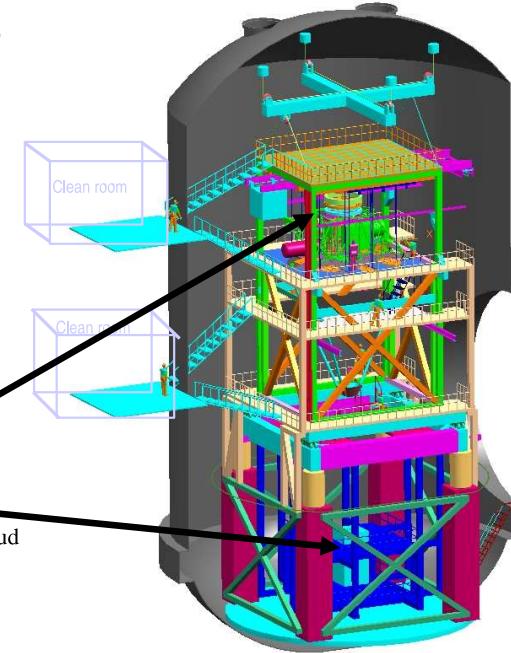
Most recent Tower Design shows an Inner Optical Tower supported by a Outer structure with Vibration Isolation at the midplane. Everything shown is in the 20K region (helium connections, etc. not shown) except clean room and lift fixture.

Current plan calls for 33KW cooldown capability, 12 KW steady state, 300-500mW N2 cooling

JSC currently has 7 KW He capability

Current plan includes 10 trucks of LN2/day during cooldown

Interferometers, Sources, Null Lens and Alignment Equipment Are in Upper and Lower Pressure Tight Enclosure Inside of Shroud



JSC "Cup Up" Test Configuration (New Proposal)

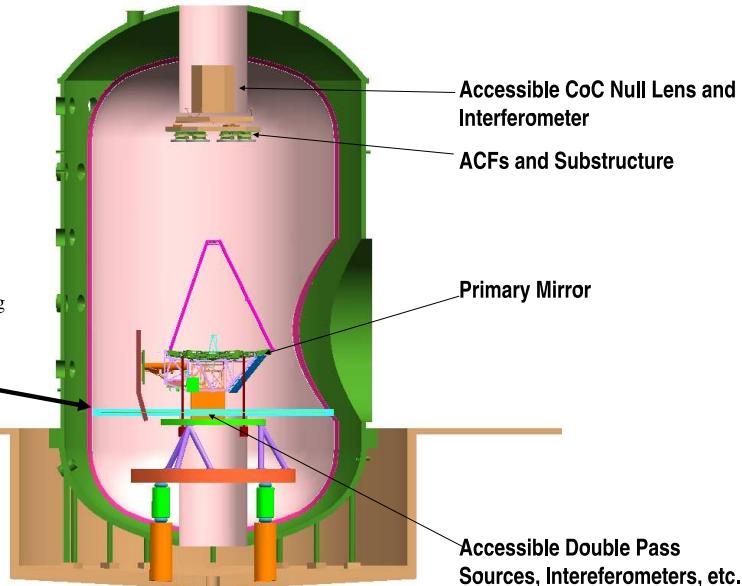


No Metrology Tower and Associated Cooling H/W. External Metrology

Two basic test options:

1. Use isolators, remove drift through fast active control + freeze test equipment jitter
 2. Eliminate vibration isolators (but use soft dampeners) to avoid drift, freeze out jitter
- Builds on successful AMSD heritage of freezing and averaging jitter, testing through windows.

Possible payload "floor" to separate ambient pressure and temperature.



