

# Abstract

Going to make LIGO the best possible ever.

# Adaptive Mode Matching in Advanced LIGO

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Spring 2017

# Preface

The era of gravitational waves astronomy was ushered in by the LIGO (Laser Interferometer Gravitational-Wave Observatory) collaboration with the detection of a binary black hole collision (Detection paper). The event that shook the foundation of space-time allowed mankind to view the cosmos in a way that had never been done previously.

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# Chapter 1

## Introduction

Say something profound here.

Structure of this thesis:

Gravitational waves and their detection

The LIGO instrument and Noise + Squeezed States of Light

Introduction to Wavefront Sensing

Experimental Mode Matching at Syracuse

Mode matching at LIGO Hanford

Future Works

### 1.1 Gravitational Waves

In 1915, Albert Einstein published his theory of general relativity [15], which contains the most complete description of gravity to date. The seminal equation in this theory, called the Einstein Field Equation is

$$G_{\mu\nu} = 8\pi T_{\mu\nu} \tag{1.1}$$

which is a set of 10 coupled second-order differential equations that are nonlinear and fully describes the interaction between space-time and mass-energy. Equation 1.1 is difficult to solve except in situations where specific approximations allow a user to find exact solutions such as spherical symmetry [7] [36]. In areas where the curvature is close to flat, the weak field approximation can be applied and the metric is described as

$$g_{\mu\nu} \cong \eta_{\mu\nu} + h_{\mu\nu} \quad (1.2)$$

where  $\eta_{\mu\nu}$  is the metric of flat space time and  $|h_{\mu\nu}| \ll 1$  is the perturbation due to a gravitational field.

By plugging in equation 1.2 into 1.1 and using empty space we obtain the familiar wave equation

$$\left(\nabla^2 - \frac{1}{c^2} \frac{\partial^2}{\partial t^2}\right) h_{\mu\nu} = 0 \quad (1.3)$$

which has a plane-wave solution of the form  $h_{\mu\nu} = A_{\mu\nu} e^{ik_\nu x^\nu}$ .

Using the gauge constraint  $h_{,\nu}^{\mu\nu} = 0$ , it follows that  $A_{\mu\nu} k^\mu = 0$  which means that the gravitational wave amplitude is orthogonal to the propagation vector.

Further imposing transverse-traceless gauge and assuming that the wave is traveling in the  $x^3$  direction, it can be shown that the complex amplitude has physical significance expressed in the matrix

$$A_{\mu\nu} = \begin{pmatrix} 0 & 0 & 0 & 0 \\ 0 & A_{xx} & A_{yx} & 0 \\ 0 & A_{xy} & -A_{xx} & 0 \\ 0 & 0 & 0 & 0 \end{pmatrix} \quad (1.4)$$

Oftentimes, the four non-zero components of equation 1.4 can be categorized into two distinct polarizations called plus and cross such that  $h_+ = A_{xx} = -A_{yy}$  and  $h_\times = A_{xy} = A_{yx}$ .

It is natural to attempt to understand the physical interpretation of equation 1.4 as an affect on the position of a free floating particle. Consider the four-velocity,  $U^\alpha$ , in the transverse traceless gauge where the coordinate itself is attached the particles and incorporates any small wiggles that would shake the coordinates. Of course, any free particles will follow the geodesic equation

$$\nabla U^\alpha = \frac{d}{d\tau} U^\alpha + \Gamma_{\mu\nu}^\alpha U^\mu U^\nu = 0 \quad (1.5)$$

where  $\Gamma_{\mu\nu}^\alpha = \frac{1}{2} g^{\gamma\alpha} (g_{\gamma\mu,\nu} + g_{\gamma\nu,\mu} - g_{\mu\nu,\gamma})$  are the famous Christoffel symbols. By evaluating the first term of the acceleration in equation 1.5,

$$\left(\frac{dU^\alpha}{d\tau}\right)_0 = -\Gamma_{00}^\alpha = \frac{1}{2} \eta_{\mu\nu} (h_{\beta 0,0} + h_{0\beta,0} + h_{00,\beta}) \quad (1.6)$$

However, comparing equation 1.4 and equation 1.6, it is clear that if the particle is initially at rest, then a moment later it is still at rest! The term "at rest" is actually used liberally here since the coordinate system



varies along with the gravitational wave.

Alternatively, one can ask if a gravitational wave passed by a pair of particles separated by length  $L$ , what would be the effect on the distance between two points? The proper distance is defined as

$$\delta l = \int g_{\mu\nu} dx^\mu dx^\nu = \int_0^L g_{xx} dx \approx |g_{xx}(x=0)|^{1/2} \approx [1 + \frac{1}{2} h_{xx}(x=0)]L \quad (1.7)$$

which shows us two very important points about the nature of gravitational waves. Firstly, the effect is very small since the length variation,  $h_{xx}$ , is a small perturbation on flat space-time. Secondly, the effect is proportional to the initial separation between the particles. This means a detector which is large will have a better chance to measure these small effects, an important point that drove the design of the Laser Interferometer Gravitational-Wave Interferometer (LIGO).

### 1.1.1 Measuring Gravitational Waves with Light

Even with the theoretical formulation of gravitational waves resolved by the 1970s [31], the possibility of detecting gravitational waves by ground-based instruments was still a controversial topic among scientists in the field [12]. During the famous Chapel Hill Conference in 1957 which included some of the great minds of the era such as Wheeler, Schwinger, and Feynman, the experimental search for gravitational waves began to take hold. One thought experiment that was proposed by Feynmann [11] where he considered a bead sliding on a string with some friction as a gravitational wave passes by. As the beads slide back and forth due to the wave described by equation 1.7, there will be some heat dissipated which means the gravitational wave must carry some energy.

There is still the issue of an incredible amount of accuracy required to measure the strain from even the most dense astrophysical objects. The earliest attempts at detecting these small signals were famously done by Joseph Weber using large resonant bars and piezoelectric transducers to extract the energy from gravitational waves at the resonant frequencies of the bars. Picture of bar detectors at LHO. However, these bars are limited by thermal noise and can only detect GWs in very narrow frequency bands.

Interferometers are devices that measure small displacements by using a laser that is split by a partially transmitting mirror (or beamsplitter), which allows 50% of the light to get reflected and 50% to be transmitted. Each of the beams travel down the arms and reflect off of mirrors and return to the beamsplitter. Upon reaching the optic, the two beams recombine and by the principle of superposition the electromagnetic waves will add linearly at the output port (or antisymmetric port). Figure[ ]. The laser beams will gather phase as they propagate down each individual arm, and when recombining, the intensity of the light will be proportional to the phase differences between each beam. This will correspond to a differential length that

is described by

$$L_- = l_x - l_y \quad (1.8)$$

As shown in equation 1.7, the effect of gravitational waves on the proper length between two free falling objects is proportional to the initial separation. From figure[GWparticles], it is intuitively clear that interferometry would be an ideal technique to detect signals from a gravitational wave. However, one can explicitly derive a ground-based interferometer's response to a GW from an astrophysical object.

Consider a gravitational wave source arbitrarily located in the sky with respect to an interferometer on Earth. By denoting the interferometer's Cartesian coordinates as  $\{\hat{x}, \hat{y}, \hat{z}\}$  with x and y located along the arms respectively such that the z-axis points directly towards the zenith (i.e. a right-handed system.). Using the well known Euler angles, a relation from the detector frame to the source frame with coordinates  $\{\hat{x}', \hat{y}', \hat{z}'\}$  can be seen in Figure[Euler].

If a gravitational wave at the source has emitted GWs with plus and cross polarizations as denoted by equation 1.4, then the detector time series can be regarded as [1] [18]

$$h(t) = F_+(\theta, \phi, \psi) h_+(t) + F_\times(\theta, \phi, \psi) h_\times(t) \quad (1.9)$$

where  $F_+(\theta, \phi, \psi)$  and  $F_\times(\theta, \phi, \psi)$  are the antenna pattern functions that project the gravitational wave amplitudes onto the detector frame.

$$F_+(\theta, \phi, \psi) = -\frac{1}{2}[1 + \cos^2(\theta)]\cos(2\phi)\cos(2\psi) - \cos(\theta)\sin(2\phi)\sin(2\psi) \quad (1.10)$$

$$F_\times(\theta, \phi, \psi) = +\frac{1}{2}[1 + \cos^2(\theta)]\sin(2\phi)\cos(2\psi) - \cos(\theta)\sin(2\phi)\sin(2\psi) \quad (1.11)$$

If the gravitational wave is located directly above the interferometer such that  $\theta = 0$  and setting  $\psi = 0$ , then the magnitude of the antenna pattern is equal to unity. Furthermore, by rotating about the  $\phi$  angle such that the detector arms align with the plus polarization, the null geodesic equation (ie the path of a photon) in the interferometer frame becomes

$$ds^2 = g_{\mu\nu}dx^\mu dx^\nu = -dt^2 + [1 + h_+]dx^2 + [1 - h_+]dy^2 + dz^2 = 0 \quad (1.12)$$

Now, if the photon is traveling along the x-arm, this means that  $dy^2 = dz^2 = 0$  and the metric equation transforms to

$$\frac{dt}{dx} = \sqrt{1 + h_+} \approx 1 + \frac{1}{2}h_+ \quad (1.13)$$

The amount of time required for the photon to reach the x-end mirror (starting at  $t = 0$ ) is equal to

$$t_1 = \int_0^{L_x} [1 + \frac{1}{2}h_+(x)]dx \quad (1.14)$$

where  $L_x$  is the total length of the x-arm. Upon returning to the beamsplitter, the photon's total time of flight for the x and y arms are, respectively,

$$t_2 = 2L_x + \frac{1}{2} \int_0^{L_x} [h_+(x) + h_+(x + L_x)]dx \quad (1.15)$$

$$t'_2 = 2L_y - \frac{1}{2} \int_0^{L_y} [h_+(y) + h_+(y + L_y)]dy \quad (1.16)$$

If the gravitational wave period is much longer than the time of flight, then  $h_+$  does not change much during the measurement, which means  $h_+(\eta_i) \approx h_+(\eta_i + L_{\eta_i}) \approx \text{constant}$ . By subtracting the flight times of the photons for each arm and setting  $L = L_x = L_y$ , the difference is proportional to the gravitational wave perturbation multiplied by the sum of arm lengths (with a factor of  $c$  to get the units right),

$$\Delta t = t_2 - t'_2 = \frac{2L}{c}h_+ \quad (1.17)$$

By recasting the expression for time of flight in terms of the phase picked up laser light as it travels through space, the differential phase shift is

$$\Delta\Phi = \Phi(t_2) - \Phi(t'_2) = \frac{4\pi}{\lambda} h_+ L \quad (1.18)$$

The equation above is simple, however, it only works for gravitational wave signals that are not frequency dependent and it assumes that the path length can be arbitrarily long. Both points are actually not true [33] but we can alleviate these discrepancies by considering a gravitational wave signal of the form  $h(t) = h_0 \exp(i2\pi f_{GW}t)$  and repeating the calculation between equations 1.14 and 1.18:

$$\Delta\Phi(t) = h(t) \tau_{RT} \frac{2\pi c}{\lambda} \text{sinc}(f_{GW}\tau_{RT}) e^{i\pi f_{GW}\tau_{RT}} \quad (1.19)$$

where  $\tau_{RT} = 2L/c$ . The response for a detector whose arm lengths are blank km have null points at the around blank Hz which means the instrument cannot be arbitrarily long, however, this point is not concerning because it is too expensive and difficult to make a terrestrial detector of this size.

How to practically measure  $\Delta\Phi$  with an interferometer is explained in section 1.2.

### 1.1.2 Detection of Gravitational Waves

The purpose of LIGO is to observe gravitational waves emanating from astrophysical objects [16], so it is natural to wonder how well a single detector can probe the universe. It is worthwhile to explicitly derive the inspiral horizon distance, which is how far a single detector can see a source comprised two equal mass compact objects optimally oriented in the sky relative to the detector. Signal-to-noise (SNR) is the level of interesting signals compared to the background noise, which can be expressed as

$$\rho = 2 \int_{-\infty}^{\infty} \frac{\tilde{s}(f)\tilde{s}^*(f)}{S_n(f)} df \quad (1.20)$$

where  $S_n(f)$  is the one-sided average power spectral density of the detector noise and  $\tilde{s}(f)$  is the Fourier transform of the detectors' response to a gravitational-wave signal.

Consider two dense objects with mass  $m_1$  and  $m_2$  rotating around each other, separated by a distance  $R$  such that quadrupole equation still holds (i.e. flat space-time with small perturbations). The detector response, in units of strain, to the gravitational waves emitted by this source is [18],

$$s(t) = \frac{\mathcal{M}}{d_L} \Theta(\theta, \phi, \psi, i) [\pi f(t) \mathcal{M}]^{2/3} \cos[\Phi(t) + \text{constant}] \quad (1.21)$$

where

$$\mathcal{M}\& = (1+z) \frac{(m_1 m_2)^{3/5}}{(m_1 + m_2)^{1/5}} \quad (1.22a)$$

$$\Theta(\theta, \phi, \psi, i)\& = 2\sqrt{[F_+(\theta, \phi, \psi)(1 + \cos^2(i))]^2 + [2F_\times \cos(i)]^2} \quad (1.22b)$$

By convention,  $\mathcal{M}$  called the chirp mass and  $d_L$  is the luminosity distance. The orientation response,  $\Theta(\theta, \phi, \psi, i)$ , is a function that depends entirely on the detector orientation relative to the angular momentum vector of the binary [Figure of a binary spiraling and a detector on earth]

where  $F_+$  and  $F_\times$  are from equation 1.10 and 1.11.

As the binary loses energy to gravitational waves, the orbit will shrink as a function of time, in turn, the orbital frequency will increase

$$f(t) = \frac{1}{\pi \mathcal{M}} \left[ \frac{5}{256} \frac{\mathcal{M}}{T - t} \right]^{3/8} \quad (1.23)$$

By defining the binary phase as  $f(t) = \frac{1}{2\pi} \frac{\partial \Phi}{\partial t}$ , it is easy to recognize that the phase also evolves with

time.

$$\Phi(t) = 2\pi \int_T^t f(t)dt = -2 \left( \frac{T-t}{5\mathcal{M}} \right)^{5/8} \quad (1.24)$$

What this means is that the gravitational wave signal from the binary will increase in amplitude and frequency as a function of time up until the coalescence time,  $T$ .

There is actually a subtle difference between the coalescence time and the moment when adiabatic approximations fail which allows usage of the quadrupole formulation.

Plugging equation 1.21 into 1.20,

$$\rho = 8\Theta(\theta, \phi, \psi, i) \frac{r_0}{d_L} \left( \frac{\mathcal{M}}{1.2M_\odot} \right)^{5/6} \zeta(f_{max}) \quad (1.25)$$

There are two important functions in the equation above that reflect the detector's performance in sensing gravitational radiation from a binary inspiral,  $r_0$  and  $\zeta(f_{max})$ .

The characteristic distance,  $r_0$  is how far the detector can see for a fixed binary's mass distribution,

$$r_0^2 = \frac{5}{192\pi^{4/3}} \left( \frac{3M_\odot}{20} \right)^{5/3} \int_0^\infty \frac{1}{S_n(f)} \frac{df}{f^{7/3}} \quad (1.26)$$

Due to the integrand's dependence on  $f^{-7/3}$ , lower frequency improvements in the detector's noise spectrum will contribute more to the distance.

$$\zeta(f_{max}) = \frac{\int_0^{2f_{max}} \frac{1}{S_n(f)} \frac{df}{f^{7/3}}}{\int_0^\infty \frac{1}{S_n(f)} \frac{df}{f^{7/3}}} \quad (1.27)$$

The sensitivity can be different for various mass distributions;  $\zeta(f_{max})$  is a normalized function that describes how well the detector's bandwidth overlaps the binary's frequency evolution. If  $f_{max}$  is higher than the minimal sensitivity frequency for the detector, then there is good overlap and  $\zeta(f_{max}) \approx 1$ . However, if the masses are sufficiently large,  $f_{max}$  will be lower because coalescence occurs before the objects reach a higher frequency regime and this will result in  $\zeta(f_{max}) \approx 0$ .

The process of two objects coalescing starts at lower frequencies and terminate at some  $f_{max}$  that depends on when the objects reach their inner-most stable circular orbit,  $f_{ISCO}$ . If  $m_1 = m_2$ , then the maximum frequency is[]

$$\begin{aligned} f_{max} &= \frac{f_{ISCO}}{1+z} \\ &= \frac{99\text{Hz}}{1+z} \frac{20M_\odot}{M} \end{aligned} \quad (1.28)$$

where  $M = m_1 + m_2$ .

