

Searching in the Dark

Chasing Magnum Opus

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1 Introduction

The detection of gravitational waves (GW) on September 14, 2015, opened a new into the Universe. Joint observations of electromagnetic (EM) counterparts to the GW signals will result in a deeper understanding of the physics involved. We anticipate EM counterparts to binary sources involving matter i.e. neutron star-neutron star (NSNS) and neutron star-black hole (NSBH) binaries. In an effort to measure these joint signals, GW detectors are working in conjunction with traditional observatories. Some of these will yield weak, nearly isotropic electromagnetic counterparts and others will not.

This report is organized as follows. Section 2 describes the developement of the chirp mass distribution, Section 3 descibes a EM followup classifier based on the data, and Section 4 will state our conclusions.

2 Chirp Mass Distribution

2.1 Density Estimation

We use a histogram to estimate the merger rate as a function of \mathcal{M}_c (See Figure 1). Since we are interested in the intrinsic rate, not just those that we detect, we weigh each point by the inverse of the spacetime volume in which we are sensitive to it, $w = 1/VT$. Binaries with a higher chirp mass are easier detect, so we do not want to count them as heavily. For a given chirp mass, we are sensitive out to a distance of

$$D(\mathcal{M}_c) = \quad (1)$$

which corresponds to a volume of

$$V(\mathcal{M}_c) = \frac{4}{3}\pi D(\mathcal{M}_c)^3. \quad (2)$$

Multiplying this by the time spent observing, $T = 0.6 \text{ yr}$, gives us the spacetime volume $(VT)(\mathcal{M}_c)$.

2.2 The Likelihood of Fitting Parameters

The likelihood of each fitting parameter is:

$$\ln [P(\{d\}|\lambda)] = \frac{-1}{2} \sum_k \frac{(r(x_k) - r_{model}(x_k))^2}{\sigma_r^2} \exp \left\{ - \int E(x) r(x_k) dx \right\} \quad (3)$$

where p_{smooth} is the smoothing prior, defined to be:

$$p_{smooth} = \exp \left\{ - \int \left[\frac{d^n(r)}{d(x)^n} \right]^2 dx \right\} \quad (4)$$

In theory, p_{smooth} can be any n^{th} derivative. To make our code robust, we define a function that takes n as an argument. The function then calls `numpy.polynomial.polyder()` to find

the n^{th} derivative. Next, we square the n^{th} derivative and integrate it between the minimum and maximum of x . Here we choose $n = 3$.

The likelihood estimator is:

$$\ln P(\{d\}|\lambda) \tag{5}$$

2.3 Model fitting

We employed a Markov Chain Monte Carlo (MCMC) script to fit the coefficients of a polynomial model for $r(M_c)$. The MCMC uses the smoothing prior as defined in section 2.2. We performed a least square fit first in order to obtain the initial guess for the coefficients of the model.

3 Classifying GW Events that have Electromagnetic Counterparts

3.1 Overview

The GW observatory, the Laser Interferometer Gravitational Wave Observatory (LIGO), can provide very rapid mass estimates of candidate GW events. Since most of these detections are mostly binary black holes and electromagnetic followup is extremely expensive, only a few events can be followed up. We have therefore trained a classifier to determine if an event will have a electromagnetic followup. We trained this classifier on the first half of the data; we simply took the mid-way point between the maximum chirp mass for the electromagnetic counterpart group and the nonelectromagnetic counterpart group. This can be seen in Figure 3.