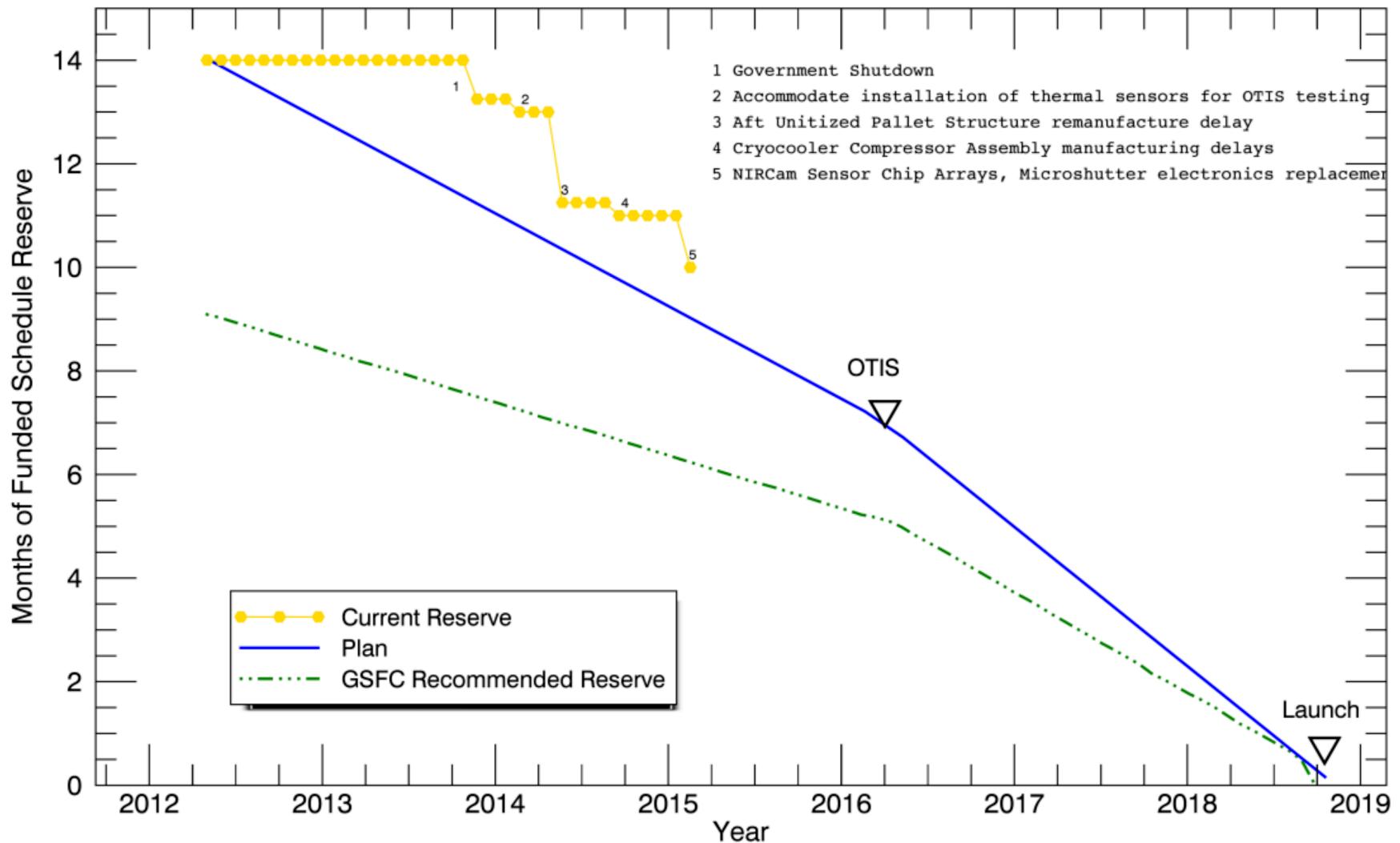
Drawing care of ITT

Page 6

JWST underwent several significant replans and risk-reduction schemes:

- ≈2003: Reduction from 8.0 to 7.0 to 6.5 meter. Ariane-V launch vehicle.
- 2005: Eliminate costly 0.7-1.0 μm performance specs (kept 2.0 μm).
- 2005: Simplification of thermal vacuum tests: cup-up, not cup-down.
- 2006: All critical technology at Technical Readiness Level 6 (TRL-6).
- 2007: Further simplification of sun-shield and end-to-end testing.
- 2008: Passes Mission Preliminary Design & Non-advocate Reviews.
- 2010, 2011: Passes Mission Critical Design Review: Replan Int. & Testing.

Project Funded Schedule Reserve



Keys to stay on schedule: 1) Sufficient Project contingency ($\gtrsim 25\%$ of total).
2) Well replanned and managed Project (starting late summer 2011).

Fiscal Year 2015 JWST HQ Milestones

Month	Milestone	FY2014 Deferral	Comment
Oct-14	1 Secondary Mirror Structure dynamics Test Readiness Review 2 ISIM Cryo-vacuum Test #2 complete 3 Flight and flight spare MIRI Cryocooler Electronics Assembly delivered to JPL 4 Johnson Space Center Optical Ground Support Equipment integration complete	• • • •	Completed 11/20/14 Completed 9/23/14 Completed 10/6/14 Completed 10/10/14
Nov-14	5 Install Engineering Development Unit Secondary Mirror Assembly onto Pathfinder 6 Johnson Space Center (JSC) Chamber A commissioning test start 7 Data Management Subsystem software Build 3 delivery		Completed 10/10/14 Completed 10/18/14 Completed 9/30/14
Dec-14	8 Demonstration model Mid-Boom Assembly thermal vacuum test start 9 Transfer Telescope Pathfinder structure ownership to GSFC 10 Flight Operations Subsystem Build 1 System Design Review 11 Proposal Planning Subsystem Build 10 delivery 12 Deliver flight Cold Head Assembly to ISIM for Cryo-vacuum test #3 13 Fine Guidance Sensor focal plane arrays ready for integration		Completed 11/19/14 Completed 10/31/14 Completed 11/20/14 Completed 11/3/14 • Completed 11/10/14 • Completed 12/9/14
Jan-15	14 Deliver Spacecraft Simulator handbook, Rev B (flight software build 1) to GSFC 15 JSC Chamber A Commissioning complete		Completed 12/11/14 Completed 11/27/14
Feb-15	16 Start formal Engineering Model Test Bed electrical integration 17 Sunshield Mid-boom Manufacturing Readiness Review 18 Sunshield Flight Layer 3 delivered to Northrop-Grumman (NGAS) 19 Deliver Telescope Pathfinder Structure to JSC 20 Observatory Operations Scripts Subsystem Build 4 delivery 21 Wavefront Sensing and Control Software Build 4 delivery	•	Completed 11/13/14 Completed 2/9/15 Completed 2/16/15 Completed 2/4/15 Completed 1/16/15 Completed 12/30/14
Mar-15	22 Qualification Sunshield Membrane Retention Device thermal vacuum test start 23 Deliver Cryocooler Jitter Attenuator Assembly to Optical Telescope Element 24 NGAS Acceptance of Spacecraft propellant tank 25 Near Infrared Instrument Detector changeouts complete		Completed 3/12/15 Delayed to April, ground support equipment issue Delayed to June due to test unit welding issue, no schedule impact Completed 4/3/2014
Apr-15	26 Start acceptance testing of flight Cryocooler Assembly and Electronics 27 Flight Observatory Deployment Tower Assembly complete 28 ISIM Vibration Testing complete 29 Start Optical Ground Support Equipment test #1 at JSC	•	Completed 3/12/15
May-15	30 Flight Cryocooler Compressor Assembly to JPL for Acceptance Test #3 Dual Thruster Module Test Readiness Review 31 Spacecraft Flight Software Build 2.2 Test Readiness Review 32 Sunshield Forward Cover Assemby shipped to NGAS 33 Deliver Flight Aft Optics System to Telescope Pathfinder 34 Data Management Subsystem Build 4 delivery	• •	Milestone deleted, due to change in thruster design, new milestone in June (#36). Moved to July, reprioritizing work for efficiencies at Nexolve
Jun-15	35 Attitude Control System test set delivery to Observatory Integration and Test 36 Propellant Mid-Course Correction testing complete 37 Delivery of new Vibration Test System to GSFC 38 ISIM Acoustic testing complete 39 Proposal Planning Subsystem Build 11 40 Thruster Module Test Readiness Review	•	Completed 2/6/15 (modified milestone to include testing post build of hardware) Completed 2/23/15
Jul-15	41 Flight spare cryo-cooler assembly to JPL for Acceptance Test #4 42 Aft Deployable ISIM Radiator build complete 43 ISIM Electro-Magnetic testing complete	•	
Aug-15	44 Deliver Spacecraft Side Equipment Panels to Observatory integration and testing 45 Deliver Reaction Wheel Assemblies to Observatory integration & testing 46 Start ISIM Cryo-vacuum Test #3 47 Start Optical Ground Support Equipment Test #2 at JSC		
Sep-15	48 Deliver Communications Antenna Bi-Axial Gimbal Assembly to Observatory integration and testing		

