

Estimating installation parameters of photovoltaic systems in northern latitudes by mathematical model fitting

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Preface

With a strong interest and background in geometry and programming, I have found the real-world problem of panel installation parameter estimation appealing from the beginning. I would like to thank William Wandji and Juha Karhu for their insights on the effects of clouds, snow, and temperature on solar PV installation data. I would also like to thank Luna for being a supportive cat.

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CHAPTER I

Introduction

This thesis examines a specific applied mathematics problem suggested by the Finnish Meteorological institute (FMI). The goal is to determine the geographic location and panel installation angles of photovoltaic solar power installations using only the power output data. The chosen method breaks down the problem of solving the geographic location and panel angles into separate algorithms. In practice, this means that the algorithms used can be less complex as they do not have to solve every unknown variable simultaneously. Additionally, visualizations of individual algorithms are more straightforward.

Using multiple algorithms also splits the parameter space into smaller spaces, thus improving the performance of fitting algorithms. The final benefit is the ability to focus solely on solving the unknown parameters. For example, if the geolocation of a system is known but the panel installation angles are not, the known geolocation can be used for panel angle estimation, resulting in higher prediction accuracy.

The applications of parameter estimation algorithms would be in improving the quality of metadata in solar PV datasets. This could have implications for solar PV research, but the existence of such algorithms poses privacy and security-related questions as well. The realization of these benefits and concerns depends on the accuracy, and to some extent, the ease of use of the algorithms.

A similar study was done by N. Haghadi et al. in 2017 [1]. The 2017 article contains results from five case studies where the standard deviation of longitude prediction errors is less than 1.5° , reaching as low as 0.08° with case study 1-2. The standard deviation of latitude predictions is higher at less than 3.5° with the best result in case study 2-2 with a standard deviation of 1.65° . Azimuth and tilt

predictions are reported as two separate angle error values with azimuth prediction errors reaching values between 4° and 27° and tilt 1.3° to 11.5° . These forementioned results are not directly comparable to the results shown in this thesis due to different datasets, geographic location and local climate, but they provide some perspective.

Another article of relevance written by M.K. Williams et al. in 2012 [2] proposes multiple methods for determining the locations and orientations of solar PV installations. Perhaps the most interesting contribution of the 2012 article is the network approach to determining geographic location. This method relies on a grid of installations with known and accurate geolocation and installation angles. According to the authors, this networked approach works up to a 10-mile accuracy when the grid of known installations is dense around the estimated installation. This could make the network approach preferable for electric companies or institutions with large amounts of data. But as of now, it does not seem usable for FMI.

CHAPTER II

Installations and datasets

A simple PV installation consists of a solar panel or a set of solar panels, an inverter and optionally a battery system. Each of these components can be configured in several different ways and how this is done depends on how installation costs, maintenance, production peak consumption hours, total output and a multitude of other factors are balanced. This variability and difference in demands results in a large variance in PV systems, but some generalizations can still be made. A typical PV installation consists of a set of PV panels which are all installed at the same installation angles. In Finland such installations are typically installed on the roofs of offices or detached houses. As the panel angles in a such system are the same, the solar panels can be modeled as a single larger solar panel in most circumstances and this makes system modeling easier.

Less common but still worth noting are PV installations with mixed installation angles. In these, individual panels or sets of panels have different installation angles. This means that modeling a such system requires the use of multiple surfaces, increasing the model complexity. A typical motivation behind a such installation would be maximizing the available surface area on the roof of a detached house or evening the power output of a PV system through a day. In rare cases PV panels may also be installed on tracking mounts which adjust the panel angles in order to maximize the output of the PV system which further complicates modeling.

The two datasets used in this thesis are from FMI. First dataset is from an installation on top of the FMI office building in Kumpula, Helsinki and it consists of two sets of fixed angle solar panels with identical installation angles. One of these sets or strings is shown in figure 2.1 and a section of data gathered from

this installation is shown in table 2.2. This data gathering is done by the inverter. The Kuopio installation is similar to the Helsinki installation in most respects with the most significant differences being the azimuth angle and different local climate. Shadowing also appears to be more common on the Kuopio installation. Parameters of the two installations are listed on table 2.3.



Figure 2.1: FMI Kumpula solar power installation string.

	Helsinki	Kuopio
Latitude	60.204°	62.892°
Longitude	24.961°	27.634°
Nominal capacity	21 kW	20.28 kW
Panel tilt	15°	15°
Panel azimuth	135°	217°
Elevation	17m	10m

Table 2.3: Parameters for the FMI's Kumpula(Helsinki) and Kuopio PV installations as listed in Böök 2020 [3].

Timestamp[UTC]	Inverter out	Inverter in	String 1	String 2
2015 – 08 – 26 03 : 34	<i>NaN</i>	<i>NaN</i>	0.5	<i>NaN</i>
2015 – 08 – 26 03 : 36	11.1	7.5	2.6	4.9
2015 – 08 – 26 03 : 37	25.4	26.1	9.8	16.3
2015 – 08 – 26 03 : 38	30.7	<i>NaN</i>	<i>NaN</i>	0.4
2015 – 08 – 26 03 : 39	46.4	44.8	20	24.8
2015 – 08 – 26 03 : 40	3.3	<i>NaN</i>	<i>NaN</i>	0.4
2015 – 08 – 26 03 : 41	29.3	18	9.1	8.9
2015 – 08 – 26 03 : 42	33.1	27.4	10.6	16.9
:	:	:	:	:
2015 – 08 – 26 12 : 42	12374.8	14619.1	7152	7467.1
2015 – 08 – 26 12 : 43	15442.2	15482.1	7708.9	7773.2
2015 – 08 – 26 12 : 44	14085.8	12898.7	6387	6511.8
:	:	:	:	:

Table 2.2: A section from FMI’s Kumpula solar site PV production data, only the timestamp and inverter output values are used by the algorithms in this thesis. All power measurements are in watts.

2.1 Simple solar PV model

Let us create a simple model for how a solar PV system functions. This helps in explaining which parameters of a PV system are relevant for later simulating the output of a PV system and where they are used.

If we assume the electronical components such as inverters, cables and the solar panels themselves to be operating at a constant efficiency independent of ambient temperature and other factors, the output of a PV system can be approximated to be $\eta * P$ where η is the combined efficiency of the solar panels, inverter and other electronic components and P is the sum of radiation reatching the solar panel surface. In a cloud free scenario we can assume P to be a function of panel angles, installation size and the position of the Sun. Figure 2.4 shows a visualization of the tree radiation types, each of which contributes to PV production.