By plugging a few mass distributions and an estimate of the sensitivity [GWINC], it is possible to study the range of a single detector. For example, a 1.4-1.4 $M_{\odot}$  binary that has optimal orientation with respect to the detector has horizon distance equal to...

Horizon distance and range, advanced LIGO projected

[Finn, Duncan Thesis] [33]

## 1.2 The LIGO Instrument

In its simplest form, the LIGO instrument is an incredibly large Michaelson interferometer. If we imagine the output as a measure of differential arm length, it becomes a natural way of detecting gravitational waves. However, to make a practical gravitational-wave observatory, the complexity will have to be extended beyond what Michelson and Morley used. The next sections will explore the various upgrades to the instrument configuration that improve the sensitivity of LIGO.

### 1.2.1 Simple Michaelson

As shown in Figure [michelsonifo], the interferometer begins with a laser which is incident on a partially-reflecting and partially-transmitting beamsplitter that sends half of the laser light down each arm. The individual beams gather phase as they propagate down the arms and then reflect off of two end mirrors and eventually recombine at the same beamsplitter. Whether the light constructively or destructively interferes will depend on their relative accumulated phase from propagating down each arm. A photodetector is placed at the output and measures the total power which converts the total amount of light into an electronic signal.

In Section 1.1.1, it was shown that the differential time of flight between photons traveling down the individual arms carry gravitational wave information. The difference in flight times,  $\Delta t$ , can be converted into how much phase,  $\Delta\phi$ , is accumulated by the photons as they propagate through space.

But the question remains, how does an interferometer explicitly measure  $\Delta\phi$ ?

If the input electric field of the interferometer is  $E_0$ , the beamsplitter will transmit  $E_0/2$  down the x-arm and reflect  $E_0/2$  down the y-arm. By setting the beamsplitter to be the origin, the two plane waves traveling down their respective arms will have gathered a phase  $\phi_i$ . Then, upon reflecting off the end mirrors and returning to the beamsplitter, each of the electric fields can be described by these equations

$$\begin{aligned} E_x &= \frac{iE_0}{2} e^{2i\phi_x} \\ E_y &= \frac{iE_0}{2} e^{2i\phi_y} \end{aligned} \tag{1.29}$$

Since the electromagnetic waves are linear, the resultant sum of waves at the antisymmetric port will be  $E_{out} = E_x + E_y$ . The photodiode at the output (or antisymmetric) port detects the integrated power which is related to the total electric field by

$$\begin{aligned} P_{AS} &= \int_{\text{Area}} I \, dA \\ &= \int_{\text{Area}} |E_x + E_y|^2 \, dA \\ &= P_{in} \cos^2(\Delta\phi) \end{aligned} \quad (1.30)$$

where  $\Delta\phi = \phi_x - \phi_y = k_x L_x - k_y L_y$  and  $P_{in} = \int_{\text{Area}} |E_0|^2 dA$  is the input power. By using equation 1.18 and the common (or average) arm length  $L_+ = \frac{L_x + L_y}{2}$ , the power due to a differential phase shift is

$$P_{AS} \approx P_{in} (1 - 2\Delta\phi) = P_{in} (1 - 2kL_+h_+) \quad (1.31)$$

There is a large DC term that is dependent on the input power and generally, it is very difficult to measure small changes in a large signal. So the next obvious method would be to shift the arms such that the output port is operating on a dark fringe, normally this is called a null-point operation. However, there are difficulties associated with this method as well. Consider shifting the phase of equation 1.30 by  $\pi/2$ , which would result in

$$P_{AS}|_{null} = P_{in} \sin^2(\Delta\phi) \approx P_{in} (kL_+h_+)^2 \quad (1.32)$$

This results in a second-order dependence on a gravitational-wave signal that is already approximated to be very small. So a good solution to the issue of how to *read* out a gravitational wave signal can be solved using radio frequency (RF) detection methods. Consider changing the interferometer input by adding an electro-optical modulator (EOM) to sinusoidally modulate the laser frequency and expanding to first order using the Bessel functions,

$$\begin{aligned} E_{in} &= E_0 e^{i(wt + \beta \cos(\Omega t))} \\ &\approx E_0 e^{iwt} [J_0(\beta) + J_1(\beta) e^{i\Omega t} + J_1(\beta) e^{-i\Omega t}] \\ &= E_{C,in} + E_{SB+,in} + E_{SB-,in} \end{aligned} \quad (1.33)$$

where  $\Omega$  and  $\beta$  are the modulation frequency and depth, respectively. The first term is commonly called the carrier field whereas the second and third terms are referred to as the (upper or lower) sidebands. Because there are multiple electric fields, it is useful to define an optical transfer function which transforms the

interferometer's input fields to its output,

$$E_{out} = E_C + E_{SB+} + E_{SB-} = \begin{pmatrix} t_C \\ t_{SB+} \\ t_{SB-} \end{pmatrix} \begin{pmatrix} E_{C,in} & E_{SB+,in} & E_{SB-,in} \end{pmatrix} \quad (1.34)$$

The carrier transfer function,  $t_C$  has already been calculated by equations 1.30 - 1.32 and the sideband transfer functions are not much different.

$$t_{SB\pm} = r_{x,\pm} e^{i\phi_{\pm,x}} - r_{y,\pm} e^{i\phi_{\pm,y}} \quad (1.35)$$

where  $\phi_{\pm,i} = (k \pm k_\Omega) \ell_i = (\frac{w+\Omega}{c}) \ell_i$ . In the current example, the sidebands and carrier fields reflect off the end mirrors identically, however, this will not be true in general when dealing with resonators that are highly frequency dependent. Plugging equation 1.35 into 1.34, the output electric field becomes

$$E_{out} = ie^{i\omega t} [J_0(\beta) k \ell_+ h_+ + J_1(\beta) \sin(k\Delta\ell + k_\Omega\Delta\ell) (e^{i\Omega t} + e^{-i\Omega t})] \quad (1.36)$$

The frequency offset of the sidebands allow them to pick up phase differently than the carrier field. Thus, by choosing a static offset between the arm lengths such that the carrier signal is on a dark fringe,  $k\Delta\ell = \pi/2$ , but the sideband are slightly off the null point (Figure []) and hence leak into the anti-symmetric port. This technique is colloquially known as the Schnupp Asymmetry and the electric field reduces to

$$E_{out} = ie^{i\omega t} [J_0(\beta) k \ell_+ h_+ + J_1(\beta) \sin(k_\Omega\Delta\ell) (e^{i\Omega t} + e^{-i\Omega t})] \quad (1.37)$$

Recall that the intensity is equal to the electric field squared,

$$\begin{aligned} I = |E_{out}|^2 &= |E_C|^2 + |E_{SB+}|^2 + |E_{SB-}|^2 \\ &+ 2\text{Re}\{ E_{SB+} E_{SB-}^* e^{2i\Omega t} \} \\ &+ 2\text{Re}\{ (E_C E_{SB-}^* + E_{SB+} E_C^*) e^{i\Omega t} \} \end{aligned} \quad (1.38)$$

The last term is referred to as the *beat note* between the carrier signal and the sidebands. It is possible to extract the term at the modulation frequency using a mixer which is an analog device that outputs the product of two inputs. Usually, the same oscillator that was used to modulate the input beam can be one



of the mixer inputs,  $\cos(\Omega t)$ , so that the demodulated signal is

$$\begin{aligned} I_{Demod} &\propto [4\pi J_0(\beta)J_1(\beta)\frac{\ell}{\lambda}\sin(k_\Omega\Delta\ell)h_+][\cos(\Omega t)\sin(\Omega t + \phi_{Demod})] \\ &\propto [4\pi J_0(\beta)J_1(\beta)\frac{\ell}{\lambda}\sin(k_\Omega\Delta\ell)h_+][\sin(\phi_{Demod}) + \sin(2\Omega t + \phi_{Demod})] \end{aligned} \quad (1.39)$$

where  $\phi_{Demod}$  is the phase that can be set by the user in order to account for extra phase shifts (ie. longer cables). After the mixer, there will be signals at DC,  $\Omega$ ,  $2\Omega$  and so on. However, the part that is linear in the gravitational wave amplitude will be at DC so a low-pass filter will allow the final signal to dominate:

$$S = 4\pi J_0(\beta)J_1(\beta)\frac{\ell}{\lambda}\sin(k_\Omega\Delta\ell)\sin(\phi_{Demod})h_+ \quad (1.40)$$

This shows that a RF detection technique will be linear in GW signal with no large DC offset. Setting the carrier on a null point means  $\Delta\ell = \frac{k_\Omega}{k}\frac{\pi}{2}$  and allows the designer to optimize the Schnupp asymmetry length to get the best signal for some modulation frequency. This type of readout scheme was used in Enhanced LIGO and is called heterodyne detection, where the sideband fields are produced by an EOM and its efficacy depends on the local oscillator's stability [19]. In contrast, the Advanced LIGO scheme uses a homodyne detection [25] method called "DC-Readout". Here the oscillator field is produced by slightly offsetting the arms away from the dark fringe and letting a small amount of carrier light through the antisymmetric port. A gravitational wave will induce sidebands on the carrier and this will allow the same mathematics as above to achieve a linear signal in gravitational wave strain. This method benefits from naturally being co-aligned and mode matched with the signal field. All techniques follow the same logic of beating the field containing useful information with a reference field to extrapolate a linear signal but the differences come from technical noise such as laser intensity fluctuations and effective quantum noise.

[4]

### 1.2.2 Fabry-Perot Cavities

There are two ways to improve the LIGO detectors: one is to increase the response from gravitational waves and the other is to decrease the noise contributions. From equation 1.18, the gravitational wave signal is proportional to the optical path length that the photon travels, which means the most straightforward method of increasing the sensitivity is to make the arms as long as possible (up to the null point described by equation 1.19). There are two methods to achieve this: a Herriott delay line or a Fabry-Perot resonator, the differences between each method is shown in Figure[FPvsDelay]. At the time of writing this thesis, all modern gravitational wave detectors use the latter method.

A Fabry-Perot cavity is an optical system comprised of two or more partially transmitting mirrors with one laser input. To create a resonator, the user must design a system so that once the laser has made one round trip within the optical system, the shape and size of the beam is the same as when it started. This may seem simple conceptually but in practice, controlling and sensing any optical cavity comes with a few challenges.

To start understanding the longitudinal degree of freedom, consider a two mirror system in Figure [FabryPerotFig] which is separated by a length  $L$  with reflection and transmission coefficients:  $r_1$ ,  $t_1$ ,  $r_2$ ,  $t_2$ . Starting with a plane wave at the input mirror with amplitude  $E_0$ , the beam will enter the cavity and propagate back between the two mirrors. The reflected, transmitted, and circulating fields [33], which are a result of the geometric series, can be described as

$$E_{REFL} = r_{FP}E_0 = \left( -r_1 + \frac{t_1^2 r_2 e^{-i2kL}}{1 - r_1 r_2 e^{-i2kL}} \right) E_0 \quad (1.41)$$

$$E_{TRAN} = t_{FP}E_0 = \left( \frac{t_1 t_2 e^{ikL}}{1 - r_1 r_2 e^{-i2kL}} \right) E_0 \quad (1.42)$$

$$E_{CIRC} = c_{FP}E_0 = \left( \frac{t_1}{1 - r_1 r_2 e^{-i2kL}} \right) E_0 \quad (1.43)$$

The fields above are all highly dependent on the round trip phase and become resonant when the cavity length is  $L = n\lambda/2$  and the circulating coefficient in the cavity is maximized such that the gain is

$$\text{Gain} = c_{FP}^2|_{\text{resonating}} = \left( \frac{t_1}{1 - r_1 r_2} \right)^2 \quad (1.44)$$

Depending on the relative reflection coefficients of the input and output mirrors, the fields on resonance will be slightly different Figure[UnderOverCritCoupled].

Notice that the fields are dependent on the cavity length and laser frequency,  $2kL$ , so one might naively determine that modulating the two parameters independently cause the same effect. However, when both are changing by large amounts, they are related by a frequency dependent transfer function

$$C(s) \frac{\Delta w}{w} = -\frac{\Delta L}{L} \quad (1.45)$$

where  $C(s) = \frac{1 - e^{-2sL/c}}{2sL/c}$  in the Laplace domain Figure[FreqRespFP]. Only when the cavity is on or near resonance, then the frequency and length variations are related by  $\frac{\Delta w}{w} = -\frac{\Delta L}{L}$ .

While sweeping through either laser frequency or cavity length and measuring the reflected (or transmitted) fields, there are features of the power spectrum which relate directly to the cavity's physical properties:

Finesse, or the line width of the resonant peak, function of  $r_1$  and  $r_2$ :

$$\mathbb{F} = \frac{\pi \sqrt{r_1 r_2}}{1 - r_1 r_2} \quad (1.46)$$

Storage Time :

$$\tau_s = \frac{L}{c\pi} \mathbb{F} \quad (1.47)$$

Cavity Pole:

$$f_{pole} = \frac{1}{4\pi\tau_s} \quad (1.48)$$

Free Spectral Range:

$$f_{FSR} = \frac{c}{2L} \quad (1.49)$$

Stability: In order to prove that an optical cavity is stable, it is useful to introduce the matrices that describe a periodic optical system which is explicitly derived in appendix A. A Fabry Perot cavity that is separated by distance  $L$  with spherical mirrors that have radii of curvature  $R_1$  and  $R_2$  will need to satisfy

$$0 \geq \left(1 + \frac{L}{R_1}\right) \left(1 + \frac{L}{R_2}\right) \geq 1 \quad (1.50)$$

in order to be geometrically stable (see Figure[BeamGeometryFP]).

Power circulating as a function of our defined parameters slightly off resonance by a length of  $\delta L$ :

$$P_{cav} = |c_{FP}|^2 = \frac{Gain}{1 + \left(\frac{2\mathbb{F}}{\pi}\right)^2 \sin^2(k\delta L)} \quad (1.51)$$

LIGO uses the frequency response of optical cavities for a number of reasons: Firstly, the round trip phase of the gravitational wave is amplified by the finesse (equation 1.46) of the 4 kilometer arm cavities, therefore increasing the sensitivity. Secondly, the input and output of the interferometer's Gaussian beam mode can be refined by only allowing the fundamental mode of the laser to resonate (input and output mode cleaners). Thirdly, Advanced LIGO uses a dual-recycled Michelson interferometer which means the symmetric port has a mirror inserted to resonate the light reflected from the Fabry Perot arms (power recycling) and another mirror shapes the frequency response of the differential cavity pole at the anti-symmetric port (signal recycling).

### Achieving Resonance in an Optical Cavity

Described above are the theoretical constructs of a FP cavity, but the question remains, how does one practically construct a resonant optical cavity? The answer comes from using a heterodyne sensing scheme similar to the one described in Section 1.2.1. Except the optical system is not a Michelson interferometer but rather a two mirror cavity. Still, the heterodyne detection method can apply to a number of different geometries such as triangular or bow-tie cavities shown in Figure[DiffResCavs].

Starting with an input laser and EOM (electro-optical modulator) that imparts upper and lower sidebands at a modulation frequency  $\Omega$ , the user injects three beams into the optical system described in Equation 1.33. When placing a photodetector on the reflection port, one should see the cavity's effect on each of the three electric fields.

$$E_{FP,out} = E_C + E_{SB+} + E_{SB-} = \begin{pmatrix} r_C \\ r_{SB+} \\ r_{SB-} \end{pmatrix} \begin{pmatrix} E_{C,in} & E_{SB+,in} & E_{SB-,in} \end{pmatrix} \quad (1.52)$$

where the reflection coefficients follow equation 1.41. Because sidebands are frequency shifted, they will effectively *see* a different cavity than the carrier and the total phase accumulated between the fields will be different. To make a good reference for the resonant carrier beam, the modulation frequency,  $k_\Omega$ , is chosen such that sidebands are be anti-resonant in the cavity.

$$r_C = -r_1 + \frac{t_1^2 r_2 e^{-i2kL}}{1 - r_1 r_2 e^{-i2kL}} \quad (1.53)$$

and

$$r_{SB\pm} = -r_1 + \frac{t_1^2 r_2 e^{-i2(k+k_\Omega)L}}{1 - r_1 r_2 e^{-i2(k+k_\Omega)L}} \quad (1.54)$$

The formalism to read out the error signal is the shown in Equation 1.38. Using a photodetector in reflection (see Figure[FPLocking]), the error signal will be linearly proportional to the laser frequency and cavity length [4].

$$\text{Error Signal} \propto \frac{L\mathbb{F}}{\lambda} \left[ \frac{\delta w}{w} + \frac{\delta L}{L} \right] \quad (1.55)$$

### 1.2.3 Fabry-Perot Michelson

Using an optical cavity to increase the number of bounces along the 4 kilometer arms will amplify the gravitational wave signal but it will also shape the frequency dependence of the sensitivity as well. This has already been alluded to in section 1.1.1 where the overall increase in arm length will tune the response function

of equation 1.19. The same formalism can apply to integrate the use of Fabry-Perot arms into a Michelson interferometer and looking at the signal at the antisymmetric port as a function of gravitational wave frequency. Previously in section 1.2.1, it was shown that the RF detection technique required beating the sidebands with a carrier field containing the gravitational wave information. However, this is not exclusively true because the signal sidebands beating with a reference carrier field could also be a valid way of measuring a signal of interest. In fact, Advanced LIGO employs a homodyne detection called DC readout [19]. The idea is that the carrier field will resonant in the arms but will be slightly modulated due to a gravitational wave. In effect, the carrier will have audio frequency sidebands imparted on the field shown in Figure[gwsb].