Milestones: How the Project reports its progress monthly to Congress.

Milestone Performance

- Since the September 2011 replan JWST reports high-level milestones monthly to numerous stakeholders

	Total Milestones	Total Milestones Completed	Number Completed Early	Number Completed Late	Deferred to Next Year
FY2011	21	21	6	3	0
FY2012	37	34	16	2	3
FY2013	41	38	20	5	3
FY2014 ♦	36	23	10	8	11
FY2015	48	25	16	5*	0

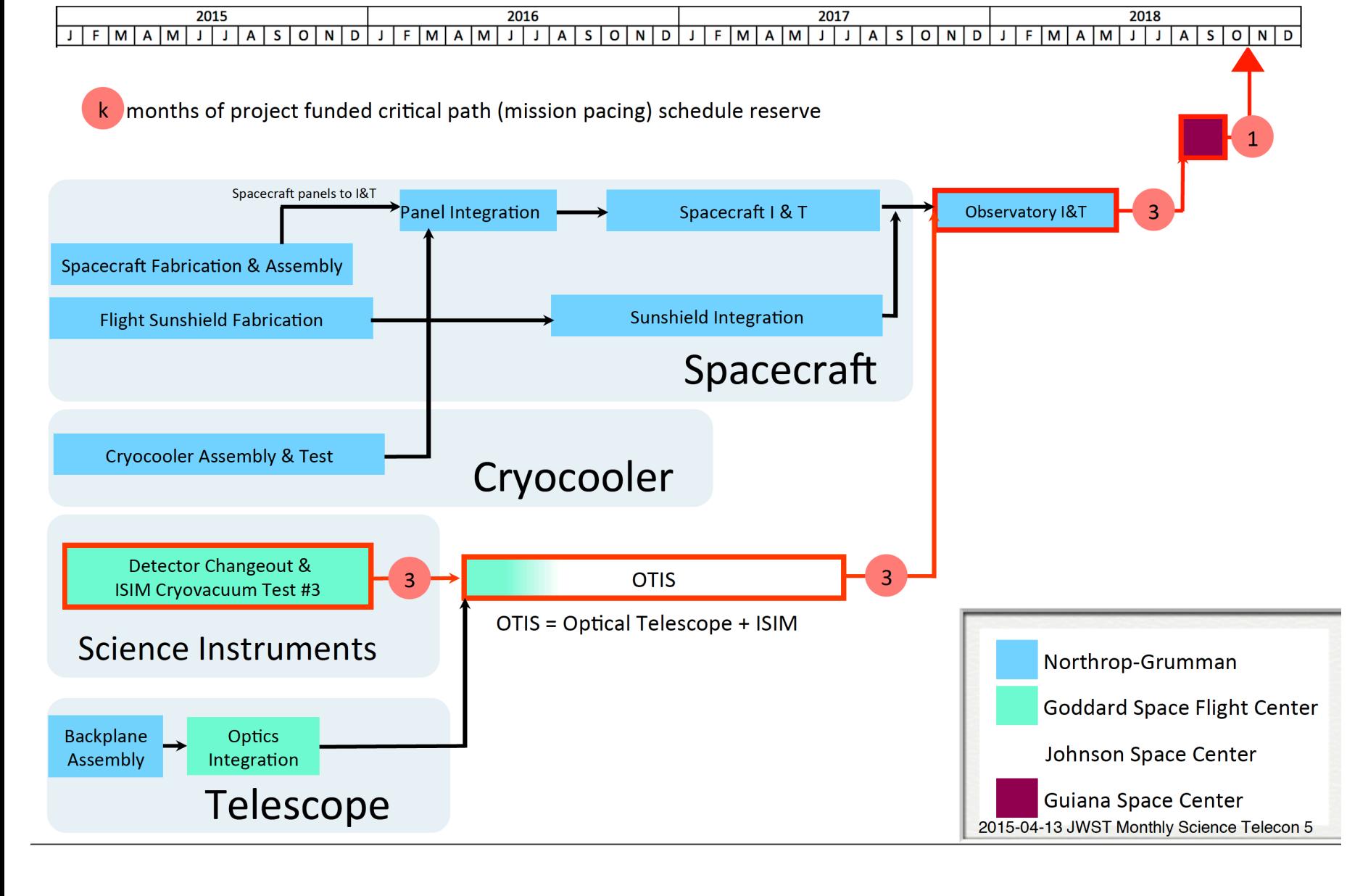
*Late milestones have been or are forecast to complete within the year. Deferred milestones are not included in the number-completed-late tally.

