

# Project 3

Kandidatnummer

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## 1 Method

We assume that the earths orbit around the sun is circular. And from circular motion and Newtons gravitational force we have that

$$F_G = G \frac{M_E M_\odot}{r^2} = \frac{M_E v^2}{r} \quad (1)$$

We now want to show that the velocity,  $v$ , of the earth can be written as

$$v^2 r = G M_\odot = 4\pi^2 \frac{AU^3}{yr^2} \quad (2)$$

Since we say it is circular orbits can we use the centripetal force to rewrite equation 2

$$M_E \omega^2 r = G \frac{M_E M_\odot}{r^2} \quad (3)$$

where  $\omega^2$  is the angular velocity of the earth. Equation 3 can now be rewritten as

$$M_E \left( \frac{2\pi}{P} \right)^2 r = G \frac{M_E M_\odot}{r^2} \quad (4)$$

(5)

Where  $P$  is a period of the earth around the sun. Now can we use Kepler's third law, which says that

the square of an orbital period  $P^2$  equals the cube of the semi-major axis of its orbiting  $a^3$

$$P^2 = a^3 \quad (6)$$

Where we substitute the circular radius  $r$  with the semi-major axis  $a$ .

## 2 Introduction

Though simulating a solar system is interesting and fun enough on it's own, it is naturally also quite useful for the study of astrodynamics. In addition, being able to simulate a solar system provides a good set of tools applicable to many other scientific areas. These tools include having a good understanding of different numerical integration methods, and being able to write a structured and fast code. In this project we wish to explore a model of our own solar system, beginning with simulating the simple two-body system including the Earth and the Sun. We will use this system to compare two different methods of numerical integration, the forward Euler method and the velocity Verlet method. This simple system also

makes a good testing ground for exploring whether our model is consistent with known physical laws such as energy conservation and Kepler's laws of planetary motion. We will also test the stability of the velocity Verlet method by including Jupiter and playing around with it's mass. From there we will include the rest of the planets in our solar system, and blah blah blag general relativity.

### 3 Theory

#### 3.1 Forward Euler

The forward Euler method is a algorithm to estimate the solution of a differential equation. The Forward Euler method wants to find the next point. To find the next point, it uses the point it is at,  $r_n$ , a small time step,  $dt$ , and the derivative of its position. Which can be expressed like this

$$y_{n+1} = y_n + y'_n \cdot dt \quad (7)$$

Where  $y_{n+1}$  is the next step,  $y_n$  is the current step,  $y'_n$  is the derived of the current step and  $dt$  is the time step. This algorithm is really based abbreviated version of a Taylor expansion, where we only expand the series one step at a time. But by only taking one step at a time, we will also get a local truncation error, which causes an error for each step we take.

$$\begin{aligned} y(t_n + dt) &= y_{n+1} \\ &= y(t_n) + y'_n \cdot dt + R(dt^2) \end{aligned} \quad (8)$$

Where  $R(dt^2)$  is the local truncation error. Since the Forward Euler method is a first-order method, will the local truncation error be proportional to the square of the step size.

#### 3.2 Velocity Verlet method

The velocity Verlet method is based on the kinematic equations for an moving object, which in our case is the earth's orbit around the sun. If we want to find the next time step for the velocity and position we do a approximation and uses Taylor-expansion

$$v_{t+dt} = v_t + \frac{dt}{2} \left( \frac{F_t}{m} + \frac{F_{t+m}}{m} \right) R(dt^3) \quad (9)$$

We can also split the equation above and perform the calculation in several steps like this:

$$v_{t+dt} = v_t + \frac{dt}{2} \left( \frac{F_t}{m} + \frac{F_{t+m}}{m} \right) R(dt^3) \quad (10)$$

We can also split the equation above and perform the calculation in several steps like this:

$$\begin{aligned} v(t + \frac{1}{2}\Delta t) &= v(t) + \frac{1}{2}a(t)\Delta t \\ x(t + \Delta t) &= x(t) + v(t + \frac{1}{2}\Delta t)\Delta t \\ a(t + \Delta t) &= f(x(t + \Delta t)) \\ v(t + \Delta t) &= v(t + \frac{1}{2}\Delta t) + \frac{1}{2}a(t + \Delta t)\Delta t \end{aligned} \quad (11)$$

#### 3.3 Conservation of angular momentum

The definition of Keepers law is that a area, where a connecting line between a planet and the sun, is always

the same for an equal time interval. Which is illustrated in the figure below:

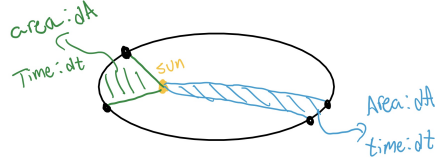


Figure 1: Keplers second law. Where the area is the same for the same time interval.

The best way to show that the angular momentum is conserved by using Keplers second law is to make a drawing.

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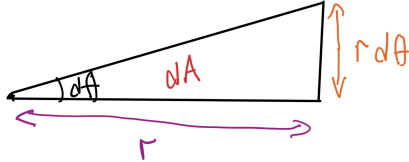


Figure 2: This square is the infinitesimal area  $dA$  that the planet has moved by a infinitesimal time interval  $dt$

Figure 2 shows us that the area is

$$dA = \frac{1}{2}r^2d\theta \quad (12)$$

We now want to find an infinitesimal area where the planet has moved around the sun, over an infinitesimal

time step, where we use equation 12.

$$\frac{dA}{dt} = \frac{1}{2}r^2\frac{d\theta}{dt} \quad (13)$$

$$= \frac{1}{2}rv_{\theta} \quad (14)$$

And the definition of angular momentum is given as

$$L = mrv_{\theta} \quad (15)$$

With equation 14 and 15 we can show that the angular momentum is conserved.

### 3.4 Testing forms of the force

Until now we have looked at the movement for the planet with the gravitational force with the square of  $r$ . We will now try to change the gravitational force where the power of  $r$  is in the range  $\beta \in [2, 3]$  The gravitational force is then given as

$$\mathbf{F}_G = -\frac{GM_{\odot}M_{planet}}{r^{\beta}}\frac{\mathbf{r}}{r} \quad (16)$$

### 3.5 Escape velocity

We will now look at a planet which begins at a distance 1 AU from the sun, and see how fast the initial velocity must be for the planet to be able to escape the sun. We find the analytical solution to compare it to our numerical solution.

The analytical expression is given as

$$|\mathbf{V}_{esc}| = \sqrt{\frac{2GM_{\odot}}{r}} \quad (17)$$

By plugging equation 2 into 16, and use that our initial position of the

planet is  $r = AU$  we get

$$| \mathbf{V}_{esc}^2 | = \frac{2 \cdot 4\pi^2 AU^3}{yr^2} \cdot \frac{1}{r} \quad (18)$$

$$= \frac{2 \cdot 4\pi^2 AU^3}{yr^2} \cdot \frac{1}{AU} \quad (19)$$

$$| \mathbf{V}_{esc} | = 2\sqrt{2}\pi \cdot \frac{AU}{yr} \quad (20)$$

## 4 Results

### 4.1 Conservation of angular momentum

If we substitute equation 15 into equation 14

$$\frac{dA}{dt} = \frac{1}{2} \frac{L}{m} \quad (21)$$

Sine  $\frac{dA}{dt}$ =constant and the mass  $m$  is constant means that the angular momentum also is constant and therefore also conserved.

## **A Calculations**

### **References**

[1] Skriv inn kilde her.