



Sterile Neutrinos and Non-Standard Interactions in Neutrino Telescopes

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Abstract

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Chapter 1

Introduction

There is no reason for us to believe that our current knowledge in particle physics is complete. The most successful framework – the Standard Model – has notable shortcomings. Neither does it explain why there is something rather than nothing, nor does it give a candidate for dark matter¹. To make matters even worse, three of the massless particles have been observed to have mass. In this thesis, we focus on the said particles – the neutrinos – and propose extensions to the Standard Model related to them. We then compare the effects of these extensions with collected data to see if our new theory more closely describes Nature or not.

Ultimately, the purpose of this excursion is to guide future research where to uncover more accurate theories. Traces of this grander theory will be present as breadcrumbs scattered in Nature. We just need to know where and how carefully to look.

1.1 Outline

This thesis is outlined as follows. In Chapter 2, we briefly review the most relevant parts of the Standard Model that will be relevant for our extensions of it. Then we present the evidence and subsequent discovery of neutrino oscillations, which ultimately led to the 2015 Nobel physics prize, along with the modifications to the Standard Model needed to accommodate the discovery. The reader is then introduced to a standard but detailed derivation of the neutrino mixing matrix and oscillation Hamiltonian, including the matter effects stemming from charged and neutral current weak interactions within the Earth.

In Chapter 3, we present an experimental procedure used to observe signals from neutrinos originating from cosmic ray interactions in the atmosphere of the Earth. Furthermore, present the two present Antarctic neutrino detectors: IceCube and

¹In the Standard Model, matter and antimatter come in pairs of one particle and one antiparticle. When interacting, antimatter *annihilates* the matter, resulting in pure light. Since the universe exists, there must have been more matter than antimatter at the Big Bang. This baryon asymmetry is unexplained in the Standard Model. Dark matter is heavy matter which does not interact with electromagnetic radiation. No such particle exists in the Standard Model.

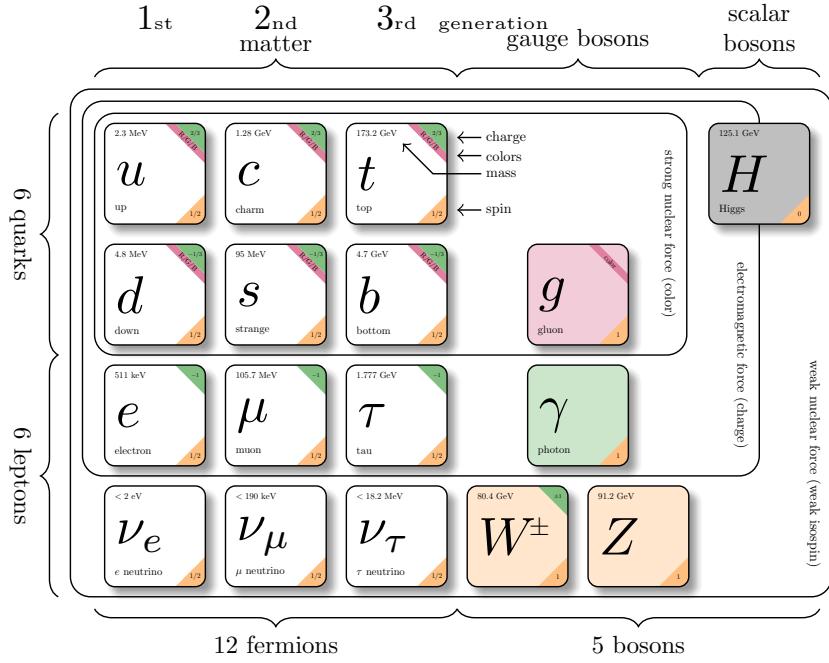
DeepCore. We also summarize a proposed detector upgrade: Precision IceCube Next Generation Upgrade (PINGU), which we will use as a forecast in our analysis.

Chapter 4 consists of two parts. In Section 4.1, we introduce a new particle: the sterile neutrino. We present how neutrino oscillations are modified by this hypothesized fourth neutrino and how these modifications would appear in IceCube if the particle is present in Nature. In Section 4.2, we completely set the sterile neutrino aside and instead turn to the interactions between neutrinos and the Earth. We again amend the Standard Model by considering exotic interactions from a higher energy theory, manifesting themselves as sub-leading modifications to the matter potential. Since these new interactions are not present in the Standard Model, we call them non-standard interactions (NSI).

Finally, Chapter 5 contains the result from our two separate extensions of the Standard Model. In Section 5.1 we present our findings for the sterile neutrino and compare those with literature. In Section 5.2, we present new constrained bounds on the NSI parameters by combining our results for IceCube, DeepCore, and PINGU.

Chapter 2

Neutrino Oscillations



2.1 The Standard Model

In order to describe the three quantizable forces of nature, we gather the mediators of each force – the vector bosons – into local (gauge) symmetry groups. Each vector boson has one corresponding generator, the set of which constitutes the group. The strong charge is mediated by eight massless gluons, which correspond to the eight independent generators of $SU(3)_C$. The weak charge is mediated by the three massive gauge bosons W^\pm and Z , and the massless photon γ , which constitute the generators of $SU(2)_L$ and $U(1)_Y$.

The subscript of each group denotes by which mechanism that force is mediated. The gluons mediate the strong force through interactions of color, emphasized with subscript C . The weak force only sees left-handed particles, which we distinguish with the subscript L . And the electroweak interaction that a particle undergoes is determined by its hypercharge Y . For example, the quarks all have a non-zero color and non-zero hypercharge, so they participate in the strong and electromagnetic interactions. If a quark is left-handed, it will also feel the weak interaction. The neutrinos on the other hand, have neither charge nor color, so they are invisible to both the strong and electromagnetic force. We express this by letting their fields transform as singlets under those symmetry groups.

Together, these three interactions make up the Standard Model gauge group $SU(3)_C \times SU(2)_L \times U(1)_Y$. This determines the form of the three coupling constants, which numerical values must be experimentally measured. Since the vector bosons

are represented by the generators, they are uniquely determined by the symmetry group. However, the scalar boson(s) and fermions are free as long as they belong to representations of the symmetry group. By this construction, the fermions have more degrees of freedom with which we can propose amendments to the model with. Even the number of fermions must be experimentally verified and can be altered from a phenomenological standpoint. In this work, we will use this leniency to introduce a new particle and examine to which extent these modifications might be supported by experimental evidence. Moreover, we will examine the possibility of adding a new completely new force to the Standard Model. As then follows from our discussion above, this means that we will add a symmetry group and thus also at least one more vector boson.

2.2 Neutrino Mixing

Japan 1998. The Super-Kamiokande experiment detects a ν_μ deficit that could not be explained by any other mechanism than the conversion of muon neutrinos to electron neutrinos. This was the first conclusive evidence of a completely new physics phenomenon: neutrino flavor oscillations. An implication of the observation of oscillations is neutrino mass, which is identically zero according to the Standard Model. Thus, we are required to extend the Standard Model to incorporate neutrino masses before we discuss the neutrino oscillations themselves.

2.2.1 Mass Generation

In the Standard Model, fermion masses are generated by the Higgs mechanism through Yukawa couplings with the fermion's right and left-handed components. All neutrinos are left-handed and all antineutrinos are right-handed. Thus, they do not pick up a Yukawa coupling, and will not undergo the Higgs mechanism, which leaves them massless. Additionally, all terms of a Lagrangian that we construct must respect the gauge invariance. This removes the possibility of the neutrino mass to be generated at loop level because any such attempt will violate the total lepton number by two units. In order to keep our theory renormalizable, we have one option left if we want to generate neutrino masses: introducing right-handed neutrinos and left-handed antineutrinos.

We consider a right-handed neutrino field, ν_R . Since the electroweak gauge group $SU(2)_L \times U(1)_Y$ only couple to left-handed particles and right-handed antiparticles, ν_R transforms as a singlet under the Standard Model symmetry group $SU(3)_C \times SU(2)_L \times U(1)_Y$. This neutrino field is *sterile* since it doesn't participate in any of the Standard Model interactions.

We extend the Standard Model by adding a right-handed component of this field with neutrino Yukawa couplings $Y_{\alpha\beta}^{\nu\nu}$ to the Higgs-lepton Yukawa Lagrangian

$$\mathcal{L}_H = - \left(\frac{v + H}{\sqrt{2}} \right) [\ell'_{\alpha L} Y_{\alpha\beta}^{\nu\ell} \ell'_{\beta R} + \nu'_{\alpha L} Y_{\alpha\beta}^{\nu\nu} \nu'_{\beta R}] . \quad (2.1)$$

Now, the Yukawa coupling matrices $Y^{\nu\ell}$ and $Y^{\nu\nu}$ are non-diagonal, with the former resulting in the charged lepton masses being non-definite. This can be remedied by diagonalizing $Y^{\nu\ell}$ with a unitary matrix V^ℓ as

$$V_{\alpha k L}^{\nu\dagger} Y_{\alpha\beta}^{\nu\ell} V_{\beta j R}^{\nu} = Y_{kj}^{\nu\ell} . \quad (2.2)$$

Similarly, we diagonalize and $Y^{\nu\nu}$ with a unitary matrix V^ν . We note that this will not yet generate neutrino masses, since no flavor rotation can make massless fields massive. Now we state a crucial difference between the properties of the charged lepton and the neutrino fields. While the charged lepton flavor eigenstate

was uniquely determined by its mass eigenstate, the neutrino flavor is a superposition of mass eigenstates. This is because neutrinos are indirectly detected via the observation of its associated charged lepton, so there is no requirement for neutrino flavor eigenstates to have a definite mass. The flavor of a neutrino is then, by definition, the flavor of the associated charged lepton. This fact is denoted as giving the mass eigenstates Latin numerals and letters, while the flavor eigenstates stay as Greek letters.

So, let the neutrino field with chirality X be denoted ν_X , with components having Latin numerals to distinguish them from the flavour components, i.e

$$\nu_{kX} = V_{k\beta X}^{\nu\dagger} \nu'_{\beta X}. \quad (2.3)$$

The diagonalized Lagrangian now takes the form

$$\begin{aligned} \mathcal{L}_H &= - \left(\frac{v + H}{\sqrt{2}} \right) [\ell'_{\alpha L} Y'^{\ell}_{\alpha\beta} \ell'_{\beta R} + \nu'_{\alpha L} Y'^{\nu}_{\alpha\beta} \nu'_{\beta R}] \\ &= - \left(\frac{v + H}{\sqrt{2}} \right) [\ell'_{\alpha L} V^{\ell}_{\alpha\beta L} Y^{\ell}_{\alpha\beta} V^{\ell\dagger}_{\alpha\beta R} \ell'_{\beta R} + \nu'_{\alpha L} V^{\nu}_{\alpha k L} Y^{\nu}_{kj} V^{\nu\dagger}_{\beta j R} \nu'_{\beta R}] \\ &= - \left(\frac{v + H}{\sqrt{2}} \right) [\ell^{\dagger}_{\alpha L} Y^{\ell}_{\alpha\beta} \ell_{\beta R} + \nu^{\dagger}_{k L} Y^{\nu}_{kj} \nu_{j R}] \\ &= - \left(\frac{v + H}{\sqrt{2}} \right) [\bar{\ell}_{\alpha L} Y^{\ell}_{\alpha\beta} \ell_{\beta R} + \bar{\nu}_{k L} Y^{\nu}_{kj} \nu_{j R}] \end{aligned} \quad (2.4)$$

By construction, Y^ℓ and Y^ν are diagonal, so we write their components as $y_\alpha^\ell \delta_{\alpha\beta}$ and $y_k^\nu \delta_{kj}$ respectively, leaving the Lagrangian as

$$\begin{aligned} \mathcal{L}_H &= - \left(\frac{v + H}{\sqrt{2}} \right) [\bar{\ell}_{\alpha L} y_\alpha^\ell \delta_{\alpha\beta} \ell_{\beta R} + \bar{\nu}_{k L} y_k^\nu \delta_{kj} \nu_{j R}] \\ &= - \left(\frac{v + H}{\sqrt{2}} \right) [\bar{\ell}_{\alpha L} y_\alpha^\ell \ell_{\alpha R} + \bar{\nu}_{k L} y_k^\nu \nu_{k R}] \\ &= - \left(\frac{v + H}{\sqrt{2}} \right) [y_\alpha^\ell \bar{\ell}_{\alpha L} \ell_{\alpha R} + y_k^\nu \bar{\nu}_{k L} \nu_{k R}] \end{aligned} \quad (2.5)$$

Now, by the introduction of the sterile field ν_R , the Dirac neutrino field is

$$\nu_k = \nu_{kL} + \nu_{kR}. \quad (2.6)$$

Multiplying ν_k with its conjugate $\bar{\nu}_k$, we get

$$\begin{aligned} \bar{\nu}_k \nu_k &= \bar{\nu}_{kL} \nu_{kL} + \bar{\nu}_{kR} \nu_{kL} + \bar{\nu}_{kL} \nu_{kR} + \bar{\nu}_{kR} \nu_{kR} \\ &= \bar{\nu}_{kL} \nu_{kR} + \bar{\nu}_{kR} \nu_{kL} \\ &= \bar{\nu}_{kL} \nu_{kR} + \text{h.c.} \end{aligned} \quad (2.7)$$

The same calculation for the charged lepton field yields the same result for ℓ_k .

Substituting this result and multiplying the Higgs vacuum expectation value v into the fields gives us

$$\begin{aligned}\mathcal{L}_H &= -\left(\frac{v+H}{\sqrt{2}}\right) [y_\alpha^\ell \bar{\ell}_\alpha \ell_\alpha + y_k^\nu \bar{\nu}_k \nu_k] \\ &= -\frac{y_\alpha^\ell v}{\sqrt{2}} \bar{\ell}_\alpha \ell_\alpha - \frac{y_k^\nu v}{\sqrt{2}} \bar{\nu}_k \nu_k - \frac{y_\alpha^\ell}{\sqrt{2}} \bar{\ell}_\alpha \ell_\alpha H - \frac{y_k^\nu}{\sqrt{2}} \bar{\nu}_k \nu_k H.\end{aligned}\quad (2.8)$$

Thus, this extension to the SM generates by the Higgs mechanism neutrino masses with terms

$$m_k = \frac{y_k^\nu v}{\sqrt{2}}, \quad (2.9)$$

where the Yukawa couplings y_k^ν needs to be experimentally determined.

2.2.2 The Mixing Matrix

Substituting the new transformation from Eq. 2.3 into the expression for the weak charged current, we get

$$\begin{aligned}j_L^\rho &= 2\bar{\nu}'_{\alpha L} \gamma^\rho \ell'_{\alpha L} \\ &= 2\bar{\nu}_{k L} V'^{\nu\dagger}_{k\alpha} V^{\ell\mu}_{\alpha\alpha} \gamma^\rho \ell_{\alpha L}\end{aligned}\quad (2.10)$$

Now, the current in Eq. 2.10 conserves lepton number, since the neutrino field with flavor α only couples to the lepton field with flavor α . Thus, neutrino interactions still conserve lepton number. However, the Higgs-lepton Yukawa Lagrangian in Eq. 2.5 violates lepton number conservation since it couples the charged lepton flavor α to the neutrino mass eigenstate k , which is a superposition of flavors. There is no transformation that leaves both the interaction and kinetic Lagrangian invariant, thus violating this symmetry.

Call $V'^{\nu\dagger}_{k\alpha} V^{\ell\mu}_{\alpha\alpha} = U_{k\alpha}^\dagger$. We will refer to the matrix U built by the components $U_{k\alpha}$ as the Pontecorvo-Maki-Nakagawa-Sakata (PMNS) matrix [1, 2]. We now have

$$j_L^\rho = 2 \sum_{\alpha} \sum_k U_{\alpha k}^\dagger \bar{\nu}_{k L} \gamma^\rho \ell_{\alpha L}. \quad (2.11)$$

What form does this new matrix U take? By construction, it provides the unitary transformation between the flavor and mass bases. Any unitary 3×3 matrix can be parametrized using three mixing angles and six phases. However, not all phases affect the charged and weak currents and are thus not observable by us. Moreover, both the Lagrangian and the currents are invariant under global $U(1)$ transformations, leaving only one physical phase for us to measure. We are now down to four degrees of freedom in the three dimensional case: three mixing angles of the form $\sin(\theta_{ij})/\cos(\theta_{ij}) = s_{ij}/c_{ij}$, and one phase of the form $e^{i\delta}$. The

mixing angle θ_{ij} specifies the amount of rotation in the $i - j$ plane, allowing an i eigenstate to mix with a j eigenstate. The physical interpretation of a mixing angle is then the degree to which the ν_i eigenstate mixes with the ν_j eigenstate. A non-zero mixing angle will then cause the flavor eigenstate ν_α to be a superposition of the mass eigenstates that we allowed to mix.

We construct the PMNS matrix using the rotation matrixes from $SO(3)$ as

$$\begin{aligned} U &= R_{23}R_{13,\delta_{\text{CP}}}R_{12} \\ &= \begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix} \begin{pmatrix} c_{13} & 0 & s_{13}e^{-i\delta_{\text{CP}}} \\ 0 & 1 & 0 \\ -s_{13}e^{i\delta_{\text{CP}}} & 0 & c_{13} \end{pmatrix} \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}. \end{aligned} \quad (2.12)$$

The term δ_{CP} determines the degree of charge-parity violation. In this work, we will always assume CP invariance for simplicity by setting $\delta_{CP} = 0^\circ$. Thus,

$$\begin{aligned} U &= \begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix} \begin{pmatrix} c_{13} & 0 & s_{13} \\ 0 & 1 & 0 \\ -s_{13} & 0 & c_{13} \end{pmatrix} \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \\ &= \begin{pmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13} \\ -s_{12}c_{23} - c_{12}s_{23}s_{13} & c_{12}c_{23} - s_{12}s_{23}s_{13} & s_{23}c_{13} \\ s_{12}s_{23} - c_{12}c_{23}s_{13} & -c_{12}s_{23} - s_{12}c_{23}s_{13} & c_{23}c_{13} \end{pmatrix}. \end{aligned} \quad (2.13)$$

In this work, we will be using the best-fit values from [3] with the exception of the CP-violating phase δ_{CP} . The 90% CL values are

$$\theta_{12} = 33.44^\circ, \quad \theta_{13} = 8.57^\circ, \quad \theta_{23} = 49.2^\circ, \quad \delta_{\text{CP}} = 0. \quad (2.14)$$

Our neglect of the issue of CP invariance leaves the PMNS matrix real, with numerical values

$$U = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} = \begin{pmatrix} 0.825 & 0.545 & 0.149 \\ -0.455 & 0.485 & 0.746 \\ 0.334 & -0.684 & 0.649 \end{pmatrix}. \quad (2.15)$$

In our study, we will only use neutrinos originating from conventional cosmic ray interactions in the atmosphere. In this so-called atmospheric neutrino flux, ν_μ are the most abundant, while the ν_τ flux is negligible. From Eq. (2.15), we see that $U_{\mu\tau}$ is the largest $U_{\mu\beta}$ term, suggesting that ν_μ primarily mixes into ν_τ , at least in vacuum. It turns out that matter effects indeed preserves this ordering, making the $\nu_\mu \rightarrow \nu_\tau$ transition the most abundant. This means we will be able to stringently constrain parameters relating to $\mu\tau$ mixing.

2.2.3 Neutrino Propagation

Since we now know how the neutrino mass and flavor eigenstates combine, and have an expression for the flavor interaction with the neutrino's charged lepton partner, we are now ready to study the flavor oscillations themselves.

Now, the Fourier expansion of the field operator $\bar{\nu}_{kL}$ in the current of Eq. 2.11 contains creation operators $a_{\nu_k}^\dagger$ of massive neutrinos with mass m_k . This means that the summation over the mass index k constructs a flavor neutrino, which interacts with the charged lepton field $\ell_{\alpha L}$. In other words, the charged current generates a flavor neutrino ν_α , which is a superposition of the mass eigenstates ν_k with weights $U_{\alpha k}^\dagger$. In the ket-formalism, we express this as

$$|\nu_\alpha\rangle = \sum_k U_{\alpha k}^\dagger |\nu_k\rangle . \quad (2.16)$$

It is the mass eigenstates $|\nu_k\rangle$ that are eigenstates of the Hamiltonian, with eigenvalues

$$E_k = \sqrt{\vec{p}^2 + m_k^2} . \quad (2.17)$$

The solution to the time-dependent Schrödinger equation

$$i \frac{d}{dt} |\nu_k(t)\rangle = H_0 |\nu_k(t)\rangle , \quad (2.18)$$

where H_0 is the vacuum Hamiltonian.

