

Project 3

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1 Method

1.1 Forward Euler

The forward Euler method is a algorithm to estimate the solution of a differential equation. The Forward Euler method wants to find the next point. To find the next point, it uses the point it is at, r_n , a small time step, dt , and the derivative of its position. Which can be expressed like this

$$y_{n+1} = y_n + y'_n \cdot dt \quad (1)$$

Where y_{n+1} is the next step, y_n is the current step, y'_n is the derived of the current step and dt is the time step. This algorithm is really based abbreviated version of a Taylor expansion, where we only expand the series one step at a time. But by only taking one step at a time, we will also get a local truncation error, which causes an error for each step we take.

$$\begin{aligned} y(t_n + dt) &= y_{n+1} \\ &= y(t_n) + y'_n \cdot dt + R(dt^2) \end{aligned} \quad (2)$$

Where $R(dt^2)$ is the local truncation error. Since the Forward Euler method is a first-order method, will

the local truncation error be proportional to the square of the step size.

1.2 Velocity Verlet method

The velocity Verlet method is based on the kinematic equations for an moving object, which in our case is the earth's orbit around the sun. If we want to find the next time step for the velocity and position we do a approximation and uses Taylor-expansion

$$v_{t+dt} = v_t + \frac{dt}{2} \left(\frac{F_t}{m} + \frac{F_{t+m}}{m} \right) R(dt^3) \quad (3)$$

We can also split the equation above and perform the calculation in several steps like this:

$$\begin{aligned} v(t + \frac{1}{2}\Delta t) &= v(t) + \frac{1}{2}a(t)\Delta t \\ x(t + \Delta t) &= x(t) + v(t + \frac{1}{2}\Delta t)\Delta t \\ a(t + \Delta t) &= f(x(t + \Delta t)) \\ v(t + \Delta t) &= v(t + \frac{1}{2}\Delta t) + \frac{1}{2}a(t + \Delta t)\Delta t \end{aligned} \quad (4)$$

The best way to show that the angular momentum is conserved by using Keplers second law is to make a drawing.

1.3 Conservation of angular momentum

The definition of Keplers law is that a area, where a connecting line between a planet and the sun, is always the same for an equal time interval. Which is illustrated in the figure below.

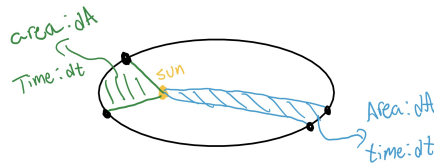


Figure 1: Keplers second law. Where the area is the same for the same time interval.

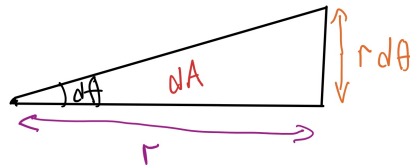


Figure 2: This square is the infinitesimal area dA that the planet has moved by a infinitesimal time interval dt

Figure 2 shows us that the area is

$$dA = \frac{1}{2} r^2 d\theta \quad (5)$$