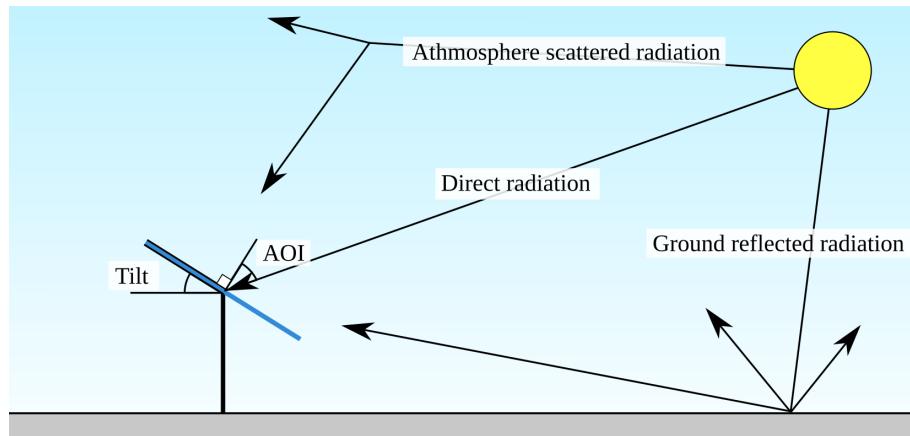


Figure 2.4: Simplified installation diagram for a PV system in ideal conditions.

This figure contains a new angle AOI or the angle of incidence which has not been previously mentioned. AOI measures the angle between the solar panel normal vector and the incoming sunlight. As AOI approaches zero, from the perspective of direct radiation the relative surface area of the solar panel is at it's maximum resulting in higher amount of radiation per unit of surface area. And when AOI is 90° or greater, the panel surface does not receive direct sunlight. The observable position of the Sun which is required for calculating the AOI is a function of date, time and geographic location.

Note that this simple model does not take panel reflectivity or the loss of efficiency

due to increased panel temperature into account and both of these topics will be discussed in later sections.

2.2 Visualizing the data

The Figure 2.5 contains a 3D point cloud generated by plotting one year of data from the FMI Helsinki dataset, revealing visible structures in the data. The clearest structure in the 3D plot is the pattern formed by the first and last non-zero power minutes and this is later used for geolocation estimation. The second structure is the dome-like shape of the point cloud. This shape can be examined by taking one-day slices from the dataset and plotting them individually, as shown in Figure 2.6. These slices are used in later sections for panel installation angle estimation.

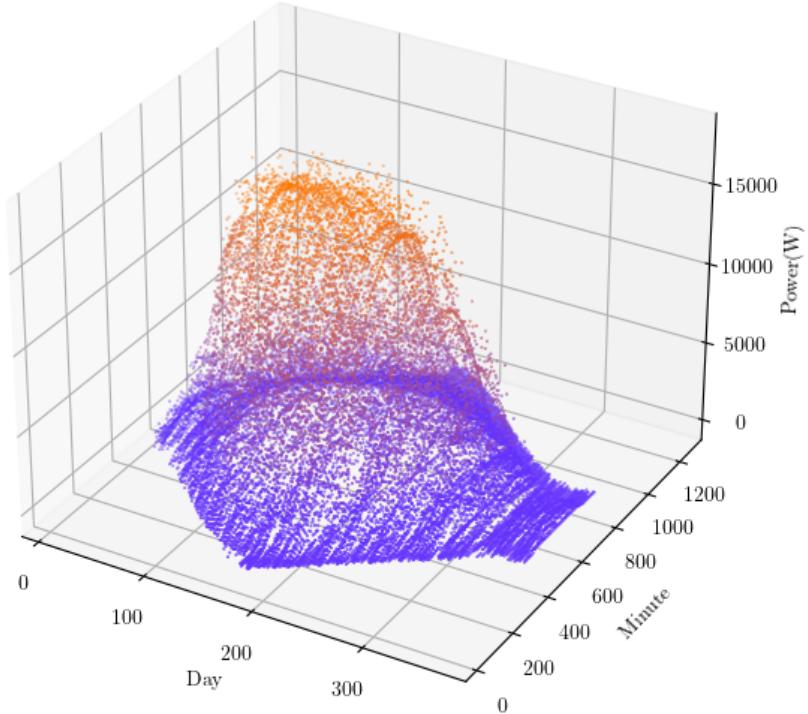


Figure 2.5: One year of data from FMI Kumpula installation as a 3D point cloud.

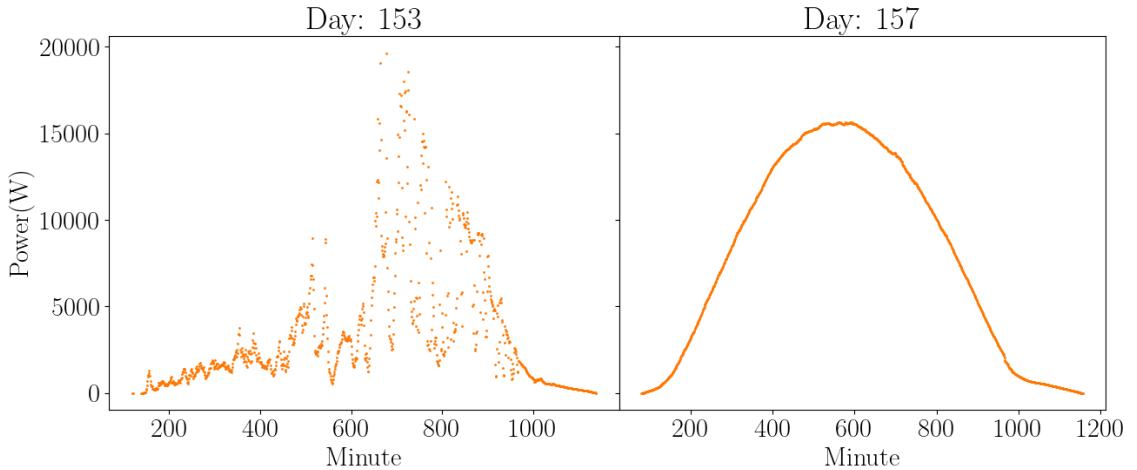


Figure 2.6: Two days from FMI Kumpula dataset with different characteristics.

The presumably cloud free day 157 in Figure 2.6 forms a skewed bell curve -like shape. This figure has a few notable features. The first and last non-zero minutes which occur somewhere around 100 and 1150 minutes can be useful for estimating the geographic location as they should correlate with sunsets and sunriseds. The peak of the plot at around 575 minutes may be a point in time during which AOI is rearching a daily minimum. And finally the knee section beginning near 1000 minutes may represent a point in time where AOI reaches 90° and direct sunlight is no longer contributing to PV production.

For the cloudy day these same traits are much harder to measure. The first and last non-zero minutes appear to be nearly the same as suggested by earlier Figure 2.5 but peak production minutes or transitions from indirect to direct sunlight are harder to clearly spot. The shape of the scatter plot is also irregular which may limit its usability. This would suggest that specific days may be preferable for analysis purposes.

2.3 Clear day detection algorithm

The ability to select days with clean power output data can be helpful for analysis purposes. The amplitude noise present in earlier Figure 2.6 would indicate that algorithmic cloud free day finding could be done by measuring how much sequential power measurements deviate from oneanother on average. The following listing describes one such algorithm.

1. Split dataset into individual days based on timestamps.
2. Create a copy of the power measurements for each day and process this copy with a low pass filter algorithm.
3. Calculate the difference between the original power values and the filtered power values. This delta value increases when high-frequency noise is present.
4. Discard days with a delta value higher than a set threshold.

The mathematically non-trivial parts here are threshold selection, difference measurement and low-pass filtering. Low-pass filtering is a term borrowed from the field of signal processing, and it refers to any algorithm that removes frequencies higher than a given limit from a signal, allowing lower frequencies to pass.

