

# Upgrades to the Fluorescence Detectors of the Pierre Auger Observatory



THE UNIVERSITY  
*of*ADELAIDE

**Tristan William Sudholz**

School of Physical Sciences  
University of Adelaide

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*Doctor of Philosophy*

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Tristan William Sudholz



# **Abstract**



# Acknowledgements



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# Nomenclature

Auger	Pierre Auger Observatory
EAS	Extensive Air Shower
NSB	Night Sky Background
PE	Photo-electron
FD	Fluorescence Detector
SD	Surface Detector
PMT	Photomultiplier Tube
FLT	First Level Trigger
eV	electron volt
ADC	Analogue-to-Digital Unit
S/N ratio	Signal to Noise Ratio



# Introduction

- Define Cosmics Rays.
- The origins of the highest energy cosmic-rays still unknown.
- First detection by Pierre Auger in 1937 and the current detector looking at these energies is the Pierre Auger Observatory.
- Hybrid experiment containing both surface detectors and fluorescence detectors
- Surface detector has nearly 100% up-time while the fluorescence detectors only have 15% up-time.
- \*\*\*\* Proposal to extend the fluorescence detector up-time. To achieve this will have to operate while the moon is above the horizon. This will increase the level NSB and will have the PMTs run under a reduced gain to compensate. \*\*\*\*
- Photomultiplier Tubes are used as pixels within the camera of the fluorescence detectors and the aim of these thesis is to quantify the characteristics of the PMT under the reduced gain and increased.
- Outline a Summary of each chapter.

Cosmic-rays are particles that originate outside of the Earth atmosphere. These particles can be photons, hadronic or leptonic in nature [ref?]. In this thesis, when mentioning cosmic-rays I will mean the hadronic component unless specified otherwise. Cosmic-rays have been measured over a large range of energies (over 6 decades in energy) and it has many interesting features have been observed in this energy spectrum. One of the longest running mysteries is what happens at the highest energy. Since the first detection of extensive air showers by Pierre Auger in 1937 [ref], many different experiments have endeavoured to solve this mystery. The Pierre Auger Observatory [ref] is currently in operation to observe cosmic-rays at the highest energies.

The Pierre Auger Observatory is a hybrid experiment consisting of both surface detectors and fluorescence detectors. (Outline location) The surface detector has a nearly 100% operation up-time ref while the fluorescence detectors only 15% operation up-time [ref]. (Outline how Auger detects cosmic-rays, just need a brief summary).

A current proposal to extend the fluorescence detector operation up-time. Extended up-time would be beneficial as the fluorescence detectors image the entire extensive air shower and would increase the number of showers observed through out yearly observation. To achieve the extended operation the fluorescence detectors would have to be operated while the moon is above the horizon. While the moon is up, this would increase the Night Sky Background level and to compensate the Photomultiplier Tubes acting as the camera pixels would have to be run under reduced gain.

- Need graph of expected variance in ADC<sup>2</sup> for the moon above the horizon for different phases.
- Want to increase the duty cycle of FD by measuring EAS under moonlight. Most likely observe under quarter to half moon. This will increased the NSB upto a factor of 10.
  - The aim of increasing the duty cycle of FD is too measure more EAS at the highest energy band ( $> 10^{19.5}$  eV).
  - Need more statistics at highest energy band to complement SD measurements.

The aim of this thesis is to quantify the characteristics of the Photomultiplier Tubes operating under this reduced gain and outline any operation strategies. Outline of each chapter is as follows:

- Chapter 1: Cosmic-rays
  - Does this work as a new line
- Chapter 2: Detection of Cosmic-Rays
  - Add text here
- Chapter 3: The Pierre Auger Observatory
  - Add text here
- Chapter 4 : EAS Selection Efficiency with Increased NSB
  - Add text here
- Chapter 5 : Quantifying Characteristics of the FD PMT
  - Add text here
- Chapter 6 : Computer Simulation of the FD PMT
  - Add text here
- Chapter 7 : Measuring Gain Variance of the FD PMT with CalA data
  - Add text here
- Chapter 8 : Laboratory Simulation of FD Shifts
  - Add text here
- Chapter 9 : Effectiveness of Cloud Camera Cuts
  - Add text here
- Chapter 10: Conclusion
  - Future Work

# Chapter 1

## Cosmic-Rays

### 1.1 History of Cosmic-Rays

First detection of ionizing radiation.

1785: Coulomb found that electroscopes can spontaneously discharge by the action of the air and not by defective insulation

1835: Faraday confirms the observation by Coulomb, with better insulation technology

1879: Crookes measures that the speed of discharge of an electroscope decreased when pressure was reduced

### 1.2 Energy Spectrum and Mass composition

Cosmic-rays have been detected over a large range of energies from GeV ( $10^9$  eV) to above EeV ( $10^{18}$  eV). Spectrum in Figure 1.1 shows the break at the knee and ankle and which type of experiments are most suited to measurement each part. Cosmic-ray spectrum starts out at  $E^{-2}$  and can be as steep as  $E^{-2.7}$  at the highest energies.

Cosmic-rays can consist of protons to iron.

CR spectrum has many features. Main features are the knee, second knee and ankle. The knee is around  $3 \times 10^{15}$ .

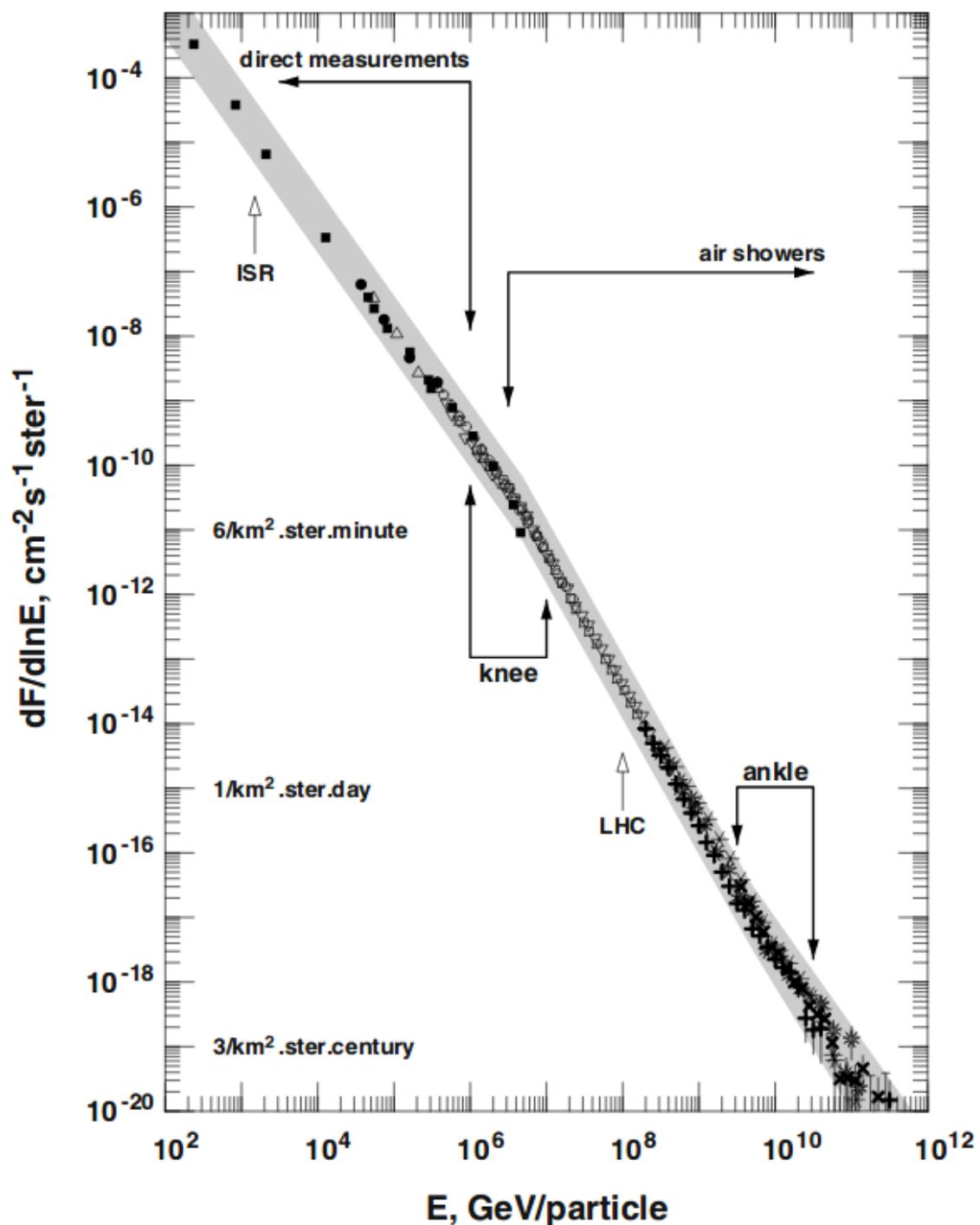
Pierre Auger Observatory measurement of isotropy that shows that the cosmic-ray spectrum changes from predominately galactic to extra-galactic at the ankle.

Predicted Greisen-Zatsepin-Kuz'min (GZK) cut-off about  $6 \times 10^{19}$ . Cosmic-rays above this energy are theorised to interact with the cosmic microwave background radiation. Greisen independently of Kuz'man and Zatsepin all predicted this energy loss.

Pierre Auger Observatory measurement of  $X_{\text{max}}$  and the second moment  $\sigma(X_{\text{max}})$  has mass composition information as well how this changes as a function of energy.

### 1.3 Production Method and Sources

- Bottom-Up Acceleration
  - Supernova explosions
  - AGN jets
  - other energetic processes



**Figure 1.1:** Measured energy spectrum of cosmic-rays from 100 GeV up to the highest detected energy.

dark matter annihilations.

- Top-Down Acceleration

Decay of massive relic particles

Typically associated with new physics beyond the standard model



# Chapter 2

## Detections of Cosmic-Rays

### 2.1 Extensive Air Showers

Use Earth's atmosphere as an interaction medium. Primary particle interacts with the molecules in the atmosphere to produce a cascade of secondary particles. This cascade of particles is referred to as an Extensive Air Shower (EAS). Hadronic primaries can produce pions, muons and other stuff. Mixture of a hadronic core with an electromagnetic component from the decay of  $\pi^0$ .

Shower profile has particles produced until energy of individual secondary particles drop below the ionization threshold. Therefore the shower will reach a point of maximum particle number then will drop off.

### 2.2 Fluorescence Production

The charge particles of EAS interact with the nitrogen molecules in the atmosphere. This interaction turns the nitrogen molecule dipole like and when the nitrogen returns to a ground state, a photon is emitted. This emitted photon is termed fluorescence light. Fluorescence light is emitted isotropically and typically in the UV band (between 300 and 400 nm). \*\*\* Show wavelength profile \*\*\*

### 2.3 Atmospheric Effects

### 2.4 Detectors and History

Early Experiments:

Volcano Ranch

Haverah Park

SUGAR

Yakutsk array is located in Russia and has been operating in different forms since 1967. The array reached a maximum collecting area of  $17 \text{ km}^2$  around 1990. Recently it has been reconfigured to have a collection area of  $8 \text{ km}^2$  to study lower energy cosmic-rays.

Akeno Gaint Air Shower Array (AGASA) is located in Tokyo, Japan. Operating at an average altitude of 667 m above sea level from 1990 to 2004. The array consist of



**Figure 2.1:** Diagram of Cosmic-ray Extensive Air Showers.

over one hundred scintillator detectors covering  $100 \text{ km}^2$  \*\*\*check this\*\*\*. The timing measurements and data collection is achieved via interconnected optical fibers.

The Fly's Eye was the first successful air fluorescence detector operating from 1981 to 1993 at the Dugway Proving Grounds in Utah, USA. Fly's Eye achieved a time averaged aperture of about  $100 \text{ km}^2 \text{sr}$  at the highest energies, considering it only operated on clear moonless nights.

HiRes improved on the Fly's Eye design by advancing resolution and sensitivity, This was achieved by increasing the telescope effective mirror area to  $3.8 \text{ m}^2$  and reducing the camera pixel angular diameter to  $1^\circ$ .

# Chapter 3

## Pierre Auger Observatory

Science Goals of the Pierre Auger Observatory is to probe the origins and characteristics of cosmic rays above  $10^{17}$  eV and to study the interactions of the most energetic particles observed in nature.

The Pierre Auger Observatory (PAO) is an hybrid detector that is located near Malargüe in the Mendoza Province, Argentina. PAO consists of 1660 Cherenkov water detector spread over 3000 km<sup>2</sup> by 24 fluorescence telescopes.

### 3.1 Surface Detector

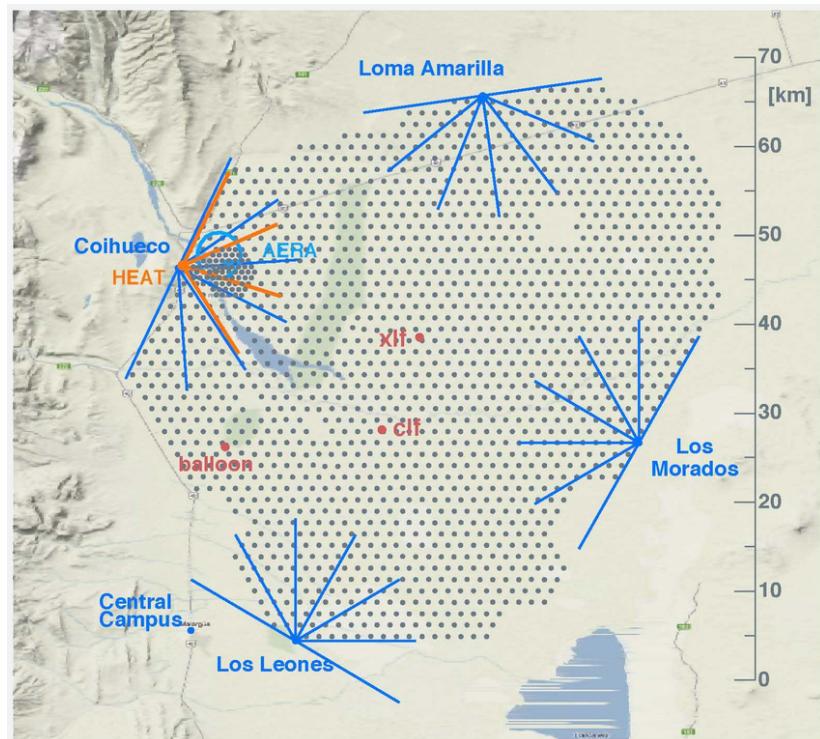
The surface array consists of 1660 water Cherenkov tanks. The majority of tanks are configured with a spacing of 1500 metres while there is a small subset of tanks in front of the fluorescence telescopes at the Coihueco site with spacing of 750 metres.

The surface array has a duty cycle of nearly 100% and the maintenance cycle is so that no more then 20 tanks are down at any one time.

#### 3.1.1 AugerPrime

### 3.2 Fluorescence Detector

There are four fluorescence detector site surrounding the surface array. At each fluorescence detector site there are six telescopes covering 180° in azimuth and 30° in elevation. At one site there are three extra telescopes with slightly greater then 90° in azimuth and cover 30° to 60° in elevation.



**Figure 3.1:** Image of layout of Pierre Auger Observatory located near Malargüe, Argentina.



**Figure 3.2:** Image of one of the fluorescence detector site (background) and one of the surface detectors (foreground).



Figure 3.3: Basic schematic of a surface detector.



Figure 3.4: Basic schematic of a fluorescence telescope.

### **3.2.1 Photomultiplier Tubes**

## **3.3 Communication System and CDAS**

## **3.4 Event Reconstruction**

### **3.4.1 Surface Detector**

### **3.4.2 Fluorescence Detector**

## **3.5 Enhancements and future upgrades**

## Chapter 4

# The Effects of Increased Night Sky Background on the Fluorescence Detector's Trigger and Reconstruction

### 4.1 Motivation

The Pierre Auger Observatory has been operational since 2004 and the collaboration is in the process of rolling out upgrades to improve the experiment. The upgrade has been named AugerPrime ([ref](#)) and is a large project to improve both the Surface Detectors (SD) and the Fluorescence Detectors (FD). The FD portion of AugerPrime involves extending the duty cycle to observe Extensive Air Shower (EAS) events while the moon is above the horizon. The main purpose for extending the FD duty cycle to include observation times while the moon is visible is to increase the statistics per year for EAS events in the highest energy bin ( $> 10^{19.5}$  eV).

In this chapter I used real data events and simulations to explore the effects of increasing the Night Sky Background (NSB) on the collaboration's reconstruction method and trigger efficiency. A real data set was used to evaluate the efficiency of the reconstruction method under different NSB intensities by artificially adding extra noise to signal traces. This method was a repetition of a study done by M. Unger [Find ref.](#). The study was used as a starting point then the results were examined in depth and found to mainly test the robustness of the collaboration's reconstruction algorithm. The analysis led to the use of the collaboration FD simulations to evaluate the effects of the increased NSB on both the FD triggers and reconstruction efficiency.

### 4.2 Background Information

#### 4.2.1 Current Fluorescence Detector Operation

Currently the FDs are operated under these guidelines: an observation run is organised for nights when the illuminated fraction of the moon is less than 70% and can have a minimum of 3 hours of operation with the moon below the horizon. The FD telescope shutters are then opened when the sun is below  $-18^\circ$  of the horizon (beginning of astronomical twilight). The FD shutters remain open while the average variance across

the FD camera is less than 100 ADC<sup>2</sup>/100 ns and the variance of individual pixels (which are photomultiplier tubes) is less than 2000 ADC<sup>2</sup>/100 ns. 100 ns is chosen as this is the smallest interval that the signal from the photomultiplier tube (PMT) anode (in electrons) is digitised to Analogue-to-Digital Units (ADCs). The FD shutters at each site will automatically close if the individual rain sensors are triggered or the measured wind speed is above 50 km/h at the individual FD sites. Using these guidelines the expected yearly FD duty cycle can be calculated. This calculation was performed by the collaboration in **DATE & ref** to estimate the theoretical the FD duty cycle and the result is shown in Table 4.1. Bad weather is mainly caused by rain, high wind speeds or lightning strikes in the FDs field of view. Failures include any software or hardware issues that prevented any of the FD telescopes from collecting data.

Theoretical up time	22%
Loss due to short nights (< 3 hrs)	-2%
Loss due to bad weather or failures	-5%
Total measurement time	15%

**Table 4.1:** Calculation of the yearly expected FD duty cycle done by the Collaboration. Bad weather is the inclusion of rain, high wind speeds and lightning strikes in the FDs field of view. Failures include any software or hardware issues that prevented any of the FD telescopes from collecting data.

Photons that enter the FD telescope pass through the optics, are converted to photo-electrons by PMT and then convert to Analogue-to-Digital units by the electronics. A basic diagram of how photons detected by a FD telescope is converted to ADC is show in Fig. 4.1. The steps are split into 3 parts: the optics, PMT and electronics. When the photon reaches the PMT it is converted to a photo-electron and the electron signal from the PMT is converted to ADC by the electronics.

The electron signal from the FD PMT is AC coupled, which means the average ADC of the NSB is zero. Instead the ADC variance is calculated and the ADC variance directly proportional to the variance in the NSB photon count. To find the variance of the ADC variance due to the NSB from the background trace the variance due to the electronic noise has to be removed first. Removal of the electronic noise is shown in Eq. 4.1.

$$[\sigma_{\text{ADC}}^2]^{\text{sky}} = [\sigma_{\text{ADC}}^2]^{\text{back}} - [\sigma_{\text{ADC}}^2]^{\text{elec}} \quad (4.1)$$

where:

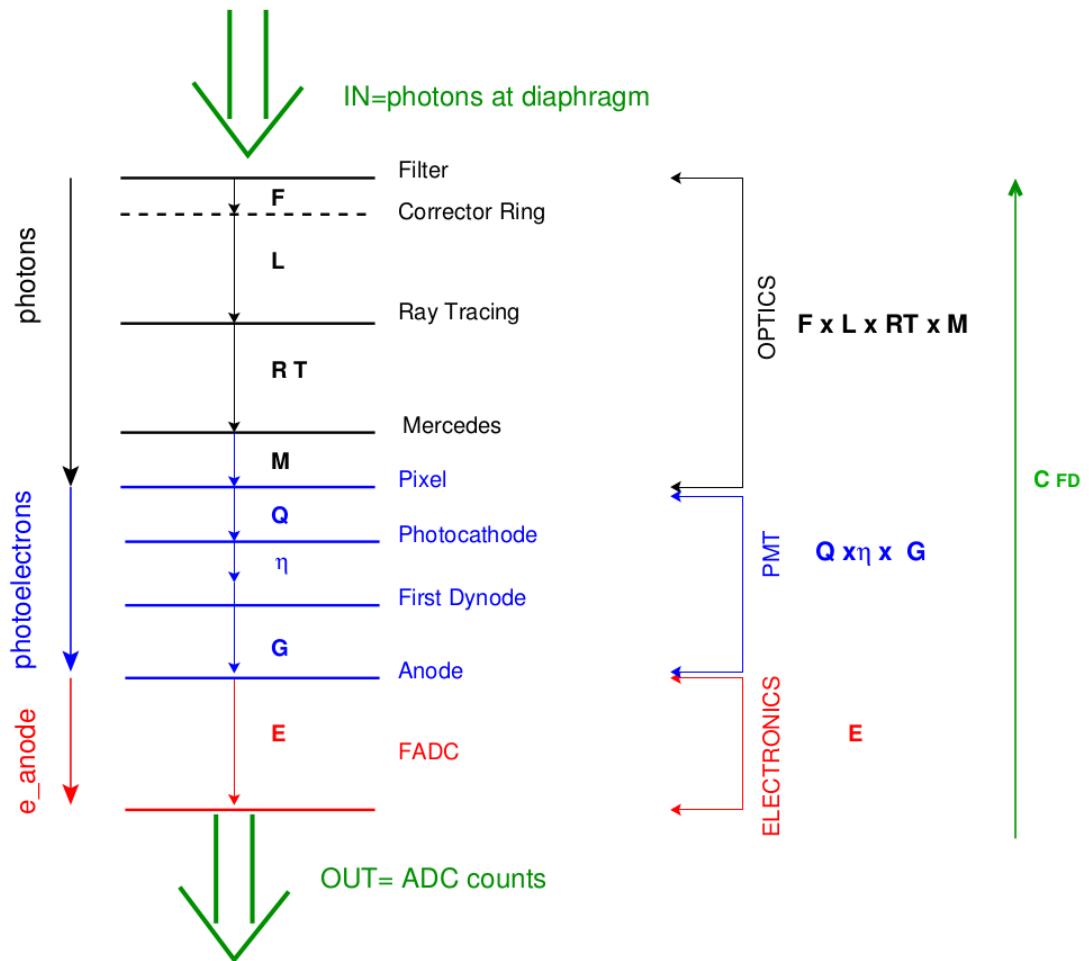
$[\sigma_{\text{ADC}}^2]^{\text{sky}}$  is the variance (ADC<sup>2</sup>/100 ns) caused by the NSB.

