



<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1



# Redesign of the Visitor Center Solar Telescope

[002.35.25.05.20-002]

Status: **Draft**

Prepared By	Organization	Date
R. Nguyen	NRAO Electronics Div.	6/21/2023

Approvals	Organization	Signatures
T. Mobraten	NRAO Electronics Division	
S. Wimbrow	NRAO Electronics Division	
D. Schafer	NRAO Electronics Division	

Released By	Organization	Signature
D. Schafer	NRAO Electronics Division	



<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

## Contents

<b>1</b>	<b>Introduction</b>	<b>2</b>
1.1	Purpose . . . . .	2
1.2	Scope . . . . .	2
<b>2</b>	<b>Related Documents and Drawings</b>	<b>2</b>
2.1	Applicable Documents . . . . .	2
2.2	Reference Documents . . . . .	2
<b>3</b>	<b>System Overview</b>	<b>2</b>
3.1	Requirements . . . . .	3
3.2	Architecture <b>(INCOMPLETE)</b> . . . . .	3
3.3	Power Distribution . . . . .	4
<b>4</b>	<b>Tuning System Gain</b>	<b>4</b>
<b>5</b>	<b>Calculating Receiver Input Power</b>	<b>6</b>
5.1	Antenna Power Measurements . . . . .	7
<b>A</b>	<b>S-Parameters <b>TEMPORARY</b></b>	<b>9</b>

## List of Figures

1	RF signal block diagram . . . . .	3
2	51 dB RF gain block . . . . .	4
3	DC power distribution block diagram (Maximum ratings) . . . . .	4
4	Output voltage of the ADL5902 as a function of input power . . . . .	5
5	Optional output scaling circuit. . . . .	5
6	Signal trace with 30 dB of gain from two ZX60-83LN-S+ LNAs. . . . .	7
7	Signal trace with 42 dB of gain and a 7.8 – 8.9 GHz filter. . . . .	8
8	LNA S-parameters. At 8.4 GHz, gain is about 15 dB for the ZX60 and 27 dB for the ALS. . . . .	9
9	Unmarked bandpass filter S-parameters. . . . .	9



<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

## 1 Introduction

### 1.1 Purpose

Located behind the Visitor Center at the Very Large Array (VLA) is a radio telescope that converts RF solar energy into a DC voltage through amplification. It was originally constructed in 2013 and is no longer functioning as of June 2023. Weathered components and lack of sufficient documentation to repair the telescope to its original state has warranted a redesign of the entire system to be more appealing to visitors and resistant to wear.

### 1.2 Scope

This document will describe the design choices made in development of the new solar telescope. Calculations and experiments to support design decisions will be demonstrated along with constraints and requirements.

## 2 Related Documents and Drawings

### 2.1 Applicable Documents

The following documents may not be directly referenced herein, but may provide necessary context or supporting material.

Ref. No.	Document Title	Rev/Doc. No.
AD01	Desiderata for Solar Observing with the EVLA	EVLAM_70
AD02	EVLA Hardware Modifications in Support of Solar Observing	EVLAM_72

### 2.2 Reference Documents

The following documents are referenced within this text:

Ref. No.	Document Title	Rev/Doc. No.
RD01	Solar Brightness Temperature and Corresponding Antenna Noise Temperature at Microwave Frequencies	<a href="#">Filehost</a>
RD02	Antenna Engineering Handbook	DSOC Bookcase
RD03	Tools of Radio Astronomy	Privately Owned
RD04	X-Band System Performance of the Very Large Array	<a href="#">Filehost</a>
RD05	10-60 GHz G/T Measurements Using the Sun as a Source—A Preliminary Study	<a href="#">Filehost</a>

## 3 System Overview

The heart of the telescope is an conical corrugated X-band horn antenna originally used to receive Voyager transmissions at 8.4 GHz; this is to be unchanged in the redesign. Power is supplied through a 15 volt linear DC power supply located inside the visitor center in order to eliminate radio frequency interference (RFI) caused by AC power and switching power supplies.

<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

As the telescope is pointed towards the sun, RF energy is received by the antenna and amplified. A rectifying circuit converts the RF power into a DC voltage and further amplifies it before outputting to a voltmeter. The DC gain should be adjusted so that as the telescope is pointed towards the sun with clear skies, the voltmeter should read a maximum 3 volts; as the telescope is pointed towards the ground or away from the sun, the voltmeter reading should decrease to zero.