The solution to Eq. 2.18 gives us the time evolution in the form of plane wave solutions:

$$|\nu_k(t)\rangle = e^{-iE_k t} |\nu_k\rangle . \quad (2.19)$$

Inserting the plane wave solution into Eq. 2.16, we get

$$|\nu_\alpha(t)\rangle = \sum_k U_{\alpha k}^\dagger e^{-iE_k t} |\nu_k\rangle . \quad (2.20)$$

Now we know how to evolve and combine the mass eigenstates to form a flavor eigenstate, but how about the reverse? We swap the index $k \rightarrow j$ in Eq. 2.16 and multiply by $U_{\alpha k}$:

$$\begin{aligned} \sum_\alpha U_{\alpha k} |\nu_\alpha\rangle &= \sum_{\alpha, j} U_{\alpha k} U_{\alpha j}^\dagger |\nu_j\rangle \\ &= \sum_j \delta_{kj} |\nu_j\rangle \\ &= |\nu_k\rangle , \end{aligned} \quad (2.21)$$

where we have used the unitarity of the leptonic mixing matrix. Eqs. 2.16 and 2.20 yield

$$\begin{aligned} |\nu_\alpha(t)\rangle &= \sum_k U_{\alpha k}^\dagger e^{-iE_k t} |\nu_k\rangle \\ &= \sum_k U_{\alpha k}^\dagger e^{-iE_k t} \left(\sum_\beta U_{\beta k} |\nu_\beta\rangle \right) \\ &= \sum_{k,\beta} U_{\alpha k}^\dagger U_{\beta k} e^{-iE_k t} |\nu_\beta\rangle . \end{aligned} \quad (2.22)$$

The probability of the flavor transition $\nu_\alpha \rightarrow \nu_\beta$ at time t is $|\langle \nu_\beta | \nu_\alpha(t) \rangle|^2$:

$$P_{\nu_\alpha \rightarrow \nu_\beta}(t) = \sum_{k,j} U_{\alpha k}^\dagger U_{\beta k} U_{\beta j}^\dagger U_{\alpha j} e^{-i(E_k - E_j)t} . \quad (2.23)$$

We assume the neutrino masses m_k to be extremely small compared to their associated energies E_k . Thus, $v \sim 1$, and $|\vec{p}| \sim E$ making the energy-dispersion relation of Eq. 2.17 to first order:

$$\begin{aligned} E_k &= \sqrt{\vec{p}^2 + m_k^2} \\ &= \vec{p}^2 \sqrt{1 + \frac{m_k^2}{\vec{p}^2}} \\ &\approx E + \frac{m_k^2}{2E} \end{aligned} \quad (2.24)$$

Hence, the exponential can be simplified, and simplifying the notation $P_{\nu_\alpha \rightarrow \nu_\beta}(t) \rightarrow P_{\alpha\beta}(t)$ we get

$$P_{\alpha\beta}(t) = \sum_{k,j} U_{\alpha k}^\dagger U_{\beta k} U_{\beta j}^\dagger U_{\alpha j} e^{-i(m_k^2 - m_j^2)t/2E} . \quad (2.25)$$

Now, our approximation $v \approx 1$ implies $x \approx t$, thus

$$\begin{aligned} P_{\alpha\beta}(x) &= \sum_{k,j} U_{\alpha k}^\dagger U_{\beta k} U_{\beta j}^\dagger U_{\alpha j} e^{-i(m_k^2 - m_j^2)x/2E} \\ &= \sum_{k,j} U_{\alpha k}^\dagger U_{\beta k} U_{\beta j}^\dagger U_{\alpha j} \exp\left(-i\frac{\Delta m_{kj}^2 x}{2E}\right), \end{aligned} \quad (2.26)$$

where we in the last step have defined the *mass-squared difference* $\Delta m_{kj}^2 = m_k^2 - m_j^2$. Since the oscillation probability depends on this quantity rather than the individual masses, it is impossible to measure the mass m_k through neutrino oscillations.

Squaring the unitarity condition $\sum_k U_{\alpha k} U_{\beta k}^\dagger = \delta_{\alpha\beta}$ yields

$$\sum_k |U_{\alpha k}|^2 |U_{\beta k}|^2 = \delta_{\alpha\beta} - 2 \sum_{k>j} \text{Re}[U_{\alpha k}^\dagger U_{\beta k} U_{\alpha j} U_{\beta j}^\dagger] \quad (2.27)$$

Since we take the PMNS matrix to be real and thus Hermitian, we drop the complex transpose in the following calculation. We have

$$\begin{aligned} P_{\alpha\beta} &= \sum_{k,j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} \exp\left(-i \frac{\Delta m_{kj}^2 x}{2E}\right) \\ &= \sum_k |U_{\alpha k}|^2 |U_{\beta k}|^2 + \sum_{k\neq j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} \exp\left(-i \frac{\Delta m_{kj}^2 x}{2E}\right) \\ &= \delta_{\alpha\beta} - 2 \sum_{k>j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} + \sum_{k\neq j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} \exp\left(-i \frac{\Delta m_{kj}^2 x}{2E}\right) \\ &= \delta_{\alpha\beta} - 2 \sum_{k>j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} + 2 \sum_{k>j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} \exp\left(-i \frac{\Delta m_{kj}^2 x}{2E}\right) \\ &= \delta_{\alpha\beta} - 2 \sum_{k>j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} \left[1 - \exp\left(-i \frac{\Delta m_{kj}^2 x}{2E}\right)\right] \\ &= \delta_{\alpha\beta} - 2 \sum_{k>j} U_{\alpha k} U_{\beta k} U_{\beta j} U_{\alpha j} \sin^2\left(\frac{\Delta m_{kj}^2 x}{4E}\right), \end{aligned} \quad (2.28)$$

which is the probability of neutrino vacuum oscillations. The calculation for antineutrinos yield the same result if one continues to assume the realness of the mixing matrix. In Fig. 2.1, we show the energy spectra of these oscillations in the low GeV range for neutrinos that travel 12 000 km¹. In the left-most panel, we see a suppressed $\nu_\mu \rightarrow \nu_e$ transition. We can understand this behavior by studying Eq. 2.28 and the numerical values of the mixing matrix in Eq. 2.15. GeV oscillations are driven by the Δm_{31}^2 mass-squared difference. Hence, the most important mixing matrix elements for $P_{\alpha\beta}$ are $U_{\alpha 3}, U_{\alpha 1}$ and $U_{\beta 3}, U_{\beta 1}$. The smallness of $U_{e 3}$ then suppresses the amplitude of $P_{\mu e}$.

this form elucidates an important aspect of neutrino oscillations. The mixing matrix elements determine the amplitude of the oscillations, while the mass-squared differences together with the ratio L/E determine the frequency. So in vacuum, the mixing angles only influence the oscillation amplitude, while the energy, trajectory, and masses of the neutrinos only influence the oscillation frequency.

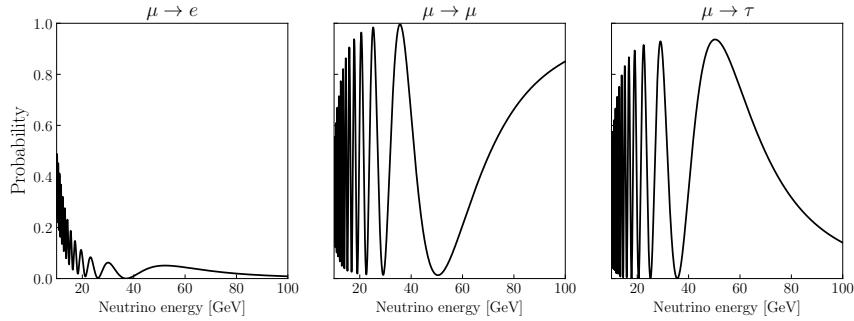
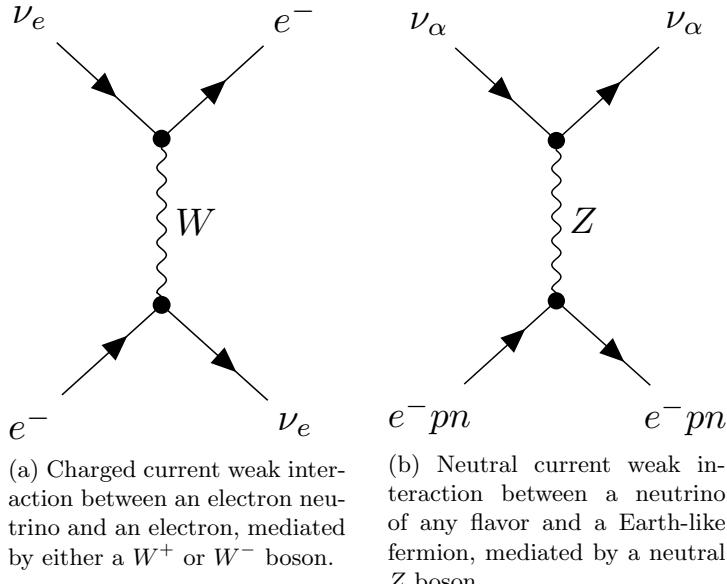
Figure 2.1: ν_μ vacuum oscillations from Eq. 2.28

Figure 2.2: Feynmann diagrams showing the two interactions that neutrinos participate in according to the Standard Model.

2.2.4 Effective Potentials

In this work, we are concerned about the interactions with the neutrino and Earth-like matter, i.e. electrons, protons, and neutrons. The possible interactions are shown in Fig. 2.2. The left panel shows that the only flavor that can undergo charged current (CC) interactions are the electron flavor. This is because the CC interaction requires a neutrino with flavor α to interact with its charged lepton partner α . Since the Earth doesn't consist of any muons or tau particles, the ν_μ and ν_τ have no charged leptons to undergo CC interactions with. The right panel shows any neutrino flavor interaction via the neutral current (NC) with Earth-like

¹This baseline is approximately the Earth diameter, and will be the maximum travel distance which we will use later on.

matter², mediated by the neutral Z boson. The interaction mediated by the W boson will give rise to an effective matter potential V_{CC} , while the Z boson is responsible for V_{NC} . Our task is now to find expressions for these.

We start with the effective Hamiltonian for the CC process. The Feynman rules for the left panel give us

$$H_{CC} = \frac{G_F}{\sqrt{2}} [\bar{\nu}_e \gamma^\rho (1 - \gamma^5) e] [\bar{e} \gamma_\rho (1 - \gamma^5) \nu_e] \quad (2.29)$$

By using the Fierz transformation

$$\mathcal{L}^{V-A}(\psi_1, \psi_2, \psi_3, \psi_4) = \mathcal{L}^{V-A}(\psi_1, \psi_4, \psi_3, \psi_2), \quad (2.30)$$

we can permute the terms inside the brackets, yielding

$$H_{CC} = \frac{G_F}{\sqrt{2}} [\bar{\nu}_e \gamma^\rho (1 - \gamma^5) \nu_e] [\bar{e} \gamma_\rho (1 - \gamma^5) e]. \quad (2.31)$$

Now, let's consider a finite volume V with electron states defined as

$$|e(p_e, h_e)\rangle = \frac{1}{2E_e V} a_e^{(h_e)\dagger}(p_e) |0\rangle, \quad (2.32)$$

i.e. using the creation operator $a_e^{(h_e)\dagger}(p_e)$ to create electron states from vacuum with momenta p_e , energy E_e , and helicity h_e . The density distribution of electrons in V is $f(E_e, T)$, which we normalize to the total number of electrons as we integrate out the momenta p_e :

$$\int dp_e^3 f(E_e, T) = N_e V = n_e \quad (2.33)$$

Here, the electron density N_e will ultimately determine the strength of the effective matter potential. To obtain the average effective Hamiltonian, project it on the

²The Earth is entirely composed of electrons, protons, and neutrons. Thus, the fundamental particles composing Earth are electrons, and up and down quarks. We refer to this as *Earth-like matter*.

electron states in Eq. 2.32 and integrate over the density and sum over the helicities:

$$\begin{aligned}
\langle H_{CC} \rangle &= \int dp_e^3 \langle e(p_e, h_e) | \times \frac{1}{2} \sum_{h_e} H_{CC} f(E_e, T) | e(p_e, h_e) \rangle \\
&= \frac{G_F}{\sqrt{2}} \int dp_e^3 \langle e(p_e, h_e) | [\bar{\nu}_e \gamma^\rho (1 - \gamma^5) \nu_e] f(E_e, T) \\
&\quad \times \frac{1}{2} \sum_{h_e} [\bar{e}(x) \gamma_\rho (1 - \gamma^5) e(x)] | e(p_e, h_e) \rangle \\
&= \frac{G_F}{\sqrt{2}} \bar{\nu}_e \gamma^\rho (1 - \gamma^5) \nu_e \int dp_e^3 f(E_e, T) \\
&\quad \times \frac{1}{2} \sum_{h_e} \langle e(p_e, h_e) | \bar{e}(x) \gamma_\rho (1 - \gamma^5) e(x) | e(p_e, h_e) \rangle . \quad (2.34)
\end{aligned}$$

First, calculate the sum using trace technology

$$\begin{aligned}
\frac{1}{2} \sum_{h_e} \langle e(p_e, h_e) | \bar{e}(x) \gamma_\rho (1 - \gamma^5) e(x) | e(p_e, h_e) \rangle &= \frac{1}{4E_e V} \sum_{h_e} \bar{u}_e^{h_e}(p_e) \gamma_\rho (1 - \gamma^5) u_e^{h_e}(p_e) \\
&= \frac{1}{4E_e V} \text{Tr} \left[\sum_{h_e} \bar{u}_e^{h_e}(p_e) u_e^{h_e}(p_e) \gamma_\rho (1 - \gamma^5) \right] \\
&= \frac{1}{4E_e V} \text{Tr} \left[(\not{p}_e + m_e) \gamma_\rho (1 - \gamma^5) \right] \\
&= \frac{(p_e)_\rho}{E_e V} . \quad (2.35)
\end{aligned}$$

Eq. 2.34 now becomes

$$\langle H_{CC} \rangle = \frac{G_F}{\sqrt{2} E_e V} \bar{\nu}_e (1 - \gamma^5) \nu_e \int dp_e^3 \not{p}_e f(E_e, T) . \quad (2.36)$$

Expand the integral, and use the fact that \vec{p}_e is odd:

$$\begin{aligned}
\int dp_e^3 \not{p}_e f(E_e, T) &= \int dp_e^3 f(E_e, T) (\gamma^0 E_e - \vec{p}_e \cdot \vec{\gamma}) \\
&= \int dp_e^3 f(E_e, T) \gamma^0 E_e \\
&= \gamma_0 E_e N_e V . \quad (2.37)
\end{aligned}$$

Inserting this into Eq. 2.36, we have

$$\begin{aligned}
\langle H_{CC} \rangle &= \frac{G_F N_e}{\sqrt{2}} \bar{\nu}_e (1 - \gamma^5) \nu_e \gamma_0 \\
&= \sqrt{2} G_F N_e \bar{\nu}_{eE} \gamma^0 \nu_{eL} , \quad (2.38)
\end{aligned}$$

where the projection operator $(1 - \gamma^5)$ in the first line ensures that only the left-hand component of the neutrino fields interact. Thus,

$$V_{CC} = \sqrt{2}G_F N_e, \quad H_{CC} |\nu_k\rangle = V_{CC} |\nu_k\rangle. \quad (2.39)$$

Here we see a crucial difference between the eigenvectors between the vacuum Hamiltonian defined in Eq. 2.16 and H_{CC} , namely that the CC interactions happen in the flavor basis rather than in the mass basis³. In other words, neutrinos propagate in their mass eigenstates but interact in their flavor eigenstate. The mixing of mass eigenstates during propagation determines if the flavor eigenstate has oscillated or not. The interactions still conserve lepton number since the interaction Lagrangian does, but the kinetic Lagrangian will not. Hence, neutrinos violate lepton number conservation during propagation.

For neutral current, we replace the electron field $e(x)$ in Eq. 2.31 by the more general fermion field $f(x)$, and the projection operator $(1 - \gamma^5)$ with $(g_V^f - g_A^f \gamma^5)$. Again, the γ^5 will cause the spacial component of p_f to disappear after integration, and the only difference between the average effective Hamiltonian for the neutral current is then the factor g_V^f :

$$V_{NC}^f = \sqrt{2}G_F N_A g_V^f. \quad (2.40)$$

Summing over the fermions, and assuming electrical neutrality and equal abundance of protons and neutrons, we have

$$\begin{aligned} V_{NC} &= \sum_{f \in e,p,n} V_{NC}^f \\ &= \sqrt{2}G_F N_A \sum_{f \in e,p,n} g_V^f \\ &= \sqrt{2}G_F N_A \left[-\frac{1}{2} + 2 \sin^2(\theta_W) + \frac{1}{2} - 2 \sin^2(\theta_W) - \frac{1}{2} \right] \\ &= -\frac{1}{\sqrt{2}}G_F N_e, \end{aligned} \quad (2.41)$$

where the electrical neutrality condition allows us to simply sum the vectorial couplings together, cancelling the electron and proton contributions (and hence, also the θ_W dependence).

We now see that both matter potentials, regardless of the interaction is of type charged or neutral, are only dependent on the electron density in the medium, N_e times a constant. So to calculate the matter oscillations in an electrically neutral medium with an equal abundance of protons and neutrons, we are only required to consider the matter density of the medium.

³A similar calculation for the NC reveals the same fact: neutrinos interact in the flavor basis.

2.2.5 Matter Oscillations

Since only ν_e undergo CC interactions in Earth-like matter, the V_{CC} potential is zero for all other flavors. However, since all flavors undergo NC interactions the total matter potential in the flavor basis is

$$V = \begin{bmatrix} V_{CC} + V_{NC} & 0 & 0 \\ 0 & V_{NC} & 0 \\ 0 & 0 & V_{NC} \end{bmatrix} = V_{CC} \delta_{\alpha e} \mathbb{I} + V_{NC} \mathbb{I}. \quad (2.42)$$

Just as in Eq. 2.18, we start with a Hamiltonian that solves the time-dependent Schrödinger equation. This time, let the Hamiltonian be

$$H = H_0 + H_I, \quad (2.43)$$

where H_0 is the Hamiltonian in vacuum, and H_I is our interaction Hamiltonian associated with our matter potentials. Let the wavefunction that describes the $\nu_\alpha \rightarrow \nu_\beta$ transition be denoted as

$$\psi_{\alpha\beta}(t) = \langle \nu_\beta | \nu_\alpha(t) \rangle, \quad (2.44)$$

i.e. the evolution of the state of a neutrino emitted at $t = 0$ with flavor α to flavor β at time t . Now using Eq. 2.16 and Eq. 2.39, we are ready to see what form our Hamiltonians take. Let us start with the vacuum Hamiltonian H_0 , and act on its Schrödinger equation with $\langle \nu_\beta |$:

$$i \frac{d}{dt} |\nu_\alpha(t)\rangle = H_0 |\nu_\alpha(t)\rangle \implies i \frac{d}{dt} \psi_{\alpha\beta} = \langle \nu_\beta | H_0 | \nu_\alpha(t) \rangle. \quad (2.45)$$

Reminding ourselves that the vacuum Hamiltonian H_0 has eigenstates in the mass basis, we write the following expression where we use the relations Eq. 2.16 and Eq. 2.21 to switch between the flavor and mass basis with the PMNS elements:

$$\begin{aligned} \langle \nu_\beta | H_0 = & \sum_k U_{\beta k} \langle \nu_k | H_0 \\ = & \sum_k U_{\beta k} E_k \langle \nu_k | \\ = & \sum_\eta \sum_k U_{\beta k} E_k U_{\eta k}^* \langle \nu_\eta |. \end{aligned} \quad (2.46)$$

Thus,

$$\begin{aligned} \langle \nu_\beta | H_0 | \nu_\alpha(t) \rangle = & \sum_\eta \sum_k U_{\beta k} E_k U_{\eta k}^* \langle \nu_\eta | \nu_\alpha(t) \rangle \\ = & \sum_\eta \sum_k U_{\beta k} E_k U_{\eta k}^* \psi_{\alpha\eta}(t). \end{aligned} \quad (2.47)$$

Using the ultrarelativistic approximation from Eq. 2.24:

$$\begin{aligned} \sum_{\eta} \sum_k U_{\beta k} E_k U_{\eta k}^* \psi_{\alpha \eta}(t) &= \sum_{\eta} \sum_k U_{\beta k} \left(p + \frac{m_k^2}{2E} \right) U_{\eta k}^* \psi_{\alpha \eta}(x) \\ &= \sum_{\eta} \sum_k U_{\beta k} \left(p + \frac{m_k^2}{2E} \right) U_{\eta k}^* \psi_{\alpha \eta}(x). \end{aligned} \quad (2.48)$$

Use the fact that $\sum_k m_k^2 = \sum_1 m_1^2 + \sum_k m_k^2 - m_1^2 = \sum_k m_1^2 + \Delta m_{k1}^2$ to pull out common terms out of the summation:

$$\begin{aligned} \sum_{\eta} \sum_k U_{\beta k} \left(p + \frac{m_k^2}{2E} \right) U_{\eta k}^* \psi_{\alpha \eta}(x) &= \sum_{\eta} \sum_k U_{\beta k} \left(p + \frac{m_1^2}{2E} + \frac{\Delta m_{k1}^2}{2E} \right) U_{\eta k}^* \psi_{\alpha \eta}(x) \\ &= \sum_{\eta} \sum_k \left(p + \frac{m_1^2}{2E} \right) U_{\beta k} U_{\eta k}^* \psi_{\alpha \eta}(x) \\ &\quad + \sum_{\eta} \sum_k U_{\beta k} \frac{\Delta m_{k1}^2}{2E} U_{\eta k}^* \psi_{\alpha \eta}(x). \end{aligned} \quad (2.49)$$

Unitarity gives $\sum_k U_{\beta k} U_{\eta k}^* = \delta_{\eta \beta}$, and the first term in the last step of Eq. 2.49 becomes

$$\sum_{\eta} \left(p + \frac{m_1^2}{2E} \right) \delta_{\beta \eta} \psi_{\alpha \eta}(x) = \left(p + \frac{m_1^2}{2E} \right) \psi_{\alpha \beta}(x). \quad (2.50)$$

Our treatment of the interaction Hamiltonian is similar except for the fact that its eigenstates lie in the flavor basis, conveniently allowing us to letting it act directly on the flavor eigenstates:

$$\begin{aligned} \langle \nu_{\beta} | H_I &= V_{\beta} \langle \nu_{\beta} | \\ &= \delta_{\beta \eta} V_{\beta} \langle \nu_{\eta} |. \end{aligned} \quad (2.51)$$

Using Eq. 2.42 in component form, we rewrite this as

$$\begin{aligned} \delta_{\beta \eta} V_{\beta} \langle \nu_{\eta} | &= \delta_{\beta \eta} (V_{CC} \delta_{\beta e} + V_{NC}) \langle \nu_{\eta} | \\ &= V_{CC} \delta_{\beta \eta} \delta_{\beta e} \langle \nu_{\eta} | + V_{NC} \langle \nu_{\beta} | \\ \implies \langle \nu_{\beta} | H_I | \nu_{\alpha} \rangle &= V_{CC} \delta_{\beta \eta} \delta_{\beta e} \langle \nu_{\eta} | \nu_{\alpha} \rangle + V_{NC} \langle \nu_{\beta} | \nu_{\alpha} \rangle \\ &= V_{CC} \delta_{\beta \eta} \delta_{\beta e} \psi_{\alpha \eta} + V_{NC} \psi_{\alpha \beta} \end{aligned} \quad (2.52)$$

Now, combining Eq. 2.49 and Eq. 2.52, we have for the full Hamiltonian

$$\langle \nu_{\beta} | H | \nu_{\alpha}(x) \rangle = \left(p + \frac{m_1^2}{2E} + V_{NC} \right) \psi_{\alpha \beta}(x) + \sum_{\eta} \sum_k \left(U_{\beta k} \frac{\Delta m_{k1}^2}{2E} U_{\eta k}^* + V_{CC} \delta_{\beta \eta} \delta_{\eta e} \right) \psi_{\alpha \eta}(x) \quad (2.53)$$

In this form, we see that the term $p + \frac{m_1^2}{2E} + V_{NC}$ is a common term to all flavor states, and does not affect the probability. It can be rotated away. Thus

$$\begin{aligned}\langle \nu_\beta | H | \nu_\alpha(x) \rangle &= \sum_\eta \sum_k \left(U_{\beta k} \frac{\Delta m_{k1}^2}{2E} U_{\eta k}^* + V_{CC} \delta_{\beta\eta} \delta_{\eta e} \right) \psi_{\alpha\eta}(x) \\ &= i \frac{d}{dx} \psi_{\alpha\beta}(x).\end{aligned}\quad (2.54)$$

If we form the vector

$$\Psi_\alpha = \begin{pmatrix} \psi_{\alpha e} \\ \psi_{\alpha\mu} \\ \psi_{\alpha\tau} \end{pmatrix}, \quad (2.55)$$

we can write the Schrödinger equation on matrix form ($i \frac{d}{dx} \Psi_\alpha = H_F \Psi_\alpha$) and compare it with Eq. 2.54 to see that the flavor Hamiltonian takes the form

$$\begin{aligned}H_F &= \frac{1}{2E} (U M^2 U^\dagger + A) \\ &= \frac{1}{2E} \left[U \begin{pmatrix} 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 \\ 0 & 0 & \Delta m_{31}^2 \end{pmatrix} U^\dagger \right] + \sqrt{2} G_F N_e \begin{pmatrix} 1 & 0 & 0 \\ 0 & 0 & 0 \\ 0 & 0 & 0 \end{pmatrix}.\end{aligned}\quad (2.56)$$

This is the three-flavor neutrino oscillation Hamiltonian that we will solve numerically to obtain the evolution of Ψ_α , whose squared components are the probabilities

$$\begin{aligned}P_\alpha = |\Psi_\alpha|^2 &= \begin{pmatrix} |\psi_{\alpha e}|^2 \\ |\psi_{\alpha\mu}|^2 \\ |\psi_{\alpha\tau}|^2 \end{pmatrix} \\ &= \begin{pmatrix} P_{\alpha e} \\ P_{\alpha\mu} \\ P_{\alpha\tau} \end{pmatrix}\end{aligned}\quad (2.57)$$

For $N_e = 0$, i.e. in vacuum, these probabilities are identical to the ones that we analytically derived in Eq. 2.28. For matter oscillations with $N_e \neq 0$, we do have closed form solutions, but they are not considered further here.