Here the filtering is done with discrete Fourier transformations (DFT) and inverse discrete Fourier transformations (IDFT). When a list of numbers is used as the input of DFT, the output is a list of ordered complex numbers, each of which represents a sine wave of a certain frequency, phase and amplitude. The sum of these wave equations forms a continuous approximation of the input values and by sampling the continuous representation, the continuous trigonometric approximation can be transformed back into discrete values. However if the complex numbers are adjusted before the IDFT operation, frequencies can be selectively modified. This means that DFT and IDFT can be used for frequency specific modification of numerical lists, low-pass filtering being one of the possibilities. In this case the low-pass filtering was accomplished by zeroing out complex numbers which do not correspond to the 6 longest frequencies, the resulting smoothening can be seen in Figure 2.7.

While this process is somewhat complicated, Fourier transformations are not the only tool for creating low pass filters. Similar results can also be achieved by locally

averaging each power value to be the average of nearest k values. Discrete Fourier transformation based methods do however have an advantage in their universality. If the 6 or 7 or n longest frequencies can be determined to be a good low-pass filter for PV power measurements, then these same frequencies should result in similar outputs no matter the temporal resolution of the power measurement data. Whereas a method based on local averages would require a different window size depending on measurement intervals.

The second component is not as complicated as the low pass filtering operation. Measuring the delta between a filtered and unfiltered set of measurements can be done by computing the discrete curve length or as was done here, measuring the absolute average deviation between filtered and unfiltered power measurement as per Equations 2.1-2.5.

$$Power = [p_0, p_1, p_2, \dots, p_n] \quad (2.1)$$

$$Power_{filtered} = [f_0, f_1, f_2, \dots, f_n] \quad (2.2)$$

$$Power_{delta} = [|p_0 - f_0|, |p_1 - f_1|, |p_2 - f_2|, \dots, |p_n - f_n|] \quad (2.3)$$

$$\text{delta}_{avg} = avg(Power_{delta}) \quad (2.4)$$

$$\text{delta}_{norm} = \text{delta}_{avg}/\max(Power) \quad (2.5)$$

The final component is threshold selection. The intermittent value delta_{avg} describes the average wattage difference between measured and low pass filtered measured power values. By definition, this delta value is dependent on noise and installation size, limiting its usability. A noise only -delta value can be calculated by normalizing the delta with the $\max(Power)$. The resulting delta_{norm} should now be comparable between installations of different sizes. Choosing to reject every day for which delta_{norm} value is higher than 0.05 would eliminate days with higher than 5% normalized noise.

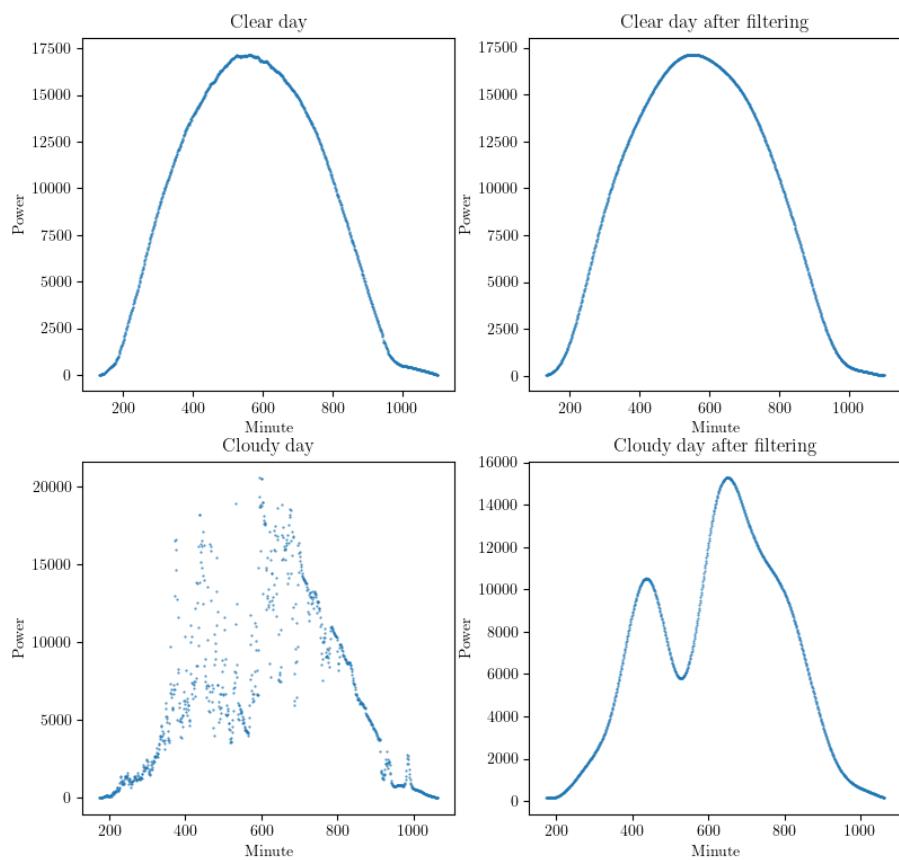


Figure 2.7: Cloud free day finder low pass filtering phase.

CHAPTER III
PV output simulation

The earlier PV model presented in section 2.1 assumes PV output to be $\eta * P$. The value of η can be difficult to accurately approximate due to differences in PV panel technologies, inverter efficiencies and other factors, but values in the range of 0.15 to 0.20 can be used as placeholders if a value is required. However for model fitting purposes η does not need to be known as the if the shape of the curve is correct, the amplitudes of the curves can be matched.

The value P has to be estimated. Physically accurate models are fairly complicated but luckily there are pre-existing tools made for solar PV modeling. The python library PVlib contains functions for estimating solar angles and different radiation types, the most relevant as of now being the POA irradiance. This plane of array irradiance is the sum of radiation reaching a $1m^2$ sized plane in clear sky conditions with given location, angles and datetime. A section of PVlib generated POA values are shown in Table 3.1.

Timestamp[UTC]	Minute	POA(W/m)
2018 – 05 – 30 00 : 00	0	0.0
2018 – 05 – 30 00 : 01	1	0.0
2018 – 05 – 30 00 : 02	2	0.0
:	:	:
2018 – 05 – 30 07 : 34	454	800.691861
2018 – 05 – 30 07 : 35	455	802.110516
2018 – 05 – 30 07 : 36	456	803.517424
:	:	:
2018 – 05 – 30 23 : 57	1437	0.0
2018 – 05 – 30 23 : 58	1438	0.0
2018 – 05 – 30 23 : 59	1439	0.0

Table 3.1: One day of simulated clear sky plane of array irradiance values. Note that the minute column is added to the table for convinience and it is redundant as minutes can be read from the timestamps.

3.1 PVlib POA evaluation

Before the PVlib POA simulations are used for parameter estimation, the simulated values should be compared with known PV measurements. Figure 3.2 contains 6 days from FMI Kumpula dataset year 2017 plotted on orange and PVlib simulated POA values with the parameters of the Kumpula installation. Power values of the POA simulations were multiplied by 19 in order to match the magnitudes of both plots.