### 3.1 Requirements

Being a visitor center attraction, the device must be engaging and appealing to visitors. In order to achieve this, the device must be at least partially transparent to show visible electronics and have a voltmeter that is driven by solar energy. The enclosure of the previous design was made of some sort of polycarbonate material which suffered from yellowing in the sun. Additionally, small computer fans with a mesh screen mounted directly to the enclosure were used to decrease operating temperature but allowed dust to accumulate over the years. In essence, the new design goals are as follows:

- (1) Utilizes a transparent enclosure or panel to view electronics, ideally one which will not yellow or fade easily.
- (2) Must be dustproof and resilient to accumulation of debris inside the enclosure while still venting heat from electronics.
- (3) Electronics must be powered by a fixed DC linear power supply and be easily serviceable.

### 3.2 Architecture (INCOMPLETE)

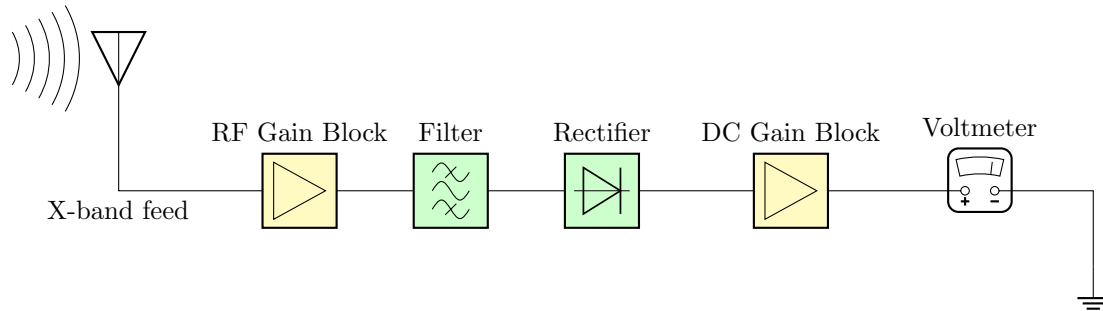


Figure 1: RF signal block diagram

Many of the RF components were either reused from the previous design or gathered at no cost. The RF gain block shown in Fig. 1 and 2 utilizes 3 low noise amplifiers (LNAs) and an attenuator. The first 2 LNAs are Mini-Circuits ZX60-83LN-S+ which provide a measured 15dB of gain at 8.4 GHz each; they are followed by a single 6 dB attenuator and an ALS-04-0149 LNA in that order. The ALS-04-0149 does not have a readily available datasheet but measured 27dB of gain at 8.4 GHz.

The bandpass filter does not have an attached datasheet but it resembles an equiripple filter and measured a 3 dB passband of about 7.8–8.9 GHz.

The rectifier and DC gain block are integrated into a PCB with SMA connectors. The former utilizes an Analog Devices ADL5902, providing a linear-in-decibel output that can be scaled by the

<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

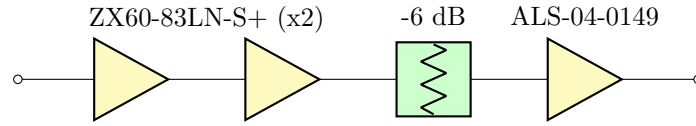


Figure 2: 51 dB RF gain block

choice of resistors on the board or the following gain element. The DC gain block utilizes an LM358 op amp with adjustable gain via potentiometer.