2.2.6 Two neutrino mixing

By working in the limit where one of the neutrino masses can be neglected, we can reduce the complexity of the analytical solution of neutrino oscillation. We now only consider the mixing between two neutrino flavors, say between ν_e, ν_μ , and ν_1, ν_2 . In other words, we let $\Delta m_{31}^2 = 0$.

The three-neutrino oscillation Hamiltonian from Eq. (2.56) now reduces to

$$\begin{aligned} H_F &= \frac{1}{2E}(UM^2U^\dagger + A) \\ &= \frac{1}{2E} \left[U \begin{pmatrix} 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 \\ 0 & 0 & 0 \end{pmatrix} U^\dagger \right] + \sqrt{2}G_F N_e \begin{pmatrix} 1 & 0 & 0 \\ 0 & 0 & 0 \\ 0 & 0 & 0 \end{pmatrix}. \end{aligned} \quad (2.58)$$

The rotation matrices R_{ij} that constitute U according to Eq. (2.12) can be simplified. The rotation R_{13} will not affect our mass matrix, since the former acts in the $1 - 3$ plane, and the latter sits in the $1 - 2$ plane. Thus, the reduced mass matrix \mathcal{M}_{12} commutes with R_{13} . Analogously, the rotation matrix R_{23} also commutes with \mathcal{M}_{12} . Moreover, since $SO(3)$ is abelian, its generators commute, i.e.

$$[R_{ij}, R_{kl}] = R_{ij}R_{kl} - R_{kl}R_{ij} = 0. \quad (2.59)$$

Thus,

$$\begin{aligned} U\mathcal{M}_{12}U^\dagger &= R_{23}R_{13}R_{12}\mathcal{M}_{12}R_{12}^\dagger R_{13}^\dagger R_{23}^\dagger \\ &= R_{12}R_{23}R_{13}\mathcal{M}_{12}R_{13}^\dagger R_{23}^\dagger R_{12}^\dagger \\ &= R_{12}\mathcal{M}_{12}R_{23}R_{13}R_{13}^\dagger R_{23}^\dagger R_{12}^\dagger. \end{aligned} \quad (2.60)$$

And due to the fact that the rotation matrices are unitary,

$$R R^\dagger = \mathcal{I}, \quad (2.61)$$

we have that

$$\begin{aligned} U\mathcal{M}_{12}U^\dagger &= R_{12}\mathcal{M}_{12}R_{23}R_{13}R_{13}^\dagger R_{12}^\dagger R_{23}^\dagger \\ &= R_{12}\mathcal{M}_{12}R_{23}R_{23}^\dagger R_{12}^\dagger \\ &= R_{12}\mathcal{M}_{12}R_{12}^\dagger. \end{aligned} \quad (2.62)$$

We have now completely removed the dependence of the mixing angle between the ν_1 and ν_3 mass eigenstates, θ_{13} , and the dependence of the mixing angle between the ν_2 and ν_3 mass eigenstates, θ_{23} . We can see this clearly by writing our reduced mixing matrix, which is only constituted by the rotation in the $1 - 2$ plane, R_{12} .

$$U = R_{12} \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \quad (2.63)$$

Let us return to the Hamiltonian with our two-neutrino mixing matrix. We have

$$\begin{aligned} H_F &= \frac{1}{2E} \left[R_{12} \begin{pmatrix} 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 \\ 0 & 0 & 0 \end{pmatrix} R_{12}^\dagger \right] + \sqrt{2}G_F N_e \begin{pmatrix} 1 & 0 & 0 \\ 0 & 0 & 0 \\ 0 & 0 & 0 \end{pmatrix} \\ &= \frac{1}{2E} \left[\begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 \\ 0 & 0 & 0 \end{pmatrix} \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix}^\dagger \right] + \sqrt{2}G_F N_e \begin{pmatrix} 1 & 0 & 0 \\ 0 & 0 & 0 \\ 0 & 0 & 0 \end{pmatrix}. \end{aligned} \quad (2.64)$$

We can reduce this down to two dimensions by extracting the upper-left block of all matrices. We now have the final two-neutrino Hamiltonian:

$$\begin{aligned} H_F &= \frac{1}{2E} \left[\begin{pmatrix} c_{12} & s_{12} \\ -s_{12} & c_{12} \end{pmatrix} \begin{pmatrix} 0 & 0 \\ 0 & \Delta m_{21}^2 \end{pmatrix} \begin{pmatrix} c_{12} & s_{12} \\ -s_{12} & c_{12} \end{pmatrix}^\dagger \right] + \sqrt{2}G_F N_e \begin{pmatrix} 1 & 0 \\ 0 & 0 \end{pmatrix} \\ &= \frac{1}{2E} \left[\begin{pmatrix} 0 & \Delta m_{21}^2 s_{12} \\ 0 & \Delta m_{21}^2 c_{12} \end{pmatrix} \begin{pmatrix} c_{12} & -s_{12} \\ s_{12} & c_{12} \end{pmatrix} \right] + \sqrt{2}G_F N_e \begin{pmatrix} 1 & 0 \\ 0 & 0 \end{pmatrix} \\ &= \frac{1}{2E} \left[\begin{pmatrix} s_{12}^2 + \sqrt{2}G_F N_e & \Delta m_{21}^2 s_{12} c_{12} \\ \Delta m_{21}^2 s_{12} c_{12} & \Delta m_{21}^2 c_{12}^2 \end{pmatrix} \right] \end{aligned} \quad (2.65)$$

This matrix has eigenvalues which we identify as the effective two-neutrino squared-mass difference $(\Delta m_{21}^2)^M$. They are

$$\pm(\Delta m_{21}^2)^M = \pm\sqrt{(\Delta m_{21}^2 \cos(2\theta_{12}) - \sqrt{2}G_F N_e)^2 + (\Delta m_{21}^2 \sin(2\theta_{12}))^2} \quad (2.66)$$

The matrix is diagonalized by the usual rotation matrix R_{12} , but with a different rotation angle. This now introduces the effective two-neutrino mixing angle θ_{12}^M :

$$U^M = R_{12}^M = R_{12}(\theta_{12}^M). \quad (2.67)$$

U^M diagonalizes H_F if and only if

$$\tan(2\theta_{12}^M) = \frac{\tan(2\theta_{12})}{1 - \frac{\sqrt{2}G_F N_e}{\Delta m_{21}^2 \cos(2\theta_{12})}}. \quad (2.68)$$

Eq. 2.68 has a pole at

$$N_e^* = \frac{\Delta m_{21}^2 \cos(2\theta_{12})}{\sqrt{2}G_F}. \quad (2.69)$$

Expressing the electron density N_e^* in terms of matter density ρ^* , electron number per nucleon Y , and Avogadro's constant N_A , we obtain

$$\rho^* = \frac{\Delta m_{21}^2 \cos(2\theta_{12})}{\sqrt{2}G_F Y N_A}. \quad (2.70)$$

Now assume that the matter density is sufficiently smooth along the trajectory x , that is $d\theta_M/dx \approx 0$. In its diagonalized form, the Schrödinger equation with two-neutrino Hamiltonian can be solved with this assumption. The solutions are the regular neutrino vacuum oscillations of Eq. 2.28 using $\theta_{13} = \theta_{23} = 0$, but with the effective parameters of Eqs. (2.66, 2.68). That is

$$P_{e\mu}^M(x) = \sin^2(2\theta_{12}^M) \sin^2\left(\frac{\Delta m_{21}^2 x}{4E}\right). \quad (2.71)$$

While two-neutrino flavor oscillation can only be assumed to hold in regions where mixing into the third flavor can be considered negligible, the fact that the mixing angle in matter depends on the squared-mass difference, and vice versa, gives us an interesting conclusion which is still valid for the full three neutrino oscillations. We previously saw that the oscillation amplitudes were entirely dependant on the mixing angles, and the oscillation frequencies were entirely dependant on the squared-mass difference along with energy, and trajectory. After introducing the matter potentials, we now see that the matter effect mixes the mixing matrix elements together with the masses.

2.2.7 Earth propagation

We now need to know how the electrons are distributed within the Earth. The Preliminary Earth Reference Model [4] gives us spherically symmetric piecewise polynomials for the Earth density in gcm^3 shown in the left panel of Fig. 2.3. We note a steep discontinuity at 3480 km where the density is nearly halved. This is the core-mantle boundary and will be visible in our oscillations.

Using a value of $Y = 0.5$ electron per nucleon, we express the matter potential as

$$\begin{aligned} V_{CC} &= \sqrt{2}G_F N_e = \sqrt{2}G_F Y N_A \rho \\ &= 3.8 \times 10^{-23} \text{ eV} \times \rho, \end{aligned} \quad (2.72)$$

a low number due to the smallness of G_F . This is plotted in the right panel of Fig. 2.3

Now, solving the Schrödinger equation with the Hamiltonian from Eq. 2.56 with the matter potential from the PREM, we obtain the nine combinations of probabilities through the Earth diameter. We note that due to our assumption of CP-

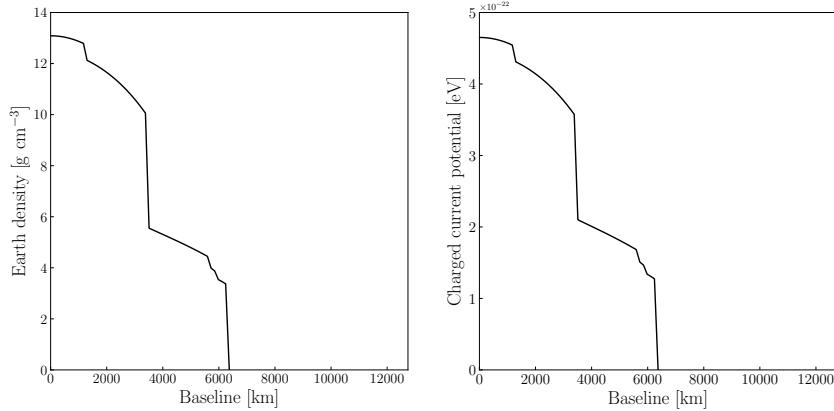


Figure 2.3: *Left panel:* The spherically symmetric Earth density according to PREM [4], as of distance from the core. *Right panel:* V_{CC} using the PREM density and 1/2 electrons per nucleon.

invariance (and thus, T-invariance), the probabilities $P_{\alpha\beta}$ and $P_{\beta\alpha}$ are equivalent. The result for GeV neutrinos is shown in Fig. 2.4.

Now we need to incorporate the *zenith angle*, here defined as the angle between the neutrino direction of travel and south. This way, neutrinos traveling through the entire diameter of the Earth up through the South Pole are defined as ‘up-going’, while neutrinos that start at the South Pole are ‘down-going’. We will mostly work with the quantity $\cos(\theta_z)$. Since we are interested in the matter effects, some studies are only looking at up-going neutrinos, i.e. neutrinos with zenith angle $-1 \leq \cos(\theta_z) \leq 0$. However, while down-going neutrinos don’t experience matter oscillations (and hardly even vacuum oscillations due to the short baseline), they can be included to reduce systematic errors. We now supplement our probability grid from Fig. 2.4 with the zenith dimension, allowing us to fully see the Earth matter effect on the oscillations in full. This is shown in Fig. 2.5.

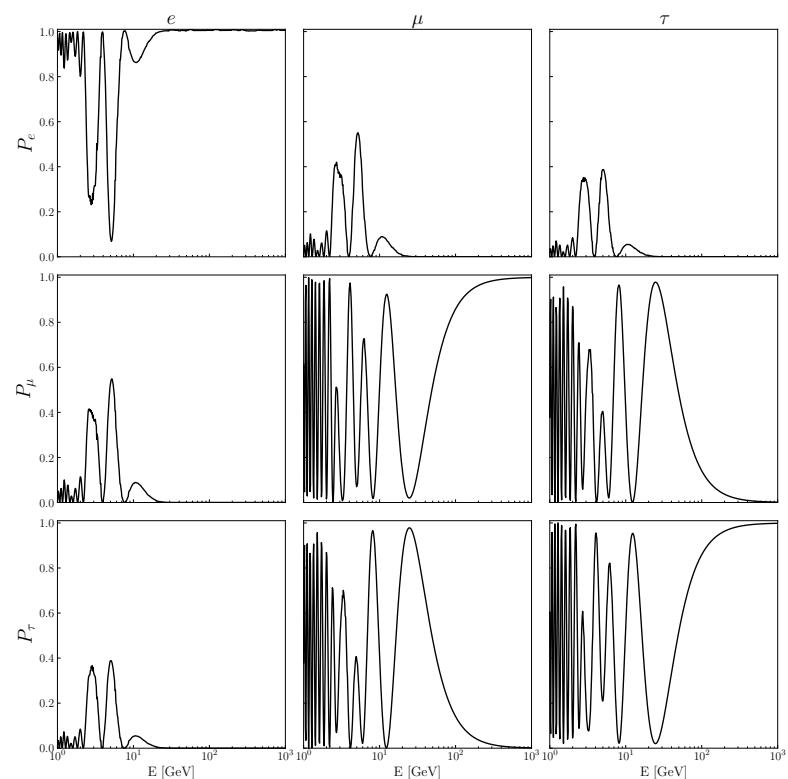


Figure 2.4

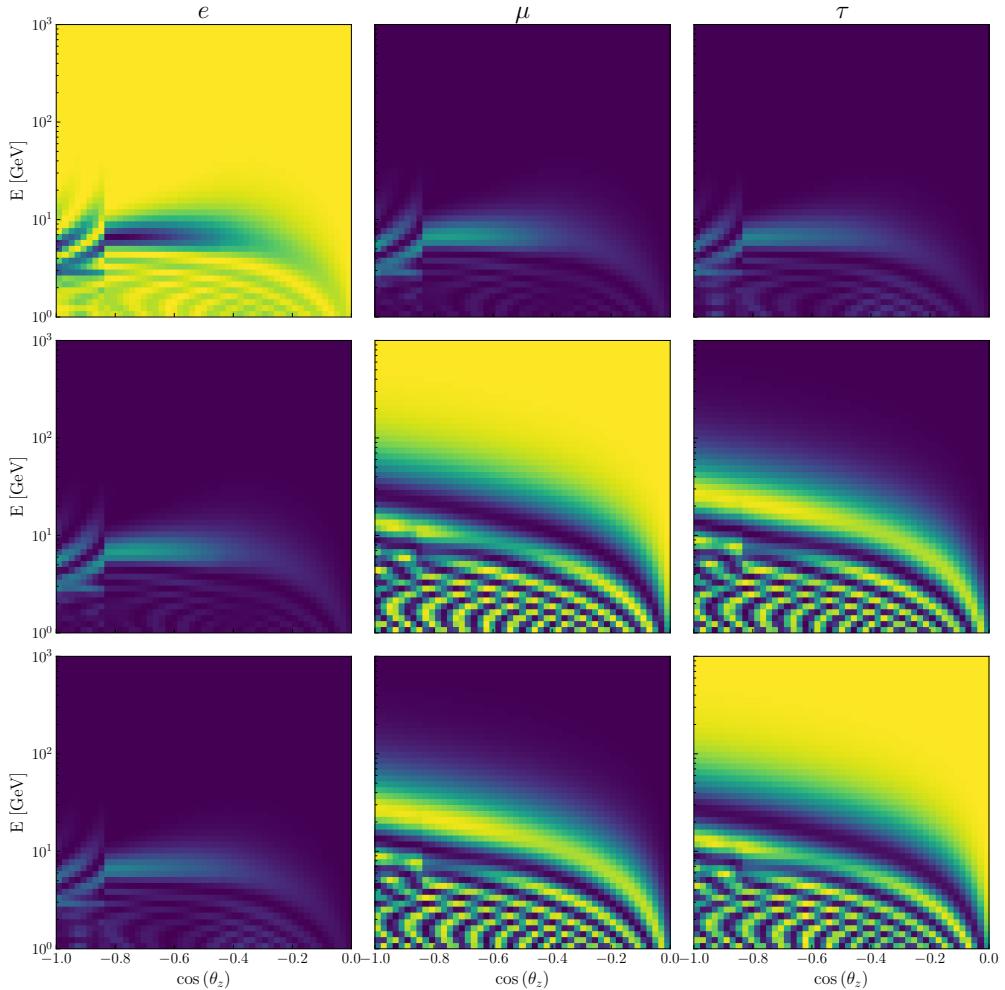


Figure 2.5: Oscillograms showing neutrino oscillations for all flavors in $E - \cos(\theta_z)$ space. Note the panels are symmetric about the diagonal panels, i.e. $P_{\alpha\beta} = P_{\beta\alpha}$ due to $\delta_{CP} = 0$. The core-mantle boundary from Fig. 2.3 is clearly displayed at $\cos(\theta_z) = -0.83$ as a sharp discontinuity for all flavors.

Chapter 3

The Antarctic Detectors

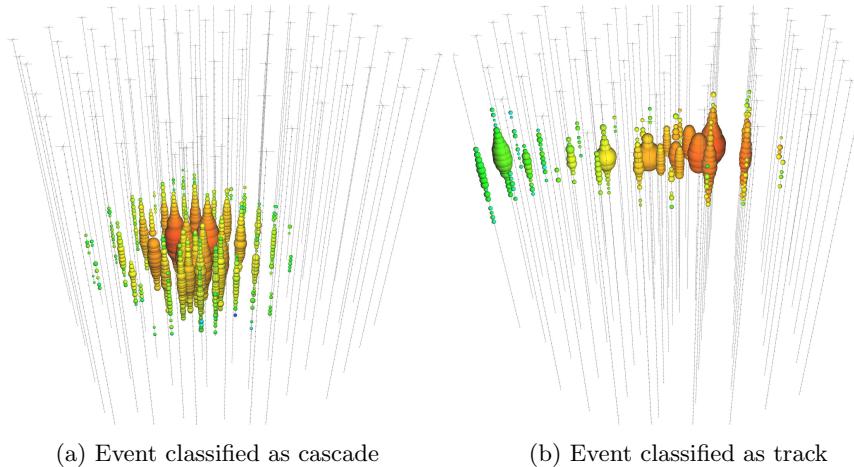


Figure 3.1: The two event types distinguished in the IceCube detector.

3.1 Neutrino detection

We always observe neutrinos indirectly through their associated charged lepton. Regardless of the type of interaction (charged current via the W boson, or neutral current via the Z), a charged lepton exits with altered properties. The lepton is then detected, and the properties of the neutrino involved in the interaction are then deduced. This deduction will be imperfect, and this introduces complexities that we will handle in Ch. 3.2.

In this work, we only study the detectors handled by the IceCube collaboration. They are of Cherenkov type, which means that they detect the secondary charged lepton by its emitted Cherenkov light, produced from its travel through the Antarctic ice. If the charged leptons interact heavily with the ice, they will travel a short distance and emit a localized flash of Cherenkov light. This event is referred to as a cascade. The neutral current interactions involve quarks, which recoils and produces showers of hadrons. Also, charged current ν_e interactions also produce cascades. A cascade event is shown in Fig. 3.1a. If the charged leptons don't interact as much in the ice, they penetrate a larger part of it, emitting light and tertiary particles as they go. This event is referred to as a track and is often due to charged current muon interactions. A track event is shown in Fig. 3.1b.