$[\sigma_{\text{ADC}}^2]^{\text{back}}$  is the variance (ADC<sup>2</sup>/100 ns) of background signal which is the combination of NSB and electronic noise.

$[\sigma_{\text{ADC}}^2]^{\text{elec}}$  is the variance (ADC<sup>2</sup>/100 ns) of the electronic noise. The variance of FD electronic noise is measured while the shutters are closed.

The typical mean ADC variance of the NSB measured by Auger ?? at Malargue is:

$$[\sigma_{\text{ADC}}^2]^{\text{sky}} = 25 \text{ ADC}^2 / 100 \text{ ns}$$



**Figure 4.1:** A diagram showing the steps that a photon passing through a FD telescope diaphragm will travel to be converted to an ADC. The steps are split into 3 parts: the optics, PMT and electronics. When the photon reaches the PMT it is converted to a photo-electron and the electron signal from the PMT anode is converted to ADC by the electronics.

The variance in  $\text{ADC}^2$  seen by the FD can be converted back to photons at the diaphragm by the FD pixel calibration constant  $C_{\text{FD}}$  (photons/ADC). The FD pixel calibration constant has been preformed by illuminating the diaphragm with a known number of photons. This was originally done by a drum calibration which used a large device that covered the entire FD diaphragm and could illuminate the whole camera with known number of diffuse photons. The drum calibration has been retired and has been replaced with an experiment called the XY scanner. The XY scanner illuminates one pixel at the time and scans over the entire camera. The advantages of the XY scanner is that it is a more compact experiment with the aim of having the ability to be deployed more often to find the absolute calibration.

$C_{\text{FD}}$  can be used to find the variance of the NSB photons at the diaphragm from the measured ADC variance. An average variance of the NSB measured by an FD telescope under standard operational guidelines is  $\sigma_{\text{ADC}}^2 = 25 \text{ ADC}^2 / 100 \text{ ns}$ . The calculations to work out the variance of the photon count ( $\text{RMS}_{\text{ph}} / 100 \text{ ns}$ ) from the measured variance in  $\text{ADC}^2$  is as follows:

$$\sigma_{\text{ph}}^2 = (C_{\text{FD}})^2 \times \sigma_{\text{ADC}}^2 \quad (4.2)$$

Using Eq. 4.2 and  $C_{\text{FD}} = 4.5 \text{ photons / ADC}$  the fluctuations on the photon count can be calculated.

$$\sigma_{\text{ph}} = 23 \text{ photons / } 100 \text{ ns}$$

where  $\sigma_{\text{ph}}$  is the fluctuations on the photon count per 100 ns.

Finding the number of photons at the diaphragm takes more steps. The first step involves converting the ADC variance to photo-electrons. This step is achieved through the combination of using Eq. 4.3 and Eq. 4.4.

$$n_{\text{pe}} = \frac{\sigma_{\text{pe}}^2}{(1 + V_G)} \quad (4.3)$$

$$\sigma_{\text{pe}}^2 = [\sigma_{\text{ADC}}^2]^{\text{sky}} / A_G^2 \quad (4.4)$$

where  $\sigma_{\text{pe}}$  is the standard deviation of the photo-electron count,  $n_{\text{pe}}$  is the photon-electron count,  $V_G$  is the PMT gain variance factor ?? and  $A_G$  is the absolute gain (ADC/photo-electron) of the PMT.  $A_G$  is equal to:

$$A_G = \eta \times G \times E \quad (4.5)$$

where

$\eta$  is the PMT collection efficiency at the first dynode.

$G$  is the PMT gain.

$E$  is the electronic conversion constant (ADC/electron).

An equivalent quantity can be used to calculate  $A_G$  and is:

$$A_G = \frac{1}{C_{\text{FD}} \times f \times Q} \quad (4.6)$$

where

$Q$  is the Quantum efficiency of the PMT turning photons into photo-electrons.

$C_{FD}$	4.5 photons/ADC
$Q$	0.29
$f$	0.494

**Table 4.2:** Typical values for the constants used to calculate  $A_G$  which is the absolute gain (ADC/photo-electron).  $C_{FD}$  is the FD pixel calibration constant,  $Q$  is the Quantum efficiency of the PMT turning photons into photo-electrons and  $f$  is the photon collection efficiency of the telescope optics.

Optics	$f$	$F \times L \times T \times R \times M$
Filter Transmission	$F$	0.83
Corrector Ring lens Transmission	$L$	0.90
Mirror Reflectivity	$R$	0.90
Camera Shadow Factor	$T$	0.79
Mercedes Collection Efficiency	$M$	0.93

**Table 4.3:** Breakdown of all the components that go into the FD optics.

$f$  is the photon collection efficiency of the telescope optics. The components that make up the telescope optics are shown in Table. 4.3.

Typical measured values for  $C_{FD}$ ,  $Q$  and  $f$  shown in Table 4.2. The PMTs used as camera pixels in the FDs are HAMAMATSU XP3062 and are operated at a gain of  $\sim 5 \times 10^4$  electrons/photo-electron.

The next step is to convert the photo-electrons to photons at the diaphragm. This step is achieved through using Eq. 4.7.

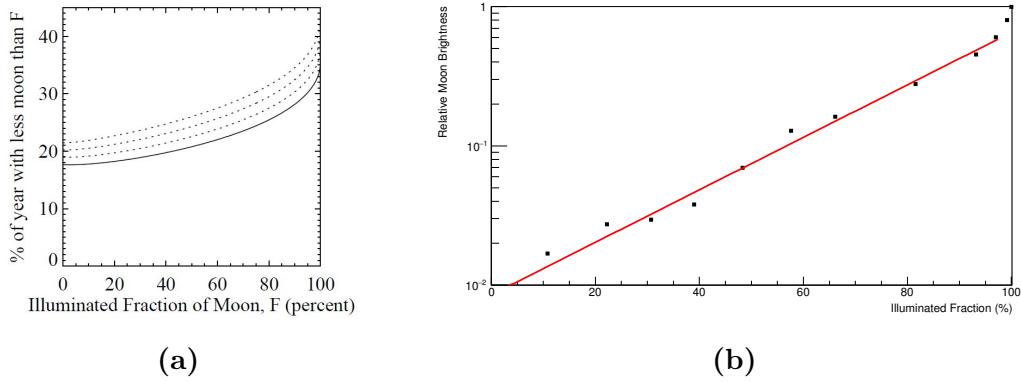
$$n_{ph} = \frac{n_{pe}}{Q \times f} \quad (4.7)$$

Therefore  $A_G$  can be calculated from Eq. 4.5 using the values from Table 4.2. From all of the equations stated above I have outlined a table showing the expected photon count at the aperture per 100 ns from the measured variance ( $ADC^2 / 100$  ns).

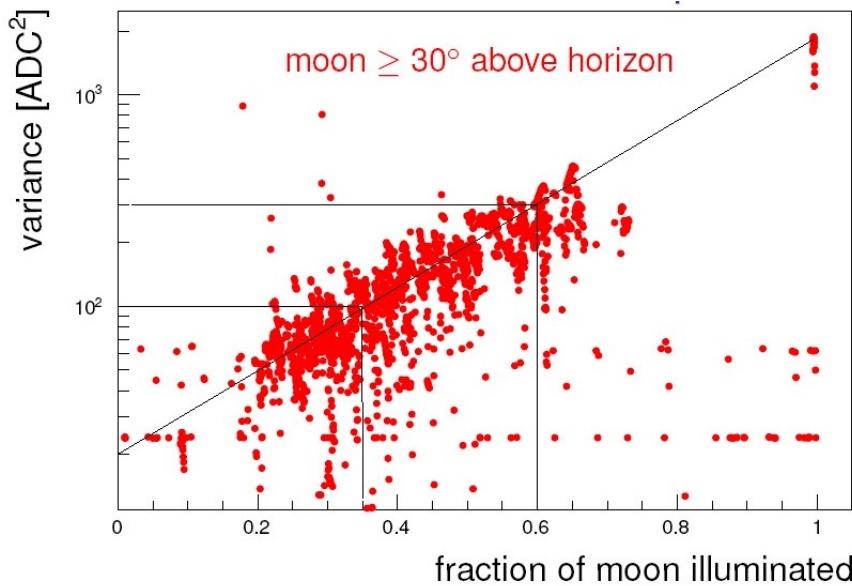
Variance ( $ADC^2 / 100$ ns)	Photons / 100 ns @ Pixel	Photons / 100 ns @ diaphragm
25	23	46
250	227	459
2500	2267	4588

#### 4.2.2 Night Sky Background and the Moon

The Night Sky Background for Auger typically refers to airglow, zodiacal light, scattered starlight and any man-made light. Moonlight is a separate source of background light as an FD shift is typically operated when the moon is below the horizon. Table 4.4 shows the expected change in average background light observed by the FD telescopes for no moon, 1st/3rd quarter moon and full moon / twilight. 1st/3rd quarter moon is either one quarter or three quarters through the moon orbital cycle and is typically 50% of the observable moon from earth is illuminated. From the table it can be seen that



**Figure 4.2:** (a) Integral Percent of nights across the year with the moon with different illuminated fractions. Only times between astronomical dusk and dawn have been included. The solid line is the astronomical twilight (sun at least  $18^\circ$  below horizon). The dotted lines represent relaxed definition of twilight with the sun being at  $15^\circ$ ,  $12^\circ$  and  $9^\circ$  below the horizon. (b) Relative brightness of the moon in the night sky compared with a full moon as a function of illuminated fraction. Images taken from [ref\(GAP1996-034\)](#).



**Figure 4.3:** Measured NSB variance from the Fluorescence Detectors at Standard PMT HV setting with different moon illuminated fractions. The moon was 30° above horizon or greater. Image taken from [ref\(Radamir\)](#).

going from no moon to quarter moon to full moon jumps by a factor of ten in the observed ADC variance. The large change in moon brightness for the different phases is due to that the light is spectral reflection not diffusion reflection.

Condition	$\sigma^2$ [ADC <sup>2</sup> /100 ns]	I <sub>a</sub> [ $\mu\text{A}$ ]
no moon	25	0.5
quarter moon	250	5
full moon/twilight	2500	50

**Table 4.4:** Expected average observed variance in ADC<sup>2</sup> and anode current in  $\mu\text{A}$  by the PMTs in the FD telescopes under different NSB conditions. No Moon is the typical conditions that the FD shift is run under.

Figure 4.2a plot the integral percent of the night time within a year with the moon less than a particular illuminated fraction. Only times between astronomical dusk and dawn has been included. The solid line shows the percent of the year is astronomical twilight (sun at least 18°below the horizon. The dotted lines represent relaxed definition of twilight with the sun being at 15°, 12° and 9° below the horizon. The line curves up slowly due to that throughout the moon orbital cycle it will be visible during the day. After first quarter and before 3rd quarter will the moon be visible for longer at night time. A combination of loosen the definition of twilight and observing under indirect moonlight could increase the duty cycle by 5% to 10% across a year.

Figure 4.2b shows the relative change in brightness of the moon depending of the observed illuminated fraction. The red line is a fitted exponential to U-Band data taken at the La Palma observatory. The graph shows the large different in moon brightness as the illuminated fraction increases which was also described in Table 4.4. The slope is comparable to actual data taken from the background files measured by the FD telescopes. The ADC variance plotted against fraction of illuminated moon is shown in Figure 4.3.

### 4.2.3 Toy Model of the effect of increased Night Sky Background on aperture

A simple measure of how well a detector will operate is to calculate the signal-to-noise (S/N) ratio. The steps that I followed was outlined in ???. The signal-to-noise ratio is denote by:

$$\text{Signal to Noise} = \frac{S}{N} \quad (4.8)$$

where S is the signal and N is the noise. The first step is to calculate the noise (N) expected from a background signal (B). First the background photo-electron count (B) seen from the night sky is calculate via:

$$B \propto \epsilon A b \Delta\Omega \Delta t$$

where  $\epsilon$  is the optical efficiency of the telescope,  $A$  is the mirror area,  $b$  is the background light flux and  $\Delta\Omega$  is the solid angle of the sky viewed by a single PMT.  $\epsilon$  is  $f \times Q$  where F is the optics factor shown in Table. 4.3 and Q is the quantum efficiency of

the photon sensitive device (I.E. photomultiplier tube). Next the noise can be found which is the square root of the background signal:

$$N = \sqrt{B} \propto \sqrt{\epsilon A b \Delta\Omega \Delta t} \quad (4.9)$$

Next step is to calculate the signal ( $S$ ) from an extensive air shower that a detector would observe.

$$S \propto \frac{\epsilon A n_+ n_\gamma c \Delta t}{4\pi R^2} e^{-R/\lambda_R} \quad (4.10)$$

where  $n_+$  is the number of charge particles in the shower viewed by the photon sensitive device,  $n_\gamma$  is the photon yield per charged particle for atmospheric scintillation,  $R$  is the distance to the shower segment,  $c\Delta t$  is the length of the shower segment, and  $\lambda_R$  is the Rayleigh attenuation length of light in the atmosphere. Substituting Eq. 4.9 and Eq. 4.10 into Eq. 4.8 and removing the approximation gives:

$$\frac{S}{N} = n_+ n_\gamma c \frac{(1 + \cos\theta)}{4\pi(R \times \sin\theta)^2} \sqrt{\frac{\epsilon A \Delta t}{b \Delta\Omega}} e^{-R/\lambda_R} \quad (4.11)$$

where  $\theta$  is the viewing angle of the EAS.

From here I use value specific to the Fluorescence Detectors to calculate the S/N ratio expected. The majority of the values used are shown in Table 4.5. The viewing angle is  $90^\circ$  as the shower viewed at  $15^\circ$  with the center of the FD camera is at  $15^\circ$  elevation. As the perpendicular segment of the shower is being observed therefore  $R = R_P$  where  $R_P$  is the closest distance that the shower axis is to the FD. The Rayleigh attenuation length ( $\lambda_R$ ) is calculated via:

$$\lambda_R = 2974 \left( \frac{\lambda}{400 \text{ nm}} \right)^4 \text{ g/cm}^2 \quad (4.12)$$

where  $\lambda$  is the chosen wavelength in nanometres. Currently  $R_P$  is in units of length where the Rayleigh attenuation length calculated in Eq. 4.12 in units of mass per length squared. A diagram of how  $R_P$  and the shower axis is related to the detector is shown in Fig. 4.4. To find the distance  $R_P$  in terms of grammage the term  $\Delta X_P$  is defined. Substituting this term into Eq. 4.11 becomes:

$$\frac{S}{N} = n_+ n_\gamma c \frac{(1 + \cos\theta)}{4\pi(R \times \sin\theta)^2} \sqrt{\frac{\epsilon A \Delta t}{b \Delta\Omega}} e^{-\Delta X_P/\lambda_R} \quad (4.13)$$

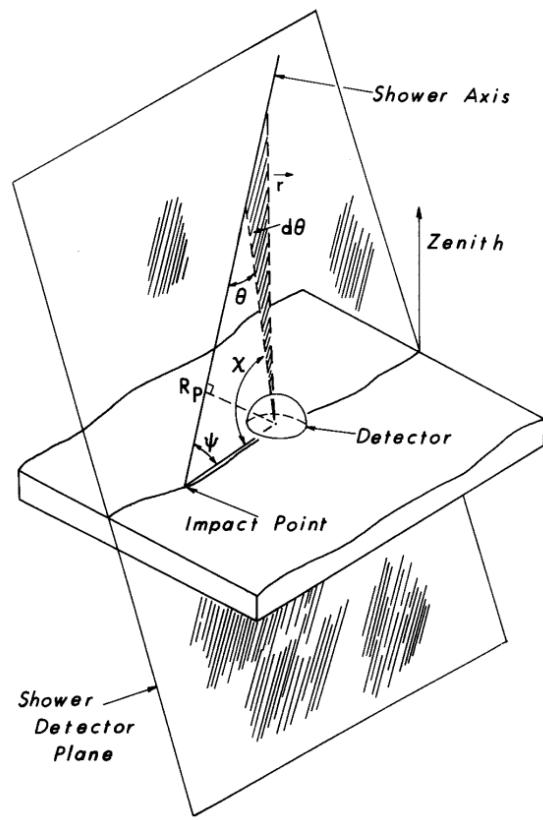
The  $\Delta X_P$  can be calculated from the distance  $R_P$  via

$$\Delta X_P = \Delta X_{\text{vertical}} / \cos 75^\circ \quad (4.14)$$

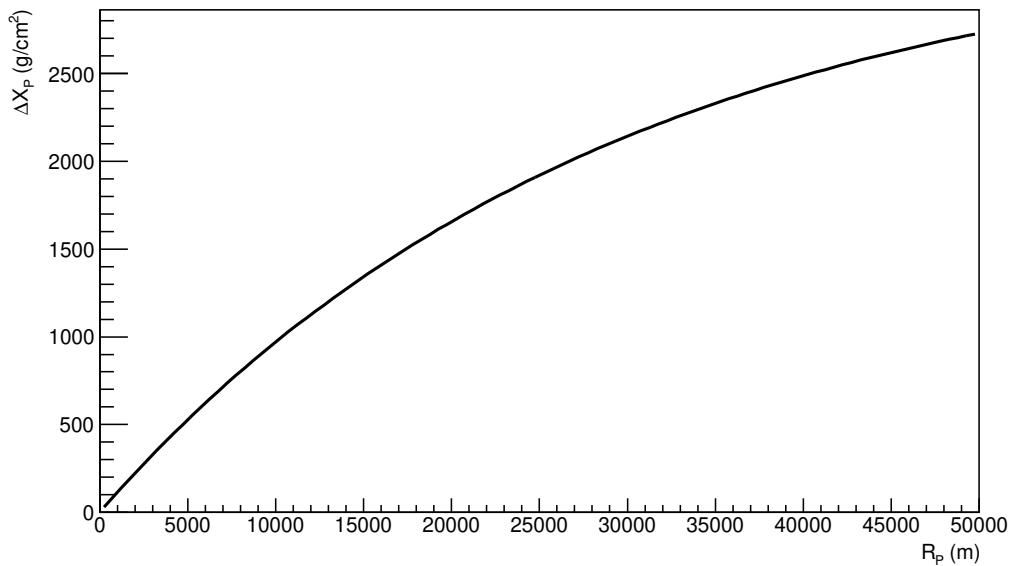
$$\Delta X_{\text{vertical}} = 860 - 860 \exp(-h/7500) \quad (4.15)$$

$$h = R_P \times \sin 15^\circ \quad (4.16)$$

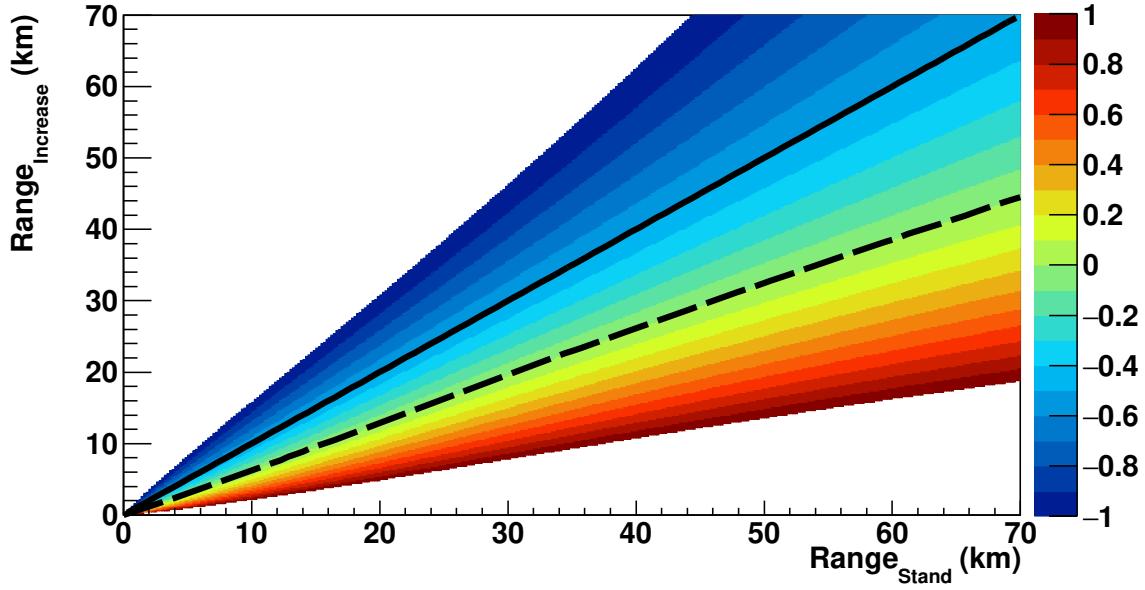
where  $h$  is the vertical distance in meters and  $\Delta X_{\text{vertical}}$  is the vertical path in  $\text{g/cm}^2$ , the shower axis is at  $15^\circ$  elevation and  $75^\circ$  is the zenith angle. To find the grammage along the path to  $R_P$  the vertical distance  $h$  is substituted and then the path  $\Delta X_P$  can be found. The relationship between  $\Delta X_P$  and  $R_P$  is shown in Fig. 4.5 for a detector viewing the sky at  $15^\circ$  elevation. The relationship is curved due to the amount of atmospheric molecules dropping off rapidly as the observed event moves further away both horizontally and vertically.