### 3.3 Power Distribution

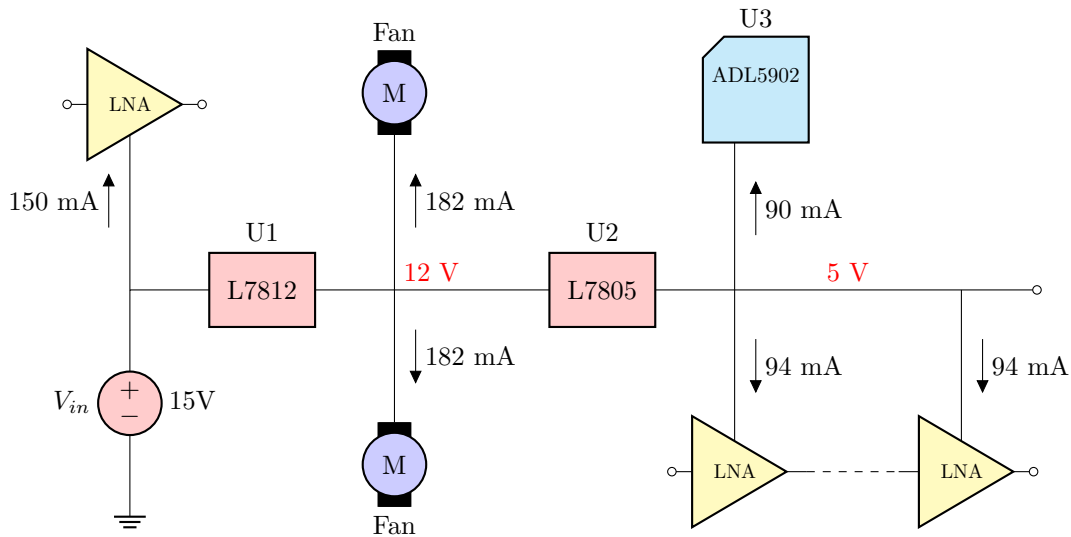


Figure 3: DC power distribution block diagram (Maximum ratings)

As part of the effort to make the telescope more modern and visually appealing, the protoboard on the previous design was replaced with a dual layer PCB which contains the voltage regulators, the RF power detector IC, and dual noninverting op amps. The 15 volt supply is used to bias the ALS-04-0149, 12 volts is used to power the fans, whereas 5 volts is used to bias the ADL5902 and the ZX60-83LN-S+ LNAs. Figure 3 shows the maximum current draw of each component to ensure that the system does not exceed the 1.5 A rating of the L78XX chips.

## 4 Tuning System Gain

Figure 4 shows the voltage-power curve of the ADL5902 at the telescope's operating frequency. Pointing the telescope at the sun will produce an input power higher than pointing it at the ground; this dynamic range is represented by the dotted lines and filled section of the curve. Adding attenuation or RF gain will shift this area to the left or right respectively. Adding DC gain

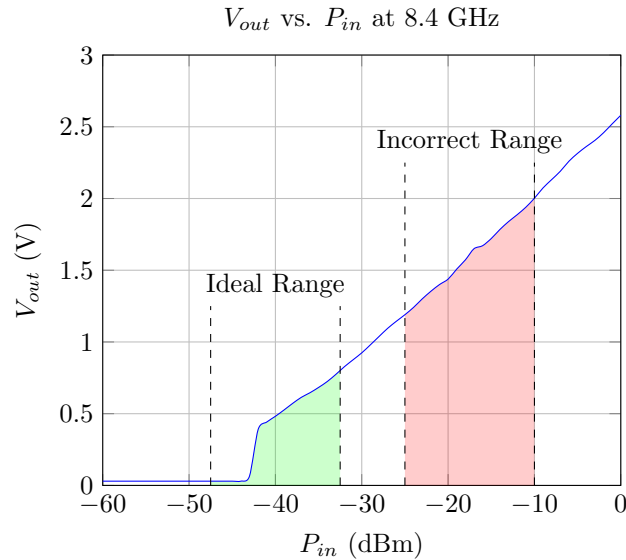


Figure 4: Output voltage of the ADL5902 as a function of input power

will increase the slope of this curve and vice versa. Ideally, the RF gain should be adjusted such that the low range is roughly 0 volts whereas the the high range is scaled to 3 volts by the DC gain block. In this case, the amplification from the DC gain block will have negligible effect on the low range. If, for some reason, additional attenuation is not possible or DC gain causes the low range to increase substantially above 0 volts, a simple solution is presented in Fig. 5. DC gain should

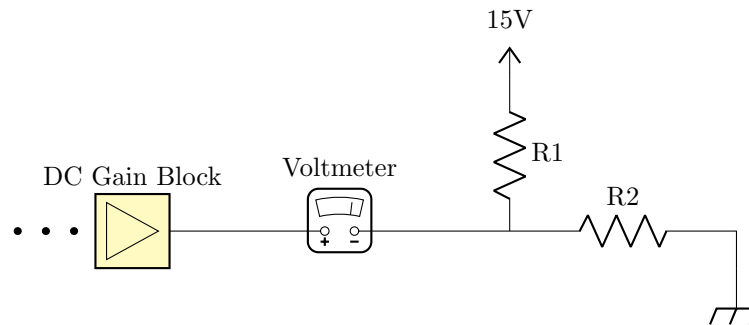


Figure 5: Optional output scaling circuit.

provide an output dynamic range of at least 3 volts, then the negative terminal of the voltmeter can be set to reference a nonzero voltage i.e the low range of the adjusted V-P curve. Note that this is limited to a maximum of 15 volts.