To detect the Cherenkov light, 60 Digital Optical Modules (DOMs) are placed on a long string up to 17 m apart. 86 of these strings are then lowered into 2.5 km deep boreholes in the ice. The holes are then sealed by refreezing the ice, resulting in a total of 5160 DOMs in a volume of approximately 1 km^3 [5].

The strings and DOMs are not spaced evenly, making some parts of the detector more sensitive to certain energy ranges than others. 8 strings packed more tightly than the other 78, making that part of the detector sensitive to neutrino energies down to single-digit GeV. Due to this part being situated deep within the ice, it

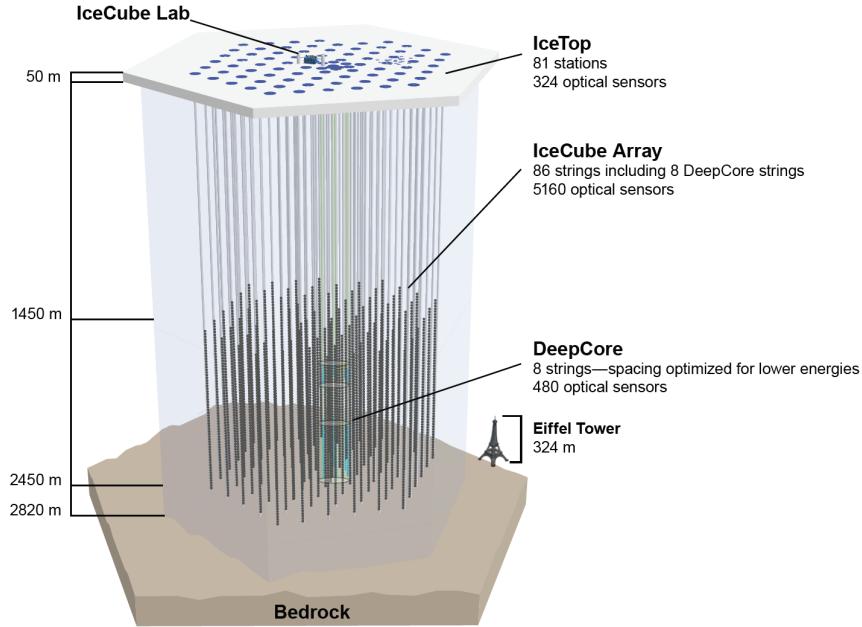


Figure 3.2: View of the full IceCube array, with the Eiffel tower for scale.

is referred to DeepCore. DeepCore will be treated as a separate and independent detector from the rest, which retains the name IceCube. A view of the current setup can be seen in Fig. 3.2. In this work, we consider DeepCore data between 5.6 GeV to 56 GeV and IceCube data in the range 0.5 TeV to 10 TeV.

In 2017, the PINGU Letter of Intent was published [6]. The ‘Precision IceCube Next Generation Upgrade’ is an upgrade that will supplement DeepCore, i.e. boosting the capabilities of neutrino detection at the GeV scale. As the PINGU upgrade is not yet financed nor built, we are not able to use any data from it. However, the collaboration has released preliminary simulations which we will use to see how the upgrade might improve results from IceCube and DeepCore. The PINGU simulations have the same structure as the DeepCore data, so our analysis referring to DeepCore will also apply to PINGU except where noted. However, we treat the PINGU detector as independent of the other experiments.

3.1.1 Atmospheric Neutrino Flux

Atmospheric neutrinos originate from cosmic rays composed of high-energy protons interacting with nuclei in the atmosphere. These interactions ultimately produces pions, which decay as

$$\begin{aligned} \pi^+ &\rightarrow \mu^+ + \nu_\mu, & \pi^- &\rightarrow \mu^- + \bar{\nu}_\mu \\ \pi^+ &\rightarrow e^+ + \nu_e, & \pi^- &\rightarrow e^- + \bar{\nu}_e. \end{aligned} \quad (3.1)$$

In the muonic decay channel, muons are emitted which will be detected by the IceCube detector. The detector's experience an uncorrelated systematic error, which comes from this *muon background*, i.e. events misclassified as muons from ν_μ interactions rather than from pion decay. For the DeepCore analysis, we will have to consider this background when calculating the events. For IceCube events, we will scan a higher energy range where the muon background can be neglected. For the PINGU events, the IceCube detector is expected to be able to act as a veto for this background. Thus, the error introduced from the muon background is expected by the collaboration to be negligible [6]. Moreover, the atmospheric flux itself is often associated with a large error. In this work, we will use a flux normalization error of 25%, and a zenith slope error of 5% [7].

The flux is provided in [8, 9], and a selection is shown in Table 3.1. The flux data is binned in $\cos(\theta_z)$. The fluxes are averaged over the azimuthal direction and solar minimum/maximum. The units of the fluxes are given as $\text{GeV}^{-1} \text{m}^{-2} \text{s}^{-1} \text{sr}^{-1}$ and are omitted from the table for clarity. We note that the fluxes for ν_τ and $\bar{\nu}_\tau$ are missing. Kaons can decay into neutral pions, which in turn can produce ν_τ , but this branching ratio is extremely small. Thus, we never have to use probabilities in the form $P_{\tau\beta}$, since we have no starting atmospheric ν_τ flux. Interpolating the data yields makes us capable of approximating all four necessary fluxes for a given true energy and true zenith. The result of ϕ_μ is shown in Fig. 3.3. We see that the flux is almost entirely zenith independent. This is due to the very small cross-section of the neutrinos. For particles interacting through the electromagnetic interaction, we would expect to see a large difference between the flux of particles traversing through the whole diameter of the Earth, compared to particles that only penetrate a small rim of the crust. Here, however, the flux at the South Pole is almost the same! This does not mean that the oscillation probability will be equal. First, the difference in the density profiles introduces wildly different matter effects. Secondly, the zenith angle determines the baseline. Completely down-going neutrinos, with $\cos(\theta_z) = 0$ will be produced at an average distance of 15 km overhead, travel down to the icy sheet of Antarctica, and then penetrate the ice down 2 km to the DOMs [7]. The energy dependence of the flux is approximately $E^{-2.7}$, so our expected event count at higher energies will suffer from this strong decline in flux.

We propagate the atmospheric neutrino flux $\phi_\alpha^{\text{atm}}(E, \theta_z)$ through the Earth. The oscillation probability $P_{\alpha\beta}$ acts as a weight to the atmospheric flux, yielding the propagated flux for flavor β at detector level as

$$\phi_\beta^{\text{det}} = \sum_\alpha P_{\alpha\beta} \phi_\alpha^{\text{atm}}, \quad (3.2)$$

where we sum over the initial lepton flavors $\alpha \in \{e, \mu, \bar{e}, \bar{\mu}\}$.

E [GeV]	ϕ_μ	$\phi_{\bar{\mu}}$	ϕ_e	$\phi_{\bar{e}}$	$\cos(\theta_z)$
27825	6.06×10^{-12}	3.17×10^{-12}	1.56×10^{-13}	1.04×10^{-13}	[-0.2, -0.1]
247707	5.94×10^{-16}	2.92×10^{-16}	1.36×10^{-17}	8.12×10^{-18}	[-0.7, -0.6]
22	3.33×10^{-2}	2.78×10^{-2}	9.57×10^{-3}	7.15×10^{-3}	[-0.3, -0.2]
432876	5.19×10^{-17}	2.32×10^{-17}	1.46×10^{-18}	9.83×10^{-19}	[-1.1, -1.0]
64280	1.58×10^{-13}	8.10×10^{-14}	3.49×10^{-15}	2.21×10^{-15}	[-0.4, -0.3]

Table 3.1: A selection of processed atmospheric South Pole fluxes from [8] by Honda et al. [9].

3.1.2 Event Reconstruction

After an event has occurred, the IceCube algorithms process the data coming from the detector to *reconstruct* the event. This means that, given the parameters recorded by the detector, what are their ‘true’ values? We are interested in two variables: energy and direction. Each event is tagged with a probable energy and zenith angle, called the reconstructed parameters E^{reco} and $\cos(\theta_z^{reco})$, which are the parameters according to the DOMs. The collaboration then uses numerous sophisticated methods to backtrack the reconstructed parameters to the true parameters. So a charged lepton hits the DOMs, and we ultimately end up with the associated neutrino’s true energy and zenith angle, E^{true} and $\cos(\theta_z^{true})$. The reconstructed parameters are what we are using to analyze the data (because this is what the detector sees), while the true parameters are used in the determination of that neutrino’s ‘actual’ flux and cross-section (because this is what Nature sees).

How do we then translate between the reconstructed and true parameters? In this work, we are using two different methods, which are based on the form of data available to us. They will be outlined in Sec. 3.2 and Sec. 3.3.

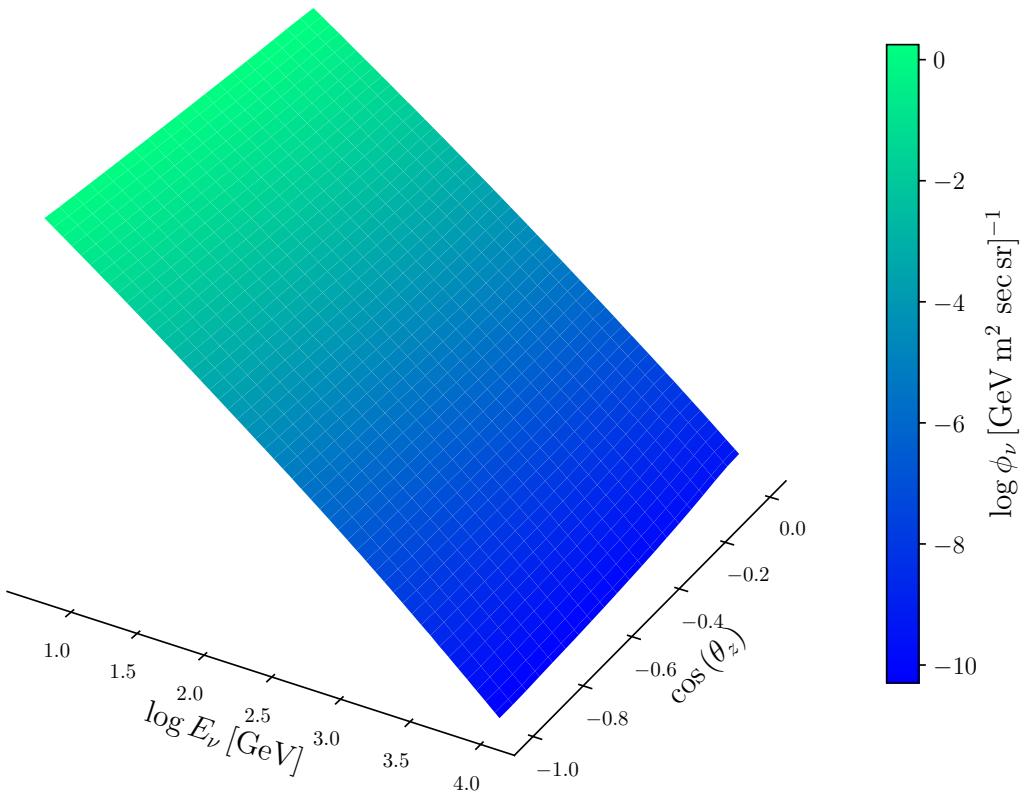


Figure 3.3: Interpolated South Pole atmospheric ν_μ flux with data from [9], displaying a near-negligible zenith angle dependence, but a strong energy dependence. The energy dependence can be parametrized to be approximated by $E^{-2.7}$.

3.2 IceCube

As the neutrinos have propagated the Earth, they arrive at the South Pole, where they interact with charged leptons in the ice. The charged lepton then emits the Cherenkov light that is detected by the array. We now are interested in the effective area A^{eff} , i.e. the cross-sectional area of the detector that the lepton is exposed to. A^{eff} depends on several parameters, some of them being detector physical volume, E^{true} , $\cos(\theta_z^{\text{true}})$, and the particle type. Fortunately, the binned A^{eff} is provided to us by the collaboration [10]. An excerpt from this data is shown in Table 3.2.

$E_{\text{min}}^{\text{true}}$ [GeV]	$E_{\text{max}}^{\text{true}}$ [GeV]	$\cos(\theta_z^{\text{true}})_{\text{min}}$	$\cos(\theta_z^{\text{true}})_{\text{max}}$	A^{eff} [m ²]
251	316	-0.92	-0.91	0.0174
794300	1000000	-0.80	-0.79	69.3600
3981	5012	-0.78	-0.77	3.1490
1585	1995	-0.07	-0.06	0.4659
398	501	-0.73	-0.72	0.0555

Table 3.2: IceCube-86 effective area from [10].

Here, A^{eff} has been averaged over A_{μ}^{eff} and $A_{\bar{\mu}}^{\text{eff}}$ by the collaboration. Thus, both μ and $\bar{\mu}$ will on average, experience the same A^{eff} in our model. Just as with the fluxes, we interpolate this in E^{true} and $\cos(\theta_z^{\text{true}})$, and show the result in Fig. 3.4. Since the IceCube array is slightly rectangular, the zenith angle affects the cross-sectional area to which the array the neutrinos are exposed to. While the flux was almost flat in $\cos(\theta_z^{\text{true}})$, the introduction of the zenith dependent A^{eff} will make the result slightly more zenith dependent than the flux itself. Energy-wise, the situation here is completely reverse from the flux. Increasing with energy, the effective area of the detector approaches its geometrical area of 1×10^6 m² but is still only in the single-digit range at TeV energies.

So now we have the physical quantities in the true parameters. But as we discussed, we need a way to translate this into the reconstructed parameters that the detector gives us. We will call the relationship between E^{reco} and E^{true} the energy resolution function, and the relationship between $\cos(\theta_z^{\text{reco}})$ and $\cos(\theta_z^{\text{true}})$ the zenith resolution function. We assume the relationship to follow a logarithmic Gaussian distribution, giving it the form

$$R(x^r, x^t) = \frac{1}{\sqrt{2\pi}\sigma_{x^r}x^r} \exp\left[-\frac{(\log x^r - \mu(x^t))^2}{2\sigma_{x^r}^2}\right]. \quad (3.3)$$

The parameters of the Gaussian are $\sigma_{x^r}(x^t)$ and $\mu(x^t)$, which are functions of the true parameters. By multiplying the Gaussian in Eq. 3.3, we are reweighing the values by the probability density of that point. This process is also called *smearing* because it effectively spreads out the data around a certain point.

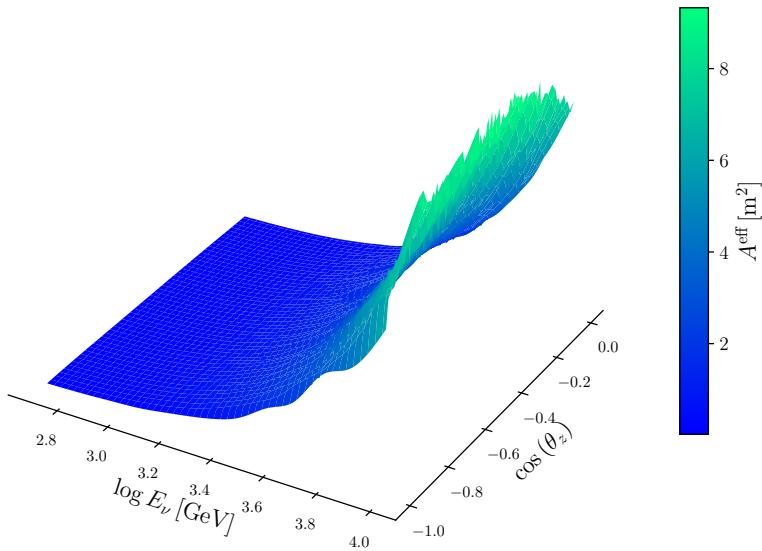


Figure 3.4: Interpolated IceCube effective area with data from [10].

So how do we then obtain $\sigma_{x^r}(x^t)$ and $\mu(x^t)$ needed to construct the Gaussian? A Monte Carlo sample publically released by the collaboration has all the ingredients that we need [11]. In Table. 3.3 we show a selection of the data. The "pdg" column refers to the Monte Carlo particle classification, where 13 is the tag for ν_μ , while -13 refers to an $\bar{\nu}_\mu$. Here we note a crucial property of the IceCube dataset that will impact our analysis: the MC released by the collaboration only includes simulated muon events.

pdg	E^{reco} [GeV]	$\cos(\theta_z^{reco})$	E^{true} [GeV]	$\cos(\theta_z^{true})$
13	1665	-0.645884	592	-0.653421
13	587	-0.373241	342	-0.424979
-13	1431	-0.177786	1169	-0.189949
-13	831	-0.807226	1071	-0.805559
13	988	-0.370746	1861	-0.367922

Table 3.3: A selection of the data found in ??

First, we let $\cos(\theta_z^{reco}) = \cos(\theta_z^{true})$ for all values. The angular resolution in IceCube for track-like events is less than 2° , making $\cos(\theta_z^{true})$ coincide with $\cos(\theta_z^{reco})$ for our study [12]. Thus, we only need to concern ourselves with the energy resolution. In Fig. 3.5, we have plotted all event counts found in the MC file, over 8 million. However, this is too much data to process efficiently, with many outliers that ultimately don't weigh in that much in the final event count. To resolve this, we have opted to train a Gaussian process regressor on the dataset, from which we can extract the predicted mean and standard deviation for a point. When doing this over E^{reco} , we sample E^{true} in the 99th percentile around the predicted mean. We then obtain the shaded band shown in Fig. 3.5.

The event rate for each bin reads

$$N_{ij} = T \sum_{\beta} \int_{(\cos \theta_z^r)_i}^{(\cos \theta_z^r)_{i+1}} d \cos \theta_z^r \int_{E_j^r}^{E_{j+1}^r} d E^r \int_0^\pi R(\theta^r, \theta^t) d \cos \theta^t \int_0^\infty R(E^r, E^t) \phi_{\beta}^{\text{det}} A_{\beta}^{\text{eff}} d E^t, \quad (3.4)$$

where T is the live time of the detector. Now this expression handles the Gaussian smearing, but we are not provided systematic error sources, DOM efficiencies, and other nuisance parameters. To correct this, we will aim to come as close as possible to the IceCube Monte Carlo, and then normalize with it. That way, we know that our null hypotheses will align while we are free to form additional hypotheses with different physics parameters.

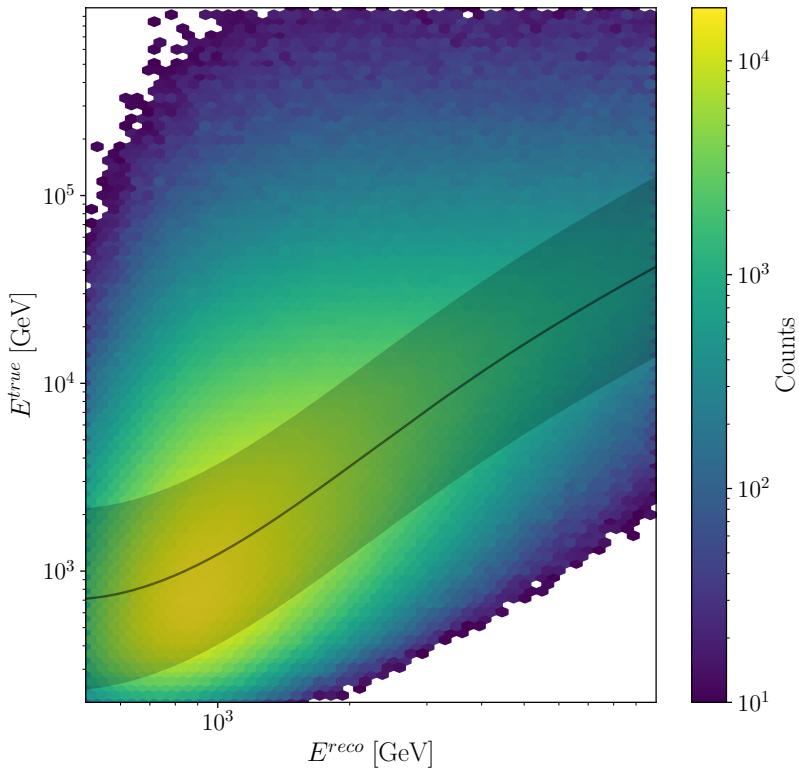


Figure 3.5: Relationship between the true and reconstructed muon energy in the IceCube MC sample [11]. Color indicates the frequency of each simulated point. Shaded area shows the 99th percentile limits predicted by the regressor trained on this set. It is within this band that then will sample the E^{true} values for each E^{reco} .

The latest available data collected and processed by the collaboration contains 305,735 muon track events, collected over eight years [12]. The data has 13 logarithmically spaced bins in $E^{reco} \in [500, 9976]$ GeV, and 20 linear bins in $\cos(\theta_z^{reco}) \in [-1, 0]$. The data is shown in Fig. 3.6.