Figure [FPMichelson] shows an input electric field of  $E_0$  incident on a 50/50 beamsplitter which then enters two orthogonal arm cavities of length  $L_x$  and  $L_y$ . Notice that the distance from the beamsplitter to the input couplers for both arms denoted by  $l_x$  and  $l_y$  are independent variables, this is to account for the DARM offset which is analogous to the Schnupp asymmetry described in section 1.2.1. After partially transmitting and reflecting at the beamsplitter, the fields will resonate in each of the arm cavities and obey the reflectivity relations described by equation 1.41. Upon recombining at the beamsplitter, the antisymmetric port electric field will be related to the arm fields by,

$$E_{AS} = \frac{1}{\sqrt{2}}[E_x^{\text{ARM}} + E_y^{\text{ARM}}] \quad (1.56)$$

where  $E_i^{\text{ARM}} = E(\omega) + [E(+\omega_{\text{GW}}) + E(-\omega_{\text{GW}})] \sin k\delta l$  is the reflection coefficient for the x and y arms which contain the carrier and GW signal sidebands. By inserting a DC microscopic length shift,  $\delta l = l_x - l_y$ , this allows for some carrier leakage field to propagate towards the antisymmetric port which creates a static reference to create the beat note with the signal sidebands. The same effect can be achieved by introducing a static offset between the x and y arms; colloquially, this is referred to as DARM offset.

For the carrier field, the amplitude of reflection is simple because the LIGO arms are a strongly overcoupled cavity operating on resonance therefore

$$E(\omega) = \frac{E_0}{\sqrt{2}} \left( -r_1 + \frac{t_1^2 r_2}{1 - r_1 r_2} \right) \approx -\frac{E_0}{\sqrt{2}} \quad (1.57)$$

The sidebands are a different story because the modulation caused by the gravitational wave very slightly move the field off resonance. Imagine the carrier circulating in the arms and then suddenly, a disturbance due to a gravitational wave displaces the end mirror at a frequency  $\omega_{\text{GW}}$  by an amount  $\Delta L$ . By using the same formalism as the expansion of a modulated field in equation 1.33, the sideband fields can be described

by,

$$\begin{aligned} E(\pm\omega_{\text{GW}}) &= \left[ \frac{t_1}{1 - r_1 r_2 e^{-2ikL}} \right] \frac{E_0}{\sqrt{2}} \left[ \frac{t_1}{1 - r_1 r_2 e^{-2i(kL \pm \frac{\omega_{\text{GW}}}{c} \Delta L)}} \right] ik\Delta L \\ &= \left[ \frac{t_1^2}{1 - r_1 r_2} \right] \frac{E_0}{\sqrt{2}} \left[ \frac{1}{1 - r_1 r_2 e^{-2i\frac{\omega_{\text{GW}}}{c} \Delta L}} \right] ik\Delta L \end{aligned} \quad (1.58)$$

The first bracket is the circulating field of the carrier signal amplified by the Fabry Perot cavity and the second bracket is the circulating field slightly off resonance due to the gravitational wave frequency offset. Setting the fields to resonate already simplifies the equation but it is also reasonable to approximate the gravitational wave as a weak signal and expand the exponential  $e^{-2i\frac{\omega_{\text{GW}}}{c} \Delta L} \approx 1 - 2i\frac{\omega_{\text{GW}}}{c} \Delta L$ . This transforms the signal sideband field into a simple form,

$$E(\pm\omega_{\text{GW}}) \approx \left[ \frac{t_1}{1 - r_1 r_2} \right]^2 \frac{E_0}{\sqrt{2}} \left[ \frac{1}{1 \pm i\frac{\omega_{\text{GW}}}{\omega_p}} \right] ik\Delta L \quad (1.59)$$

where  $\omega_p = \frac{1-r_1 r_2}{r_1 r_2} \frac{c}{2L}$  is the differential pole frequency.

The photodiode signal at the antisymmetric port is proportional to the beat note between the carrier and signal sideband that is demodulated by a phase  $\phi_D$ ,

$$\begin{aligned} S &\propto [E(\omega)E^*(+\omega_{\text{GW}}) + E(-\omega_{\text{GW}})E^*(\omega)] \sin(k\delta l) \sin(\phi_D) \\ &\propto E_0^2 \frac{2\pi L}{\lambda} \left[ \frac{t_1}{1 - r_1 r_2} \right]^2 \left[ \frac{1}{1 - i\frac{\omega_{\text{GW}}}{\omega_p}} \right] \sin(k\delta l) \sin(\phi_D) h_{\text{GW}} \end{aligned} \quad (1.60)$$

where the length disturbance  $\Delta L$  was transformed to originate from a gravitational wave signal  $k\Delta L = \frac{2\pi L}{\lambda} h_{\text{GW}}$ . The response in this particular interferometer setup to the presence of differential length motion is linearly proportional to the gravitational wave amplitude as well as the input power  $P_0 = E_0^2$ , but there is also a frequency dependent component which is the same as a single Fabry-Perot arm that acts like a low pass filter above the corner frequency.

### 1.2.4 Power-Recycled Fabry-Perot Interferometers

**Power Recycling** If the interferometer is operating such that the 4 km arms are exactly different in arms a pi over two times the wavelength, then the intensity of the light at antisymmetric port will be close to null. This means the power from the arms will reflect back towards the input laser. Ref[] shows the effect of adding a partially reflecting mirror to increase the optical gain of the Michaelson for the sidebands and carrier fields.

References: [Meers, Kiwamu]

### 1.2.5 Dual-Recycled Fabry-Perot Interferometers

One of the biggest changes made between initial and advanced LIGO was the addition of a signal recycling mirror at the antisymmetric port shown in Figure []. This extra mirror allows an extra degree of freedom in shaping the sensitivity curve such.

Figure comparing the three cases of increased sensitivity.

References: [Kiwamu]

### 1.2.6 Fundamental Noise Sources

Ref: Evan Hall, GWINC The proceeding sections describe ways to increase the response of LIGO to gravitational waves; equally as important is the science of characterizing and reducing the noise contributions from everything else.

Noise budget:

#### Quantum Noise

One fundamental noise source that is limiting LIGO's sensitivity comes from the fluctuations of quantum vacuum entering the anti-symmetric port and coupling to the input laser. A quantum mechanical description of an interferometer was constructed by Caves [8][9] [10], where he used electric field operators to show that vacuum fluctuations are the cause of radiation pressure and shot noise in an interferometer.

**Quantum states:** It is well known in physics [37] [22] that a solution to the quantum harmonic oscillator in the energy eigenbasis employs the annihilation and creation operators,  $\hat{a}^\dagger$  and  $\hat{a}$ , to factorize the Hamiltonian

$$\hat{H} = \hbar\omega(\hat{N} + 1/2) \quad (1.61)$$

where  $\hat{N} = \hat{a}^\dagger \hat{a}$  is the number operator. When using this formalism to create a coherent electromagnetic field, it is useful to define a unitary operator that displaces the vacuum state [20]:

$$\hat{D} \equiv \hat{D}(\alpha) \equiv \exp(\alpha \hat{a}^\dagger - \alpha^* \hat{a}) = e^{-\frac{|\alpha|^2}{2}} e^{\alpha \hat{a}^\dagger} e^{\alpha^* \hat{a}} \quad (1.62)$$

$$\hat{D}(\alpha) = \hat{D}^{-1}(\alpha) = \hat{D}(-\alpha)$$

$$\hat{D}^\dagger \hat{a} \hat{D} = \hat{a} + \alpha \quad (1.63)$$

$$\hat{D}^\dagger \hat{a}^\dagger \hat{D} = \hat{a}^\dagger + \alpha^*$$

$$|\alpha\rangle = \hat{D} |0\rangle = e^{-\frac{|\alpha|^2}{2}} e^{\alpha \hat{a}^\dagger} |0\rangle \quad (1.64)$$

### Radiation Pressure

One might naively think that power fluctuations in the laser cause radiation pressure effects on the test masses which will result in noise. However, if the 50/50 beamsplitter is perfect, then the momentum transfer to each test mass will be a common length change and will not vary the intensity at the antisymmetric port (or the symmetric port for that matter).

The concept of quantum radiation pressure arises from considering plane wave waves entering the interferometer from both the symmetric and anti-symmetric ports. This method is similar to the input-output methods of section 1.2.1, however, the difference being that the beamsplitter will couple the electric fields from the input laser and quantum vacuum.

To start, consider the electric fields combining at the beamsplitter from both ports to strike the mirrors, respectively,

$$E_x = \frac{1}{\sqrt{2}} \left[ iE_0 + E_{AS,in} \right] \quad (1.65a)$$

$$E_y = \frac{1}{\sqrt{2}} \left[ E_0 + iE_{AS,in} \right] \quad (1.65b)$$

The intensities hitting each test mass will be

$$|E_x|^2 = \frac{1}{2} \left[ |E_0|^2 + |E_{AS,in}|^2 + i(E_0 E_{AS,in}^* - E_0^* E_{AS,in}) \right] \quad (1.66a)$$

$$|E_y|^2 = \frac{1}{2} \left[ |E_0|^2 + |E_{AS,in}|^2 - i(E_0 E_{AS,in}^* - E_0^* E_{AS,in}) \right], \quad (1.66b)$$

The differential momentum transfer between the two masses will be equal to the differences in intensities:

$$\begin{aligned} \mathbf{P} &= \frac{2\hbar\omega}{c} \left( |E_x|^2 - |E_y|^2 \right) \\ &= \frac{2\hbar\omega}{c} \left( E_0 E_{AS,in}^* - E_0^* E_{AS,in} \right) \\ \Rightarrow \hat{\mathbf{P}} &= \frac{2\hbar\omega}{c} \left( \hat{a}_1^\dagger \hat{a}_2 - \hat{a}_2^\dagger \hat{a}_1 \right) \end{aligned} \quad (1.67)$$

The last part of equation 1.67 replaces the classical electric fields with creation and annihilation operators for the symmetric input mode ( $\hat{a}_1^\dagger, \hat{a}_1$ ) and antisymmetric input mode ( $\hat{a}_2^\dagger, \hat{a}_2$ ) modes. Recall that there is a well established convention denoting the wave function of the two dimensional modes of the quantum harmonic oscillator [20],  $|\alpha, \beta\rangle$ , where  $\alpha$  denotes the input symmetric mode and  $\beta$  refers to the input antisymmetric mode.

$$|\alpha, 0\rangle = \hat{D}_1(\alpha) |0, 0\rangle \quad (1.68)$$



The interesting results arising from this formulation is the expectation value

$$\langle \hat{\mathbf{P}} \rangle = \langle \alpha, 0 | \hat{\mathbf{P}} | \alpha, 0 \rangle = 0 \quad (1.69)$$

And the variance

$$\begin{aligned} \Delta \mathbf{P}^2 &= \langle \hat{\mathbf{P}}^2 \rangle - \langle \hat{\mathbf{P}} \rangle^2 \\ &= \left( \frac{2\hbar\omega}{c} \right)^2 \langle \alpha, 0 | \hat{a}_1^\dagger \hat{a}_2 \hat{a}_1 \hat{a}_2^\dagger + \hat{a}_2^\dagger \hat{a}_1 \hat{a}_2 \hat{a}_1^\dagger - \hat{a}_2^\dagger \hat{a}_1 \hat{a}_1 \hat{a}_2^\dagger - \hat{a}_1^\dagger \hat{a}_2 \hat{a}_2 \hat{a}_1^\dagger | \alpha, 0 \rangle \\ &= \left( \frac{2\hbar\omega}{c} \right)^2 |\alpha|^2 \end{aligned} \quad (1.70)$$

For some time interval, input laser power consisting of  $\langle N \rangle = |\alpha|^2$  photons for a time interval  $\Delta T$  is,

$$P_{in} = \frac{\hbar\omega}{\Delta T} |\alpha|^2 \quad (1.71)$$

To solve for the amplitude spectral density of the displacement, Newton's second law can be applied in the frequency domain

$$M(2\pi f)^2 \tilde{x}(f) = \tilde{F}(f) = \frac{\Delta \mathbf{P}}{\Delta T} \quad (1.72)$$

The force spectrum is white because the impacting photons arrive randomly.

Solving for the displacement,

$$\tilde{x}_{RP}(f) = \sqrt{\frac{\hbar\omega}{\Delta T} P_{in}} \frac{1}{2Mc(\pi f)^2} \quad (1.73)$$

The noise spectral density for  $M = 40kg$ ,  $P_{in} = 125W$ , and  $\lambda = 1064nm$

$$\tilde{h}_{RP}(f) = 2.04 \times 10^{-20} \frac{1}{f^2} \left[ \frac{\text{Strain}}{\sqrt{\text{Hz}}} \right] \quad (1.74)$$

**Shot Noise** Another way that quantum fluctuations can vary the antisymmetric port output is by adding phase noise. Imagine holding the test masses rigidly such that the only effects on the output light is due to a phase change in the laser light in the arms. By using the same formulation for radiation pressure but propagating the fields back to the beam splitter, equations 1.65 will have extra phase,

$$E_{x,out} = \frac{1}{\sqrt{2}} \left[ iE_0 + E_{AS,in} \right] e^{-i2kL_x} \quad (1.75a)$$

$$E_{y,out} = \frac{1}{\sqrt{2}} \left[ E_0 + iE_{AS,in} \right] e^{-i2kL_y} \quad (1.75b)$$

Antisymmetric port output electric field is

$$\begin{aligned} E_{AS,out} &= \frac{1}{\sqrt{2}}(E_{x,out} + iE_{y,out}) \\ &= ie^{-ikL_x - ikL_y} [\cos(\Delta\phi)E_0 - \sin(\Delta\phi)E_{AS,in}] \end{aligned} \quad (1.76)$$

where  $\Delta\phi = k(L_x - L_y)$  Output intensity,

$$\begin{aligned} \mathbf{I}_{AS,out} &= |E_{AS,out}|^2 \\ &= \left[ \cos^2(\Delta\phi)|E_0|^2 + \sin^2(\Delta\phi)|E_{AS,in}|^2 - \sin(\Delta\phi)\cos(\Delta\phi)[E_0E_{AS,in}^* + E_0^*E_{AS,in}] \right] \\ \Rightarrow \hat{\mathbf{I}}_{AS,out} &= \left[ \cos^2(\Delta\phi)a_1^\dagger a_1 + \sin^2(\Delta\phi)a_2^\dagger a_2 - \sin(\Delta\phi)\cos(\Delta\phi)[a_1^\dagger a_2 + a_2^\dagger a_1] \right] \end{aligned} \quad (1.77)$$

The expectation value for the intensity using an input coherent laser with  $\alpha$  and quantum vacuum input at the antisymmetric port is

$$\begin{aligned} \langle \hat{\mathbf{I}} \rangle &= \langle \alpha, 0 | \hat{\mathbf{I}}_{AS,out} | \alpha, 0 \rangle \\ &= \cos^2(\Delta\phi)|\alpha|^2 \end{aligned} \quad (1.78)$$

which matches the classical description of a Michelson output.

