

Chapter 1

Neutrinos and the Large Scale Structure

In this chapter, we will briefly discuss the effects that neutrinos have on the matter power spectrum. For a more in-depth discussion of all the effects that neutrinos have on our universe, we refer to the main source of this chapter [?]. In the first section, we will give a phenomenological overview of the effect of massive neutrinos on the large-scale structure(LSS). Then we will remark on the effect that additional massless neutrinos have on the matter power spectrum. These discussions will also be representative of other simple dark matter models. Simply put any relativistic, collisionless, thermalized particle acts like a neutrino on the large-scale structure of the universe.

1.1 Massive Neutrinos

When discussing the effect of the massive neutrinos on the power spectrum the main effect can be very roughly described with three main phenomenological statements:

- Massive neutrinos stop clustering on scales larger than their free streaming scale that is related to their velocity.
- Massive neutrinos lead to an overall steplike suppression of the power spectrum by an amount proportional to their mass.
- Massive neutrinos change the redshift of matter-to-radiation equality essentially shifting the peak of the matter power spectrum.

The first effect leads to a suppression of the power spectrum for scales smaller than some scale k_* . This can be understood from the following reasoning. The free-streaming scale represents a scale under which collisionless particles can not be confined, if we define it analogously to the Jeans wavenumber we find a free-streaming wavenumber

$$k_{\text{fs}} = \sqrt{\frac{3}{2}} \frac{a H}{c_\nu}. \quad (1.1)$$

For relativistic neutrinos, c_ν is one while for nonrelativistic neutrinos we can find the thermal velocity by dividing the mean momentum by the neutrino mass. This leads to

$$c_\nu = \approx \frac{8.78 \cdot 10^{-3}}{a} \frac{0.06 \text{ eV}}{m_\nu}$$

$$k_{\text{fs}} \approx 4.66 \cdot 10^{-3} \frac{m_\nu}{0.06 \text{ eV}} a^2 \frac{H}{H_0} h \text{ Mpc}^{-1}. \quad (1.2)$$

During matter domination for the relativistic case, the free streaming wavenumber falls with η^{-1} and for the nonrelativistic case the wavenumber grows with η . This means that at the transition from relativistic to nonrelativistic we find the largest scale at which neutrinos stopped clustering. We denote this scale with k_{min} (we will call it the minimum clustering scale), and it is precisely where the steplike suppression of the neutrinos starts. Given that the transition to the nonrelativistic regime happens during matter domination, k_{min} can be approximated by

$$k_{\text{min}} \approx 7.96 \cdot 10^{-3} \sqrt{\frac{\Omega_m}{0.3}} \sqrt{\frac{m_\nu}{0.06 \text{ eV}}} h \text{ Mpc}^{-1}. \quad (1.3)$$

To describe the second effect we first need to dissect our matter perturbation into baryons, CDM and massive neutrinos. We can write

$$\delta_m = f_c \delta_c + f_b \delta_b + f_\nu \delta_\nu. \quad (1.4)$$

The factors f_X are the fractional contribution of the species X to the total energy density of the universe. The equations of motion drive the neutrino perturbations to an equilibrium similar to the cold constituents. However, due to the high velocity of the neutrinos, there is non-negligible pressure and anisotropic stress. This leads to

- (a) a small-scale suppression of the neutrino perturbation that is approximately given by

$$\delta_\nu \propto \left(\frac{k_{\text{fs}}}{k} \right)^2 \delta_c$$

on scales smaller than the free streaming scale.

- (b) a scale-dependent growth rate that is reduced compared to the cold matter constituents until the equilibrium has been reached.

For the following discussion, we will only look at what will happen to the matter perturbations for scales much larger than the minimum clustering scale. Because of the two reasons stated above, on these scales, we use that at late times typical $|\delta_\nu| \ll \delta_c$. We can now expand equation 1.4 to find

$$\langle \delta_m(k) \delta_m^*(k') \rangle \approx (f_c + f_b)^2 \left\langle \frac{(f_c \delta_c(k) + f_b \delta_b(k))(f_c \delta_c^*(k') + f_b \delta_b^*(k'))}{(f_c + f_b)^2} \right\rangle$$

$$:= (f_c + f_b)^2 \langle \delta_{cb}(k) \delta_{cb}^*(k') \rangle \quad (1.5)$$

$$\Longleftrightarrow P_{mm}(k) \approx (1 - f_\nu)^2 P_{cb}(k), \quad (1.6)$$

where we have defined the perturbation of the CDM+baryon field δ_{cb} and it's respective power spectrum. After the baryon drag epoch, i.e. after baryons decouple from the radiation field, the equations of motion for this new density contrast are simplified to

$$\delta_{cb}'' + \frac{a'}{a} \delta_{cb}' - 4a^2 \pi G \langle \rho_c + \rho_b \rangle \delta_{cb} = 0, \quad (1.7)$$

where the prime denotes derivatives with respect to conformal time.