This plot shows that in presumably cloud free conditions during days 123, 136 and 141 the POA values and real PV output curves have matching shapes with minor deviations. And during days 70, 150 and 240 clouds, snow or other factors have caused noise into power generation resulting in deviation between the measurements and the PVlib POA simulation. This noise appears to result in both increases and decreases in expected power generation. Decreases are fairly intuitive as clouds can block direct solar radiation from reaching panel surfaces but positive noise can be unintuitive and it has implications for model fitting. Well positioned clouds can reflect additional sunlight onto the panels or snow near panels may increase ground

reflectivity. This noise characteristic means that local maximums or minimums of power measurements can not be used to approximate the expected power output in normal conditions.

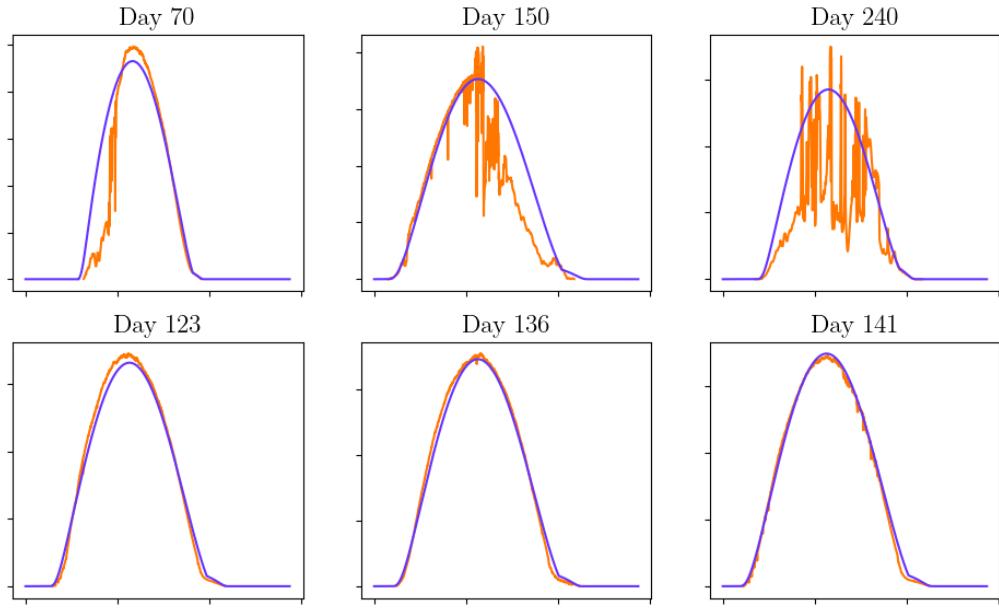


Figure 3.2: Plots comparing POA simulations with real PV measurements from FMI Kumpula dataset year 2017.

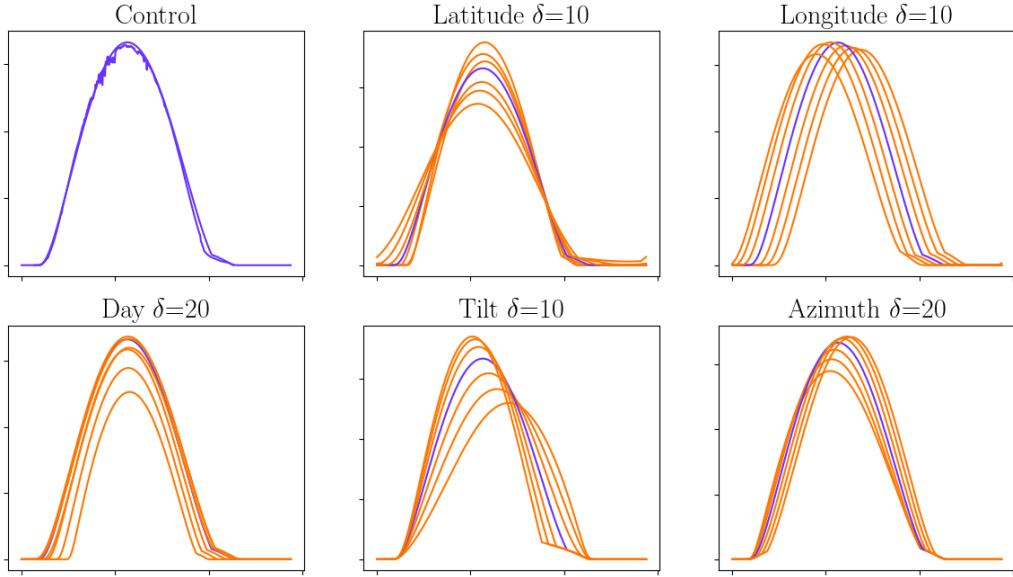


Figure 3.3: Influence of changes in PVlib simulation parameters on generated power output curves. Control shows FMI Helsinki measurements and simulation with the same parameters as the Helsinki installation. Simulated power values are multiplied by 19 in order to match values on y-axis.

By varying the different simulation parameters as shown in Figure 3.3, we can examine the relationships between parameters and power generation curves. This can help us understand if there are usable patterns in the data. In the best case scenario each of the simulation function inputs would affect one measurable property in the irradiance plots and their relationship would be bijective. To give an example, if the peak power minute was isolated from all other parameters than the longitude and the relationship between longitude and peak power minute was linear, it would be possible to solve the peak power minute to longitude function with just a few plane of array irradiance simulations.

In the exact opposite case where every measurable property of irradiance plots is affected by every input parameter, solving the parameters would be much harder or even impossible. For example if all of the parameters influenced the same traits to different extents and the system was not bijective, multiple parameter combinations could result in the same simulated power graph. In a such system there would not be a single solution but rather a set of possible solutions.

The problem of solving installation parameters would appear to be somewhere in between the two extremes. The longitude parameter would seem to shift the curve along the time axis whereas tilt and azimuth parameters do not affect the first or last non-zero minutes but they do affect the shape of the curve. Observations of parameter to trait interactions are listed in Table 3.4.

Parameter	Traits affected
Latitude	Shape, first and last minute times
Longitude	First and last minute times
Tilt	Shape
Azimuth	Shape

Table 3.4: Function input to observed trait table.