<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

## 5 Calculating Receiver Input Power

The first step to determine RF gain will be to find the range of antenna input power when pointed at the sun; note that this may vary daily and annually depending on solar cycle[RD01, Fig. 2]. Instantaneous power received by the antenna from a celestial source in watts,  $P_R$ , is defined by Eq. 1

$$P_R = S A_e \Delta f \quad (1)$$

where  $S$  is the source flux density in  $\text{W m}^{-2} \text{Hz}^{-1}$ ,  $A_e$  is the effective aperture area in  $\text{m}^2$ , and  $\Delta f$  is the receiver bandwidth in Hz[RD02, eq. (41-2)]. Effective aperture can be calculated using measurements of the antenna as shown in Eq. 2

$$\begin{aligned} A_e &= \frac{\lambda^2}{4\pi} G \\ &= \frac{\pi d^2}{4} \eta_a \end{aligned} \quad (2)$$

where  $G$  is linear antenna gain,  $d$  is the diameter of the circular horn at the aperture, and  $\eta_a$  is the unitless aperture efficiency[RD02, pg. (15-27)][RD03, eq. (5.58)]. An ideal horn antenna has an  $\eta_a$  of 0.522. The calculations for this design will use an approximate value of 0.5<sup>1</sup>. The antenna diameter at the opening measured to be 34 cm, resulting in an effective area of 0.0454  $\text{m}^2$  or -13.4 dB  $\text{m}^2$ .

Receiver bandwidth  $\Delta f$  can be defined by the bandwidth of the bandpass filter in Fig. 1, thus, the received power  $P_R$  is more representative of channel power without any amplification, rather than antenna output power, as the antenna is a wideband feed with a higher bandwidth than the filter. In this case, receiver bandwidth is defined to be 1100 MHz or 90.4 dB Hz.

Lastly, values for solar flux density  $S$  must be found within the operating frequency. Ho *et al.*[RD01] measured mean, minimum, and maximum values for solar flux density at 2800 MHz and 8800 MHz but only numbers for the former are shown in the text. However, they claim that solar flux density at 8800 MHz are higher by a factor of 2.17 on average, thus, the mean, minimum, and maximum values at 2800 MHz can be multiplied by 2.17 to achieve approximate values for  $S$ . Additionally, solar brightness temperature shown in [RD01, Fig. 5] can be used to calculate solar flux density via [RD01, eq. (2)]. The former method shows higher dynamic range and is shown in Table 5 in solar flux units<sup>2</sup>.

	$S_{\min}$	$S_{\mu}$	$S_{\max}$
2.8 GHz	70 SFU	150 SFU	280 SFU
8.8 GHz	152 SFU	326 SFU	608 SFU

Table 1: Solar Flux Density  $S$  at 8.8 GHz extrapolated by multiplying  $S$  in [RD01] at 2.8 GHz by a factor of 2.17.

With values for  $S$ ,  $A_e$ , and  $\Delta f$  found, we can now calculate the antenna channel power  $P_R$  with equation 1 to be -87.9 dBm at  $S_{\mu}$ . Calculating  $P_R$  at  $S_{\min}$  and  $S_{\max}$  show a roughly  $\pm 3$  dBm variation in  $P_R$ . Note that this calculation assumes that solar flux density is constant across the entire frequency range of the receiver. This is a limitation of available information on solar flux density; in practice,  $S$  generally increases with frequency[RD05].

<sup>1</sup>The x-band horn measured an estimated aperture efficiency of  $0.62 \pm 0.03$  but only when mounted to the VLA dish[RD04]. This design will assume the antenna to be a standard horn antenna.

<sup>2</sup>One solar flux unit (SFU) is equal to  $10^{-22} \text{ W m}^{-2} \text{Hz}^{-1}$ .

<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

## 5.1 Antenna Power Measurements

In order to confirm these calculations, the antenna was connected to an Anritsu Model MS2720T/720 Spectrum Analyzer centered at 8.4 GHz spanning 100 MHz and pointed towards the sun.

