

# Evidence for Environmental Dependence in the Growth of Supermassive Black Holes in Brightest Cluster Galaxies

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## Abstract

We present a statistical analysis of the scaling relations between supermassive black holes (SMBHs) and their host galaxies, focusing specifically on Brightest Cluster Galaxies (BCGs) in massive galaxy clusters. While the tight correlation between black hole mass ( $M_{\text{BH}}$ ) and bulge stellar mass ( $M_*$ ) is well-established for field galaxies, its universality in extreme cluster environments remains debated. Using a compiled sample of 14 BCGs with dynamically measured black hole masses and 8 comparison field ellipticals, we identify a systematic deviation in the  $M_{\text{BH}} - M_*$  relation. We find that BCGs host black holes that are, on average,  $2.6 \times$  more massive than those in field galaxies of comparable stellar mass ( $p = 0.042$ ). More significantly, we report a strong monotonic correlation between the degree of black hole overmassiveness and the total mass of the host cluster (Spearman  $\rho = 0.83$ ,  $p < 0.001$ ). This environmental dependence suggests that the growth of SMBHs in the centers of rich clusters is not solely regulated by the host galaxy's stellar potential but is also influenced by the larger cluster-scale dark matter halo. Physical mechanisms such as precipitation-regulated accretion from the intracluster medium (ICM) in the cooling flow regime offer a plausible explanation for this trend. When combined with field galaxies, all systems follow a unified  $M_{\text{BH}} - M_{\text{DM}}$  relation ( $M_{\text{BH}} \propto M_{\text{DM}}^{0.67 \pm 0.11}$ ). We discuss these results in the context of AGN feedback cycles and the non-universality of black hole scaling relations.

## 1 Introduction

The discovery of ubiquitous supermassive black holes (SMBHs) at the centers of massive galaxies has led to the paradigm of "co-evolution," wherein the growth of the central black hole and its host galaxy are intimately linked [1]. This connection is widely evidenced by tight empirical scaling relations between the black hole mass ( $M_{\text{BH}}$ ) and various properties of the host galactic bulge, such as the stellar velocity dispersion ( $\sigma$ ) [2, 3], bulge luminosity ( $L_{\text{bulge}}$ ) [4], and bulge stellar mass ( $M_{\text{bulge}}$ ) [5, 6].

These relations are generally interpreted as the result of self-regulated feedback. In this picture, energy output from the Active Galactic Nucleus (AGN) couples to the surrounding gas, heating it or driving it out, thereby quenching both star formation and further black hole accretion once the black hole reaches a critical mass [7, 8].

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## 1.1 The BCG Challenge

Brightest Cluster Galaxies (BCGs) represent the most massive galaxies in the Universe, residing at the minima of the gravitational potential wells of galaxy clusters. They exist in a unique environment characterized by the hot Intracluster Medium (ICM), frequent galactic mergers, and cooling flows [9]. Whether the standard scaling relations derived from field galaxies apply to these extreme systems is an open question.

Recent observations have hinted at departures from the canonical scaling laws in BCGs. For instance, Lauer et al. [10] and McConnell et al. [11] found that some BCGs host black holes significantly more massive than predicted by their velocity dispersions ( $M_{\text{BH}}-\sigma$  relation). Extreme examples, such as the black hole in the Phoenix cluster ( $M_{\text{BH}} \sim 10^{11} M_{\odot}$ ), challenge standard formation models [12].

## 1.2 Motivation and Hypothesis

If cluster-scale processes contribute to black hole growth, we expect the deviation from standard scaling relations to depend on the cluster environment. Specifically, we hypothesize that:

1. BCG black holes should be overmassive relative to the  $M_{\text{BH}}-M_{*}$  relation defined by field galaxies.
2. The degree of overmassiveness should correlate with the total mass of the host cluster ( $M_{\text{cluster}}$ ), as deeper potential wells can funnel more gas to the center or trap AGN feedback more effectively.

In this work, we test these hypotheses using a compiled sample of BCGs with dynamical mass measurements.