♦ Milestone accounting in FY2014 was complicated by the government shutdown and multicomponent milestones

FY14: 8 milestones late by 1 month due to Oct 13 Government shutdown.

FY15: 4/5 of the “Lates” not on critical path, causing no launch delay.

Simplified Schedule



Path forward to Launch (in Oct. 2018): 10 months schedule reserve.
 Instruments+detectors & Optical Telescope Element remain on critical path.



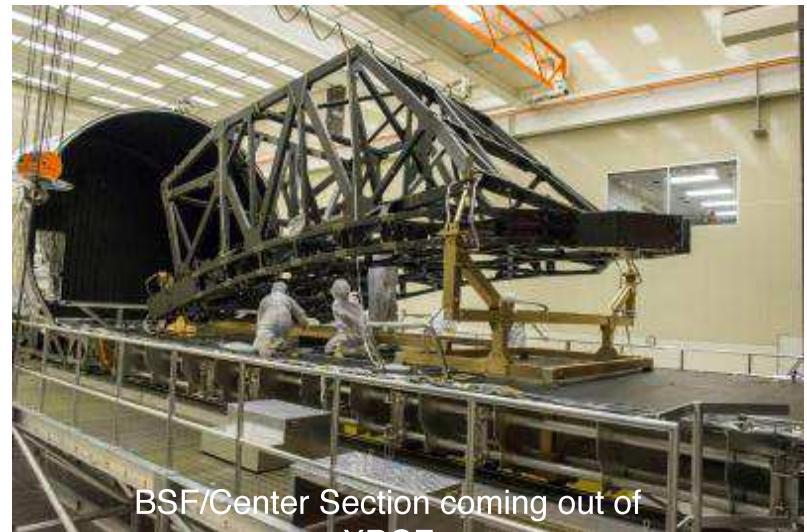
Backplane Support Frame, Center Section, & Wings



- Center Section is complete
- Wings and cryo cycling is complete
- BSF assembly is complete
- Integration of the BSF to Center Section Complete
 - Cryo Cycling at MSFC XRCF complete



BSF and Center Section



BSF/Center Section coming out of XRCF

2014: Flight back-plane ready to receive mirrors starting in 2015.



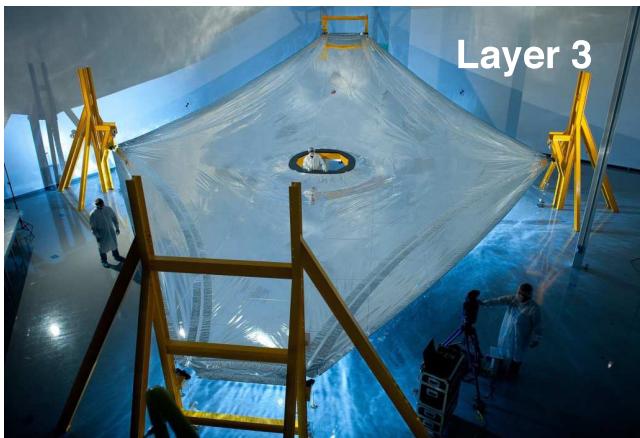
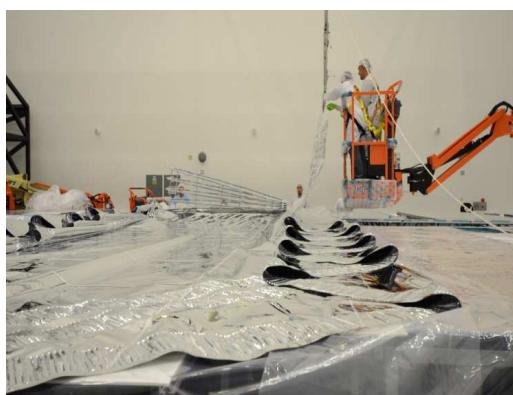
Sunshield Template Membrane Work Completed

Templates Verify Design/Manufacturing Prior to Flight Build

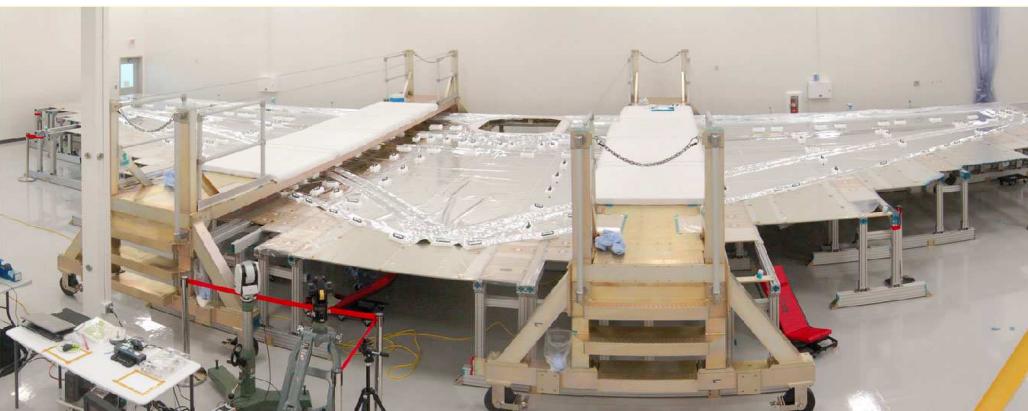


- All Template Layers Completed
- Preparing for flight article manufacturing
- First two Flight Manufacturing Readiness Reviews Completed
- Membrane pull out test complete

Stringing Operations



Template
Layers 3-5



Hole Tool Operations

Flight sunshield to be completed & tested by 2015 at Northrop (CA).