The third term is starting to look like the Friedmann–Lemaître equation minus the effect of massive neutrinos. If we add them we can write

$$\delta_{cb}'' + \frac{a'}{a} \delta_{cb}' - \frac{3}{2} (1 - f_\nu) (a')^2 \delta_{cb} = 0. \quad (1.8)$$

If we assume that we are deep into matter domination such that the bulk of neutrinos are already non-relativistic. The neutrino density thus scales with a^{-3} . At this time the background evolution matches the neutrinoless case, and we can write that $a \propto \eta^2$. If we insert the conformal time into the equality we can solve the differential equation to find

$$\delta_{cb}'' + \frac{2}{\eta} \delta_{cb}' - \frac{6}{\eta^2} (1 - f_\nu) \delta_{cb} = 0. \quad (1.9)$$

$$\implies \delta_{cb} \propto \eta^{\alpha_\pm} \quad \text{with} \quad \alpha_\pm = -\frac{1}{2} \pm \frac{1}{2} \sqrt{25 - 24f_\nu} \quad (1.10)$$

This gives us one growing mode and a decaying mode. The result we find is different from the result of the neutrinoless universe where for the growing mode we find $\delta \propto a$. When we also use that neutrinos contribute very little to the total energy density, i.e. $f_\nu \ll 1$, we can find an approximate solution

$$\delta_{cb} = A_- a^{-\frac{3}{2} + \frac{3}{5}f_\nu} + A_+ a^{1 - \frac{3}{5}f_\nu}. \quad (1.11)$$

If we neglect the decaying mode, inserting the result for δ_{cb} into the Poisson equation we also find

$$k^2 \Psi = -\frac{3}{2} (1 - f_\nu) (a')^2 \delta_{cb} \propto a^{-\frac{3}{5}f_\nu}. \quad (1.12)$$

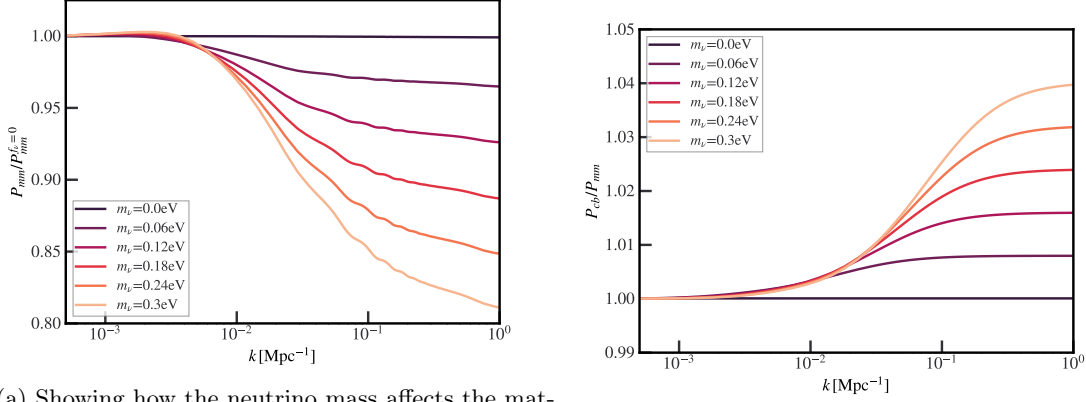
We can conclude that due to the massive neutrinos, we see that not only does the density contrast of clustering matter grow slower, but also that the metric perturbations slowly decay. From numerical calculations, we can find a rough estimate of the suppression on scales much smaller than the minimal clustering scale. For $f_\nu \ll 1$ the overall suppression of the power spectrum can be estimated with

$$\frac{P_{cb}}{P_{cb}^{f_\nu=0}} \approx 1 - 6f_\nu, \quad (1.13)$$

or for the total matter power spectrum $P_{mm}/P_{mm}^{f_\nu=0} \approx 1 - 8f_\nu$. We illustrate the suppression of the power spectrum and the difference between the CDM+baryon spectrum and the total matter power spectrum in figure 1.1. In the figure we have fixed the total nonrelativistic matter density to match the low redshift expansion history between the different models this means reducing the cold dark matter density. These changes also lead to further suppression of the power spectrum on the smallest scales, essentially tilting the plateau further.

The last effect of massive neutrinos is the shifting of the redshift of equality. When fixing the

Figure 1.1: The effect of changing N_{eff} when fixing different quantities as explained in the Text. The Ratios were multiplied with a factor to better differentiate between them.