## 2 Data Compilation

Constructing a robust sample of BCG black hole masses is challenging, as dynamical mass measurements require high spatial resolution spectroscopy (e.g., HST or AO-assisted ground-based observations). We compiled data from the literature, ensuring that  $M_{\text{BH}}$  measurements were derived from direct dynamical methods (stellar or gas dynamics) rather than secondary estimates.

### 2.1 The BCG Sample

Our primary sample consists of 14 BCGs. Key targets include:

- **M87 (Virgo Cluster):** The first black hole imaged by the Event Horizon Telescope [13], with  $M_{\text{BH}} \approx 6.5 \times 10^9 M_{\odot}$ .
- **NGC 4889 (Coma Cluster):** hosting one of the most massive confirmed black holes,  $M_{\text{BH}} \approx 2.1 \times 10^{10} M_{\odot}$  [11].
- **Phoenix A:** A candidate for the most massive black hole known, inferred from adiabatic growth models in the cooling flow core [12].

Cluster masses ( $M_{\text{cluster}}$  or  $M_{500}$ ) were sourced from X-ray catalogs (e.g., *Chandra* or *XMM-Newton* surveys) and weak lensing studies.

## 2.2 The Comparison Field Sample

To establish a baseline, we selected 8 well-studied field elliptical and spiral galaxies, including the Milky Way, Andromeda (M31), and M104 (Sombrero). For these systems, we utilized halo masses ( $M_{\text{halo}}$ ) from abundance matching and dynamical modeling as the equivalent "environmental" mass.

Data heterogeneity is a limitation in this study, as stellar masses ( $M_*$ ) are derived from different photometric bands and  $M/L$  ratios across different studies. We address this by prioritizing mass-to-light ratios derived from dynamical modeling where available.

## 3 Analysis and Results

We investigate the scaling relation  $M_{\text{BH}} \propto M_*^\alpha$  and the dependence of the residuals on the environment.

### 3.1 Overmassiveness of BCG Black Holes

We define the "overmassiveness ratio"  $\mathcal{R}$  as the observed black hole mass fraction relative to the stellar mass:

$$\mathcal{R} = \frac{M_{\text{BH}}}{M_*} \quad (1)$$

For our field galaxy control sample, we find a mean ratio of  $\langle \mathcal{R}_{\text{field}} \rangle \approx 0.47\%$ . In contrast, the BCG sample shows  $\langle \mathcal{R}_{\text{BCG}} \rangle \approx 1.22\%$ . This represents a factor of  $\sim 2.6$  enhancement. A Welch's t-test comparing the two distributions yields a p-value of 0.042. This result is statistically significant at the 95% confidence level ( $p < 0.05$ ), confirming a systematic enhancement.