Telescope Assembly Ground Support Equipment



Ambient Optical Alignment Stand
GSFC



Hardware has been installed
at GSFC approximately 8
weeks ahead of schedule



March 2012 NAC
Science Meeting



Landing a mirror onto
backplane simulator



Chamber doesn't look so big anymore!



NORTHROP GRUMMAN

Ball

EXELIS

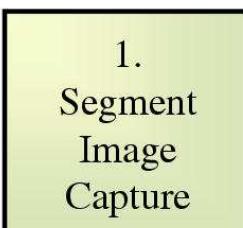
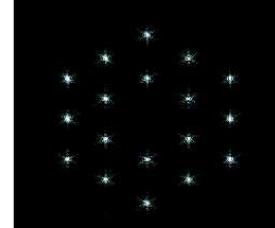
ATK
ALLIANT TECHSYSTEMS



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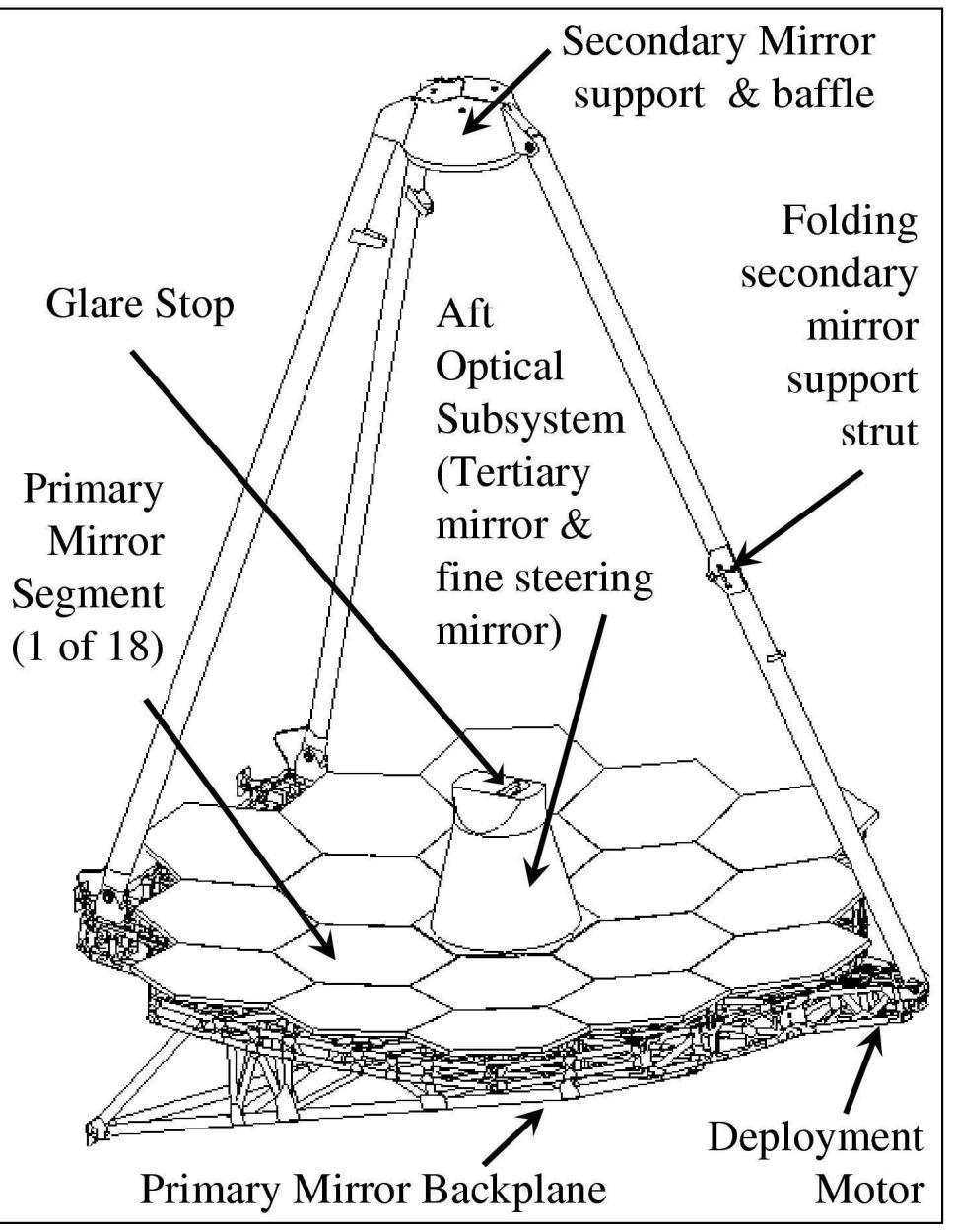
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July 2015: OTIS — World's largest TV chamber readied to test JWST.