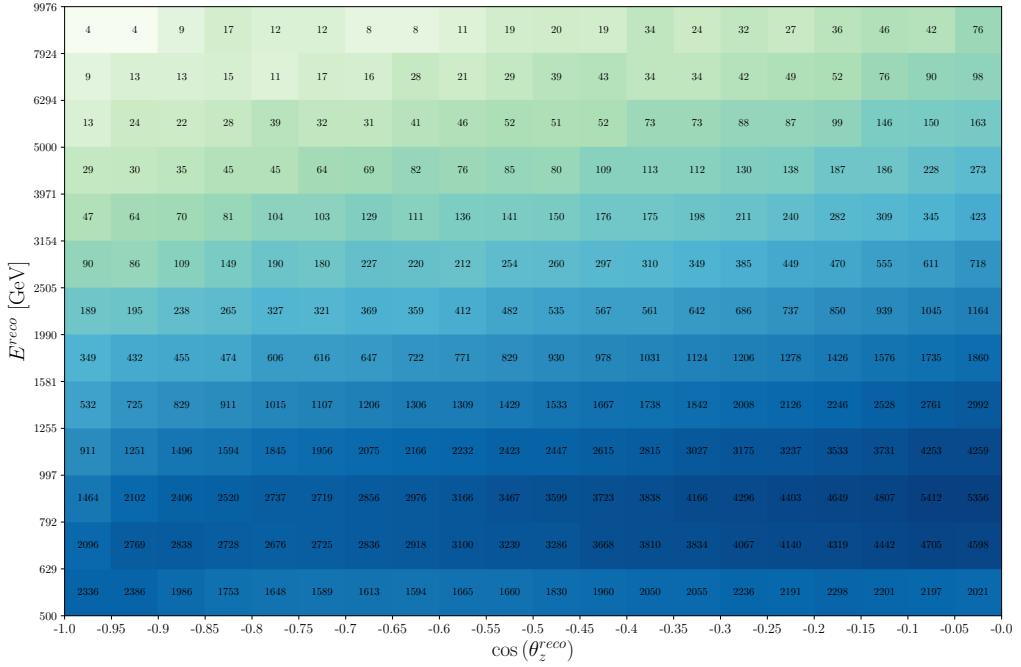


Figure 3.6: IceCube track events from [12].

Monte Carlo normalization

Independent researchers outside of the IceCube collaboration will not be able to more precisely simulate the detector. The IceCube Monte Carlo is a complex and proprietary machinery, so our goal in this section is to come as close as we can to their Monte Carlo simulations. After we are confident that our code displays the same overall features as the ‘official’, we normalize our results N_{ij}^{sim} as

$$N_{ij} = \frac{N_{ij}^{\text{null}}}{N_{ij}^{\text{MC}}} N_{ij}^{\text{sim}}. \quad (3.5)$$

For each bin i, j , we then obtain a correction factor that contains information that we are unable to obtain or sufficiently incorporate. One example of such information is the systematic errors of the DOMs. Recent IceCube data releases do not include such information. Since the systematic errors are affecting the event count on a bin-by-bin basis, they can in theory drastically modify the binned results. Another example of an error source that will be remedied by this method is the flux. We are using a fairly simple model of the atmospheric flux that excludes atmospheric prompt and astrophysical fluxes. The IceCube collaboration uses several different flux models which are initialized by a parametrization of the cosmic ray flux.¹

In Fig. 3.7, we present the Icecube Monte Carlo obtained from their 2020 sterile analysis [12], along with our null hypothesis times a constant factor. We deemed

¹Included in the cosmic ray models are e.g. the pion to kaon ratio, which are often used as a nuisance parameter. By not being able to include this in our error analysis, our method will be limited to only consider the overall flux normalization, rather than the components that produce the flux in the first place.

these shapes to be satisfactory, thus allowing us to multiply Eq. 3.4 by the correction factors of Eq. 3.5. We now arrive to our final event count

$$N_{ij} = \frac{N_{ij}^{\text{null}}}{N_{ij}^{\text{MC}}} T \sum_{\beta} \int_{(\cos \theta_z^r)_i}^{(\cos \theta_z^r)_{i+1}} d \cos \theta_z^r \int_{E_j^r}^{E_{j+1}^r} d E^r \int_{E_{min}^t}^{E_{max}^t} R(E^r, E^t) \phi_{\beta}^{\text{det}} A_{\beta}^{\text{eff}} d E^t , \quad (3.6)$$

with E_{min}^t , E_{max}^t , and $R(E^r, E^t)$ taken from the Gaussian process regressor. Thus,

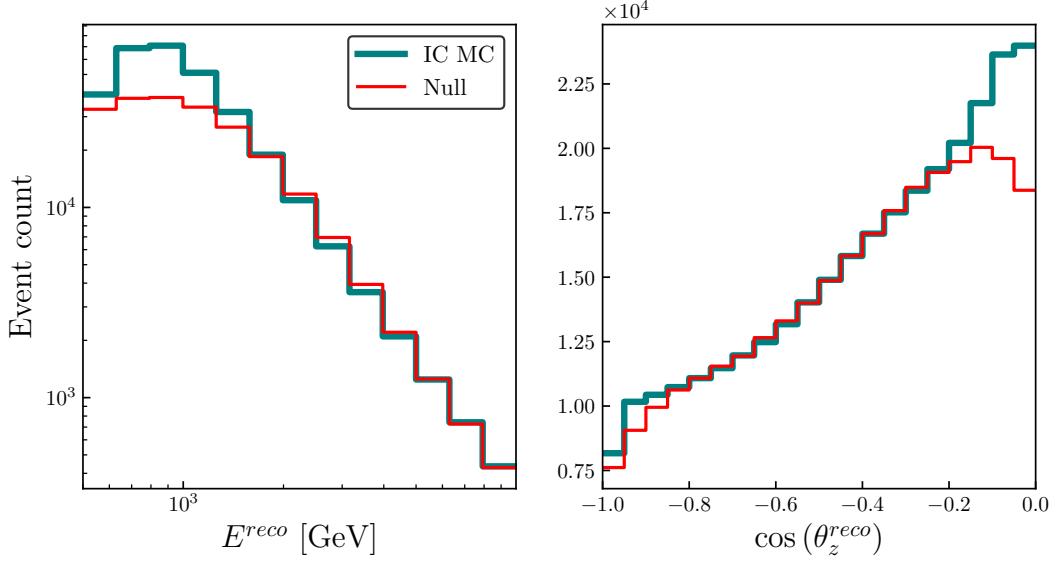


Figure 3.7: IceCube Monte Carlo, binned in E^{reco} and $\cos(\theta_z^{reco})$. We compare this with our simulations shown as ‘Null’ in the plots. We see a large deviation for high zenith values, but since these neutrinos have traversed less matter and have a smaller baseline, we deem this to be acceptable.

we are now able to sufficiently approximate the IceCube Monte Carlo, which makes us able to run simulations based on different physics scenarios.

3.3 DeepCore

In this part, we use the publically available DeepCore data sample [13] which is an updated version of what was used by the IceCube collaboration in a ν_μ disappearance analysis [14].

The detector systematics include ice absorption and scattering, and overall, lateral, and head-on optical efficiencies of the DOMs. They are applied as correction factors using the best-fit points from the DeepCore 2019 ν_τ appearance analysis [15].

The data include 14901 track-like events and 26001 cascade-like events, both divided into eight $\log_{10} E^{reco} \in [0.75, 1.75]$ bins, and eight $\cos(\theta_z^{reco}) \in [-1, 1]$ bins. Each event has a Monte Carlo weight $w_{ijk,\beta}$, from which we can construct the event count as

$$N_{ijk} = C_{ijk} \sum_{\beta} w_{ijk,\beta} \phi_{\beta}^{\text{det}}, \quad (3.7)$$

where $C_{k\beta}$ is the correction factor from the detector systematic uncertainty and $\phi_{\beta}^{\text{det}}$ is defined as Eq. 3.2. We have now substituted the effect of the Gaussian smearing by treating the reconstructed and true quantities as a migration matrix.

The oscillation parameters used on our DeepCore simulations are from the best-fit in the global analysis in [3]: $\theta_{12} = 33.44^\circ$, $\theta_{13} = 8.57^\circ$, $\Delta m_{21}^2 = 7.42 \text{ eV}^2$, and we marginalize over Δm_{31}^2 and θ_{23} .

3.4 PINGU

The methodology behind the PINGU simulations is the same as with our DeepCore study . We use the public MC [16], which allows us to construct the event count as in Eq. 3.7. However, since no detector systematics is yet modeled for PINGU, the correction factors C_{ijk} are all unity. As with the DeepCore data, the PINGU Monte Carlo is divided into eight $\log_{10} E^{reco} \in [0.75, 1.75]$ bins, and eight $\cos(\theta_z^{reco}) \in [-1, 1]$ bins for both track- and cascade-like events. We generate ‘data’ by predicting the event rates at PINGU with the following best-fit parameters from [3], except for the CP-violating phase which is set to zero for simplicity.

$$\begin{aligned} \Delta m_{21}^2 &= 7.42 \times 10^{-5} \text{ eV}^2, & \Delta m_{31}^2 &= 2.517 \times 10^{-3} \text{ eV}^2, \\ \theta_{12} &= 33.44^\circ, & \theta_{13} &= 8.57^\circ, & \theta_{23} &= 49.2^\circ, & \delta_{\text{CP}} &= 0. \end{aligned} \quad (3.8)$$

Chapter 4

Beyond the 3ν Picture

$$\mathcal{L}_{\text{NC}} = -2\sqrt{2}G_F \epsilon_{\alpha\beta}^{fX} (\bar{\nu}_\alpha \gamma^\mu P_L \nu_\beta)$$

4.1 The Sterile State

In 1996, the LSND experiment reported an excess of $\bar{\nu}_e$ events from an $\bar{\nu}_\mu$ beam [17]. Nine years later, MiniBooNE not only reproduced the $\bar{\nu}_e$ anomaly but observed an excess in the ν_e events too. Together with the so-called reactor and gallium anomalies, these reports suggested that the three massive neutrino framework could be amended. By introducing a fourth heavy mass state ν_4 , both appearance and disappearance anomalies could be minimally accommodated. However, we know from the decay width of the Z boson that it only can interact with three flavor species¹, so this fourth mass state cannot be interacting weakly. In other words, it needs to transform as a singlet under the broken gauge group $SU(2)_L \times U(1)_{EM}$. We now distinguish between the three original neutrino flavors (e , μ , and τ) and the new fourth flavor by calling the former *active* neutrinos and the latter *sterile* due to its non-interacting behavior.

The experiments listed above indicate that the mass-squared difference of the sterile neutrino is in the eV² scale, while the two others are three and five magnitudes smaller. To remind us of this hierarchy, we write $3 + 1$, since the sterile state proposed is heavier than the active states.

Two neutrino physics parameters probed by CMB lensing observations are the effective number of neutrinos N_{eff} , and the sum of neutrino masses $\sum m_i$. The 2018 results from the Planck Collaboration in [19] constrain the parameters to

$$N_{\text{eff}} = 2.96^{+0.34}_{-0.33}, \quad \sum m_i < 0.12 \text{ eV} \quad (95\%), \quad (4.1)$$

consistent with the Standard Model with oscillations² prediction of $N_{\text{eff}} = 3.045$ degrees [20]. Assuming the Planck data, one thermalized sterile neutrino ($N_{\text{eff}} = 4$) is excluded at 6σ . Thus, thermalized sterile neutrinos are strongly in tension with cosmological observations if one does not consider some special mechanism. So if a fourth neutrino species is present in nature, it will have ramifications not only on neutrino physics itself, but on our understanding of the CMB as well.

4.1.1 The Hamiltonian

The inclusion of the sterile state in the Hamiltonian is straightforward. We extend the PMNS matrix with terms for the mixing between the new flavor (U_{si}) and mass ($U_{\alpha 4}$) eigenstates:

$$U_{4\text{gen}} = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} & U_{e4} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} & U_{\mu 4} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} & U_{\tau 4} \\ U_{s1} & U_{s2} & U_{s3} & U_{s4} \end{pmatrix}, \quad (4.2)$$

¹Direct measurements of the invisible Z decay yields $N = 2.92 \pm 0.05$ [18].

²Of course, the Standard Model *without* neutrino masses dictates $N_{\text{eff}} = 3$ and $\sum m_i \equiv 0$.

where a common parametrization of this new mixing matrix is

$$U_{4gen} = R_{34}R_{24}R_{14}U_{3gen}. \quad (4.3)$$

This parametrization does not affect the physics, but determines the structure of the mixing angles.

The mass matrix extends analogously:

$$M_{4gen}^2 = \begin{pmatrix} 0 & 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 & 0 \\ 0 & 0 & \Delta m_{31}^2 & 0 \\ 0 & 0 & 0 & \Delta m_{41}^2 \end{pmatrix}. \quad (4.4)$$

Now, the interaction with matter requires a careful reconsideration of the matter potential. We start off with the unaltered potential matrix from Eq. 2.42. Just as with the PMNS matrix, we extend this to 4×4 :

$$\begin{pmatrix} V_{CC} + V_{NC} & 0 & 0 & 0 \\ 0 & V_{NC} & 0 & 0 \\ 0 & 0 & V_{NC} & 0 \\ 0 & 0 & 0 & 0 \end{pmatrix}. \quad (4.5)$$

Now we need to include the terms that describes the matter potential felt by the sterile flavor state. Recalling our discussion above, we remind ourselves that the sterile neutrino by definition does not participate in any interaction³. Thus, all potential terms involving the sterile state are zero. In other words, the potential matrix in Eq. 4.5 is complete, save for the usual subtraction by a constant times the identity matrix:

$$\begin{aligned} V_{4gen} &= \begin{pmatrix} V_{CC} + V_{NC} & 0 & 0 & 0 \\ 0 & V_{NC} & 0 & 0 \\ 0 & 0 & V_{NC} & 0 \\ 0 & 0 & 0 & 0 \end{pmatrix} - V_{NC} \begin{pmatrix} 1 & 0 & 0 & 0 \\ 0 & 1 & 0 & 0 \\ 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 1 \end{pmatrix} \\ &= \sqrt{2}G_F \begin{pmatrix} N_e & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & -N_n/2 \end{pmatrix} \\ &= \sqrt{2}G_F N_e \begin{pmatrix} 1 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & -1/2 \end{pmatrix}. \end{aligned} \quad (4.6)$$

³The exception to this is of course gravity. The sterile neutrino is not massless.

where we have assumed electrical neutrality in the last step, yielding $N_e = N_n$. Thus, the final Hamiltonian with a fourth sterile neutrino is

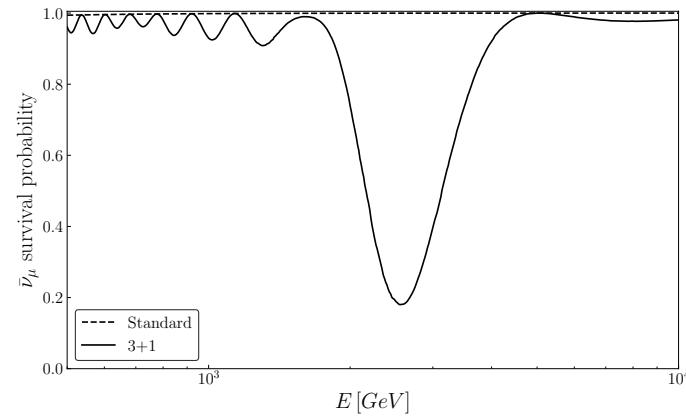
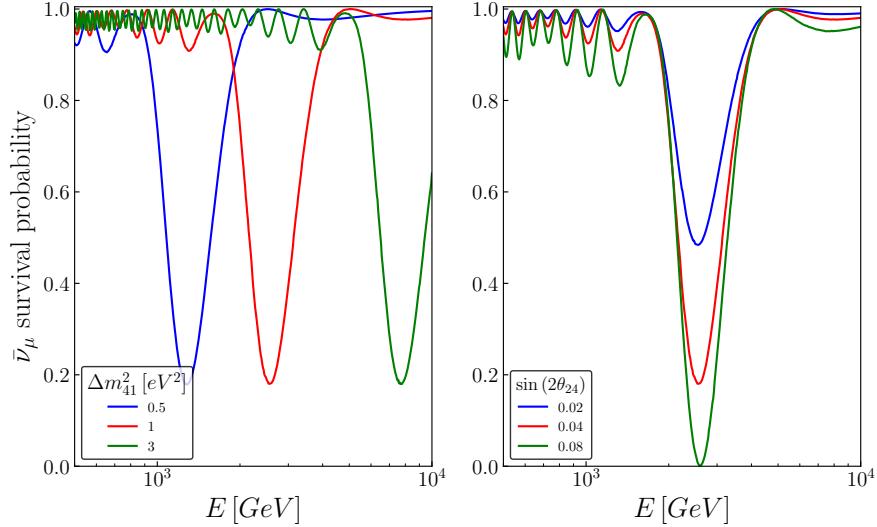
$$H_{4gen} = \frac{1}{2E} \left[U \begin{pmatrix} 0 & 0 & 0 & 0 \\ 0 & \Delta m_{21}^2 & 0 & 0 \\ 0 & 0 & \Delta m_{31}^2 & 0 \\ 0 & 0 & 0 & \Delta m_{41}^2 \end{pmatrix} U^\dagger \right] + \sqrt{2} G_F N_e \begin{pmatrix} 1 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & -1/2 \end{pmatrix}. \quad (4.7)$$

Now, looking at Eqs. 4.3 and 4.4, we see that we have introduced four new parameters: Δm_{41}^2 , θ_{14} , θ_{24} , θ_{34} . In our analysis, we reduce the number of parameters by only considering Δm_{41}^2 and θ_{24} , which are the most important for the ν_μ oscillations. In fact, one can suppress the e channel by letting $\theta_{14} = 0$ in the GeV to TeV range, thus using a two-flavor neutrino approximation.

4.1.2 Sterile Signals

Since the new neutrino does not interact weakly, how do we then detect its signal? If the sterile mixing angle θ_{i4} is non-zero, we allow the sterile mass state to mix with the active state i . The most interesting case is when $\theta_{24} \neq 0$, which for $\Delta m_{41}^2 \sim \text{eV}^2$ gives rise to a resonant disappearance in the for $P_{\bar{\mu}\bar{\mu}}$, shown in Fig. 4.1a.

So for a eV-scale sterile neutrino and a non-zero θ_{24} , we expect a TeV $\bar{\nu}_\mu$ disappearance. The resonance dip is affected by the value of Δm_{41}^2 and θ_{24} as shown in Fig. 4.1b. We see how the value of Δm_{41}^2 shifts the peak, while the mixing angle adjusts its strength. In the $3 + 0$ scenario, $P_{\bar{\mu}\bar{\mu}} = 1$ in this range. Thus, if we are able to see a $\bar{\nu}_\mu$ disappearance in the TeV region, we are in a good position to argue for a eV 2 scale Δm_{41}^2 .

(a) $P_{\bar{\mu}\bar{\mu}}$ resonant disappearance(b) Shifting and amplification of the resonance by altering Δm_{41}^2 and θ_{24} , respectively.

4.2 Non-Standard Interactions

We have no reason to believe that the Standard Model gauge group is the complete picture. We have already seen how the failure of the Standard Model to predict neutrino masses and flavor oscillations force us to amend it. Just as the electroweak theory $SU(2)_L \times U(1)_Y$ is spontaneously broken to $U(1)_{EM}$, a higher order theory at a different energy scale with completely different properties might undergo a similar spontaneous symmetry break at high energies, producing at lower energies the Standard Model that we know. In fact, just as the fermion masses originate from the electroweak symmetry breaking, the neutrino masses might be generated from another broken symmetry resulting in the Standard Model gauge group. In this sense, the Standard Model might be considered an effective low energy theory, which yields impressive results in some areas, but failing in others.

Compared to other fermions, the lightness of neutrino masses⁴ along with their sparse SM interactions might indicate that these particles provide the best starting point for us to probe new physics, in which exotic neutrino interactions might occur [21]. We call these interactions non-standard interactions (NSI) in order to distinguish them from the ‘standard’ interactions of the Standard Model.

Following the approach of the Standard Model that the group generators uniquely determine the gauge bosons, a different gauge theory will have different interactions than those we presently know. These new interactions can be parametrized as model-independent four-fermion effective operators [22, 23]. Following the discussion in [24], we see that the NSI parameters ϵ resulting from six-dimensional operators have the scale

$$\epsilon \propto \frac{m_W^2}{m_\epsilon^2} \sim \frac{10^{-2}}{m_\epsilon^2} \quad (4.8)$$

in TeV, so the new interactions generated at a mass scale of $m_\epsilon = 1$ TeV will produce parameters in the order of 10^{-2} , two magnitudes below the standard matter effect. Thus, if we assume the new interactions to arise from a higher-energy theory above electroweak scale, we then predict that the parameters contribute with at most a factor of 10^{-2} to the standard matter effect, decreasing quadratically.

4.2.1 Altering the Matter Potential

Up until now, we have only considered weak neutrino interactions with electrons, protons, and neutrons. We can phenomenologically allow these interactions to include the up and down quarks which are present in the Earth as the fundamental

⁴The extreme lightness might be explained by a mass generation at a higher order theory.

components of neutrons and protons, as seen in the Lagrangians

$$\begin{aligned}\mathcal{L}_{\text{CC}} &= -2\sqrt{2}G_F \epsilon_{\alpha\beta}^{ff'X} (\bar{\nu}_\alpha \gamma^\mu P_L \ell_\beta) (\bar{f}' \gamma_\mu P_X f) \\ \mathcal{L}_{\text{NC}} &= -2\sqrt{2}G_F \epsilon_{\alpha\beta}^{fX} (\bar{\nu}_\alpha \gamma^\mu P_L \nu_\beta) (\bar{f} \gamma_\mu P_X f),\end{aligned}$$

where CC denotes the charged current interaction with the matter field $f \neq f' \in \{u, d\}$, and NC denotes the neutral current interaction with $f \in \{e, u, d\}$. The CC NSI effect affects neutrino production and detection, and will not be considered further. NC NSI affect the matter potential, and is thus of interest to us.