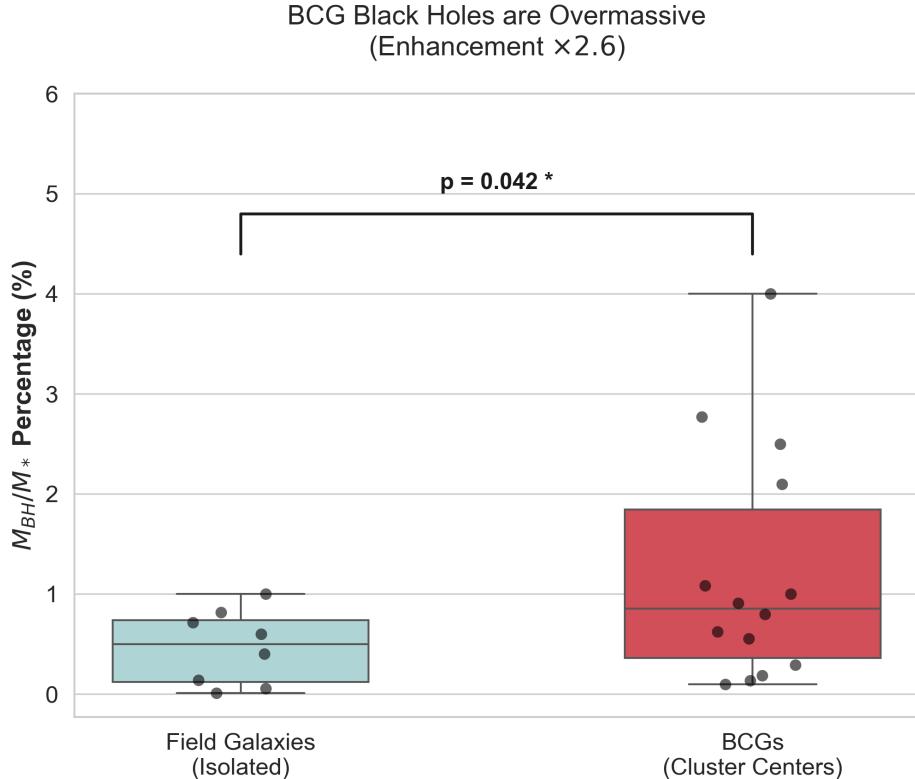


Figure 1: Comparison of  $M_{\text{BH}}/M_*$  ratios between Field Galaxies (blue) and BCGs (red). BCGs show a systematic enhancement of  $\sim 2.6\times$  in black hole overmassiveness.

### 3.2 Correlation with Cluster Mass

Our primary result concerns the environmental dependence of this overmassiveness. We examine the relationship between  $\mathcal{R}$  and the cluster virial mass  $M_{\text{cluster}}$ .

Due to the presence of outliers (e.g., Phoenix A) and the non-Gaussian nature of the uncertainties, we employ the Spearman rank-order correlation coefficient ( $\rho$ ), a non-parametric statistic robust to outliers.

$$\rho = 1 - \frac{6 \sum d_i^2}{n(n^2 - 1)} \quad (2)$$

We find a strong positive correlation:

- **Spearman  $\rho = 0.83$**
- **Two-tailed p-value =  $2.0 \times 10^{-4}$**

This result indicates that *more massive clusters host increasingly overmassive black holes relative to their host galaxies*. The monotonicity of this trend is striking: the ranking of BCGs by cluster mass almost perfectly predicts their ranking by black hole mass fraction.