3.1.1 Influence of different longitudes

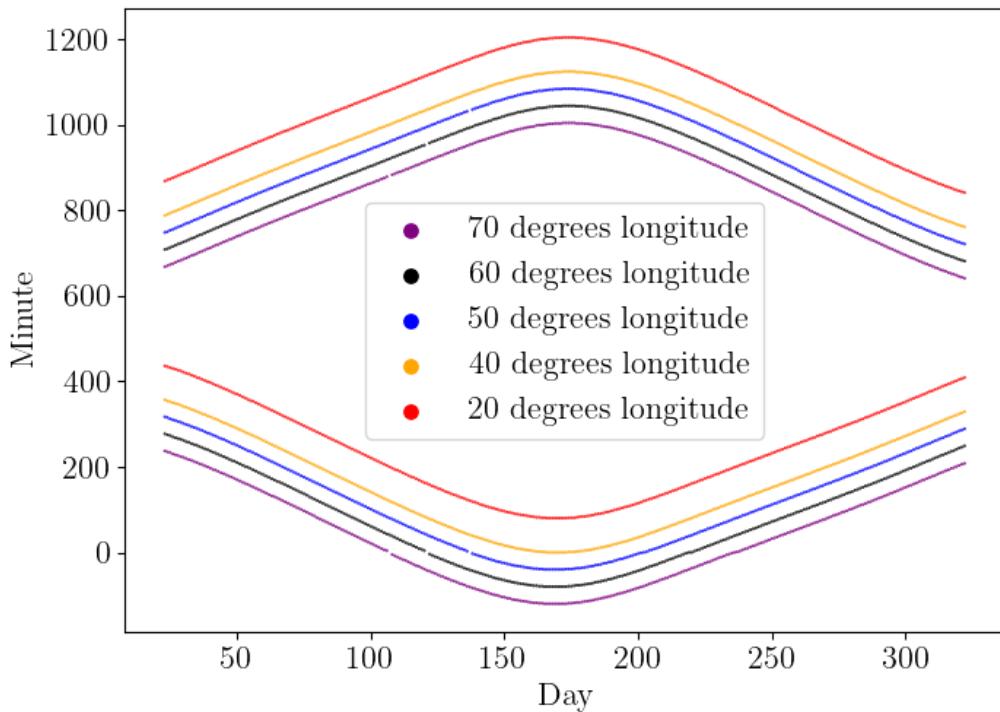


Figure 3.5: First and last non-zero minutes of each day from year long simulations at different longitudes.

Based on earlier observations listed in Table 3.4, solving the longitude of installations would seem like a sensible starting point. The figure comparing the effects of different parameters seemed to suggest that the relationship between longitude and significant minute times is very close to linear and the same is seen here in Figure 3.5. In Hagdadi 2017 [1] and in Williams 2012 [2] this relationship was used in order to determine the geographic longitude. The algorithms used by both of the articles relies on calculating an approximation for the time of the solar noon based on the average of the first and last minutes, this solar noon minute is then translated into a geographic longitude coordinate.

3.1.2 Influence of different latitudes

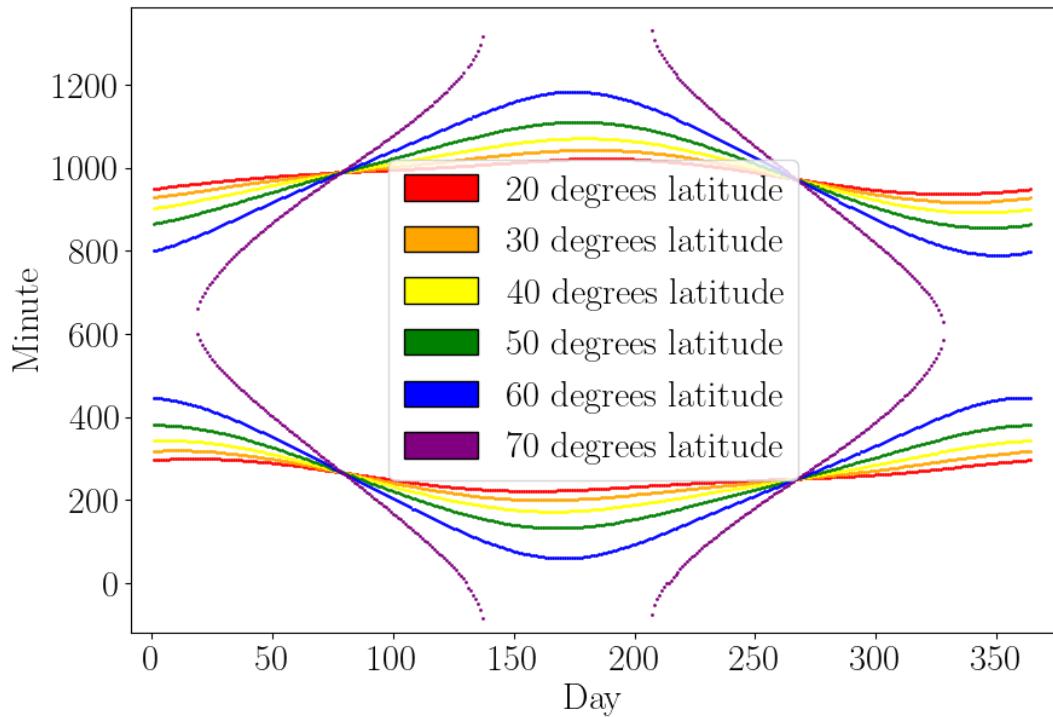


Figure 3.6: First and last non-zero power minutes of each day from year long POA simulations at different northern latitudes.

The latitude simulations in Figure 3.6 show that the day length stays fairly consistent at lower latitudes but with latitudes of 50°N and higher, the day length becomes significant. This variation in day lengths can be consistently measured near equinoxes as day lengths are always well defined.

3.2 Increasing the accuracy of solar PV simulations

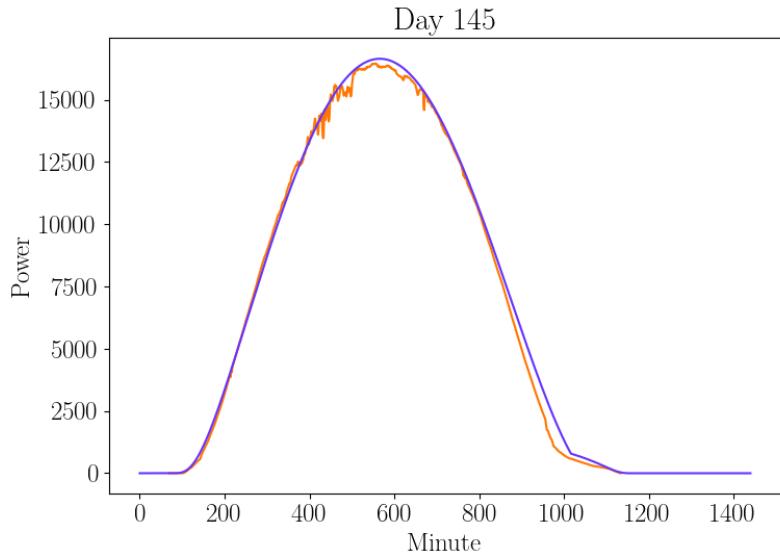


Figure 3.7: Figure comparing Kumpula PV measurements on orange and PVlib simulated clear sky POA irradiance(purple). POA power values were scaled with a multiplier in order to match curve magnitudes.

Figure 3.7 and the earlier multi-day comparison Figure 3.2 both suggest that PVlib POA estimates are a fairly good approximation of clear sky PV output. However when examined more closely, there are some deviations between the POA simulations and real PV outputs. The PVlib simulations suggest a higher than achieved output near 580 and 1000 minutes.

At 1000 minutes this is likely a result of reflective losses as the radiation reaching a panel surface is not equal to absorbed radiation and with south-east facing panels the direct component of solar radiation would be reaching the panel surfaces at shallow angle, increasing the fraction of radiation reflected away from the panel surface. This means that a more accurate PV model would have to take panel reflections into account in order to correct for this discrepancy between the model and the PV measurements.

Near 580 minutes the panels are absorbing a high amount of radiation. If we assume the solar panels to be operating at approximately 20% efficiency, 80% of incoming radiation is transformed into heat and the efficiency of PV panels and

many other electrical devices is dependent on operating temperature. This would also explain the negative noise present during the peak output hours as gusts of wind may be enough to cool the panels enough to cause noise patterns in the PV data. This noise is difficult to estimate, but panel temperature modeling could be included in a better PV model and this may result in higher estimation accuracy near peak production times.

