

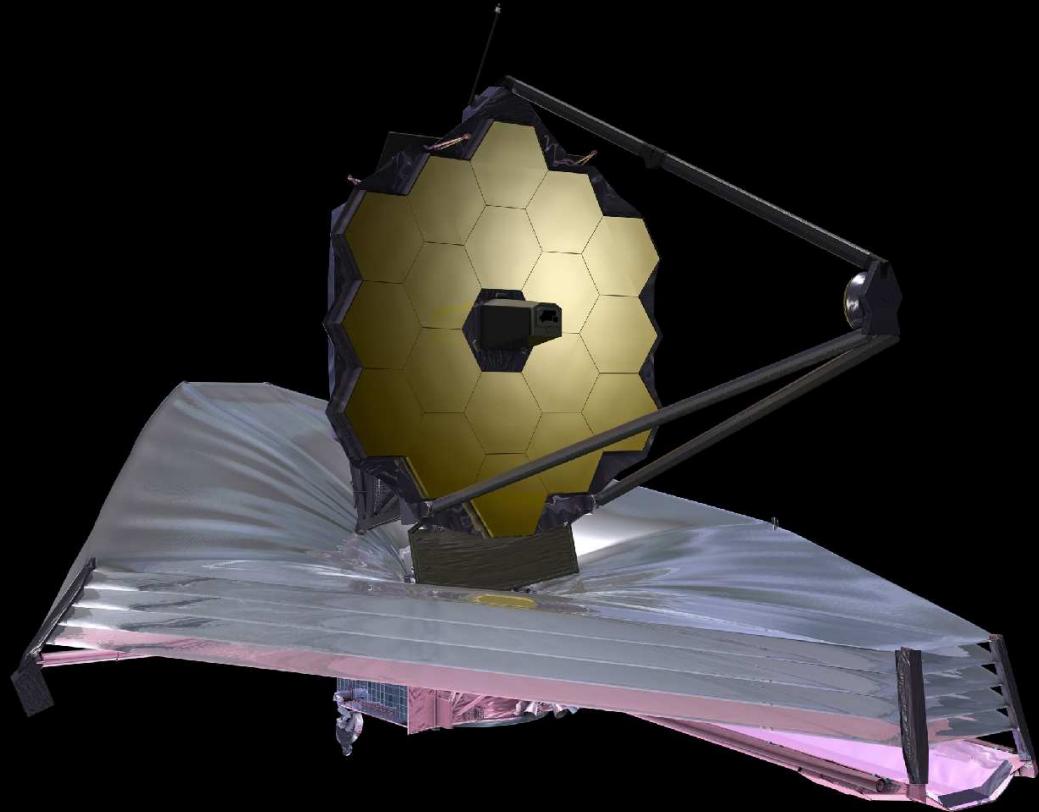
# How will JWST measure First Light, Galaxy Assembly & Supermassive Blackhole Growth: New Frontier after HST

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(Ex) ASU Grads: N. Hathi, H. Kim, M. Mechtley, R. Ryan, M. Rutkowski, B. Smith, & A. Straughn



Colloquium at INAF, Osservatorio Astronomico di Roma, Rome, Italy

Tuesday March 11, 2014. All presented materials are ITAR-cleared.

# Outline

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- (1) Recent key aspects of the Hubble Space Telescope (HST) project.
  - (2) Measuring Galaxy Assembly and Supermassive Black-Hole Growth: including first WFC3  $z \simeq 6$  QSO host galaxy detection ...
  - (3) Brief Update on the James Webb Space Telescope (JWST)?
  - (4) How can JWST measure the Epochs of First Light & Galaxy Assembly, and Supermassive Black-Hole Growth?
  - (5) Summary and Conclusions.
- [● (6) How can JWST measure Star-formation & Earth-like exoplanets?]



Edwin P. Hubble (1889–1953) — Carnegie astronomer

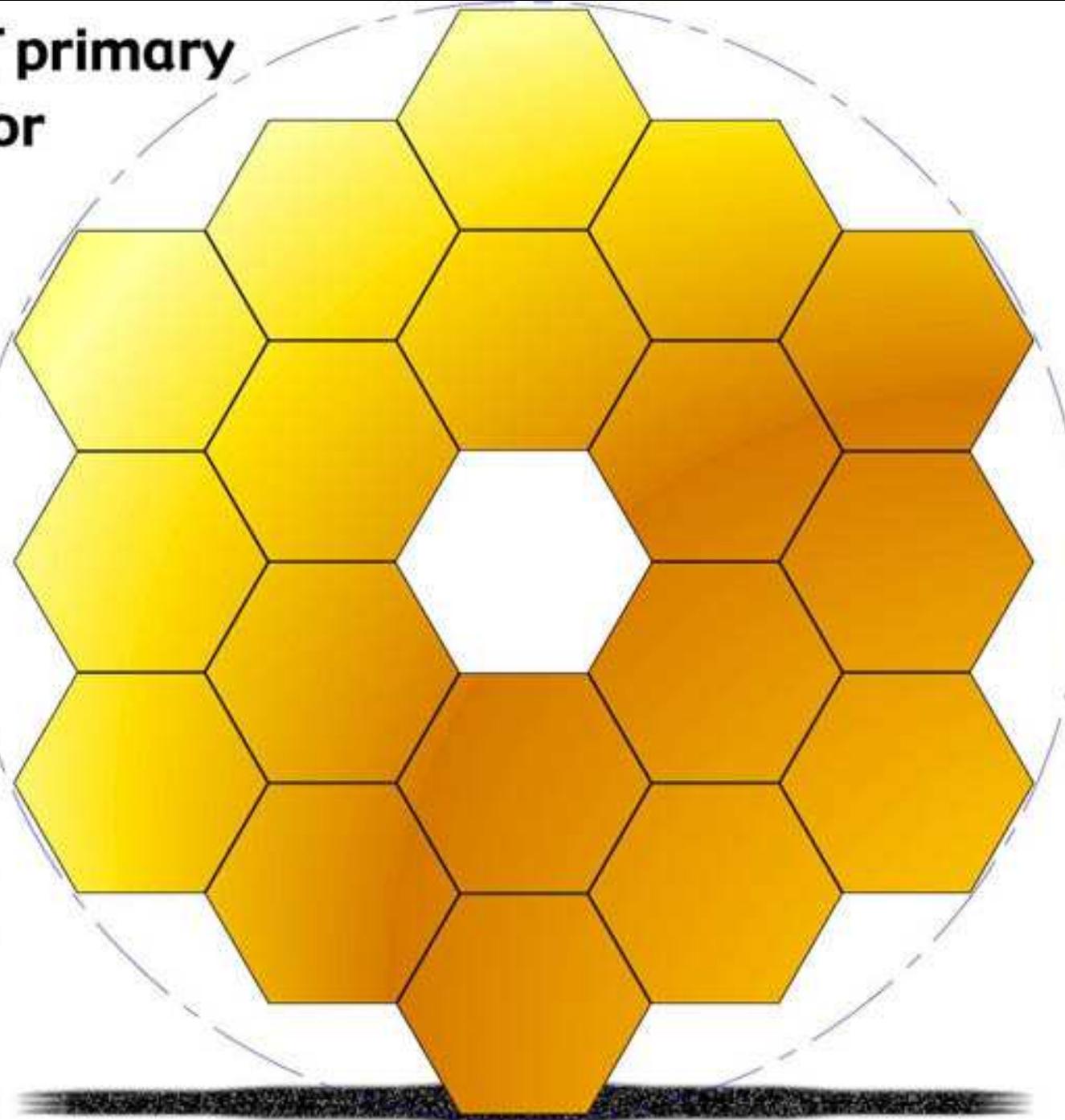


James E. Webb (1906–1992) — Second NASA Administrator

Hubble: Concept in 1970's; Made in 1980's; Operational 1990– $\gtrsim$ 2014.

JWST: The infrared sequel to Hubble from 2018–2023 (–2029?).

**JWST primary  
mirror**

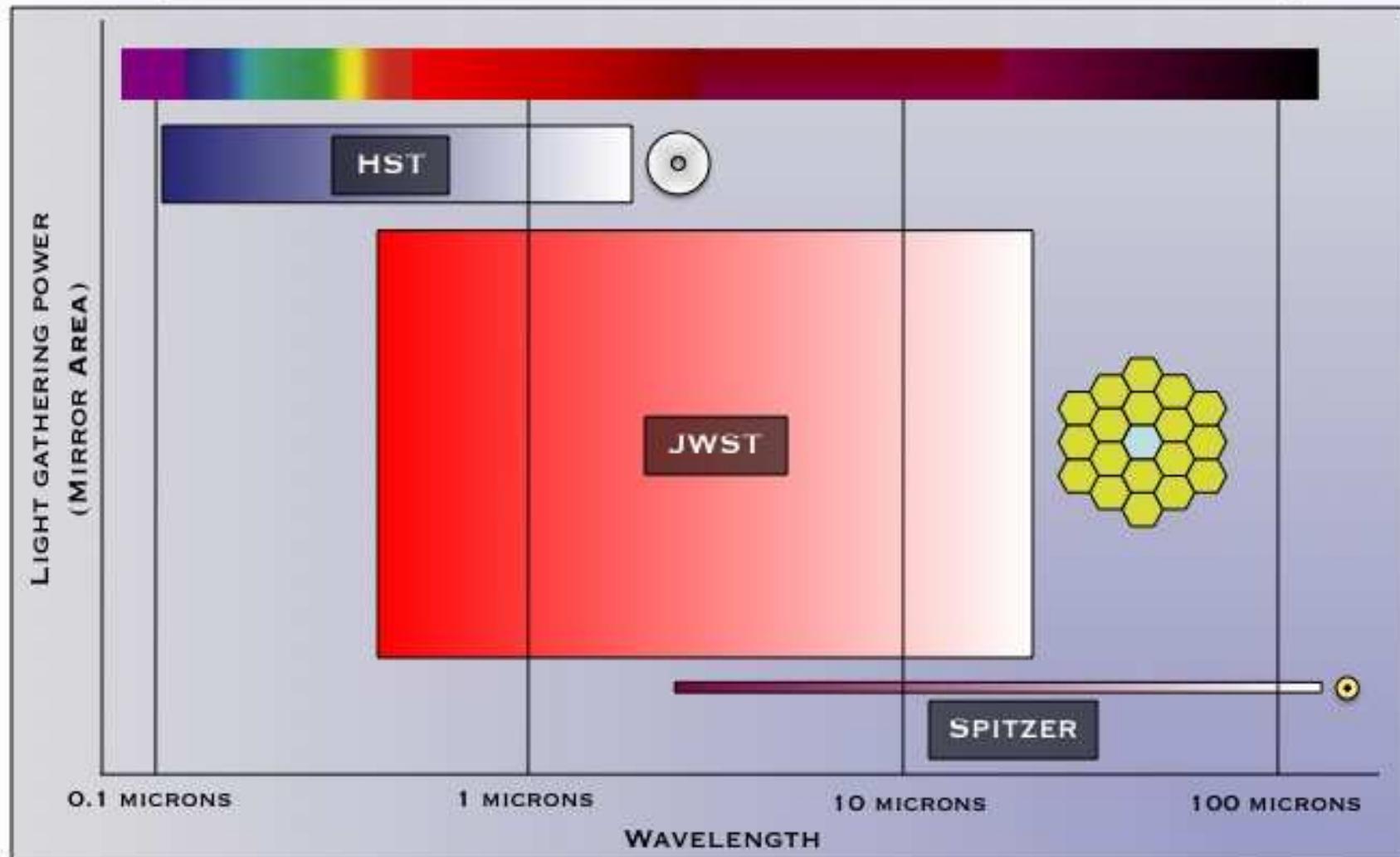


**Hubble primary  
mirror**



JWST  $\simeq 2.5 \times$  larger than Hubble, so at  $\sim 2.5 \times$  larger wavelengths:  
JWST has the same resolution in the near-IR as Hubble in the optical.

# THE JAMES WEBB SPACE TELESCOPE



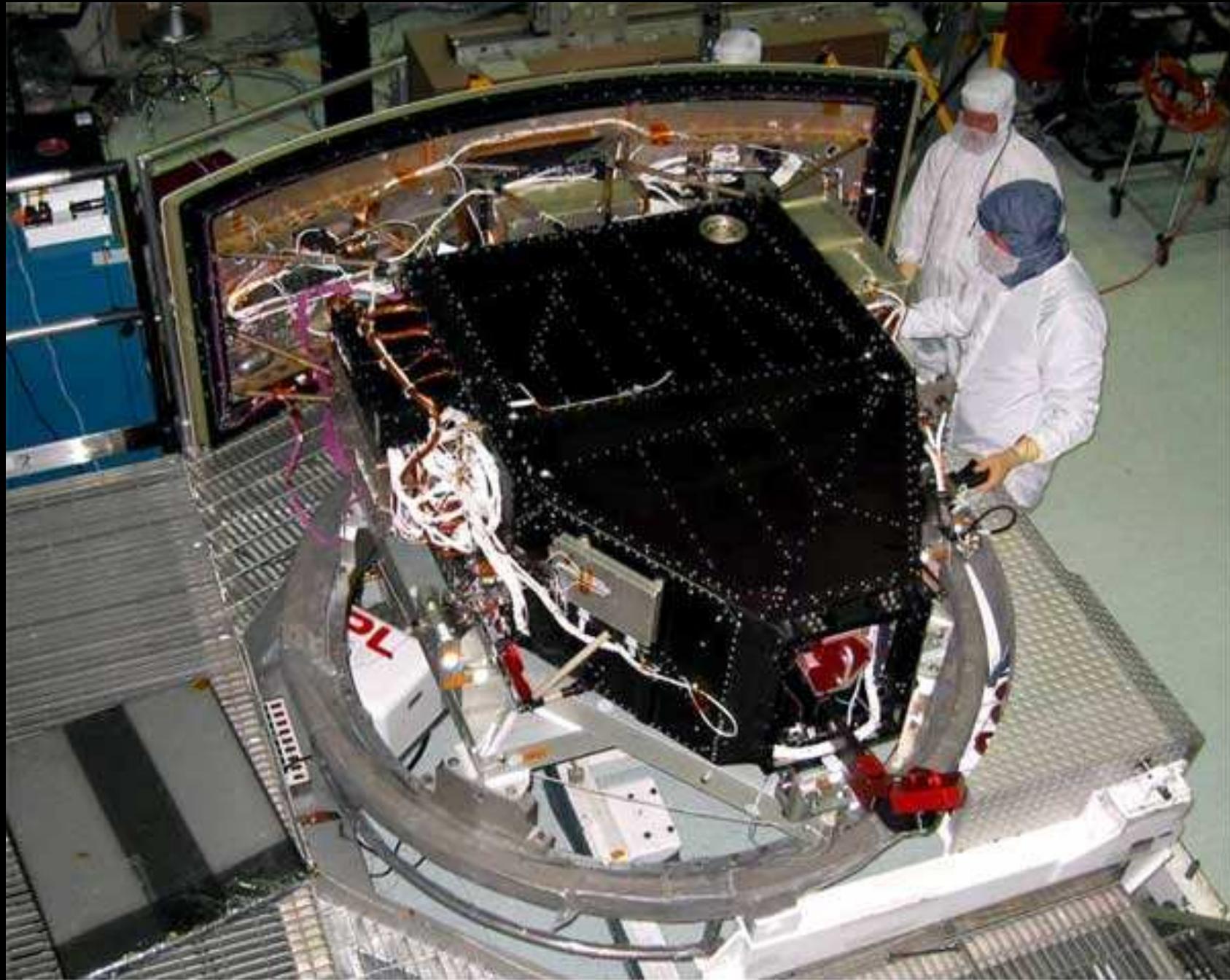
## LIGHT GATHERING POWER

$$\text{JWST} = 25 \text{ m}^2; \text{ HUBBLE} = 4.5 \text{ m}^2; \text{ SPITZER} = 0.6 \text{ m}^2$$

JWST is the perfect near-mid-IR sequel to HST and Spitzer:

- Vastly larger  $A(\times\Omega)$  than HST in UV-optical and Spitzer in mid-IR.

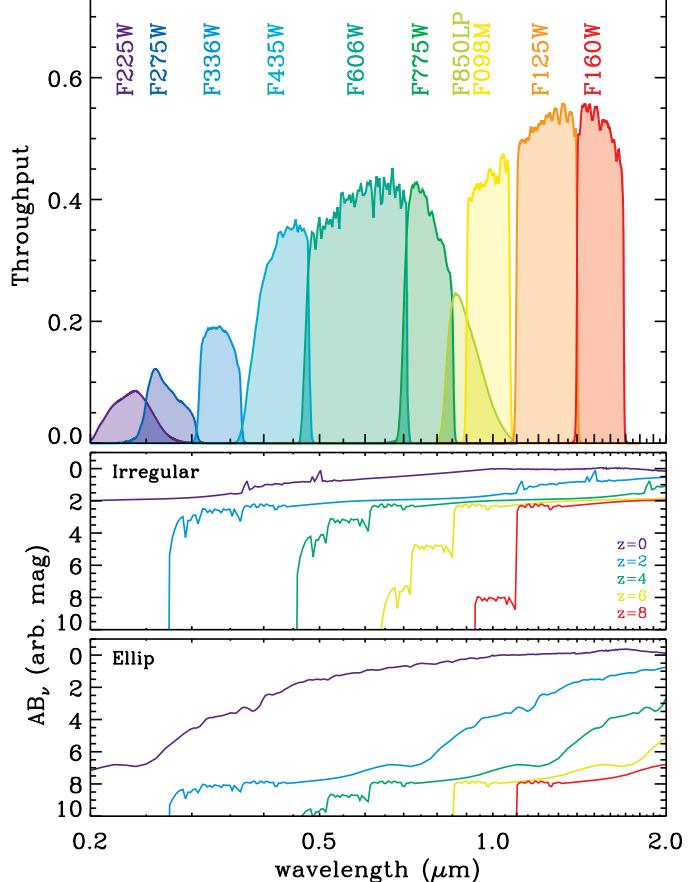
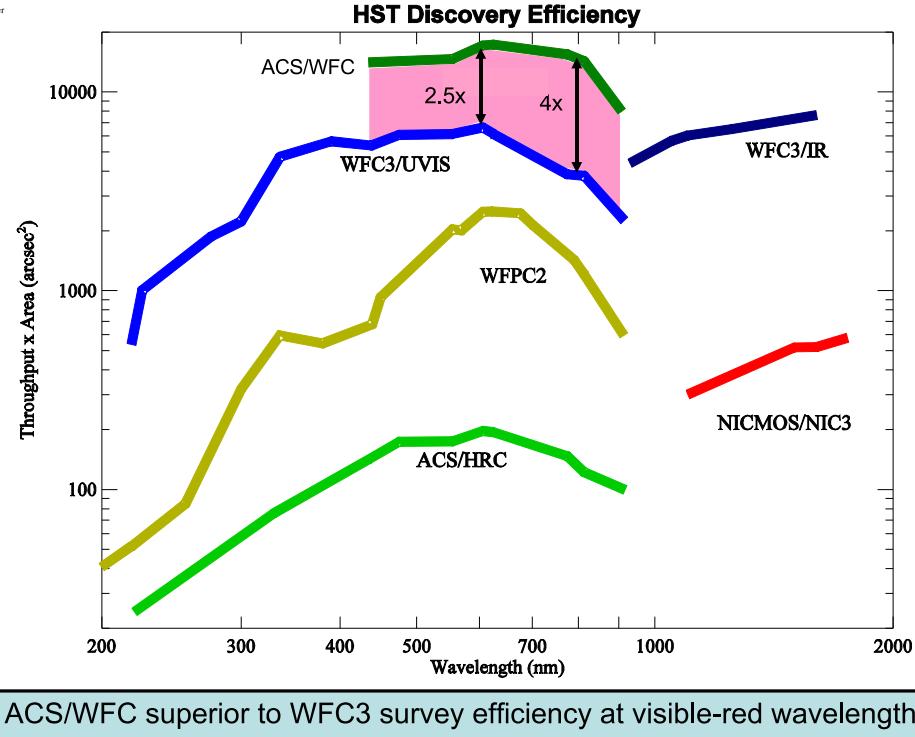
# (1) WFC3: Hubble's new Panchromatic High-Throughput Camera



HST WFC3 and its IR channel: a critical pathfinder for JWST science.



## Role of ACS in HST Post-SM4 Imaging Capability



WFC3/UVIS channel unprecedented UV–blue throughput & areal coverage:

- QE  $\gtrsim 70\%$ ,  $4\text{k} \times 4\text{k}$  array of  $0\farcs04$  pixel, FOV  $\simeq 2\farcs67 \times 2\farcs67$ .

WFC3/IR channel unprecedented near-IR throughput & areal coverage:

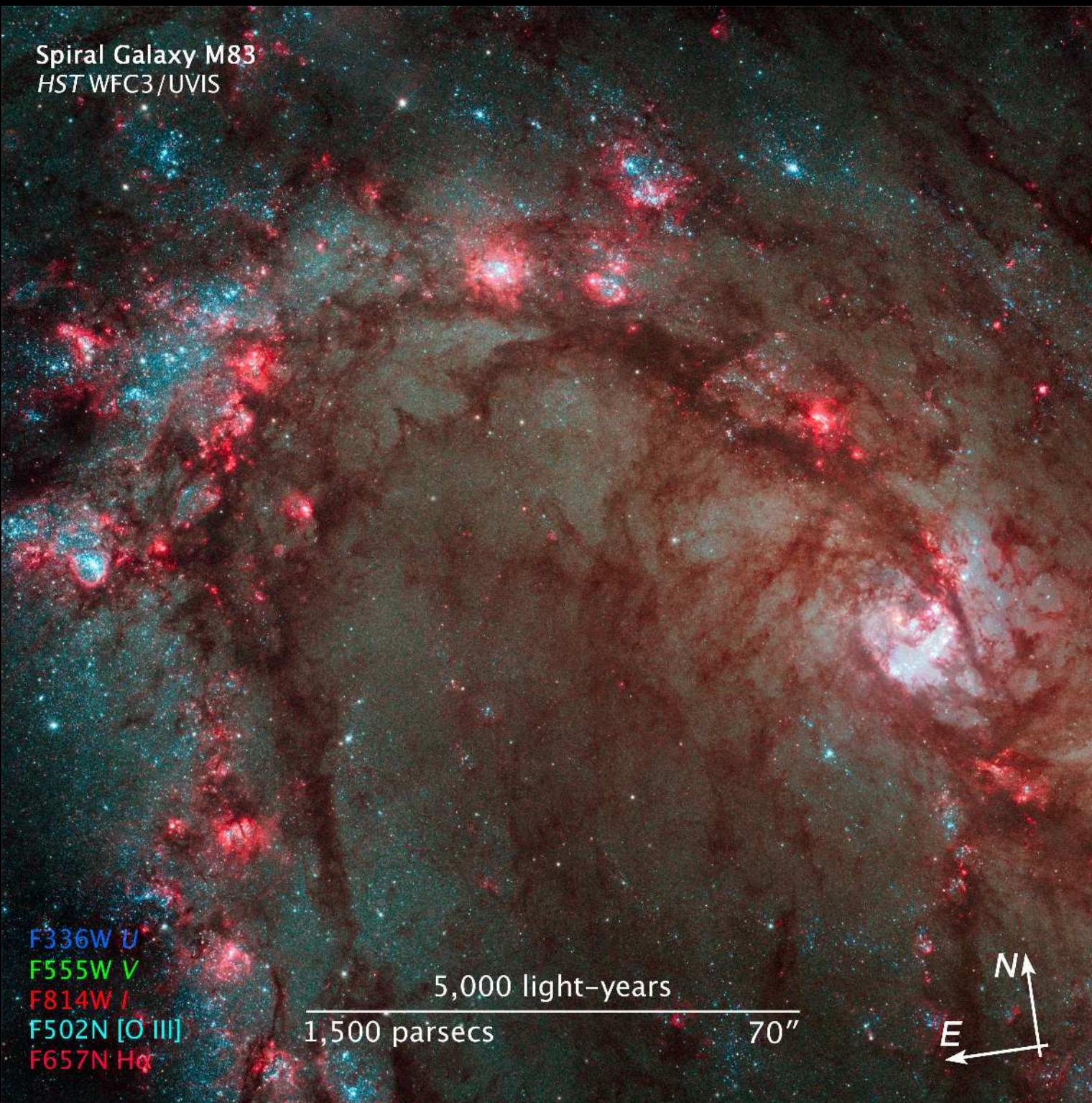
- QE  $\gtrsim 70\%$ ,  $1\text{k} \times 1\text{k}$  array of  $0\farcs13$  pixel, FOV  $\simeq 2\farcs25 \times 2\farcs25$ .

$\Rightarrow$  WFC3 opened major new parameter space for astrophysics in 2009:

WFC3 filters designed for star-formation and galaxy assembly at  $z \simeq 1\text{--}8$ .

- HST WFC3 and its IR channel a critical pathfinder for JWST science.

## (2) Measuring (Nearby) Galaxy Assembly and Supermassive Black-Hole Growth.

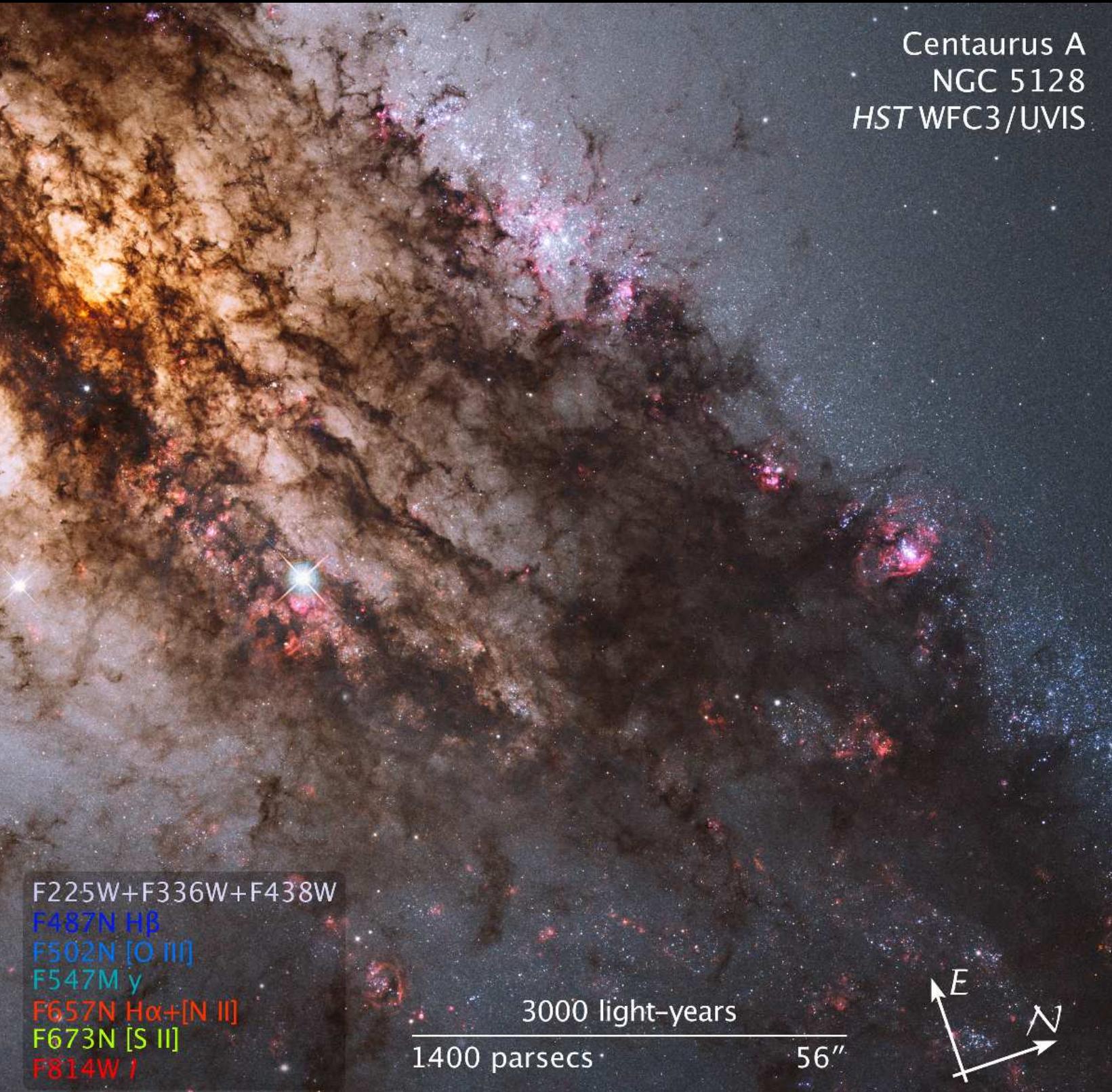


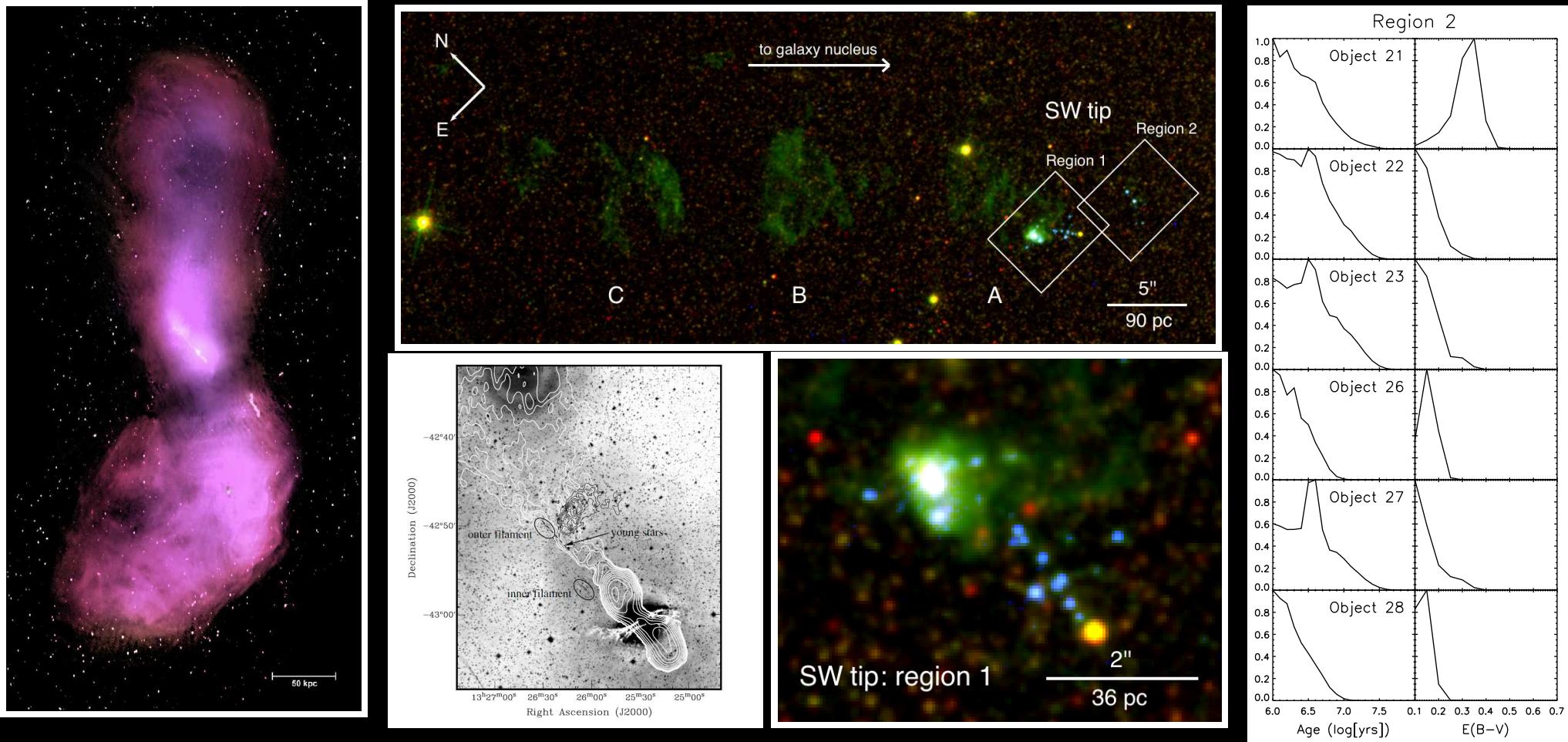


*"For God's sake, Edwards. Put the laser pointer away."*

The danger of having Quasar-like devices too close to home ...