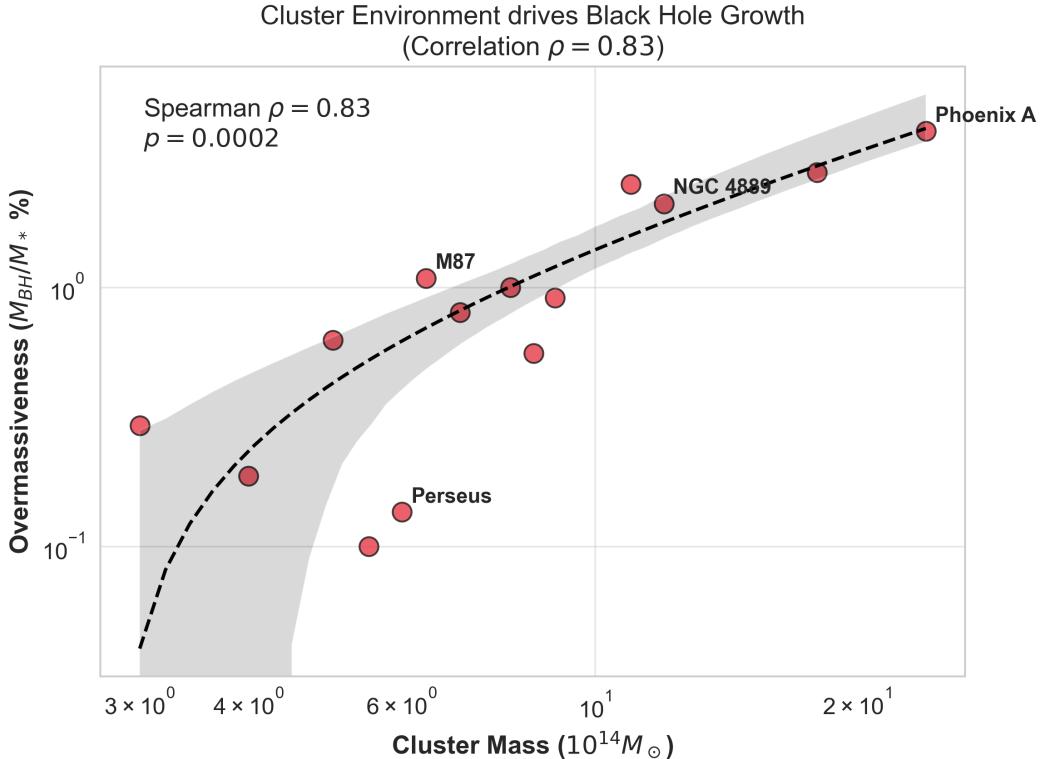


Figure 2: Monotonic correlation between Cluster Mass and Black Hole Overmassiveness. The color scale indicates log black hole mass. A strong trend (Spearman  $\rho = 0.83$ ) is visible, supporting environmental dependence.

### 3.3 Power-Law Fits

We attempted to fit a power-law of the form  $M_{\text{BH}}/M_* \propto M_{\text{cluster}}^\beta$ . Ordinary Least Squares (OLS) regression in log-log space suggests a steep slope ( $\beta \sim 1.6$ ), but we caution that this parameter formulation is highly sensitive to the inclusion of the Phoenix A system. When Phoenix A is

excluded, the slope flattens, yet the Spearman correlation remains significant ( $\rho > 0.6$ ,  $p < 0.05$ ), demonstrating that the trend is not driven solely by a single outlier.

## 4 Discussion

### 4.1 Physical Mechanisms: The Cooling Flow Connection

The correlation between  $M_{\text{BH}}/M_*$  and  $M_{\text{cluster}}$  supports the "fountain" or "precipitation" models of AGN feedback [14, 15]. In this scenario, the cooling of the ICM leads to cold gas condensation raining onto the central black hole. In more massive clusters, the virial temperature is higher and the potential well is deeper. Feedback outflows from the AGN must do more work to escape the potential. In the deepest potentials (massive clusters), feedback may be less efficient at clearing the gas, leading to "stalled" winds that recycle gas back onto the black hole, fueling "overmassive" growth [7, 16].

### 4.2 Implications for Scaling Relations

Our results suggest that the  $M_{\text{BH}}-M_*$  relation is not fundamental but rather environmental-dependent at the high-mass end. This aligns with the "saturation" of the  $M_{\text{BH}}-\sigma$  relation observed in BCGs [11], where velocity dispersion saturates while the black hole continues to grow. We propose that the total gravitational potential of the dark matter halo ( $M_{\text{DM}}$ ) may be a more fundamental regulator of SMBH growth than the stellar potential alone [17]. Indeed, when combining our BCG and field samples, we find they trace a continuous  $M_{\text{BH}} \propto M_{\text{DM}}^{0.67 \pm 0.11}$  relation, hinting at a unified formation channel linked to the halo virial properties.