3.2.1 Improved PV model

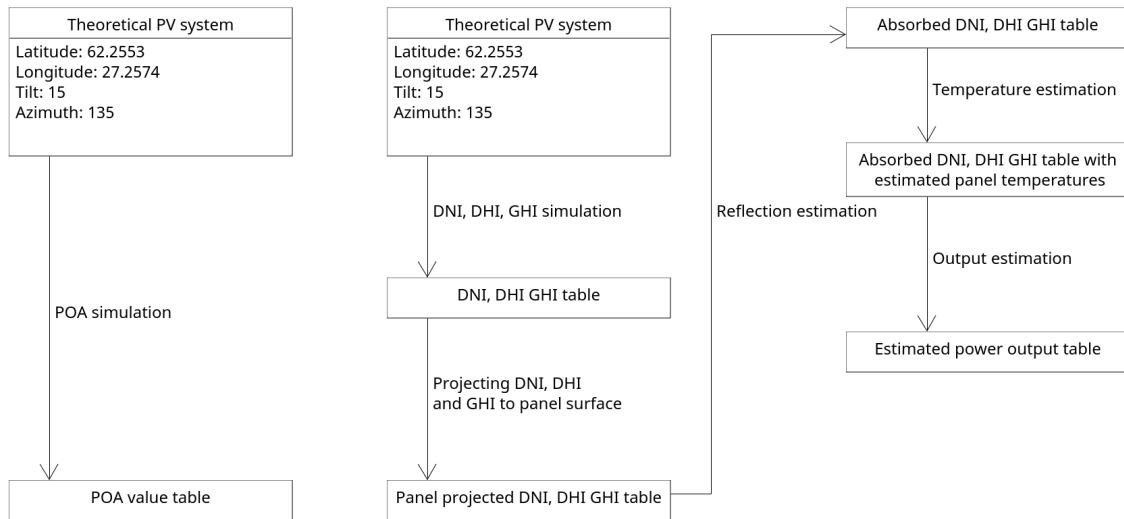


Figure 3.8: Original POA-based PV model and the improved model with reflection and temperature estimation.

Figure 3.8 contains a comparison of steps required for PVlib POA based solar PV output estimation and the steps of a more physically accurate PV model. While the steps required by the methods may seem different, they are very similar.

When the PVlib POA model is used, PVlib internally calculates estimates for DNI, DHI and GHI, each of which is a component of solar irradiance, and then transposes the irradiances to the solar panel surface. The sum of the radiation types after transposition is the plane of array irradiance or POA. However if we use PVlib to generate a table of the three radiation components, we can then transpose the radiation types to the panel surface with our own models for transpositions, reflection estimation, panel temperatures and the output of the PV system. The three irradiance types were mentioned earlier as direct radiation, atmosphere scattered radiation and ground reflected radiation but in meteorology they go by the following names.

- DNI is direct normal irradiance which represents direct sunlight received by a 1m^2 -sized plane with AOI of 0 degrees. Note that DNI consists of only direct sunlight and atmospherically scattered light or ground reflections are not components of DNI.

- DHI is diffuse horizontal irradiance which represents the irradiance reaching a shaded $1m^2$ -sized plane with tilt of 0 degrees. DHI represents atmosphere scattered radiation and DHI can be measured with an instrument which blocks direct sunlight.
- GHI is global horizontal irradiance and it represents the amount of irradiance per $1m^2$ -sized plane with tilt of 0 degrees. GHI is made up of direct irradiance and atmosphere scattered irradiance. GHI combined with albedo can be used to estimate the ground reflected irradiance.

3.2.2 Panel surface transpositions

The first step in the improved model is panel surface transposition. Sandia National Laboratories, the original author of PVlib, suggests the following two equations for DNI and GHI transposition and five alternative models for DHI transposition. Perez 1990 model [4] was chosen for DHI transposition due to the use of the model by a co-worker at FMI and inclusion in Sandia suggested projection models [5]. Chosen DNI and GHI transposition models are shown below. DHI Perez transposition model utilizes number tables and as a function included in PVlib was used, the mathematical notation for the method was left out. See thesis source code for the implementations.

DNI transposition [6]

$$DNI_{poa}(DNI, AOI) = DNI * \cos(AOI) \quad (3.1)$$

Where AOI is the angle of incidence.

GHI transposition [7]

$$GHI_{poa}(GHI, albedo, tilt) = GHI * albedo * \frac{1 - \cos(tilt)}{2} \quad (3.2)$$

Where

$albedo = 0.151$. This represents ground reflectivity near the solar PV installation. The value varies significantly depending on the area, season and time of day. A fixed value is used as terrain, weather and snow cover are assumed to be unknown.

$tilt$ is the tilt angle of the PV installation, this is 15° for both FMI installations.

3.2.3 Reflection estimation

The following Solar irradiance reflection equations originate from 2001 Martin and Ruiz paper [8]. This study used a physical contraption in which solar panels were rotated under a light beam from a solar simulator. Values from the reflective losses equations represent the fraction of irradiance reflected away from the panels e.g. a $DNI_{reflected}$ value of 0.24 would tell us that 24% of irradiance is lost due to reflections and 76% is absorbed.

Reflected fraction of DNI

$$DNI_{reflected} = \exp\left(-\frac{1}{a_r}(c_1 p_1 + c_2 p_1^2)\right) \quad (3.3)$$

Reflected fraction of DHI

$$DHI_{reflected} = \exp\left(-\frac{1}{a_r}(c_1 p_2 + c_2 p_2^2)\right) \quad (3.4)$$

Reflected fraction of GHI

$$GHI_{reflected} = \frac{\exp(-\cos(\alpha)/a_r) - \exp(-1/a_r)}{1 - \exp(-1/\alpha_r)} \quad (3.5)$$

Where

$a_r = 0.159$ Empirical reflectance constant for polycrystalline silicon solar PV panels.

$c_1 = \frac{4}{3\pi}$ Fitting parameter 1.

$c_2 = -0.074$ Fitting parameter 2.

$p_1 = \sin(\text{tilt}) + \frac{\text{tilt} - \sin(\text{tilt})}{1 - \cos(\text{tilt})}$

$p_2 = \sin(\text{tilt}) + \frac{\pi - \text{tilt} - \sin(\text{tilt})}{1 + \cos(\text{tilt})}$

3.2.4 Absorbed irradiance

Now that reflective losses can be estimated for each radiation source, the amount of radiation absorbed by the PV panels can be calculated with the following equations. Resulting values are in watts.

Absorbed DNI

$$DNI_{absorbed} = DNI_{poa}(1 - DNI_{reflected}) \quad (3.6)$$

Absorbed DHI

$$DHI_{absorbed} = DHI_{poa}(1 - DHI_{reflected}) \quad (3.7)$$

Absorbed GHI

$$GHI_{absorbed} = GHI_{poa}(1 - GHI_{reflected}) \quad (3.8)$$

Absorbed radiation

$$POA_{absorbed} = DNI_{absorbed} + DHI_{absorbed} + GHI_{absorbed} \quad (3.9)$$

3.2.5 Panel temperature estimation

The previously calculated $POA_{absorbed}$ can be used to estimate the panel temperature by using the following King et al. 2004 model [9]. This King model requires air temperature, wind speed and absorbed radiation to be known. As air temperature nor wind speed are known, dummy values can be used instead. Note that the King model does not take thermal capacity of the panels into account and thus the model responds to changes faster than the real temperature of PV panels.