Centaurus A  
NGC 5128  
*HST WFC3/UVIS*

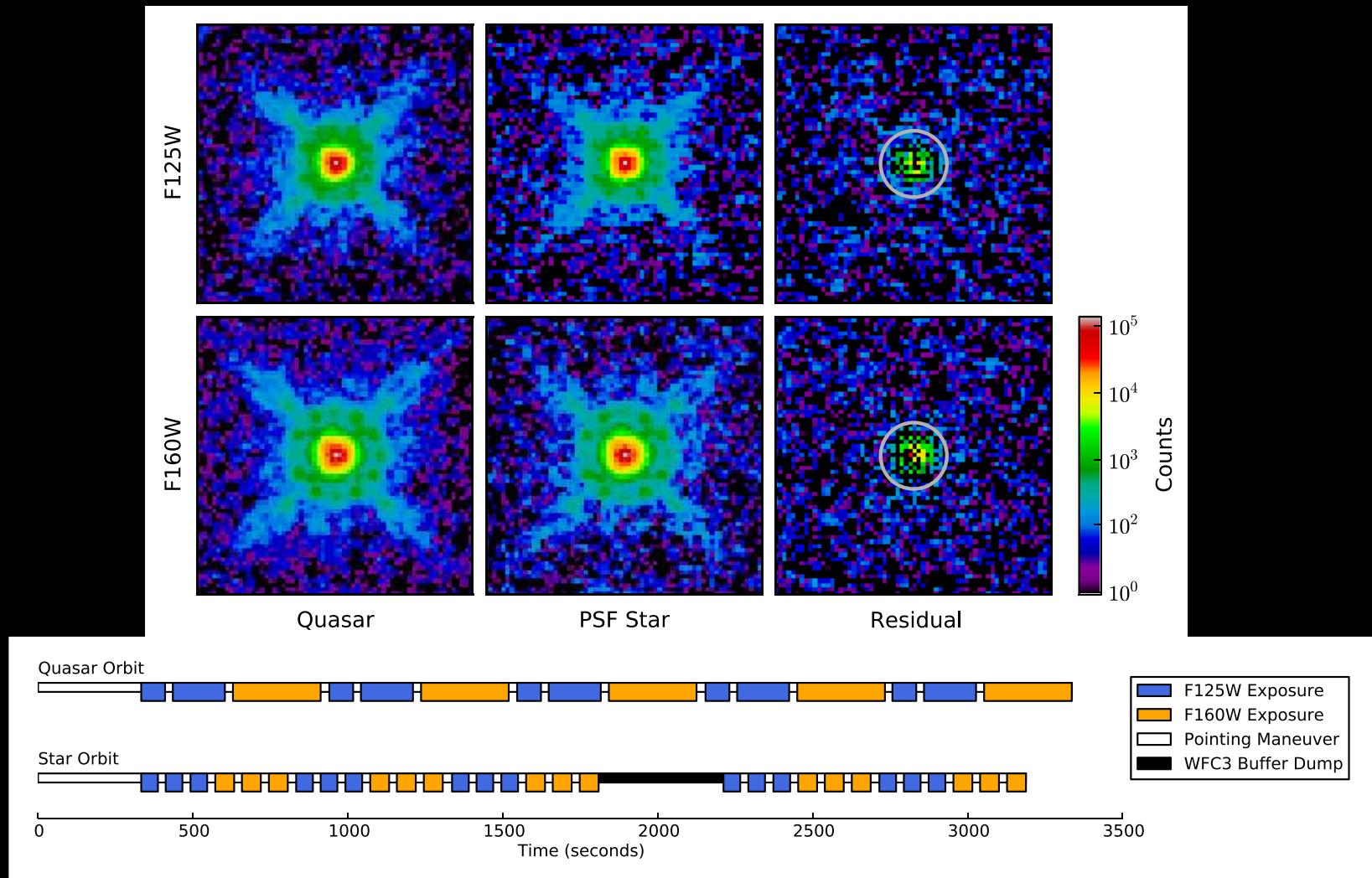




- [Left] CSIRO/ATNF 1.4 GHz image of Cen A (Feain<sup>+</sup> 2009).  
 Fermi GeV source (Yang<sup>+</sup> 12); & Auger UHE Cosmic Rays (Abreu<sup>+</sup> 2010).
- [Middle] SF in Cent A jet's wake (Crockett<sup>+</sup> 2012, MNRAS, 421, 1602).
- [Right] Well determined ages for young ( $\sim 2$  Myr) stars near Cen A's jet.

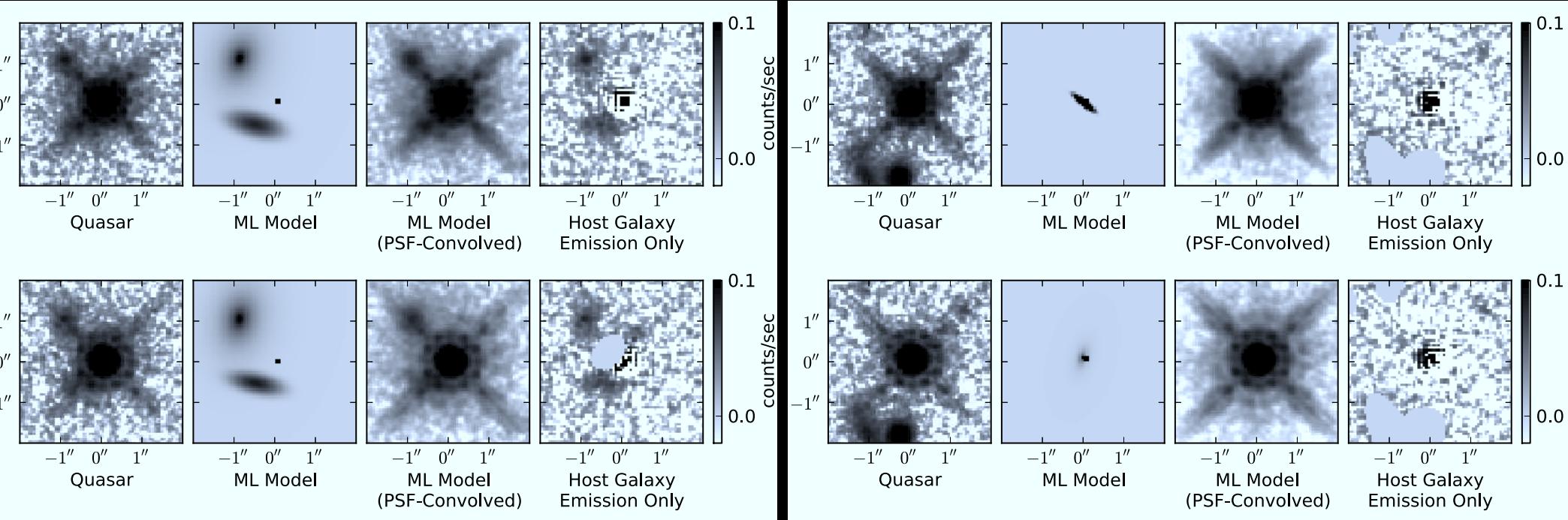
- JWST will trace older stellar pops and SF in much dustier environments.
- We must do all we can with HST in the UV–blue before JWST flies.

## (2) HST WFC3 observations of QSO host galaxies at $z \simeq 6$ (age $\lesssim 1$ Gyr)



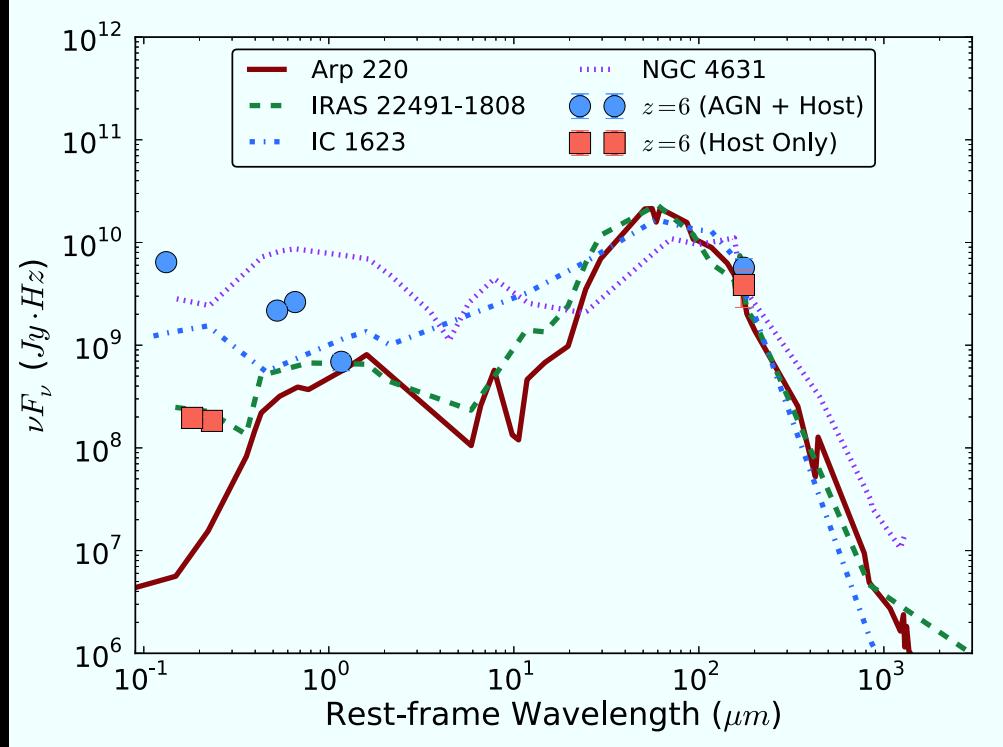
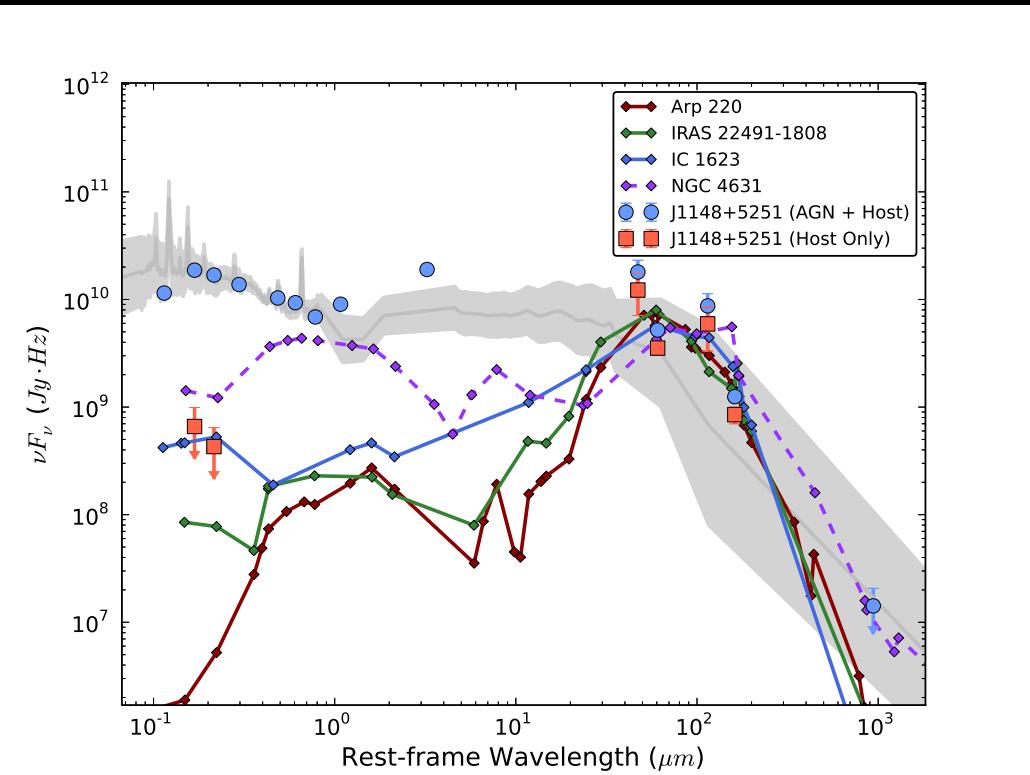
- Careful contemporaneous orbital PSF-star subtraction: Removes most of “OTA spacecraft breathing” effects (Mechtley ea 2012, ApJL, 756, L38).
- PSF-star ( $AB \simeq 15$  mag) subtracts  $z=6.42$  QSO ( $AB \simeq 18.5$ ) nearly to the noise limit: NO host galaxy detected  $100 \times$  fainter ( $AB \gtrsim 23.5$  at  $r \gtrsim 0\farcs3$ ).

## (2) WFC3: First detection of one QSO Host Galaxy at $z \simeq 6$ (Giant merger?)



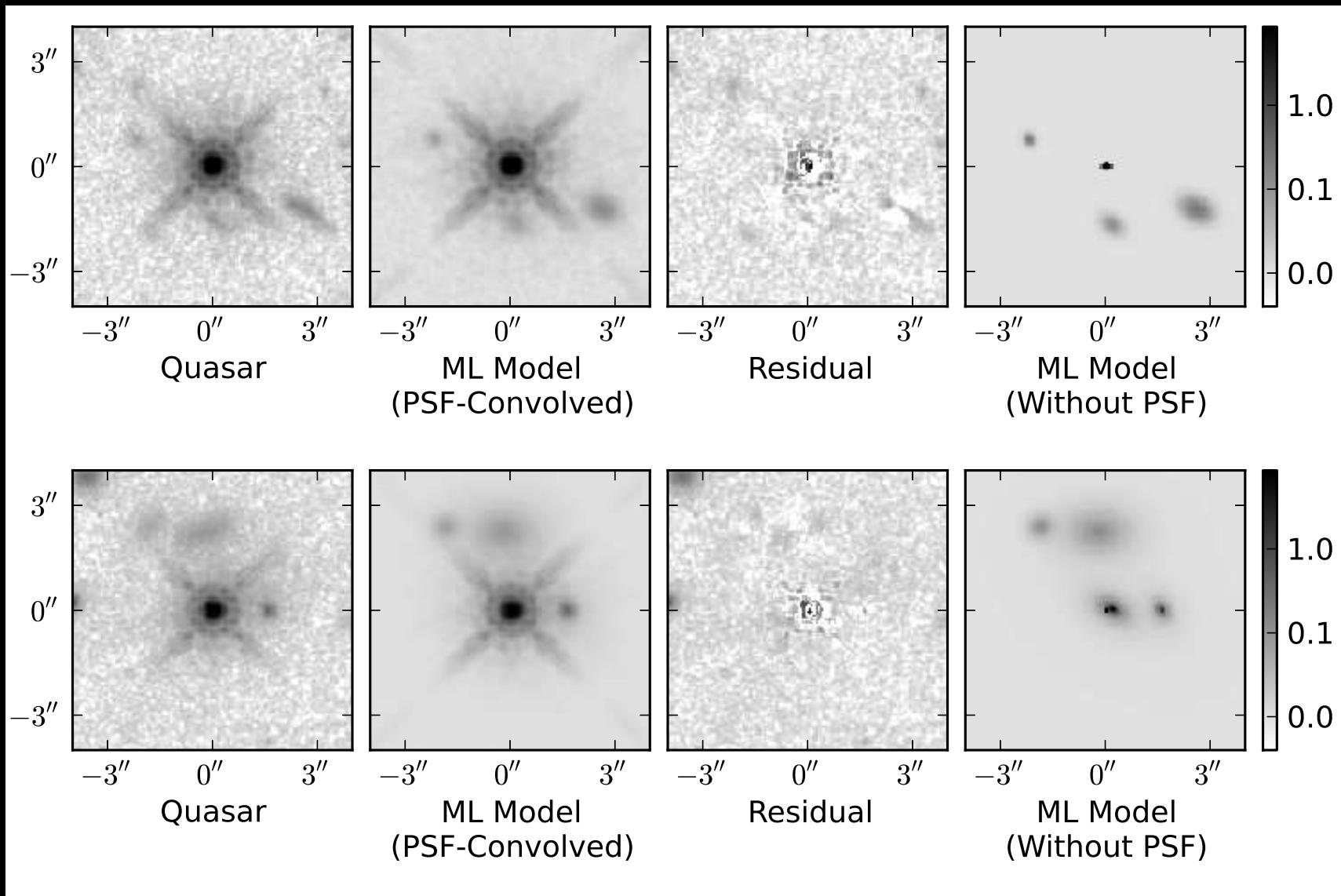
- Monte Carlo Markov-Chain of observed PSF-star + Sersic ML light-profile. Gemini AO images to pre-select PSF stars (Mechtley<sup>+</sup> 2014).
- First solid detection out of four  $z \simeq 6$  QSOs [2 more to be observed].
- One  $z \simeq 6$  QSO host galaxy: Giant merger morphology + tidal structure??
- Same J+H structure! Blue UV-SED colors:  $(J-H) \simeq 0.19$ , constrains dust.
  - IRAS starburst-like SED from rest-frame UV–far-IR,  $A_{FUV} \sim 1$  mag.
  - $M_{AB}^{host}(z \simeq 6) \lesssim -23.0$  mag, i.e.,  $\sim 2$  mag brighter than  $L^*(z \simeq 6)$ !
- $\Rightarrow z \simeq 6$  QSO duty cycle  $\lesssim 10^{-2}$  ( $\lesssim 10$  Myrs); 1/4 QSO's close to Magorrian.

## (2) HST WFC3 observations of dusty QSO host galaxies at $z \simeq 6$



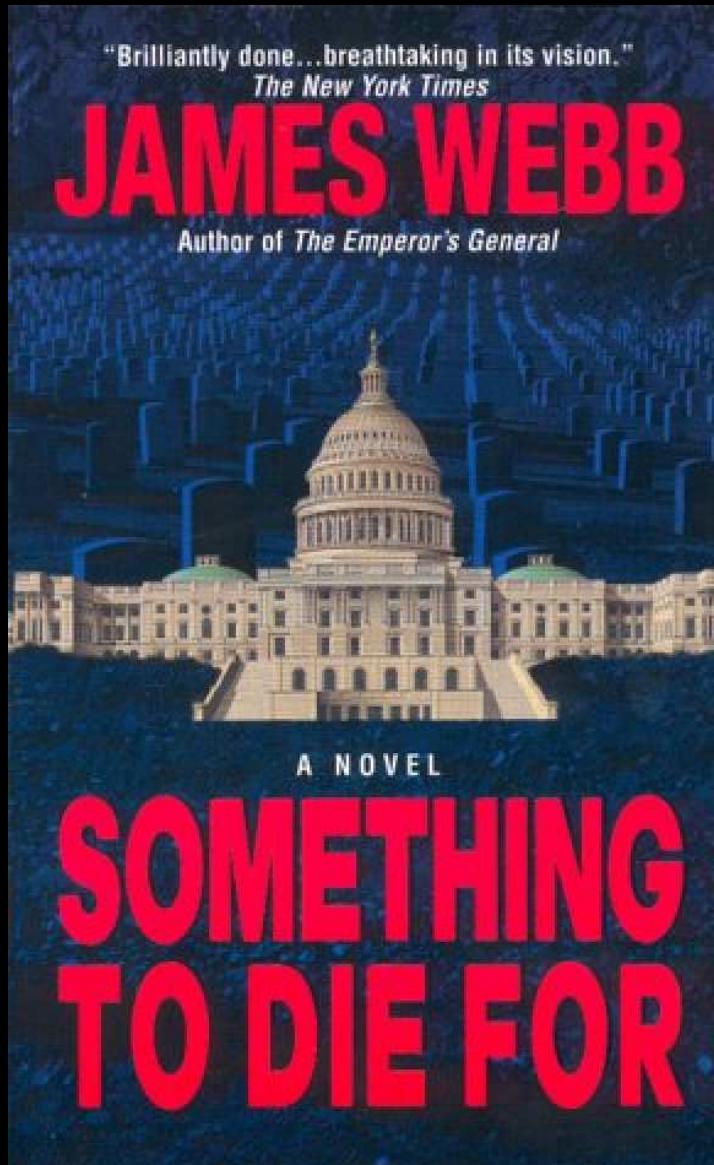
- Blue dots:  $z \simeq 6$  QSO SED, Grey: Average radio-quiet SDSS QSO spectrum at  $z \gtrsim 1$  (normalized at  $0.5\mu\text{m}$ ). Red:  $z \simeq 6$  host galaxy (WFC3+submm).
- Nearby fiducial galaxies (starburst ages  $\lesssim 1$  Gyr) normalized at  $100\mu\text{m}$ :  
 [LEFT] Rules out  $z = 6.42$  spiral or bluer host galaxy SEDs for 1148+5251.  
 (U)LIRGs & Arp 220s permitted (Mechtley et al. 2012, ApJL, 756, L38).  
 [RIGHT] Detected QSO host has IRAS starburst-like SED from rest-frame UV–far-IR,  $A_{FUV}(\text{host}) \sim 1$  mag (Mechtley 2013 PhD; et al. 2014).
- JWST Coronagraphs can do this  $10\text{--}100 \times$  fainter (& for  $z \lesssim 20$ ,  $\lambda \lesssim 28\mu\text{m}$ ).

## (2) WFC3 observations of QSO host galaxies at $z \approx 2$ (evidence for mergers?)



- Monte Carlo Markov-Chain runs of observed PSF-star + Sersic ML light-profile models: merging neighbors (some with tidal tails?; Mechtley, Jahnke, Koekemoer, Windhorst et al. 2014).
- JWST Coronagraphs can do this 10–100× fainter (& for  $z \lesssim 20$ ,  $\lambda \lesssim 28\mu\text{m}$ ).

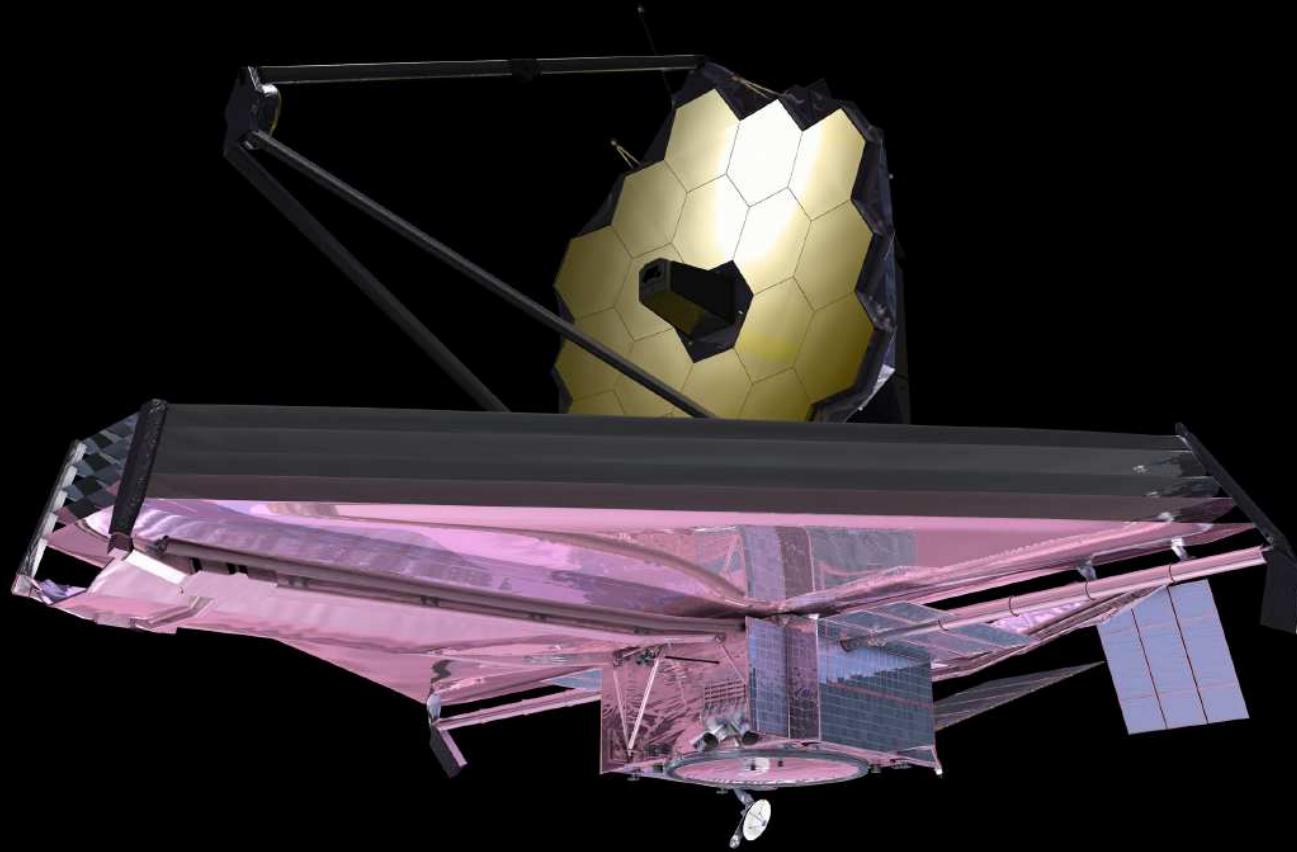
### (3) Brief Update of the James Webb Space Telescope (JWST).



To be used by students & scientists after 2018 ... It'll be worth it.

(RIGHT) Life-size JWST prototype on the Capitol Mall, May 2007.

### (3) Brief Update of the James Webb Space Telescope



- A fully deployable 6.5 meter ( $25 \text{ m}^2$ ) segmented IR telescope for imaging and spectroscopy at  $0.6\text{--}28 \mu\text{m}$  wavelength, to be launched in Fall 2018.
- Nested array of sun-shields to keep its ambient temperature at 40 K, allowing faint imaging (AB=31.5 mag) and spectroscopy.

# THE JAMES WEBB SPACE TELESCOPE

## JWST LAUNCH

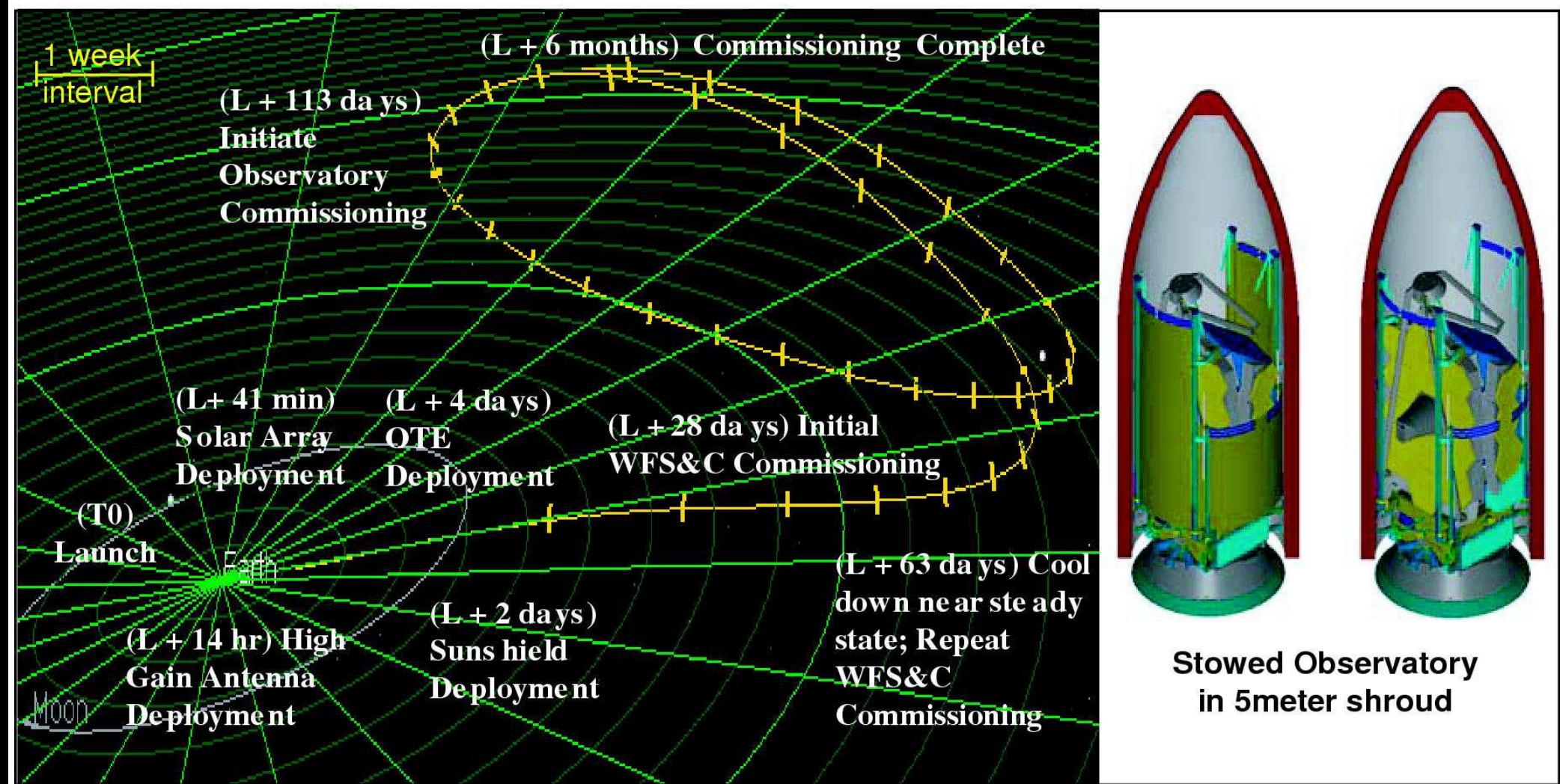
- LAUNCH VEHICLE IS AN ARIANE 5 ROCKET, SUPPLIED BY ESA
- SITE WILL BE THE ARIANESPACE'S ELA-3 LAUNCH COMPLEX NEAR KOUROU, FRENCH GUIANA



ARIANESPACE - ESA - NASA

- The JWST launch weight will be  $\lesssim 6500$  kg, and it will be launched to L2 with an ESA Ariane-V launch vehicle from Kourou in French Guiana.