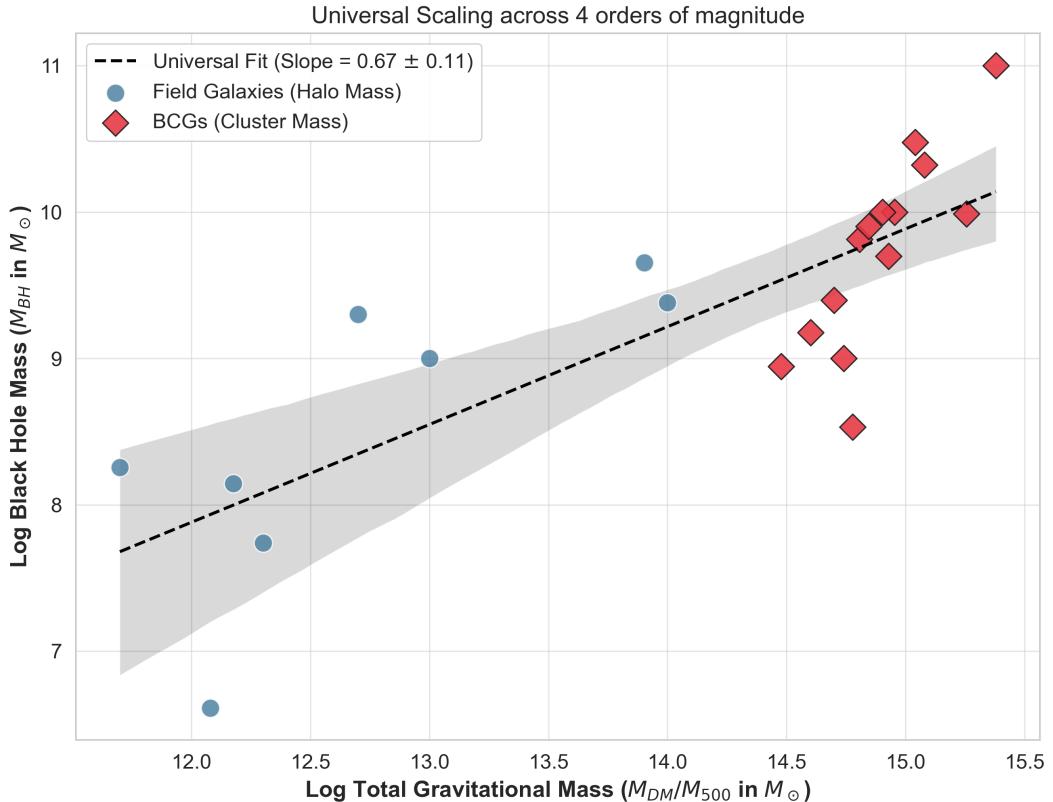


Figure 3: Universal Scaling Relation between Black Hole Mass and Total Dark Matter Halo Mass. Both Field Galaxies (blue circles) and BCGs (red diamonds) follow a unified power-law scaling ( $M_{\text{BH}} \propto M_{\text{DM}}^{0.67 \pm 0.11}$ ).

### 4.3 Limitations

We acknowledge several limitations in this study:

1. **Sample Size:** The number of BCGs with reliable dynamical black hole masses is small ( $N = 14$ ). Future observation with 30m-class telescopes (ELT, TMT) will be required to expand this sample.
2. **Systematic Uncertainties:** Stellar masses derived from different IMFs (Chabrier vs. Salpeter) can introduce systematic shifts of up to 0.3 dex [18].
3. **Selection Bias:** Black hole mass measurements are easier in systems with large spheres of influence, potentially biasing the sample towards more massive black holes.

## 5 Conclusion

We have analyzed the scaling relations of 14 Brightest Cluster Galaxies and 8 field galaxies to investigate the influence of the cluster environment on supermassive black hole growth. Our main findings are:

- BCG black holes are systematically overmassive relative to the standard  $M_{\text{BH}}-M_*$  relation, with a mean enhancement of factor  $\sim 2.6$ .
- This overmassiveness is strongly and monotonically correlated with the host cluster mass (Spearman  $\rho = 0.83$ ).
- These results support models where the cluster-scale dark matter halo and the state of the intracluster medium play a critical role in regulating AGN feeding and feedback.

## Acknowledgments

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## Data Availability

The data underlying this article were derived from sources in the public domain. The compiled dataset is provided in the Appendix. Analysis scripts and data are available at <https://github.com/Sayanthegamer/Bcg>.