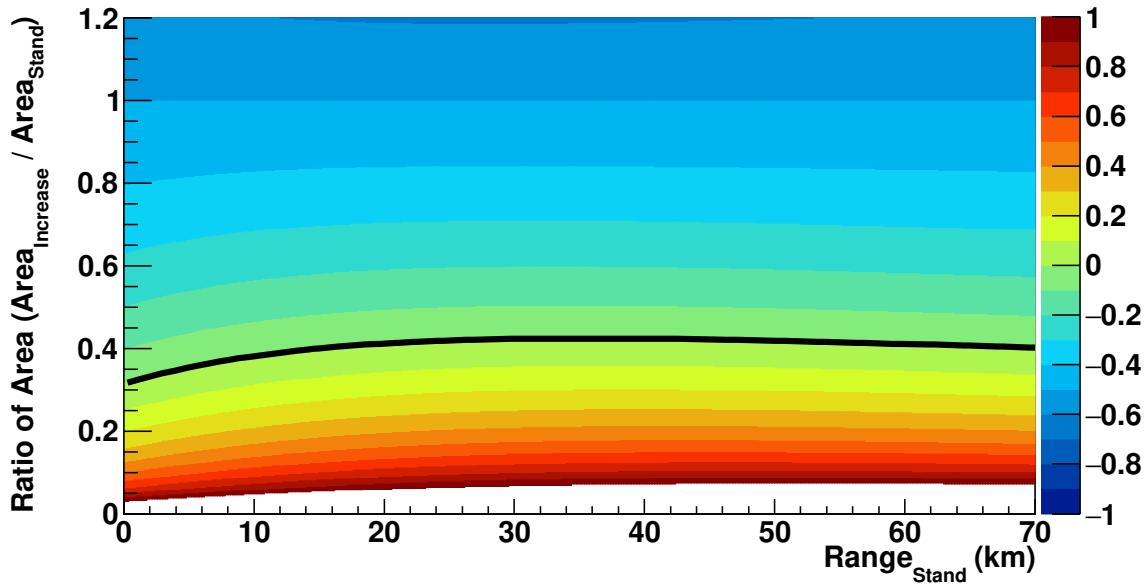
**Figure 4.4:** Diagram showing the parameters used. Also show how they relate to the plane of the shower axis and to the position of the detector.



**Figure 4.5:** How the relationship between the distance to the shower axis and grammage along the path. In this case the distance to the shower axis is denoted  $R_P$  and the grammage is denoted  $\Delta X_P$ .  $R_P$  is the closest distance that the shower axis is to the detector. The angle of the shower is  $15^\circ$



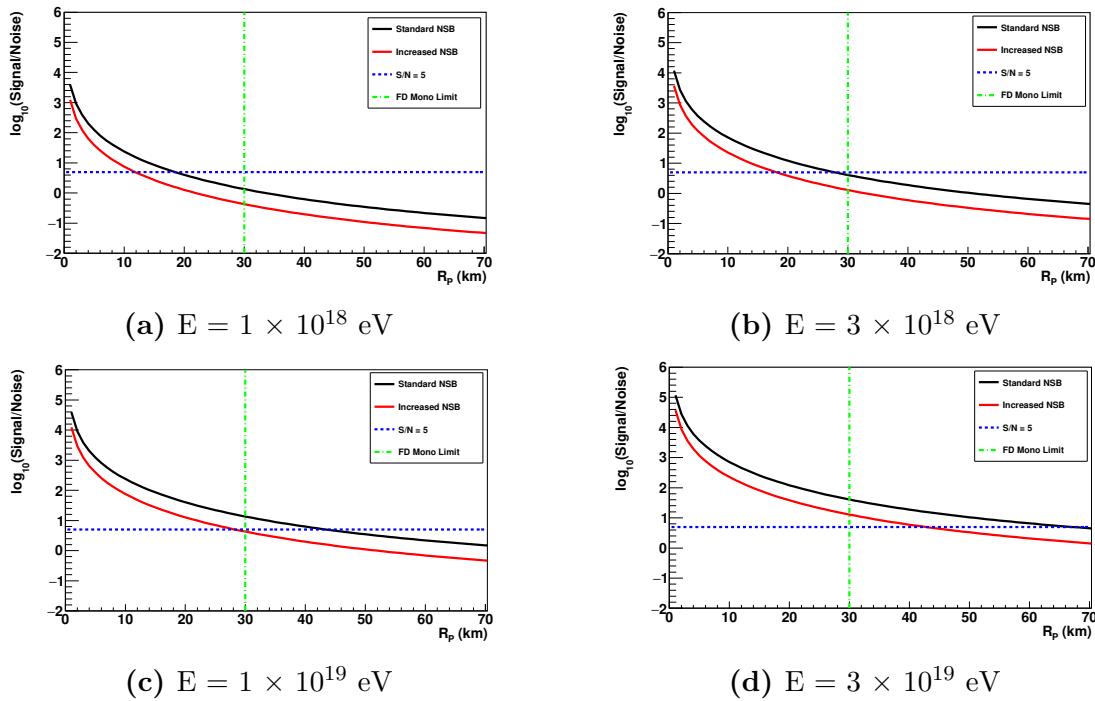
**Figure 4.6:** The image of the calculated Signal-to-Noise ratio the NSB at the Standard value and Increased NSB by a factor of 10 as a function of distance for an individual detector. The dashed black line denotes the distances where the Signal-to-Noise ratio is the same for the different NSB levles. The solid black line denotes where the distances are equal.



**Figure 4.7:** The image shows how the ratio of the detector viewing area changes with increasing the NSB by a factor of 10 as a function of distance from the detector. The solid black line denotes where the Signal-to-Noise ratio is the same.

optical efficiency	$\epsilon$	0.135
mirror area	$A$	3.8 m <sup>2</sup>
Standard background light flux	$b_1$	5e5 photons / m <sup>2</sup> / sr / $\mu$ s
Increased background light flux	$b_2$	5e6 photons / m <sup>2</sup> / sr / $\mu$ s
solid angle of a single PMT	$\Delta\Omega$	5.38e-4 sr
viewing angle	$\theta$	90°
speed of light	$c$	3e2 m / $\mu$ s
number of charged particles	$n_+$	Particle Energy / 1e9
Fluorescence yield	$n_\gamma$	4 photons / m / charged particle
Chosen wavelength	$\lambda$	350 nm
Rayleigh attenuation length	$\lambda_R$	1743 g / cm <sup>2</sup>

**Table 4.5:** Values specific to the Fluorescence Detectors.



**Figure 4.8:** Using the FD values in Table 4.5 and Eq. 4.11 to calculate the Signal-to-Noise ratio from specific values for the FD, NSB and conversion's from shower energy to photons to photo-electrons. The black solid line represents the S/N ratio at typical NSB flux seen at Auger, the red solid line represents the S/N ratio if the NSB flux was increased by a factor of 10, the blue dash line represents the single pixel trigger threshold, and the green dashed line represents the viewing limit of a single FD telescope before an event is closer to another FD telescope.

The result of calculating the theoretical of the Signal-to-Noise Ratio is shown in Fig. 4.6 and Fig. 4.7. The scale for both is shown in  $\text{Log}_{10}$  scale so the colour range above and below one is equidistant.  $\text{Range}_{\text{Stand}}$  is the expected S/N ratio for the range with standard observed NSB and  $\text{Range}_{\text{Increase}}$  is the expected S/N for the range with observed NSB increased by a factor of 10. Fig. 4.6 shows the relationship between the S/N ratio for  $\text{Range}_{\text{Stand}}$  and  $\text{Range}_{\text{Increase}}$ . The solid black line represents the S/N ratio for when the distances are equal. The difference between ratio at the solid black line is  $1/\sqrt{10}$ . The dashed black line represents where the S/N ratio are equal. It can be seen that  $\text{Range}_{\text{Increase}}$  is smaller than  $\text{Range}_{\text{Stand}}$  which is expected when the NSB has been increased for  $\text{Range}_{\text{Increase}}$ . Fig. 4.7 shows how much the observing area changes from  $\text{Range}_{\text{Stand}}$  to  $\text{Range}_{\text{Increase}}$ . The solid black line denotes where the S/N ratio is the same.

Table 4.5 and Eq. 4.11 were used to produce Fig. 4.8 which shows how the S/N ratio changes versus distance for cosmic-rays of four different primary energies. The black solid line represents the S/N ratio at typical NSB flux seen at Auger, the red solid line represents the S/N ratio if the NSB flux was increased by a factor of 10, the blue dash line represents the single pixel trigger threshold, and the green dashed line represents the viewing limit of a single FD telescope before an event is closer to another FD telescope. The figures show that for showers with energies above  $3 \times 10^{19}$  eV that increasing the NSB by a factor of 10 would expect no impact on the FD collecting area. For showers energies of  $1 \times 10^{19}$  eV and below it would be expected that the FD collecting area would be reduced as the NSB is increased.

## 4.3 Increasing Night Sky Background in Real Data to evaluate Reconstruction Efficiency

### - Aim

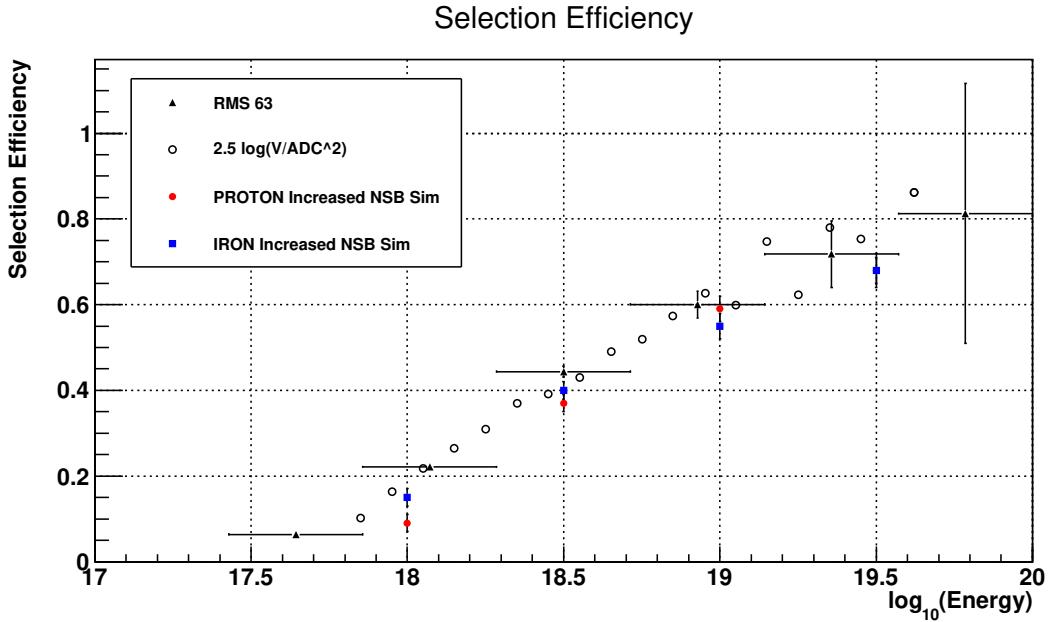
To test the effect of increasing the NSB levels in real data to evaluate the effects on reconstruction efficiency. The reconstruction tool that Auger employs is a internally developed software package called OffLine. OffLine takes selected shower events and goes through the process of finding signals within pixels, and calculating energy and arrival direction. The real data set was the showers used for analysis for ICRC 2015.

### 4.3.1 Method

#### - Take shower set

- add artifical noise to the raw pixel traces
- Use OffLine to reconstruct shower with the increased noise
- Compare Selection Efficiency, Energy resolution and bias, and Xmax resolution and bias.

To evaluate the effects of increased NSB levels on the reconstruction efficiency the data set used for ICRC 2015 had artifical noise add to the raw pixel traces. Once the desired NSB level has been added OffLine is used to reconstruct the shower events again. The results from the standard NSB levels and the increased NSB levels can be compare. I mainly focused on the comparison between Efficiency, Energy resolution and bias, and Xmax resolution and bias. This method allows for any NSB flux levels to be investigated.



**Figure 4.9:** The selection efficiency with my repetition against the data points from M. Unger. The method and data sets are similar with the NSB used differing slightly.

The Efficiency is comparison of events that are successfully reconstructed at standard NSB compared to the number of events that are successfully reconstructed at the increased NSB levels. The Efficiency was calculated with the equation used:

$$\text{Efficiency}_{\text{Data}} = N'_{\text{SelectData}} / N^0_{\text{SelectData}} \quad (4.17)$$

where  $N^0_{\text{SelectData}}$  is the number of selected events at the standard NSB level and  $N'_{\text{SelectData}}$  is the number of selected events at the increased NSB level.

After the Efficiency was calculated the bias and resolution for Xmax and energy was determined. For real data, the bias is the relative change in the mean of the distributions at increased NSB to the mean of the distributions at standard NSB, both after reconstruction and selection cuts. The bias calculations for real data become:

$$\Delta E_{\text{Data}} = \frac{E_{\text{IncreasedNSB}} - E_{\text{StandardNSB}}}{E_{\text{StandardNSB}}} \quad (4.18)$$

$$\Delta X_{\text{maxData}} = X_{\text{maxIncreasedNSB}} - X_{\text{maxStandardNSB}} \quad (4.19)$$

The energy and Xmax resolution is calculated via:

$$\sigma_{\text{resData}} = \left( \frac{1}{N} \sum \sigma_i^2 \right)^{1/2} \quad (4.20)$$

where  $\sigma_i^2$  is the reconstructed uncertainty on the individual shower events.

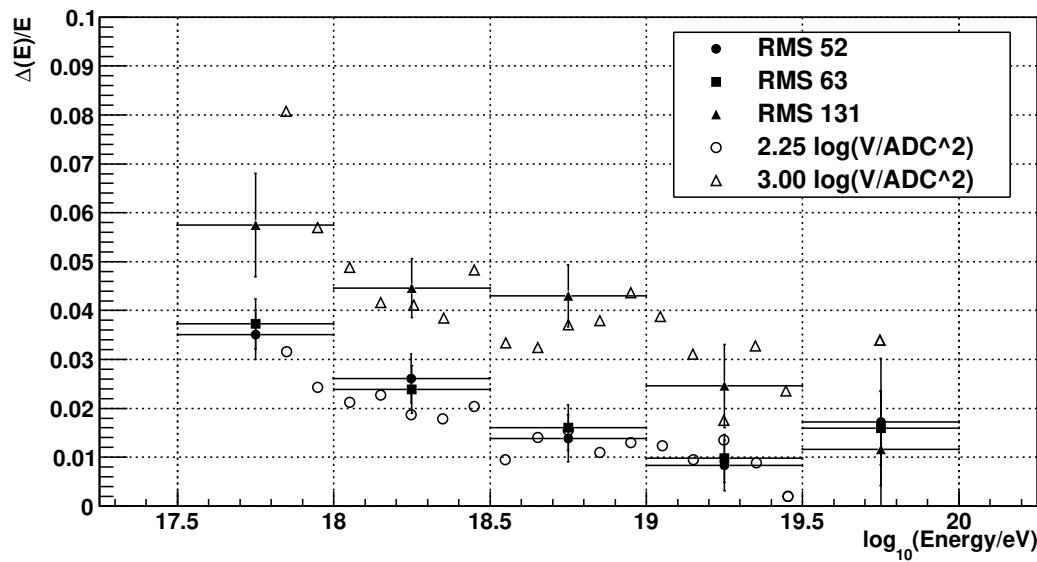


Figure 4.10: Energy Bias from increasing NSB in Real Data.

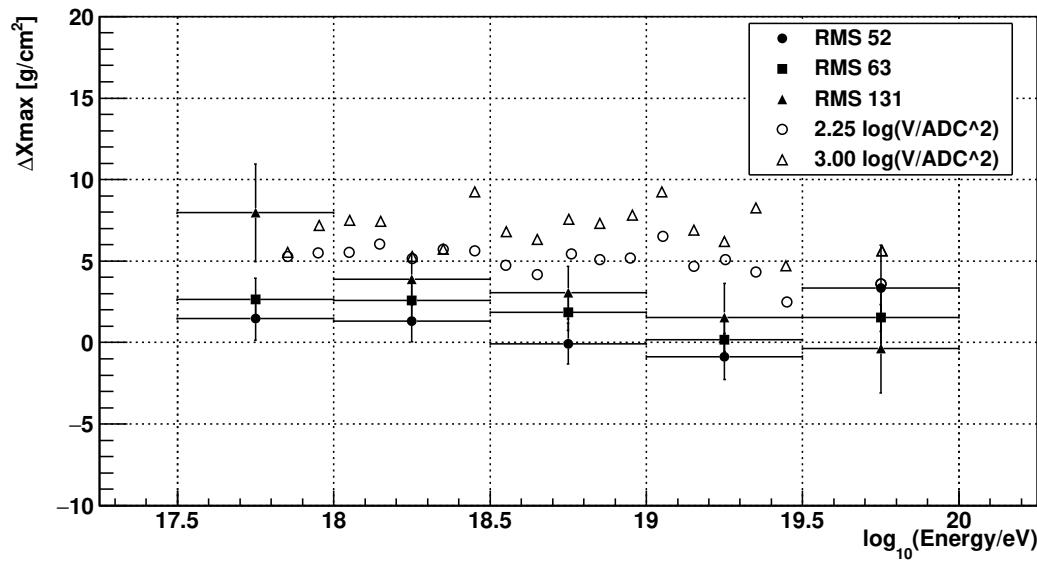
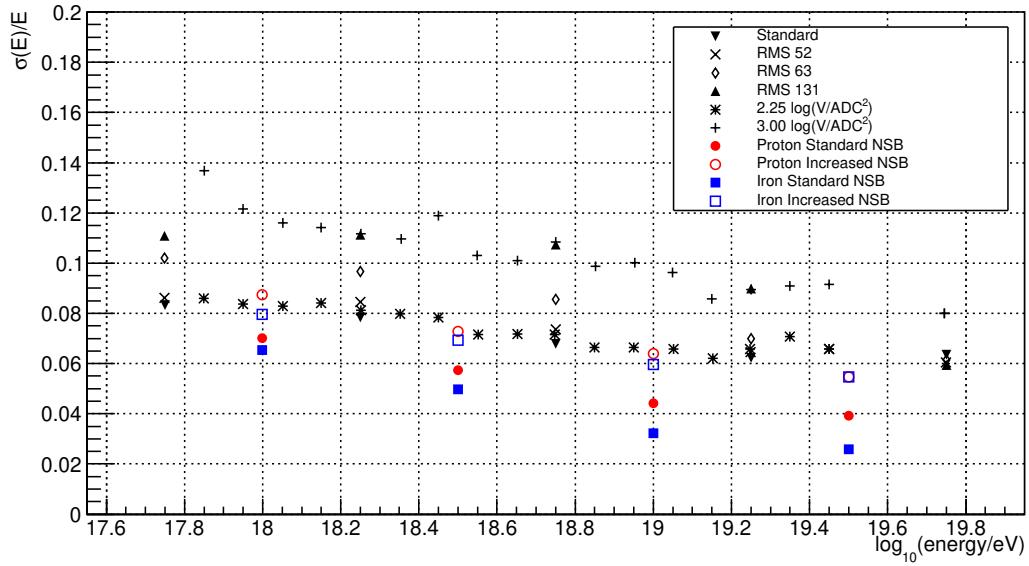
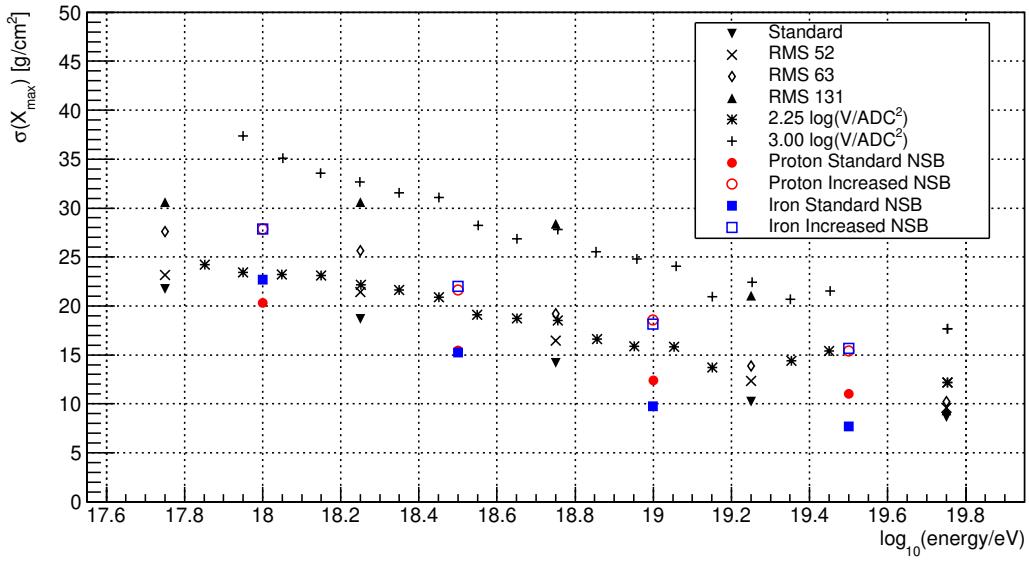


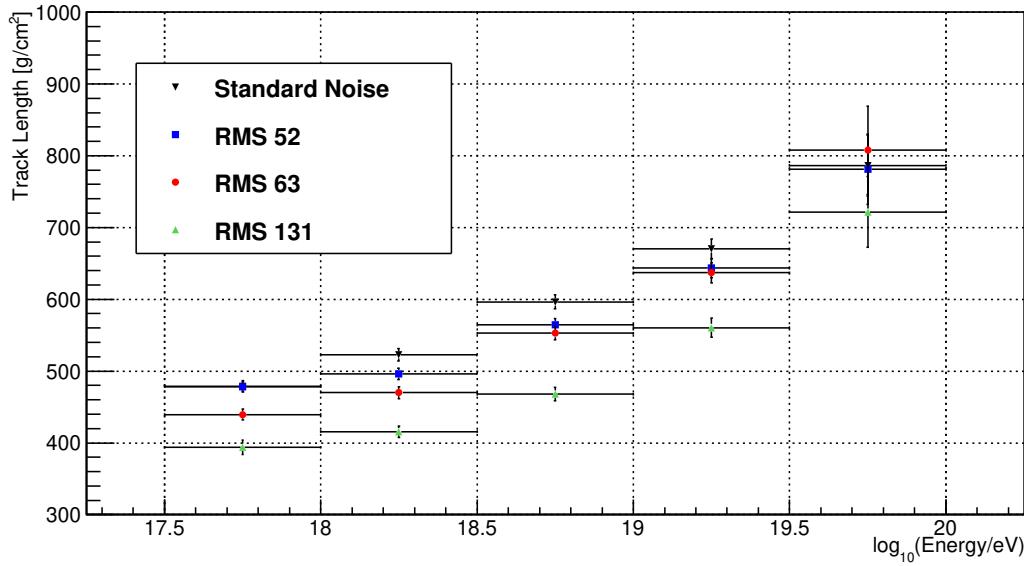
Figure 4.11: Xmax Bias from increasing NSB in Real Data.