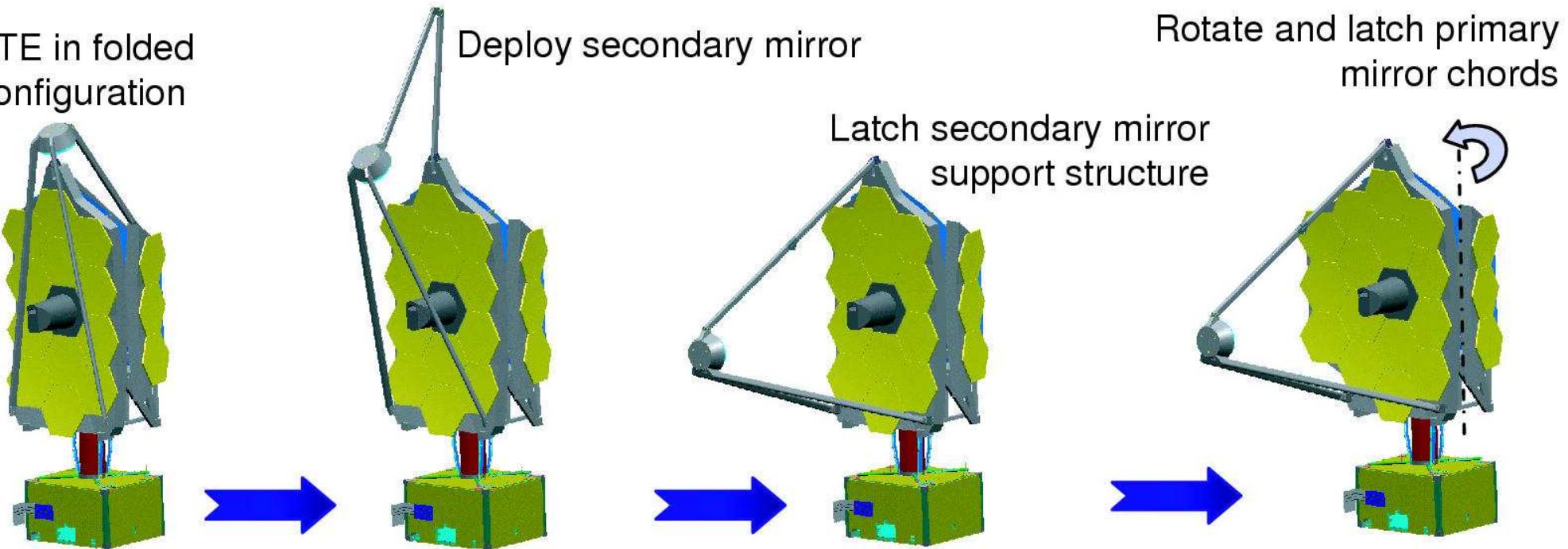
### (3a) How will JWST travel to its L2 orbit?



- After launch in 2018 with an ESA Ariane-V, JWST will orbit around the Earth–Sun Lagrange point L2, 1.5 million km from Earth.
- JWST can cover the whole sky in segments that move along with the Earth, observe  $\gtrsim 70\%$  of the time, and send data back to Earth every day.

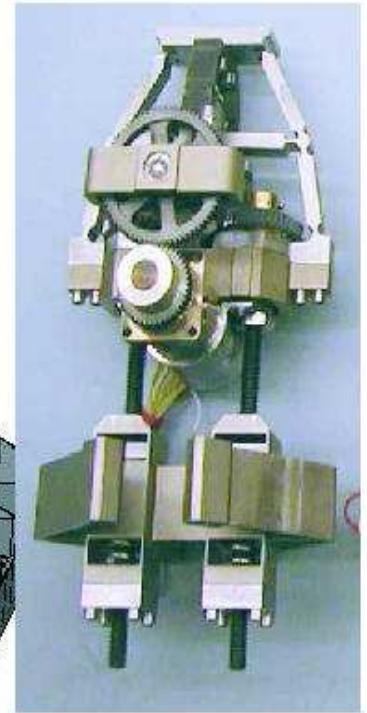
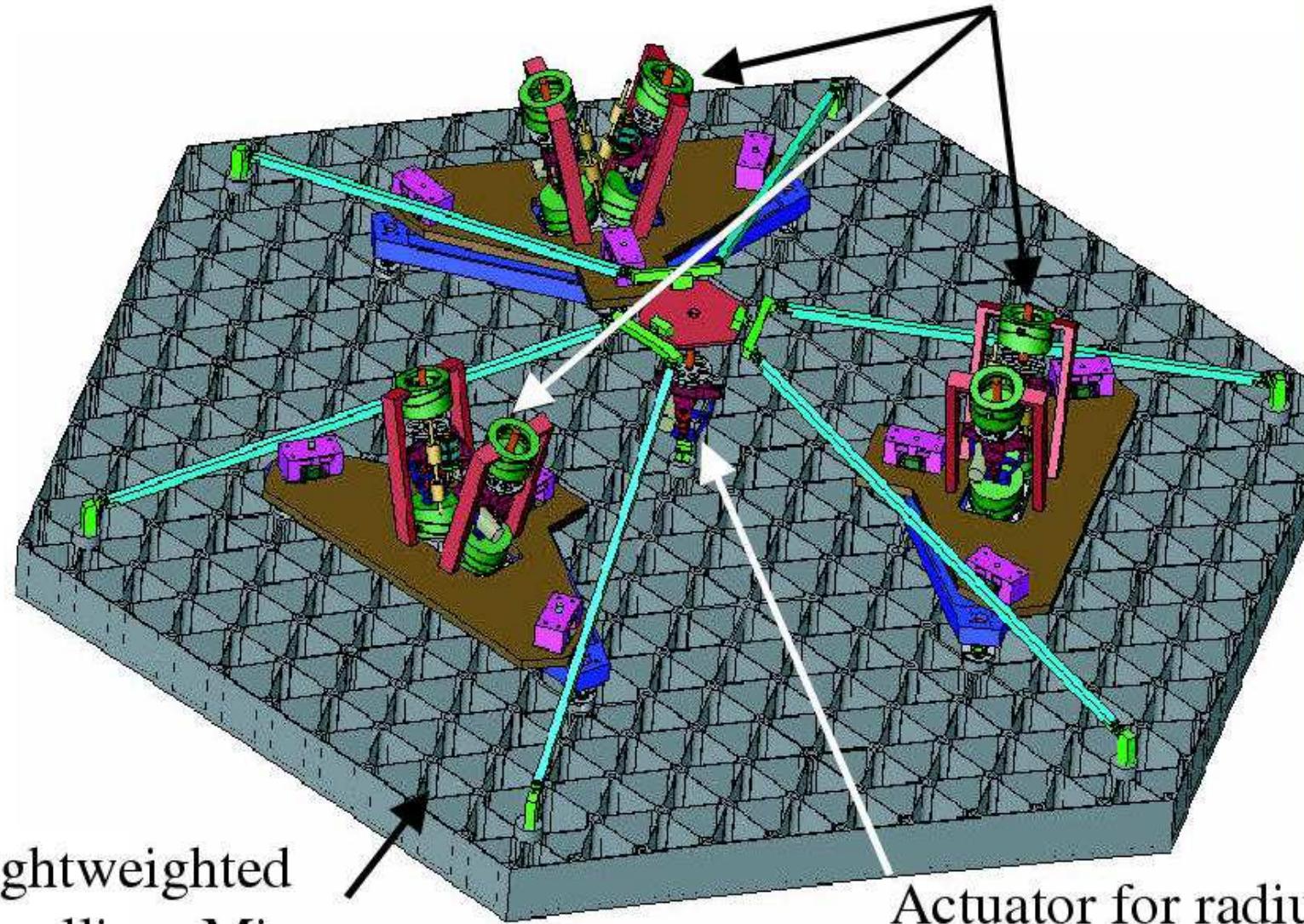
- (3b) How will JWST be automatically deployed?

OTE in folded configuration



- During its two month journey to L2, JWST will be automatically deployed, its instruments will be cooled, and be inserted into an L2 orbit.
- The entire JWST deployment sequence will be tested several times on the ground — but only in 1-G: Component and system tests in Houston.
- Component fabrication, testing, & system integration is on schedule: 18 out of 18 flight mirrors completely done, and meet the 40K specifications.

## Actuators for 6 degrees of freedom rigid body motion



Actuator development unit

Lightweighted  
Beryllium Mirror

Actuator for radius  
of curvature adjustment

Active mirror segment support through “hexapods”, similar to Keck.

Redundant & doubly-redundant mechanisms, quite forgiving against failures.

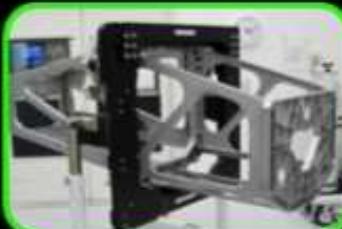


# JWST Hardware Status

Primary Mirror Segment



Aft Optics System



PM Flight Backplane



Tertiary Mirror



Fine Steering Mirror

Secondary Mirror Pathfinder Strut



Secondary Mirror Hexapod



Secondary Mirror

ISIM Flight Bench



Membrane Mgmt



Pathfinder Membrane



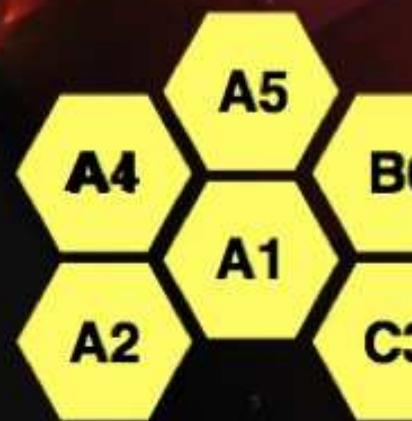
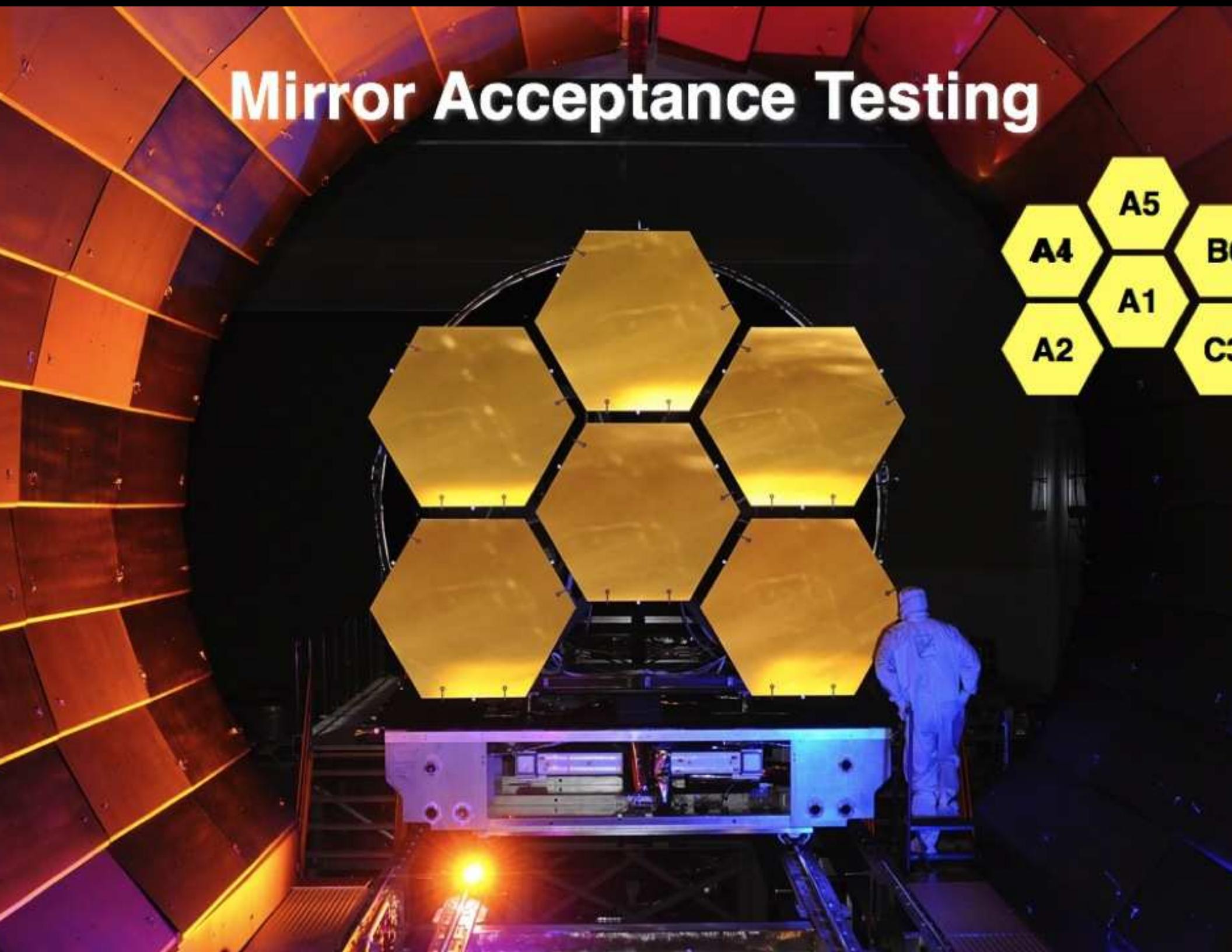
Spacecraft computer Test Unit



Mid-boom Test

80% of launch mass designed<sup>3</sup> and built as of Jan. 2014.

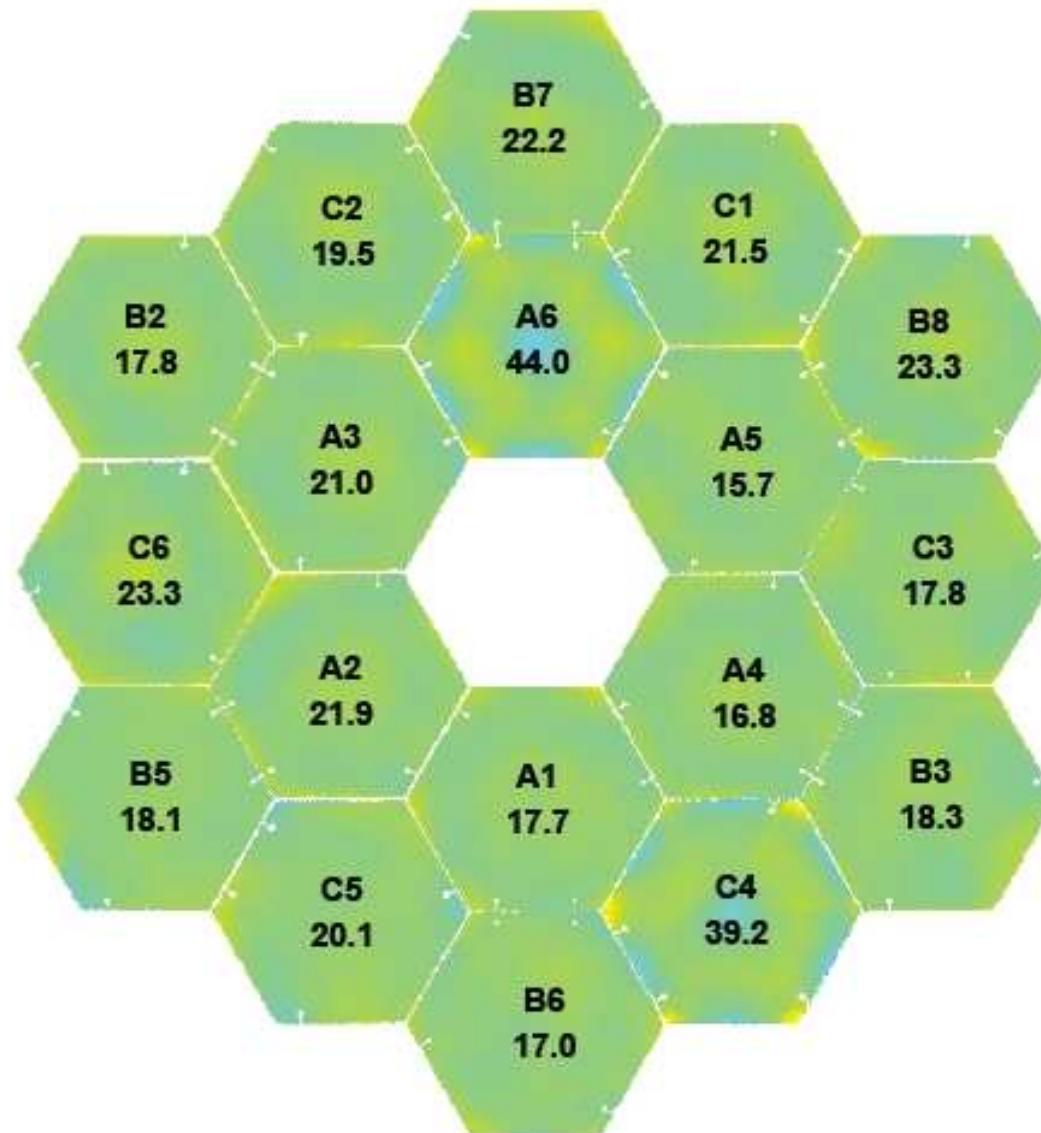
# Mirror Acceptance Testing





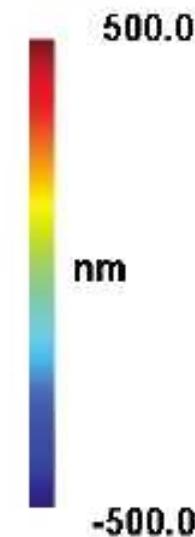


# Primary Mirror Composite



RMS:  
**23.2 nm**

PV:  
**515.5 nm**

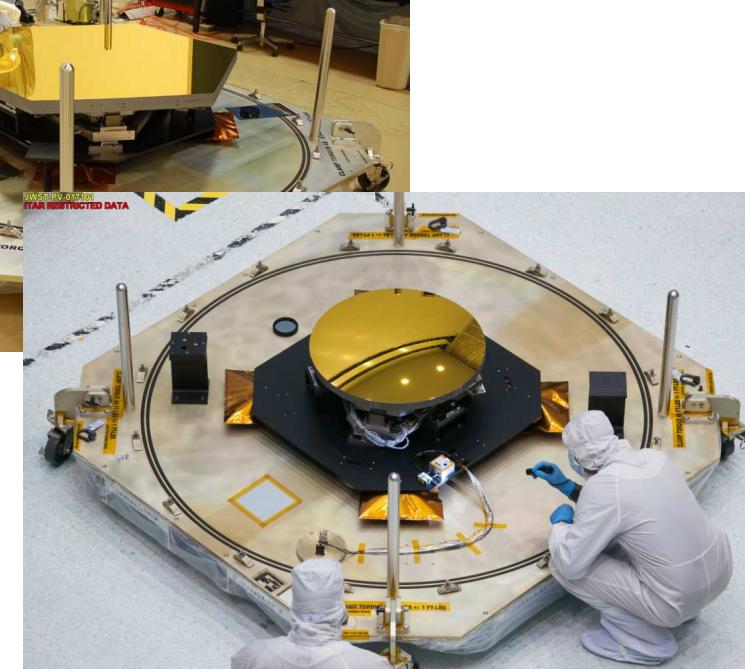
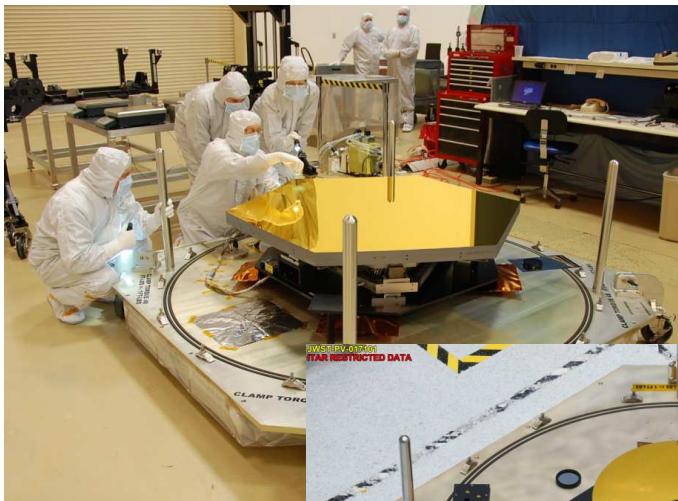




# Mirror Status



- **15 flight primary mirrors and the flight secondary mirror are at GSFC in storage**
  - All spares were at GSFC in storage (SM spares, 3 PMSA spares)
    - 2 EDU mirrors sent back to Ball for gear motor rework
  - All flight gear motor refurbishment is complete
  - All flight mirrors will be at GSFC by end of year, needed in 2015



Jan 2014: All 18 flight mirrors now delivered to NASA GSFC (MD).



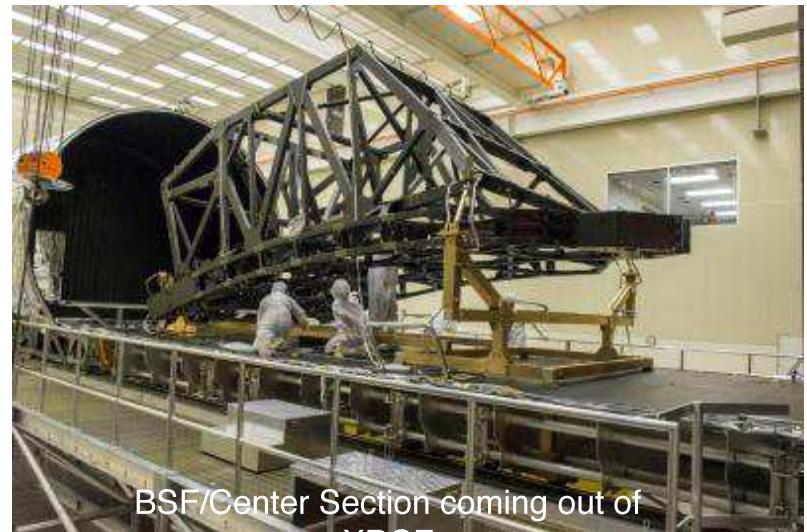
# Backplane Support Frame, Center Section, & Wings



- Center Section is complete
- Wings and cryo cycling is complete
- BSF assembly is complete
- Integration of the BSF to Center Section Complete
  - Cryo Cycling at MSFC XRCF complete



BSF and Center Section



BSF/Center Section coming out of XRCF

Jan 2014: Flight back-plane ready to receive mirrors in 2014.

# Telescope Assembly Ground Support Equipment



Ambient Optical Alignment Stand  
GSFC



Hardware has been installed  
at GSFC approximately 8  
weeks ahead of schedule



March 2012 NAC  
Science Meeting



Landing a mirror onto  
backplane simulator



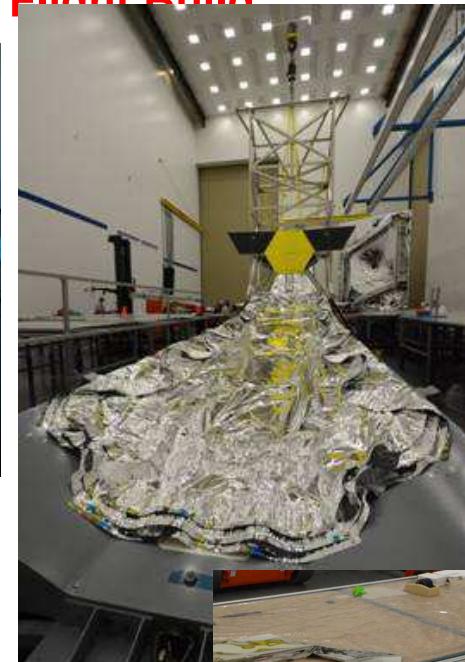
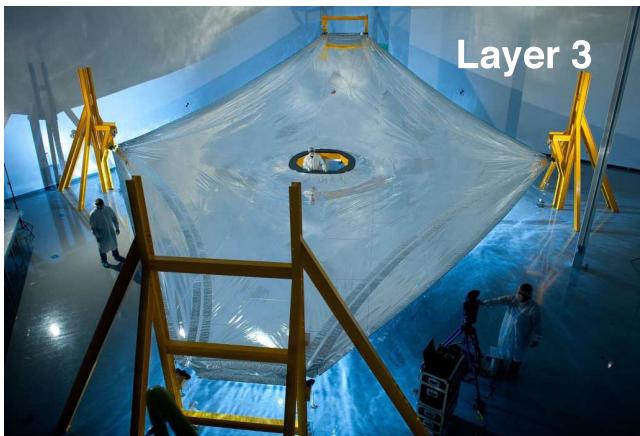
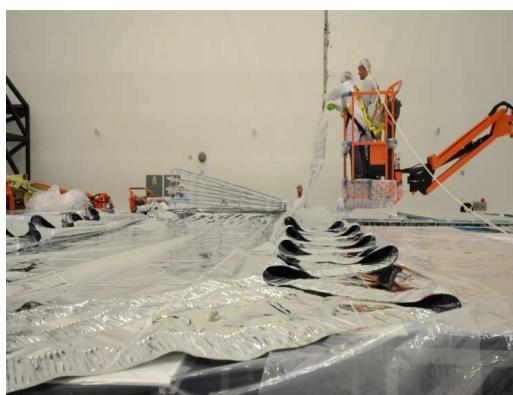
# Sunshield Template Membrane Work Completed

Templates Verify Design/Manufacturing Prior to Flight Build



- All Template Layers Completed
- Preparing for flight article manufacturing
- First two Flight Manufacturing Readiness Reviews Completed
- Membrane pull out test complete

Stringing Operations



Template  
Layers 3-5



Hole Tool Operations

Flight sunshield to be completed & tested by 2015 at Northrop (CA).

# JWST instrument update: US (UofA, JPL), ESA, & CSA.

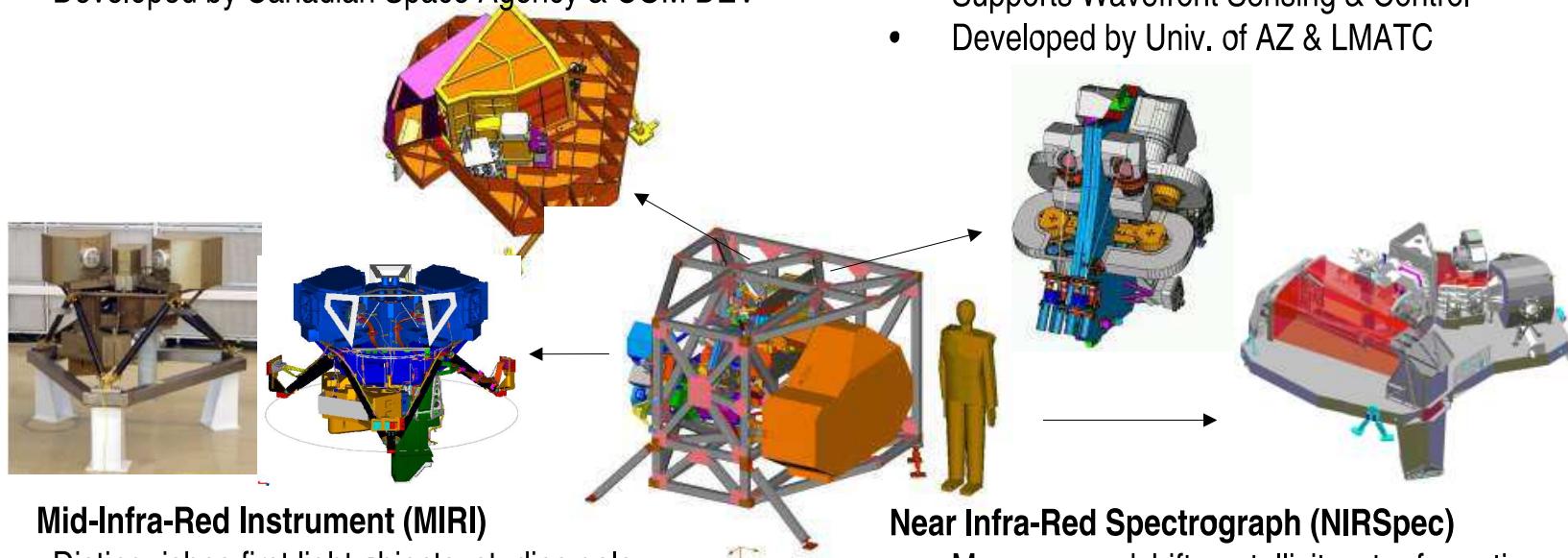


## Instrument Overview



### Fine Guidance Sensor (FGS)

- Ensures guide star availability with >95% probability at any point in the sky
- Includes Narrowband Imaging Tunable Filter
- Developed by Canadian Space Agency & COM DEV

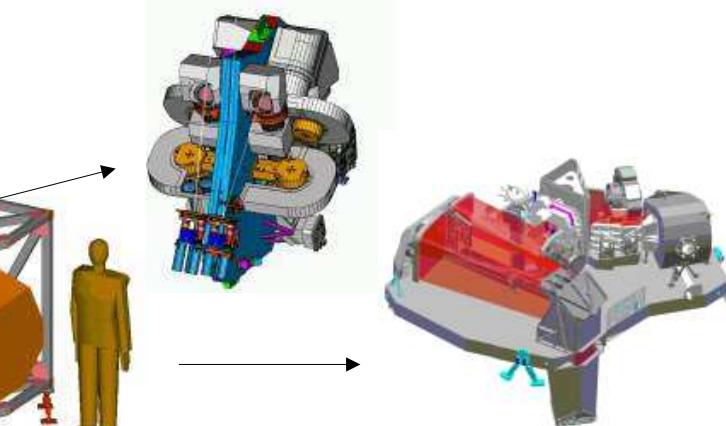


### Mid-Infra-Red Instrument (MIRI)

- Distinguishes first light objects; studies galaxy evolution; explores protostars & their environs
- Imaging and spectroscopy capability
- 5 to 27 microns
- Cooled to 7K by Cyro-cooler
- Combined European Consortium/JPL development

### Near Infra-Red Camera (NIRCam)

- Detects first light galaxies and observes galaxy assembly sequence
- 0.6 to 5 microns
- Supports Wavefront Sensing & Control
- Developed by Univ. of AZ & LMATC