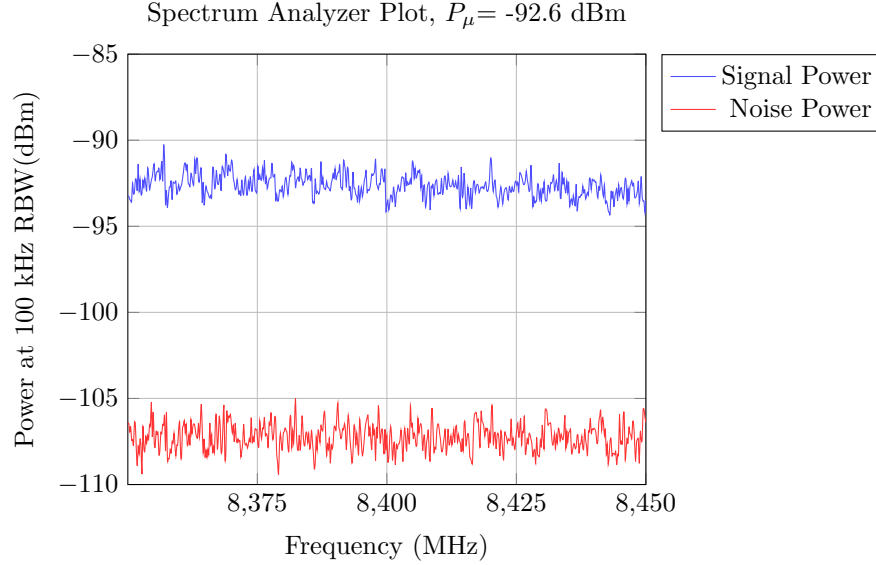


Figure 6: Signal trace with 30 dB of gain from two ZX60-83LN-S+ LNAs.

The Anritsu is not capable of performing channel power measurements, so, assuming the mean signal power  $P_\mu$  in Fig. 6 can be extrapolated to the entire receiver bandwidth, channel power is calculated as follows:

$$\begin{aligned}
 P_R &= P_\mu - G_{dB} - (RBW)_{dB} + (BW)_{dB} \\
 &= -92.6 \text{ dBm} - 30 - 50 + 90.4 \\
 &= -82.2 \text{ dBm}
 \end{aligned}$$

This measurement is about 5 dB greater than calculated but is within the degree of accuracy necessary for this design. It's possible this error was caused by assuming equal power throughout the entire passband, or due to the fact that this operating frequency is outside the bandwidth of the amplifiers used, thus resulting in nonlinear gain across frequency. Note that the noise power in Fig. 6 was measured with a  $50\Omega$  load at the spectrum analyzer input and no amplifiers.

In another experiment on a different day, traces were taken with the telescope pointed at both the sun and ground in a cloudy and sunny condition. This time with 42 dB of gain from a ZX60-83LN-S+ and an ALS-04-0149 in that order, along with the 7.8 – 8.9 GHz bandpass filter between them.





<b>Title:</b> Redesign of the Visitor Center Solar Telescope	<b>Owner:</b> Remy Nguyen	<b>Date:</b> June 21, 2023
<b>NRAO Doc #:</b> [002.35.25.05.20-002]		<b>Version:</b> 1