(a) Showing how the neutrino mass affects the matter power spectrum. For the comparison we fixed $\{h, \Omega_b, \Omega_m\}$

(b) Amplification of the CDM+baryon power spectrum for different neutrino masses

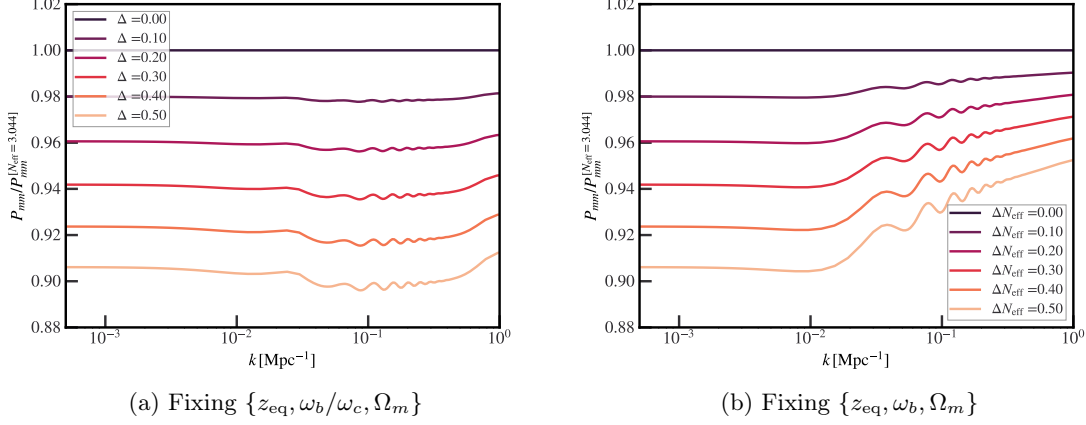
density parameters of radiation, total matter, and baryons additional neutrinos shift the redshift of equality by a factor of $(1 - f_\nu)$. This is due to the neutrinos being ultra-relativistic at the time of equality and thus not contributing to the matter density. The shift of the redshift of equality leads to a shift in the peak of the matter power spectrum as it is given from the wavenumber that crossed the Hubble horizon precisely at equality. The slight wiggles in the ratio come from a change in the expansion history at earlier times shifting the scale of the BAO. We note that when we combine our *Euclid* probes with *Planck*, the CMB experiments strongly constrain the angular scale θ_s , the redshift of equality z_{eq} , and the physical density parameter of baryons, ω_b . It can be shown that when fixing these quantities a change in the neutrino mass only leads to a more scale-independent suppression and a slight shift in the phase of the BAOs.

1.2 Massless Neutrinos

The effect of additional massless neutrinos on the matter power spectrum can be separated into background effects and perturbation effects. For the latter we can already guess that these will be small. The main contribution of massless neutrinos to the matter power spectrum must happen during radiation domination as they become negligible during matter domination. During radiation domination on scales smaller than the Hubble horizon, it can be shown that the massless neutrinos have very little effect on the evolution of cold perturbations. This is because the massless neutrino perturbations crossing the Hubble horizon show an oscillatory behaviour. When decomposing the equations of motion for the perturbations into fast oscillating modes and slowly growing modes, we see that cold matter is primarily determined by slow modes. As a consequence, we can treat the massless neutrino perturbations as free streaming one matter domination starts.

The background effects of additional massive neutrinos have to be analysed a bit more carefully. Depending on what quantities are fixed other quantities need to change as well. Firstly we start

Figure 1.2: The effect of changing N_{eff} when fixing different quantities as explained in the Text. The Ratios were multiplied with a factor to better differentiate between them.



with the parametrization of additional massless relics, for this, we use the parameter N_{eff} . We write, that the density of total radiation is given as

$$\rho_r = \rho_\gamma \left[1 + \frac{7}{8} \left(\frac{4}{11} \right)^{4/3} N_{\text{eff}} \right] \quad (1.14)$$

This means, that when fixing the density parameters of matter a change of N_{eff} coincides with a change of the redshift of equality, shifting the peak of the power spectrum. In order to fix the redshift of equality we need to scale ω_m as ω_γ is very tightly constrained by Firas.

The overall amplitude of the matter power spectrum is related to the amplitude of scalar perturbations A_s as well as the total matter density parameter Ω_m . The latter is proportional to the physical density of total matter by $\omega_m = \Omega_m h^2$. If we also want to fix the amplitude of the matter power spectrum without changing A_s we have to scale h accordingly.

Next, if we want to subtract that effect then we need to rescale the matter density. Depending on if we switch ω_b or ω_c this changes the ratio of ω_b/ω_c . This ratio is crucial in the large-scale amplitude of the power spectrum as well as the amplitude of the BAOs. To cancel this effect we need to equally scale the CDM and baryon density. What would be left in this set of fixed variables is a change in ω_b , that changes the phase of the BAOs. We have no longer the freedom to fix this effect.

We could have also decided to fix ω_b this would need us to change the ratio of cold dark matter to baryons and suppress the large-scale plateau. We illustrate the effects of changing N_{eff} for both cases in figure 1.2.