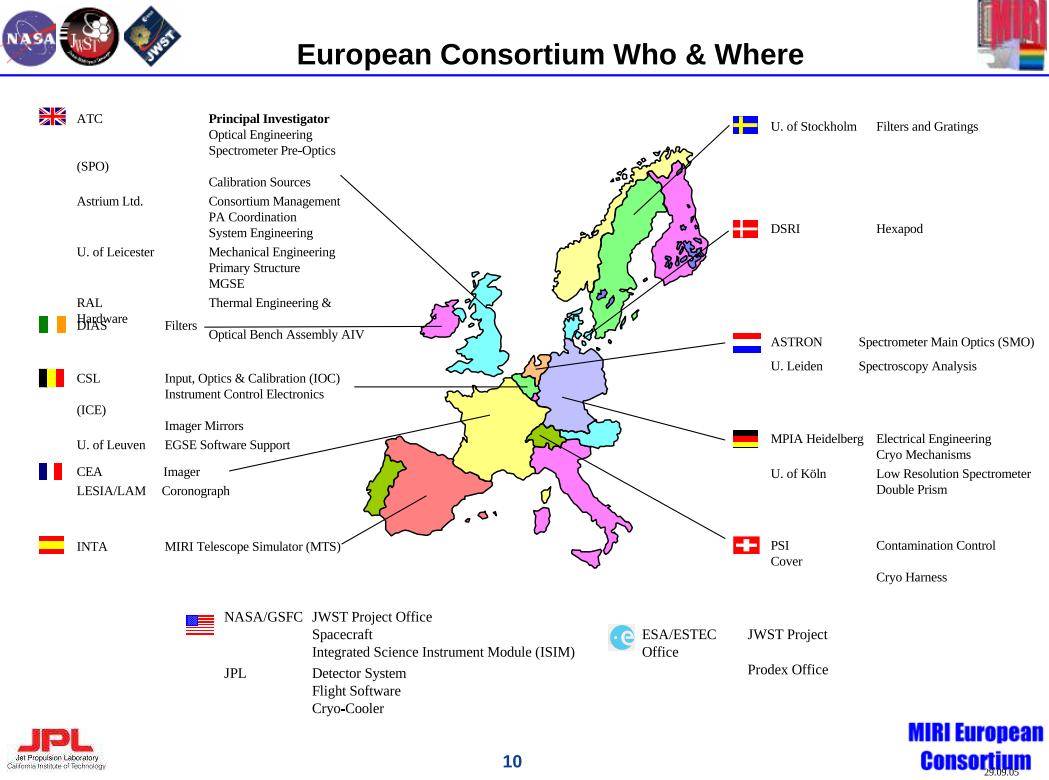
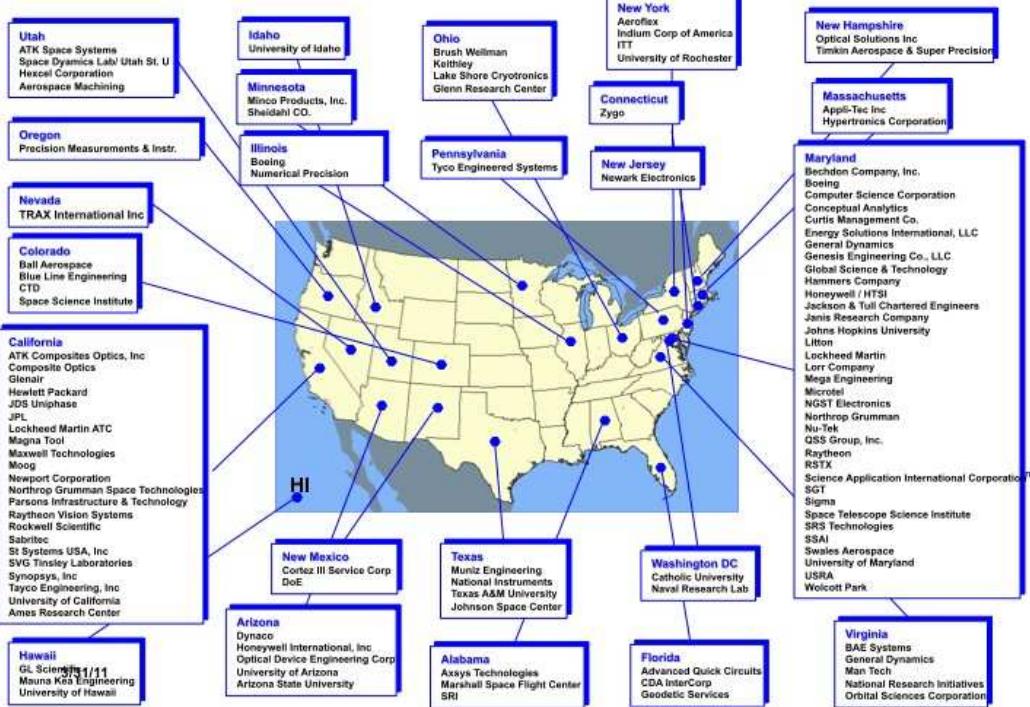


### Near Infra-Red Spectrograph (NIRSpec)

- Measures redshift, metallicity, star formation rate in first light galaxies
- 0.6 to 5 microns
- Simultaneous spectra of >100 objects
- Developed by ESA & EADS with NASA/GSFC Detector & Microshutter Subsystems

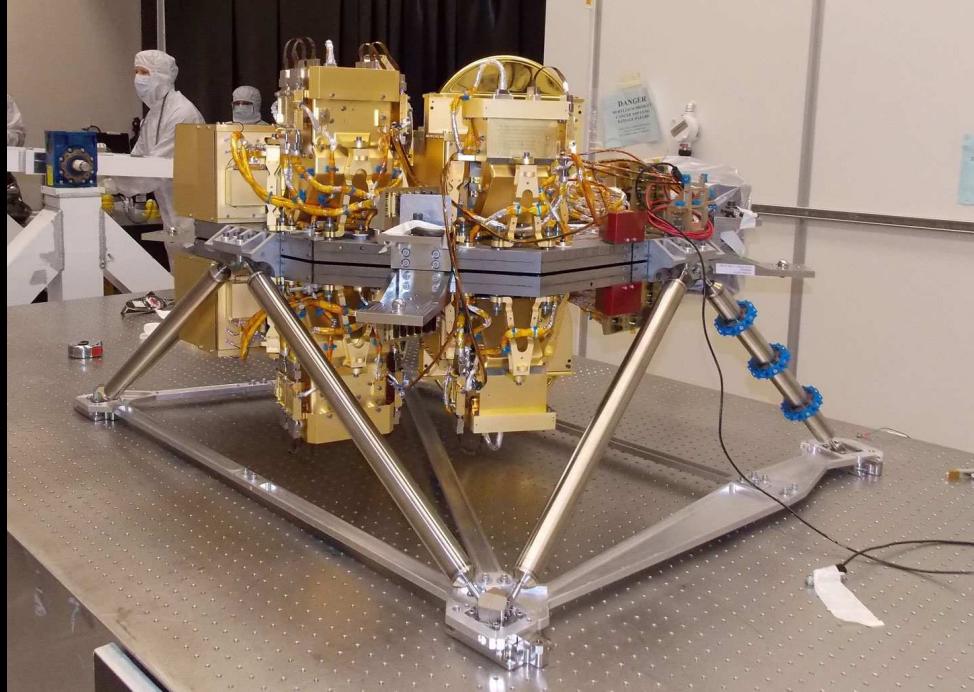
All delivered: MIRI 05/12; FGS 07/12; NIRCam 07/13, NIRSpec 9/13!

# JWST: A Product of the Nation



- JWST hardware made in 27 US States:  $\gtrsim 80\%$  of launch-mass finished.
- Ariane V Launch & NIRSpec provided by ESA; & MIRI by ESA & JPL.
- JWST Fine Guider Sensor + NIRISS provided by Canadian Space Agency.
- JWST NIRCam made by UofA and Lockheed.

Thank you, Italy, for your contributions to JWST!

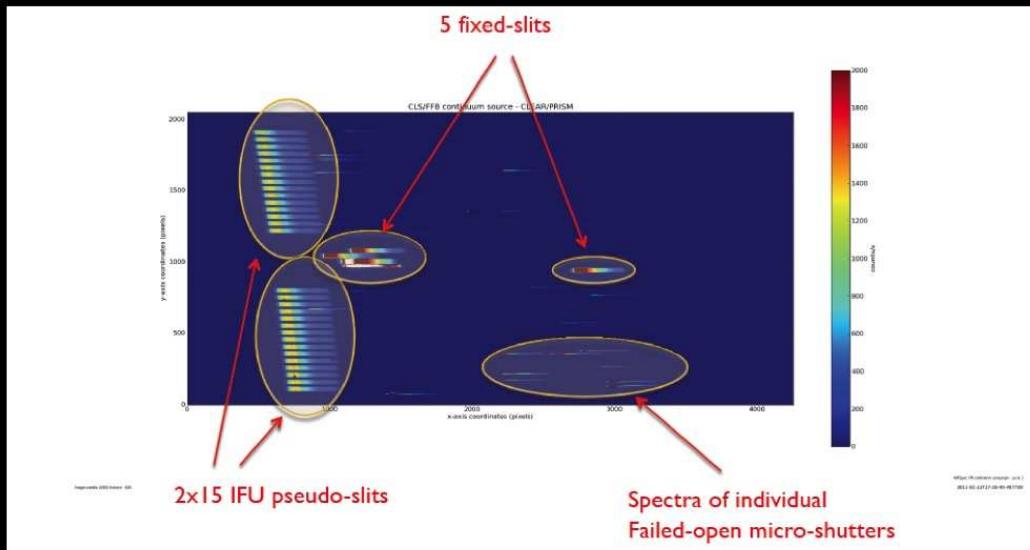


JWST's short-wavelength ( $0.6\text{--}5.0\mu\text{m}$ ) imagers:

- NIRCam — built by UofA (AZ) and Lockheed (CA).
- Fine Guidance Sensor (&  $1\text{--}5\mu\text{m}$  grisms) — built by CSA (Montreal).
- FGS includes very powerful low-res Near-IR grism spectrograph (NIRISS).
- FGS delivered to GSFC 07/12; NIRCam delivered July 28, 2013!



## Flight NIRSpec First Light

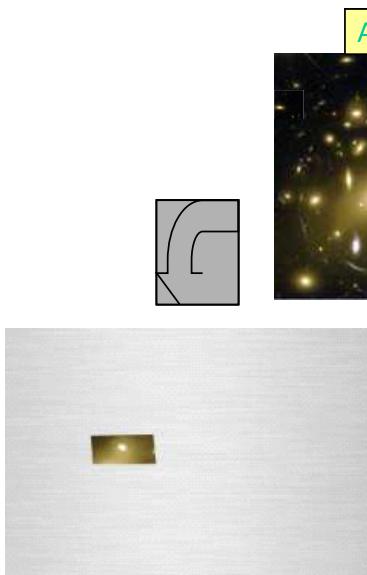


JWST's short-wavelength ( $0.6\text{--}5.0\mu\text{m}$ ) spectrograph:

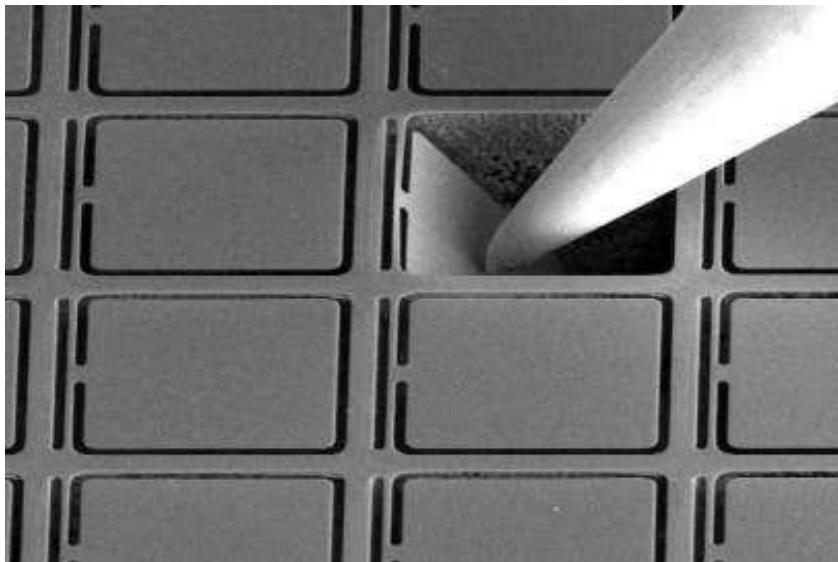
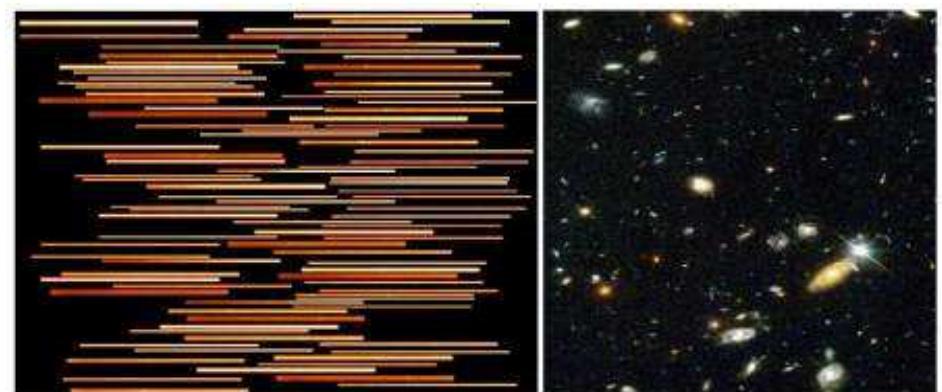
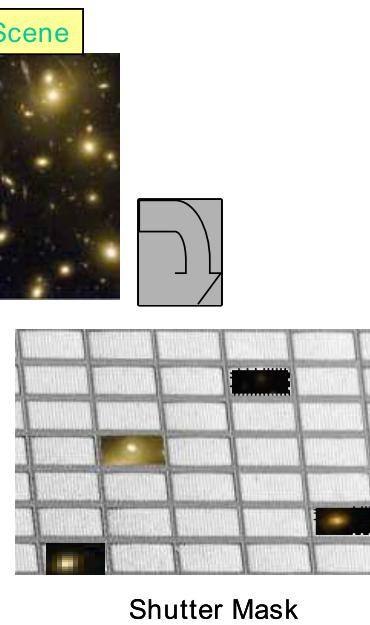
- NIRSpec — built by ESA/ESTEC and Astrium (Munich).
- Flight build completed and tested with First Light in Spring 2011.

NIRSpec delivered to NASA/GSFC in Sept. 2013.

# Micro Shutters

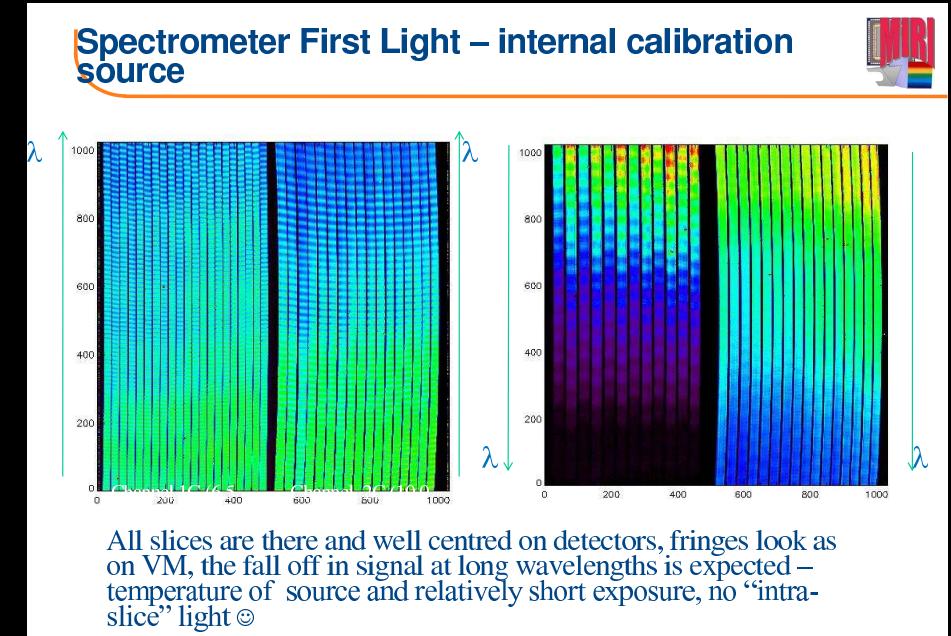


Metal Mask/Fixed Slit





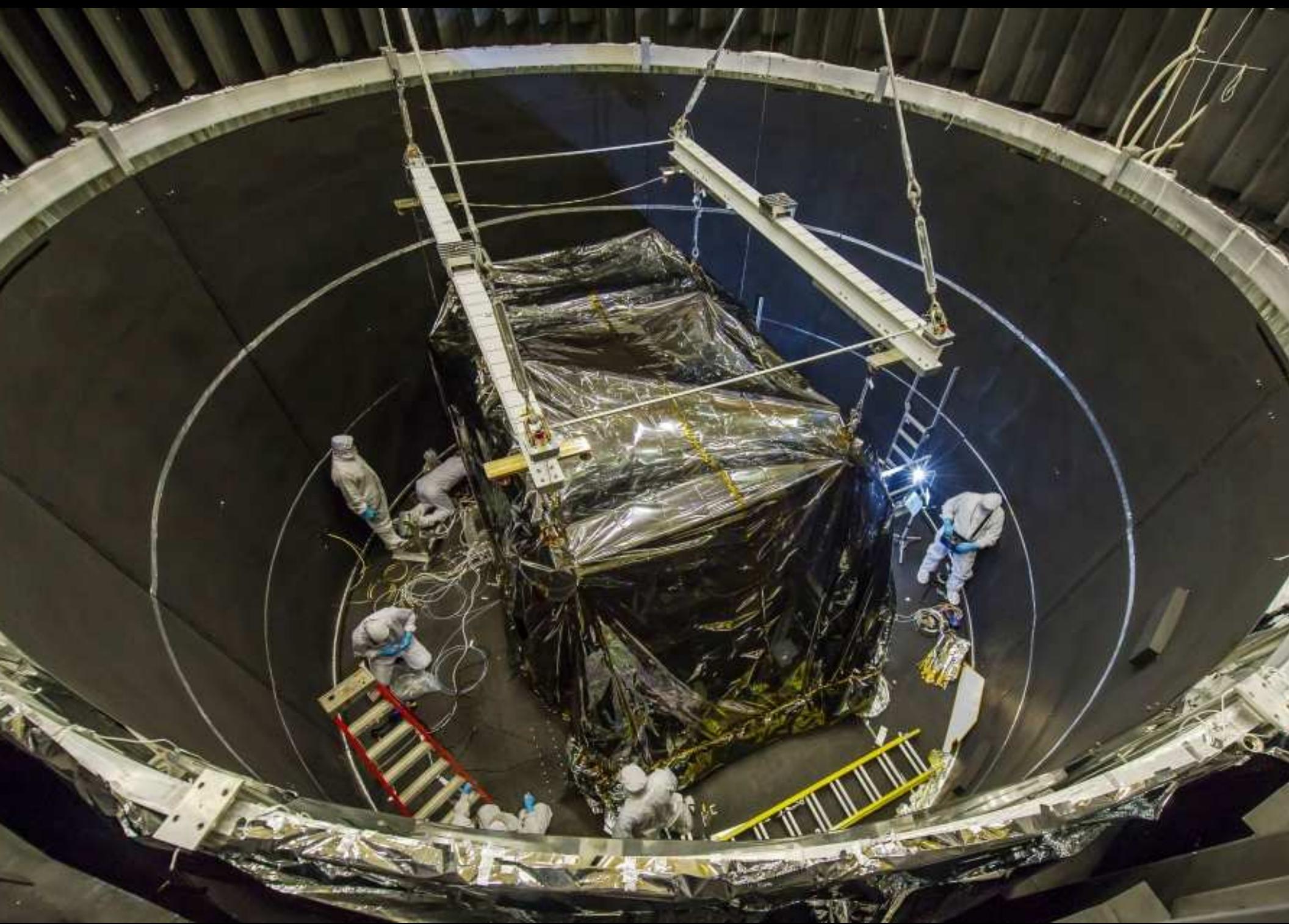
## Flight MIRI



JWST's mid-infrared ( $5\text{--}29\mu\text{m}$ ) camera and spectrograph:

- MIRI — built by ESA consortium of 10 ESA countries & NASA JPL.
- Flight build completed and tested with First Light in July 2011.

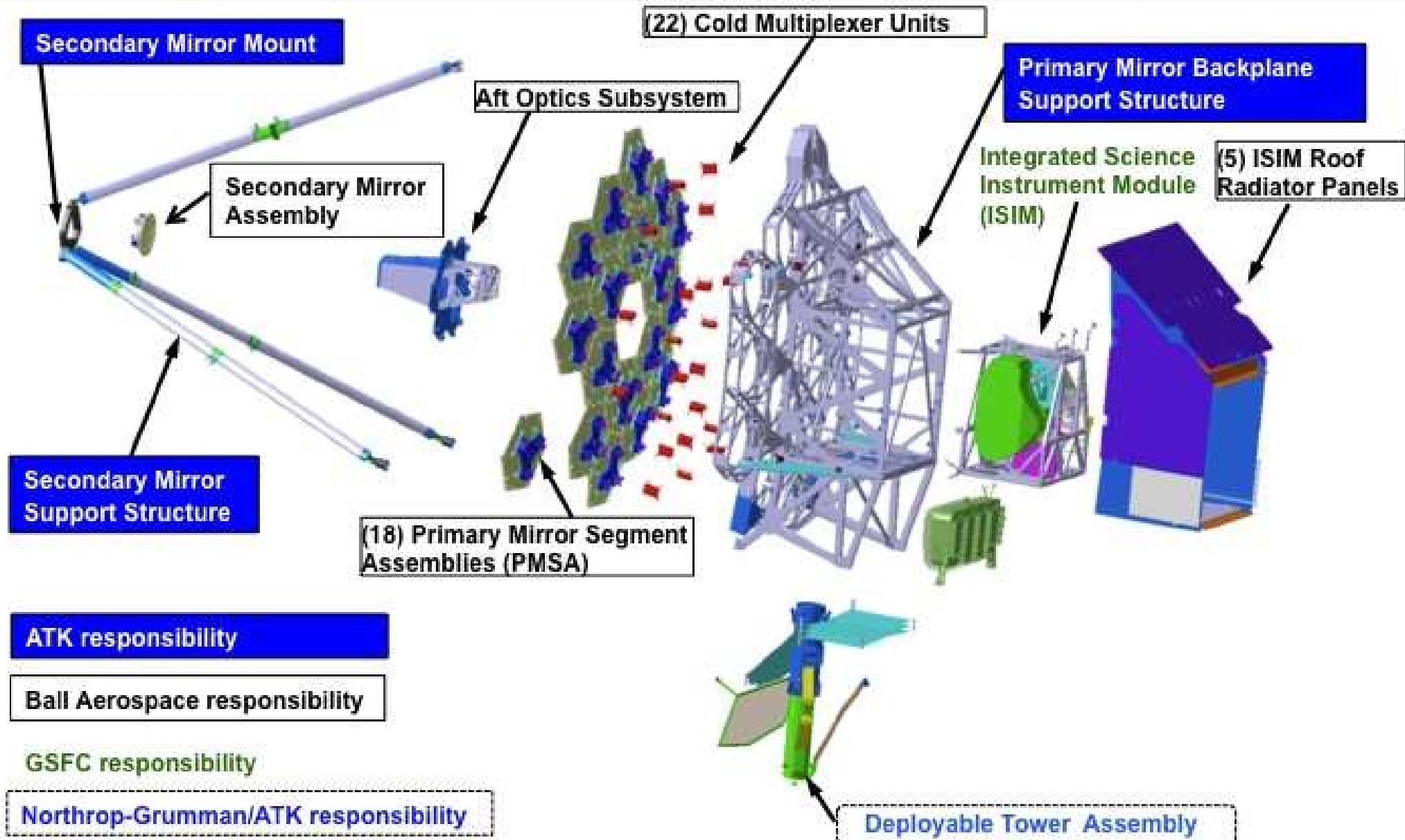
MIRI delivered to NASA/GSFC in May 2012.



Aug. 2013: Actual Flight ISIM (with MIRI and FGS) lowered into OSIM.

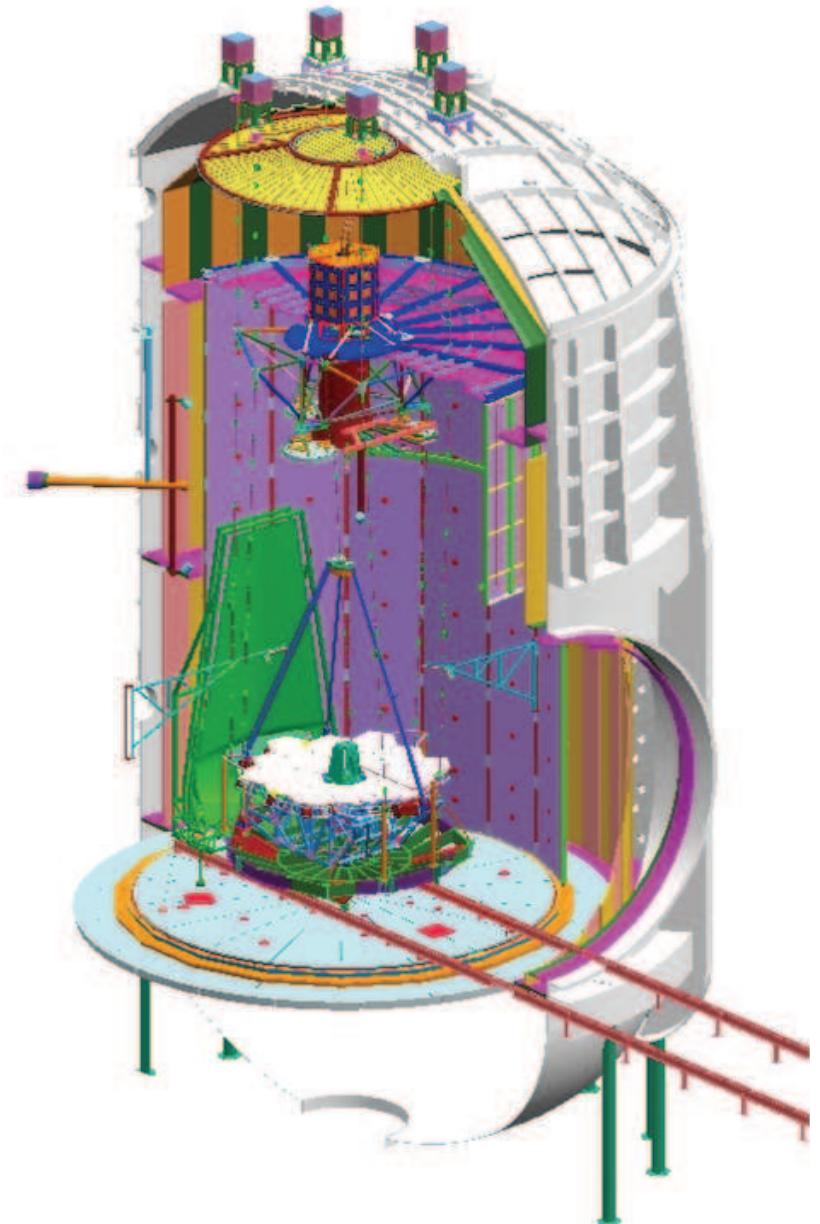
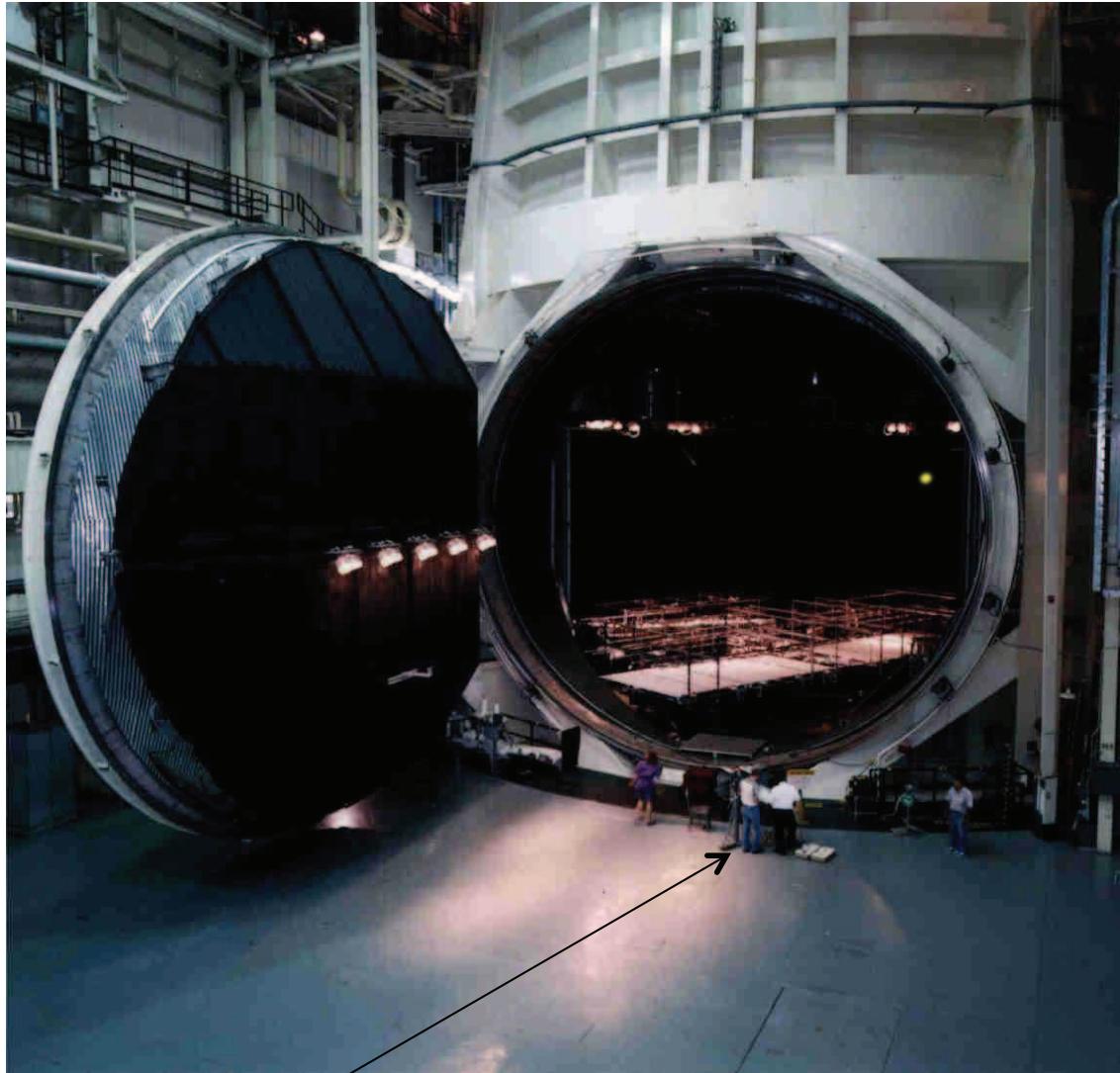


# TELESCOPE ARCHITECTURE





# OTE Testing – Chamber A at JSC



Notice people for scale

Will be the largest cryo vacuum test chamber in the world

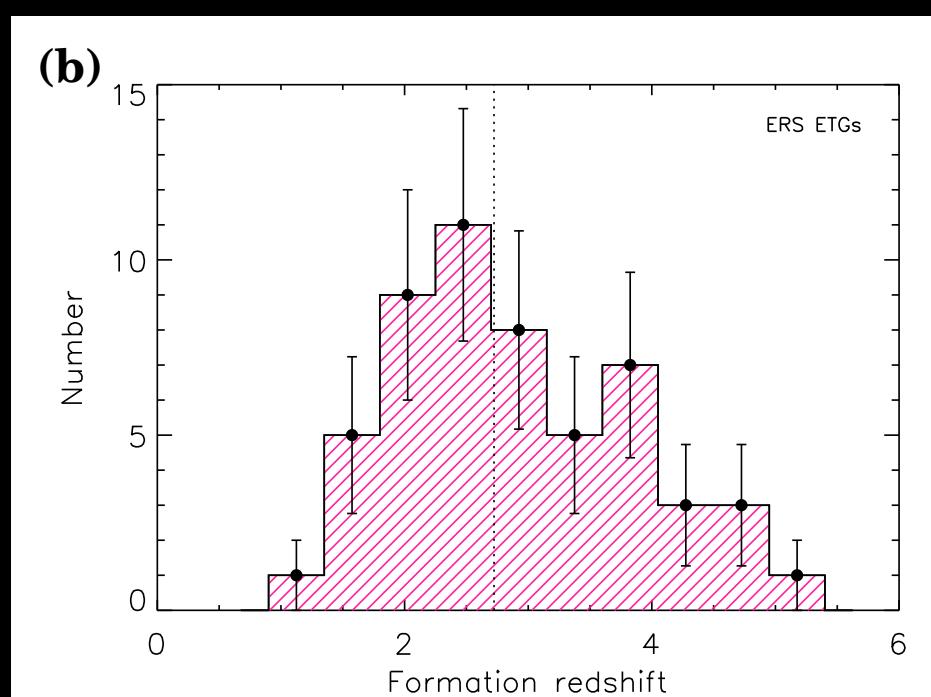
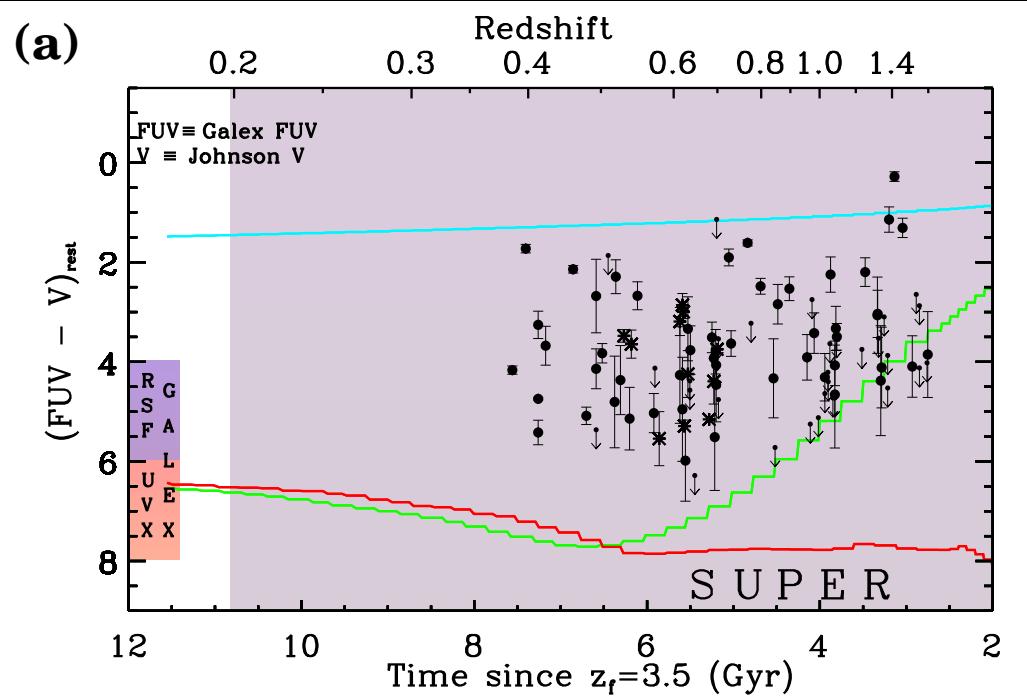
OTIS: Largest TV chamber in world: will test whole JWST in 2015–2016.