Panel temperature

$$T_{panel} = T_{air} + POA_{absorbed}e^{C_a + C_b * w} \quad (3.10)$$

Where

T_{air} is air temperature in $^{\circ}C$.

$C_a = -3.47$ Model fitting constant.

$C_b = -0.0594$ Model fitting constant.

w is wind speed at panel elevation.

As air temperature and wind speed are unknown, dummy values have to be used for both variables. Air temperatures of $20^{\circ}C$ and wind speed of 2m/s were used as dummy values.

3.2.6 Output estimation

Now that the absorbed radiation and panel temperature are known, the output can be estimated with the following Huld et al. 2010 model [10].

PV output model

$$P_{output} = P_{rated} P_n \eta_{rel} \quad (3.11)$$

Where

P_{rated} is the rated power of the system in kW.

$$P_n = \frac{POA_{absorbed}}{1000}$$

$$\eta_{rel} = 1 + k_1 \ln(P_n) + k_2 \ln(P_n)^2 + T_{diff} k_3 + k_4 \ln(P_n) + k_5 \ln(P_n)^2 + k_6 T_{diff}^2$$

In η_{rel} the variables k_1 to k_6 are fitting parameters and T_{diff} is $T_{panel} - 25^\circ C$.

k1	k2	k3	k4	k5	k6
-0.017162	-0.040289	-0.004681	0.000148	0.000169	0.000005

If the variable P_{rated} is not known, it can be estimated by using an area matching algorithm.

Area matching

$$P_{rated} \approx \sum_{t=0}^{1439} P_{measured}(t) / \sum_{t=0}^{1439} P_{simulated}(t) \quad (3.12)$$

Where

$P_{simulated}$ is simulated output of a $1m^2$ installation for a clear day.

$P_{measured}$ is the measured PV output during a clear day.

Visualization of area matching by using a single cloud free day from Helsinki dataset and a PVlib POA simulation is included in Figure 3.9.

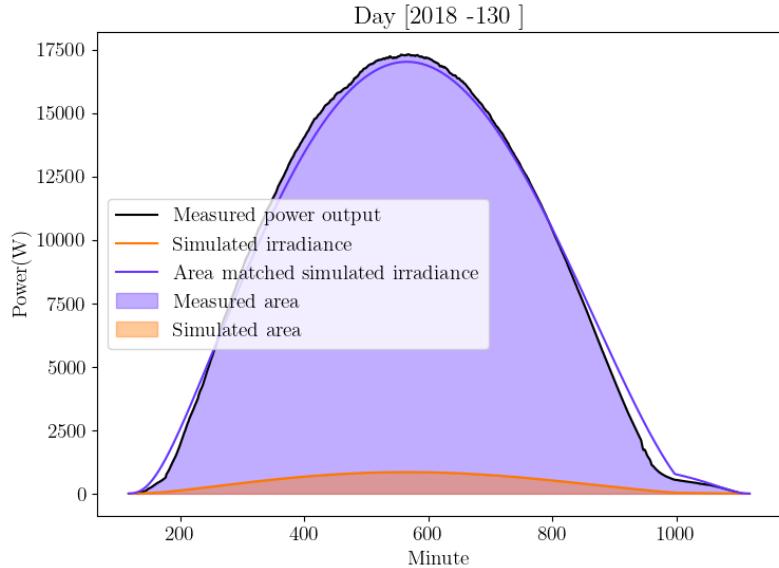


Figure 3.9: Measurement data from FMI Helsinki dataset, POA irradiance simulation was computed with FMI Kumpula coordinates and installation angle parameters.

3.2.7 Model improvement results

The models can be compared with the original measurement data with the following Equation 3.13.

Model delta

$$Delta_{model} = \sum_{t=0}^{1439} |P_{measured}(t) - P_{simulated}(t)| / 1440 \quad (3.13)$$

The $Delta_{model}$ value represents a normalized per minute deviation between the tested model and the measured data. Normalization by division with 1440 results in low delta values as this assumes that the system is constantly generating power regardless of the time of day and thus the resulting delta value is somewhat misleading but for comparison between models these values are still useful.

Delta values for detected cloud free days for FMI Helsinki installation are on average 244W for the POA model and 145W for the improved model based on a sample of 22 cloud free days. For FMI Kuopio the POA model achieved a delta of 487W whereas the delta of the improved model was 320W with a 42 day dataset. This would indicate that the improved PV model is a closer approximation for clear sky PV output than the PVlib POA model.

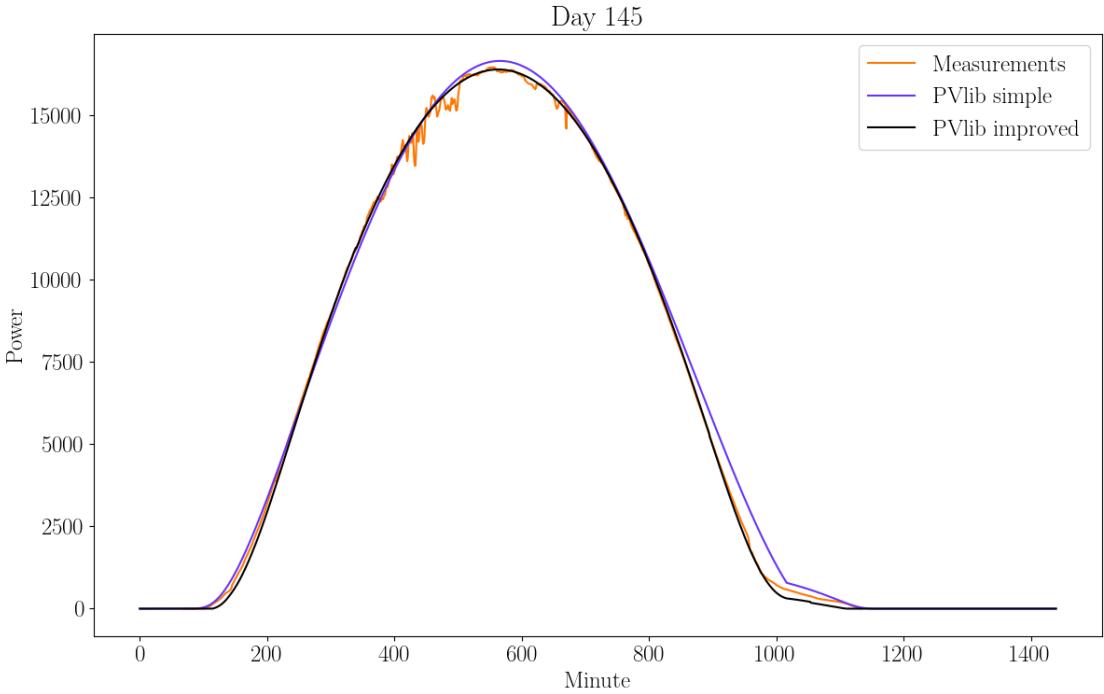


Figure 3.10: Comparison of PVlib POA and the improved PV power generation model on a day from Helsinki dataset.