**Figure 4.12:** Energy Resolution using both Smearing Method data and simulated showers.



**Figure 4.13:** Xmax Resolution using both Smearing Method data and simulated showers.



**Figure 4.14:** Track length calculated from events with standard NSB and those with increased NSB fluxes.

### 4.3.2 Results

### 4.3.3 Discussion

## 4.4 Increasing the Night Sky Background in Simulations to evaluate Trigger and Reconstruction Efficiency

### 4.4.1 Method

The Efficiency was calculated with the equation used:

$$\text{Efficiency} = N'_{\text{Select}} / N^0_{\text{Select}} \quad (4.21)$$

where  $N^0_{\text{Select}}$  is the number of selected events at the standard NSB level and  $N'_{\text{Select}}$  is the number of selected events at the increased NSB level.

For simulated data the bias is the relative change in the mean of the distributions after the full simulation by EAS events going through an atmosphere with a specified NSB photon field, the FD telescopes optics, trigger, reconstruction and selection cuts compared with Monte-Carlo truth.

$$\Delta E_{\text{Sim}} = \frac{E_{\text{recon}} - E_{\text{true}}}{E_{\text{true}}} \quad (4.22)$$

$$\Delta X_{\text{maxSim}} = X_{\text{maxrecon}} - X_{\text{maxtrue}} \quad (4.23)$$

The energy and Xmax resolution is calculated via:

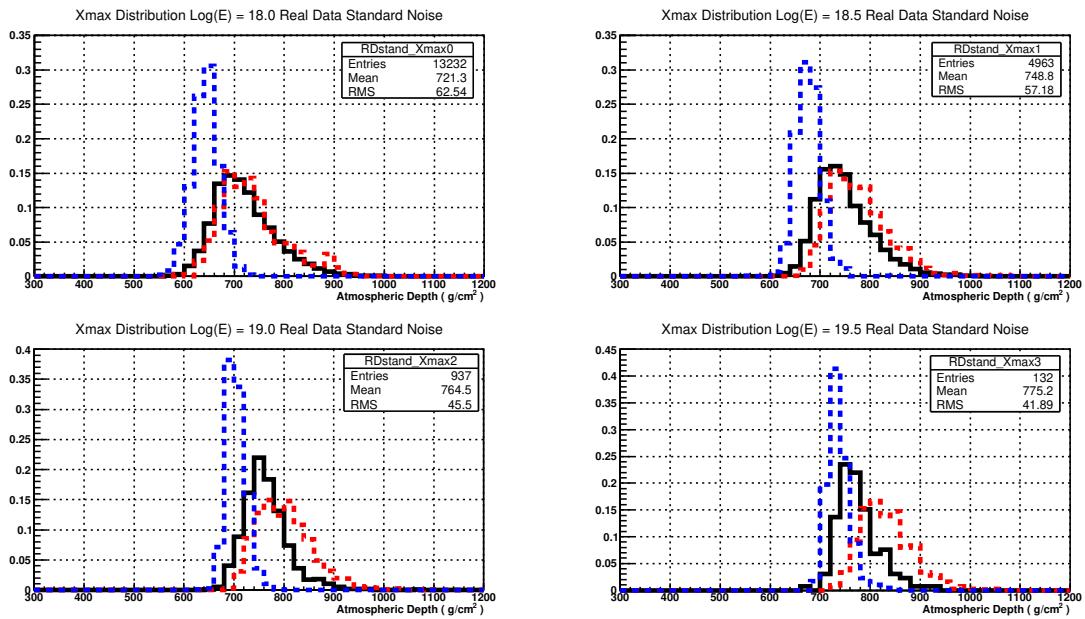
$$\sigma_{\text{resSim}} = \left( \frac{1}{N} \sum \frac{1}{\sigma_i^2} \right)^{1/2} \quad (4.24)$$

**4.4.2 Comparison of Simulated Data to Real Data**

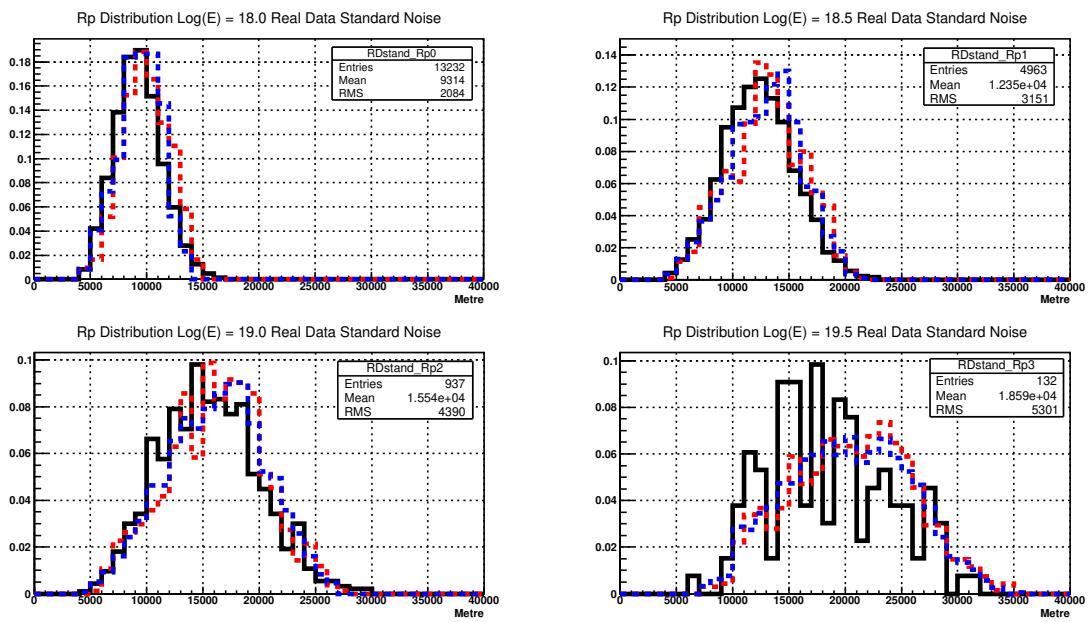
**4.4.3 Results**

**4.4.4 Discussion**

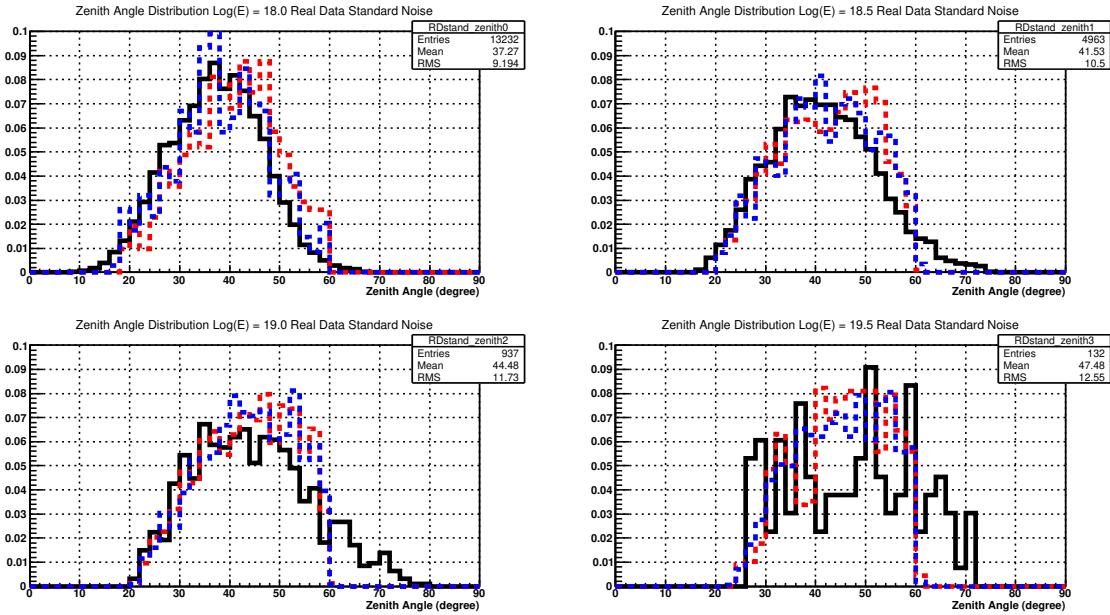
**4.5 Conclusion/Summary**



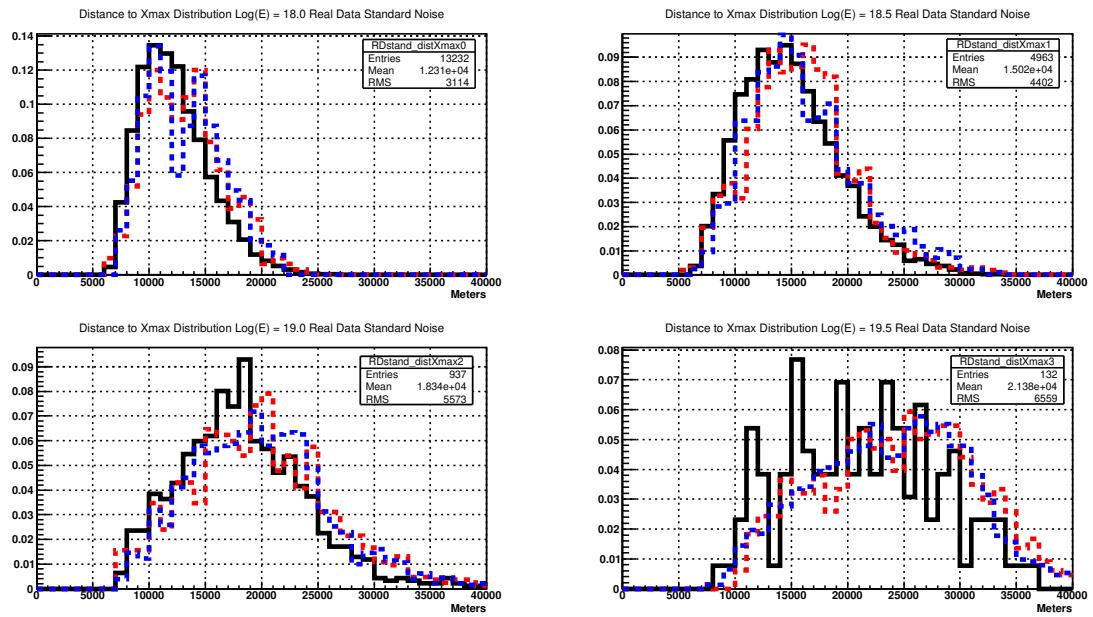
**Figure 4.15:** Distribution of Xmax with Real Data and simulation of proton and iron showers.



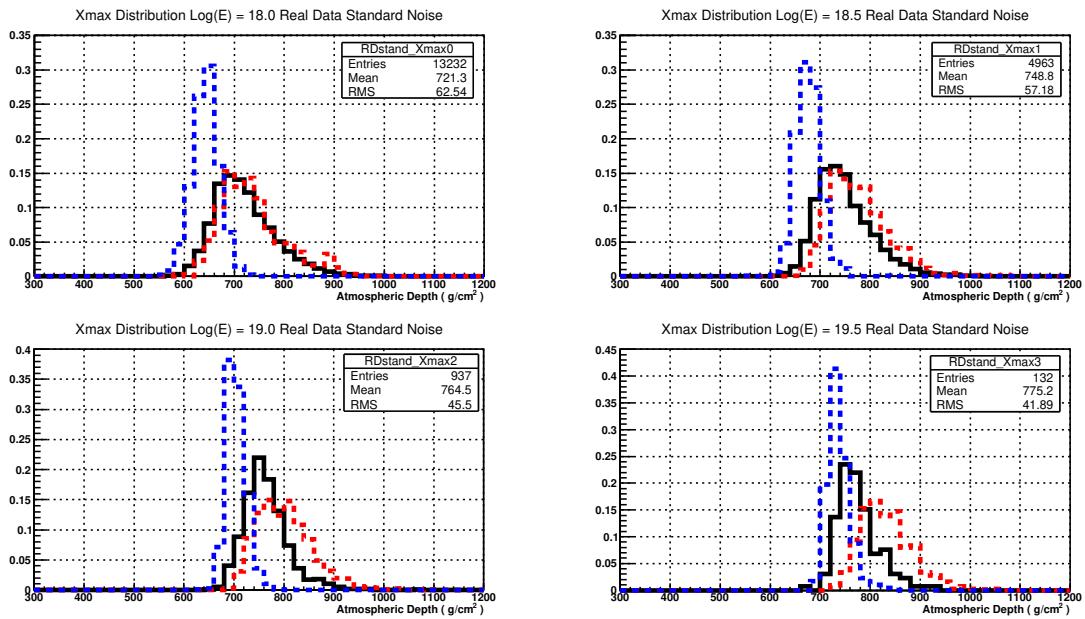
**Figure 4.16:** Distribution of Rp with Real Data and simulation of proton and iron showers.



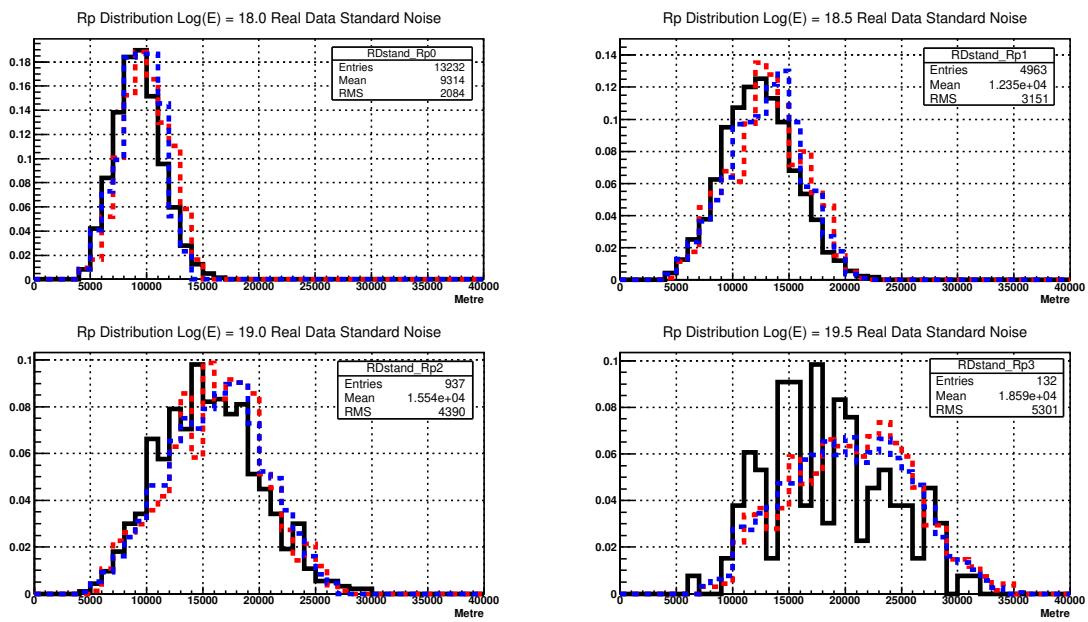
**Figure 4.17:** Distribution of Zenith angle with Real Data and simulation of proton and iron showers.



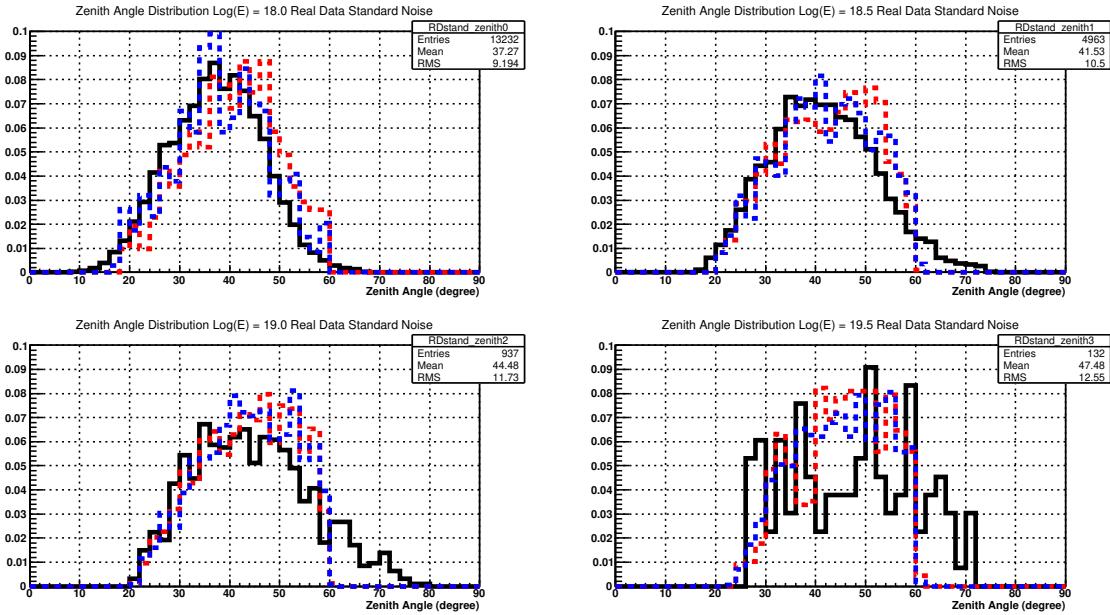
**Figure 4.18:** Distribution of Distance to Xmax with Real Data and simulation of proton and iron showers.



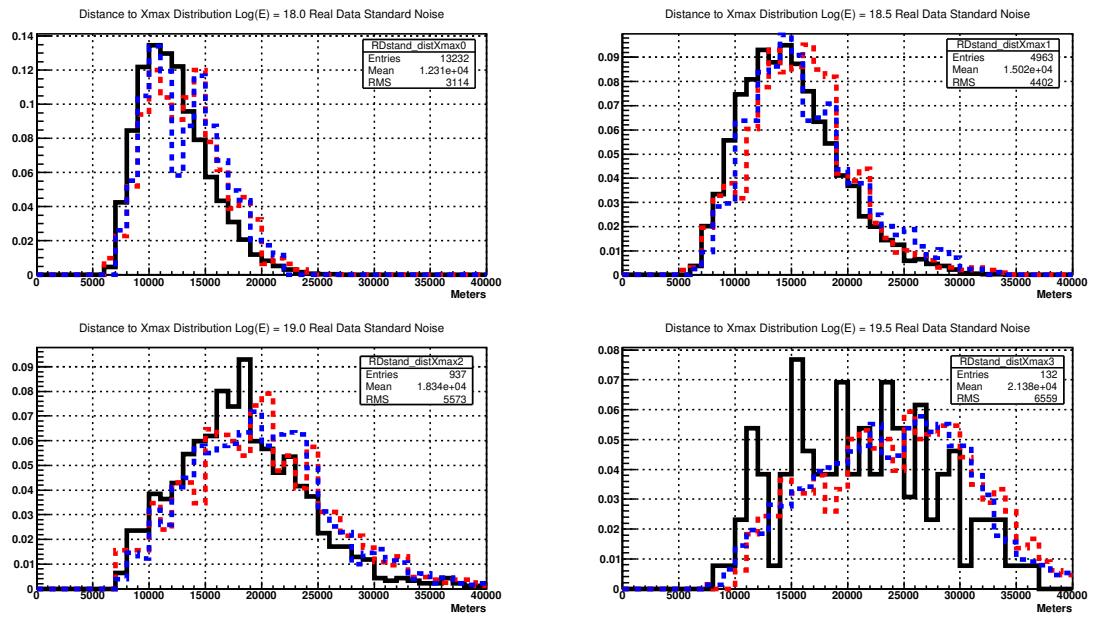
**Figure 4.19:** Distribution of Xmax with Real Data and simulation of proton and iron showers.



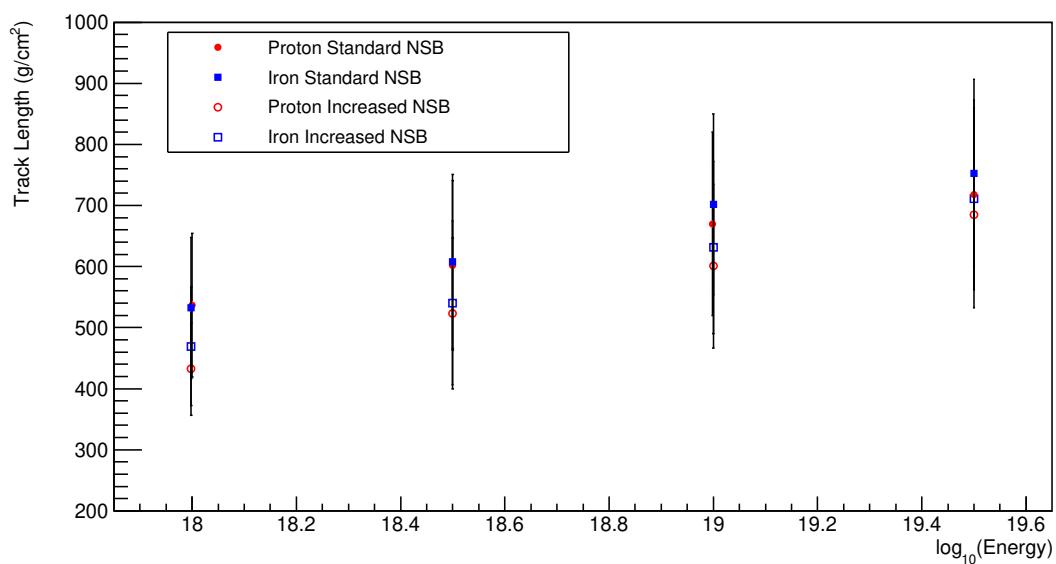
**Figure 4.20:** Distribution of Rp with Real Data and simulation of proton and iron showers.



**Figure 4.21:** Distribution of Zenith angle with Real Data and simulation of proton and iron showers.



**Figure 4.22:** Distribution of Distance to Xmax with Real Data and simulation of proton and iron showers.



**Figure 4.23:** Track length using simulation of proton and iron CONEX showers.

## Chapter 5

# Quantifying Characteristics of the Fluorescence Detector Photomultiplier Tube

Characterising the PMT at 600V and 900V

- Using the characteristics of the PMT at 900V as a baseline
- Measure linearity
- ND filters vs Two LED method
- temperature effects

### 5.1 Motivation

- Start off with moonlight explanation
  - Looking at effects of increased NSB on equipment
  - What type of PMT used

The Fluorescence Detectors (FDs) at Auger use Photomultiplier Tube (PMT) as camera pixels within the telescopes to view the Extensive Air Showers. The PMTs are operated at a gain of  $5 \times 10^4$  electrons per photo-electron. This gain relates to a High Voltage about 900V. The investigation into running the Fluorescence Detectors under moonlight has lead to the thinking of running the PMTs under decreased PMT gain. The PMT gain is expected to decrease by a factor of 10 to compensate. Therefore the PMT High Voltage would be dropped to about 600V.