#### (4) How can JWST measure Galaxy Assembly and SMBH/AGN Growth?

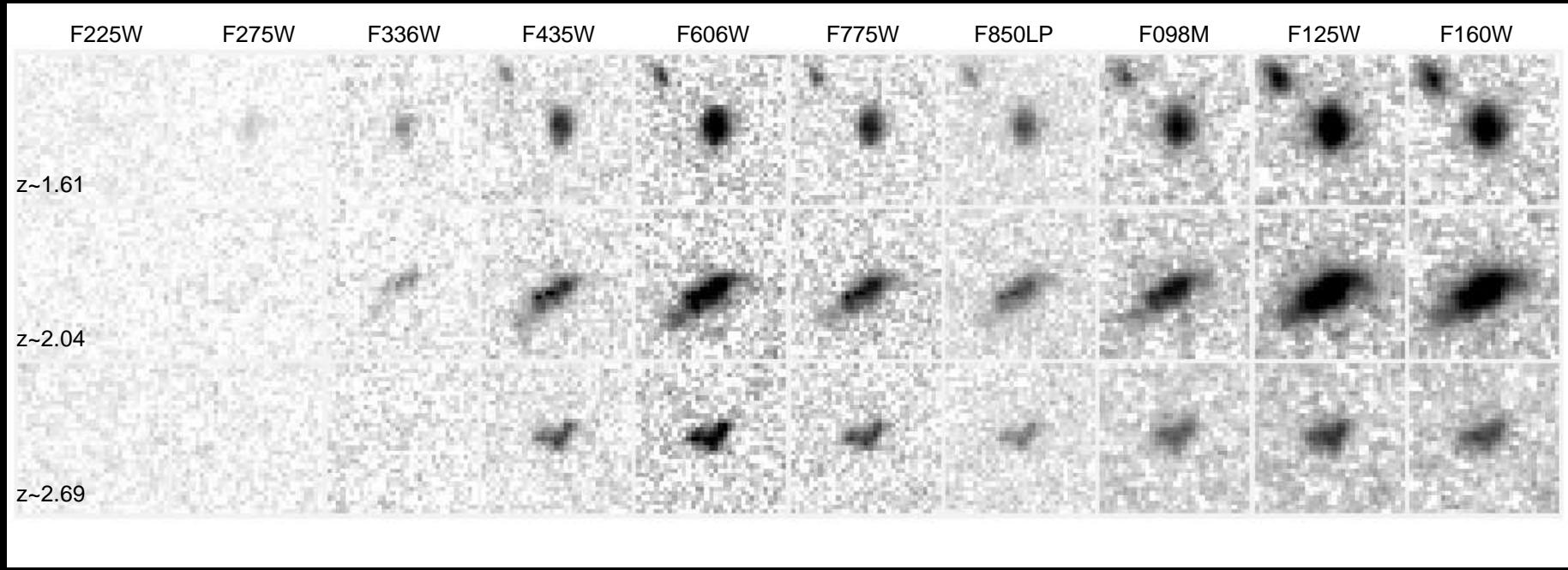


10 filters with HST/WFC3 & ACS reaching AB=26.5–27.0 mag (10- $\sigma$ ) over 40 arcmin<sup>2</sup> at 0.07–0.15" FWHM from 0.2–1.7 $\mu$ m (UVUBVizYJH). JWST adds 0.05–0.2" FWHM imaging to AB $\simeq$ 31.5 mag (1 nJy) at 1–5 $\mu$ m, and 0.2–1.2" FWHM at 5–29 $\mu$ m, tracing young+old SEDs & dust.

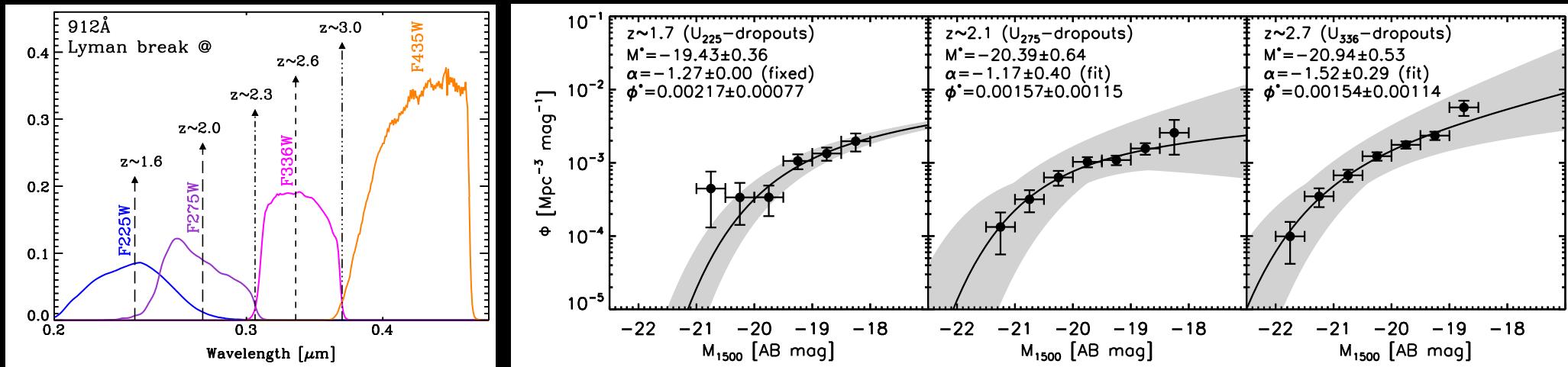
# Rest-frame UV-evolution of Early Type Galaxies since $z \lesssim 1.5$ .



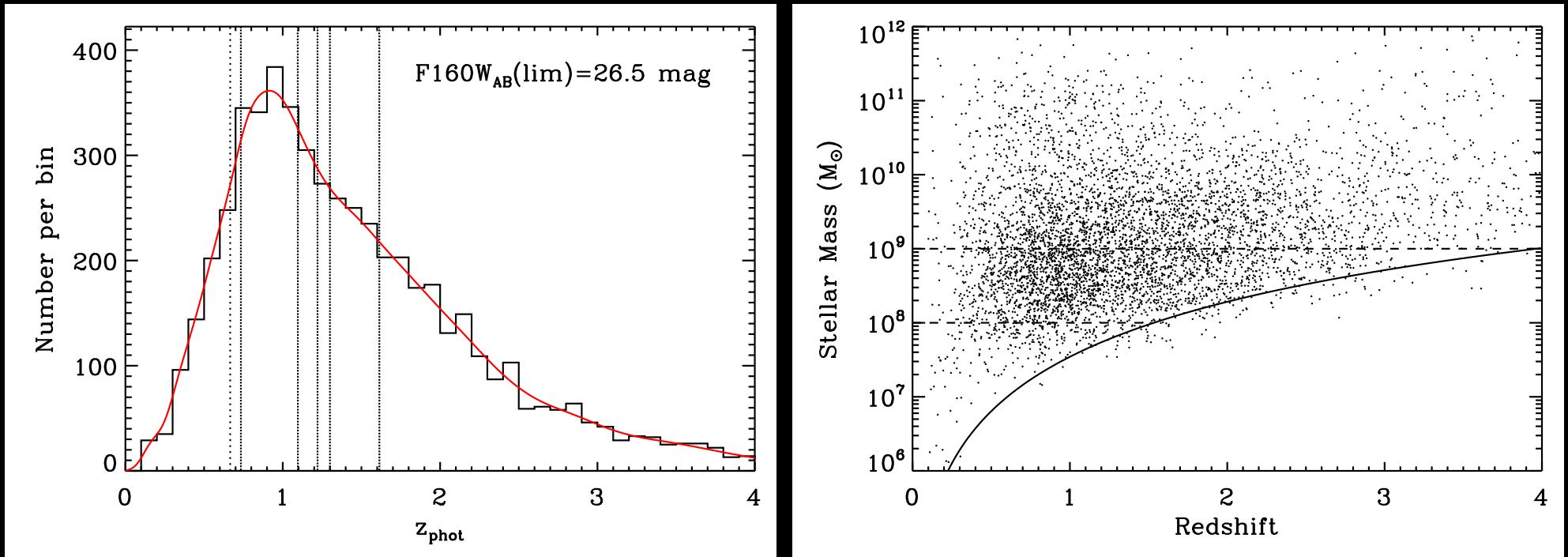
- 10-band WFC3 ERS data measured rest-frame UV-light in nearly all early-type galaxies at  $0.3 \lesssim z \lesssim 1.5$  (Rutkowski et al. 2012, ApJS, 199, 4).
- ⇒ Most ETGs have continued residual star-formation after they form.
- Can determine their  $N(z_{form})$ , which resembles the cosmic SFH diagram (*e.g.*, Madau et al. 1996). This can directly constrain the process of galaxy assembly and down-sizing (Kaviraj, Rutkowski et al. 2012, MNRAS).
- JWST will extend this to all redshifts with Balmer+4000Å-break ages.



## Lyman break galaxies at the peak of cosmic SF ( $z \simeq 1\text{-}3$ ; Hathi<sup>+</sup> 2010, 2013)



- JWST will similarly measure faint-end LF-slope evolution for  $1 \lesssim z \lesssim 12$ . (e.g., Bouwens<sup>+</sup> 2010, 2013; Hathi<sup>+</sup> 2010, 2013; Oesch + 2010; 2013; Ellis<sup>+</sup> 2012; Robinson<sup>+</sup> 2013).



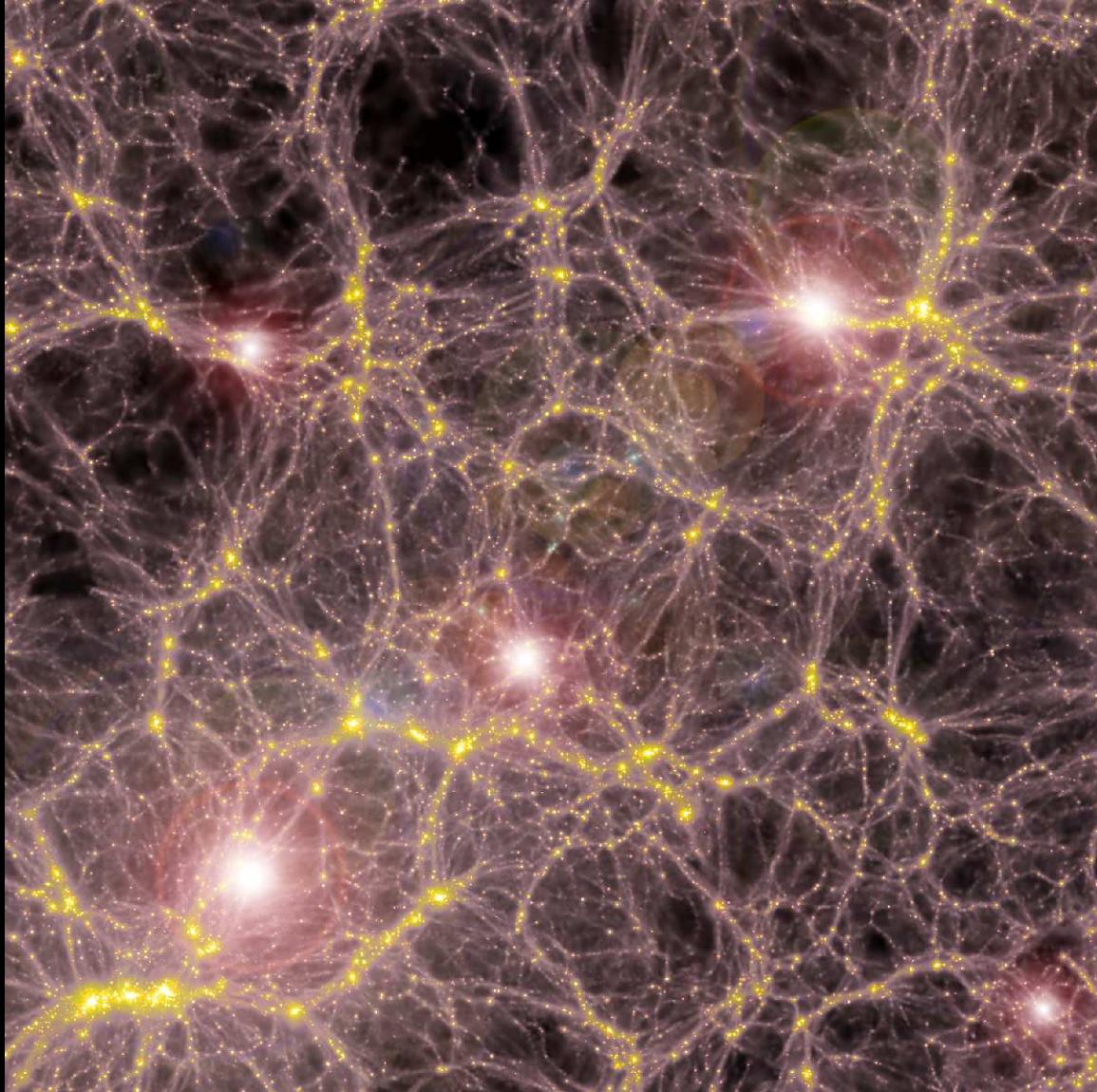
WFC3 ERS 10-band redshift estimates accurate to  $\lesssim 4\%$  with small systematic errors (Hathi et al. 2010, 2013), resulting in a reliable  $N(z)$ .

- Measure masses of faint galaxies to AB=26.5 mag, tracing the process of galaxy assembly: downsizing, merging, (& weak AGN growth?).  
⇒ Median redshift in (medium-)deep fields is  $z_{\text{med}} \simeq 1.5\text{--}2$ .

ERS shows WFC3's new panchromatic capabilities on galaxies at  $z \simeq 0\text{--}7$ .

- HUDF shows WFC3  $z \simeq 7\text{--}9$  capabilities (Bouwens<sup>+</sup> 2010; Yan<sup>+</sup> 2010).
- JWST will trace mass assembly and dust content  $\lesssim 5$  mag deeper from  $z \simeq 1\text{--}12$ , with nanoJy sensitivity from 0.7–5  $\mu\text{m}$ .

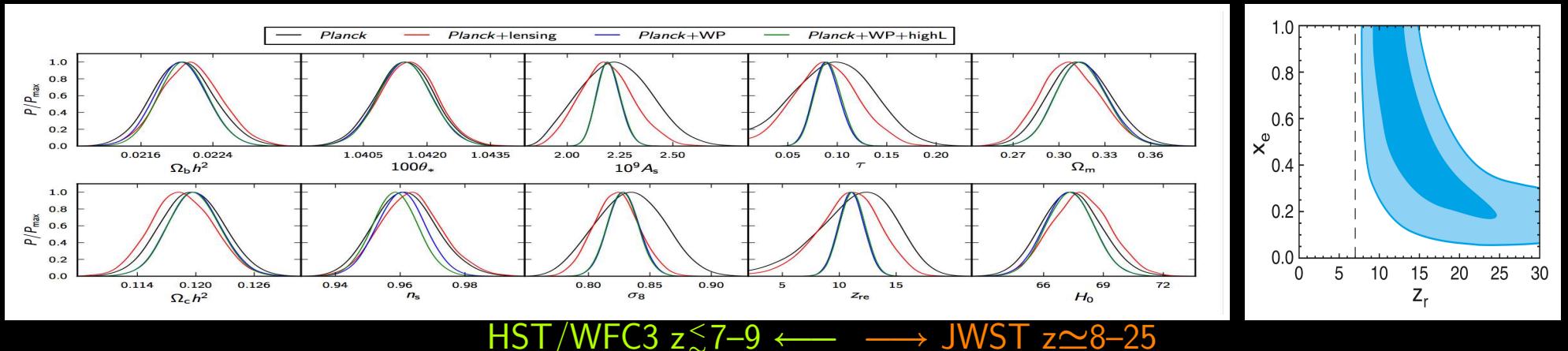
## (4a) How will JWST Observe First Light and Reionization?



- Detailed Hydrodynamical models (e.g., V. Bromm) suggest that massive Pop III stars may have reionized universe at redshifts  $z \lesssim 10-30$  (First Light).
- A this should be visible to JWST as the first Pop III stars and surrounding (Pop II.5) star clusters, and perhaps their extremely luminous supernovae at  $z \simeq 10 \rightarrow 30$ .

We must make sure we theoretically understand the likely Pop III mass-range, their IMF, their duplicity and clustering properties, their SN-rates, etc., with accurate predictions before JWST flies.

# Implications of the WMAP year-9 & Planck13 results for JWST science:



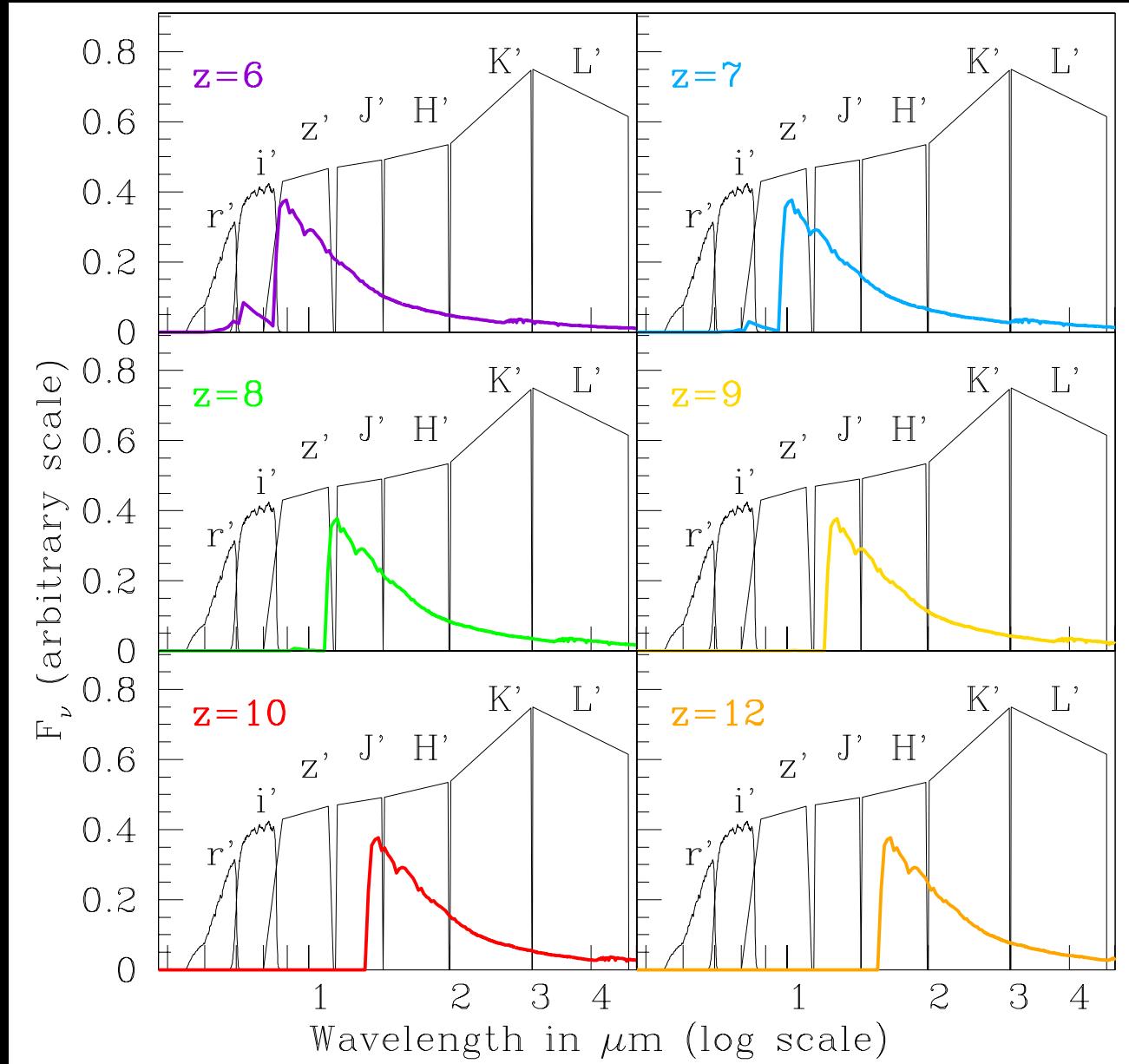
The year-9 WMAP data provided better foreground removal (Komatsu<sup>+</sup> 2011; Hinshaw<sup>+</sup> 2012; but see: Planck XVI 2013.)

$\implies$  First Light & Reionization occurred between these extremes:

- (1) Instantaneous at  $z \simeq 11.1 \pm 1.1$  ( $\tau = 0.089 \pm 0.013$ ), or:
- (2) Inhomogeneous & drawn out: starting at  $z \gtrsim 20$ , peaking at  $z \lesssim 11$ , ending at  $z \simeq 7$ . The implications for HST and JWST are:
- HST/ACS has covered  $z \lesssim 6$ , and WFC3 is covering  $z \lesssim 7-9$ .
- For First Light & Reionization, JWST will survey  $z \simeq 8$  to  $z \simeq 15-20$ .

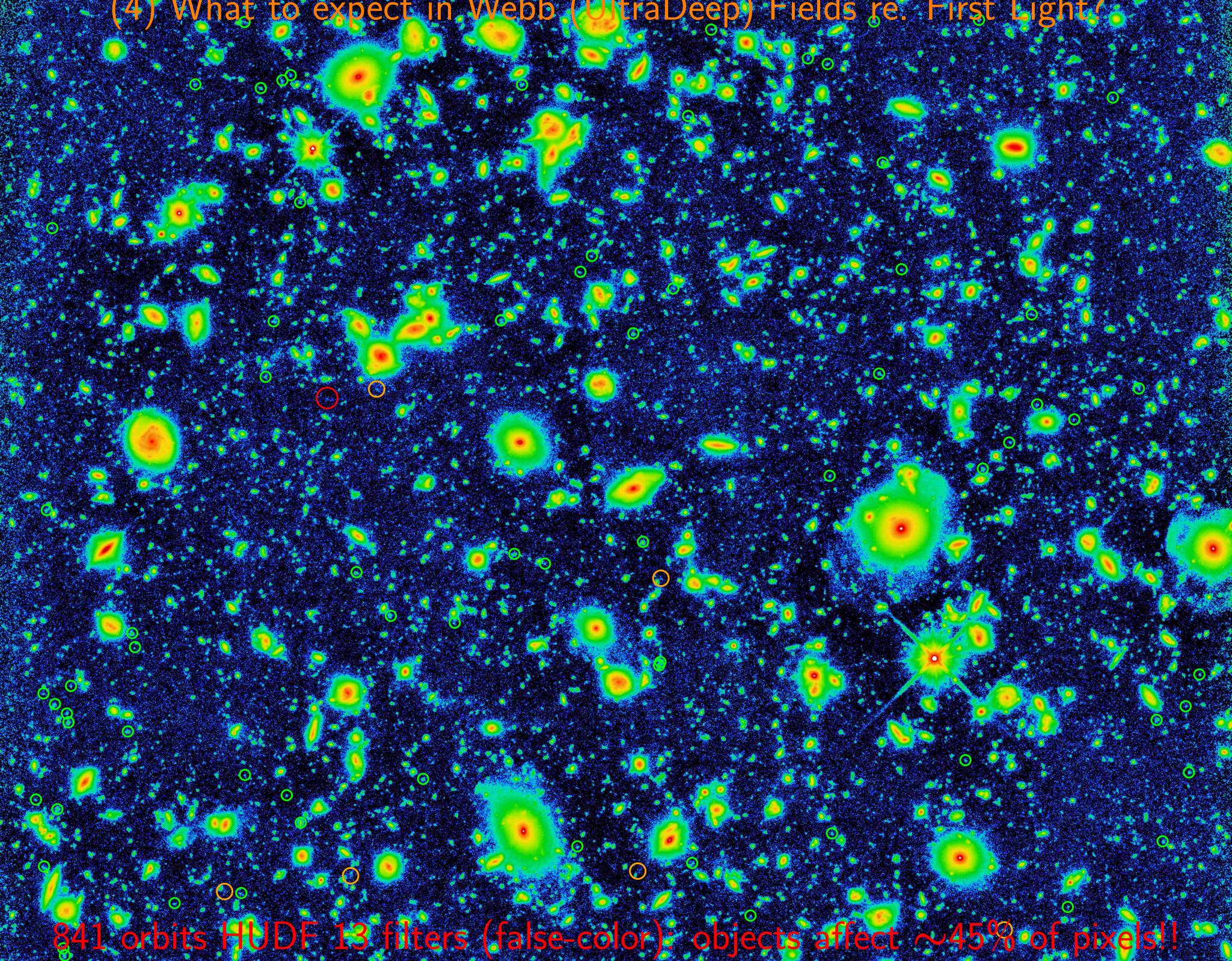
Question: If Planck- $\tau \downarrow \lesssim 0.08$  (TBD, Planck14), then how many reionizers will JWST see at  $z \simeq 10-20$ ?

#### (4) How will JWST measure First Light & Reionization?



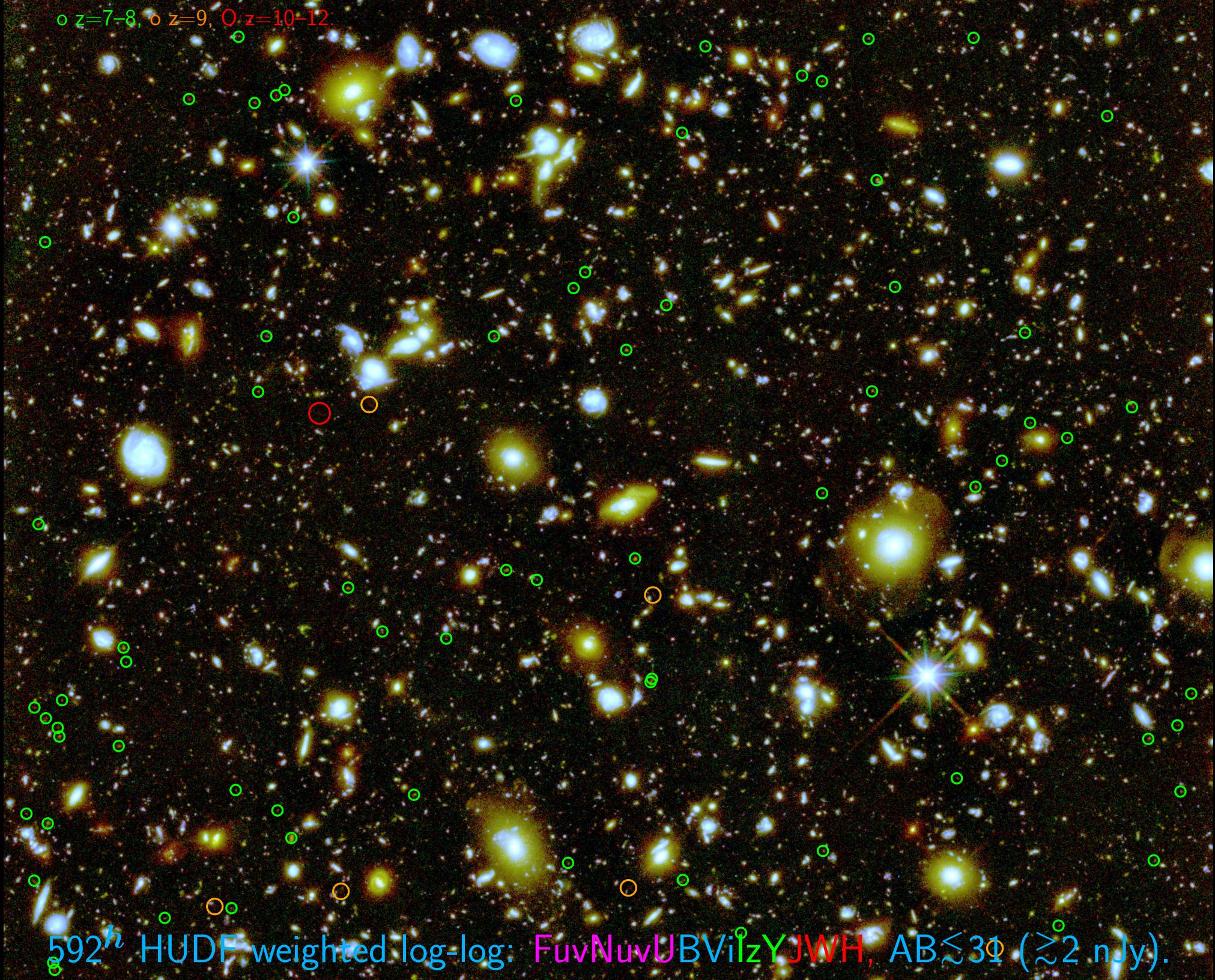
- Can't beat redshift: to see First Light, must observe near–mid IR.  
⇒ This is why JWST needs NIRCam at  $0.8\text{--}5 \mu\text{m}$  and MIRI at  $5\text{--}28 \mu\text{m}$ .