Following the same methods to calculate the radiation pressure, photon number variance is

$$\begin{aligned} \Delta \mathbf{I} &= \sqrt{\langle \hat{\mathbf{I}}^2 \rangle - \langle \hat{\mathbf{I}} \rangle^2} \\ &= |\alpha| [\cos^2(\Delta\phi)] \end{aligned} \quad (1.79)$$

$$I = N \cos^2(\Delta\phi) \quad (1.80)$$

$$\frac{\partial I}{\partial \phi} = -2|\alpha|^2 \cos(\Delta\phi) \sin(\Delta\phi) \quad (1.81)$$

$$\delta\phi = \sqrt{\frac{\hbar\omega}{P_{in}}} \cot(\Delta\phi) \quad (1.82)$$

Here  $\delta\phi$  is the microscopic change in phase due to shot noise, whereas  $\Delta\phi$  is the DC offset in the arm lengths to begin with. So in general, the shot noise contribution is dependent on the amount of light present at the anti-symmetric port and this will vary depending on what type of read out that is implemented.

$$\tilde{h}_{SN}(f) = 8.2 \times 10^{-22} \left[ \frac{\text{Strain}}{\sqrt{\text{Hz}}} \right] \quad (1.83)$$

Unsurprisingly, the variance is proportional to the amount of photons present and the phase difference between the interferometer arms. Interestingly, there are a few subtleties associated with measuring the noise, it is assumed here that the interferometer readout is a simple photodetector located at the antisymmetric port. However, is it possible to measure the light at two different phases (90 degrees apart) and subtract the results. Also, in section 1.2.1, there is a freedom to choose the readout scheme of the interferometer (heterodyne or homodyne) which will also affect the overall quantum noise.

Losses!?!?!?!?!?

### Seismic Noise

Reference:[Kissel Thesis] Seismic noise will be the low frequency barrier for all terrestrial gravitational-wave detectors. By using seismic isolation platforms and quadruple suspensions, the noise contribution can be attenuated for frequencies larger than 1Hz.

Compare low and high noise microseisms [5] The Advanced LIGO's solution to the low frequency wall is comprised of seismic isolation platforms and multi-level suspensions.

### Thermal Noise

Brownian motion [14] is the random movement of particles suspended in a fluid. The LIGO test masses, which make up the 4 kilometer long Fabry Perot cavities, are a large macroscopic object but their constituent atoms are excited by the ambient temperature and have associated random motion. Those atoms are interconnected and allow for the propagation of an infinite number of elastic wave modes that contain random relative phases and the linear superposition of the modes create distortions that result in thermal noise.

There are two important points that the Fluctuation-Dissipation theorem is based off:

1) When in thermal equilibrium, a system's response to random fluctuations will be the same as when a small force applied. 2) When a system dissipates energy, there exists a reverse process that allows energy to flow back into the system.

A good way to start analyzing thermal noise is to consider a simple harmonic oscillator with mass  $m$ , that is in a thermal bath with a damping force,  $F = -b\dot{x}(t)$ .

$$\ddot{x}(t) + \gamma\dot{x}(t) + \omega_0^2 x(t) = F/m \quad (1.84)$$

Here,  $\gamma = b/m$  and the force,  $F$ , can be described by white noise. In this case, the force represents work being done on the system by the external world whereas the damping factor is the release of the system's energy to the outside world. By taking the Fourier transform, the equation of motion becomes

$$\tilde{x}(f) = \frac{\tilde{F}(f)}{m} \frac{1}{\omega^2 + i\gamma\omega + \omega_0^2} \quad (1.85)$$

This implies that the spectral density for viscous damping is

$$S_{Vis} = \frac{S_F}{m^2} \frac{1}{(\omega^2 - \omega_0^2)^2 + (\gamma\omega)^2} \quad (1.86)$$

Because the noise is due to a thermal bath, the integrated power over all frequencies must satisfy this equation

$$\frac{1}{2} \int_{-\infty}^{\infty} S_{Vis}(\omega) d\omega = \frac{k_B T}{m\omega_0^2} \quad (1.87)$$

where  $k_B$  is the familiar Boltzmann's constant and  $T$  is the ambient temperature. Solving the integral and plugging the result for  $S_F$  back into the equation 1.86 gives the power spectrum for viscous damping,

$$S_{Vis} = \frac{4k_B T \gamma}{m} \frac{1}{(\omega^2 - \omega_0^2)^2 + (\gamma\omega)^2} \quad (1.88)$$

This equation shows that for three different frequency regimes, the response can vary wildly.

$$\omega \ll \omega_0 \rightarrow S_{vis} = \frac{4k_B T}{mQ\omega_0^3} \quad (1.89a)$$

$$\omega \approx \omega_0 \rightarrow S_{vis} = \frac{4k_B T}{m\omega_0^3} Q \quad (1.89b)$$

$$\omega \gg \omega_0 \rightarrow S_{vis} = \frac{4k_B T}{mQ} \frac{\omega_0}{\omega^4} \quad (1.89c)$$

where  $Q = \omega_0/\gamma$  is commonly known as the quality factor of a resonator. Qualitatively, this is directly proportional to the resonance height and its thinness. Figure [visthermalnoise] shows the resultant frequency dependent power spectrum for a few different quality factors, for higher  $Q$  systems the amount of noise contribution from thermal noise is reduced at frequencies above the resonance. Another commonly used definition of the quality factor is the amount of energy loss per cycle, if the  $Q$  of a system is high, then the amount of energy loss per oscillation is very small and will take a longer time returning to equilibrium. A good description of the quality factor and its application to gravitational wave detectors can be found in section 7.5 of [33].

As system's mode number increases, the previous analysis becomes very difficult to solve analytically. However, the connection between a damped harmonic oscillator's energy loss and power spectrum can be described generically using the Fluctuation Dissipation Theorem,

$$S_x(f) = \frac{k_B T}{\pi^2 f^2} |\operatorname{Re}(Y(f))| \quad (1.90)$$

where  $Y(f)$  is the complex mechanical admittance (inverse of the impedance) of a system. The physical interpretation of equation 1.90 can be viewed as a tunnel of energy between a system of interest and the outside world. If there is an ability to couple energy from one system to another via some process, then the reverse must be true as well. This subtle but powerful statement can be applied generally. For example, a resistor which is able to dissipate thermal energy when current is flowing through it will convert random external temperature fluctuations into stray currents that results in *Johnson Noise*.

The mechanical admittance of a system is defined as  $Y(f) = \dot{x}(f)/F(f)$ , which can be written down directly for a damped harmonic oscillator using equation 1.85,

$$\begin{aligned} \frac{\dot{x}(f)}{F} &= \frac{i\omega}{m} \frac{1}{\omega_0^2 - \omega^2 + i\gamma\omega} \\ &= \frac{i\omega}{m} \frac{(\omega_0^2 - \omega^2 - i\gamma\omega)}{(\omega_0^2 - \omega^2)^2 + (\gamma\omega)^2} \end{aligned} \quad (1.91)$$

and by plugging this into the fluctuation dissipation theorem, the power spectral density matches equation 1.88 for the viscous damping. To extend the usage further, consider the thermal noise from a solid which has some thermoelastic dissipation from an internal mode given by  $F_{diss} = -k(1 + \phi_L)x(t)$  where the phase term comes from the lag between applied force and the system's response to said force. By solving Newton's second law with this new damping term, the displacement spectrum becomes

$$\tilde{x}(f) = \frac{\tilde{F}}{m} \frac{1}{(1 + i\phi_L)\omega_0^2 - \omega^2} \quad (1.92)$$

Unlike the previous example of viscous damping, the integral for thermoelastic dissipation only holds for some frequency regime but implementing the fluctuation dissipation theorem leads directly to the noise power spectrum,

$$S_{TD}(f) = \frac{4k_B T \omega_0^2 \phi_L}{m\omega} \frac{1}{(\omega^2 - \omega_0^2)^2 + (\omega_0 \phi_L)^2} \quad (1.93)$$

The frequency response will also be different compared to viscoelastic dissipation,

$$\omega \ll \omega_0 \rightarrow S_{vis} = \frac{4k_B T}{mQ\omega\omega_0^2} \quad (1.94a)$$

$$\omega \approx \omega_0 \rightarrow S_{vis} = \frac{4k_B T}{m\omega_0^3} Q \quad (1.94b)$$

$$\omega \gg \omega_0 \rightarrow S_{vis} = \frac{4k_B T}{mQ} \frac{\omega_0^2}{\omega^5} \quad (1.94c)$$

The structural damping formalism can also be applied to the suspension fibers which hold the 40 kilogram test masses[[21], [33]] which yield a noise spectral density that contains resonances at a variety of frequencies. The lower two modes consist of the pendulum and rocking modes, while starting at 500 hertz the violin modes will be seen. In fact, the violin modes and their harmonics could get rung up from earthquakes and saturate the output mode cleaner's DC photodiodes so they must be actively damped in order to proceed towards nominal low noise.

The largest contribution to the thermal noise for current interferometers do not stem from viscous damping, thermoelastic damping of the substrate, or internal friction of the suspension fibers, but rather from the thermal noise contributed by the dielectric coatings [[23]] which cause phase noise at the high reflectivity surface of the mirrors.

Thermal Noise [[35]]

Suspension Thermal Noise [21]

Substrate Thermal Noise [[33]]

Coating Thermal Noise [[23]]

Thermal-Optical Noise[[17]]

Residual Viscous damping[[33]]

## Newtonian Noise

Although there are some noise sources that can be reduced using increasingly complicated techniques such as various levels of seismic isolation or quantum nondemolition devices as shown above. [[34] and [26]]. Newtonian noise caused by fluctuating gravitational fields from the wave motion of the ground is one that cannot easily be reduced but possibly be measured and fed-forward or subtracted off-line.[[13]]

## 1.3 Squeezed States of Light

Virtually all undergraduate quantum mechanics textbooks include a section on the harmonic oscillator [[37]]. An interesting result when solving the Schrodinger equation using creation ( $\hat{a}^\dagger$ ) and annihilation  $\hat{a}$  operators is the existence of a non-zero energy ground state.

This effect leads to Quantum Noise, which is a fundamental source that can be improved by modifying

the quantum vacuum using correlated photons and injecting the states of light into the antisymmetric port of the interferometer. Within the LIGO community, this procedure of modifying quantum vacuum is called squeezing. Caves analytically derived the effects of quantum noise on the interferometer as well as the improvement due to

Losses!?!?!?!?!? References: [Caves, Dwyer, Kwee, Miao]

## Chapter 2

# The Fundamentals of Mode-Matching

Theory section of modematching. Gaussian beams, define relevant quantities, Gouy Phase!

When dealing with length sensing degrees of freedoms such as section 1.2.2, the simple plane wave approximation is sufficient in describing the dynamics, however, when trying to understand the misalignment and mode-mismatch signals, it is necessary to incorporate Gaussian beams and associated their higher order modes.

### 2.1 Gaussian Beam Optics

Consider the famous Maxwell's equations:

$$\begin{aligned}\nabla \times \mathbf{E} &= -\frac{\partial \mathbf{B}}{\partial t}, \\ \nabla \cdot \mathbf{B} &= 0, \\ \nabla \times \mathbf{B} &= \mu \mathbf{J} + \frac{1}{c^2} \frac{\partial \mathbf{E}}{\partial t} \\ \nabla \cdot \mathbf{E} &= \frac{\rho}{\epsilon}\end{aligned}\tag{2.1}$$

Concentrating on the electric field in vacuum, we arrive at the Helmholtz Equation

$$(\nabla^2 + k^2)\mathbf{U}(\mathbf{r}, t) = 0\tag{2.2}$$

where  $k = \frac{2\pi\nu}{c}$  is the wave number and  $\mathbf{U}(\mathbf{r}, t)$  is the complex amplitude which can describe either the electric or magnetic fields. There are a variety of solutions to equation 2.2 which include the plane and spherical waves [32]. These two types of solutions are the extremes when considering the angle and spatial



distribution as a function of propagation. The plane wave which has been used in chapter 1 is a beam whose rays have no variance in the spatial direction as it propagates through space, whereas the spherical wave starts at a point and spreads as a function of distance from the origin. Somewhere in between these two extrema is the paraxial wave which is a wave pattern that has its rays at small angles relative to the direction of propagation (see Figure[raysoflight]). It is possible to express the paraxial solution to equation 2.2 as a plane wave with a modulated complex envelope  $\mathbf{A}(r)$

$$\mathbf{U}(\mathbf{r}) = \mathbf{A}(\mathbf{r})e^{-ikz} \quad (2.3)$$

However, in order for the paraxial wave to exist,  $\mathbf{U}(r)$  must satisfy the Helmholtz equation or alternatively, the complex amplitude portion must satisfy a separate differential equation known as the paraxial Helmholtz Equation. By imposing the constraints which force the envelope to vary slowly with respect to the z-axis within the distance of one wavelength  $\lambda = 2\pi/k$ ,

$$\left| \frac{\partial^2 \mathbf{A}}{\partial z^2} \right| \ll \left| k \frac{\partial \mathbf{A}}{\partial z} \right| \quad (2.4a)$$

$$\left| \frac{\partial^2 \mathbf{A}}{\partial z^2} \right| \ll \left| k \frac{\partial^2 \mathbf{A}}{\partial x^2} \right| \quad (2.4b)$$

$$\left| \frac{\partial^2 \mathbf{A}}{\partial z^2} \right| \ll \left| k \frac{\partial^2 \mathbf{A}}{\partial y^2} \right| \quad (2.4c)$$

the partial differential equation which arises is called the Paraxial Helmholtz Equation:

$$\nabla_T^2 A(r) - i2k \frac{\partial A(r)}{\partial z} = 0 \quad (2.5)$$

where  $\nabla_T^2 = \frac{\partial^2}{\partial x^2} + \frac{\partial^2}{\partial y^2}$  is the transverse Laplacian. A simple solution for equation 2.5 is the complex paraboloidal wave

$$A(\mathbf{r}) = \frac{A_0}{q(z)} e^{\frac{-ikr^2}{2q(z)}}, \quad q(z) = z + iz_0 \quad (2.6)$$

where  $z_0$  is the Rayleigh range and is directly proportional to the square of the waist size. In order to separate the amplitude and phase portions of the wave, it is useful to rewrite  $q(z)$  as

$$\frac{1}{q(z)} = \frac{1}{R(z)} - i \frac{\lambda}{\pi W^2(z)} \quad (2.7)$$

Plugging equation 2.7 into 2.6 leads directly to the complex amplitude for a Gaussian Beam

$$U(r) = A_0 \frac{W_0}{W(z)} e^{-\frac{r^2}{W^2(z)}} e^{-ikz - ik \frac{r^2}{2R(z)} + i\phi(z)} \quad (2.8)$$

where

$$W(z) = W_0 \sqrt{1 + \left(\frac{z}{z_0}\right)^2} \quad (2.9a)$$

$$R(z) = z \left[ 1 + \left(\frac{z}{z_0}\right)^2 \right] \quad (2.9b)$$

$$\phi(z) = \tan^{-1} \left( \frac{z}{z_0} \right) \quad (2.9c)$$

$$W_0 = \sqrt{\frac{\lambda z_0}{\pi}} \quad (2.9d)$$

### Gouy Phase

The Gouy phase given by equation 2.9c is the phase lag that occurs as a function of propagation along the  $\hat{z}$ -axis which ranges from  $\phi(z) = \pm\pi/2$  for  $z = \pm\text{inf}$ . Physically, this manifests itself as the difference between the laser beam behaving like a plane wave when  $z \approx 0$  (or near field) and eventually evolving into a spherical wave when  $z \approx \pm\text{inf}$  (or far field) (see Figure[GouyPhaseFig]). Although this phenomenon is interesting mathematically, most texts do not consider how to practically use or measure the Gouy phase. In LIGO technical notes and electronic logs, the Gouy phase of sensors or actuators are used to describe where along the laser beam path the piece of equipment is located. This will indicate what degree of freedom is being observed or adjusted, respectively. For example, in ray optics [32] a laser beam has two degrees of freedom, the displacement and angle from the optical axis (see Figure[Gouyphase]). If an actuator such as a piezo-electric transducer with an attached reflective mirror is placed near the origin or focus of the laser beam, then the controlled degree of freedom is only the angle. Alternatively, if the actuator is placed in the far field then the controlled degree of freedom is almost entirely displacement. This is also true of sensors at various points along the  $\hat{z}$ -axis to determine the exact alignment of the optical beam.

In practice, it is impossible to sense the true waist of an optical system because it could be inside a Fabry-Perot cavity. So the clever designer must insert a pick-off beam to sample the transmission or reflection of the interested optical system and use a combination of lenses to create a telescope for the sensors and infer the degrees of freedom. Then, by rotating to the two-dimensional space created by the sensors to represent the interested space spanned by the optical system's degrees of freedom (see Figure[RotDofs]); colloquially, this is known as a sensing matrix. Using this method, sensors do not have to be exactly at the near or far field because it is only required that the two sensors are separated by  $\pi/2$  in Gouy phase, in fact, it is usually easier to place one sensor at  $+\pi/4$  and the other at  $-\pi/4$ .

Unfortunately, there is no transducer that directly measures the Gouy phase so in order to *determine* the Gouy phase of a particular sensor or actuator, one must model the beam propagation and fit the waist location.