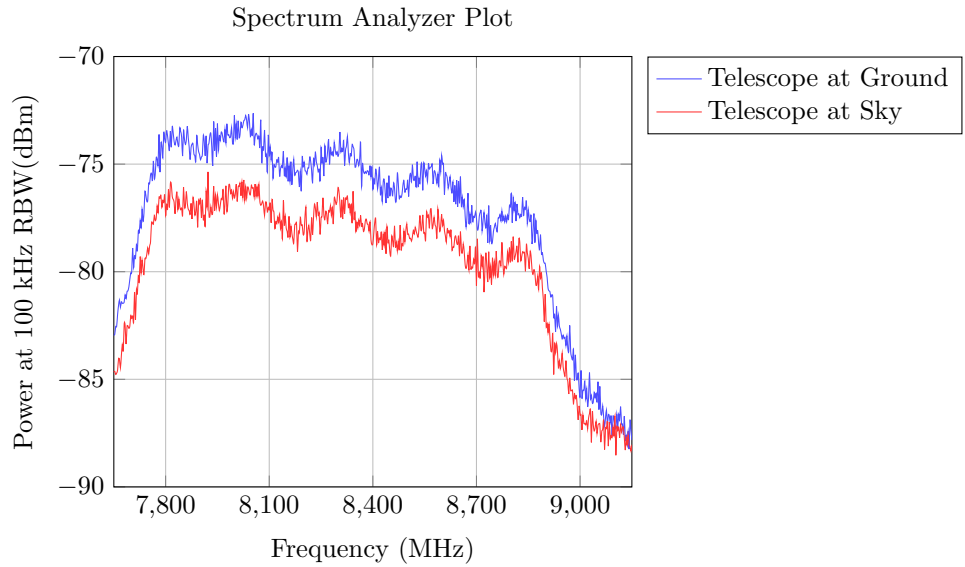


Figure 7: Signal trace with 42 dB of gain and a 7.8 – 8.9 GHz filter.

Figure 7 shows a roughly 2 dB drop in power when the telescope moved from the ground towards the horizon/sky; this power reading also did not change when pointing directly at the sun versus anywhere else in the sky. Additionally, the right and left-hand side ports on the antenna feed did not make a difference to the readings. The ripples in the trace are caused by the filter response and the downward slope is due to reaching the end of each LNA's operating frequency.

Power calculations conducted in the same manner as the trace in Fig. 6 show a channel power of about -80 dB in the passband.

## A S-Parameters **TEMPORARY**

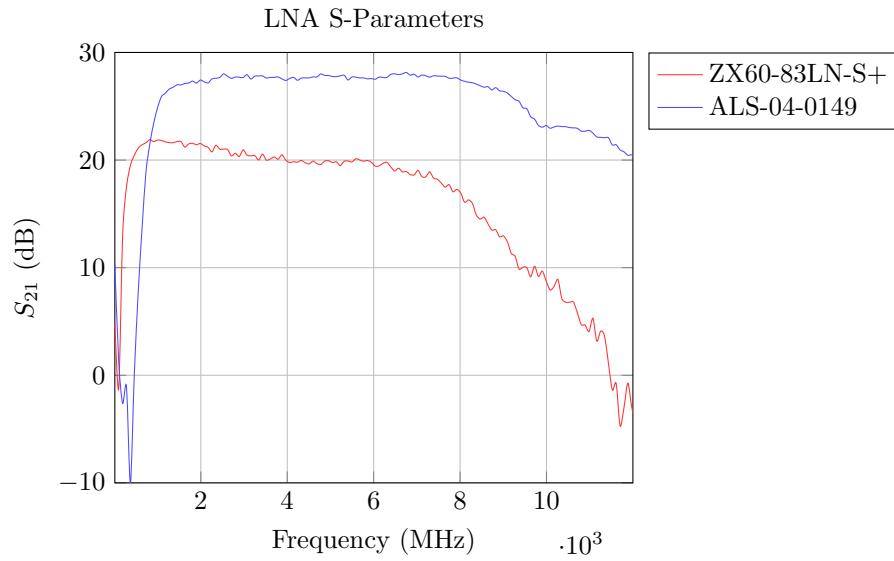


Figure 8: LNA S-parameters. At 8.4 GHz, gain is about 15 dB for the ZX60 and 27 dB for the ALS.

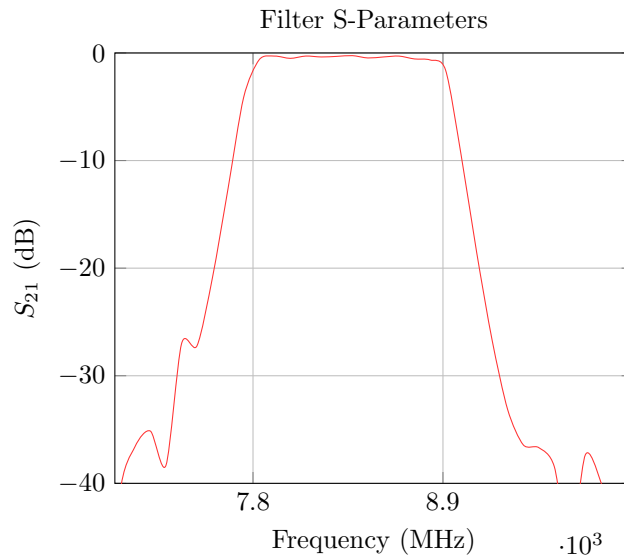


Figure 9: Unmarked bandpass filter S-parameters.