#### (4) What to expect in Webb (UltraDeep) Fields re: First Light?



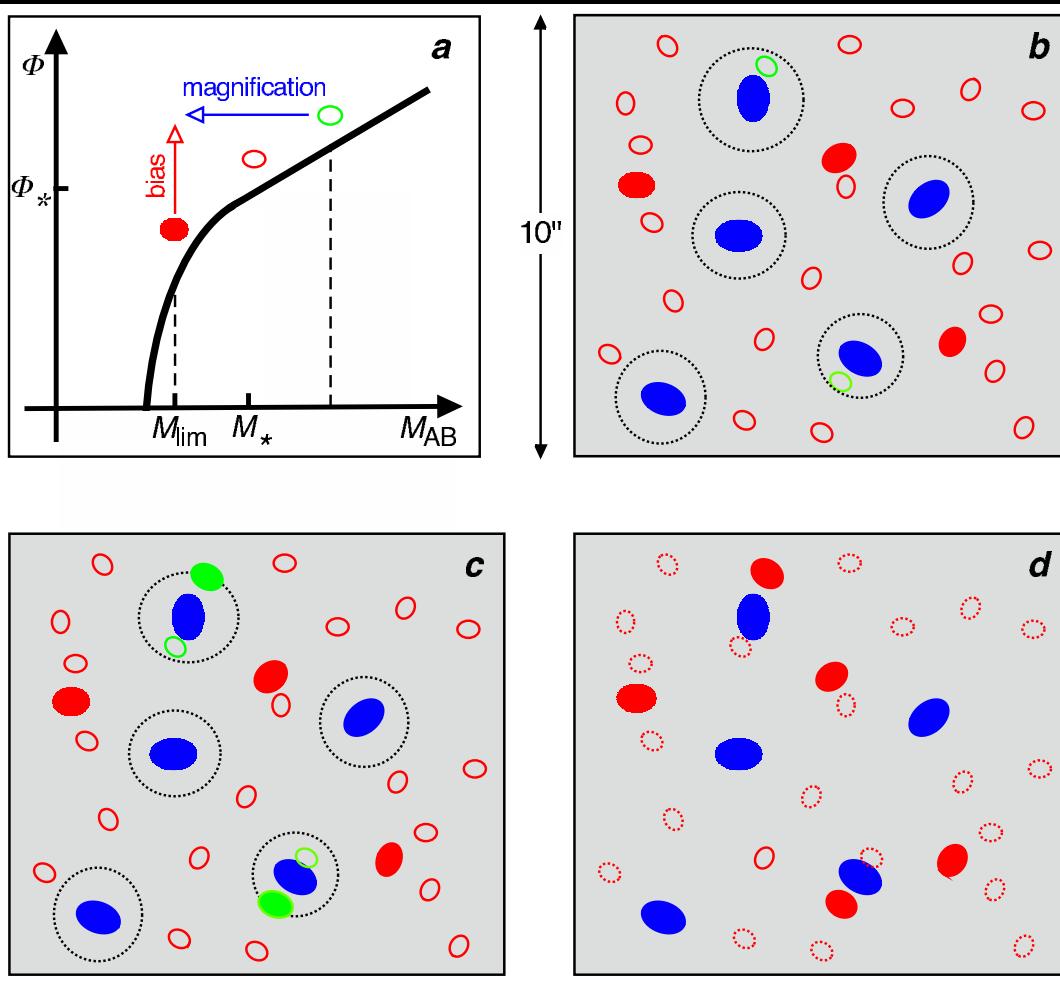
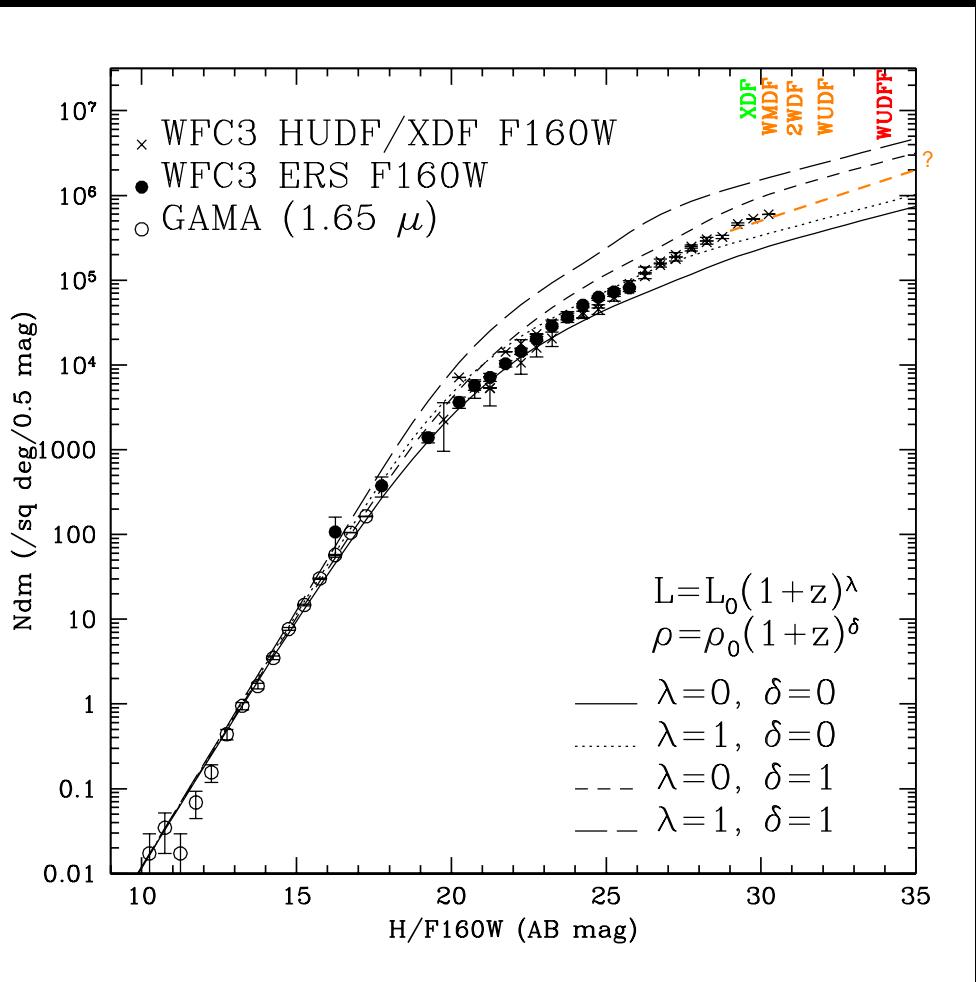
841 orbits HUDF 13 filters (false-color): objects affect  $\sim 45\%$  of pixels!!

○  $z=7-8$ , ○  $z=9$ , ○  $z=10-12$ .



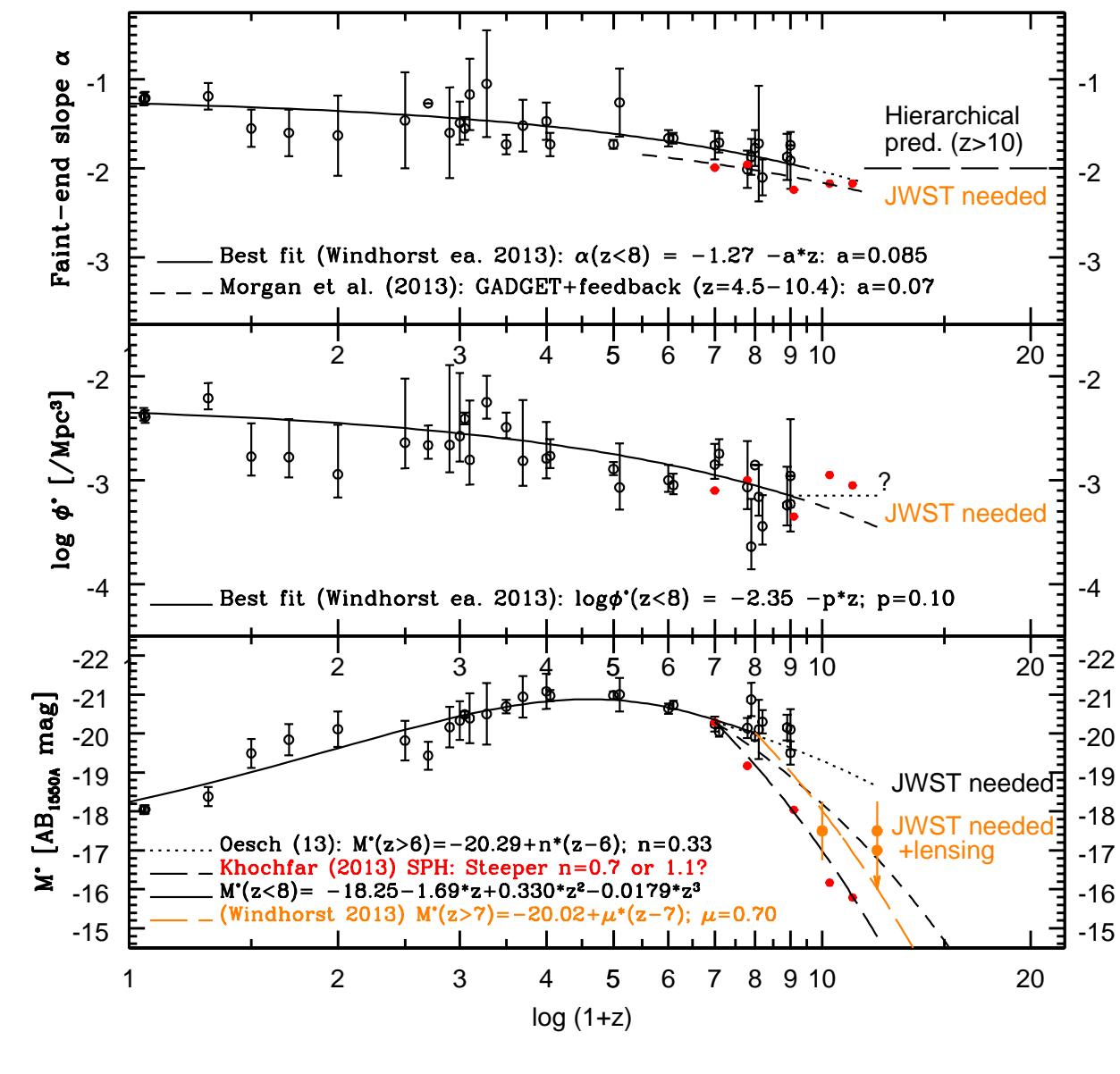
592<sup>h</sup> HUDF weighted log-log: FuvNuvUBVi<sup>z</sup>YJWH, AB $\lesssim 3\Phi$  ( $\gtrsim 2$  nJy).

# HUDF WFC3 IR Galaxy Counts: What to expect in Webb (UltraDeep) Fields?



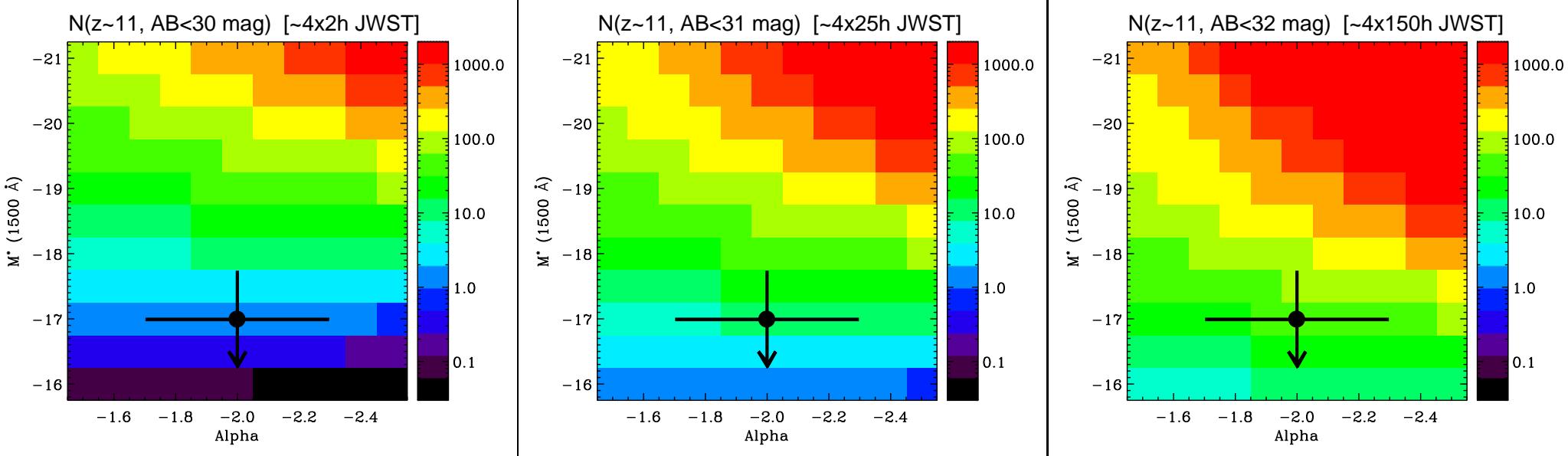
$1.6\mu\text{m}$  counts (Windhorst<sup>+</sup>2011). [F150W, F225W, F275W, F336W, F435W, F606W, F775W, F850LP, F105W, F125W, F140W not shown].

- Faint-end near-IR count-slope  $\simeq 0.12 \pm 0.02 \text{ dex/mag} \iff$
- Faint-end LF-slope ( $z_{\text{med}} \simeq 1.6$ )  $\alpha \simeq -1.4 \Rightarrow$  reach  $M_{AB} \simeq -14 \text{ mag.}$
- WUDF (---) can see  $AB \lesssim 32$  objects:  $M_{AB} \simeq -15$  (LMCs) at  $z \simeq 11$ .
- Lensing will change the landscape for JWST observing strategies.



Evolution of Schechter LF: faint-end LF-slope  $\alpha(z)$ ,  $\Phi^*(z)$  &  $M^*(z)$ :

- For JWST  $z \gtrsim 8$ , expect  $\alpha \lesssim -2.0$ ;  $\Phi^* \lesssim 10^{-3}$  ( $\text{Mpc}^{-3}$ ) (Oesch+ 11).
  - HUDF: Characteristic  $M^*$  may drop below -18 or -17.5 mag at  $z \gtrsim 10$ .
- ⇒ Will have significant consequences for JWST survey strategy.

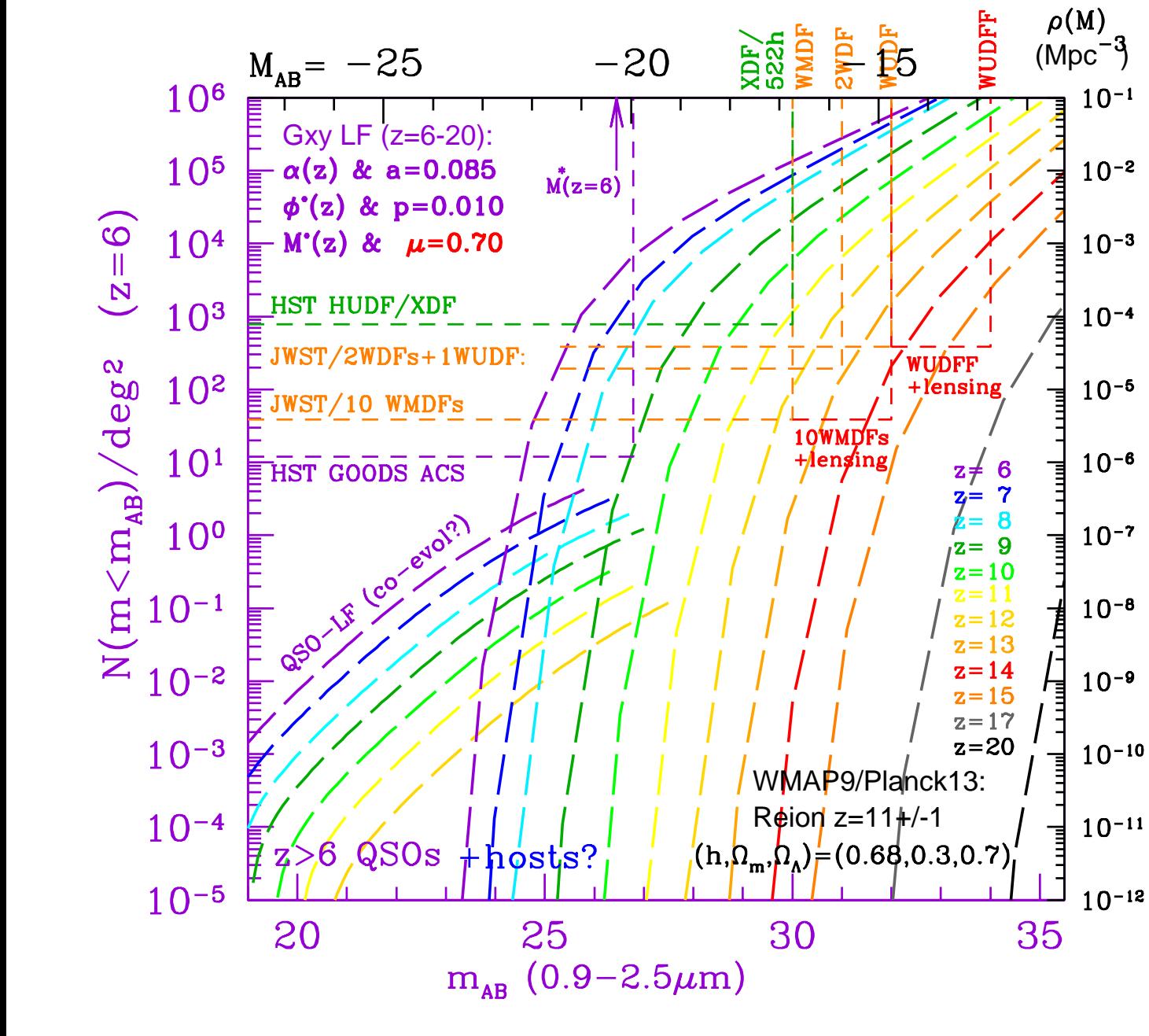


What do the 6 possible  $z \simeq 9$  and single  $z \gtrsim 10$  HUDF candidate mean?

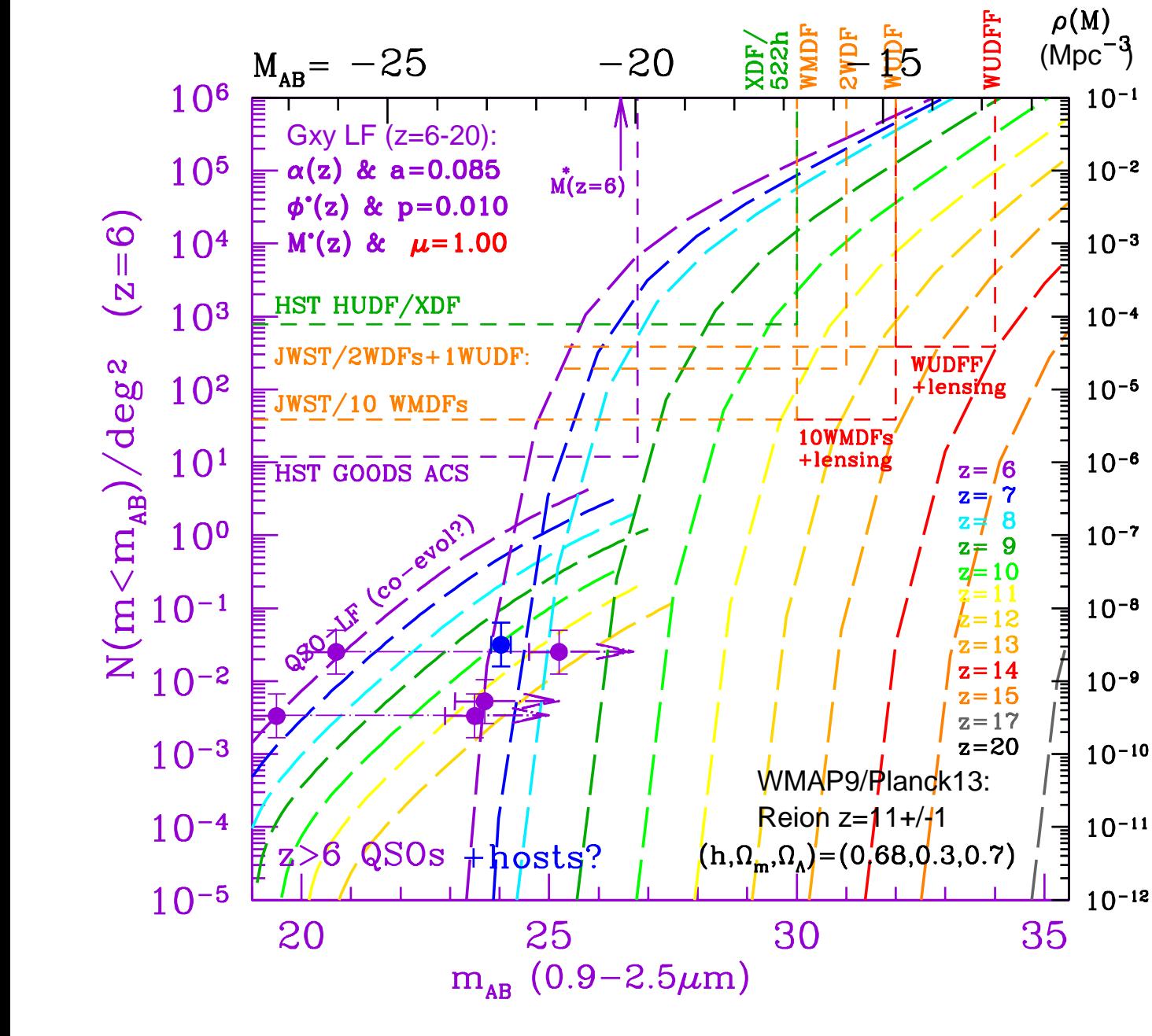
Integrate Schechter LFs with  $\alpha(z)$ ,  $\Phi^*(z)$  and  $M^*(z)$ :  $\lesssim 45\%$  sky-coverage by  $AB \lesssim 30$  objects (Koekemoer<sup>+</sup>13). Cosmic Variance  $\gtrsim 30\%$ .

For any  $\alpha(z \gtrsim 9-10)$ , implies  $M^*(z \gtrsim 10) \gtrsim -17.5$  mag (fainter!), so plan:

- (1) [Left] Webb “Medium-Deep” Fields (**WMDF**) ( $10 \times 4 \times 2h$  RAW): Expect few  $z \simeq 10-12$  objects to  $AB \lesssim 30$  mag, so plan lensing targets.
- (2) [Middle] Webb Deep Field (**WDF**) ( $4 \times 25h$  7-filt NIRCam GTO): Expect 8–25 objects at  $z \simeq 10-12$  to  $AB \lesssim 31$  mag.
- (3) [Right] Webb UltraDeep Field (**WUDF**) ( $4 \times 150h$ ; NIRCam DD?): Expect 30–90 objects to  $AB \lesssim 32$  mag, many more if lensing targets.



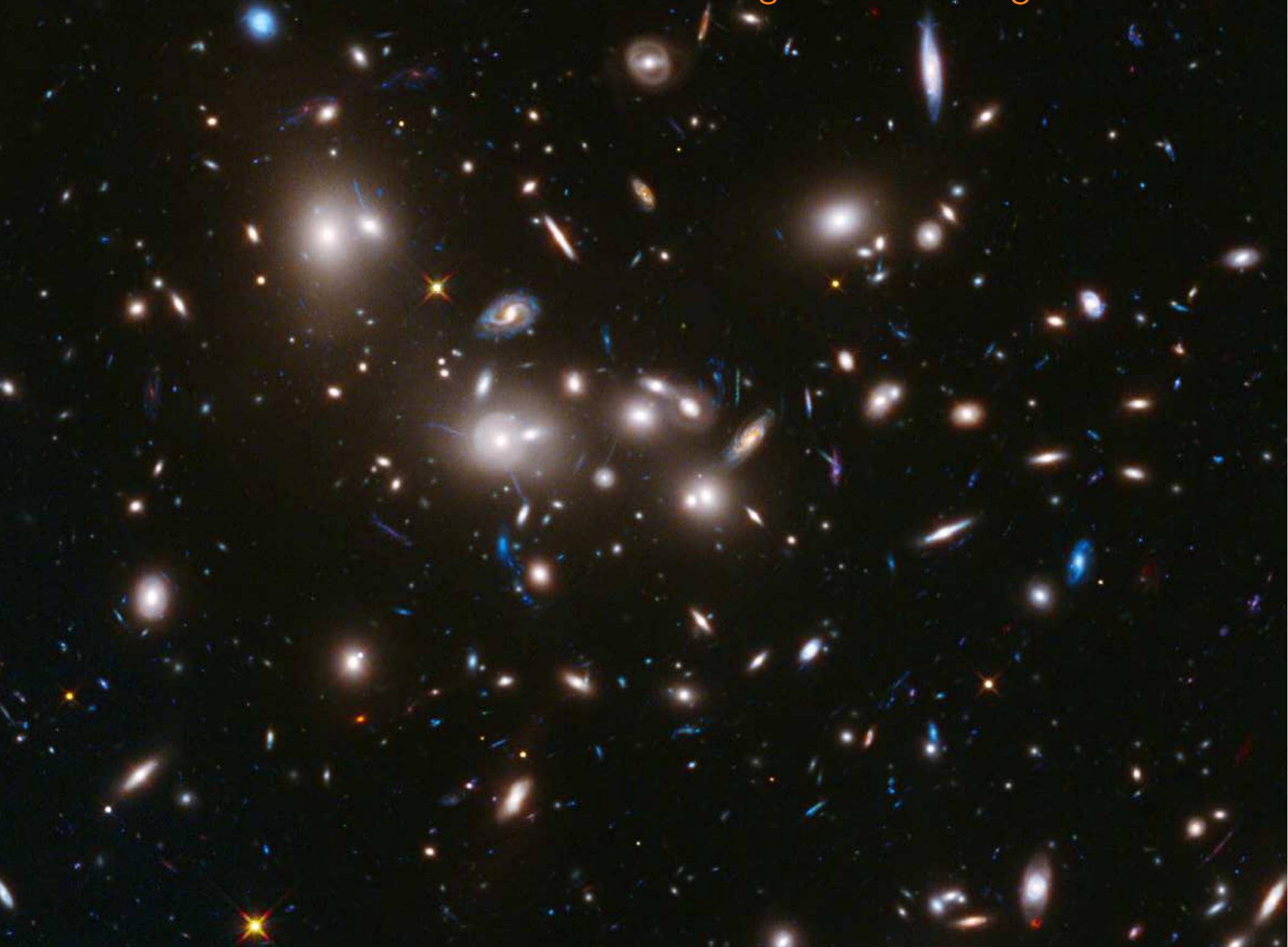
- Schechter LF ( $z \lesssim 6 \lesssim 20$ ) with  $\alpha(z)$ ,  $\Phi^*(z)$ ,  $M^*(z)$  above &  $\mu=0.70$ .  
 Area/Sensitivity for: HUDF/XDF, 10 WMDFs, 2 WDFs, & 1 WUDF.  
 • Will need lensing targets for WMDF–WUDFF to see  $z \simeq 14\text{--}16$  objects.



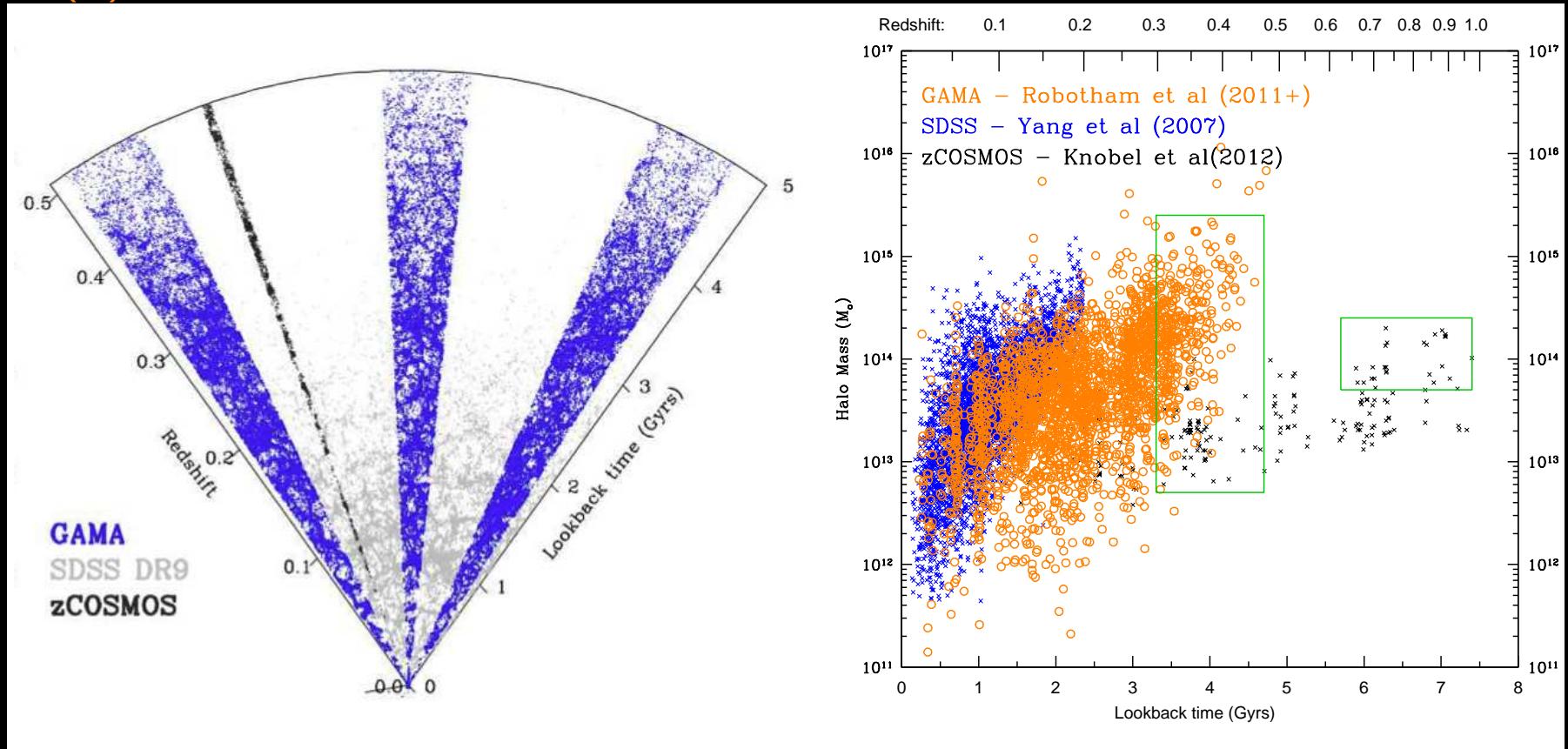
Same as pg. 15, but pessimistic  $M^*(z)$  evolution parameter:  $\mu=1.0$ .

- If so, JWST surveys would need lensing to see most  $\gtrsim 11$  objects.
  - Add  $z \simeq 6$  QSO host galaxy limits (or fluxes) by Mechtley<sup>+</sup> (2012, 2013).

HST Frontier Field A2744: JWST needs lensing to see First Light at  $z \gtrsim 10-15$ !



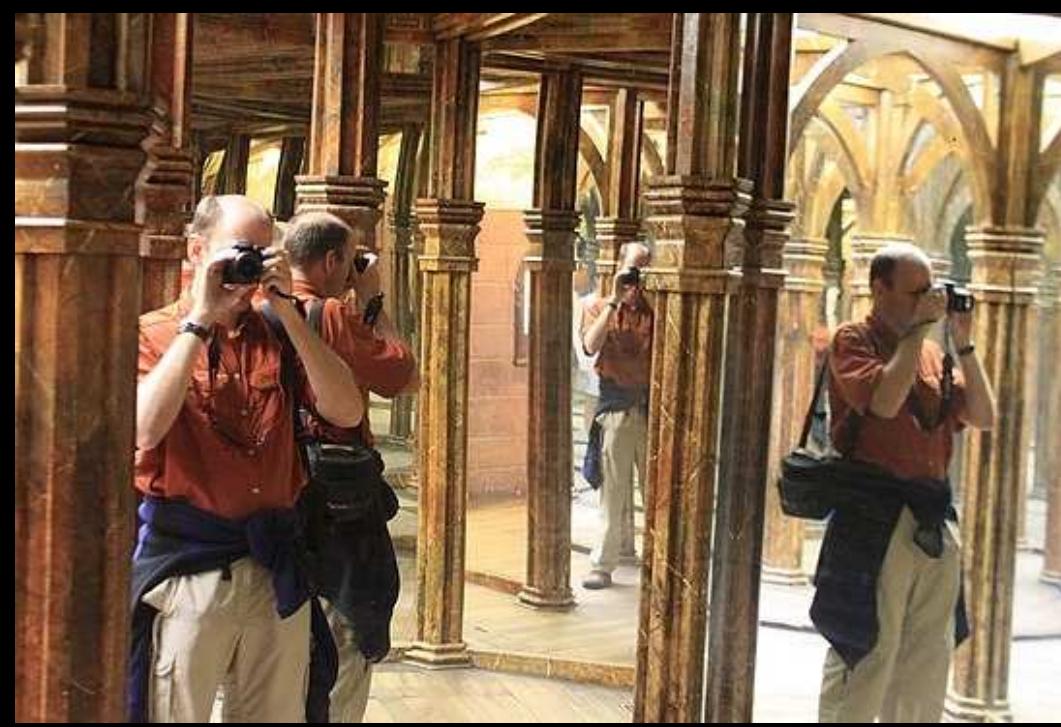
#### (4) Gravitational Lensing to see First Light population at $z \gtrsim 10$ .