The aim is to use the measurements done at 900V as the baseline of expected performance and then observe the repeated measurements at 600V. We have a good understanding of the characteristics of the PMT operating at 900V due to this being the nominal voltage that the PMTs in the FD telescopes are run at. I explore the characteristics at 600V as we are expecting to reduce the PMT gain by a factor of 10 and this related to an approximate voltage drop to 600V. This may be an over estimate but we can be confident if the PMT voltage used is greater than 600V then the characteristics will be the same.

## 5.2 PMT Linearity

One of the main characteristics of a Photomultiplier Tube (PMT) to measure is its linearity. A PMT linearity is where if a light intensity is doubled the PMT response doubles as while. It is important to quantify this region to be confident that if the light intensity changes the PMT response follows.

I look at two methods to measure the PMT linearity - using Neutral Density filters and the Two LED method.

### 5.2.1 Neutral Density Filters

The method of using Neutral Density filters involves employing filters to reduce the light intensity reaching the PMT cathode by a known amount. The filters were used singularly and in combination to get different Optical Transmission. The filters are label with optical density which is different to optical depth. To convert optical density to transmission:

$$\text{Optical Depth} = -\log_{10} \left( \frac{\text{Transmission}}{100} \right) \quad (5.1)$$

To get different transmission values, the filters can be stack and the optical depth values just add together (IE. if using two filters one with ND 0.4 and ND 0.1 the total ND will be 0.5).

I repeated a study done by Privitera et. al. (1990) [ref.](#) on Auger PMTs. I wanted to use the result done at an High Voltage value of 900V as my baseline then continue the study to observe the result at an High Voltage value of 600V.

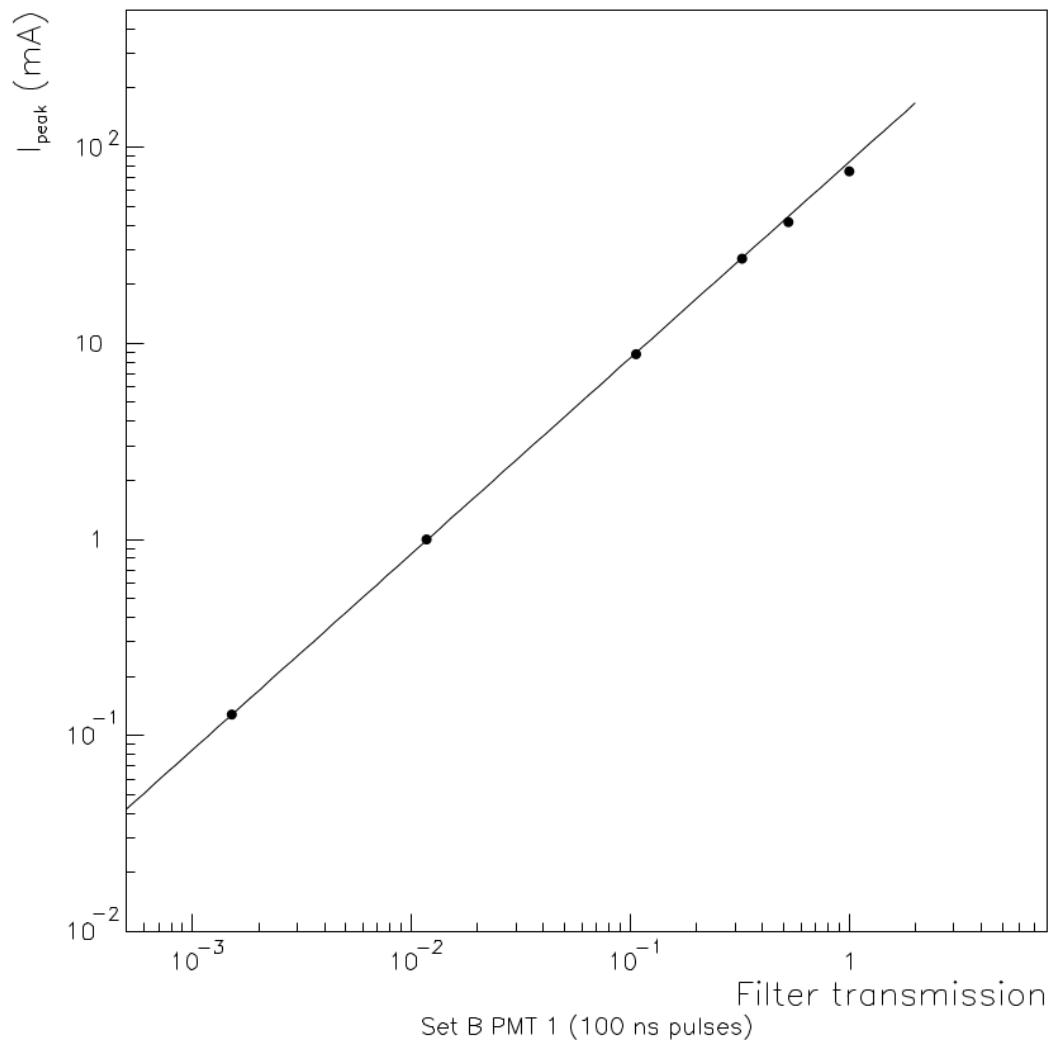
Some of the disadvantages of using Neutral Density filters it that you are restricted to fixed units of transmission depending on which filters you have access too. Also you are relying on the fact that the optical density doesn't change over time and is not effected by handing and storage. There is also some uncertainty on the optical density value quote.

### 5.2.2 Two LED Method

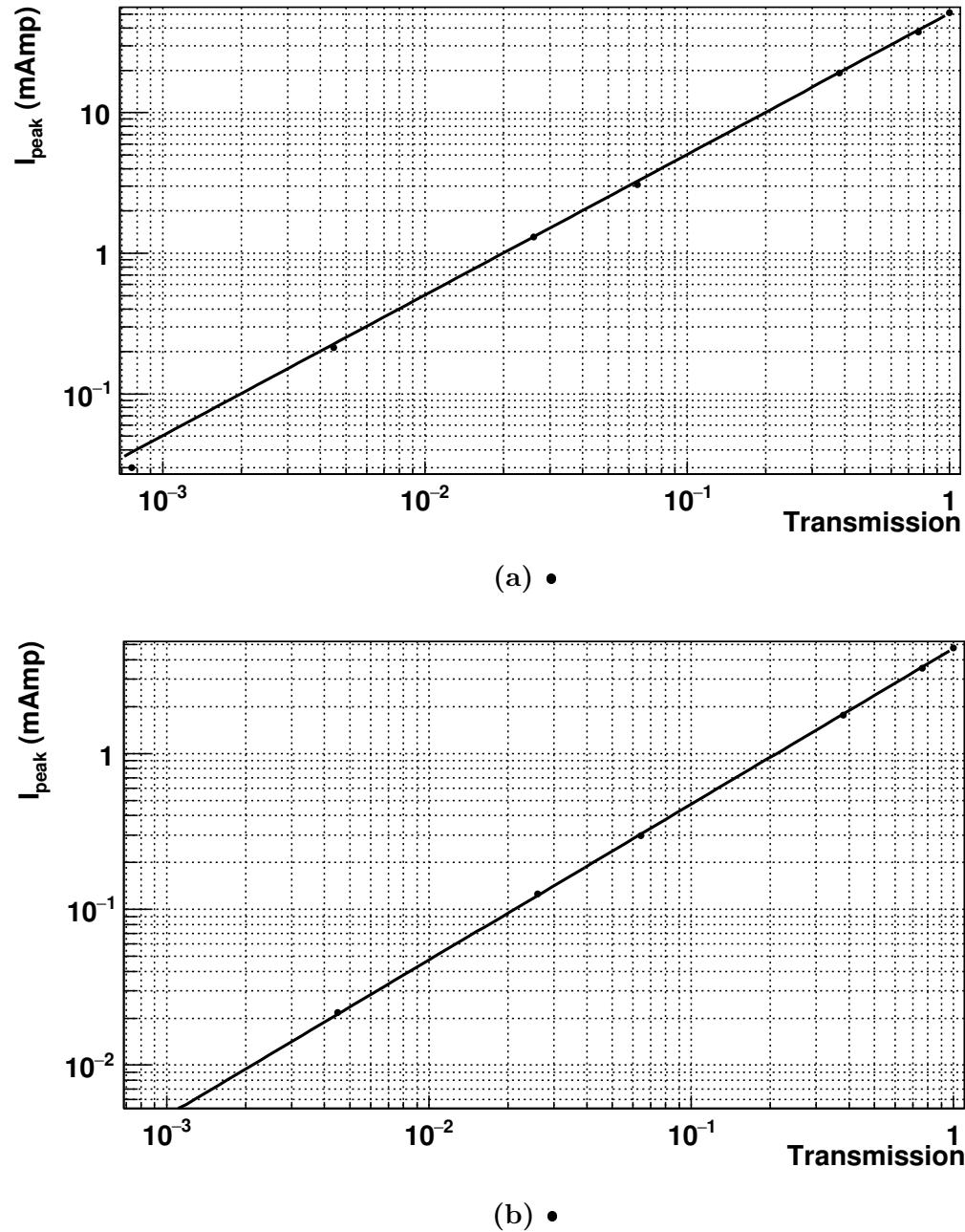
Another method investigate was the two LED method. This method was first used in the original measurement of the SD PMTs [ref.](#) To measure the linearity, two LEDs are pointed at the PMT and with one set of observations taken with the LEDs on separately then one set taken with the LEDs on at the same time. A graphical representation is shown in Fig. 5.3 of the set-up used to measure the PMT linearity at both 900V and 600V. Two pulse generators are used to drive the LEDs separately and to allow the two pulses to overlap when needed to measure the PMT response to the LEDs on at the same time. To find the linearity a ratio is plotted against peak pulse ampere. the ratio is calculated via:

$$\text{Ratio} = \frac{\text{LED}_{1+2}}{\text{LED}_1 + \text{LED}_2} \quad (5.2)$$

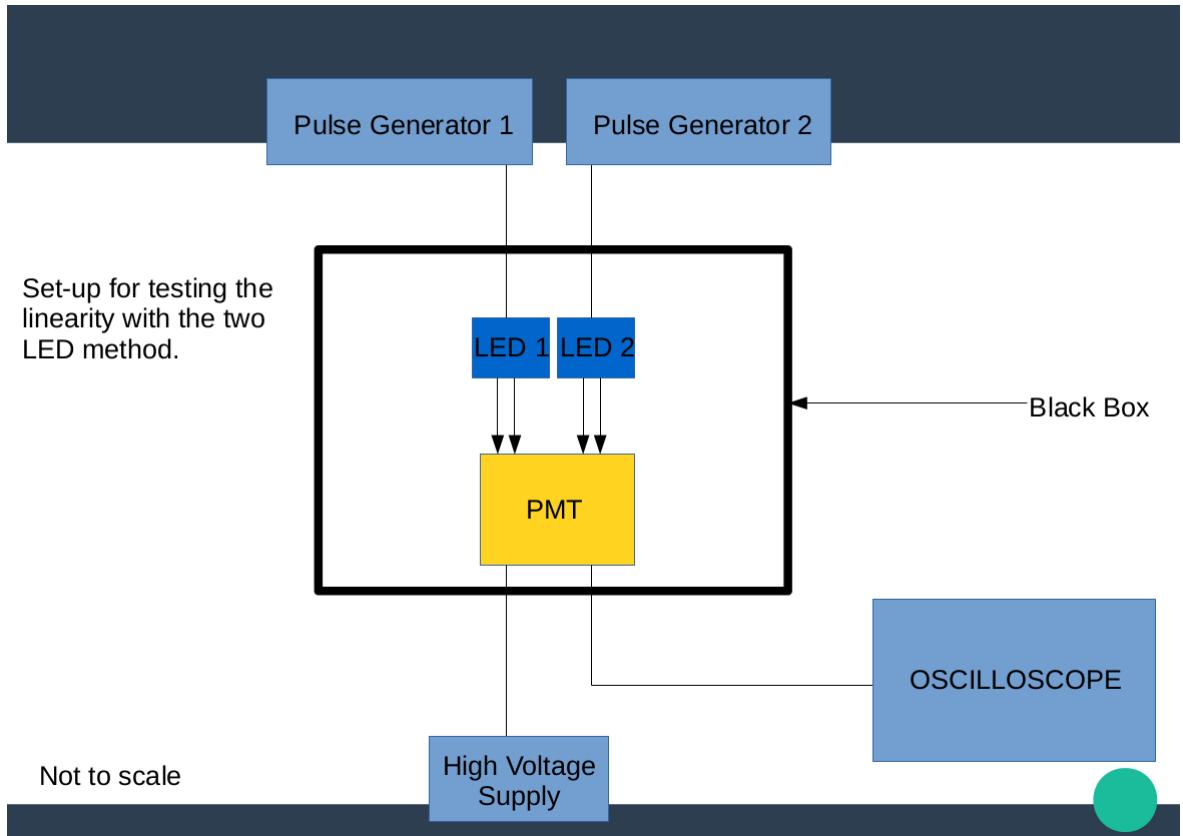
where  $\text{LED}_{1+2}$  is the measurement when both LEDs are on at the same time, and  $\text{LED}_1$  and  $\text{LED}_2$  is the measurement when they are on separately. The principle employed here is that while the intensity is within the linearity region the peak of  $\text{LED}_{1+2}$  should equal the addition of  $\text{LED}_1$  and  $\text{LED}_2$  separately. Therefore the ratio should be one while both peak anode ampere from the LEDs on separately and LEDs on together are within the region of linearity. As the light intensity is increased the LEDs on



**Figure 5.1:** Previous PMT linearity test done by Privitera et. al. 1999 with Neutral density filters. The linearity of the PMT is shown by peak of the current against the filter transmission.



**Figure 5.2:** Neutral density method at 900V and 600V.



**Figure 5.3:** A graphical representation of the Two LED Setup used in the lab at University of Adelaide.

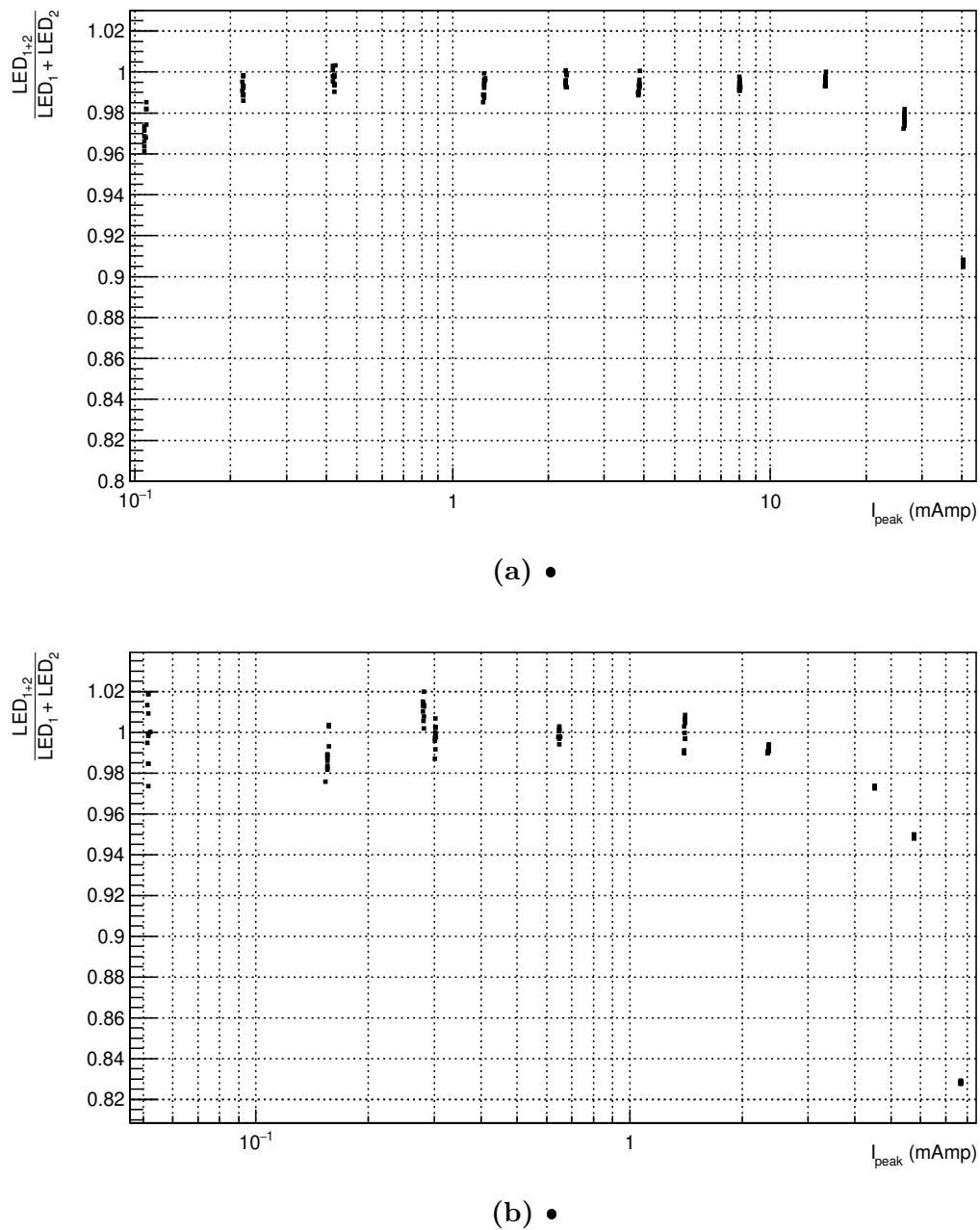
together will first move out of the region of linearity first. Typically this means that the peak ampere measured from both LEDs on at the same time will be less than the combination of the peak ampere when measured separately. This will cause the ratio to dip below one.

The advantages of this method is that any LED levels can be used and do not have to worry about the two LEDs linearity. As long as the two LED intensities are different it will produce a different response in the PMT.

The result of the two LED Method is shown in Fig. 5.4 for both 900V and 600V. The x-axis is the peak ampere of the measured pulses when both LEDs are on at the same time. The plots show that while the peak anode currents change by a close factor of 10, the gain of the PMT has also been reduced by a similar factor. This seems to indicate that the cathode is the limiting factor for the PMT linearity for this type. Maybe quote theoretical peak anode current for a  $10^{21}$  eV shower.

### 5.3 Effects of Temperature on PMT Gain

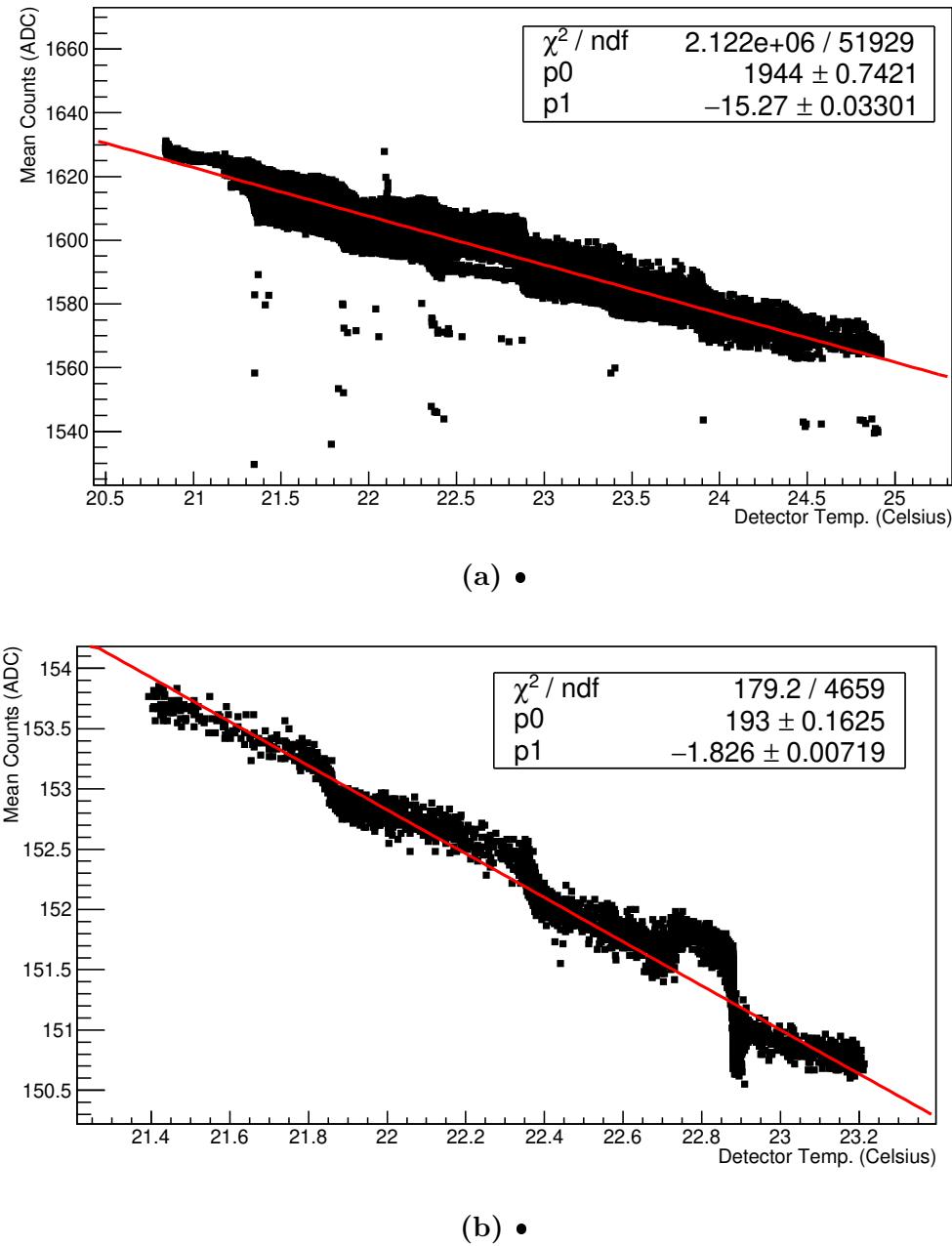
Quoted on any PMT data sheet is the expected change in gain as a function of energy. The XP3062 PMT used in the FD telescopes is quoted to be  $-0.2\% \text{ } \text{K}^{-1}$ . I measured this value in the lab at University of Adelaide at both 900V and 600V. The set-up to monitor the gain of the PMT involved a LED inside a copper oven pointed at the cathode pulsed once per second, and a temperature gauge. I was able to control the laboratory temperature to see the gain change over a variety of temperature values.