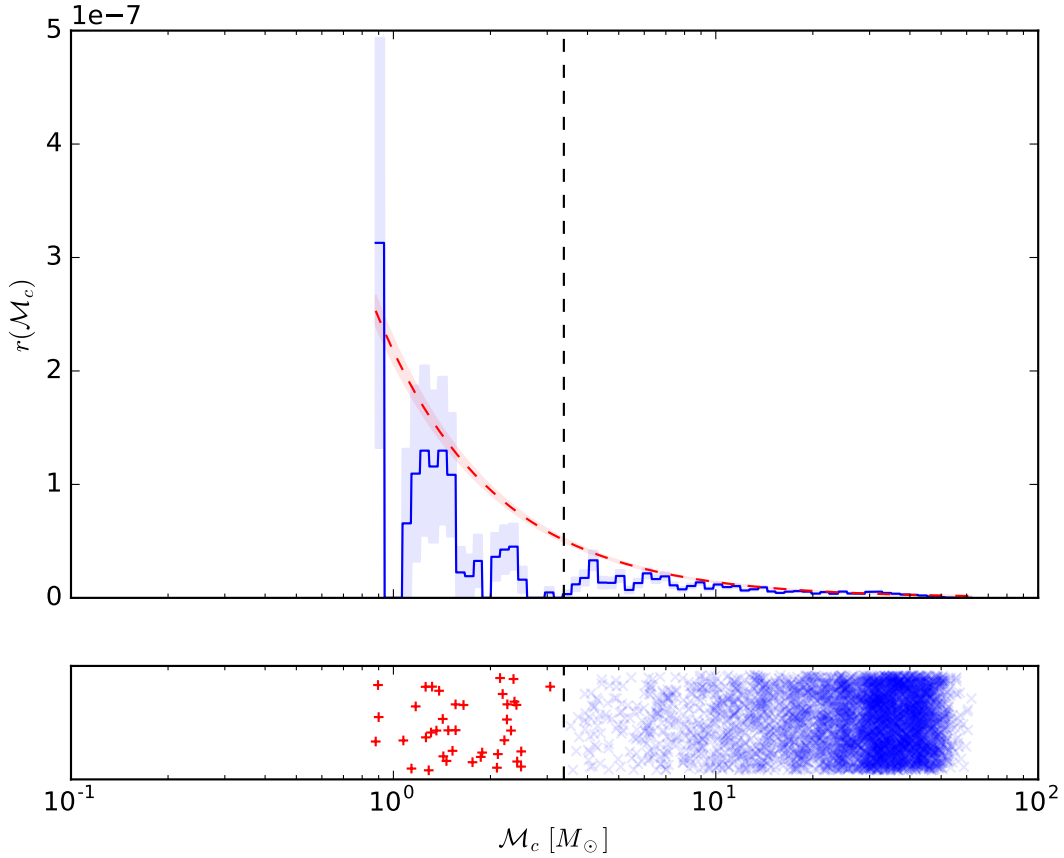


Figure 1: Estimated rate of compact binary mergers, based on 5000 synthetic observations. Blue line is weighted histogram fit. Red curve is power law fit. Shaded regions are $1\text{-}\sigma$ error bars. Vertical dashed line is boundary between events with counterparts and without.

3.2 Method

The classifier was constructed simply by taking the minimum chirp mass event of the other group (no electromagnetic counterparts) and the maximum chirp mass event of the electromagnetic counterpart group and finding the mid-point between those two events. This classifier was trained from the first half of the dataset. This is visualized by the filled points in Figure 3.