<i>First light NIRCam</i>	<i>After Step 1</i>	<i>Initial Capture</i>	<i>Final Condition</i>
		18 individual 1.6-m diameter aberrated sub-telescope images PM segments: < 1 mm, < 2 arcmin tilt SM: < 3 mm, < 5 arcmin tilt	PM segments: < 100 μm , < 2 arcsec tilt SM: < 3 mm, < 5 arcmin tilt
2. Coarse Alignment Secondary mirror aligned Primary RoC adjusted		Primary Mirror segments: < 1 mm, < 10 arcsec tilt Secondary Mirror : < 3 mm, < 5 arcmin tilt	WFE < 200 μm (rms)
3. Coarse Phasing - Fine Guiding (PMSA piston)		WFE: < 250 μm rms	WFE < 1 μm (rms)
4. Fine Phasing		WFE: < 5 μm (rms)	WFE < 110 nm (rms)
5. Image-Based Wavefront Monitoring		WFE: < 150 nm (rms)	WFE < 110 nm (rms)

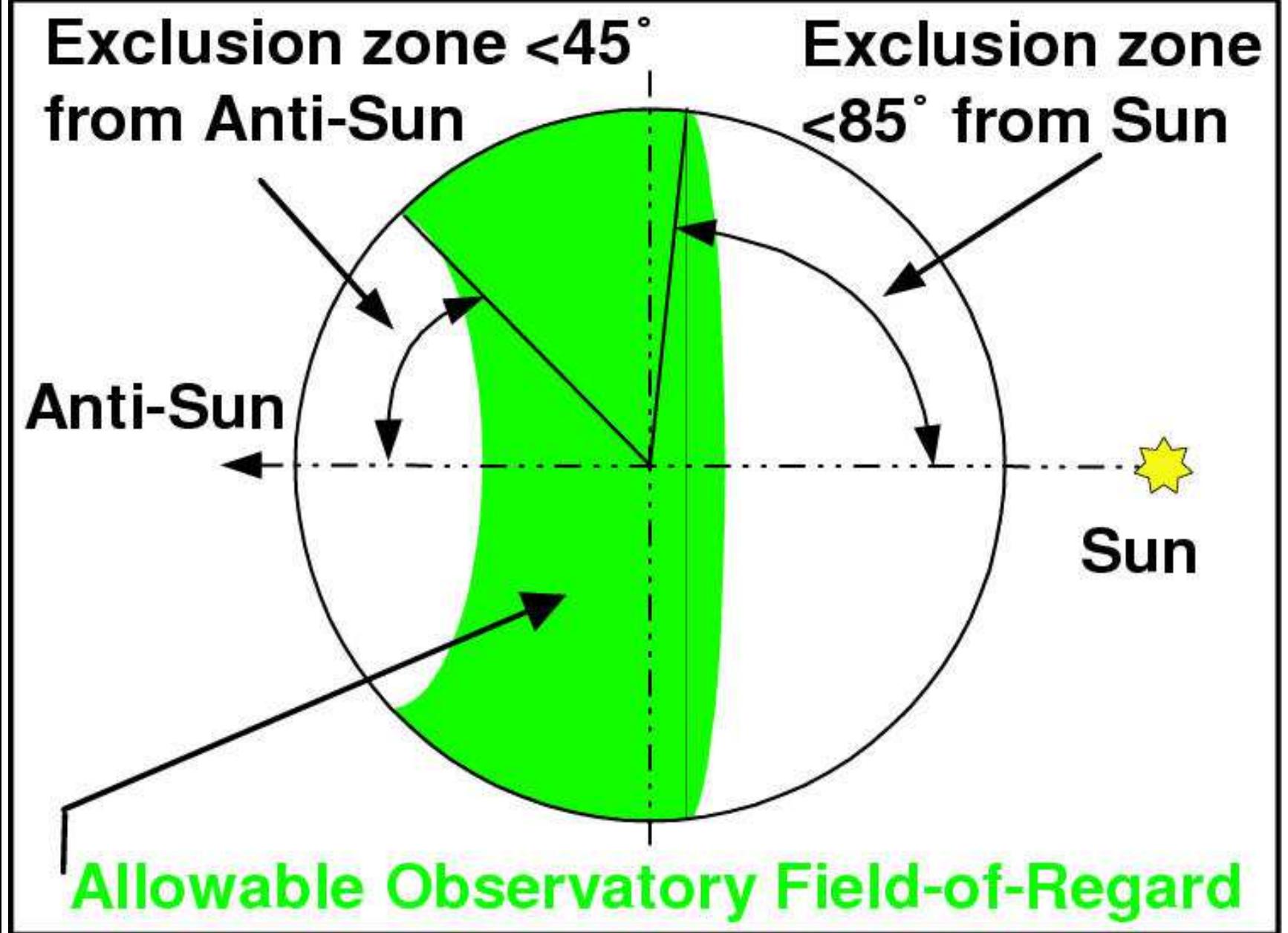
JWST's Wave Front Sensing and Control is similar to the Keck telescope.

In L2, need WFS updates every 10 days depending on scheduling/illumination.



Wave-Front Sensing tested hands-off at 40 K in 1-G at JSC in 2015-2016.