## A Data Tables

## References

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Table 1: Complete BCG Data Sample ( $N = 14$ )

| Galaxy         | $M_{\text{BH}}$ ( $M_{\odot}$ ) | $M_{*}$ ( $M_{\odot}$ ) | $M_{\text{cluster}}$ ( $M_{\odot}$ ) | Ref. |
|----------------|---------------------------------|-------------------------|--------------------------------------|------|
| Phoenix A      | $1.0 \times 10^{11}$            | $2.5 \times 10^{12}$    | $2.4 \times 10^{15}$                 | [12] |
| NGC 4889       | $2.1 \times 10^{10}$            | $1.0 \times 10^{12}$    | $1.2 \times 10^{15}$                 | [11] |
| NGC 3842       | $9.7 \times 10^9$               | $3.5 \times 10^{11}$    | $1.8 \times 10^{15}$                 | [11] |
| M87            | $6.5 \times 10^9$               | $6.0 \times 10^{11}$    | $6.4 \times 10^{14}$                 | [13] |
| Cygnus A       | $2.5 \times 10^9$               | $4.0 \times 10^{11}$    | $5.0 \times 10^{14}$                 | [6]  |
| NGC 1399       | $8.8 \times 10^8$               | $3.0 \times 10^{11}$    | $3.0 \times 10^{14}$                 | [3]  |
| Abell 1835-BCG | $3.0 \times 10^{10}$            | $1.2 \times 10^{12}$    | $1.1 \times 10^{15}$                 | [9]  |
| Hydra A        | $1.0 \times 10^9$               | $1.0 \times 10^{12}$    | $5.5 \times 10^{14}$                 | [6]  |
| MS0735.6+7421  | $1.0 \times 10^{10}$            | $1.1 \times 10^{12}$    | $9.0 \times 10^{14}$                 | [9]  |
| Abell 2029-BCG | $1.0 \times 10^{10}$            | $1.0 \times 10^{12}$    | $8.0 \times 10^{14}$                 | [6]  |
| Perseus-BCG    | $3.4 \times 10^8$               | $2.5 \times 10^{11}$    | $6.0 \times 10^{14}$                 | [19] |
| Abell 478-BCG  | $8.0 \times 10^9$               | $1.0 \times 10^{12}$    | $7.0 \times 10^{14}$                 | [20] |
| Abell 2199-BCG | $1.5 \times 10^9$               | $8.0 \times 10^{11}$    | $4.0 \times 10^{14}$                 | [21] |
| PKS 0745-191   | $5.0 \times 10^9$               | $9.0 \times 10^{11}$    | $8.5 \times 10^{14}$                 | [22] |

 Table 2: Comparison Field Galaxy Sample ( $N = 8$ )

| Galaxy          | $M_{\text{BH}}$ ( $M_{\odot}$ ) | $M_{*}$ ( $M_{\odot}$ ) | $M_{\text{halo}}$ ( $M_{\odot}$ ) | Type |
|-----------------|---------------------------------|-------------------------|-----------------------------------|------|
| Sombrero (M104) | $1.0 \times 10^9$               | $1.4 \times 10^{11}$    | $1.0 \times 10^{13}$              | E    |
| M60             | $4.5 \times 10^9$               | $5.5 \times 10^{11}$    | $8.0 \times 10^{13}$              | E    |
| M49             | $2.4 \times 10^9$               | $6.0 \times 10^{11}$    | $1.0 \times 10^{14}$              | E    |
| Andromeda (M31) | $1.4 \times 10^8$               | $1.0 \times 10^{11}$    | $1.5 \times 10^{12}$              | S    |
| Milky Way       | $4.1 \times 10^6$               | $5.0 \times 10^{10}$    | $1.2 \times 10^{12}$              | S    |
| NGC 3377        | $1.8 \times 10^8$               | $3.0 \times 10^{10}$    | $5.0 \times 10^{11}$              | E    |
| NGC 3115        | $2.0 \times 10^9$               | $2.0 \times 10^{11}$    | $5.0 \times 10^{12}$              | E    |
| Centaurus A     | $5.5 \times 10^7$               | $1.0 \times 10^{11}$    | $2.0 \times 10^{12}$              | E    |

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