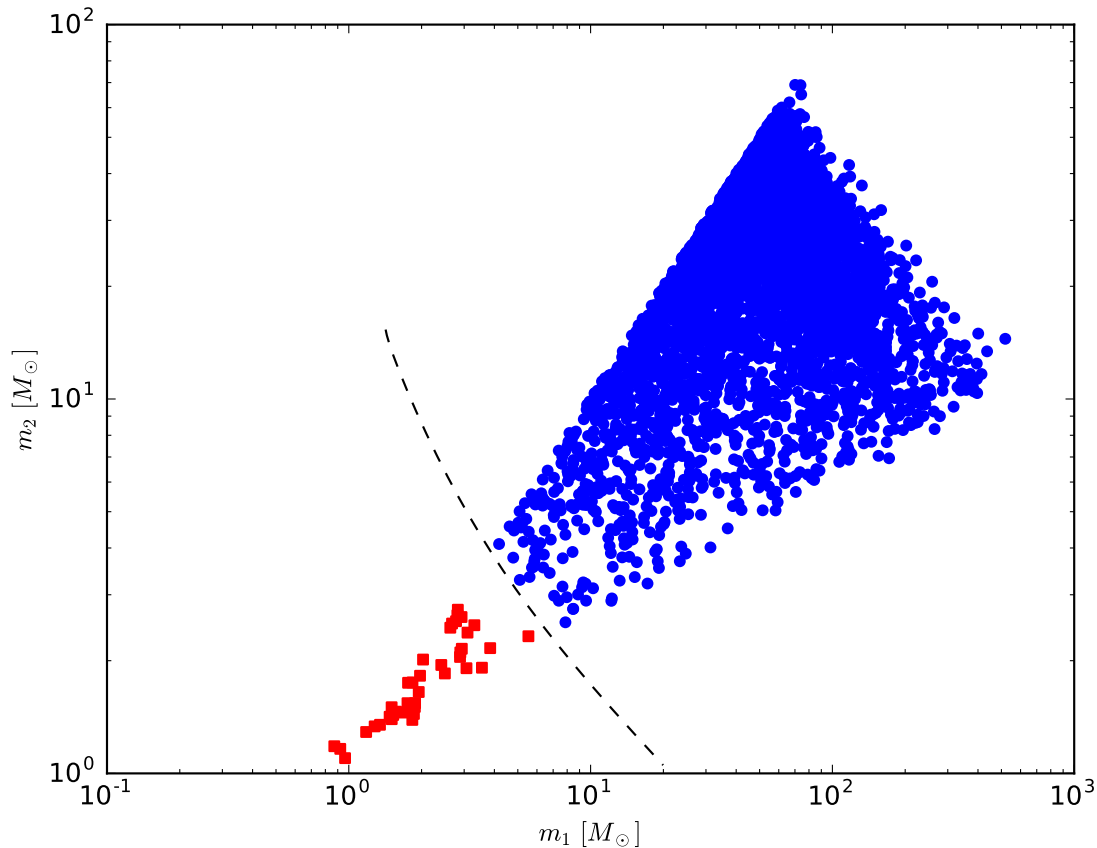


Figure 2

3.3 Results

The classifier correctly classified the two groups without any contamination. More importantly, this was also the case when classifying the full data set. As you can see in Figure 3, the classifier correctly classified the two groups without any contamination. In the Table 1, you can see numerically that the classifier worked completely. Listed are the chirp mass for the maximum electromagnetic counterpart event and the minimum of the other group along with the chirp mass of the line that divides the group. The classifier shows a clear distinction between the two groups.

Figure 1 shows the rate vs the chirp mass with the dividing line from the classifier overplotted. This correlates to the two hump structure in the graph that represents the

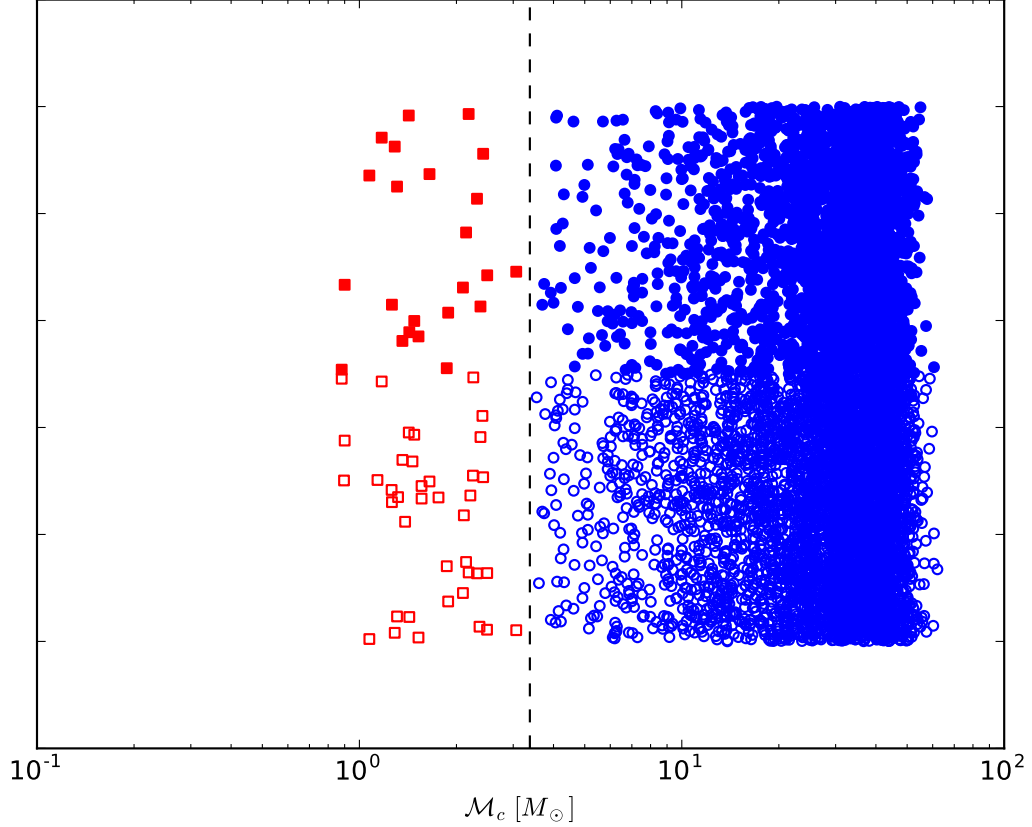


Figure 3: This shows a selected set of the data. The red points represent the events with electromagnetic counterparts, and the blue points represent the events without electromagnetic counterparts. The filled points represent the train dataset, and the open points represent the rest of the data. The vertical line indicates the division between the two groups indicated by the classifier.

two groups (electromagnetic counterparts and nonelectromagnetic counterparts). Figure 2 shows a similar correlation; the dividing clearly separates the two groups in m_1 - m_2 parameter space.

4 Conclusions

We created a classifier to distinguish between GW events with and without electromagnetic counterparts. This classifier was trained using half of the data. The classifier correctly

Table 1: Chirp masses of the maximum event from the electromagnetic counterpart group, the minimum event from nonelectromagnetic counterpart group, and the vertical dividing line.

3.545980	3.065123	3.378711
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classified both groups completely without any contamination for both the training set and the full dataset. Figure 1 and 2 shows the correlation between 1D and 2D mass distributions and the classification of the two groups.

References

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