What are the best lenses in 2018: Rich clusters or (compact) galaxy groups?

[Left] Redshift surveys: SDSS  $z \lesssim 0.25$  (Yang<sup>+</sup> 2007), GAMA  $z \lesssim 0.45$  (Robotham<sup>+</sup> 2011), and zCOSMOS  $z \lesssim 1.0$  (Knobel<sup>+</sup> 2012).

- GAMA: 22,000 groups  $z \lesssim 0.45$ ; 2400 with  $N_{spec} \gtrsim 5$  (Robotham<sup>+</sup> 11).
- $\lesssim 10\%$  of GAMA groups compact for lensing (Konstantopoulos<sup>+</sup> 13).
- Large group sample to identify optimal lens-candidates for  $z \gtrsim 6$  sources.



Two fundamental limitations may determine ultimate JWST image depth:

- (1) Cannot-see-the-forest-for-the-trees effect [Natural Confusion limit]:  
Background objects blend into foreground because of their own diameter  
⇒ Need multi- $\lambda$  deblending algorithms.
- (2) House-of-mirrors effect [“Gravitational Confusion”]: Most First Light objects at  $z \gtrsim 12-14$  may need to be found by cluster or group lensing.  
⇒ Need multi- $\lambda$  object finder that works on sloped backgrounds  
⇒ If  $M^*(z \gtrsim 10) \gtrsim -18$ , need to use & model gravitational foreground.

## (5) Conclusions

- (1) HST set stage to measure galaxy assembly in the last 12.7-13.0 Gyrs.
  - Most  $z \simeq 6$  QSO host galaxies faint (dusty?), with 1 exception:  $L >> L^*$ .
- (2) JWST passed Preliminary & Critical Design Reviews in 2008 & 2010. Management replan in 2010-2011. No technical showstoppers thus far:
  - More than 80% of JWST H/W built or in fab, & meets/exceeds specs.
- (3) JWST is designed to map the epochs of First Light, Reionization, and Galaxy Assembly & SMBH-growth in detail. JWST will determine:
  - Formation and evolution of the first star-clusters after 0.2 Gyr.
  - How dwarf galaxies formed and reionized the Universe after 1 Gyr.
- (4) JWST will have a major impact on astrophysics this decade:
  - IR sequel to HST after 2018: Training the next generation researchers.
  - JWST will define the next frontier to explore: the Dark Ages at  $z \gtrsim 20$ .

# SPARE CHARTS

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- References and other sources of material shown:

<http://www.asu.edu/clas/hst/www/jwst/> [Talk, Movie, Java-tool]

<http://www.asu.edu/clas/hst/www/ahah/> [Hubble at Hyperspeed Java–tool]

<http://www.asu.edu/clas/hst/www/jwst/clickonHUDF/> [Clickable HUDF map]

<http://www.jwst.nasa.gov/> & <http://www.stsci.edu/jwst/>

<http://ircamera.as.arizona.edu/nircam/>

<http://ircamera.as.arizona.edu/MIRI/>

<http://www.stsci.edu/jwst/instruments/nirspec/>

<http://www.stsci.edu/jwst/instruments/fgs>

Gardner, J. P., et al. 2006, *Space Science Reviews*, 123, 485–606

Mather, J., & Stockman, H. 2000, *Proc. SPIE* Vol. 4013, 2

Windhorst, R., et al. 2008, *Advances in Space Research*, 41, 1965

Windhorst, R., et al., 2011, *ApJS*, 193, 27 ([astro-ph/1005.2776](#)).

# Northrop Grumman Expertise in Space Deployable Systems

- Over 45 years experience in the design, manufacture, integration, verification and flight operation of spacecraft deployables
- 100% mission success rate, comprising over 640 deployable systems with over 2000 elements



### Baseline "Cup Down" Tower Configuration at JSC (Before)



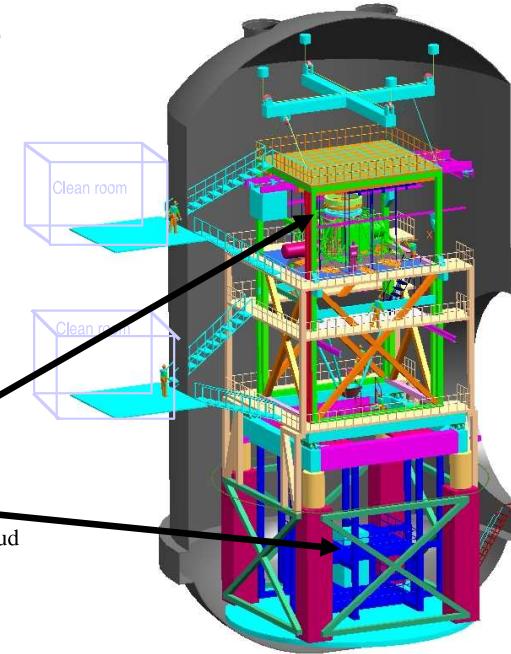
Most recent Tower Design shows an Inner Optical Tower supported by a Outer structure with Vibration Isolation at the midplane. Everything shown is in the 20K region (helium connections, etc. not shown) except clean room and lift fixture.

Current plan calls for 33KW cooldown capability, 12 KW steady state, 300-500mW N2 cooling

JSC currently has 7 KW He capability

Current plan includes 10 trucks of LN2/day during cooldown

Interferometers, Sources, Null Lens and Alignment Equipment Are in Upper and Lower Pressure Tight Enclosure Inside of Shroud



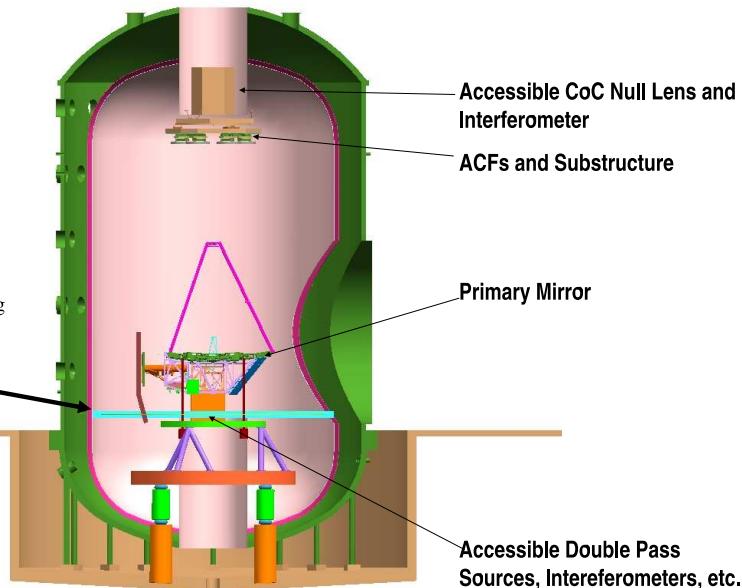
### JSC "Cup Up" Test Configuration (New Proposal)



No Metrology Tower and Associated Cooling H/W. External Metrology

Two basic test options:

1. Use isolators, remove drift through fast active control + freeze test equipment jitter
  2. Eliminate vibration isolators (but use soft dampeners) to avoid drift, freeze out jitter
- Builds on successful AMSD heritage of freezing and averaging jitter, testing through windows.



Drawing care of ITT

Page 6

JWST underwent several significant replans and risk-reduction schemes:

- ≈2003: Reduction from 8.0 to 7.0 to 6.5 meter. Ariane-V launch vehicle.
- 2005: Eliminate costly 0.7-1.0  $\mu\text{m}$  performance specs (kept 2.0  $\mu\text{m}$ ).
- 2005: Simplification of thermal vacuum tests: cup-up, not cup-down.
- 2006: All critical technology at Technical Readiness Level 6 (TRL-6).
- 2007: Further simplification of sun-shield and end-to-end testing.
- 2008: Passes Mission Preliminary Design & Non-advocate Reviews.
- 2010, 2011: Passes Mission Critical Design Review: Replan Int. & Testing.

# Fiscal Year 2014 HQ Milestones

Assumes JWST is appropriated in FY2014 the full President's budget request of new obligation authority (NOA).

Month	Milestone	Comment
Oct-13	1 Primary Mirror Backplane Support Structure Cryogenic Testing Readiness Review	Completed 9/10
	2 Mirror Deployment Electronics Unit Manufacturing Readiness Review	Completed 10/8
Nov-13	3 Jet Propulsion Lab. (JPL) Cryogenic Test Chamber Readiness Review	Delayed: pulse tube, cooler shield issues
	4 Johnson Space Center (JSC) Telescope and ISIM support structure fabrication complete	Completed 11/4
Dec-13	5 Spacecraft Critical Design Review Complete	Delayed to 1/14 [shutdown]
	6 MIRI Cryocooler Flight Cold Head Assembly delivered to ISIM	Delayed 1/21/2014
	7 JSC Clean Room ready to receive ground support equipment	Delayed to 1/14 [shutdown]
	8 Complete ISIM cryogenic-vacuum risk reduction test	Concluded 11/13/2013, but not all tests completed because of shutdown
Jan-14	9 Delivery of last Primary Mirror Segment to GSFC	Completed 12/16
	10 Observatory Operations software scripts Build 3 Complete	
	11 New detector focal plane arrays for NIRCam ready for integration into instrument	Completed 11/20
Feb-14	12 Secondary Mirror Mount delivery	
	13 MIRI Cryocooler flight electronics delivered to JPL	Completed 11/22
	14 Final Data Management Subsystem Design Review	
	15 Flight NIRCam and NIRSpec ready for integration into ISIM	Delayed to 3/14 [shutdown]
Mar-14	16 Spacecraft Solar Array Manufacturing Readiness Review	
	17 JSC Chamber A Telescope ground support equipment test #1 design review	
Apr-14	18 Telescope actuators electronics drive unit delivery	
	19 Flight MIRI cryocooler assembly delivered to JPL	
	20 MIRI Cryocooler Flight Refrigerant Line Deployment Assembly delivered to integration and testing	
	21 Sunshield Membrane Cover Assembly Manufacturing Readiness Review	
	22 MIRI cryocooler Test Readiness Review	
May-14	23 Updated Observatory Commissioning Plan (rev C) delivery	
	24 Start acceptance testing of flight cryocooler assembly and associated electronics	
	25 Start cryo-vacuum test with fully integrated ISIM ("CV2")	Delayed to 6/14 [shutdown]
Jun-14	26 Flight spare MIRI cryocooler assembly delivered to JPL	
	27 JSC Chamber A bake-out and cryogenic proof testing complete	
	28 Hardware ready for MIRI cryo cooler test #3: checkout complete	
Jul-14	29 Spacecraft Mid-Course Correction Thruster Final Assembly complete	
	30 Proposal Planning Subsystem build 9 complete	
	31 Sunshield Mid-boom and Stem assembly Manufacturing Readiness Review	
Aug-14	32 Spacecraft Flight Software Build 2.2 Test Readiness Review	
	33 NIRSpec and FGS/NIRISS new Focal Plane Arrays ready for integration	Delayed to 9/14 [shutdown]
	34 JSC cryogenic test telescope and ISIM test ground support equipment integration complete	
Sep-14	35 Complete cryo-vacuum test of fully integrated ISIM ("CV2")plete	Delayed to 10/14 [shutdown]
	36 NIRSpec new microshutters ready for integration	Delayed to 10/14 [shutdown]

Blue font denotes milestones accomplished ahead of schedule, orange font denotes milestones accomplished late.

Milestones: How the Project reports its progress monthly to Congress.

# Milestone Performance

- Since the September 2011 replan JWST reports high-level milestones monthly to numerous stakeholders

	Total Milestones	Total Milestones Completed	Number Completed Early	Number Completed Late	Deferred to Next Year
FY2011	21	21	6	3	0
FY2012	37	34	16	2	3
FY2013	41	38	20	5	3
FY2014	36	7	5	10*	0

\*Late milestones have been or are forecast to complete within the year.  
Shutdown related delayed milestones included in this tally

7 out of 10 FY14 milestones late by 1 month due to Government shutdown.

None of these are on the critical path, so caused no launch delay.

V3 (anti-spacecraft)

V1  
V2

Secondary mirror

Cassegrain focus

Fine Steering Mirror

f/#: 20.0

Effective Focal Length: 131.4 m

PM diameter = 6.6 m (circumscribed circle)

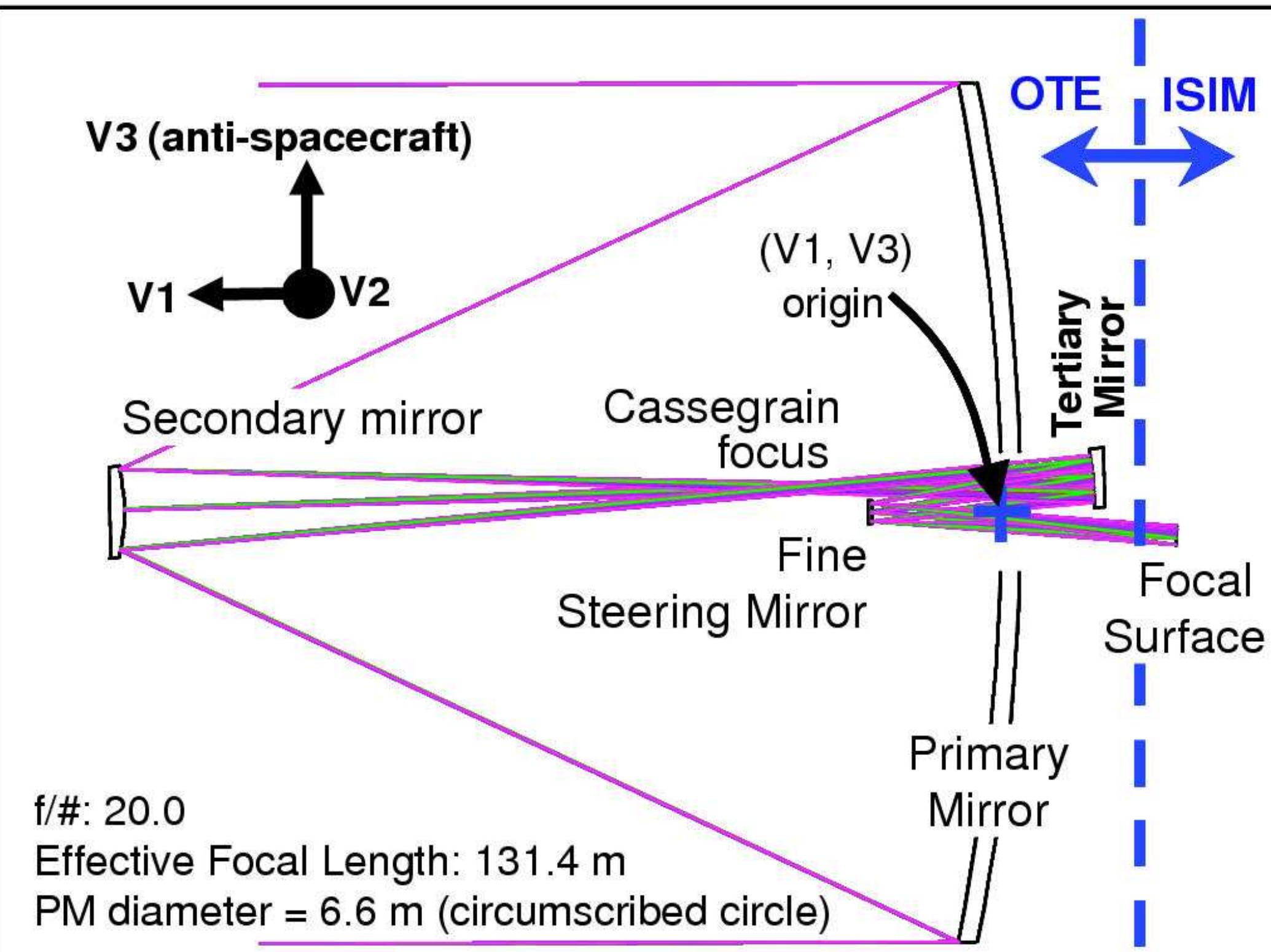
(V1, V3)  
origin

OTE  
ISIM

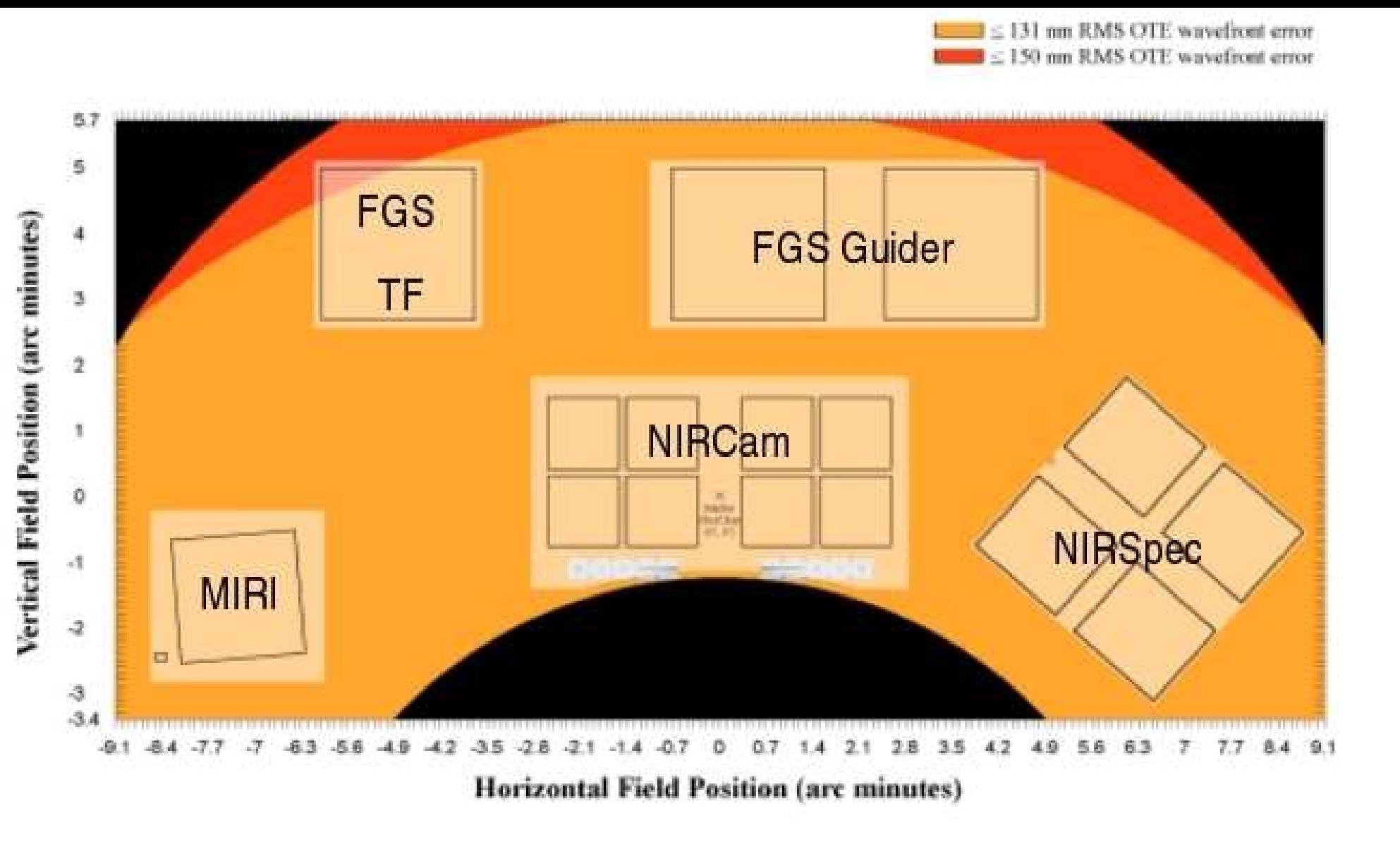
Tertiary  
Mirror

Focal  
Surface

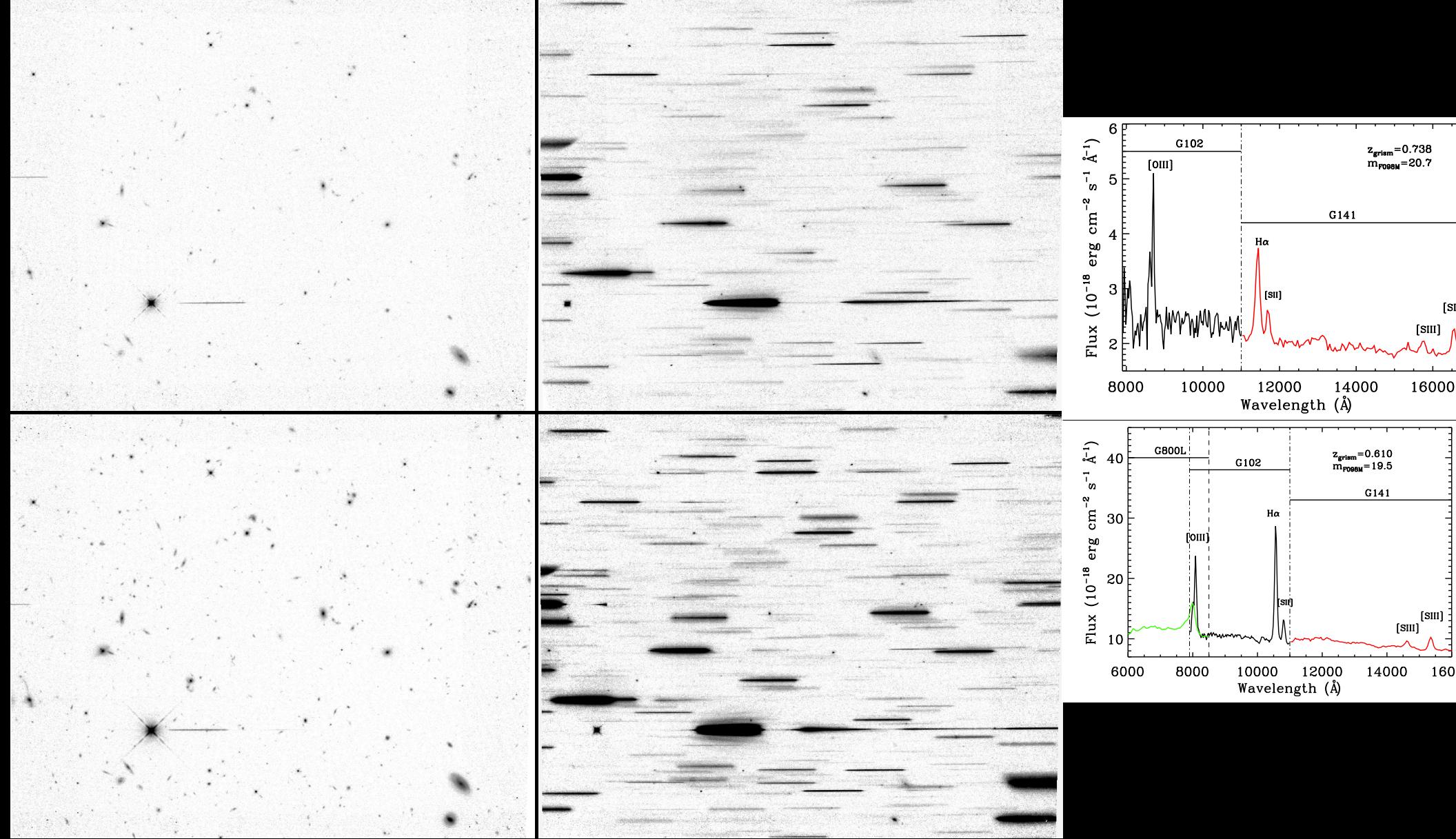
Primary  
Mirror



- (3c) What instruments will JWST have?



- All JWST instruments can in principle be used in parallel observing mode:
- Currently only being implemented for parallel *calibrations*.



HST/WFC3 G102 & G141 grism spectra in GOODS-S ERS (Straughn<sup>+</sup> 2010)

IR grism spectra from space: unprecedented new opportunities in astrophysics.

- JWST will provide near-IR grism spectra to AB $\lesssim$ 29 mag from 2–5.0  $\mu$ m.

## (4b) Predicted Galaxy Appearance for JWST at redshifts $z \simeq 1-15$

NGC 3310



ESO0418-008



UGC06471-2



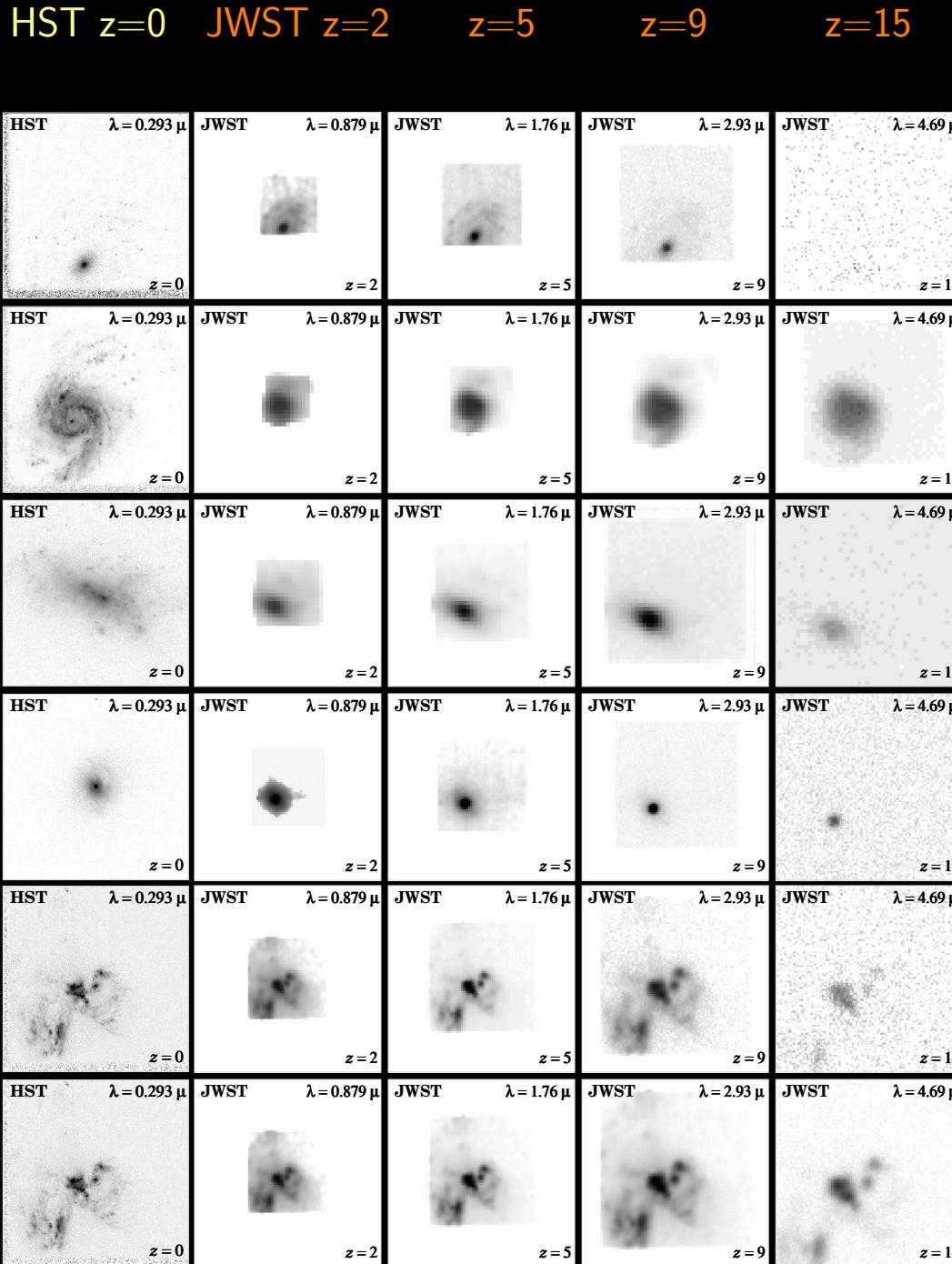
### Ultraviolet Galaxies

HST • WFPC2

NASA and R. Windhorst (Arizona State University) • STScI-PRC01-04

- The rest-frame UV-morphology of galaxies is dominated by young and hot stars, with often significant dust imprinted (Mager-Taylor et al. 2005).
- High-resolution HST ultraviolet images are benchmarks for comparison with very high redshift galaxies seen by JWST.

## (4b) Predicted Galaxy Appearance for JWST at redshifts $z \simeq 1$ –15



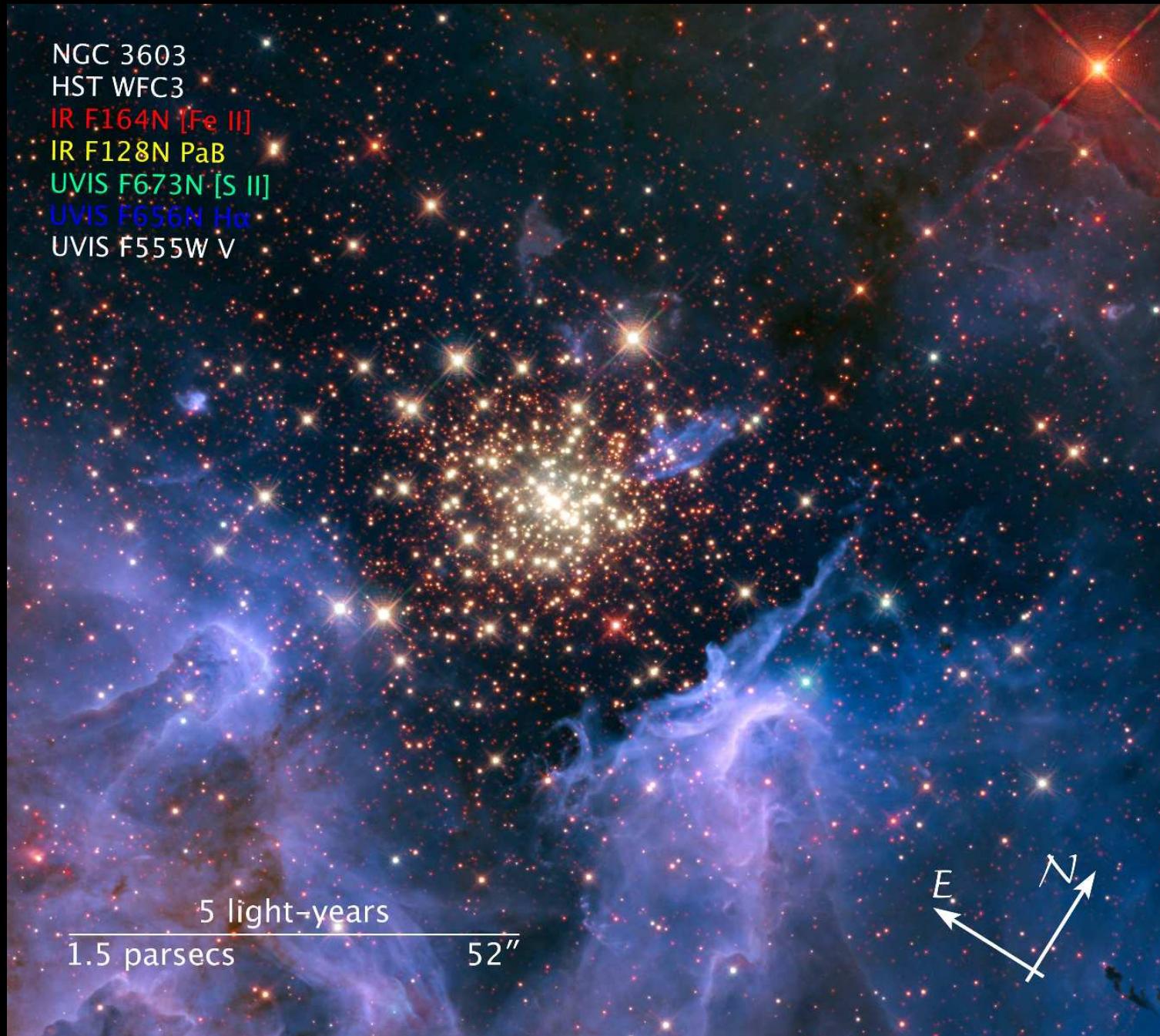
With Hubble UV-optical images as benchmarks, JWST can measure the evolution of galaxy structure & physical properties over a wide range of cosmic time:

- (1) Most spiral disks will dim away at high redshift, but most formed at  $z \lesssim 1$ –2.