### Hermite-Gauss Modes

The fundamental Gaussian beam is not the only solution which can be used to solve equation 2.5. In fact, there exists a complete set of solutions that can solve the paraxial Helmholtz Equation in rectangular coordinates, which are referred to as the Hermite Gauss modes

$$U_{mn}(x, y, z) = A_{mn} \left[ \frac{W_0}{W(z)} \right] \mathbb{G}_m \left( \frac{\sqrt{2}x}{W(z)} \right) \mathbb{G}_n \left( \frac{\sqrt{2}y}{W(z)} \right) \times \exp \left\{ -ikz - \frac{ik(x^2 + y^2)}{2R(z)} + i(m + n + 1)\phi(z) \right\} \quad (2.10)$$

where,

$$\mathbb{G}(u) = \mathbb{H}(u) \exp(-u^2/2) \quad (2.11)$$

and  $\mathbb{H}(u)$  are the well known Hermite polynomials. It is important to mention that the Gouy phase of the complex amplitude is different than the fundamental Gaussian beam by a factor of  $(m + n + 1)$ . Also, the intensity distribution of these higher order modes are much different. Both of these facts will become extremely important in the following wavefront sensing discussion (Figure[HGModes]).

The extra Gouy phase that a higher order mode accrues as a function of  $z$  shows up as in the frequency or cavity length domain as one sweeps through a full free spectral range using an offset in the length or frequency error signal, see Figure[FSRwHOM]. This can be directly used to measure the higher order mode separation and becomes important if the cavity has a low finesse which means the fundamental Gaussian mode has a spread comparable in frequency to the first order mode then it is easy to hop from one to the other.

Figure[FSRwHOM] The difference between left and right plot is representative of the PRC and SRC. The former has a high finesse and a relatively large round trip Gouy phase where as the latter has a low finesse and very little Gouy phase. An error signal could get confused by which lock point to use if a transient moves through the system and causes a mode hop.

### Laguerre Modes

Another complete set of alternative solutions to equation 2.5 exists which are called the Laguerre-Gauss modes

$$\begin{aligned}
V_{\mu\nu}(\rho, \theta, z) &= A_{\mu\nu} \left[ \frac{W_0}{W(z)} \right] \mathbb{L}_{\nu}^{\mu} \left( \frac{\sqrt{2}x}{W(z)} \right) \\
&\times \exp \left\{ -ikz - \frac{ik\rho^2}{2R(z)} + i(\mu + 2\nu + 1)\phi(z) \right\}
\end{aligned} \tag{2.12}$$

where  $\mathbb{L}_{\nu}^{\mu} \left( \frac{\sqrt{2}x}{W(z)} \right)$  is the Laguerre polynomial function. Both equations 2.10 and 2.12 are able to fully describe any complex electromagnetic amplitude; and because they both form complete sets, there is a rotation which can map from one basis to the other [ref Bond and Biejergasern]

$$U_{\mu\nu}^{LG}(x, y, z) = \sum_k^N i^k b(n, m, k) U_{N-k, k}^{HG}(x, y, z) \tag{2.13}$$

where

$$b(n, m, k) = \sqrt{\left( \frac{(N-k)!k!}{2^N n!m!} \right)} \frac{1}{k!} \frac{d^k}{dt^k} [(1-t)^m (1+t)^m] |_{t=0} \tag{2.14}$$

### 2.1.1 Misalignment and Higher Order Modes

Morrison and Anderson [2] [27] derived a simplistic way of how small misalignments and mode mismatch in cavities can couple the fundamental Gaussian beam into various higher order modes. This is done by taking a linear cavity and using its perfectly matched Gaussian beam as a reference, and then varying the input electric field with small perturbations and expanding in terms of the higher order modes. So long as the mismatches are small, it is possible to consider only the first few terms of the expansion which have gained most of their power from the fundamental mode.

Consider the first three Hermite-Gauss modes of equation 2.10 in one dimension and normalized to set the total optical power to unity (derived in Appendix[]):

$$\begin{aligned}
U_0(r) &= \left( \frac{2}{\pi w^2(z)} \right)^{1/4} e^{-r^2/w^2(z)} \\
U_1(r) &= \left( \frac{2}{\pi w^2(z)} \right)^{1/4} \frac{2r}{w(z)} e^{-r^2/w^2(z)} \\
U_2(r) &= \left( \frac{2}{\pi w^2(z)} \right)^{1/4} \frac{1}{\sqrt{2}} \left( \frac{4r^2}{w^2(z)} - 1 \right) e^{-r^2/w^2(z)}
\end{aligned} \tag{2.15}$$

#### Beam Axis Tilted

If the input beam into an optical cavity is tilted by an angle  $\alpha$  with respect to the nominal cavity axis as shown in Figure [], the wave front of the input beam will have an extra phase propagation relative to the cavity that is approximately proportional to  $e^{ik\alpha r}$ . By implementing the small angle approximation, which

is valid if the misalignment is much smaller than the divergence angle of the fundamental mode  $k\alpha r \ll 1$ , the resultant input beam is

$$\Psi \approx U_0(r)e^{ik\alpha r} \approx U_0(r)(1 + ik\alpha r) = U_0(r) + \frac{ik\alpha w(z)}{\sqrt{2\pi}}U_1(r) \quad (2.16)$$

Here the factor associated with the first higher order mode is complex, indicating there is a 90 degree phase difference between the fundamental and off axis mode.

### Beam Axis Displaced

If the input beam is displaced in the transverse direction by a quantity  $\Delta r$ , the resultant waveform will be

$$\begin{aligned} \Psi &= U_0(r + \Delta r) \\ &= \left( \frac{2}{\pi w^2(z)} \right)^{1/4} e^{-(r+\Delta r)^2/w^2(z)} \\ &= \left( \frac{2}{\pi w^2(z)} \right)^{1/4} e^{-(r^2+2r\Delta r+\Delta r^2)/w^2(z)} \\ &\approx \left( \frac{2}{\pi w^2(z)} \right)^{1/4} e^{-r^2/w^2(z)} e^{-2r\Delta r/w^2(z)} \\ &\approx \left( \frac{2}{\pi w^2(z)} \right)^{1/4} e^{-r^2/w^2(z)} \left( 1 - \frac{2r\Delta r}{w^2(z)} \right) \\ &= \left( U_0(r) - \sqrt{\frac{2}{\pi}} \frac{\Delta r}{w(z)} U_1(r) \right) \end{aligned} \quad (2.17)$$

Similarly to a tilted input beam axis, the displaced beam axis couples power to the first higher order mode, however, the latter does not have a 90 degree phase difference previously seen in the former. This point is of extreme importance when trying to discern between the two effects as shown in Section []. Although comparing the two cases in Figure, one can already see the difference between the wavefronts in the near field,  $z \ll z_R$ , and the far field,  $z \gg z_R$ .

In the near field, there is no phase difference due to a displaced beam, but there is one for a tilted beam. Conversely, in the far field, there is no phase difference due to a tilted beam, but there is one from a displaced beam. In order to implement a closed loop feedback system, the wavefront sensors discussed in Section 2.2 will use this precise logic to extract an error signal.

### 2.1.2 Mode Mismatch and Higher Order Modes

#### Waist Size Shifted

By considering the effect of evaluating the fundamental mode at the waist position,  $z = 0$ , but changing the waist size by a small amount  $\epsilon$ , it is possible to see coupling into higher order modes by expanding to first order.

$$\begin{aligned}
\Psi &= U_0(r, w(z) = w_0/(1 + \epsilon)) \\
&= \left(\frac{2}{\pi w_0^2}\right)^{1/4} \sqrt{1 + \epsilon} \ e^{-r^2(1+\epsilon)^2/w_0^2} \\
&\approx \left(\frac{2}{\pi w_0^2}\right)^{1/4} (1 + \epsilon/2) \ e^{-r^2/w_0^2} \ e^{-2r^2\epsilon/w_0^2} \\
&\approx \left(\frac{2}{\pi w_0^2}\right)^{1/4} (1 + \epsilon/2) \ e^{-r^2/w_0^2} \ (1 - 2r^2\epsilon/w_0^2) \\
&\approx \left(\frac{2}{\pi w_0^2}\right)^{1/4} \left(1 + 2\epsilon\left(\frac{1}{4} - \frac{r^2}{w_0^2}\right)\right) \ e^{-r^2/w_0^2} \\
&= U_0(r) - \frac{\epsilon}{\sqrt{2}} U_2(r)
\end{aligned} \tag{2.18}$$

Changing the waist size by a small amount will couple the fundamental mode to the in-phase second order Hermite Gauss mode.

#### Waist Position Shifted

To repeat the process from above with a waist position shift, it is useful to start with a more general equation that includes the phase that is gained from including the radius of curvature,

$$\Psi = \left(\frac{2}{\pi w^2(z)}\right)^{1/4} e^{-r^2/w^2(z)} e^{-ikr^2/2R(z)} \tag{2.19}$$

where  $R(z)$  is from equation 2.9b. It is also useful to approximate the shift in waist position along the longitudinal axis is small compared to the Rayleigh range of the beam,  $\Delta z \ll z_0$ , which leaves the waist size approximately the same and the radius of curvature inversely proportional to the shift.

$$w^2(\Delta z) = w_0^2 \left[1 + \left(\frac{\Delta z}{z_0}\right)^2\right] \approx w_0^2 \tag{2.20a}$$

$$R(\Delta z) = \Delta z \left(1 + \left(\frac{z_0}{\Delta z}\right)^2\right) \approx \frac{z_0^2}{\Delta z} \tag{2.20b}$$

Plugging the equations above into 2.19,

$$\begin{aligned}
\Psi &\approx \left(\frac{2}{\pi w_0^2}\right)^{1/4} e^{-r^2/w_0^2} e^{-ikr^2\Delta z/2z_0^2} \\
&\approx \left(\frac{2}{\pi w_0^2}\right)^{1/4} e^{-r^2/w_0^2} \left(1 - \frac{ikr^2\Delta z}{2z_0^2}\right) \\
&= U_0(r) - \left(\frac{2}{\pi w_0^2}\right)^{1/4} e^{-r^2/w_0^2} \frac{ikr^2\Delta z}{2z_0^2} \\
&= U_0(r) - i\frac{\Delta z}{2kw_0^2} \left(4U_2(r) + U_0(r)\right)
\end{aligned} \tag{2.21}$$

The equations above show that a fundamental Gaussian mode that is shifted in waist position will couple power to the second order Hermite Gauss mode. Although changes in the waist size or position couple power to the same mode, they differ by a 90 degrees in phase as denoted by the extra factor of  $i$  in the coupling coefficient. By recognizing the two effects are in different quadrature phases will allow a user to design a system to distinguish between the different types of physical couplings, this is shown in Section 2.2.

In order to be physically valid one would need to consider the full two dimensional space so that the equation would encapsulate the full transverse mode, however, the x and y components would follow the exact same derivation. On that point, it is important to note that only the mode mismatch couplings from either a varying waist position or size has higher order modes that are circularly symmetric.

## 2.2 Wavefront Sensing

Heterodyne detection via modal decomposition of the full electric field allows the use of wavefront sensors to extract an error signal from the optical system. Hefetz et.al [24] created a formalism to describe the use of wavefront sensors by creating frequency sidebands which accumulate a different Gouy phase than the electric field at the carrier frequency when passed through the optical system. By observing the demodulated signal of the intensity, it is possible to obtain a linear signal that corresponds to a physical misalignment or mode mismatch.

Fundamentally, the purpose of wavefront sensing is to detect the content of higher order modes due to physical disturbances of the optical cavity (ie. mode mismatch or misalignment). In other words, it is examining the difference between the incoming beam's and the cavity's eigenmodes.

Consider a general equation for an electric field which is a linear combination of all higher order modes of the complex amplitude

$$E(x, y, z) = \sum_{m,n}^{\infty} a_{mn} U_{mn}(x, y, z) \tag{2.22}$$

where  $U_{mn}(x, y, z)$  are the eigenmodes described in equation 2.10 (or 2.12) and  $a_{mn}$  is the complex

amplitude. It is also convenient in the following analysis to use vectors when describing the composition of the electric fields.

$$|E(x, y, z)\rangle = \begin{pmatrix} E_{00} \\ E_{01} \\ E_{10} \\ E_{20} \\ E_{02} \end{pmatrix} \quad (2.23)$$

When creating a theory that involves laser beams, it is useful to define operators that are important in describing physical situations. For example, laser beams propagate through space and pick up phase according to equation 2.10 which can be represented by the spatial propagation operator,

$$\hat{P}_{mn,kl} = \delta_{mn}\delta_{kl} \exp[-ik(z_2 - z_1)]\exp[i(m + n + 1)\phi(z)] \quad (2.24)$$

However, it is useful to compare how the fundamental Gaussian mode propagates compared to the higher order modes,

$$\hat{\eta}_{\mu\nu} = \begin{pmatrix} e^{i\phi} & 0 & 0 & 0 & 0 \\ 0 & e^{2i\phi} & 0 & 0 & 0 \\ 0 & 0 & e^{2i\phi} & 0 & 0 \\ 0 & 0 & 0 & e^{3i\phi} & 0 \\ 0 & 0 & 0 & 0 & e^{3i\phi} \end{pmatrix} \quad (2.25)$$

From the above diagonal elements, it is clear that the higher order modes have an extra phase compared to the fundamental 00 mode, this effect will be extremely important on how an error signal can be derived from the optical system.

$$|E(x, y, z_2)\rangle = \hat{M}(x, y, z_1, z_2) |E(x, y, z_1)\rangle \quad (2.26)$$

where  $\hat{M}(x, y, z_1, z_2)$  is the misalignment operator. Since we are using the paraxial approximation, the z-components of the misalignment operator are small so we can approximate  $\hat{M}(x, y, z_1, z_2) \approx \hat{M}(x, y)$  and the expectation value is

$$M_{mn,kl} = \langle U_{mn}(x, y, z_1) | M(x, y) | U_{kl}(x, y, z_2) \rangle \quad (2.27)$$



where the product is an integral over the transverse space  $\iint_{D(x,y)} dx dy$

$$\hat{\Theta}_{\mu\nu} = \begin{pmatrix} 1 & 2i\theta_x & 2i\theta_y & 0 & 0 \\ 2i\theta_x & 1 & 0 & 0 & 0 \\ 2i\theta_y & 0 & 1 & 0 & 0 \\ 0 & 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 0 & 1 \end{pmatrix} \quad (2.28)$$

$$\hat{\mathbb{D}}_{\mu\nu} = \begin{pmatrix} 1 & \alpha_x/\omega_0 & \alpha_y/\omega_0 & 0 & 0 \\ \alpha_x/\omega_0 & 1 & 0 & 0 & 0 \\ \alpha_y/\omega_0 & 0 & 1 & 0 & 0 \\ 0 & 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 0 & 1 \end{pmatrix} \quad (2.29)$$

$$\hat{\mathbb{Z}}_{\mu\nu} = \begin{pmatrix} 1 & 0 & 0 & \Delta z_x & \Delta z_y \\ 0 & 1 & 0 & 0 & 0 \\ 0 & 0 & 1 & 0 & 0 \\ \Delta z_x & 0 & 0 & 1 & 0 \\ \Delta z_y & 0 & 0 & 0 & 1 \end{pmatrix} \quad (2.30)$$

where  $\Delta z_{(x,y)} = \frac{i}{\sqrt{2}} \frac{\lambda b}{2\pi\omega_0}$

$$\hat{\mathbb{Z}}_{0,\mu\nu} = \begin{pmatrix} 1 & 0 & 0 & \Delta z_{0,x} & \Delta z_{0,y} \\ 0 & 1 & 0 & 0 & 0 \\ 0 & 0 & 1 & 0 & 0 \\ \Delta z_{0,x} & 0 & 0 & 1 & 0 \\ \Delta z_{0,y} & 0 & 0 & 0 & 1 \end{pmatrix} \quad (2.31)$$

where  $\Delta z_{0,(x,y)} = \frac{1}{\sqrt{2}} \frac{\omega' - \omega_0}{\omega_0}$

Figure: Beat frequency between two modes.

**Example: Fabry Perot Cavity**

**Example: Simple Michaelson**

## 2.3 Effects of Mode-matching on Squeezing

Still not clear to me. In the most simple sense, a loss is akin to coupling quantum vacuum into a squeezed state. References: [Miao, Sheon, Kimble, Dwyer]

## Chapter 3

# Simulating Mode-Matching with Finesse

### 3.1 How it works

[29] Summary of how Finesse works (input output matrix), how it handles HOMs

### 3.2 Finesse Simulations

#### 3.2.1 ALIGO Design with FC and Squeezer

#### 3.2.2 Looking at just Modal Change

#### 3.2.3 QM Limited Sensitivity

### 3.3 Results

- \* Signal recycling cavity mismatches

- \* Mismatches before the OMC

- \* Mismatch contour graph: Comparing all of ALIGO cavities

- \* Optical Spring pops up at 7.4 Hz in the Signal-to-Darm TF, re-run with varying SRM Trans which should.