We have no independent sensitivity for the neither chirality nor flavor type of ϵ^X , so we sum over these and study the effective matter NSI parameter $\epsilon_{\alpha\beta}$:

$$\epsilon_{\alpha\beta} = \sum_{X \in \{L, R\}} \sum_{f \in \{e, u, d\}} \frac{N_f}{N_e} \epsilon_{\alpha\beta}^{fX}. \quad (4.9)$$

Our matter study will be wholly confined to the interior of the Earth, where we assume electrical neutrality and equal distribution of neutrons and protons, we get $N_u/N_e \simeq N_d/N_e \simeq 3$. Also we assume the components $\epsilon_{\alpha\beta}$ to be real. Thus,

$$\epsilon_{\alpha\beta} = \sum_X \epsilon_{\alpha\beta}^{eX} + 3(\epsilon_{\alpha\beta}^{uX} + \epsilon_{\alpha\beta}^{dX}) \quad (4.10)$$

Now, $\epsilon_{\alpha\beta}$ enters the Hamiltonian as entries of a potential-like matrix. In Eq. 4.11, $A_{CC}\text{diag}(1, 0, 0)$ is our familiar matter potential from the Standard Model. There is also our new term, $A_{CC}\epsilon$, which contains the components $\epsilon_{\alpha\beta}$:

$$\begin{aligned}H &= \frac{1}{2E} [UM^2U^\dagger + A_{CC} \text{diag}(1, 0, 0) + A_{CC}\epsilon] \\ &= \frac{1}{2E} \left[UM^2U^\dagger + A_{CC} \begin{pmatrix} 1 + \epsilon_{ee} & \epsilon_{e\mu} & \epsilon_{e\tau} \\ \epsilon_{\mu e} & \epsilon_{\mu\mu} & \epsilon_{\mu\tau} \\ \epsilon_{\tau e} & \epsilon_{\tau\mu} & \epsilon_{\tau\tau} \end{pmatrix} \right]. \quad (4.11)\end{aligned}$$

In the limit $\epsilon_{\alpha\beta} \rightarrow 0$, we recover the standard interaction Hamiltonian from Eq. 2.56. We can draw several conclusions from this form of the Hamiltonian. Any nonzero off-diagonal element $\epsilon_{\alpha\beta}, \alpha \neq \beta$ will make the NSI violate lepton flavor, just as the off-diagonal elements of U does in the SM. Moreover, since the SM potential has the same order in A_{CC} as the NSIs, any $\epsilon_{\alpha\beta} \sim 1$ will make the new matter effect be the same order as the SM effect.

We have two more modifications to the matrix ϵ . First, all terms of the Hamiltonian must of course be Hermitian, thus

$$\begin{pmatrix} \epsilon_{ee} & \epsilon_{e\mu} & \epsilon_{e\tau} \\ \epsilon_{\mu e} & \epsilon_{\mu\mu} & \epsilon_{\mu\tau} \\ \epsilon_{\tau e} & \epsilon_{\tau\mu} & \epsilon_{\tau\tau} \end{pmatrix} = \begin{pmatrix} \epsilon_{ee} & \epsilon_{e\mu} & \epsilon_{e\tau} \\ \epsilon_{e\mu} & \epsilon_{\mu\mu} & \epsilon_{\mu\tau} \\ \epsilon_{e\tau} & \epsilon_{\mu\tau} & \epsilon_{\tau\tau} \end{pmatrix}. \quad (4.12)$$

Now we have reduced the possible number of NSI parameters from 9 down to 6.

By combining oscillation data and neutrino-nucleon scattering at the COHERENT experiment, 90 % CL ranges of some of the NSI parameters were constrained in [25] to

$$\begin{aligned} -0.090 &< \epsilon_{\tau\tau} < 0.38 \\ -0.01 &< \epsilon_{\mu\tau} < 0.009 \\ -0.073 &< \epsilon_{e\mu} < 0.044 \\ -0.15 &< \epsilon_{e\tau} < 0.13. \end{aligned} \quad (4.13)$$

4.2.2 IceCube Signal

In our analysis of IceCube, we are constrained to muon track events. Thus, we are not able to test any theory which does not modify $P_{\alpha\mu}$. Moreover, the IceCube data is available in the range 500 GeV to 10 TeV range, where any rapid oscillations have settled down.

Since all standard matter potentials are diagonal, the elements $\epsilon_{\alpha\beta}$, $\alpha = \beta$ will directly adjust the matter potential felt by flavor α . The off-diagonal terms have a more interesting theoretical implication as they open up matter interactions across flavors. Remember, in the Standard model, we are restricted to weak interactions that conserve lepton number. However, an off-diagonal NSI parameter allows flavor transitions during matter interactions. Thus, the off-diagonal elements constitute new sources of flavor violation. Moreover, $\epsilon_{\alpha\beta}$ modifies the $\nu_\alpha \rightarrow \nu_\beta$ transition independently of the mixing matrix. Thus, we can in theory boost a flavor transition that is suppressed by the mixing matrix. Moreover, since we have 6 NSI parameters, and the PMNS matrix is parametrized with only 3 parameters, we have more degrees of freedom when adjusting the probabilities using the NSI matrix. Each term in the PMNS matrix consists of at least two of the three mixing angles, making them much more dependent on each other.

As discussed in Eq. 2.15, the atmospheric $\nu_\mu \rightarrow \nu_\tau$ transition will be the most abundant, making $\epsilon_{\mu\tau}$, $\epsilon_{\mu\mu}$, $\epsilon_{\tau\tau}$ the most suitable NSI parameters to constrain from muon events. As we will see, $\epsilon_{e\mu}$ is also a candidate, albeit a weaker one.

In Fig. 4.2, we see how the introduction of $\epsilon_{\mu\tau} = 0.02$ alters the ν_μ and $\bar{\nu}_\mu$ survival probabilities for neutrinos that traverse the entire Earth diameter (i.e. $\cos(\theta_z^{true}) = -1$). $\epsilon_{\mu\tau}$ does not dramatically change neither amplitude nor frequency of the probabilities. Instead, it seems to stretch or compress the oscillations. Since the only difference between the way neutrinos and antineutrinos interact with matter is the sign of the potential, the probability for ν_μ with positive $\epsilon_{\alpha\beta}$ is identical to the probability for $\bar{\nu}_\mu$ with negative $\epsilon_{\alpha\beta}$. Thus, the dashed line in the right panel not only shows the survival probability for $\bar{\nu}_\mu$ with $\epsilon_{\mu\tau} = 0.02$, but also the survival probability for ν_μ with $\epsilon_{\mu\tau} = -0.02$. Hence, we note that $\epsilon_{\mu\tau} > 0$ stretches

(compresses) $P_{\mu\mu}$ for neutrinos (antineutrinos), while $\epsilon_{\mu\tau} < 0$ compresses (stretches) $P_{\mu\mu}$ for neutrinos (antineutrinos).

The value of $\epsilon_{\tau\tau}$ affects neither $P_{\mu\mu}$ nor $P_{\bar{\mu}\bar{\mu}}$, in the IceCube region above 500 GeV. Hence, we will not be able to say anything about $\epsilon_{\tau\tau}$ in our IceCube study. Comparing the probabilities in Fig. 4.2 with $\epsilon_{\tau\tau} = 0.05$ with the ones for $\epsilon_{\mu\tau} = 0.02$ in Fig. 4.2, we see that even though we let $\epsilon_{\tau\tau}$ take 2.5 times the value of $\epsilon_{\mu\tau}$, its effect on $P_{\mu\mu}$ is smaller. The weakening of the $P_{\bar{\mu}\bar{\mu}}$ resonance will be visible in DeepCore, but we should expect a less stringent constraint due to the weakness of the effect compared to $\epsilon_{\mu\tau}$.

Thus, we will use IceCube to constrain $\epsilon_{\mu\tau}$ only.

Moving on to $\epsilon_{e\mu}$ and Fig. 4.3, we see that both probabilities has shifted downwards for $E^{true} > 500$ GeV. In Fig. 4.3, we see that the muon channel remains largely unaffected of the value of $\epsilon_{e\tau}$ as we expected. The exception of this lies in the DeepCore region of rapid oscillations, where mixing is more violent.

4.2.3 DeepCore and PINGU Signals

Now we repeat our probability analysis but for the DeepCore/PINGU region of 5.6 GeV to 56 GeV. As we previously saw, we have rapid oscillations, which means that ‘indirect’ modifications (i.e. $\epsilon_{e\tau}$ will affect the $P_{\mu\mu}$ channel) will be more apparent, since all flavors are involved to a greater degree compared with the more stable region above 500 GeV, where many oscillations have settled down.

Another feature of our DeepCore study includes the fact that we now have access to cascade events, in which ν_e and ν_τ are more abundant. Thus, we are no longer constrained to the μ channel alone, but we can now find interesting features in the other channels too. However, we remember that the ν_μ flux is still the most abundant.

Fig. 4.2 shows that $\epsilon_{\mu\tau}$ affects both $P_{\mu\mu}$ and $P_{\bar{\mu}\bar{\mu}}$ over the whole energy range. Since IceCube also sees this, we hope to be able to boost the constraining of $\epsilon_{\mu\tau}$ by combining the two experiments.

Regarding $\epsilon_{\tau\tau}$ in Fig. 4.3, the signal mainly shows in the $P_{\bar{\mu}\bar{\mu}}$ channel as a weaker resonance dip in the 20 GeV region. Thus, DeepCore/PINGU alone will be used to constrain this parameter.

$\epsilon_{e\mu}$ in Fig. 4.3 shows weak resonance dampening for both ν_μ and $\bar{\nu}_\mu$.

For $\epsilon_{e\tau}$, we see a similar resonance dampening in $P_{\mu\mu}$ as we did with $\epsilon_{\tau\tau}$ for $P_{\bar{\mu}\bar{\mu}}$. Hence, we should be able to see the $\epsilon_{e\tau}$ effect in DeepCore/PINGU, but remember that we now have the option to look at the other flavor channels.

Fig. 4.4 shows the electron neutrino and antineutrino survival probabilities, and here we see a clear difference when turning on $\epsilon_{e\tau}$.

We plot the event pull $(N_{NSI} - N_{SI})/\sqrt{N_{SI}}$ where $N_{(N)SI}$ are the numbers of expected events assuming (non-)standard interactions in Fig. 4.5. This gives the

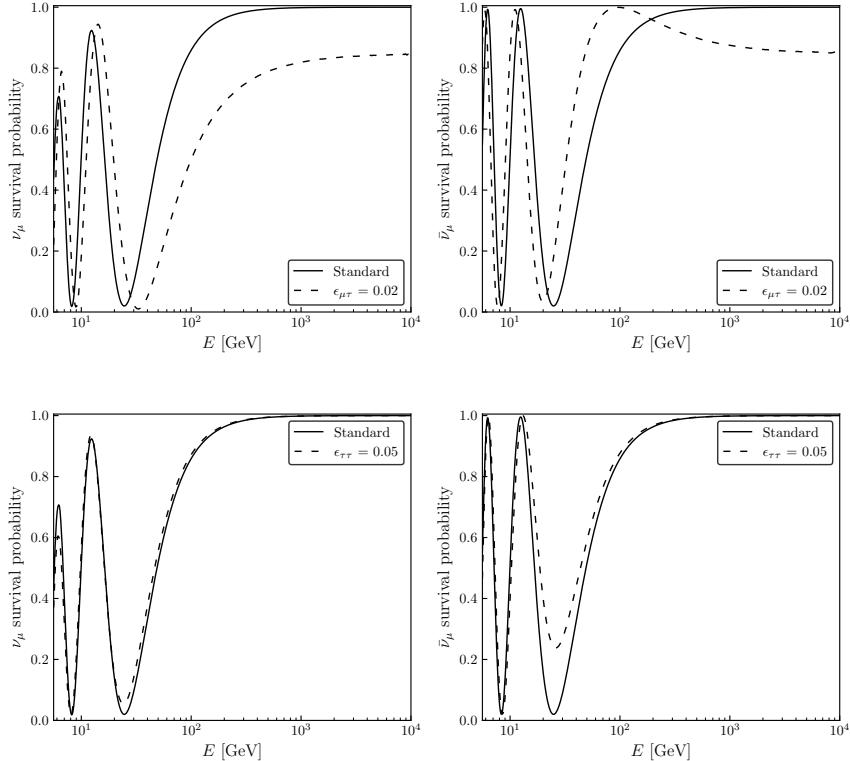


Figure 4.2: *Top panel:* Muon neutrino and antineutrino survival probabilities for $\cos(\theta_z^{true}) = -1$ when $\epsilon_{\mu\tau} = 0.02$. All other NSI parameters are fixed to zero. In the GeV range, $\epsilon_{\mu\tau}$ shifts the oscillations to the right for ν_μ , and to the left for $\bar{\nu}_\mu$. At TeV energies, both probabilities simply get shifted down, resulting in a net reduction of track events. *Bottom panel:* Muon neutrino and antineutrino survival probabilities for $\cos(\theta_z^{true}) = -1$ when $\epsilon_{\tau\tau} = 0.05$. All other NSI parameters are fixed to zero. $\epsilon_{\tau\tau}$ does not affect the probabilities above 100 GeV and this parameter is thus unable to be constrained by tracks in IceCube in our study. However, the dampening of the $\bar{\nu}_\mu$ survival probability will be visible to DeepCore and PINGU, since it occurs within their energy ranges.

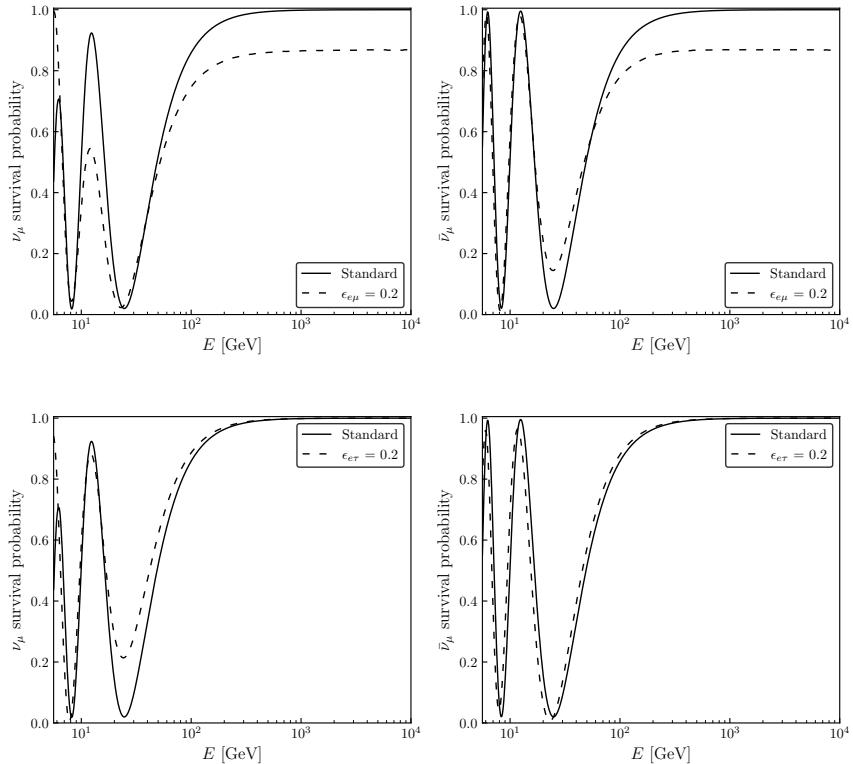


Figure 4.3: Muon neutrino and antineutrino survival probabilities for $\cos(\theta_z^{true}) = -1$. *Top panel:* $\epsilon_{e\mu} = 0.2$. All other NSI parameters are fixed to zero. Instead of the oscillations shifting to the left or right, $\epsilon_{e\mu}$ dampens an oscillation peak for low GeV. At TeV, the probability is shifted down, just as with $\epsilon_{\mu\tau}$. *Bottom panel:* $\epsilon_{e\tau} = 0.2$. All other NSI parameters are fixed to zero. Here, we see a very weak shifting and a ν_μ dampening at low GeV, and again, no visible effect at TeV energies.

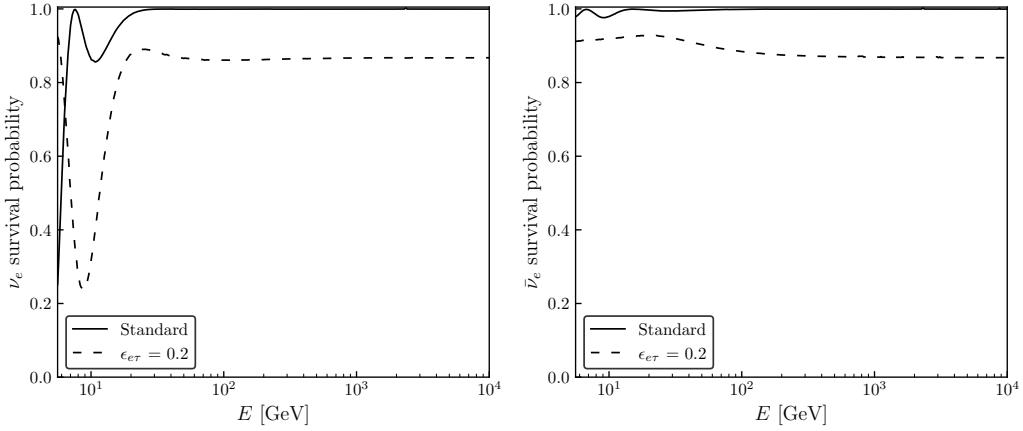


Figure 4.4: Electron neutrino and antineutrino survival probabilities for $\cos(\theta_z^{true}) = -1$ when $\epsilon_{e\tau} = 0.2$. All other NSI parameters are fixed to zero. Compared to the ν_μ and $\bar{\nu}_\mu$ plots, here we see a strong difference across the whole energy range.

normalized difference in the number of expected events at the detector, and illustrates the expected sensitivity of DeepCore for the NSI parameters.

Now we will compare the probability plots with the final event counts in DeepCore. It is not enough to only study the effect at probability level, since that is in true quantities. Since the detector data comes in reconstructed quantities, it might be the case that a feature at probability level gets smeared away in the reconstruction, and not showing up at all in the data. In general, features showing at probability level are often washed out at reconstruction if they occur in the rapidly oscillating part of single-digit GeV energies.

The top right panel shows a surplus of DeepCore track events at 20 GeV to 30 GeV for through-going neutrinos when we turn on $\epsilon_{\tau\tau}$. This is what we expected from the bottom right panel in Fig. 4.2, where we saw that we had an increase in $\bar{\nu}_\mu$ survival probability at 20 GeV to 40 GeV.

The third right panel shows the expected event pull for $\epsilon_{e\mu}$. Comparing with the probabilities in Fig. 4.3, we see that the surplus of the $\bar{\nu}_\mu$ survival in the 20 GeV to 30 GeV and the deficit above 40 GeV is intact, while the 10 GeV ν_μ deficit is completely washed out.