**Figure 5.4:** Two LED method at 900V and 600V.

The LED was inside a copper oven to maintain an independent constant temperature so the number of photons emitted could be considered to be fixed. The PMT was inside a box that highly light proof - the top was held down with screws, then two layers of black electrical tape was applied to the lid joinings and then a black out curtain was draped over the top. Only the temperature was allowed to vary over the measurement period. Therefore any changes in the absolute ADC measurements are related to changes in the PMT gain.

Measurements for temperature vs PMT gain was taken at both High Voltage values of 900V and 600V. The plots in Fig. 5.5 the effects of temperature on the PMT can be clearly seen at both high voltage values. For either high voltage value the measured PMT gain change per K was  $-0.6\%\text{K}$  needs uncertainty. This maybe different to the quoted data sheet but is not unreasonable. The quoted data sheet is an average of many PMTs while this is a measurement of one particular PMT. What was interested was that this value did not deviate if using the PMT at 900V or 600V.



**Figure 5.5:** Showing the how PMT gain correlates with surrounding temperature.

# Chapter 6

## Computer Simulation of Fluorescence Detector Photomultiplier Tubes

Simulating the FD PMT under differing NSB and for different reasons.

- Theoretical value for Gain Variance
- PMT Gain Variance
- Show both for flat distribution and Gaussian variations for dynodes
- Results
- FD FLT under increased NSB

### 6.1 Motivation

- PMT Theory/Background

### 6.2 Method and Theory

For a deeper understanding of how the PMT gain variance works I used a toy model Monte Carlo to simulate electrons moving throughout the PMT dynode chain. Gain Variance is the Gaussian broadening of the measured anode signal. The number of electrons that are emitted from each dynode follows a Poisson distribution. When an electron hits a dynode the number of electrons emitted follow a probability distribution that is Poission in shape. As electrons travel further down the dynode chain and the number of electrons emitted per dynode increase the distribution becomes more Gaussian like.

To simulate the Gain Variance within a PMT I set-up my toy model to match the number of stages with the XP3062 PMT. I was only looking at the single photo-electron case (ie. where only a single electron is emitted from the cathode.). If there was no extra broadening then the distribution would be dictated by the Poisson distribution. For each electron that hits a dynode a random number generator is used to determine the number of electrons that leave the dynode.

- Lot more detail. Through how the toy model is produced.

**Draw flow chart of toy model Monte Carlo.**

### 6.2.1 Theoretical value of Gain Variance?

I investigated different scenarios - firstly each dynode able to emit electrons with uniform Poisson distribution across its surface and secondly each dynode having some sort of irregularity across its surface. The irregularity was added by shifting the mean of the Poisson distribution by sampling from a random Gaussian distribution. The gain variance was investigated at different PMT gain values. The PMT gain values picked was equivalent to a high voltage across the PMT of 1300V, 900V and 600V. The high voltage of 1300V was picked as a direct gain variance measurement was preformed by **some group within Auger**, 900V is the typical high voltage used when the FD telescopes are operating and 600V is the approximate high voltage that is expected to be used when the FD telescopes are observing under increased NSB conditions.

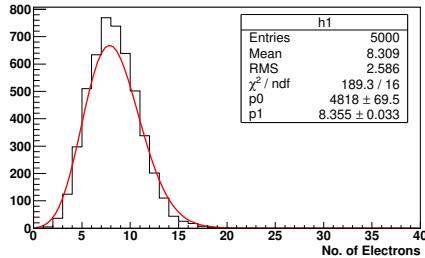
## 6.3 Results of PMT Gain Variance Simulation

### 6.4 Simulation of Gain Variance Method

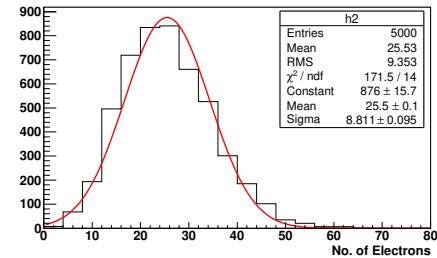
- Show off histogram of simulated Gain Variance Ratios
- Show that there's a natural spread
- some of the calculated ratio's will be less then one due to spread in measure variances
- differences in methods? Pairs vs averages method should return the same results
- benefits of combining noise traces over using just 140 bins per trace.

### 6.5 FD FLT under different NSB levels

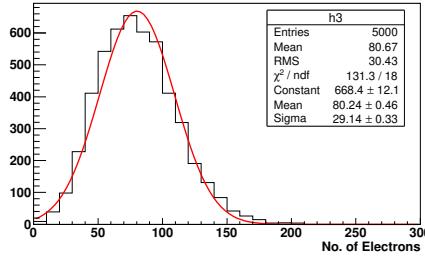
Looked at study to simulate the effects of increased NSB on the First Level Trigger (FLT). The Auger maintains the FLT to trigger at 100 Hz range. From simulations with different NSB levels it can be seen trigger threshold above the mean would be required to maintain the expected trigger level. The NSB of  $2.71\backslash 100$  ns is the typical NSB level observed at the site. From there I picked 5 increased NSB levels to show how threshold above the mean changes.



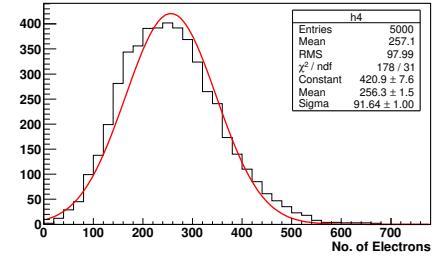
(a) Distribution of simulated electrons leaving dynode 1. The red line is a fitted Poisson distribution.



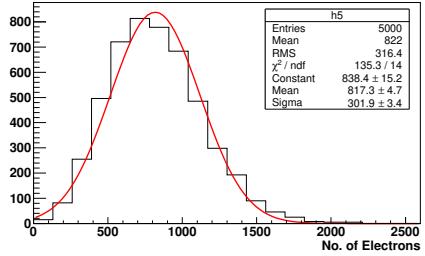
(b) Distribution of simulated electrons leaving dynode 2. The red line is a fitted Gaussian distribution.



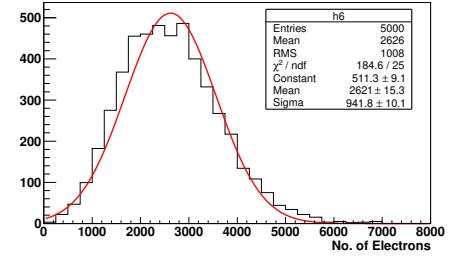
(c) Distribution of simulated electrons leaving dynode 3. The red line is a fitted Poisson distribution.



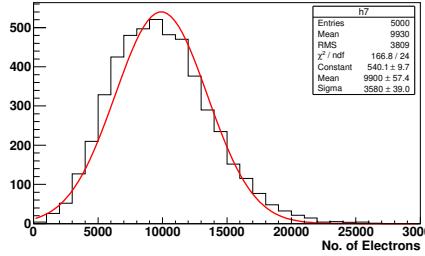
(d) Distribution of simulated electrons leaving dynode 4. The red line is a fitted Gaussian distribution.



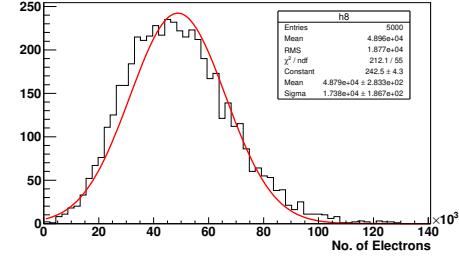
(e) Distribution of simulated electrons leaving dynode 5. The red line is a fitted Poisson distribution.



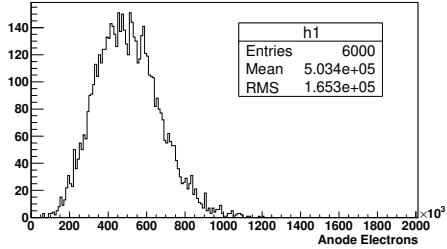
(f) Distribution of simulated electrons leaving dynode 6. The red line is a fitted Gaussian distribution.



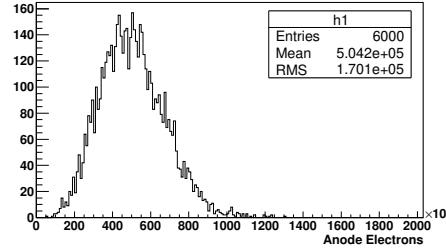
(g) Distribution of simulated electrons leaving dynode 7. The red line is a fitted Poisson distribution.



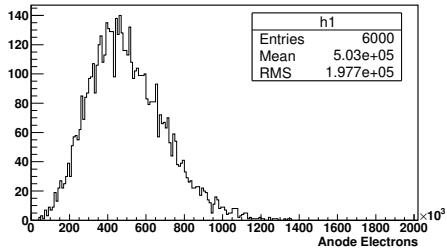
(h) Distribution of simulated electrons leaving dynode 8. The red line is a fitted Gaussian distribution.



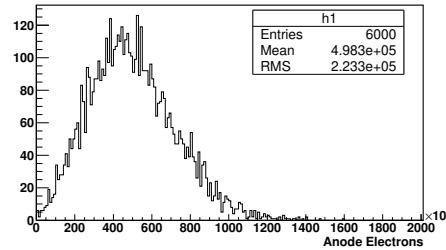
(a) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^5$ . Poisson fluctuations at each stage only. No each Gaussian broadening at any dynode.



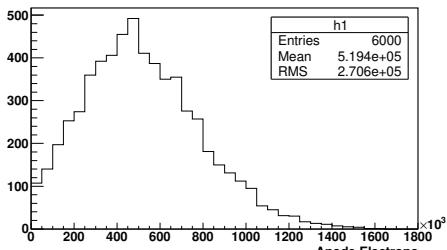
(b) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^5$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 10%.



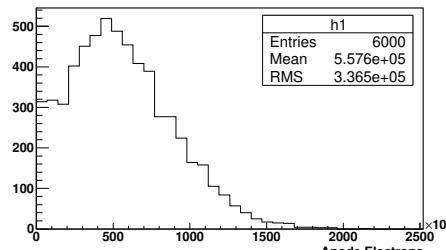
(c) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^5$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 20%.



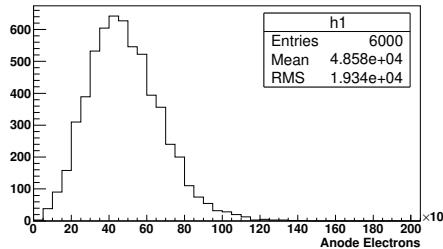
(d) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^5$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 30%.



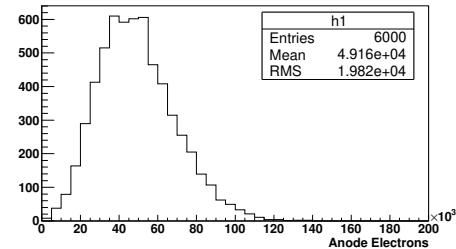
(e) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^5$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 40%.



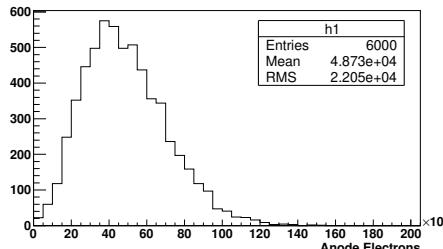
(f) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^5$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 50%.



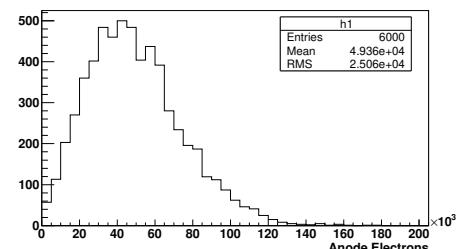
**(a)** Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^4$ . Poisson fluctuations at each stage only. No each Gaussian broadening at any dynode.



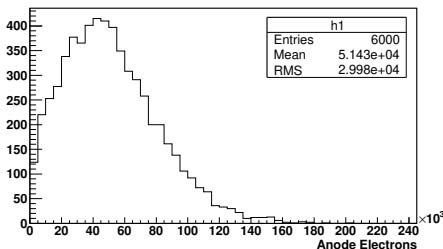
**(b)** Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^4$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 10%.



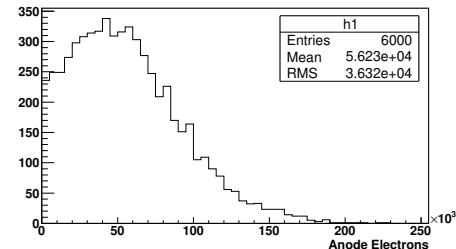
**(c)** Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^4$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 20%.



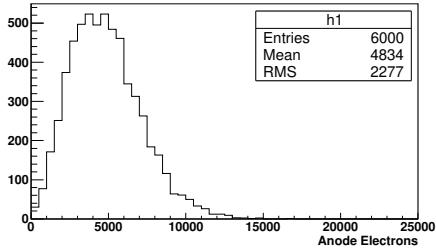
**(d)** Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^4$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 30%.



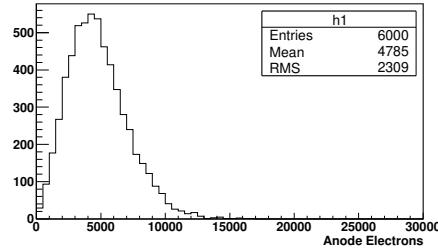
**(e)** Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^4$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 40%.



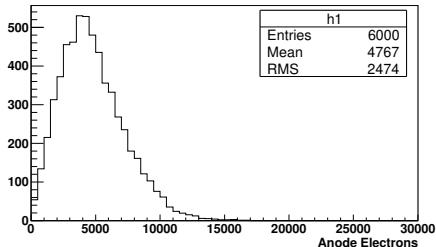
**(f)** Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^4$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 50%.



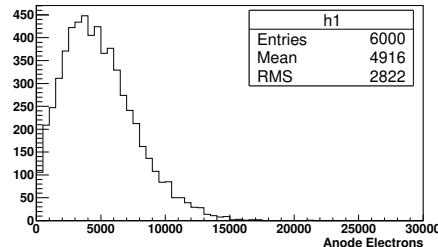
(a) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^3$ . Poisson fluctuations at each stage only. No each Gaussian broadening at any dynode.



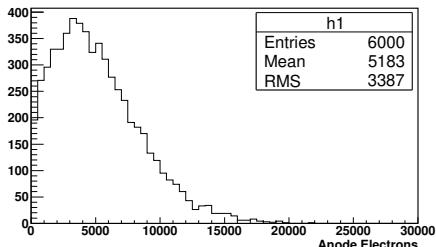
(b) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^3$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 10%.



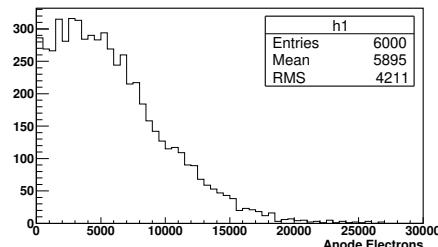
(c) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^3$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 20%.



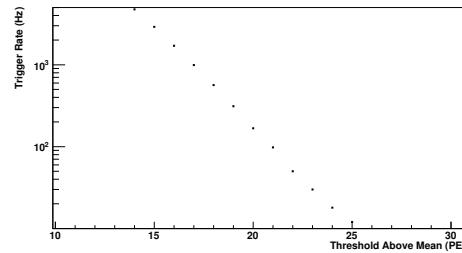
(d) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^3$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 30%.



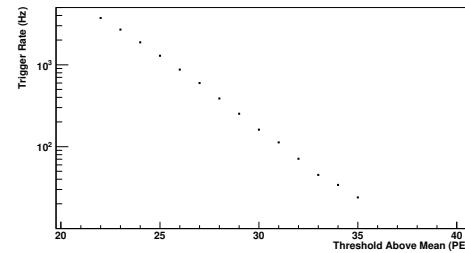
(e) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^3$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 40%.



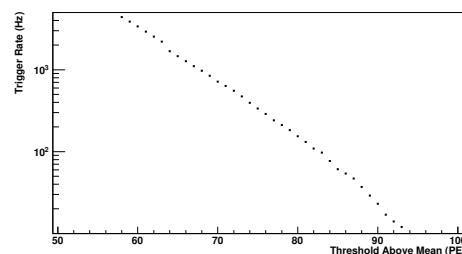
(f) Simulated observed signal from PMT anode with 8 stages and gain of  $5 \times 10^3$ . Poisson fluctuations at each stage only. Added Gaussian broadening at each dynode of 50%.



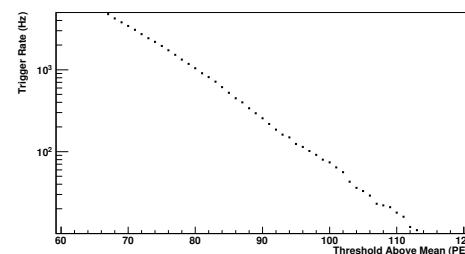
(a) FLT simulation with NSB of 2.71 pe / 100ns.



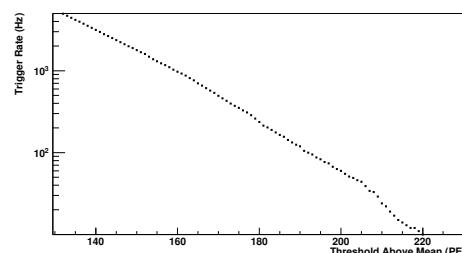
(b) FLT simulation with NSB of 6.60 pe / 100ns.



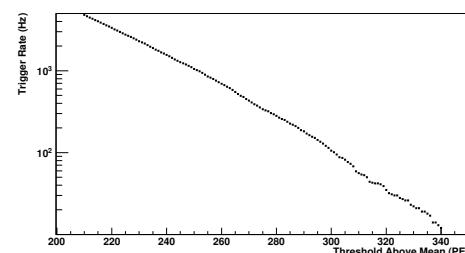
(c) FLT simulation with NSB of 46.8 pe / 100ns.



(d) FLT simulation with NSB of 65.8 pe / 100ns.



(e) FLT simulation with NSB of 263 pe / 100ns.



(f) FLT simulation with NSB of 657 pe / 100ns.



## Chapter 7

# Measuring Fluorescence Detector Photomultiplier Gain Variance with CalA Data

Measuring Gain Variance of FD PMT with CalA Data

- Measuring Gain Variance in the lab did not work. Equipment was not sensitive enough to the low current.
- There was issues with calibrating the LED light source with another PMT (QE curve and wavelength response not the same?)
- Using Low/Standard measurements of CalA to find Gain Variance Ratio
- Two different methods
- Bootstrap method to find uncertainties on Method 2

### 7.1 Motivation

- Background

- Why?

The absolute value of the gain variance cannot be found but using the CalA data a relative change can be found. This is useful for the collaboration simulations as a Gain variance is coded for the PMT at standard voltage settings. Finding out the relative change in gain variance would be useful to be used for simulations of the FD PMTs at a lower voltage settings (eg. 600V).

### 7.2 Using CalA to measure relative changes in Gain Variance

I am using CalA data from the FD telescopes to measure the relative changes in PMT gain variance as the gain was changed by a factor of 10. CalA data is calibration data used to monitor any changes in PMT gain as a function of time. CalA is performed at the beginning and end of a nightly observation while FDs are operated. Pulses of light

are piped via fibre optics cables which is pointed at each of the Fd telescope cameras. Each CalA run is a set of 50 pulses which have a width of approximately  $60 \mu\text{s}$ .

#### show image of CalA pulse

One of the values that can be calculated from the CalA is denoted  $K_V$ . The value  $K_V$  is calculated via:

$$K_V = \frac{\text{Mean}}{\text{Sigma}^2} = \frac{10}{2 \times G(1 + V_G) \times F} \quad (7.1)$$

The Mean is the average ADC count of the observed CalA pulse seen in the FD pixel. The  $\text{Sigma}^2$  is the variance calculated around a fit to the signal in  $\text{ADC}^2$ . The signal has a slope due to the effects of a capacitor used to remove the DC component of the signal. The slope is proportional to the time constant of the capacitor employed.  $G$  is the PMT gain,  $F$  is the noise equivalent bandwidth (Hz) and  $V_G$  is the PMT gain variance.

The method used to measure the ratio in gain variance is to take the calculated means and sigmas from the pulses and then finding a ratio between the  $K_V$  and Gains at the two different voltage settings.

$$K = \frac{\text{Mean}}{\text{Sigma}^2} = \frac{10}{2 \times G(1 + V_G) \times F} \quad (7.2)$$

$$\frac{(K_V)_{\text{Low}}}{(K_V)_{\text{Stand}}} = \frac{\text{Mean}_{\text{Low}}}{\text{Sigma}_{\text{Low}}^2} \div \frac{\text{Mean}_{\text{Stand}}}{\text{Sigma}_{\text{Stand}}^2} \quad (7.3)$$

$$\frac{\text{Mean}_{\text{Low}}}{\text{Sigma}_{\text{Low}}^2} \div \frac{\text{Mean}_{\text{Stand}}}{\text{Sigma}_{\text{Stand}}^2} = \frac{G_{\text{Stand}}(1 + V_G)_{\text{Stand}}}{G_{\text{Low}}(1 + V_G)_{\text{Low}}} \quad (7.4)$$

$$\frac{G_{\text{Stand}}}{G_{\text{Low}}} = \frac{\text{Mean}_{\text{Stand}}}{\text{Mean}_{\text{Low}}} \quad (7.5)$$

$$\frac{(1 + V_G)_{\text{Low}}}{(1 + V_G)_{\text{Stand}}} = \frac{\text{Sigma}_{\text{Low}}^2 \times \text{Mean}_{\text{Stand}}^2}{\text{Sigma}_{\text{Stand}}^2 \times \text{Mean}_{\text{Low}}^2} \quad (7.6)$$

## 7.3 Electronic Noise

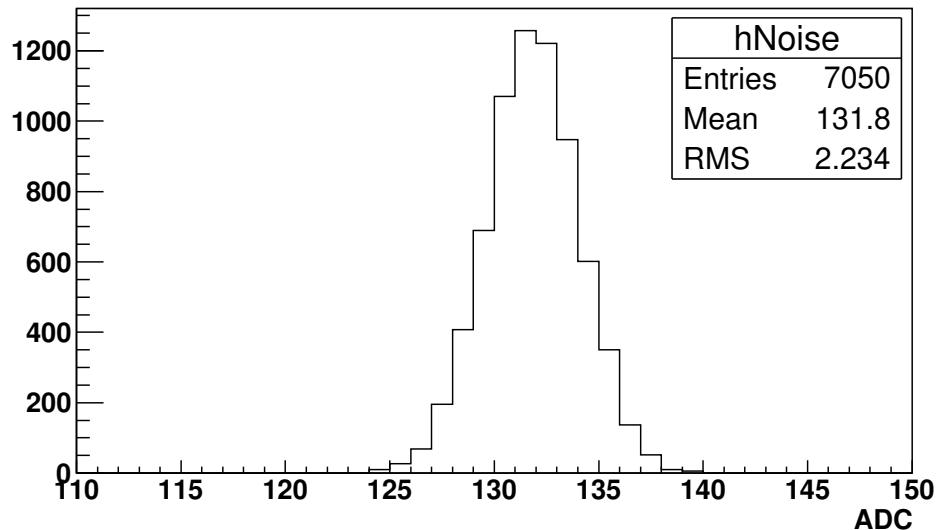
I investigated the consistency of the electronic noise across a set of 50 CalA traces. The consistency was looked at as there is only about 140 bins of noise before each of the signal pulse starts. Therefore finding an accurate mean and variance of the electronic noise on individual pulses is difficult. An accurate measurement of the electronic noise mean and variance was required as the measured change in the relative gain variance ratio was only a few percent.