## Chapter 4

# Experimental Mode Matching Cavities at Syracuse

In conjunction with Sandoval et al, the adaptive modematching table top experiment was able to show the feasibility of a fully dynamical system.

### 4.1 Adaptive Mode Matching

Reque Real time digital system and model.

Signal chain.

### 4.2 Actuators

#### 4.2.1 Thermal Lenses

Fabian's work and UFL paper.

#### 4.2.2 Translation Stages

### 4.3 Sensors

Recall in Section 2.2 that the RF modal decomposition technique relies on comparing the Gouy phase of the higher order modes to the fundamental Gaussian mode using an array of RF photodiodes to extract an error signal. The complication arises when using this method to extract the beat note between the fundamental

mode and symmetric donut mode because the photodetector arrays must match the higher order mode geometry. This is easier to deal with when trying to sense angular distortions from a misaligned cavity because the 01 and 10 modes can be sensed with a split photodetector, however, using that a quadrant photodiode will not work to sense mode mismatch due to the cylindrical symmetry of the 02 and 20 modes. To date, Advanced LIGO uses no RF sensors to detect mode matching so the next sections will provide sensing methods which can be used for dynamic closed loop feedback control.

### 4.3.1 Bullseye Photodiodes (BPD)

It is clear that using a quadrant photodetector to measure the symmetric  $U_{20} + U_{02}$  mode which arises from mismatch is futile because the integrated power on each side will be exactly the same due to donut mode symmetry. So it is simplest to try subtracting the inner beam power from the outer ring of the electric field in order to measure a phase difference using a specialized RF photodiode called a Bullseye Photodetector shown in Figure [ ].

#### BPD Calibration

To use the BPDs, there are a few steps required in designing the optical setup which is different than the standard PDH or WFS method. Most notably, the beam size incident on the BPD must be tuned such that the zero crossing of the  $(U_{20} + U_{02})$  matches the boundary between the inner and outer segments. This condition is met if  $\omega_0 = \sqrt{2}r_0$  and Appendix C shows the power ratio of the outer to inner segments when this is true,

$$\text{Power Ratio} = \frac{P_2 + P_3 + P_4}{P_1} = \frac{e^{-2r_0^2/\omega_0^2}}{1 - e^{-2r_0^2/\omega_0^2}} \approx 0.582 \quad (4.1)$$

Picture of BPD

Pitch and Yaw sensing matrix

Demodulation phase

Another constraint on using this method is that the Gouy phase separation between successive BPDs must be close to 45 or 135 degrees so that the error signals from each BPD can be orthogonalized, this is in contrast to angular wavefront sensors which require 90 degree Gouy phase separation.

### 4.3.2 Mode Converters

The bullseye photodiodes can be difficult to calibrate and manufacture so a particularly interesting method of sensing mode mismatch is to convert the fields with a cylindrical telescope. Using lenses which contain separate curvatures for each direction (x or y), one being flat and the other is curved with focal length,  $f$ .

The idea is to break the cylindrical symmetry of the donut mode (Figure []) and convert the beam into a pringle mode, then, an error signal can be extracted using the radio frequency quadrant photodiode that the angular wavefront sensors employ

### Vector formalism for laser beams and optical cavities

Consider an optical cavity that is longitudinally resonant on the TEM<sub>00</sub> mode using the Pound-Drever-Hall technique shown in Section 1.2.2 and also locked with angular wavefront sensors shown in Section 2.2. When there is a small mismatch between the waist size and position of the input beam relative to the cavity, this is equivalent to coupling the TEM-00 mode into higher order modes. Since mode matching is only concerned with coupling to the second order modes, the resultant field in reflection ( $r_{FP}$ ) of a Fabry-Perot resonator will have the form

$$|U_{refl}\rangle = r_{FP} \begin{pmatrix} U_{00} \\ 0 \\ 0 \end{pmatrix} + \epsilon \begin{pmatrix} 0 \\ U_{20} \\ U_{02} \end{pmatrix} \quad (4.2)$$

$$\epsilon = \frac{1}{\sqrt{2}} \left( \frac{\delta w}{w_0} + i \frac{\delta z}{z_R} \right) \quad (4.3)$$

where  $\delta w$  and  $\delta z$  are the mismatches in waist size and position, respectively.

O'Neil et al [30] wrote down a formalism that explicitly showed the effects of cylindrical telescopes on the full range of Hermite Gauss modes. An arbitrary mode converter is comprised of two cylindrical lenses that can be rotated about the axis of propagation shown in Figure []:

$$\hat{M} = \hat{R}\hat{C} \quad (4.4)$$

where  $\hat{R}$  is a rotation operator and  $\hat{C}$  is the mode converting operator. Since mode matching primarily couples power into the 02 and 20 modes of the cavity, we can use a 3x3 matrix. For rotations about the axis of propagation, the operator is

$$\hat{R}_{ij} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos(\Delta\theta) & \sin(\Delta\theta) \\ 0 & -\sin(\Delta\theta) & \cos(\Delta\theta) \end{pmatrix} \quad (4.5)$$

For the mode converter, the operator is

$$\hat{C}_{ij} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & e^{-i\Omega_{20}} & 0 \\ 0 & 0 & e^{i\Omega_{02}} \end{pmatrix} \quad (4.6)$$

where  $\Omega_{mn} = (m + \frac{1}{2}) \tan^{-1} \left( \frac{d}{z_{R,x}} \right) + (n + \frac{1}{2}) \tan^{-1} \left( \frac{d}{z_{R,y}} \right)$  and  $d$  is the distance from one cylindrical lens to the waist.  $\Omega_{mn}$  is the amount of extra phase that a higher order mode will experience due to the mode converter. It is important to note, the matrices above show that the Gaussian beam is only astigmatic within the region of between the cylindrical lenses and unchanged outside of the telescope. If the beam reflected from the cavity is well mode matched to the cylindrical telescope, then a mode converter will introduce an astigmatism and vary the separate Rayleigh ranges [3],

$$\frac{z_{R,x}}{z_{R,y}} = \frac{1 + d/f}{1 - d/f} \quad (4.7)$$

Of course, the tuning of  $\theta$ ,  $d$  and  $f$  is left up to the optical designer's choice. The obvious selection for  $\theta$  should rotate the output beam such that the pringle mode intensities are centered on the individual quadrant photodiodes. The choice of separation distance  $d$  and focal length  $f$  are a bit more subtle. In Figure [] from O'Neil [30], the only conversion that transforms the diagonal HG mode into a symmetric donut mode is with  $\Delta\Omega = \pi/2$ . This implies the converse is true if one desires to transform the donut mode into an HG mode of the same order. Making this choice of phase propagation will automatically constrain the optical setup,

$$\frac{\pi}{2} = 2 \left[ \tan^{-1} \left( \frac{d}{z_{R,x}} \right) - \tan^{-1} \left( \frac{d}{z_{R,y}} \right) \right] \quad (4.8)$$

$$\Rightarrow \sqrt{2} - 1 = \frac{d}{z_{R,y}} = \frac{z_{R,x}}{d} \quad (4.9)$$

The equation above only shows a relation between the Rayleigh ranges and the separation. However, by imposing mode matching conditions it is possible also constrain the focal length of the cylindrical lenses as well.

$$f = \left[ \frac{1}{R_x(d)} - \frac{1}{R_y(d)} \right]^{-1} = \sqrt{2} d \quad (4.10)$$

### Extracting information from mode

Mueller et al [28] showed that the error signal from mode mismatch could be extracted by using an RF detection scheme. Using this formalism, the error signal on a quadrant photodetector from a mismatched

cavity after a mode converter is

$$\begin{aligned}
S &\propto \sin(\Omega t) \text{Im} \left\{ \epsilon^* \left[ \int_{A1,A3} \hat{M}_{00}^\dagger \langle U_{00} | [\hat{M}_{20}^\dagger |20\rangle + \hat{M}_{02}^\dagger |02\rangle] - \int_{A2,A4} \hat{M}_{00}^\dagger \langle U_{00} | (\hat{M}_{20}^\dagger |20\rangle + \hat{M}_{02}^\dagger |02\rangle) \right] \right\} \\
&\propto \sin(\Omega t) \text{Im} \left\{ \epsilon^* \left[ \int_{A1,A3} \langle U_{00} | (|20\rangle - |02\rangle) - \int_{A2,A4} \langle U_{00} | (|20\rangle - |02\rangle) \right] \right\}
\end{aligned} \tag{4.11}$$

In the above equation,  $\hat{C}_{ij}$  was chosen with  $\Delta\Omega = \pi/2$  and  $\hat{R}_{ij}$  should rotate the beams such that the intensities in Figure ModeConverter are aligned with the quadrant photodiodes.

### 4.3.3 DC Mode Matching

The benefit to using the RF sensing scheme for generating an error signal is the automatic reduction of the higher order mode content which can use actuators to adjust the incoming mode of the beam to whatever the cavity requires. This is particularly useful in areas where the laser power density is high and absorption on the high reflectivity surfaces cause thermal distortions in the cavity mode. However, the output mode cleaner remains generally fixed in its eigenbasis such that the nominal input mode is static; in fact, the current alignment scheme uses DC quadrant photodiodes to feed back to the OMC suspensions and the control loop offsets are tuned to minimize the 10/01 modes. If there was a low-noise way to continually measure the beam size at two different Gouy phases prior to entering the output mode cleaner, then there could be an error signal which would be used for thermal actuators directly after the signal recycling cavity. This can be achieved with bullseye sensors that read power ratio relatively quickly. Another interesting way of measuring the mode is using a slowly rotating razor blade with known angular frequency in front of a single photodiode and inferring the error function which will directly give the beam size.



## Chapter 5

# Mode Matching Cavities at LIGO Hanford

Reference[Dan Hoake, Kognelik and Li, aLOGS]

Modeling: Finesse vs Alamode

Measurements: Mode profiling and Cavity Scanning

Misalignments described in Chapter 2 can be actively suppressed using the modal decomposition technique which leaves reduces the amount of odd order modes present in the cavity. This leaves the even order modes which arise from modal mismatch to be the dominate sources of optical losses assuming mirrors with a small amount of scatter loss( $< 100$  ppm).

### 5.1 Mode Matching IFO to OMC: Single Bounce vs DRMI

#### 5.1.1 SR3 Heater

Beckhoff code

#### 5.1.2 SRM Heater

Test set up in OSB optics lab - Mode Profiling: Single Bounce - OMC Scanning

### 5.2 Mode Matching Squeezer to OMC

- Mode Profiling - OMC Scanning - Astigmatism: expected in LIGO and measured!

### 5.3 Mode Matching as a Function of DARM Offset

### 5.4 Modal Contrast Defect

Generally, the contrast defect is defined as the ratio of power between the antisymmetric port and the reflected port when locked on a dark fringe. In other words, it is the amount of junk light that is present in the interferometer when light between the two arms do not perfectly interfere with each other. This junk light can be the symptom of various causes, for example, an imbalance of reflectivity between ITMX and ITMY will cause non-perfect destructive interference at the antisymmetric port and a camera would see a TEM00 beam when locked on length. Another cause of contrast defect could be from misalignment between the ITMs or beamsplitter which will result in seeing a TEM 01/10 mode. However, if both of the aforementioned causes are fixed with a combination of stringent design specifications for the reflectivity and alignment loops closed to minimize angular jitter, then the contrast defect will be dominated by mode mismatch which can be fixed by a combination of ring heaters and CO2 lasers. The picture gets even more complicated when adding in the absorption for individual optics and introducing multiple Fabry-Perot cavities which will treat the sidebands and carrier fields differently. One of the main goals for the Thermal Compensation System is to correct the cold and hot interferometer differences in radii of curvature. A useful equation that relates the contrast defect to the fields incident on the  $\hat{x}$  and  $\hat{y}$  side of the beamsplitter,

$$CD = \frac{P_{AS}}{P_{REFL}} \Big|_{Dark} = \frac{1}{2} \left( 1 - \int \psi_X \psi_Y dA \right) \quad (5.1)$$

where the amplitude overlap integral can be directly calculated from equation D.6.

In the modal picture, the contrast defect can be defined using the zeroth eigenmode of each arm and then expanded to project the X-arm's basis onto Y-arm using higher order Laguerre-Gauss modes,  $LG_y^{00} \rightarrow LG_x^{00} + \alpha LG_x^{10}$ . Where  $\alpha = \frac{1}{\sqrt{2}} \left( \frac{\Delta\omega_0}{\omega_0} + i \frac{\Delta z}{z_R} \right)$  is the amount of higher order mode coupling due to mismatch in beam size and location, respectively (see Chapter 2).

$$\begin{aligned} CD &\equiv \frac{P_{AS}}{P_{REFL}} \\ &= \frac{|LG_x^{00} - LG_y^{00}|^2}{|LG_x^{00} + LG_y^{00}|^2} \\ &= \frac{|LG_x^{00}|^2 + |LG_x^{00} + \alpha LG_x^{10}|^2 - 2 \operatorname{Re}(LG_x^{00} [LG_x^{00*} + \alpha LG_x^{10}])}{|LG_x^{00}|^2 + |LG_x^{00} + \alpha LG_x^{10}|^2 + 2 \operatorname{Re}(LG_x^{00} [LG_x^{00*} + \alpha LG_x^{10}])} \\ &\approx \frac{\alpha^2}{4} \\ &\approx \frac{1}{8} \left[ \left( \frac{\Delta\omega_0}{\omega_0} \right)^2 + \left( \frac{\Delta z}{z_R} \right)^2 \right] \end{aligned} \quad (5.2)$$

Mismatch between the arm cavities can stem from a few different sources such as mismatches between the test masses radii of curvature on the high reflectivity surfaces will cause the resonant modes to be different between the X-arm and Y-arm. In addition to modes being different, there exists a static (or cold) substrate lens in each of the compensation plates. This static lens becomes important for two reasons: Firstly, the sidebands in DRMI are promptly rejected by the arm cavities. Secondly, the gravitational wave sidebands which are generated and resonant in the arms will see the substrate lensing as it travels to the beamsplitter.

### 5.4.1 Simple Michelson Contrast Defect

The simple Michelson can give the first estimate of the contrast defect when starting to commission the interferometer, however, it will not be directly used in nominal low noise since the dual-recycled Michelson is implemented. By propagating the input mode cleaner beam to each of the input test masses and taking a single bounce at the high reflectivity surface, one can approximate the resultant mode overlap. At this point the carrier and sideband fields follow the same ABCD matrix transfer function, so only one calculation is needed to estimate the contrast defect. A keen reader will notice that this model will not take the Schnupp asymmetry into account which allows the  $\hat{x}$ -direction beam to travel an extra 8 centimeters further than the  $\hat{y}$ , however, this effect will only change the end result by approximately 10%. In fact, for this interferometer configuration, the dominate source of mismatch will be from the prompt reflection off the HR surfaces of the ITMs where most of the phase change occurs.

Figure[]

### 5.4.2 Dual Recycled Michelson Contrast Defect

The contrast defect for DRMI is similar to the simple Michelson analysis, however, there is an added complexity in that the resonance of the power recycling cavities (PRX and PRY) make up the mode content at the beamsplitter. This means that the power recycling mirror has to be mode matched well to the ITMs and this also means the field inside DRMI will see the substrate lensing.

### 5.4.3 Fabry-Perot Contrast Defect

The contrast defect from the two 4km arms will be dominated by the resonant mode of the individual arm cavities and will not see the static substrate lens distortions. The reasoning for this is subtle and described in Section 5.5.1.

## 5.5 Hot vs. Cold Interferometers

In the previous section, the contrast defect calculations did not include the self-heating thermal effects that will be present when the interferometer reaches its nominal state. When fully operational, the arm cavities can have a few hundred kilowatts of laser power resonating while in observing mode. An approximate but useful description of the total arm power in each arm is

$$P_{\text{ARM}} \approx \frac{1}{2}(g_{\text{PRC}} * g_{\text{ARM}} * P_{\text{IN}}) \quad (5.3)$$

where  $g_{\text{PRC}} \approx 30$  is the power recycling cavity gain and  $g_{\text{ARM}} \approx 250$  is the single arm cavity gain. For O3, the intended input power,  $P_{\text{IN}}$  will reach approximately 50 W which means  $P_{\text{ARM}} \approx 187,500$  W. A fraction of that power, approximately .5 PPM, will be absorbed by the high reflectivity surface, hence, creating an additional thermal lens which changes the radius of curvature. Additionally, the thermo-refractive index will change as a function of absorbed temperature [38]. By changing the curvature of the test masses, the resonant Gaussian mode will also change its profile. For the substrate lensing, the sideband signals will be most adversely affected because the carrier field will see a cancellation of the distortion from the promptly reflected beam (see Section [])

### 5.5.1 Thermal Lensing

As seen in Section 1.2.5, the sideband and carrier frequencies propagate differently in the interferometer, which means their fields will see different lensing from thermal effects. To understand how the fields change quantitatively, the easiest way is to invoke the ABCD transfer matrix approach for the carrier and sideband fields individually.