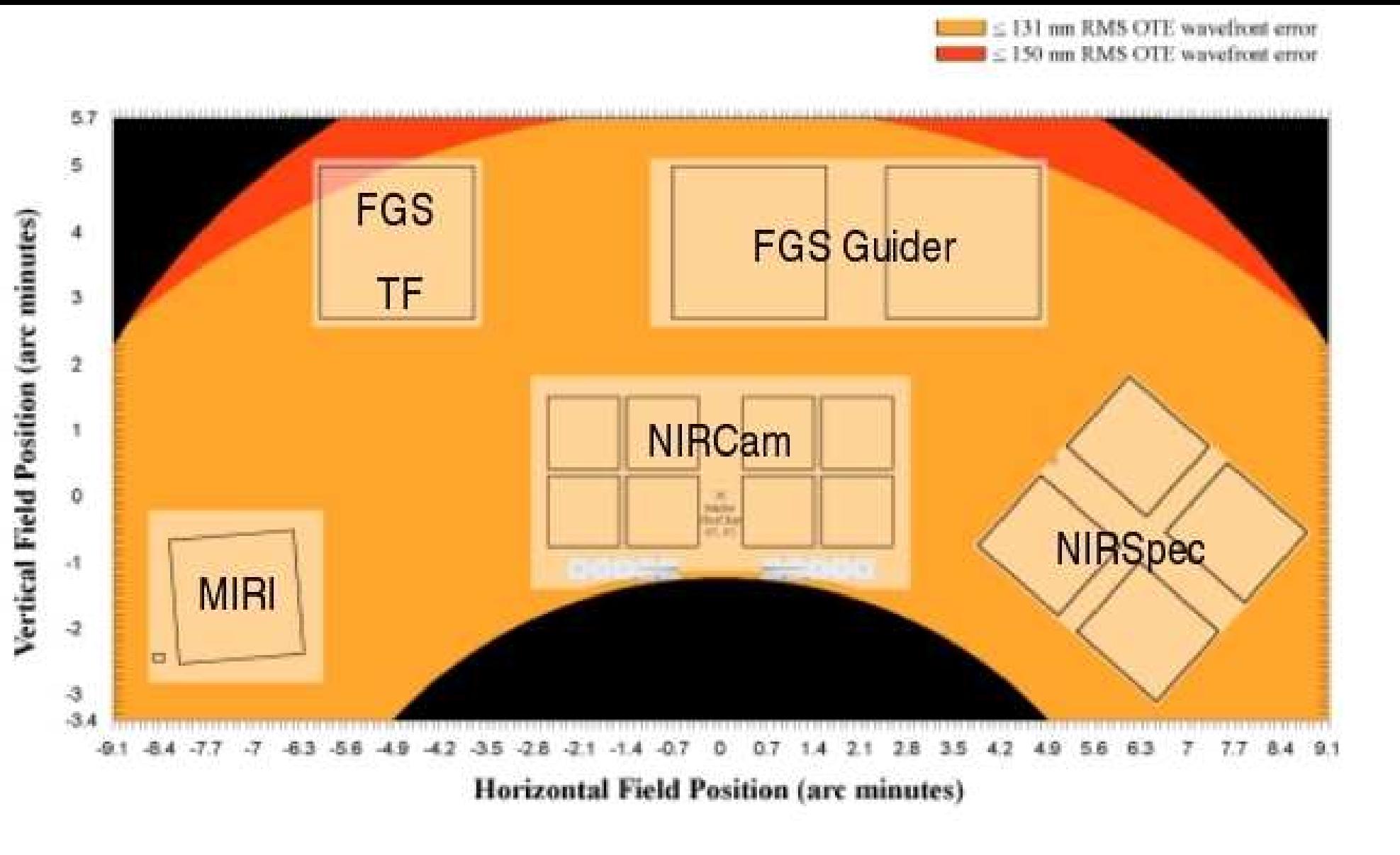
Ball 1/6 scale-model for WFS: produces diffraction-limited 2.0 μm images.



JWST can observe North/South Ecliptic pole targets continuously:

- 1000-hr JWST projects swap back/forth between NEP/SEP targets.
- JWST gets the very best reaction wheels (Rockwell Collins; Heidelberg).

- (3c) What instruments will JWST have?



- All JWST instruments can in principle be used in parallel observing mode:
- Currently only being implemented for parallel *calibrations*.