#### Need to show electronic noise as function of time.

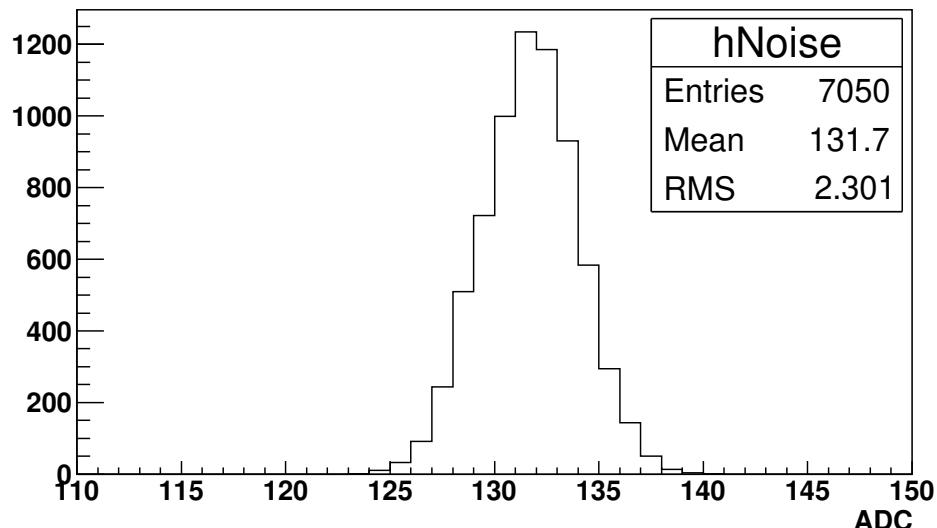
An example of a pixel electronic noise trace is shown in Fig. ???. It can be seen in this figure that stitching all of the noise bins from the 50 traces is remarkably stable. This result allows for all the noise for a single PMT pixel to be placed into a histogram to find a more accurate value for the mean and variance.

## 7.4 Pairs Method

For both standard and lower voltage settings, 50 sets of pulses are recorded for the CalA analysis. The pair method involves taking single CalA shots from standard and



(a) Standard HV



(b) Lower HV

**Figure 7.1:** Sample of the observed electronic noise observed for a single pixel within Los Leones telescope 4. Electronic noise outside of the PMT so will be the same separate from the HV setting across the PMT.

lower voltage settings and fitting an exponential to the signal. The fitted exponential is used to find the mean value at the top of the signal and the variance around the fit. For each pixel 50 values for the Gain Variance ratio was found and this was repeated for the 440 pixels within the FD telescope camera.

Fig. 7.3 shows the average ADC found across a camera for a FD telescope. The pattern follows how the camera is illuminated by the LED pointing at it. It shows that the spot is brightest in the middle and the intensity drops off towards the edges **need to add in diagram of labelled pixel no. across a FD camera**. There are uncertainties on the averages, but are smaller than the points displayed on the figures. The variance measured in Fig. 7.5 shows an expected pattern also. The variance is proportional to the mean and is expected to follow a similar shape. This will not be exact but a good indicator of whether the variance was calculated correctly.

### 7.4.1 Results

Fig. 7.6 and Fig. 7.7 shows a demonstration of using the Pairs Method to calculate the Gain Variance Ratio. The CalA data used was taken on the night of the 11-06-2016 by Los Leones Mirror 4. The mean value in Fig. 7.6 shows a 3.7% change in the ratio which translates to a 10% change in the actual value of the PMT gain variance.

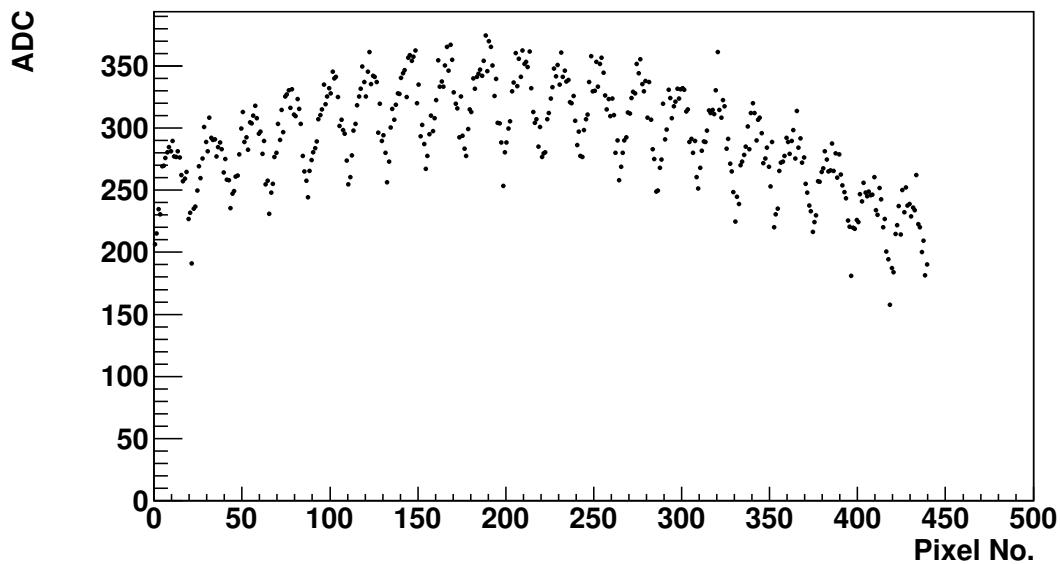
Fig. 7.7 shows the gain variance ratio measured per pixel. There can be seen no major structure or pattern. This is desired as the Gain variance is tied to individual PMTs and it would not be a great sign if the ratio followed a pattern like what was seen for the means and variances.

## 7.5 Averaging Sets of Traces Method

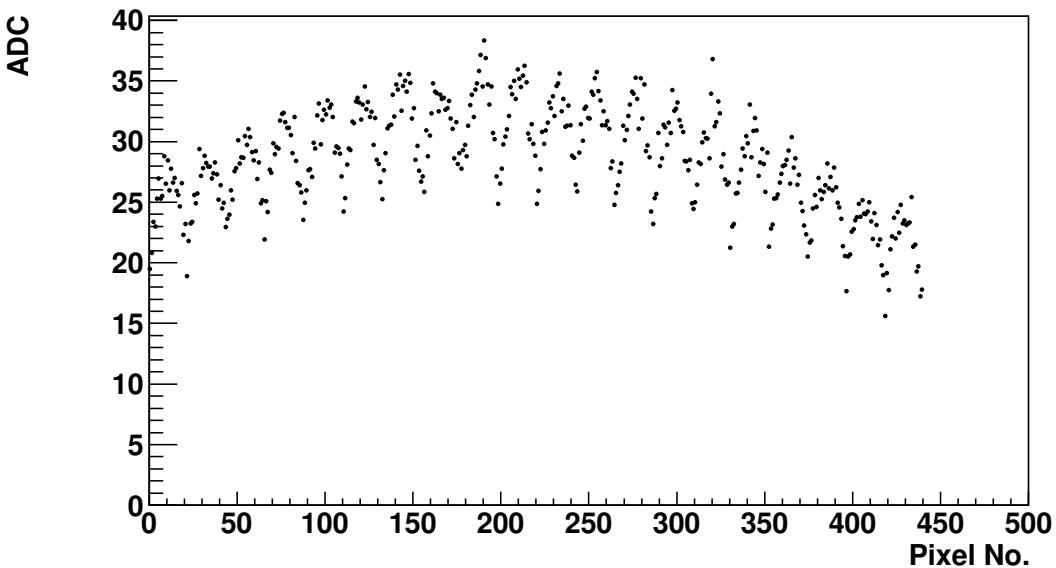
Instead of calculating the Gain Variance ratio on pairs of CalA traces then finding an average value the 50 traces for each set are stacked. This forms an average trace consisting of the 50 traces. To find the mean and variances of the noise a linear line of form  $f(x) = x$  is fitted to the first 140 bins. The fitted  $x$  is used as the mean and the variance is calculated around this value. Next an exponential is fitted to the signal. The value of the fit at the top of the signal is used as the mean while the variance is calculated around the fitted exponential.

### 7.5.1 Results

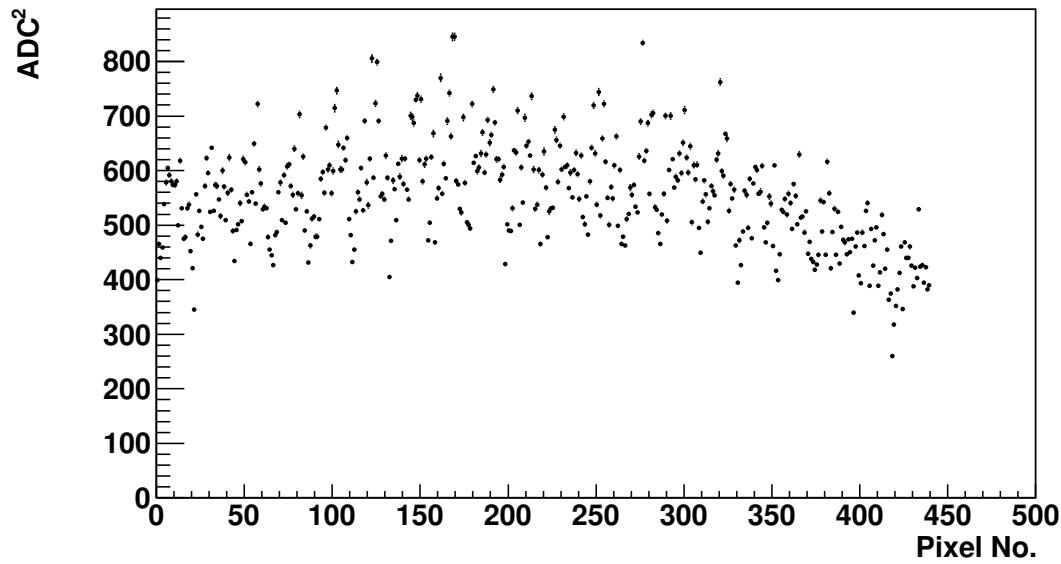
There is a larger spread in the histogram of calculated Gain Variance ratio than when the Pairs Method was employed. One good property is that no visible structure or pattern can be seen. There are definitely some extreme values of the ratio measured. To see what could cause the large outliers I had a look at the reduced chi-square of the fitted function to the noise and signal. There was no obvious problems when looking at the chi-squares as a function of pixel number for the signal. Examples are shown in Fig. 7.14 and Fig. 7.15. Examining the reduced chi-square of the fits to the noise it was seen that there were some large values. This was shown in Fig. 7.16 and Fig. 7.17. What was causing these weird fits was one or more of the noise bins having a larger deviation than expected from the measured average. **Add in plot showing off this behaviour**. This led to the process of investigating whether Least Trimmed Squares would be useful.



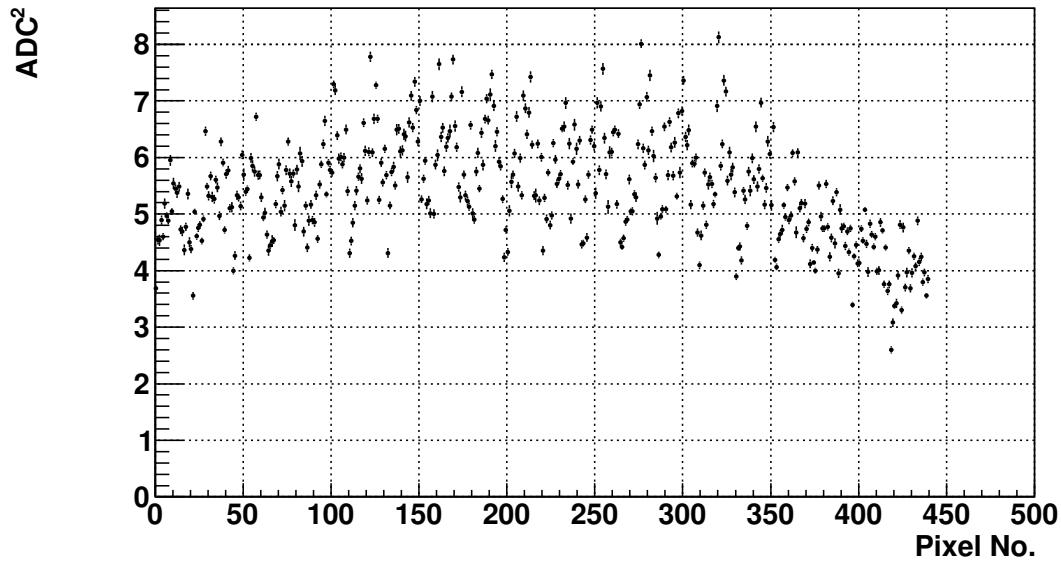
**Figure 7.2:** Mean measured at Standard HV for Los Leones Mirror 4. CalA data taken on the 11-06-2016.



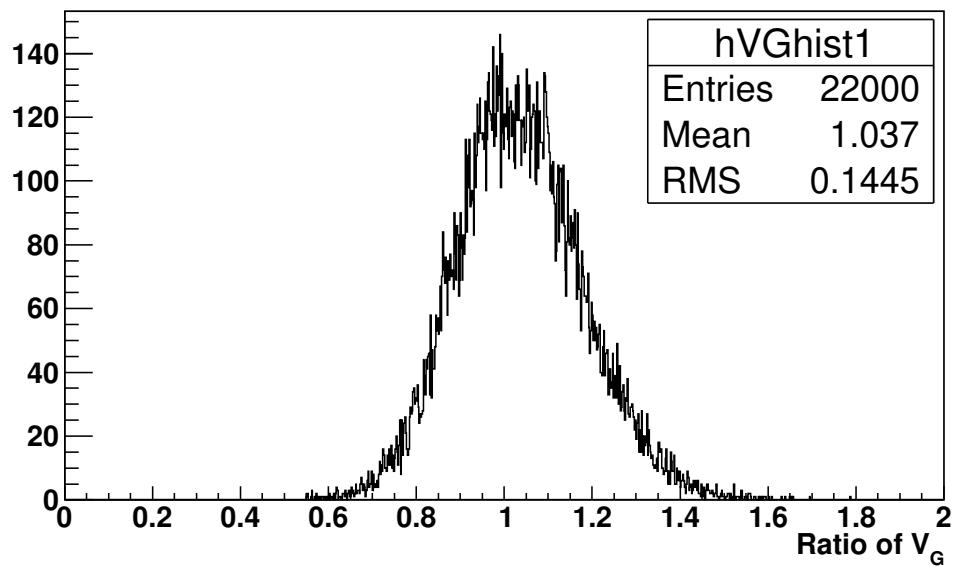
**Figure 7.3:** Mean measured at Lower HV for Los Leones Mirror 4. CalA data taken on the 11-06-2016.



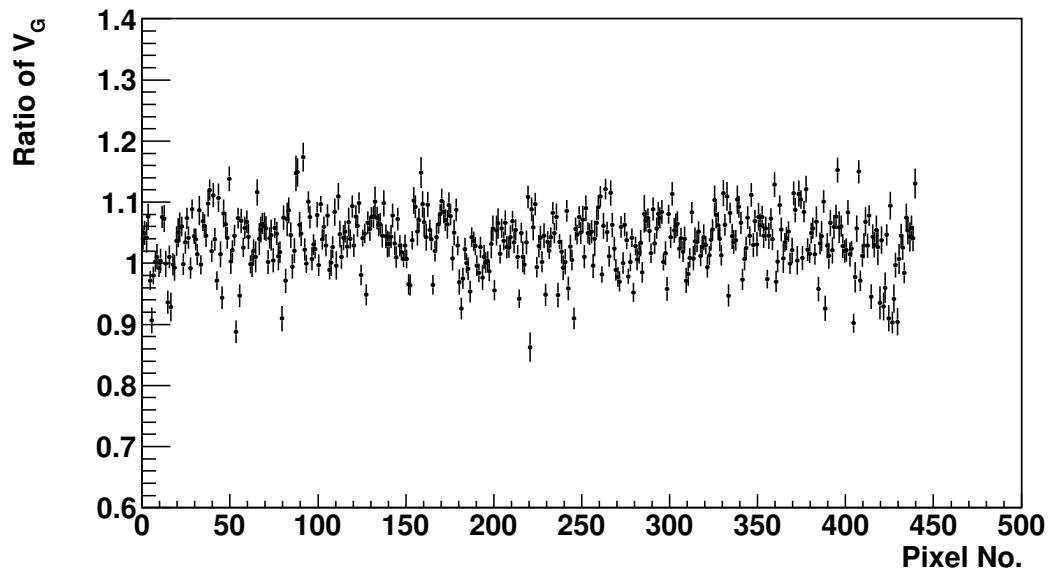
**Figure 7.4:** Variance measured at Standard HV for Los Leones Mirror 4. CalA data taken on the 11-06-2016.



**Figure 7.5:** Varaince measured at Lower HV for Los Leones Mirror 4. CalA data taken on the 11-06-2016.



**Figure 7.6:** Histogram of the all the pairs methods for Los Leones Mirror 4. CalA data was taken on the 11/06/2016 at both Standard and Lower gain settings. There are 50 traces for each of the 440 pixels recorded at both gain settings.