#### Carrier

Consider Figure [], a two-mirror optical system which has an input beam,  $E_{\text{in}}$ , that has a portion promptly reflected off the input mirror to create  $E_{\text{prompt}}$

$$|E_{\text{prompt}}^{\text{carrier}}\rangle = \hat{M}_{\text{carrier}}^{\text{prompt}} |E_{\text{in}}^{\text{carrier}}\rangle \quad (5.4)$$

**Sidebands****GW Signal**

As mentioned in Section 1.2.1, LIGO's gravitational wave readout currently employs a DC readout scheme that extracts the signal by beating the carrier field with the sidebands created by the gravitational wave.

## Chapter 6

# Solutions for Detector Upgrades

A solution to the current mode matching situation will be solved by further expanding active wavefront control described in Section [6]. A full modal picture, sensors and actuator

### 6.1 Sensors

- \* Bullseye photodetectors

- \* Operation: range (in terms of watts and percentage mismatch)

### 6.2 High Sensitivity Telescope

- \* Translation stages: General principles, rigging, sensitivity

- \* Mechanical description (Solidworks designs)

- \* Constraints (range, vacuum, alignment, integration)

- \* Electronics

- \* Software

- \*Implementation: rigg something better, in situ measurements.

### 6.3 A Global Active Wavefront Control Scheme

To truly implement a system that solves the mode matching challenges presented in LIGO, an active control scheme that implements both sensors and actuators will be required to span the degrees of freedom that are present in the interferometer.

### 6.3.1 Hard and Soft Mode Matching

Similar to the wavefront sensors already being used in LIGO at the time of this writing, the active wavefront system should be able to rotate the phase of their actuators and sensors to the common and differential basis. Also, the extra two degrees of freedom required are the hard and soft modes for thermal lensing which curve the ITMs and ETMs in a way which vary the waist size and position orthogonally.

# Appendices



## Appendix A

# Resonator Formulas

Equation 1.50 describes the stability condition for a two mirror Fabry-Perot cavity. It is worthwhile to derive the criterion for geometric stability from the ray matrix tools commonly used in optics.

Consider two plane waves traveling in space shown by figure [], they differ by two quantities: the axial and angular separations,  $y$  and  $\theta$ , respectively. These two quantities can be transformed via these optical matrices:

Lens

$$\hat{F}_i = \begin{pmatrix} 1 & 0 \\ -\frac{1}{f_i} & 1 \end{pmatrix} \quad (\text{A.1})$$

Curved Mirror

$$\hat{M}_i = \begin{pmatrix} 1 & 0 \\ \frac{2}{R_i} & 1 \end{pmatrix} \quad (\text{A.2})$$

Space

$$\hat{D}_i = \begin{pmatrix} 1 & d_i \\ 0 & 1 \end{pmatrix} \quad (\text{A.3})$$

Periodic Fabry-Perot with two mirrors is

$$\begin{pmatrix} y_{m+1} \\ \theta_{m+1} \end{pmatrix} = \hat{M}_{FP} \begin{pmatrix} y_m \\ \theta_m \end{pmatrix} \quad (\text{A.4})$$

where

$$\hat{M}_{FP} = \hat{M}_i \hat{D}_i \hat{M}_i \hat{D}_i \quad (\text{A.5})$$

is the optical transfer matrix. The goal is to find a geometric condition that is dependent on the optical transfer matrix which keeps the axial displacement from diverging.

$$\begin{pmatrix} y_{m+1} \\ \theta_{m+1} \end{pmatrix} = \begin{pmatrix} A & B \\ C & D \end{pmatrix} \begin{pmatrix} y_m \\ \theta_m \end{pmatrix} \quad (\text{A.6})$$

which means

$$\theta_m = \frac{y_{m+1} - A y_m}{B} \quad (\text{A.7a})$$

$$\theta_{m+1} = \frac{y_{m+2} - A y_{m+1}}{B} \quad (\text{A.7b})$$

Solving for  $y_{m+2}$

$$y_{m+2} = (A + D) y_{m+1} - \det(\hat{M}_{FP}) y_m \quad (\text{A.8})$$

Assuming a geometrical solution where  $y_m = y_o h^m$  and plugging into the equation above,

$$h^2 = (A + D) h - \det(\hat{M}_{FP}) \quad (\text{A.9})$$

which is a quadratic equation that has two solutions and can be further simplified if the index of refraction for the entire system remains constant such that  $\det(\hat{M}_{FP}) = 1$ . Plugging back into  $y_m$  and doing some algebra

$$y_m \propto \sin(m\phi) \quad (\text{A.10})$$

where  $\phi = \cos^{-1}(\frac{A+D}{2})$ , which is also referred to as the round trip Gouy phase of the cavity. In order for  $y_m$  to be harmonic instead of hyperbolic and hence confined, this condition must be met

$$\frac{|A + D|}{2} \leq 1 \quad (\text{A.11})$$

By actually calculating the terms of  $\hat{M}_{FP}$  and doing even more algebra, it is clear that

$$0 \geq \left(1 + \frac{L}{R_1}\right) \left(1 + \frac{L}{R_2}\right) \geq 1 \quad (\text{A.12})$$

which is what was stated in equation 1.50.

There is a simpler and less algebraic way to reach the same conclusion by looking at the Rayleigh range

of a finite Gaussian beam for a simple cavity. In Table II of Kogelnik and Li ref[Kogelnik], there is an expression for the Rayleigh range

$$\begin{aligned} z_R^2 &= \frac{L(R_1 - L)(R_2 - L)(R_1 + R_2 - L)}{(R_1 + R_2 - 2L)^2} \\ &= \frac{g_1 g_2 (1 - g_1 g_2)}{(g_1 - g_2 - 2g_1 g_2)^2} \end{aligned} \tag{A.13}$$

If the Rayleigh range is a real number, then once again, equation 1.50 must be true.

Effect of higher order modes into the cavity, mode scanning. All comes from round trip Gouy phase. -  
RT Gouy Phase - HOM Coupling

## Appendix B

# Hermite Gauss Normalization

According to equation 2.10, the higher order modes in the Hermite Gauss basis has the intensity profile,

$$I_{mn}(x, y, z) = |A_{mn}|^2 \left[ \frac{W_0}{W(z)} \right]^2 \mathbb{G}_n^2 \left( \frac{\sqrt{2}x}{W(z)} \right) \mathbb{G}_n^2 \left( \frac{\sqrt{2}y}{W(z)} \right) \quad (\text{B.1})$$

It is useful to normalize the first few lowest order modes with respect to the total optical power since the Gaussian beam will couple to them the most due either misalignment or mode mismatch as seen in section [3](#).

In one dimension, the total optical power for the first 3 modes are

$$\begin{aligned} P_0(x, y, z) &= \int_{-\infty}^{\infty} |A_0|^2 \left[ \frac{W_0}{W(z)} \right]^2 e^{-2x^2/w^2(z)} dx \\ P_1(x, y, z) &= \int_{-\infty}^{\infty} |A_1|^2 \left[ \frac{W_0}{W(z)} \right]^2 \frac{8x^2}{w^2(z)} e^{-2x^2/w^2(z)} dx \\ P_2(x, y, z) &= \int_{-\infty}^{\infty} |A_2|^2 \left[ \frac{W_0}{W(z)} \right]^2 \left( \frac{8x^2}{w^2(z)} - 2 \right)^2 e^{-2x^2/w^2(z)} dx \end{aligned} \quad (\text{B.2})$$

In two dimensions, the total optical power for the first 3 modes are

$$\begin{aligned} P_{00}(x, y, z) &= \int_{-\infty}^{\infty} \int_{-\infty}^{\infty} |A_{00}|^2 \left[ \frac{W_0}{W(z)} \right]^2 e^{-2x^2/w^2(z)} e^{-2y^2/w^2(z)} dx dy \\ P_{10}(x, y, z) &= \int_{-\infty}^{\infty} \int_{-\infty}^{\infty} |A_{10}|^2 \left[ \frac{W_0}{W(z)} \right]^2 \frac{8x}{w^2(z)} e^{-2x^2/w^2(z)} e^{-2y^2/w^2(z)} dx dy \\ P_{20}(x, y, z) &= \int_{-\infty}^{\infty} \int_{-\infty}^{\infty} |A_{20}|^2 \left[ \frac{W_0}{W(z)} \right]^2 \left( \frac{8x^2}{w^2(z)} - 2 \right)^2 e^{-2x^2/w^2(z)} e^{-2y^2/w^2(z)} dx dy \end{aligned} \quad (\text{B.3})$$

By setting the equations above to unity, the normalization factors become

$$\begin{aligned}
A_0 &= \left( \frac{2}{\pi w_0^2} \right)^{1/4} \\
A_1 &= \left( \frac{2}{\pi w_0^2} \right)^{1/4} \frac{1}{\sqrt{2}} \\
A_2 &= \left( \frac{2}{\pi w_0^2} \right)^{1/4} \frac{1}{\sqrt{8}}
\end{aligned} \tag{B.4}$$

$$\begin{aligned}
A_{00} &= \sqrt{\frac{2}{\pi w_0^2}} \\
A_{10} &= \sqrt{\frac{1}{\pi w_0^2}} \\
A_{20} &= \sqrt{\frac{1}{4\pi w_0^2}}
\end{aligned} \tag{B.5}$$

Therefore the normalized modes are

## Appendix C

# Bullseye and Quadrant Photodiode Characterization

Both RF and DC segmented photodiodes are widely used in LIGO's sensing schemes, however, the quadrant photodiodes (QPDs) are more readily implemented. However, recently a bullseye photodiode has been installed upstream of the pre-mode cleaner to sense the size and angular jitter noise contribution from the high powered oscillator. It is useful to have angular sensors both upstream and downstream of optical cavities in order to narrow down where jitter could be coupling into the optical path. That being said, to calibrate the two types of sensors in common units, it is easiest to use units of beam diameters.

### C.1 Quadrant Photodiodes (QPD)

Calculating the QPD response to angular jitter is relatively straight forward. Beginning with a Gaussian beam in rectangular coordinates that is displaced along the horizontal axis by a small value  $\Delta x$  and expanding to first order:

$$\begin{aligned} I(x, y) &= \left( \frac{2}{\pi w^2} \right) e^{-2 \frac{(x+\Delta x)^2 + y^2}{w^2}} \\ &\approx \left( \frac{2}{\pi w^2} \right) e^{-2 \frac{x^2 + y^2}{w^2}} \left( 1 - \frac{4x\Delta x}{w^2} \right) \end{aligned} \tag{C.1}$$

	Seg 1	Seg 2	Seg 3	Seg 4
Pit	1	1	-2	0
Yaw	-1	1	0	0
Wid	1	1	1	-1
Sum	1	1	1	1

Table C.1: Bullseye photodiode sensing matrix.

Converting to polar coordinates where  $x = r \cos \theta$  and  $y = r \sin \theta$  then integrating over the individual segments to get the power,

$$\begin{aligned}
P &= \left( \frac{2}{\pi w^2} \right) \int_0^\infty \int_{\theta_1}^{\theta_2} e^{-2\frac{r^2}{w^2}} \left( 1 - \frac{4r \cos \theta}{w^2} \Delta x \right) r dr d\theta \\
&= \frac{1}{2} - [\cos \theta_2 - \cos \theta_1] \frac{8\Delta x}{\pi w^4} \int_0^\infty e^{-2\frac{r^2}{w^2}} r^2 dr \\
&= \frac{1}{2} - [\cos \theta_2 - \cos \theta_1] \frac{\Delta x}{2\pi w^2} \left( \sqrt{2\pi} w \operatorname{erf}\left(\frac{\sqrt{2}r}{w}\right) - 4re^{-2\frac{r^2}{w^2}} \right) \Big|_0^\infty \\
&= \frac{1}{2} - [\cos \theta_2 - \cos \theta_1] \frac{\Delta x}{w} \sqrt{\frac{1}{2\pi}}
\end{aligned} \tag{C.2}$$

where  $\theta_1$  and  $\theta_2$  are the limits from  $\pm\pi/2$  for the right half and  $\mp\pi/2$  for the left half. Denoting the right and left segment as  $P_2$  and  $P_1$ , respectively, it is possible to subtract the halves and obtain a calibrated DC signal in units of beam diameter.

$$S = P_2 - P_1 = 2\sqrt{\frac{2}{\pi}} \frac{\Delta x}{w} \tag{C.3}$$

It is trivial to repeat the calculation for vertical displacements so it is left for the reader to complete.

## C.2 Bullseye Photodiodes

Similar to the QPD calibration above, it is possible to express pitch and yaw displacements that are normalized to beam diameters. Additionally, the bullseye's geometry will give insight on the beam size jitter. A sensing matrix for these degrees of freedom is realized in Table C.2

### C.2.1 Width

The bullseye's calibration is determined in a similar manner to the QPD however, the sensitivity will inherently depend on the beam size. To calculate what beam size is optimal, first consider a power integral for a

cylindrically symmetric Gaussian beam,

$$\begin{aligned} \text{Power} &= \int_A^B |A_{00}|^2 e^{-\frac{2r^2}{\omega^2}} 2\pi r dr \\ &= -|A_{00}|^2 \frac{\pi\omega^2}{2} e^{-\frac{2r^2}{\omega^2}} \Big|_A^B \end{aligned} \quad (\text{C.4})$$

Plugging in limits for the equation where  $R$  is the boundary between the inner and outer segments,

$$P_{in} = \text{Power} \Big|_0^R = |A_{00}|^2 \frac{\pi\omega^2}{2} [1 - e^{-\frac{2R^2}{\omega^2}}] \quad (\text{C.5})$$

$$P_{out} = \text{Power} \Big|_R^\infty = |A_{00}|^2 \frac{\pi\omega^2}{2} [e^{-\frac{2R^2}{\omega^2}}] \quad (\text{C.6})$$

The error signal comes from subtracting the inner segment from the outer,

$$S_{WID} = P_{in} - P_{out} \quad (\text{C.7})$$

Minimizing the derivative with respect to the beam size will determine the optimal width which gives the best sensitivity to beam size change,

$$\begin{aligned} \frac{\partial S_{WID}}{\partial \omega} &\approx \frac{8R^2}{\omega^2} \left(1 - \frac{2R^2}{\omega^2}\right) = 0 \\ \Rightarrow \omega &= \sqrt{2}R \end{aligned} \quad (\text{C.8})$$

When determining the beam size, it is possible to measure the beam size directly by fitting the power ratio.