Visible to JWST at very high  $z$  are:

- (2) Compact star-forming objects (dwarf galaxies).
- (3) Point sources (QSOs).
- (4) Compact mergers & train-wrecks.

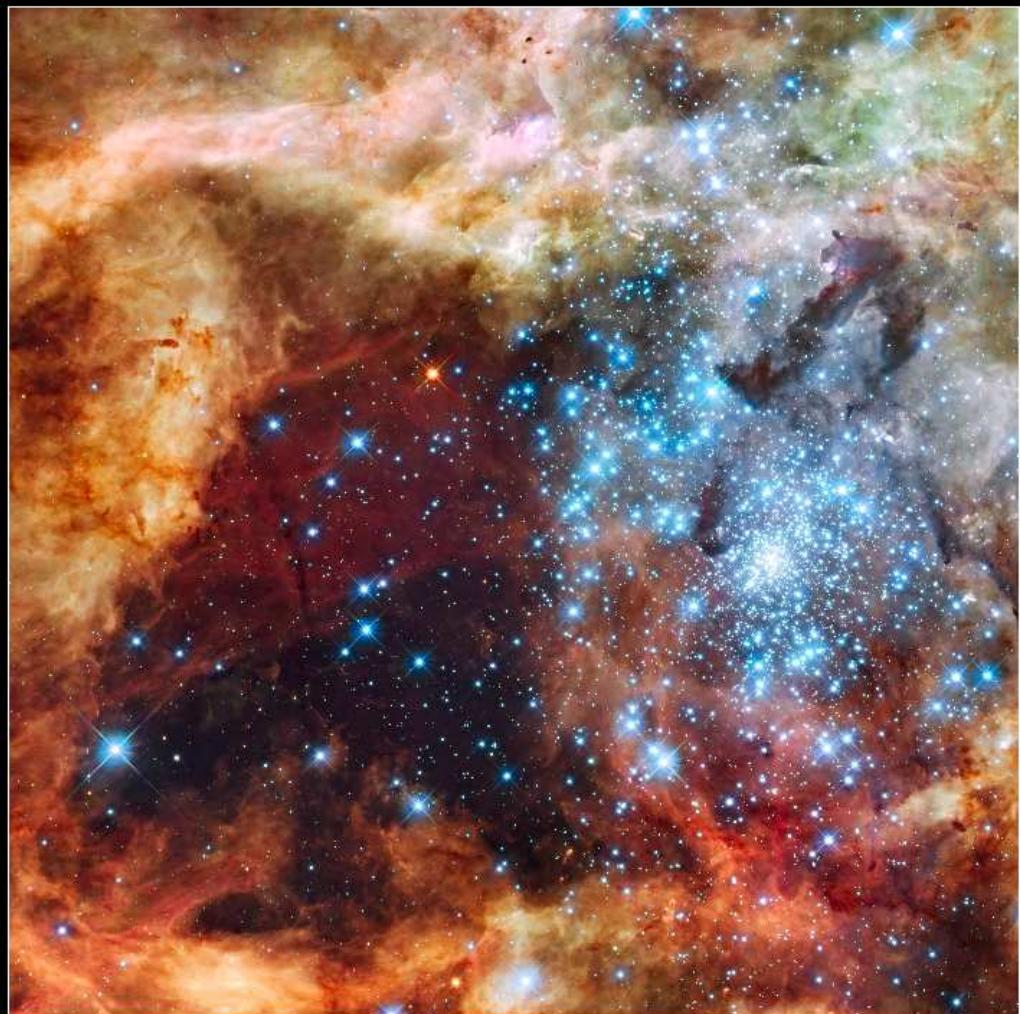
## (6) How can JWST measure Star-Formation and Earth-like exoplanets?



NGC 3603: Young star-cluster triggering star-birth in “Pillars of Creation”

Visible

Infrared



## 30 Doradus Nebula and Star Cluster

*Hubble Space Telescope • WFC3/UVIS/IR*

NASA, ESA, F. Paresce (INAF-IASF, Italy), and the WFC3 Science Oversight Committee

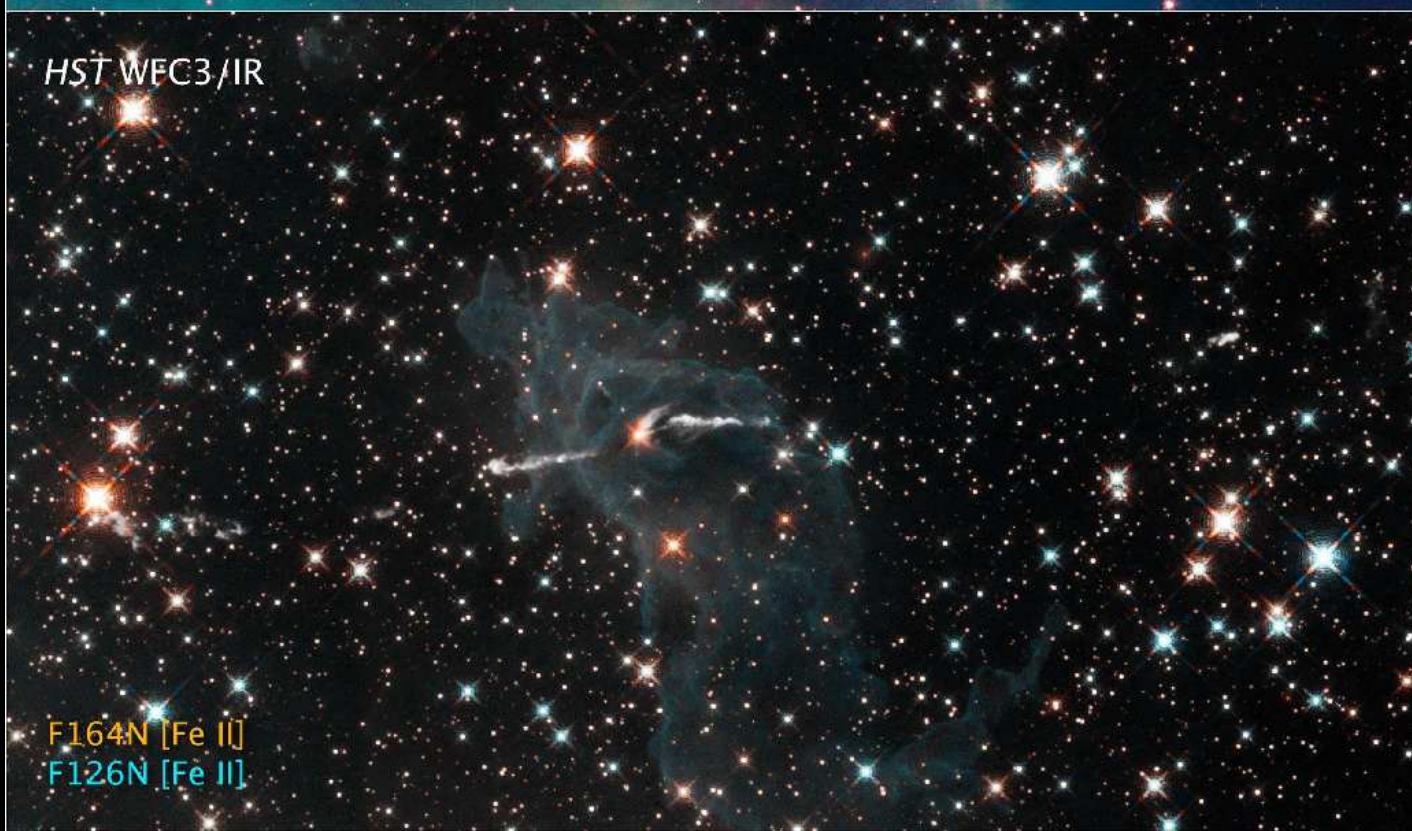
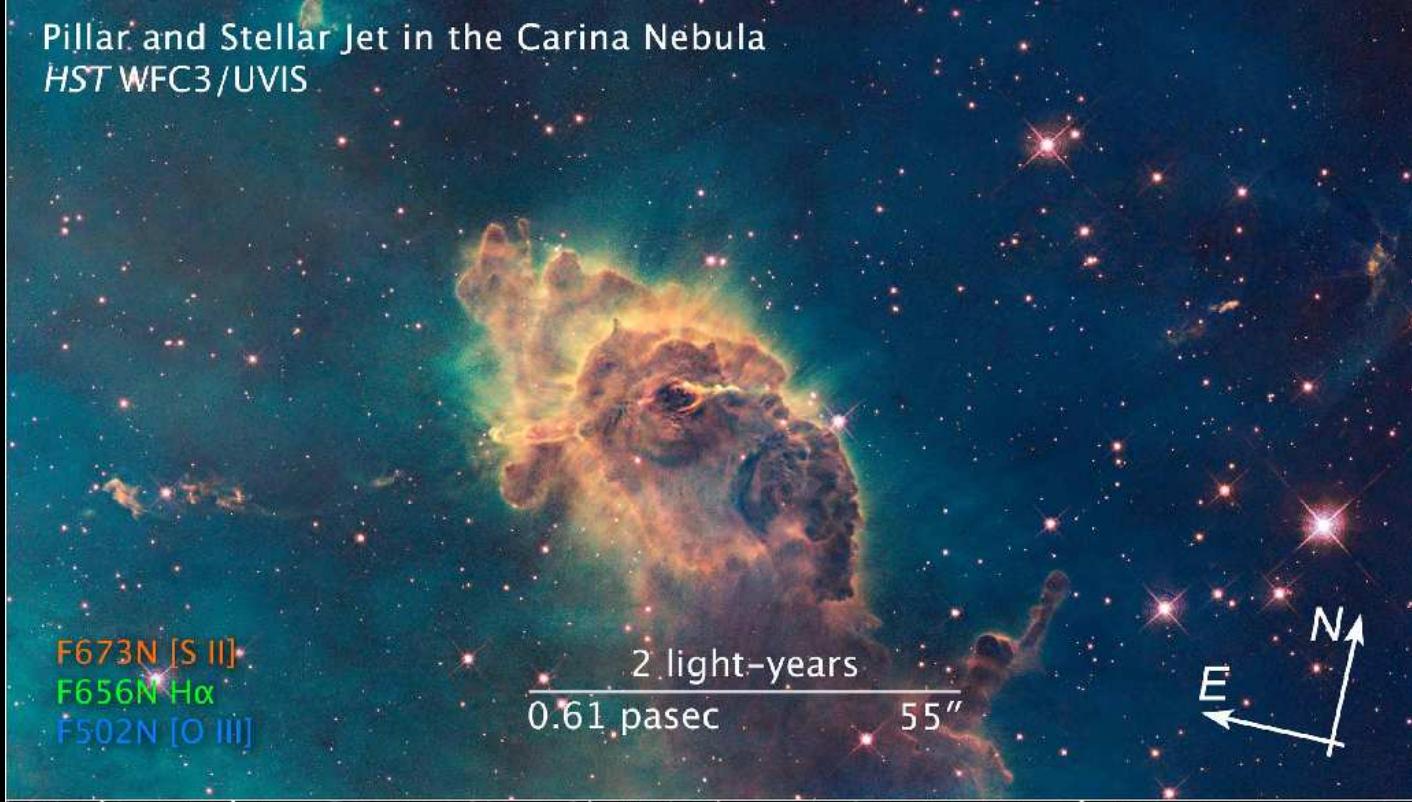
STScI-PRC09-32b

30 Doradus: Giant young star-cluster in Large Magellanic Cloud (150,000 ly), triggering birth of Sun-like stars (and surrounding debris disks).

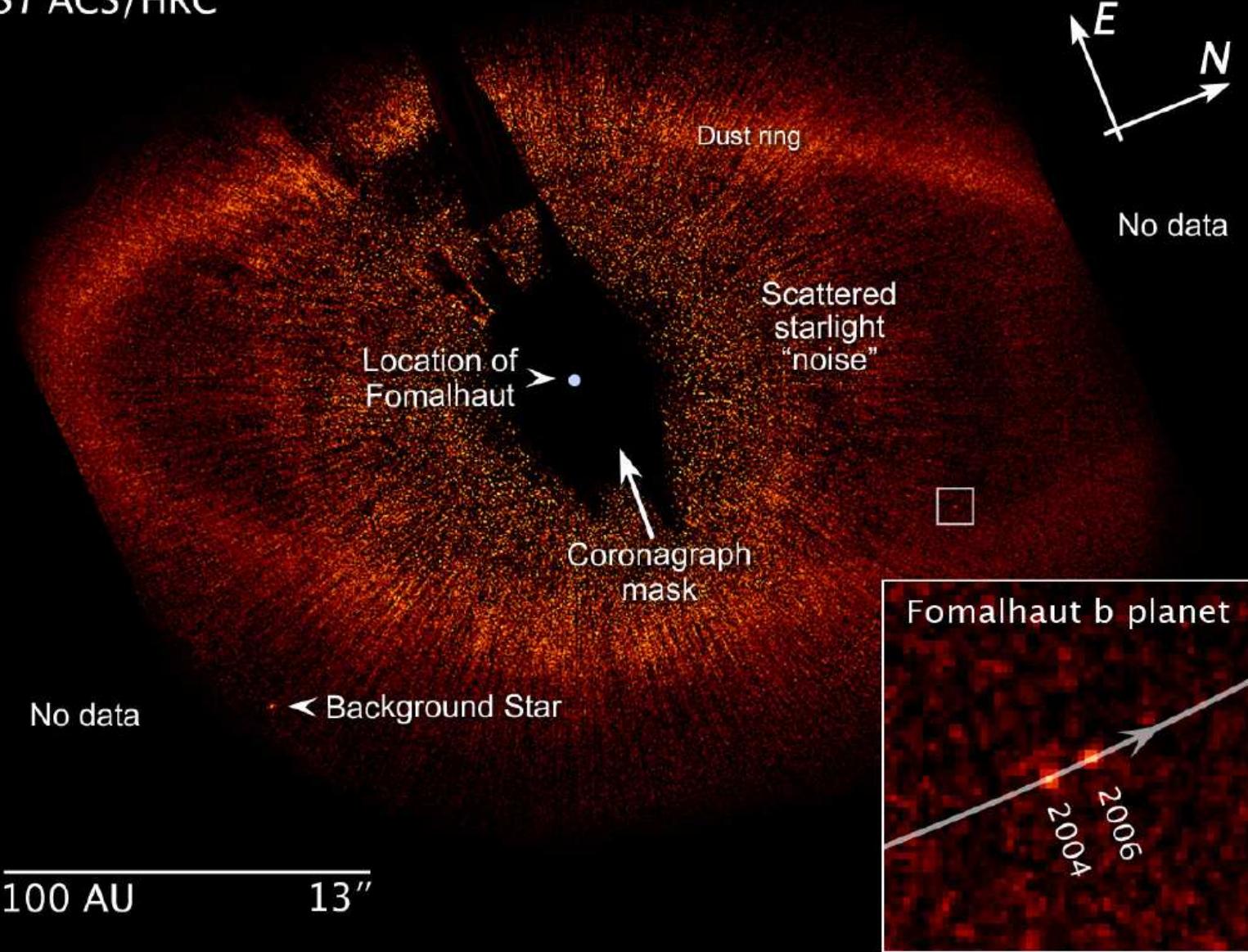




Pillar and Stellar Jet in the Carina Nebula  
HST WFC3/UVIS

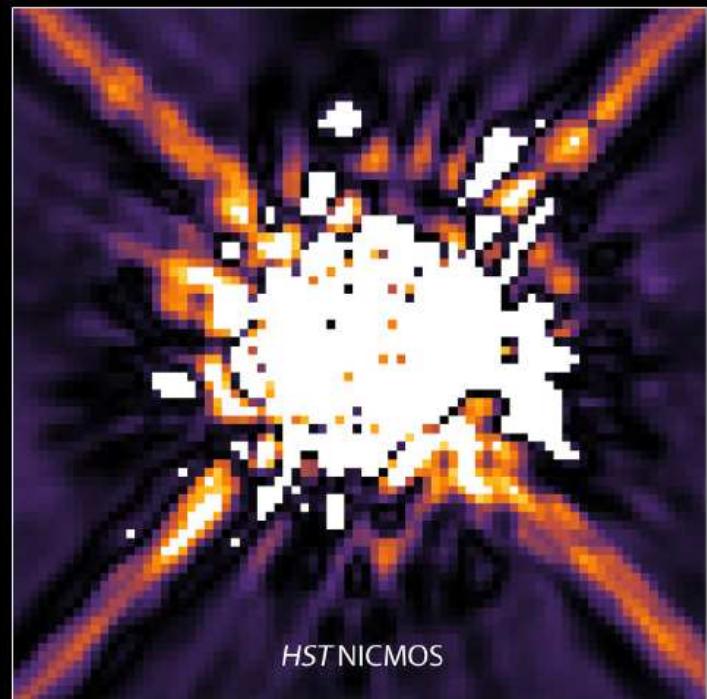


Fomalhaut  
HST ACS/HRC

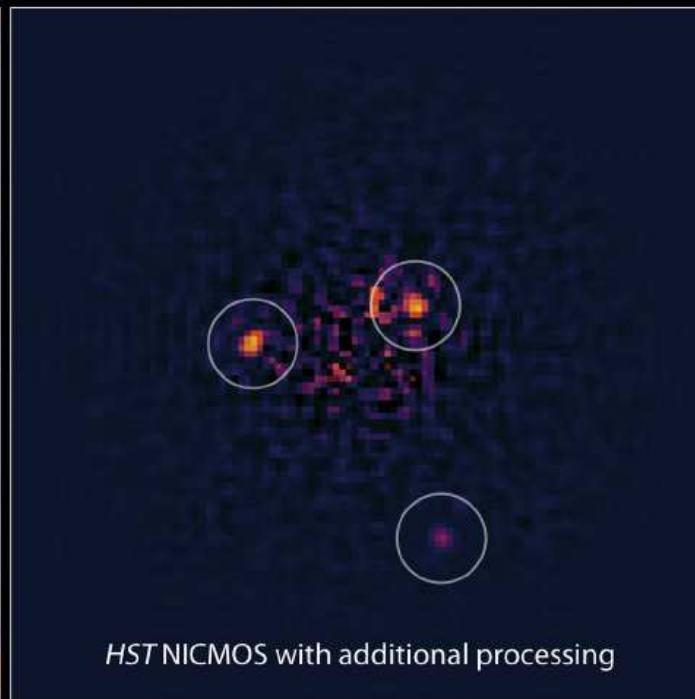


HST/ACS Coronagraph imaging of planetary debris disk around Fomalhaut:  
First direct imaging of a moving planet forming around a nearby star!  
JWST can find such planets much closer in for much farther stars.

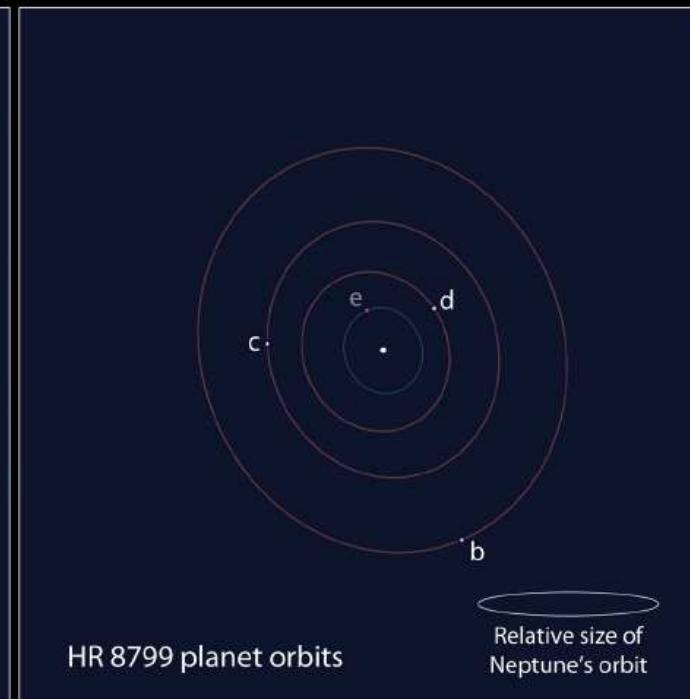
## Exoplanet HR 8799 System



NASA, ESA, and R. Soummer (STScI)



HST NICMOS with additional processing



HR 8799 planet orbits

STScI-PRC11-29

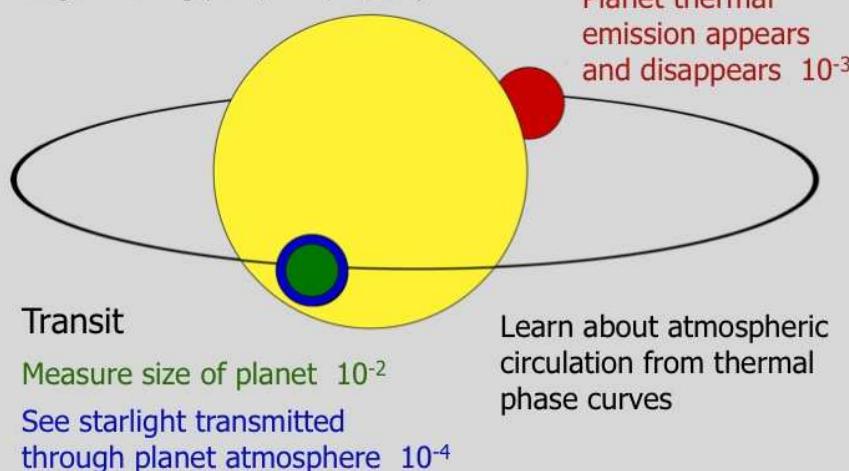
HST/NICMOS imaging of planetary system around the (carefully subtracted) star HR 8799: Direct imaging of planets around a nearby star.

Press release: <http://hubblesite.org/newscenter/archive/releases/2011/29/>

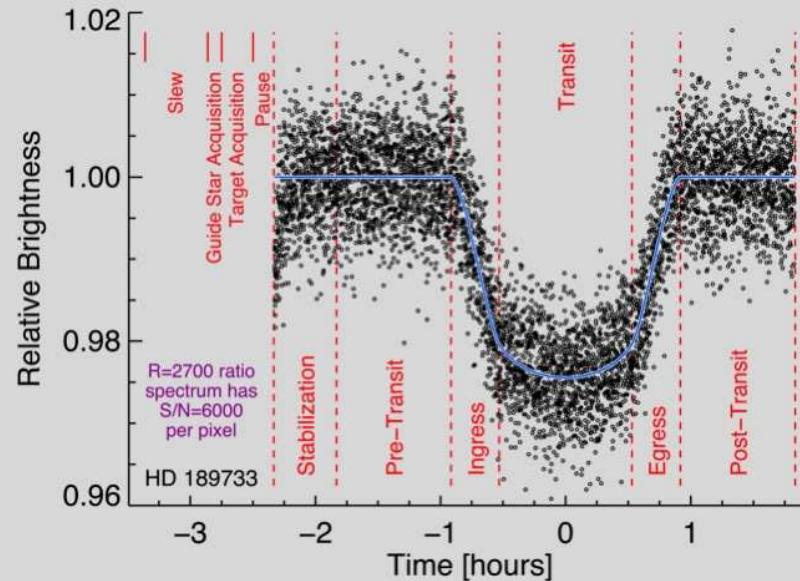
JWST can find such planets much closer in for much farther-away stars.

## Schematic of Transit and Eclipse Science

Seager & Deming (2010, ARAA, 48, 631)



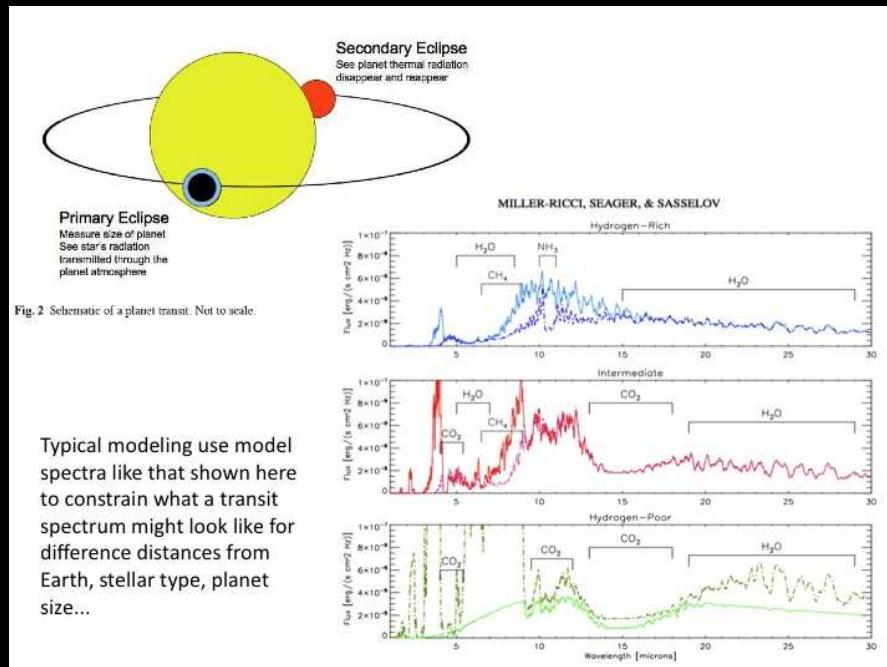
## Timeline of a Transit Observation



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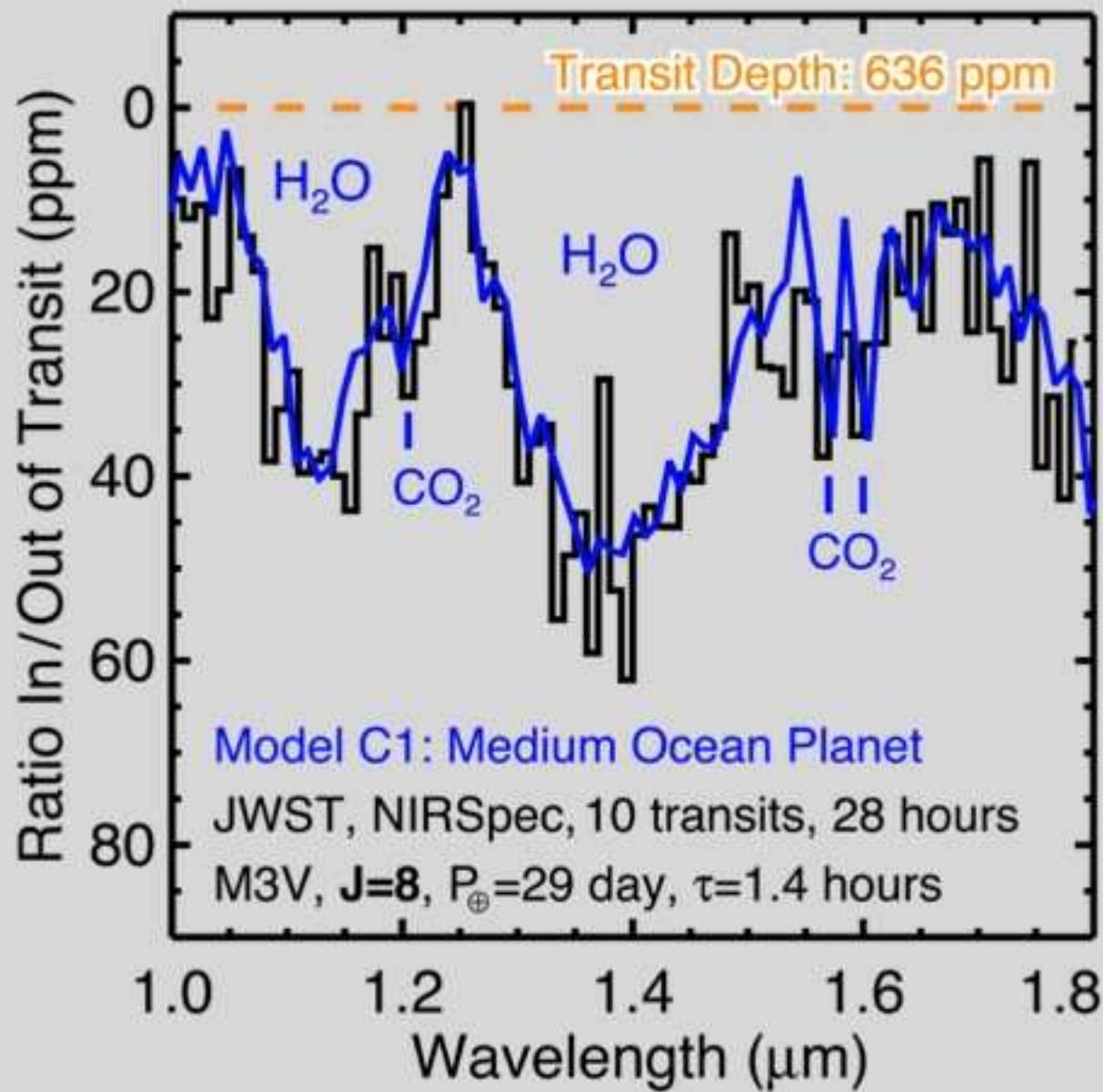
13

JWST can do very precise photometry of transiting Earth-like exoplanets.



JWST IR spectra can find water and CO<sub>2</sub> in (super-)Earth-like exoplanets.

# Transit Spectrum of Habitable “Ocean Planet”



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JWST IR spectra can find water and  $\text{CO}_2$  in transiting Earth-like exoplanets.

B, I, J AB-mag vs.  
half-light radii  $r_e$   
from RC3 to HUDF  
limit are shown.

All surveys limited by  
by SB (+5 mag dash)

Deep surveys bounded  
also by object density.

Violet lines are gxy  
counts converted to  
to natural conf limits.

Natural confusion  
sets in for faintest  
surveys ( $AB \gtrsim 25$ ).  
Will update for JWST.