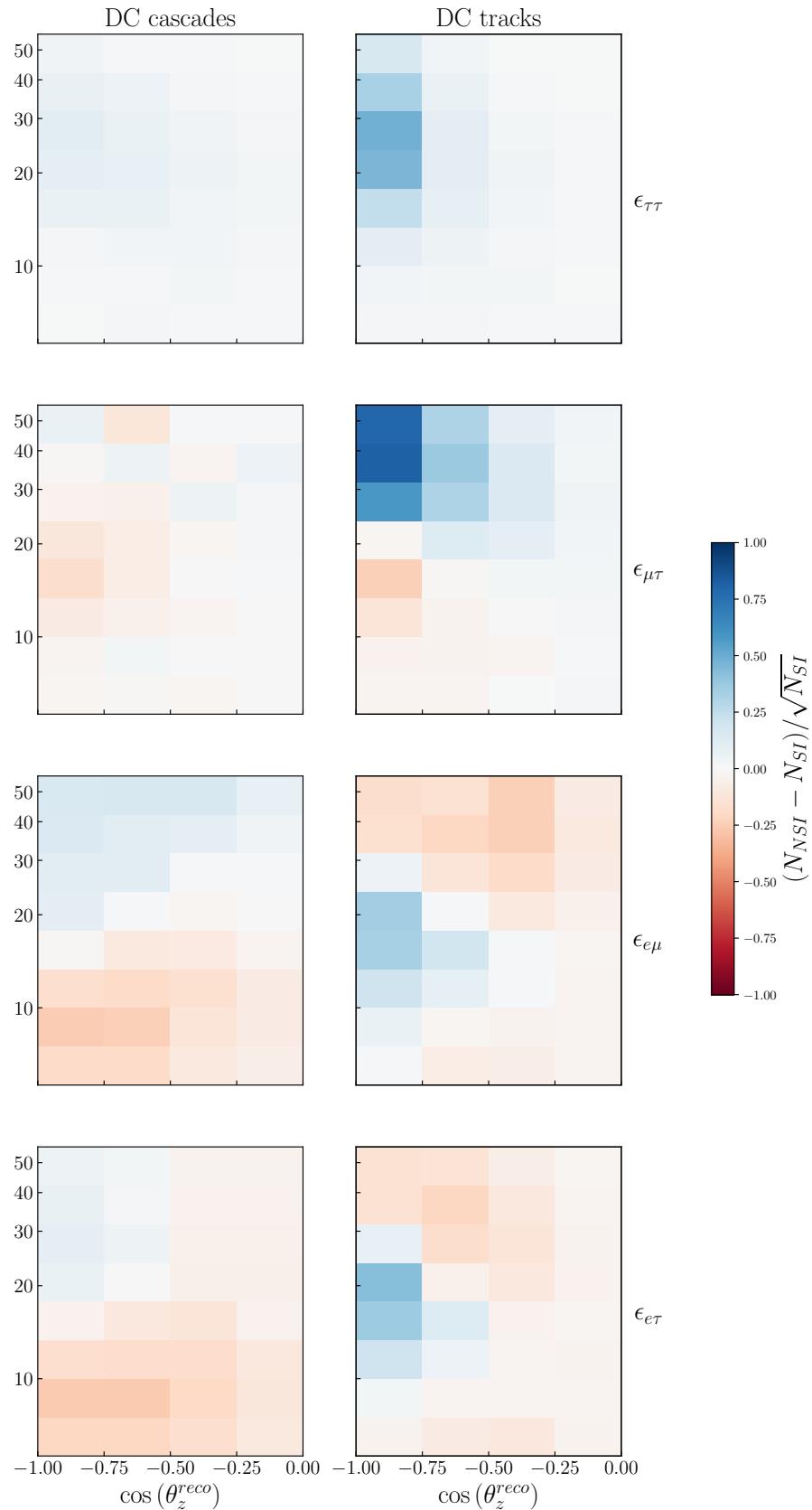
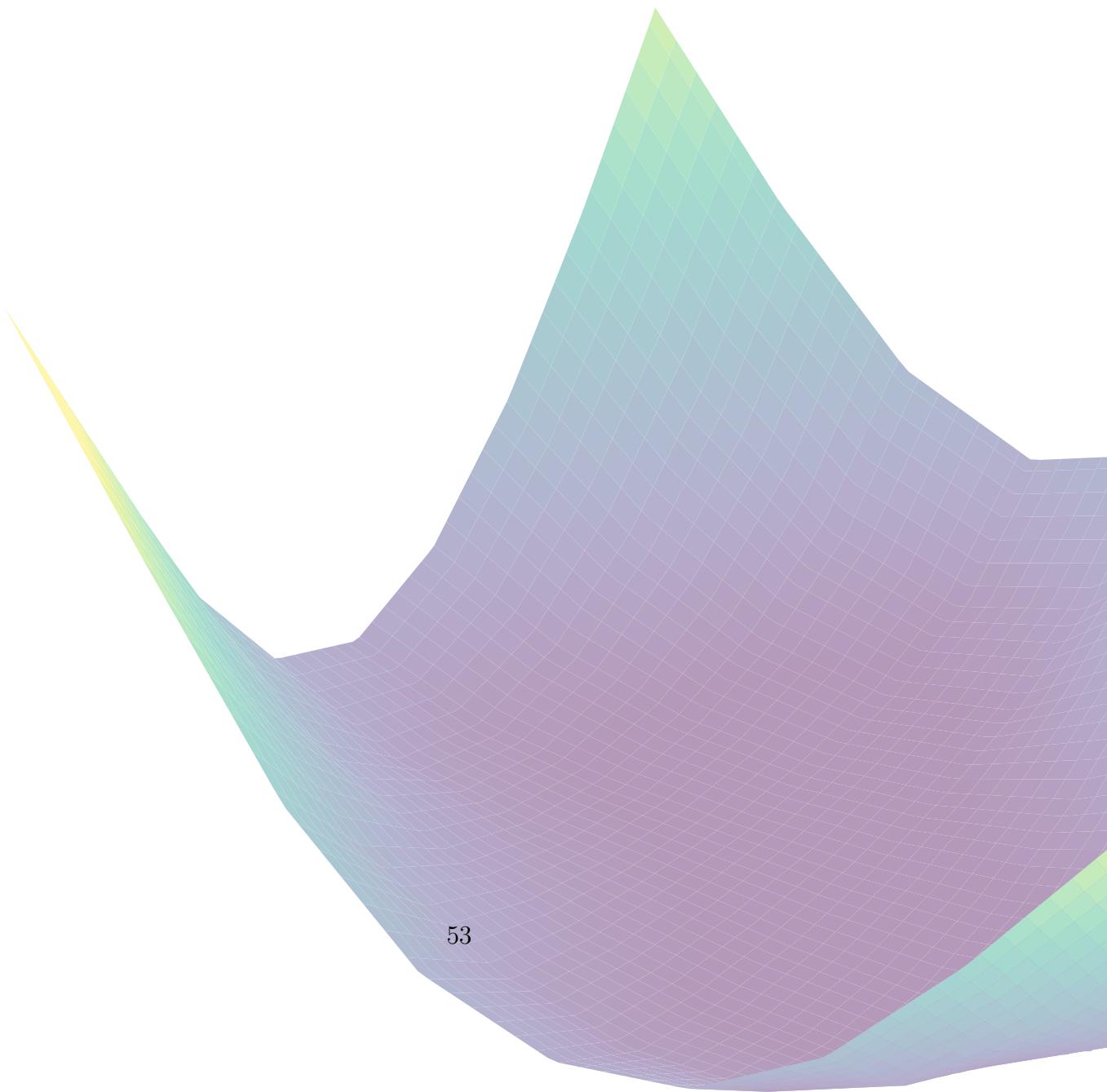


Figure 4.5: Plots of the quantity $(N_{NSI} - N_{SI}) / \sqrt{N_{SI}}$, displaying which reconstructed bins the NSI signal is most clear.

Chapter 5

Results



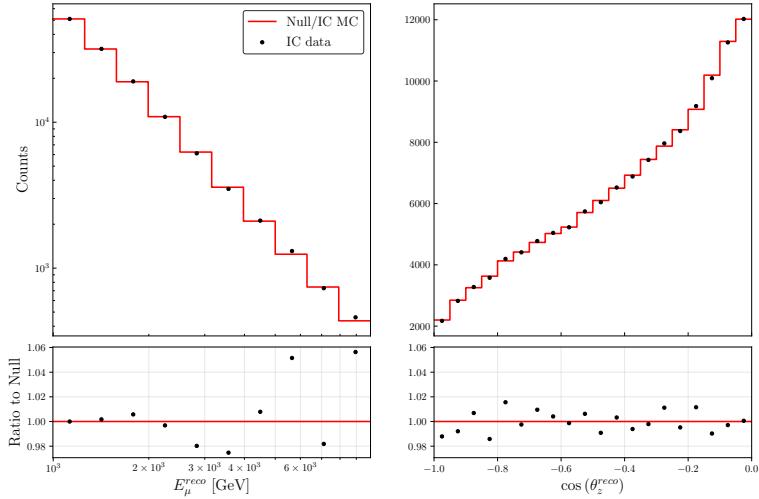


Figure 5.1: How the null hypothesis compares to the IceCube data from [12].

5.1 The Sterile Hypothesis

With the null (3ν) hypothesis normalized to the IceCube Monte Carlo as in Eq. 3.6, we are now in good shape to study the sterile effect on the probabilities, and how that compares to data. First, let us see how the collected data from Fig. 3.6 deviates from the predicted 3ν oscillations. Fig. 5.1 shows an impressive agreement to the standard 3ν oscillation picture, with the largest deviations of 6% for neutrinos with high single-digit TeV values. In zenith, the data only deviates $\pm 2\%$. Recalling the TeV disappearance from Fig. 4.1a, we see that we have a similar deficiency found in the data at 4 TeV. Thus, we expect the best fitting sterile hypothesis to include a Δm_{41}^2 that places the resonant disappearance in the same region. From Fig. 4.1b, we see that a mass of $\Delta m_{41}^2 = 1.5 \text{ eV}^2$ achieves this.

5.1.1 χ^2 Minimization

For our analyses, we define our χ^2 as

$$\chi^2(\hat{\theta}, \alpha, \beta) = \sum_{ij} \frac{(N_{ij}^{\text{th}} - N_{ij}^{\text{data}})^2}{(\sigma_{ij}^{\text{data}})^2 + (\sigma_{ij}^{\text{syst}})^2} + \frac{(1 - \alpha)^2}{\sigma_\alpha^2} + \frac{\beta^2}{\sigma_\beta^2} \quad (5.1)$$

where we minimize over the model parameters $\hat{\theta}$, the penalty terms α and β . N_{ij}^{th} is the expected number of events from theory, and N_{ij}^{data} is the observed number of events in that bin. We set $\sigma_\alpha = 0.25$ as the atmospheric flux normalization error, and $\sigma_\beta = 0.05$ as the zenith angle slope error [7]. The observed event number has

an associated Poissonian uncertainty $\sigma_{ij}^{\text{data}} = \sqrt{N_{ij}^{\text{data}}}$. For IceCube, the event count takes the form

$$N_{ij}^{\text{th}} = \alpha [1 + \beta(0.5 + \cos(\theta_z^{\text{reco}})_i)] N_{ij}(\hat{\theta}), \quad (5.2)$$

with $N_{ij}(\hat{\theta})$ from Eq. 3.6. Here, the term $\beta(0.5 + \cos(\theta_z^{\text{reco}})_i)$ allows the event distribution to rotate with angle β around the median zenith angle of $\cos(\theta_z^{\text{reco}}) = -0.5$. We also have an uncorrelated systematic error $\sigma_{ijk}^{\text{syst}}$. We set $\sigma_{ijk}^{\text{syst}} = f \sqrt{N_{ijk}^{\text{data}}}$, where f is a factor of our own choosing.

The minimization of Eq. 5.1 simply returns a value for each set of model parameters. We use this value to quantify to what extent our theoretical simulations N_{ij}^{th} agree with the data N_{ij}^{data} within the error bounds provided by σ_a, σ_b , and f . We then select the set of χ^2 values which are the smallest and our allowed parameters are then their associated $\hat{\theta}$. We then translate the χ^2 distribution by subtracting the best-fit point, and analyze $\Delta\chi^2 = \chi^2 - \chi^2_{\min}$. From the χ^2 -distribution, one can show that the 90% confidence level has a value of 2.71 for two degrees of freedom. When slicing through the two-dimensional grid of $\Delta\chi^2$ at this level, we then obtain a contour plot that tells us what parameter values are within our 90% confidence level and what regions are not.

5.1.2 Sterile Mass and Mixing

Let us start with looking at the best-fit event distribution resulting from the χ^2 minimization. Fig. 5.4 now contains the best-fit event distribution assuming the sterile hypothesis. The best-fit values are $\Delta m_{41}^2 = 0.01 \text{ eV}^2$ and $\sin(2\theta_{24})^2 = 0.67$ ($\theta_{24} = 27.5^\circ$). The contour plot shown in Fig. 5.3 divides the parameter space into two regions. To the right of the boundary, the $\Delta\chi^2$ has values above the confidence level, meaning that those parameter pairs can be excluded at a certain confidence level.

By tuning the scaling f of the uncorrelated systematic error, we can shift the contour. We extract the contour from the IceCube sterile analysis [12], and compare in Fig. 5.2. We see that the contour in the higher Δm_{41}^2 region above 1 eV^2 is very sensitive to the tuning, while the lower region between $1 \times 10^{-1} \text{ eV}^2$ to 1 eV^2 remains fixed. We put $f = 35\%$ and compare our results to the Giunti et al. [?] in Fig. 5.3. Here, we see that our lower fixed region between $1 \times 10^{-1} \text{ eV}^2$ to 1 eV^2 is closer to the IceCube result than Giunti et al, but that our contour slightly diverges above 1 eV^2 .

We now in Fig. 5.4 plot the best-fit sterile hypothesis with the data from Fig. 5.1, to finally see how the best sterile hypothesis compares to data.

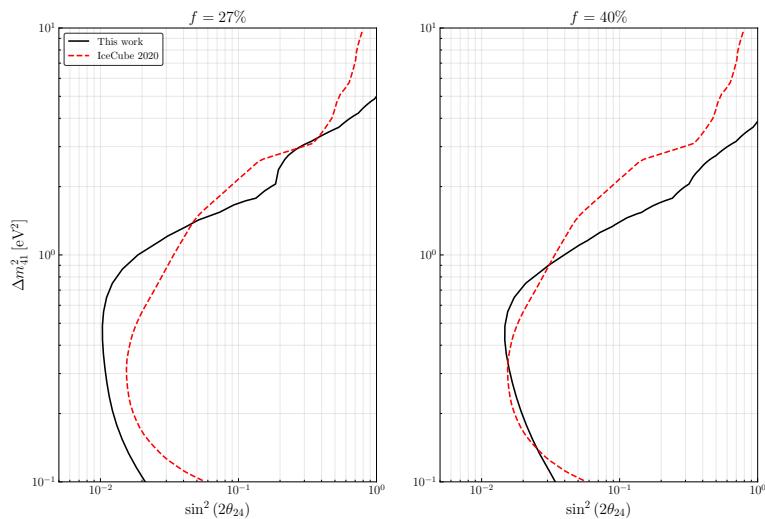


Figure 5.2: The 90% CL contour with two different systematic error scalings, compared with the official IceCube result from [12]

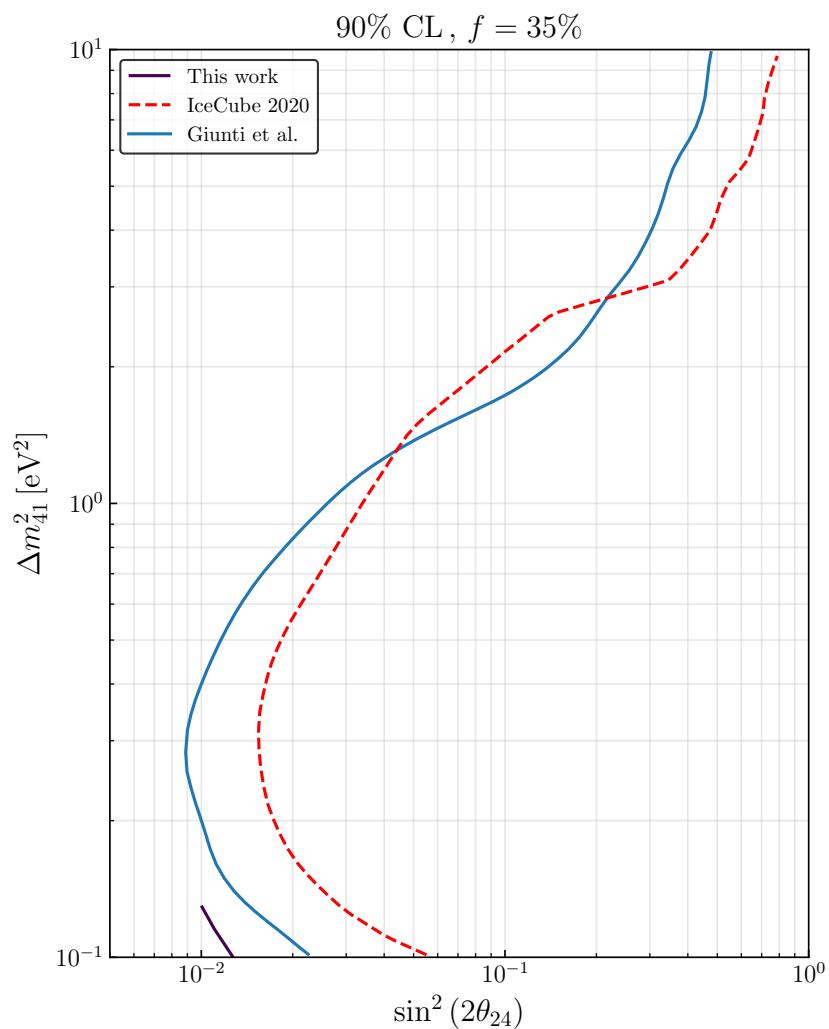


Figure 5.3: The 90% CL contour using $f = 35\%$, compared with the official IceCube result from [12] and the result from [?]

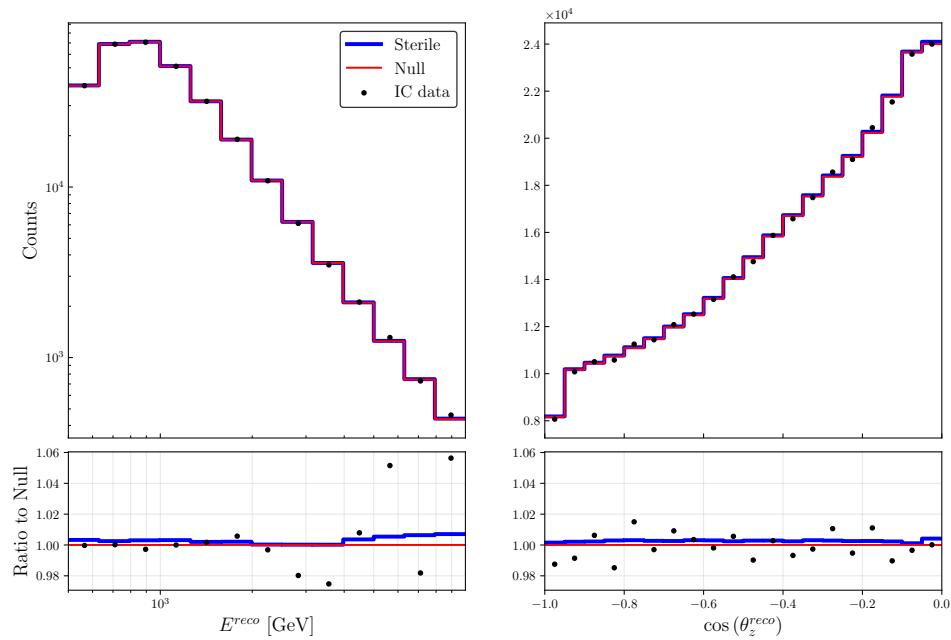


Figure 5.4: The predicted binned event count assuming our best-fit sterile hypothesis, compared with the null and data.

5.2 Non-Standard Interactions

5.2.1 Constraining the NSI parameters

In this section, we will constrain the four NSI parameters $\epsilon_{\tau\tau}$, $\epsilon_{\mu\tau}$, $\epsilon_{e\mu}$, and $\epsilon_{e\tau}$ by considering the detectors separately as well as jointly. For our analyses, we define our χ^2 as

$$\chi^2(\hat{\theta}, \alpha, \beta, \kappa) = \sum_{ijk} \frac{(N_{ijk}^{\text{th}} - N_{ijk}^{\text{data}})^2}{(\sigma_{ijk}^{\text{data}})^2 + (\sigma_{ijk}^{\text{syst}})^2} + \frac{(1-\alpha)^2}{\sigma_\alpha^2} + \frac{\beta^2}{\sigma_\beta^2} \quad (5.3)$$

where we minimize over the model parameters $\hat{\theta} \in \{\Delta m^2, \theta_{23}, \epsilon\}$, the penalty terms α and β , and the free parameter κ . N_{ijk}^{th} is the expected number of events from theory, and N_{ijk}^{data} is the observed number of events in that bin.

In our simulations of N_{ijk}^{th} , we set all standard oscillation parameters to their current best-fit values of Eq. 3.8, except for Δm^2 and θ_{23} which we vary over their 3σ limits $2.435 \times 10^{-3} \text{ eV}^2$ to $2.598 \times 10^{-3} \text{ eV}^2$ and 40.1° to 51.7° , respectively.

We set $\sigma_\alpha = 0.25$ as the atmospheric flux normalization error, and $\sigma_\beta = 0.05$ as the zenith angle slope error [7]. The observed event number has an associated Poissonian uncertainty $\sigma_{ijk}^{\text{data}} = \sqrt{N_{ijk}^{\text{data}}}$. For IceCube, the event count takes the form

$$N_{ij}^{\text{th}} = \alpha [1 + \beta(0.5 + \cos(\theta_z^{\text{reco}})_i)] N_{ij}(\hat{\theta}), \quad (5.4)$$

with $N_{ij}(\hat{\theta})$ from Eq. 3.4. Here, we allow the event distribution to rotate around the median cosine-zenith of $\cos(\theta_z^{\text{reco}}) = -0.5$.

For DeepCore and PINGU, the event count takes the form

$$N_{ijk}^{\text{th}} = \alpha [1 + \beta \cos(\theta_z^{\text{reco}})_i] N_{ijk}(\hat{\theta}) + \kappa N_{ijk}^{\mu_{\text{atm}}}, \quad (5.5)$$

with $N_{ijk}(\hat{\theta})$ from Eq. 3.7. $N_{ijk}^{\mu_{\text{atm}}}$ is the muon background, which is left to float freely in the DeepCore analysis. The background at PINGU can be considered negligible to first order [16], and we thus put $\kappa = 0$ when calculating the PINGU χ^2 . For DeepCore and PINGU, the median cosine-zenith is $\cos(\theta_z^{\text{reco}}) = 0$, and we allow the event count to rotate around this point.

We treat the uncorrelated systematic uncertainties differently for each detector. For IceCube, we set $\sigma_{ijk}^{\text{syst}} = f \sqrt{N_{ijk}^{\text{data}}}$. We consider best, normal, and worst-case scenarios in IceCube using $f = 5\%$, 10% , and 15% respectively. For PINGU, we use the same form but instead use $f = 0\%$, 3% , and 5% . For DeepCore, we use the provided systematic error distribution which accounts for uncertainties in the finite MC statistics and the data-driven muon background estimate [13]. This is summarized in Table 5.1.

Experiment	Best case	Baseline	Worst case
IceCube	5%	10%	15%
PINGU	0%	3%	5%

Table 5.1: Our definition of the best, baseline, and worst case scenarios considered in each experiment, modelled by $\sigma_{ijk}^{\text{syst}} = f \sqrt{N_{ijk}^{\text{data}}}$ with f from the table. We do not consider different DeepCore scenarios because her systematic error distribution is already provided in the data release [13].

For the joint analysis, we follow the parameter goodness-of-fit prescription [26] and construct the joint χ^2 as

$$\chi_{\text{joint}}^2 = \sum_{\text{exp}} \chi_{\text{exp}}^2 - \chi_{\text{exp,min}}^2 \quad (5.6)$$

with test statistic $\chi_{\text{joint,min}}^2$. The $\Delta\chi_{\text{joint}}^2$ is then $\Delta\chi_{\text{joint}}^2 = \chi_{\text{joint}}^2 - \chi_{\text{joint,min}}^2$.

After the oscillation parameters have been marginalized out, we plot $\Delta\chi^2$ for each of the four NSI parameters in Fig. 5.5. The results are shown in Fig. 5.6 and summarized in Tables 5.2 and 5.3.

Comparing the PINGU and the DeepCore results in Fig. 5.5, we note that the best-fit for each NSI parameter for the PINGU experiment is expected to be zero. This is because the ‘data’ we generated during the PINGU simulations assume no NSI since they have yet to be observed in nature. This introduces a non-NSI bias in all joint analyses which include PINGU since PINGU has stronger statistics than DeepCore and will thus pull the joint χ^2 towards $\epsilon = 0$.