**Figure 7.7**

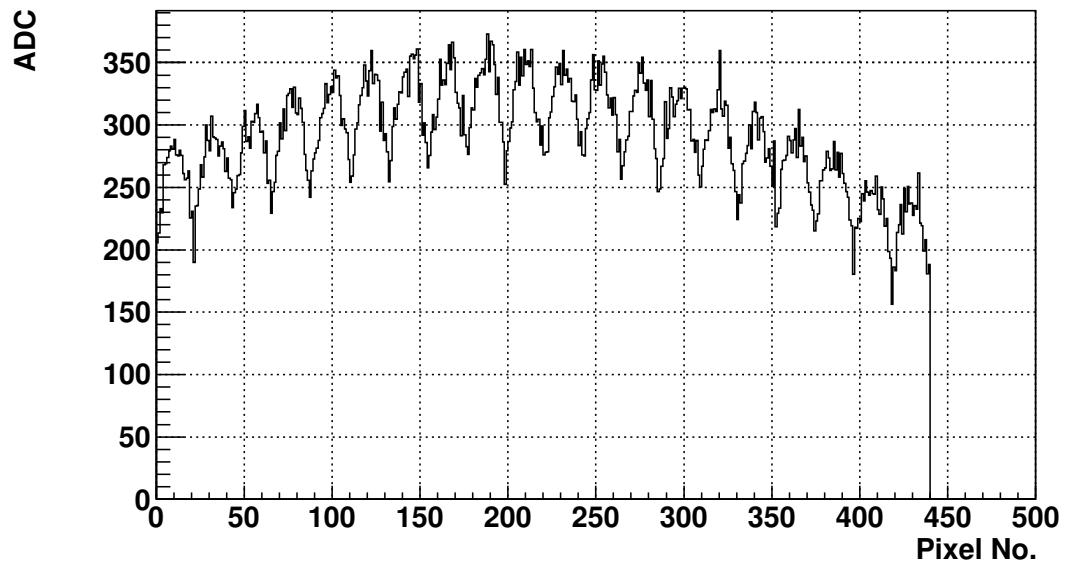


Figure 7.8

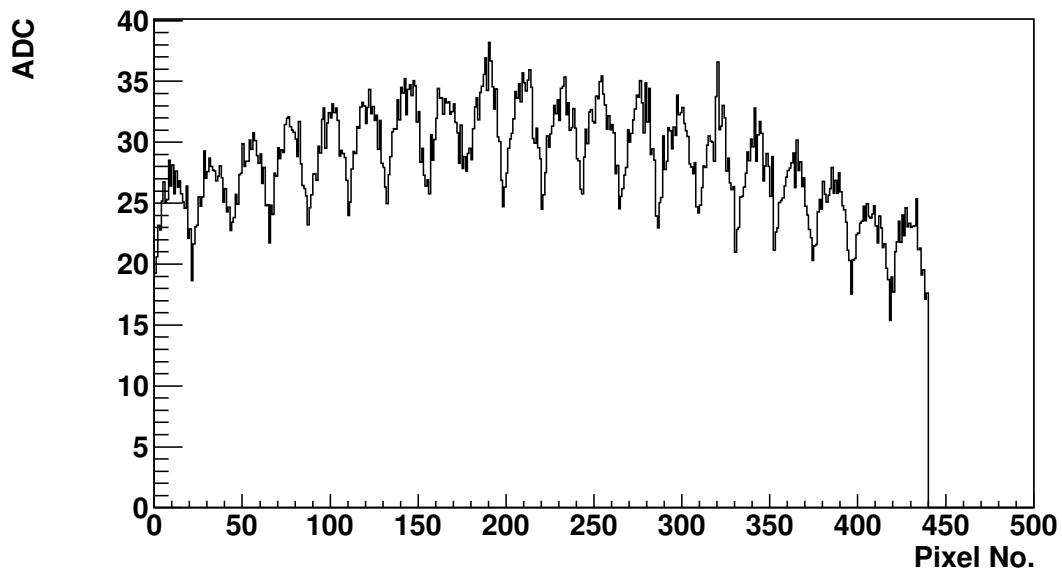


Figure 7.9

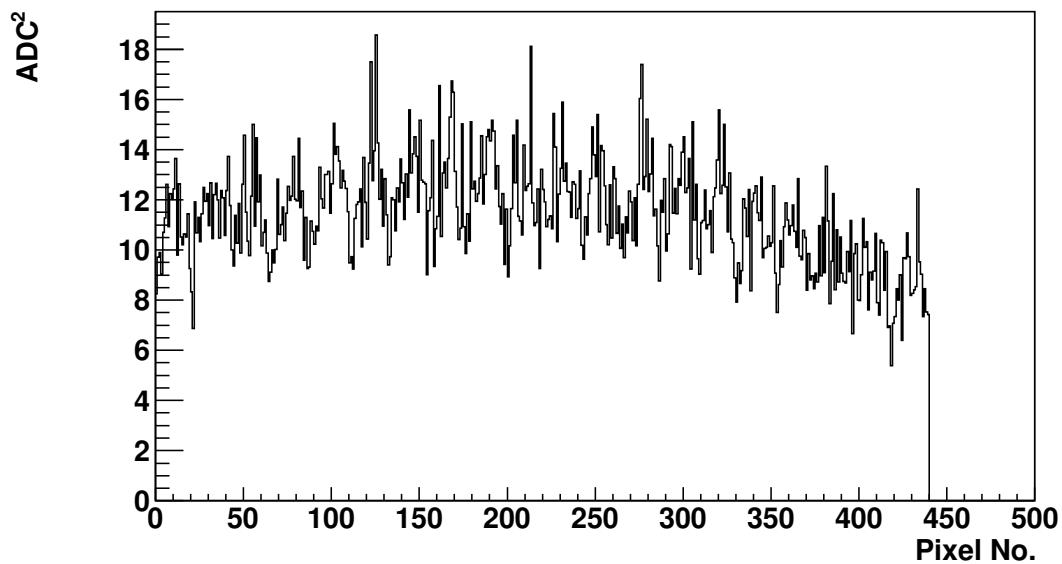


Figure 7.10

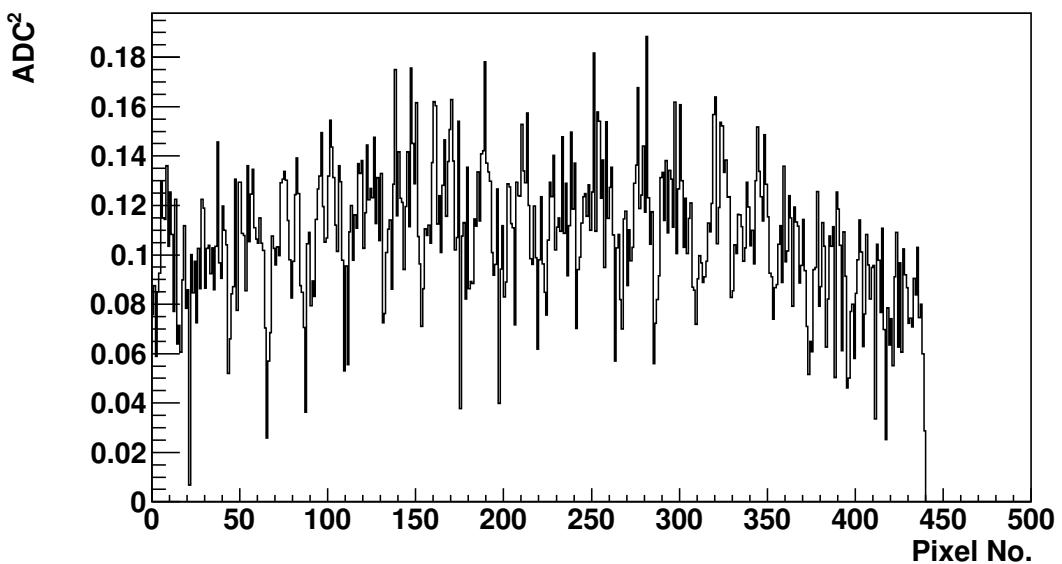


Figure 7.11

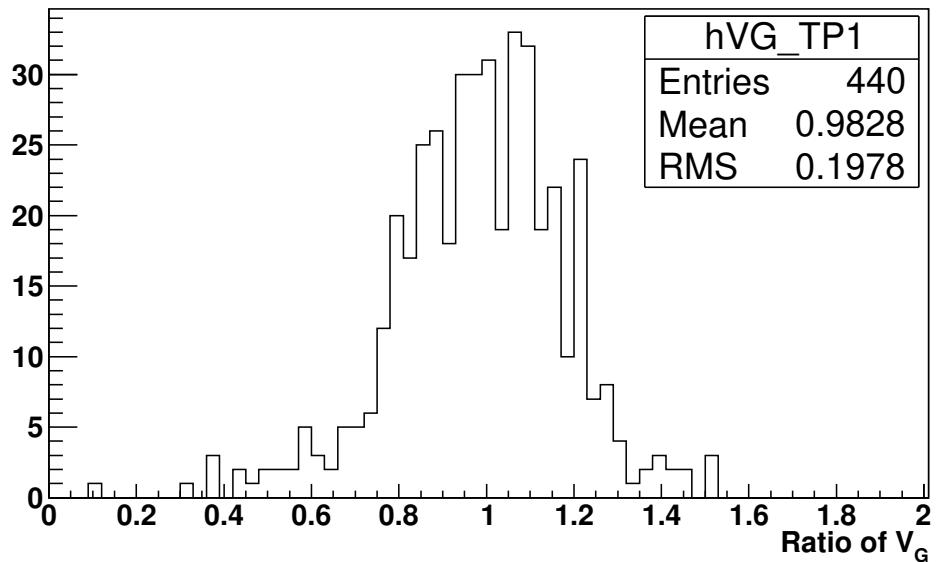


Figure 7.12

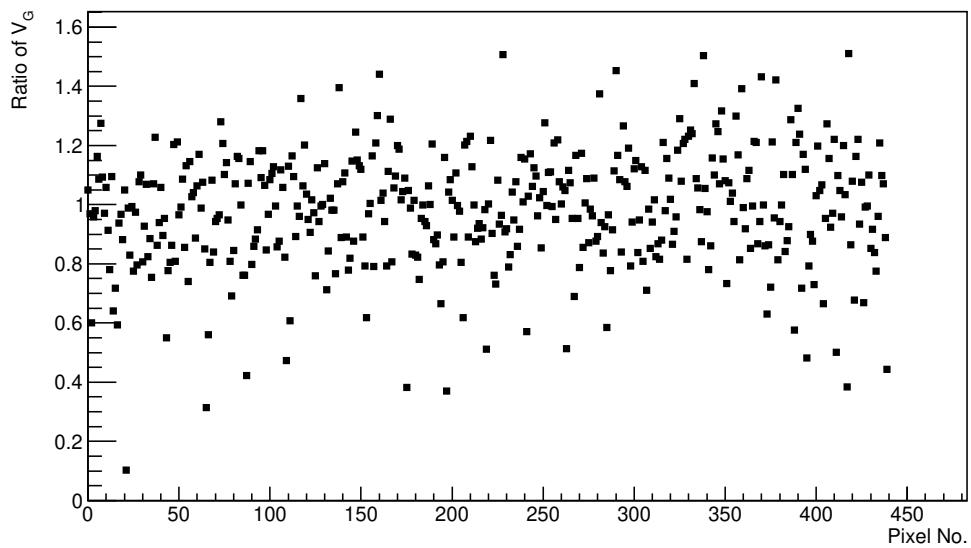
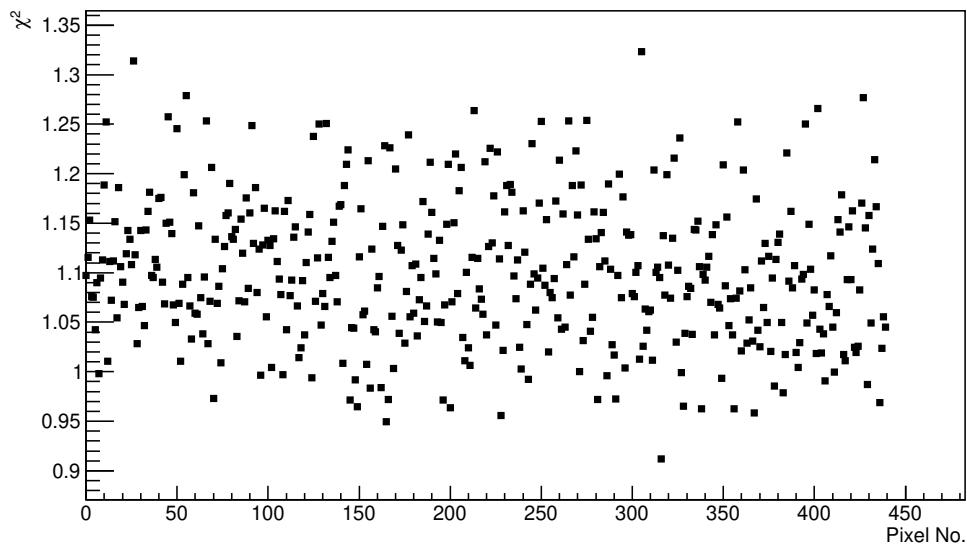
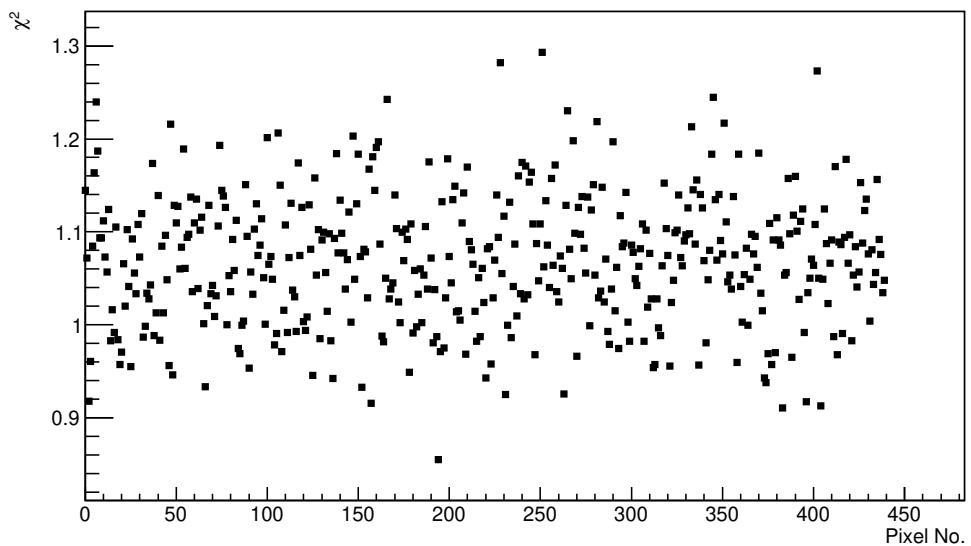


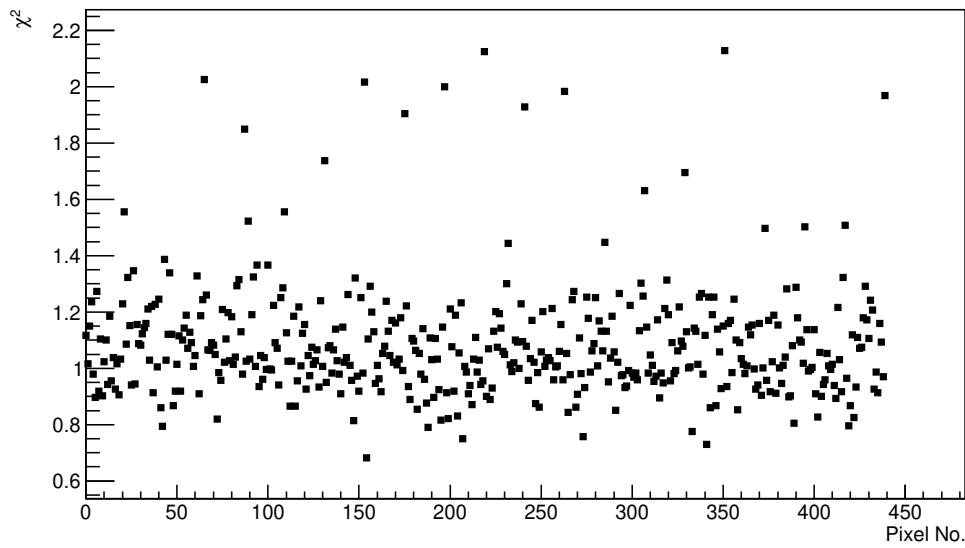
Figure 7.13



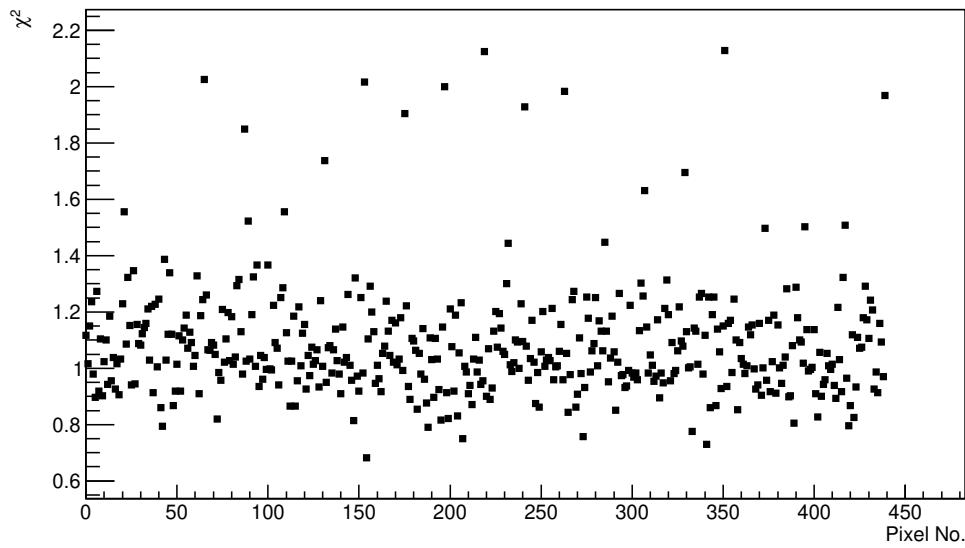
**Figure 7.14:** Reduced Chi-square for fitted exponential on the signal for CalA events measured at Standard HV.



**Figure 7.15:** Reduced Chi-square for fitted exponential on the signal for CalA events measured at Lower HV.



**Figure 7.16:** Reduced Chi-square for fitted line on the signal for CalA events measured at Standard HV.



**Figure 7.17:** Reduced Chi-square for fitted line on the signal for CalA events measured at Lower HV.

## **7.6 Result of Averaging Sets of Traces Method with Least Trimmed Squares**

Least Trimmed Square is the method that involves removing points that have the greatest sigma away from an initial fit. A point is removed one at a time with the fit repeated and the reduced chi-square checked. This method is repeated until the reduced chi-square is below a threshold.

## **7.7 Result of Averaging Sets of Traces Method using Noise Distribution**

## **7.8 Attempts to measure Gain Variance directly in the Lab**

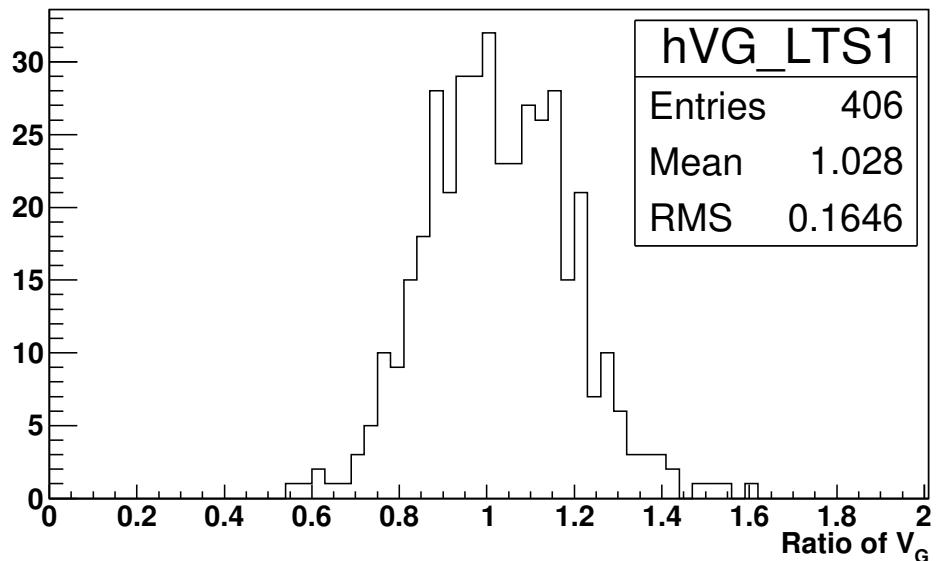


Figure 7.18

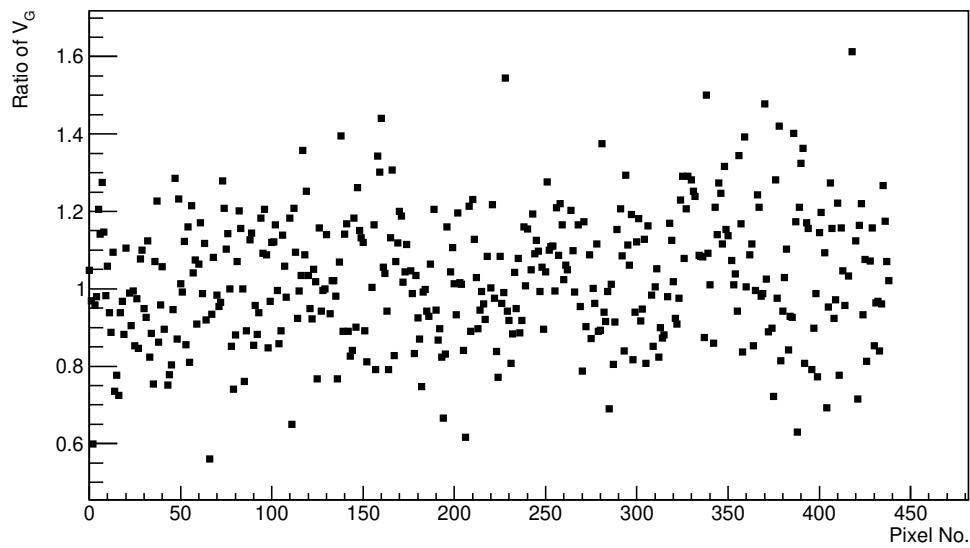


Figure 7.19

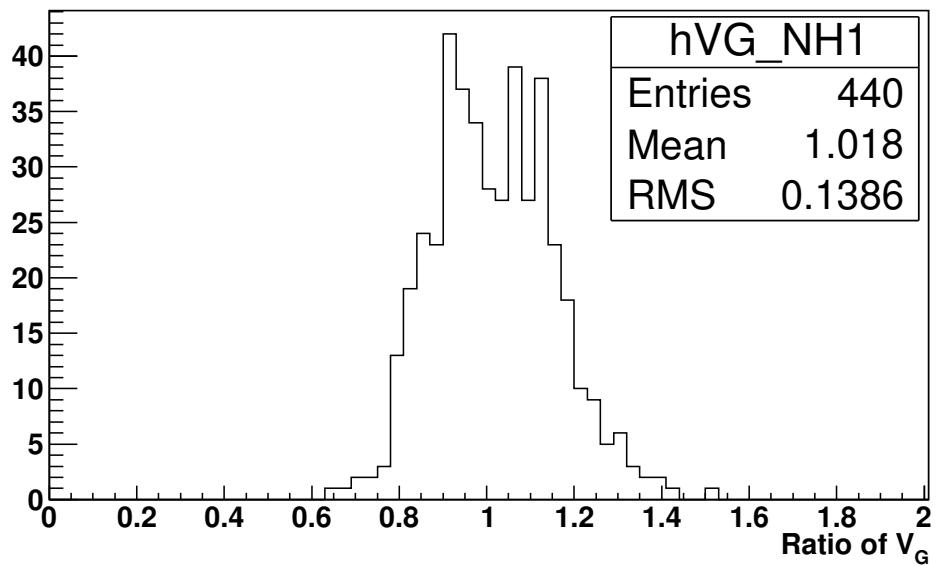


Figure 7.20

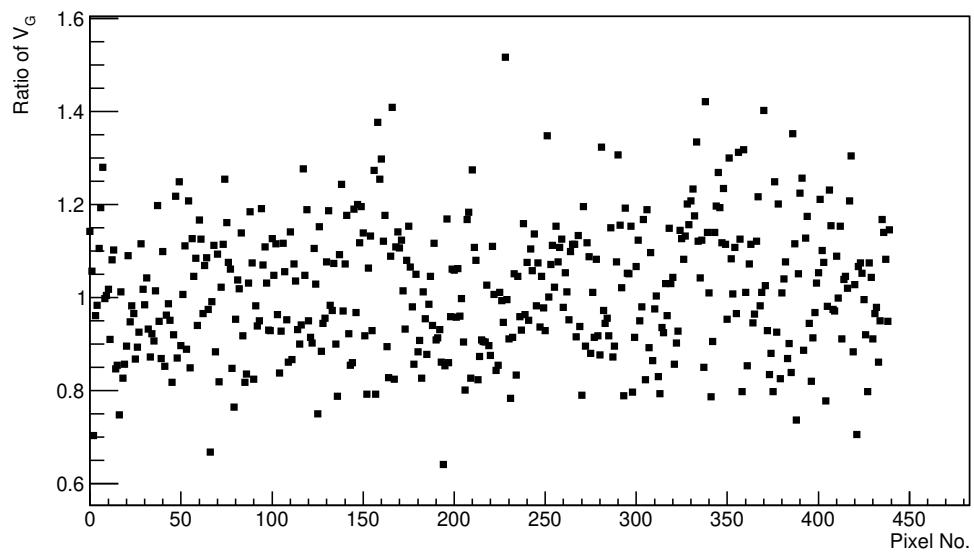


Figure 7.21



## Chapter 8

# Laboratory Simulation of FD shift

Laboratory Simulation of FD shift under differing NSB levels.

- Measurements for both 900V and 600V (900V used as baseline)
- Different length shifts
- Changing NSB
- measuring how the relevant Gain changes throughout a run



## Chapter 9

# Evaluation of Cloud Camera Cuts

First look into the effectiveness of the Cloud Camera cuts on PAO Golen Hybrid data.

- Are we being too conservative?
- Effects on Xmax, Zenith and Rp distributions



# **Chapter 10**

## **Conclusion**

### **10.1 Future Work**



# Bibliography

- [1] W. Heitler, *The quantum theory of radiation*. Oxford: (Clarendon Press, Ed. 3, Oxford), 1954.