(4b) Predicted Galaxy Appearance for JWST at redshifts $z \simeq 1-15$

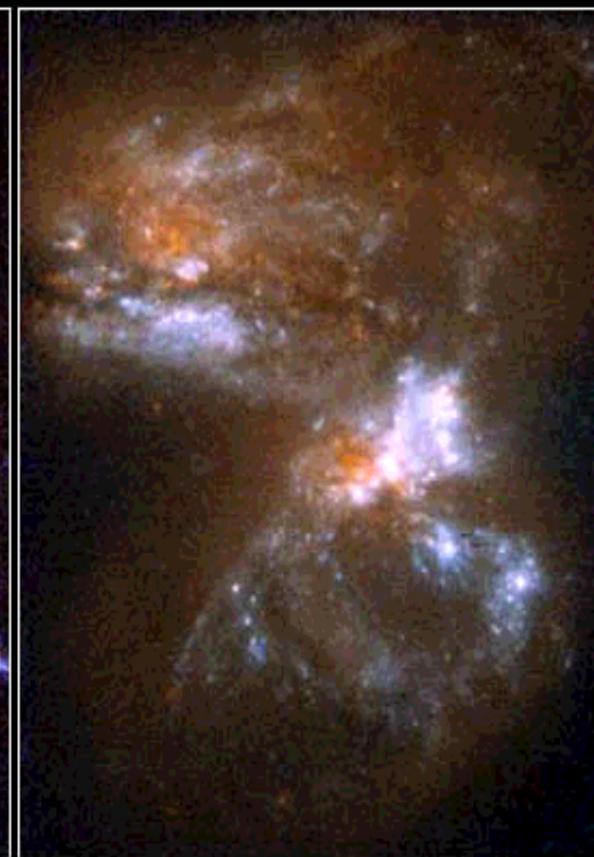
NGC 3310



ESO0418-008



UGC06471-2



Ultraviolet Galaxies

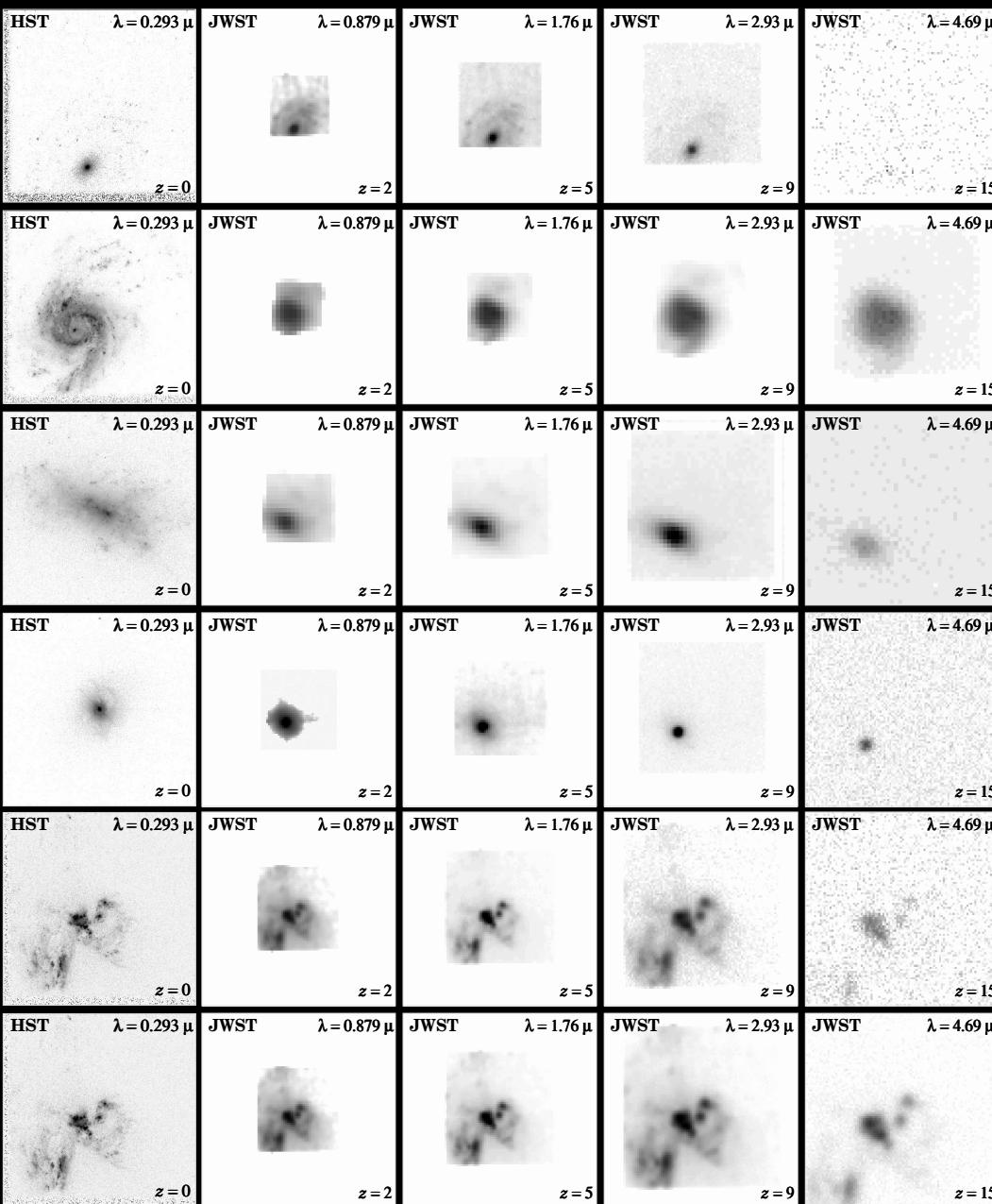
NASA and R. Windhorst (Arizona State University) • STScI-PRC01-04

HST • WFPC2

- The rest-frame UV-morphology of galaxies is dominated by young and hot stars, with often significant dust imprinted (Mager-Taylor et al. 2005).
- High-resolution HST ultraviolet images are benchmarks for comparison with very high redshift galaxies seen by JWST.

(4b) Predicted Galaxy Appearance for JWST at redshifts $z \simeq 1$ –15

HST $z=0$ JWST $z=2$ $z=5$ $z=9$ $z=15$



With Hubble UV-optical images as benchmarks, JWST can measure the evolution of galaxy structure & physical properties over a wide range of cosmic time:

- (1) Most spiral disks will dim away at high redshift, but most formed at $z \lesssim 1$ –2.

Visible to JWST at very high z are:

- (2) Compact star-forming objects (dwarf galaxies).
- (3) Point sources (QSOs).
- (4) Compact mergers & train-wrecks.

B, I, J AB-mag vs.
half-light radii r_e
from RC3 to HUDF
limit are shown.

All surveys limited by
by SB (+5 mag dash)

Deep surveys bounded
also by object density.

Violet lines are gxy
counts converted to
to natural conf limits.

Natural confusion
sets in for faintest
surveys ($AB \gtrsim 25$).
Will update for JWST.