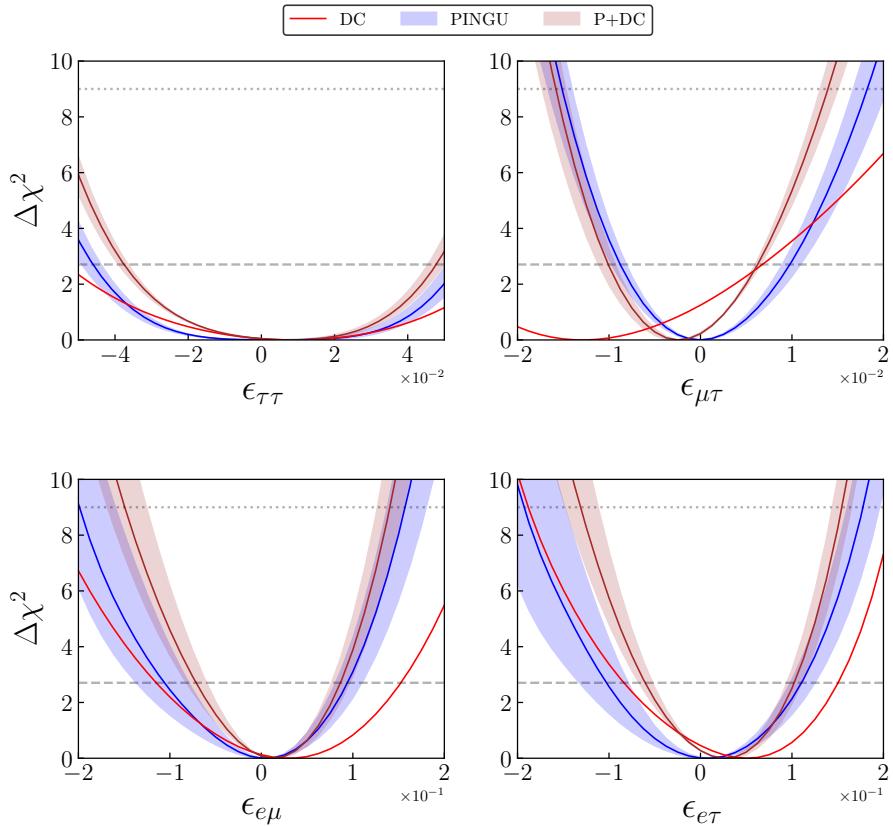


Figure 5.5: Confidence regions for PINGU and DeepCore scenarios listed in Table 5.1, with their joint $\Delta\chi^2$ in maroon. Δm_{31}^2 and θ_{23} have been marginalized out, and all other NSI parameters not shown in each panel are fixed to zero. IceCube tracks only reveal $\epsilon_{\mu\tau}$, and are displayed separately in Fig. 5.6. Dotted lines are the 90% and 3 σ CL levels.

Parameter	Best 90% CL	Best 3σ
IceCube		
$\epsilon_{\mu\tau}$	[-0.008, 0.009]	[-0.014, 0.017]
DeepCore		
$\epsilon_{\tau\tau}$	[-0.044, 0.051]	[-0.062, 0.069]
$\epsilon_{\mu\tau}$	[-0.008, 0.009]	[-0.014, 0.017]
$\epsilon_{e\mu}$	[-0.079, 0.081]	[-0.16, 0.14]
$\epsilon_{e\tau}$	[-0.079, 0.098]	[-0.15, 0.16]
IceCube + DeepCore		
$\epsilon_{\mu\tau}$	[-0.029, 0.007]	[0.026]

Parameter	Baseline 90% CL	Baseline 3σ
IceCube		
$\epsilon_{\mu\tau}$	[-0.009, 0.01]	[-0.015, 0.018]
DeepCore		
$\epsilon_{\tau\tau}$	[-0.046, 0.054]	[-0.065]
$\epsilon_{\mu\tau}$	[-0.009, 0.01]	[-0.015, 0.018]
$\epsilon_{e\mu}$	[-0.11, 0.094]	[-0.20, 0.16]
$\epsilon_{e\tau}$	[-0.10, 0.11]	[-0.19, 0.18]
IceCube + DeepCore		
$\epsilon_{\mu\tau}$	[-0.029, 0.007]	[0.026]

Parameter	Worst 90% CL	Worst 3σ
IceCube		
$\epsilon_{\mu\tau}$	[-0.01, 0.011]	[-0.017, 0.02]
DeepCore		
$\epsilon_{\tau\tau}$	[-0.049, 0.057]	[-0.07]
$\epsilon_{\mu\tau}$	[-0.01, 0.011]	[-0.017, 0.02]
$\epsilon_{e\mu}$	[-0.14, 0.11]	[-0.23, 0.18]
$\epsilon_{e\tau}$	[-0.13, 0.13]	[-0.23, 0.12]
IceCube + DeepCore		
$\epsilon_{\mu\tau}$	[-0.029, 0.007]	[0.026]

Table 5.2: IceCube and DeepCore results from the $\Delta\chi^2$ in Fig. 5.6. Δm_{31}^2 and θ_{23} have been marginalized out, and all other NSI parameters other than the one shown for each row are set to zero. Best, baseline, and worst refer to the systematic uncertainty scenarios considered as in Table 5.1.

Parameter	Best 90% CL	Best 3σ
PINGU		
$\epsilon_{\tau\tau}$	[-0.054, 0.067]	[]
$\epsilon_{\mu\tau}$	[-0.029, 0.007]	[0.026]
$\epsilon_{e\mu}$	[-0.12, 0.15]	[-0.23, 0.24]
$\epsilon_{e\tau}$	[-0.084, 0.15]	[-0.19, 0.21]
DeepCore + PINGU		
$\epsilon_{\tau\tau}$	[-0.036, 0.046]	[-0.056, 0.064]
$\epsilon_{\mu\tau}$	[-0.009, 0.006]	[-0.015, 0.013]
$\epsilon_{e\mu}$	[-0.06, 0.077]	[-0.126, 0.127]
$\epsilon_{e\tau}$	[-0.052, 0.095]	[-0.112, 0.144]
IceCube + DeepCore + PINGU		
$\epsilon_{\mu\tau}$	[-0.009, 0.006]	[-0.015, 0.013]
Parameter	Baseline 90% CL	Baseline 3σ
PINGU		
$\epsilon_{\tau\tau}$	[-0.054, 0.067]	[]
$\epsilon_{\mu\tau}$	[-0.029, 0.007]	[0.026]
$\epsilon_{e\mu}$	[-0.12, 0.15]	[-0.23, 0.24]
$\epsilon_{e\tau}$	[-0.084, 0.15]	[-0.19, 0.21]
DeepCore + PINGU		
$\epsilon_{\tau\tau}$	[-0.038, 0.048]	[-0.058, 0.067]
$\epsilon_{\mu\tau}$	[-0.01, 0.006]	[-0.016, 0.014]
$\epsilon_{e\mu}$	[-0.071, 0.086]	[-0.149, 0.141]
$\epsilon_{e\tau}$	[-0.061, 0.103]	[-0.131, 0.155]
IceCube + DeepCore + PINGU		
$\epsilon_{\mu\tau}$	[-0.01, 0.006]	[-0.016, 0.014]
Parameter	Worst 90% CL	Worst 3σ
PINGU		
$\epsilon_{\tau\tau}$	[-0.054, 0.067]	[]
$\epsilon_{\mu\tau}$	[-0.029, 0.007]	[0.026]
$\epsilon_{e\mu}$	[-0.12, 0.15]	[-0.23, 0.24]
$\epsilon_{e\tau}$	[-0.084, 0.15]	[-0.19, 0.21]
DeepCore + PINGU		
$\epsilon_{\tau\tau}$	[-0.039, 0.05]	[-0.06]
$\epsilon_{\mu\tau}$	[-0.011, 0.007]	[-0.017, 0.015]
$\epsilon_{e\mu}$	[-0.082, 0.097]	[-0.17, 0.16]
$\epsilon_{e\tau}$	[-0.067, 0.11]	[-0.15, 0.17]
IceCube + DeepCore + PINGU		
$\epsilon_{\mu\tau}$	[-0.011, 0.007]	[-0.017, 0.015]

Table 5.3: PINGU and joint results from the $\Delta\chi^2$ in Fig. 5.5. Δm_{31}^2 and θ [23] have been marginalized out, and all other NSI parameters other than the one shown for each row are set to zero. Best, baseline, and worst refer to the systematic uncertainties given in Table 5.1.

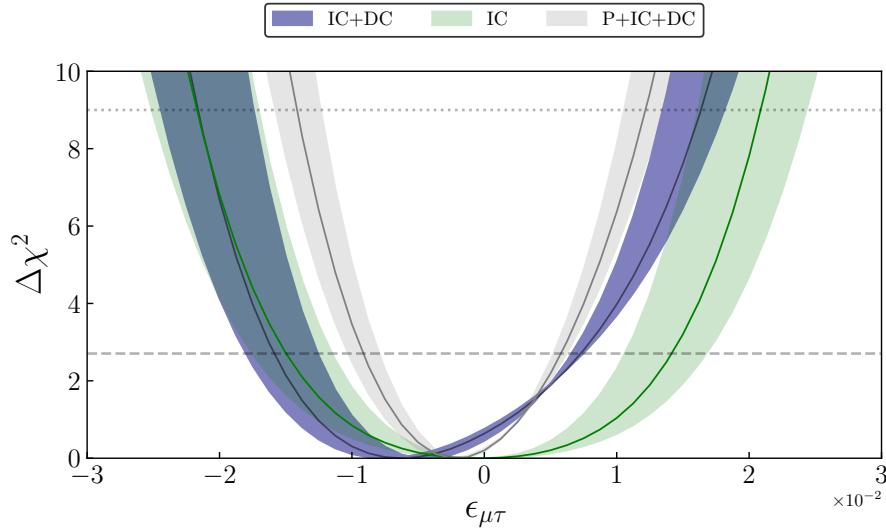


Figure 5.6: IceCube $\epsilon_{\mu\tau}$ $\Delta\chi^2$ regions for scenarios as defined in Table 5.1. Δm_{31}^2 and θ_{23} and have been marginalized out, and all other NSI parameters other than $\epsilon_{\mu\tau}$ are fixed to zero. Dotted lines are the 90% and 3 σ CL levels.

Parameter	Best fit		
	Δm^2	θ_{23}	ϵ
DeepCore			
$\epsilon_{\tau\tau}$	2.435	47.84	0.0125
$\epsilon_{\mu\tau}$	2.435	43.97	-0.005
$\epsilon_{e\mu}$	2.435	43.97	0
$\epsilon_{e\tau}$	2.435	43.97	0.05
IceCube			
$\epsilon_{\mu\tau}$	2.435	51.70	0
IceCube + DeepCore			
$\epsilon_{\mu\tau}$	2.517	43.97	-0.01

Table 5.4: Best fit points for Δm_{31}^2 and θ_{23} are given in units of 10^{-3}eV^2 and degrees, respectively.

There is no reason for us to assume that only one NSI parameter exists in Nature, unless we impose a symmetry on the gauge group which generates NSI. For example, the matter potential matrix in our Hamiltonian from Eq. (2.42) is diagonal since the interaction Lagrangian respects lepton number conservation. The PMNS matrix, on the other hand, has off-diagonal elements because the kinetic Lagrangian breaks this symmetry. So ideally, we would simulate a grid where all NSI parameters are allowed to vary, but this is not feasible. Thus, we take we set $\epsilon_{\tau\tau} = 0$ and let $\epsilon_{\mu\tau}, \epsilon_{e\mu}, \epsilon_{e\tau}$ vary, along with Δm_{31}^2 and θ_{23} . We let $\epsilon_{\tau\tau} = 0$ because we saw that it did not influence the other parameters. We then marginalize out the standard oscillation parameters and one of the NSI parameters and plot the remaining two in Fig. 5.8. We see that the pairs $\epsilon_{e\mu} - \epsilon_{\mu\tau}$ and $\epsilon_{e\tau} - \epsilon_{\mu\tau}$ are symmetrical. Hence, we can see that no relationship exists between them. With $\epsilon_{e\tau} - \epsilon_{e\mu}$, however, the contours are assuming a different shape. The contour allows positive values of $\epsilon_{e\tau}$ and $\epsilon_{e\mu}$ to a greater degree than mixed or negative values.

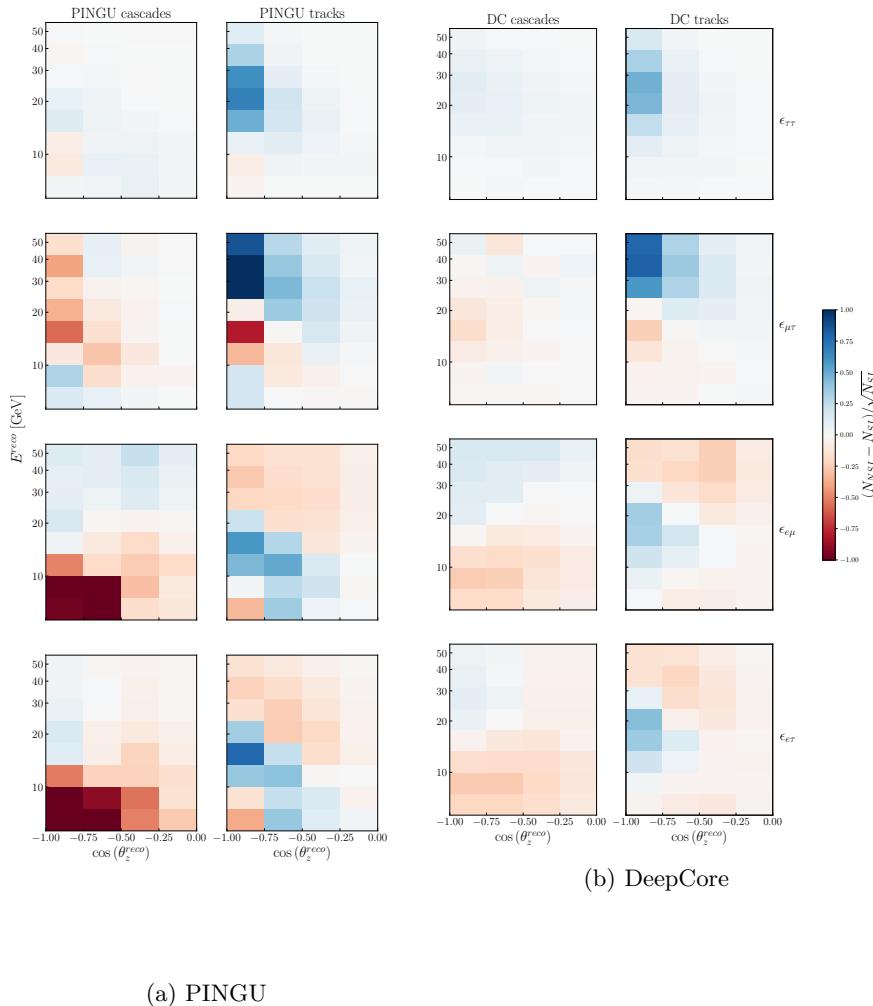


Figure 5.7: Expected pulls of the form $(N_{NSI} - N_{SI})/\sqrt{N_{SI}}$ for PINGU and DeepCore after 3 years.

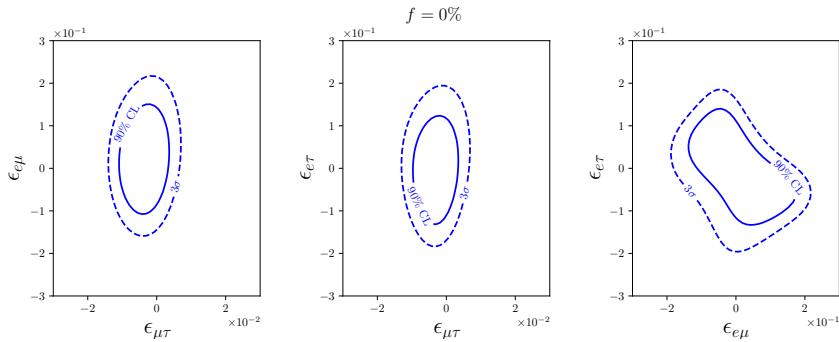


Figure 5.8: Allowed regions for some of the NSI parameters after three years of PINGU data, assuming no uncorrelated systematic error. All plots are made with $\epsilon_{\tau\tau} = 0$. The two standard oscillation parameters Δm_{31}^2 and θ_{23} along with the one NSI parameter not shown have been marginalized out.

Bibliography

- [1] B. Pontecorvo, Mesonium and antimesonium Vol: 33.
URL <https://www.osti.gov/biblio/4343073>
- [2] Z. Maki et al., Remarks on the Unified Model of Elementary Particles 28 (5) 870–880. doi:[10.1143/PTP.28.870](https://doi.org/10.1143/PTP.28.870).
- [3] I. Esteban et al., The fate of hints: Updated global analysis of three-flavor neutrino oscillations 2020 (9) 178. doi:[10.1007/JHEP09\(2020\)178](https://doi.org/10.1007/JHEP09(2020)178).
- [4] A. M. Dziewonski and D. L. Anderson, Preliminary reference Earth model 25 (4) 297–356. doi:[10.1016/0031-9201\(81\)90046-7](https://doi.org/10.1016/0031-9201(81)90046-7).
- [5] C. Weaver, Evidence for Astrophysical Muon Neutrinos from the Northern Sky.
- [6] The IceCube Gen2 Collaboration, Letter of Intent: The Precision IceCube Next Generation Upgrade (PINGU). arXiv:[1401.2046](https://arxiv.org/abs/1401.2046).
- [7] M. Honda et al., Calculation of atmospheric neutrino flux using the interaction model calibrated with atmospheric muon data.doi:[10.1103/PhysRevD.75.043006](https://doi.org/10.1103/PhysRevD.75.043006).
- [8] M. Honda, HKKM2014 flux table one year average.
URL <http://www.icrr.u-tokyo.ac.jp/~mhonda/nflux2014/index.html>
- [9] M. Honda et al., Atmospheric neutrino flux calculation using the NRLMSISE-00 atmospheric model 92 (2) 023004. doi:[10.1103/PhysRevD.92.023004](https://doi.org/10.1103/PhysRevD.92.023004).
- [10] The IceCube Collaboration, All-sky point-source IceCube data: Years 2010–2012. doi:[10.21234/B4F04V](https://doi.org/10.21234/B4F04V).
- [11] The IceCube Collaboration, Search for sterile neutrinos with one year of IceCube data.
URL <https://icecube.wisc.edu/data-releases/2016/06/search-for-sterile-neutrinos-with-one-year-of-icecube-data/>
- [12] The IceCube Collaboration, Searching for eV-scale sterile neutrinos with eight years of atmospheric neutrinos at the IceCube Neutrino Telescope 102 (5) 052009. doi:[10.1103/PhysRevD.102.052009](https://doi.org/10.1103/PhysRevD.102.052009).

- [13] The IceCube Collaboration, Three-year high-statistics neutrino oscillation samples. [doi:10.21234/ac23-ra43](https://doi.org/10.21234/ac23-ra43).
- [14] The IceCube Collaboration, Measurement of Atmospheric Neutrino Oscillations at 6–56 GeV with IceCube DeepCore 120 (7) 071801. [doi:10.1103/PhysRevLett.120.071801](https://doi.org/10.1103/PhysRevLett.120.071801).
- [15] The IceCube Collaboration, Measurement of atmospheric tau neutrino appearance with IceCube DeepCore 99 (3) 032007. [doi:10.1103/PhysRevD.99.032007](https://doi.org/10.1103/PhysRevD.99.032007).
- [16] The IceCube Collaboration, IceCube Upgrade Neutrino Monte Carlo Simulation. [doi:10.21234/qfz1-yh02](https://doi.org/10.21234/qfz1-yh02).
- [17] LSND Collaboration et al., Evidence for $\{\nu_\mu \rightarrow \nu_e\}$ Oscillations from the LSND Experiment at the Los Alamos Meson Physics Facility 77 (15) 3082–3085. [doi:10.1103/PhysRevLett.77.3082](https://doi.org/10.1103/PhysRevLett.77.3082).
- [18] Particle Data Group et al., Review of Particle Physics 2020. [doi:10.1093/ptep/ptaa104](https://doi.org/10.1093/ptep/ptaa104).
- [19] P. Collaboration, Planck 2018 results. VI. Cosmological parameters 641 A6. [arXiv:1807.06209](https://arxiv.org/abs/1807.06209), [doi:10.1051/0004-6361/201833910](https://doi.org/10.1051/0004-6361/201833910).
- [20] P. F. de Salas and S. Pastor, Relic neutrino decoupling with flavour oscillations revisited 2016 (07) 051. [doi:10.1088/1475-7516/2016/07/051](https://doi.org/10.1088/1475-7516/2016/07/051).
- [21] M. B. Gavela et al., Minimal Flavour Seesaw Models 2009 (09) 038–038. [arXiv:0906.1461](https://arxiv.org/abs/0906.1461), [doi:10.1088/1126-6708/2009/09/038](https://doi.org/10.1088/1126-6708/2009/09/038).
- [22] J. Salvado et al., Non-standard interactions with high-energy atmospheric neutrinos at IceCube 2017 (1) 141. [arXiv:1609.03450](https://arxiv.org/abs/1609.03450), [doi:10.1007/JHEP01\(2017\)141](https://doi.org/10.1007/JHEP01(2017)141).
- [23] Y. Farzan and M. Tortola, Neutrino oscillations and Non-Standard Interactions. [arXiv:1710.09360](https://arxiv.org/abs/1710.09360).
- [24] T. Ohlsson, Status of non-standard neutrino interactions 76 (4) 044201. [arXiv:1209.2710](https://arxiv.org/abs/1209.2710), [doi:10.1088/0034-4885/76/4/044201](https://doi.org/10.1088/0034-4885/76/4/044201).
- [25] P. Coloma et al., A COHERENT enlightenment of the neutrino Dark Side 96 (11) 115007. [arXiv:1708.02899](https://arxiv.org/abs/1708.02899), [doi:10.1103/PhysRevD.96.115007](https://doi.org/10.1103/PhysRevD.96.115007).
- [26] M. Maltoni and T. Schwetz, Testing the statistical compatibility of independent data sets 68 (3) 033020. [arXiv:hep-ph/0304176](https://arxiv.org/abs/hep-ph/0304176), [doi:10.1103/PhysRevD.68.033020](https://doi.org/10.1103/PhysRevD.68.033020).