For example, if the constraint that  $\omega = \sqrt{2}R$  then the power ratio is,

$$\text{DC Power Ratio} = \frac{P_{out}}{P_{in}} = \frac{e^{-2R^2/\omega_0^2}}{1 - e^{-2R^2/\omega_0^2}} \approx 0.582 \quad (\text{C.9})$$

### C.2.2 Pitch

To calibrate the pitch signal on a bullseye, first consider a Gaussian that is displaced in the vertical direction,

$$\begin{aligned} I^{\text{Pit}}(x, y) &= \left(\frac{2}{\pi w^2}\right) e^{-2\frac{x^2 + (y + \Delta y)^2}{w^2}} \\ &\approx \left(\frac{2}{\pi w^2}\right) e^{-2\frac{x^2 + y^2}{w^2}} \left(1 - \frac{4x\Delta y}{w^2}\right) \end{aligned} \quad (\text{C.10})$$



The integrated power for a given segment is,

$$\begin{aligned} P_{\theta_1 \rightarrow \theta_2}^{\text{Pit}} &= \left( \frac{2}{\pi w^2} \right) \int_R^\infty \int_{\theta_1}^{\theta_2} e^{-2\frac{r^2}{w^2}} \left( 1 - \frac{4r \sin \theta}{w^2} \Delta y \right) r dr d\theta \\ &= \left( \frac{\theta_2 - \theta_1}{2\pi} \right) e^{-2\frac{R^2}{w^2}} + (\cos \theta_2 - \cos \theta_1) \frac{1}{\sqrt{2\pi}} \frac{\Delta y}{\omega} \left[ \text{erfc} \left( \frac{\sqrt{2}R}{\omega} \right) \sqrt{\frac{8}{\pi}} \frac{R}{\omega} e^{-2\frac{R^2}{w^2}} \right] \end{aligned} \quad (\text{C.11})$$

The error signal from a pitch displacement following Table C.2 is,

$$S^{\text{Pit}} = P_{seg1}^{\text{Pit}} + P_{seg2}^{\text{Pit}} - 2P_{seg3}^{\text{Pit}} = 3\sqrt{\frac{3}{2\pi}} \frac{\Delta y}{\omega} \left[ \text{erfc} \left( \frac{\sqrt{2}R}{\omega} \right) + \sqrt{\frac{8}{\pi}} \frac{R}{\omega} e^{-2\frac{R^2}{w^2}} \right] \quad (\text{C.12})$$

### C.2.3 Yaw

Using C.1 and repeating the mathematics above, the power for an individual segment is

$$\begin{aligned} P_{\theta_1 \rightarrow \theta_2}^{\text{Yaw}} &= \left( \frac{2}{\pi w^2} \right) \int_R^\infty \int_{\theta_1}^{\theta_2} e^{-2\frac{r^2}{w^2}} \left( 1 - \frac{4r \cos \theta}{w^2} \Delta x \right) r dr d\theta \\ &= \left( \frac{\theta_2 - \theta_1}{2\pi} \right) e^{-2\frac{R^2}{w^2}} + (\cos \theta_2 - \cos \theta_1) \frac{1}{\sqrt{2\pi}} \frac{\Delta x}{\omega} \left[ \text{erfc} \left( \frac{\sqrt{2}R}{\omega} \right) \sqrt{\frac{8}{\pi}} \frac{R}{\omega} e^{-2\frac{R^2}{w^2}} \right] \end{aligned} \quad (\text{C.13})$$

Plugging in the angles to get the signal response in terms of beam radius,

$$S^{\text{Yaw}} = P_{seg1}^{\text{Yaw}} - P_{seg2}^{\text{Yaw}} = \frac{3}{\sqrt{2\pi}} \frac{\Delta x}{\omega} \left[ \text{erfc} \left( \frac{\sqrt{2}R}{\omega} \right) + \sqrt{\frac{8}{\pi}} \frac{R}{\omega} e^{-2\frac{R^2}{w^2}} \right] \quad (\text{C.14})$$

### C.2.4 RF

$$\begin{aligned} P_{RF} &= \int_A^B |A_{01}|^2 \left( 1 - \frac{2r^2}{\omega_0^2} \right) e^{-\frac{2r^2}{\omega_0^2}} 2\pi r dr \\ &= -|A_{01}|^2 \frac{\pi}{2} \omega_0^2 e^{-\frac{2r^2}{\omega_0^2}} \left( 1 + \frac{4r^4}{\omega_0^4} \right) \Big|_A^B \end{aligned} \quad (\text{C.15})$$

$$P_{in} = P_{RF} \Big|_0^{r_0} = -|A_{01}|^2 \frac{\pi}{2} \omega_0^2 \left( e^{-\frac{2r_0^2}{\omega_0^2}} \left( 1 + \frac{4r_0^4}{\omega_0^4} \right) - 1 \right) \quad (\text{C.16})$$

$$P_{out} = P_{RF} \Big|_{r_0}^\infty = -|A_{01}|^2 \frac{\pi}{2} \omega_0^2 e^{-\frac{2r_0^2}{\omega_0^2}} \left( 1 + \frac{4r_0^4}{\omega_0^4} \right) \quad (\text{C.17})$$

$$\text{RF Power Ratio} = \frac{P_{out}}{P_{in}} = \frac{e^{-2r_0^2/\omega_0^2}}{1 - e^{-2r_0^2/\omega_0^2}} \approx 2.7844 \quad (\text{C.18})$$

## Appendix D

# Overlap of Gaussian Beams

Referenced in section When a cavity is mode mismatched to an incoming laser field, the amount of power loss from scattering to higher order modes is quantified by the spatial overlap integral between the TEM00 cavity eigenmode and TEM00 of the input beam. An arbitrary Gaussian integral is defined as,

$$\begin{aligned} |A(r)\rangle &= \frac{A_0}{q(z)} e^{\frac{-ikr^2}{2q(z)}} \\ &= \frac{A_0}{q(z)} e^{\frac{-ikr^2(z-iz_0)}{2|q(z)|^2}} \end{aligned} \quad (\text{D.1})$$

where  $A_0$  is a real amplitude,  $q(z) = z + iz_0$  is the complex beam parameter,  $k$  is the wave number, and  $r$  is the radial variable in the transverse direction. Then normalizing to unity,

$$\langle A(r)|A(r)\rangle = \frac{|A_0|^2}{z^2 + z_0^2} \int_0^\infty e^{\frac{-kr^2 z_0}{|q(z)|^2}} 2\pi r dr = 1 \quad (\text{D.2})$$

$$A_0 = \sqrt{\frac{kz_0}{\pi}} \quad (\text{D.3})$$

For a Gaussian beam with arbitrary q-parameters,

$$|A_i\rangle = \frac{A_{0,i}}{q_i} e^{\frac{-ikr^2(z-iz_0)}{2|q_i|^2}} \quad (\text{D.4})$$

where  $z_{0,i}$  is the waist size of one particular beam, the overlap integral for the amplitude becomes

$$\langle A_1|A_2\rangle = 2i \frac{z_{0,1}z_{0,2}}{q_1 - q_2^*} \quad (\text{D.5})$$

So the power overlap is:

$$\text{Power Overlap} = |\langle A_1 | A_2 \rangle|^2 = 4 \frac{z_{0,1} z_{0,2}}{|q_1 - q_2^*|^2} \quad (\text{D.6})$$

Essentially, mode matching an optical system only requires the designer to match the q-parameter of the incoming beam to the cavity, so it makes sense that the final power overlap depends only on the waist size and location.

## Appendix E

# Extending the ABCD Formalism

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# Bibliography

- [1] J. Abadie et al. Sensitivity Achieved by the LIGO and Virgo Gravitational Wave Detectors during LIGO’s Sixth and Virgo’s Second and Third Science Runs. *arXiv*, 2012. 8
- [2] D. Z. Anderson. Alignment of resonant optical cavities. *Appl. Opt.*, 23(17):2944–2949, Sep 1984. doi: 10.1364/AO.23.002944. URL <http://ao.osa.org/abstract.cfm?URI=ao-23-17-2944>. 34
- [3] M. Beijersbergen, L. Allen, H. van der Veen, and J. Woerdman. Astigmatic laser mode converters and transfer of orbital angular momentum. *Optics Communications*, 96(1):123 – 132, 1993. ISSN 0030-4018. doi: [https://doi.org/10.1016/0030-4018\(93\)90535-D](https://doi.org/10.1016/0030-4018(93)90535-D). URL <http://www.sciencedirect.com/science/article/pii/003040189390535D>. 45
- [4] E. D. Black. An introduction to pound–drever–hall laser frequency stabilization. *American Journal of Physics*, 69(1):79–87, 2001. doi: 10.1119/1.1286663. URL <https://doi.org/10.1119/1.1286663>. 15, 18
- [5] D. G. Blair. *The detection of gravitational waves*. Cambridge University Press, 2005. 25
- [6] A. F. Brooks, B. Abbott, M. A. Arain, G. Ciani, A. Cole, G. Grabeel, E. Gustafson, C. Guido, M. Heintze, A. Heptonstall, M. Jacobson, W. Kim, E. King, A. Lynch, S. O’Connor, D. Ottaway, K. Mailand, G. Mueller, J. Munch, V. Sannibale, Z. Shao, M. Smith, P. Veitch, T. Vo, C. Vorvick, and P. Willems. Overview of advanced ligo adaptive optics. *Appl. Opt.*, 55(29):8256–8265, Oct 2016. doi: 10.1364/AO.55.008256. URL <http://ao.osa.org/abstract.cfm?URI=ao-55-29-8256>. 52
- [7] S. M. Carroll. *Spacetime and geometry: an introduction to general relativity*. Pearson, 2003. 5
- [8] C. M. Caves. Quantum-mechanical noise in an interferometer. *Phys. Rev. D*, 23:1693–1708, Apr 1981. doi: 10.1103/PhysRevD.23.1693. URL <https://link.aps.org/doi/10.1103/PhysRevD.23.1693>. 21
- [9] C. M. Caves and B. L. Schumaker. New formalism for two-photon quantum optics. i. quadrature phases

- and squeezed states. *Phys. Rev. A*, 31:3068–3092, May 1985. doi: 10.1103/PhysRevA.31.3068. URL <https://link.aps.org/doi/10.1103/PhysRevA.31.3068>. 21
- [10] C. M. Caves, K. S. Thorne, R. W. P. Drever, V. D. Sandberg, and M. Zimmermann. On the measurement of a weak classical force coupled to a quantum-mechanical oscillator. i. issues of principle. *Rev. Mod. Phys.*, 52:341–392, Apr 1980. doi: 10.1103/RevModPhys.52.341. URL <https://link.aps.org/doi/10.1103/RevModPhys.52.341>. 21
- [11] J. L. Cervantes-Cota, S. Galindo-Uribarri, and G. F. Smoot. A brief history of gravitational waves. *Universe*, 2(3), 2016. ISSN 2218-1997. doi: 10.3390/universe2030022. URL <http://www.mdpi.com/2218-1997/2/3/22>. 7
- [12] H. Collins. *Gravitys shadow: the search for gravitational waves*. University of Chicago Press, 2004. 7
- [13] J. C. Driggers, J. Harms, and R. X. Adhikari. Subtraction of newtonian noise using optimized sensor arrays. , 86(10):102001, Nov. 2012. doi: 10.1103/PhysRevD.86.102001. 28
- [14] A. Einstein. Investigations on the theory of the brownian movement. *Ann. der Physik*, 1905. URL [http://www.physik.fu-berlin.de/~kleinert/files/eins\\_brownian.pdf](http://www.physik.fu-berlin.de/~kleinert/files/eins_brownian.pdf). 25
- [15] A. Einstein. Zur Elektrodynamik bewegter Körper. (German) [On the electrodynamics of moving bodies]. *Annalen der Physik*, 322(10):891–921, 1905. doi: <http://dx.doi.org/10.1002/andp.19053221004>. 5
- [16] R. W. et al. Ligo document m890001-v3, 1989. URL <https://dcc.ligo.org/LIGO-M890001/public>. 10
- [17] M. Evans, S. Ballmer, M. Fejer, P. Fritschel, G. Harry, and G. Ogin. Thermo-optic noise in coated mirrors for high-precision optical measurements. *Phys. Rev. D*, 78:102003, Nov 2008. doi: 10.1103/PhysRevD.78.102003. URL <https://link.aps.org/doi/10.1103/PhysRevD.78.102003>. 28
- [18] L. S. Finn. Binary inspiral, gravitational radiation, and cosmology. *Phys. Rev.*, D53:2878–2894, 1996. doi: 10.1103/PhysRevD.53.2878. 8, 10
- [19] P. Fritschel, R. Bork, G. González, N. Mavalvala, D. Ouimette, H. Rong, D. Sigg, and M. Zucker. Readout and control of a power-recycled interferometric gravitational-wave antenna. *Appl. Opt.*, 40(28):4988–4998, Oct 2001. doi: 10.1364/AO.40.004988. URL <http://ao.osa.org/abstract.cfm?URI=ao-40-28-4988>. 15, 19
- [20] C. C. Gerry and P. Knight. *Introductory quantum optics*. Cambridge University Press, 2005. 21, 22



- [21] G. I. González and P. R. Saulson. Brownian motion of a mass suspended by an anelastic wire. *The Journal of the Acoustical Society of America*, 96(1):207–212, 1994. doi: 10.1121/1.410467. URL <https://doi.org/10.1121/1.410467>. 28
- [22] D. J. Griffiths. *Introduction to quantum mechanics*. Pearson Prentice Hall, 2005. 21
- [23] G. M. Harry, A. M. Gretarsson, P. R. Saulson, S. E. Kittelberger, S. D. Penn, W. J. Startin, S. Rowan, M. M. Fejer, D. R. M. Crooks, G. Cagnoli, J. Hough, and N. Nakagawa. Thermal noise in interferometric gravitational wave detectors due to dielectric optical coatings. *Classical and Quantum Gravity*, 19(5): 897, 2002. URL <http://stacks.iop.org/0264-9381/19/i=5/a=305>. 28
- [24] Y. Hefetz, N. Mavalvala, and D. Sigg. Principles of calculating alignment signals in complex resonant optical interferometers. *J. Opt. Soc. Am. B*, 14(7):1597–1606, Jul 1997. doi: 10.1364/JOSAB.14.001597. URL <http://josab.osa.org/abstract.cfm?URI=josab-14-7-1597>. 37
- [25] S. Hild, H. Grote, J. Degallaix, S. Chelkowski, K. Danzmann, A. Freise, M. Hewitson, J. Hough, H. Lück, M. Prijatelj, K. A. Strain, J. R. Smith, and B. Willke. Dc-readout of a signal-recycled gravitational wave detector. *Classical and Quantum Gravity*, 26(5):055012, 2009. URL <http://stacks.iop.org/0264-9381/26/i=5/a=055012>. 15
- [26] S. A. Hughes and K. S. Thorne. Seismic gravity-gradient noise in interferometric gravitational-wave detectors. *Phys. Rev. D*, 58:122002, Nov 1998. doi: 10.1103/PhysRevD.58.122002. URL <https://link.aps.org/doi/10.1103/PhysRevD.58.122002>. 28
- [27] E. Morrison, B. J. Meers, D. I. Robertson, and H. Ward. Automatic alignment of optical interferometers. *Appl. Opt.*, 33(22):5041–5049, Aug 1994. doi: 10.1364/AO.33.005041. URL <http://ao.osa.org/abstract.cfm?URI=ao-33-22-5041>. 34
- [28] G. Mueller, Q. ze Shu, R. Adhikari, D. B. Tanner, D. Reitze, D. Sigg, N. Mavalvala, and J. Camp. Determination and optimization of modgit ste matching into optical cavities by heterodyne detection. *Opt. Lett.*, 25(4):266–268, Feb 2000. doi: 10.1364/OL.25.000266. URL <http://ol.osa.org/abstract.cfm?URI=ol-25-4-266>. 45
- [29] U. of Birmingham. *Finesse Manual*, 2014. URL <http://www.gwoptics.org/finesse/>. 41
- [30] A. T. O’Neil and J. Courtial. Mode transformations in terms of the constituent hermite–gaussian or laguerre–gaussian modes and the variable-phase mode converter. *Optics Communications*, 181(1): 35 – 45, 2000. ISSN 0030-4018. doi: [https://doi.org/10.1016/S0030-4018\(00\)00736-7](https://doi.org/10.1016/S0030-4018(00)00736-7). URL <http://www.sciencedirect.com/science/article/pii/S0030401800007367>. 44, 45

- [31] F. A. E. Pirani. Republication of: On the physical significance of the riemann tensor. *General Relativity and Gravitation*, 41(5):1215–1232, May 2009. ISSN 1572-9532. doi: 10.1007/s10714-009-0787-9. URL <https://doi.org/10.1007/s10714-009-0787-9>. 7
- [32] B. E. A. Saleh and M. C. Teich. *Fundamentals of Photonics*. Wiley-Interscience, 2009. 30, 32
- [33] P. Saulson. *Fundamentals of Gravitational Wave Detectors*. World Scientific, 2016. 9, 12, 16, 26, 28
- [34] P. R. Saulson. Terrestrial gravitational noise on a gravitational wave antenna. *Phys. Rev. D*, 30:732–736, Aug 1984. doi: 10.1103/PhysRevD.30.732. URL <https://link.aps.org/doi/10.1103/PhysRevD.30.732>. 28
- [35] P. R. Saulson. Thermal noise in mechanical experiments. *Phys. Rev. D*, 42:2437–2445, Oct 1990. doi: 10.1103/PhysRevD.42.2437. URL <https://link.aps.org/doi/10.1103/PhysRevD.42.2437>. 28
- [36] B. F. Schutz. *A first course in general relativity*. Cambridge University Press, 2009. 5
- [37] R. Shankar. *Principles of Quantum Mechanics*. Springer Verlag, 2014. 21, 28
- [38] W. Winkler, K. Danzmann, A. Rüdiger, and R. Schilling. Heating by optical absorption and the performance of interferometric gravitational-wave detectors. *Phys. Rev. A*, 44:7022–7036, Dec 1991. doi: 10.1103/PhysRevA.44.7022. URL <https://link.aps.org/doi/10.1103/PhysRevA.44.7022>. 50