Figure 3.10 shows what the improvement looks like in a best case scenario. The improved model can be seen to approximate the power output with comparable or higher accuracy than the POA model during the tested day. This is especially notable during peak production hours and near 1000 minutes where AOI is high.

Despite the good performance of the improved model, the POA model still has uses due to faster computational speed. If for example only the first and last non-zero minutes matter, it may not make sense to use a physically more accurate model. In the next chapters which of these models is being used will be mentioned on case by case basis.

CHAPTER IV

Estimating geographic location

In order to evaluate the performance of longitude and latitude estimation functions, it may prove useful to be able to translate the error values from degrees to kilometers. The following Equations 4.1 and 4.2 can be used to approximate the deltas of longitude and latitude estimation functions in kilometers. Note that these functions are only approximations as they rely on the assumption that the Earth is a perfect sphere and not an irregular ellipsoid.

Latitudinal distance to kilometers(Distance on North-South axis)

$$Distance_{latitudinal}(lat_d) = (40000\text{km}/360^\circ) * lat_d \quad (4.1)$$

Where

lat_d is the distance between two points in degrees latitude.

40000km is an approximation for Earths circumference.

Longitudinal distance to kilometers at given latitude(Distance on East-West axis)

$$Distance_{longitudinal}(lon_d, lat) = (40000\text{km}/360^\circ) * \cos(lat) * lon_d \quad (4.2)$$

Where

lon_d is the distance in degrees longitude.

lat is the latitude for which the distance is calculated.

As long as the deviations are small enough and highly accurate error values are not needed, the total error in absolute terms can be estimated by using the latitudinal

and longitudinal distances as the x and y coordinates on a cartesian plane and computing the euclidean distance between the origin and resulting point.

4.1 Estimating geographic longitude

As mentioned in Sections 3.1.2 and 3.1.1, the geographic location of a PV system has a strong correlation to the timing of the first and last non-zero measurements of each day whereas the influence of tilt and facing parameters appears to be nonexistent. The relationship would seem to be so clear that without further analysis it would be tempting to use fairly simplistic mathematical models for these estimations. The following longitude estimation Function 4.3 can be derived with two assumptions. These assumptions are that solar noon occurs at 12:00 or 720 minutes at longitude 0° each day and at 6:00 or 360 minutes at 90°. Rest of the values can then be linearly interpolated. Note that here solar noon refers to the midpoint between the first and last non-zero minute which is different from astronomical solar noon which occurs nearly at the same time.

As only the first and last non-zero minute times are relevant for longitude and latitude estimation, PVlib POA model is used for both longitude and latitude estimation.

Naive solar noon to longitude equation

$$\text{Longitude}(sn) = 180^\circ - \frac{360^\circ}{1440} * sn \quad (4.3)$$

Where sn is the approximated solar noon minute calculated by taking the average of first and last non-zero power generation minute of a day.

The simplicity of Equation 4.3 makes the equation appealing, but the assumption of solar noon occurring at 720 minutes should still be verified. In Figure 4.1 solar noons can be seen to occur at around 720 minutes at longitude 0° but they can also be observed occurring 15 minutes earlier or later than that. The cause for this pattern is a combination of Earth's axial tilt and elliptical orbit [11]. This 15 minute delta would translate into an error range of $(\pm 15/1440) * 360^\circ = \pm 3.75^\circ$ degrees or approximately $\pm 200\text{km}$ at the latitudes of Helsinki according to the Equation 4.1.

Knowing that the PV installation is within a 400 kilometer wide slice may in some cases be enough for determining the country in which the PV installation

is located in, but for most other purposes this level of accuracy is unlikely to be valuable. Fortunately the naive model can be improved upon by taking the solar noon timing variation into account.

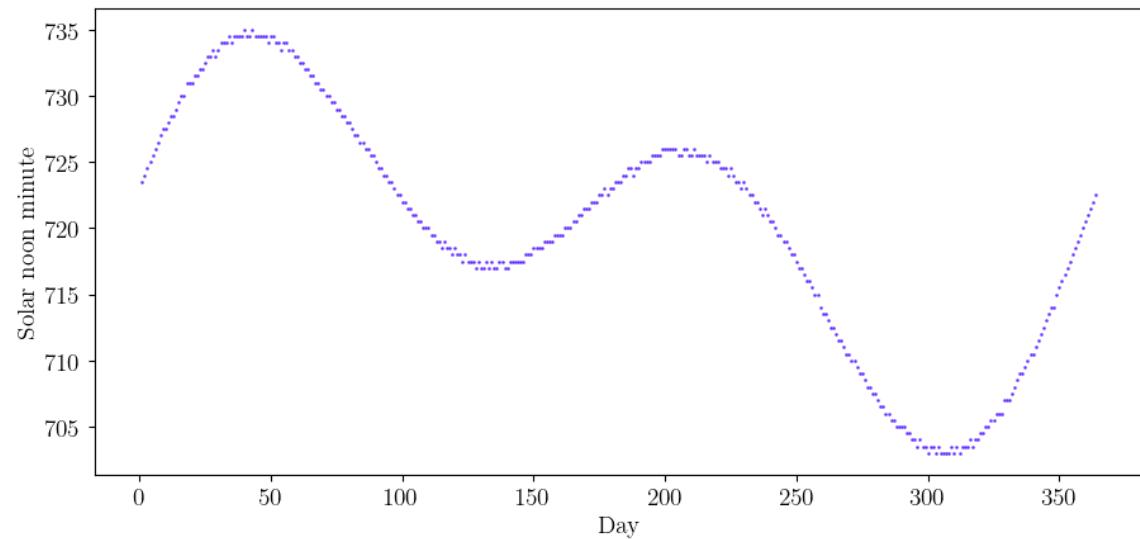


Figure 4.1: Approximations of solar noon minutes based on PVlib POA function at longitude 0° for year 2023.

Improved longitude estimation function

$$\text{Longitude}(sn) = \frac{360}{1440} (sn_{poa} - sn) \quad (4.4)$$

Where sn is the solar noon estimate based on measurement data and sn_{poa} is the simulated solar noon at 0° longitude. The new function parameter sn_{poa} compensates for the variation seen in Figure 4.1.

The improved algorithm should no longer have a systematic error of up to 15 minutes after the difference between solar time and UTC time variation has been taken into account. In addition to correcting for the irregular solar noon timing, the algorithm can be improved even further by using it on larger sections of data and averaging the results, or alternatively, the algorithm could be applied only on selected cloud-free days where the expected errors are likely to be smaller.

If the unfiltered multi-day approach is used, choosing the right day range is crucial. If the range is too narrow, a single outlier value can distort the results significantly, however if the whole year is used, certain periods of the year may contain more noise than others and thus their use could decrease the accuracy of the results. The two scatterplots in Figure 4.2 show that the data quality from the very first and last days of the year seem to be significantly worse than the data from the longest days of the year. Based on these visualizations, days inside the range 100th to 280th would seem best suited for first and last minute sensitive analysis algorithms for both Helsinki and Kuopio installations.

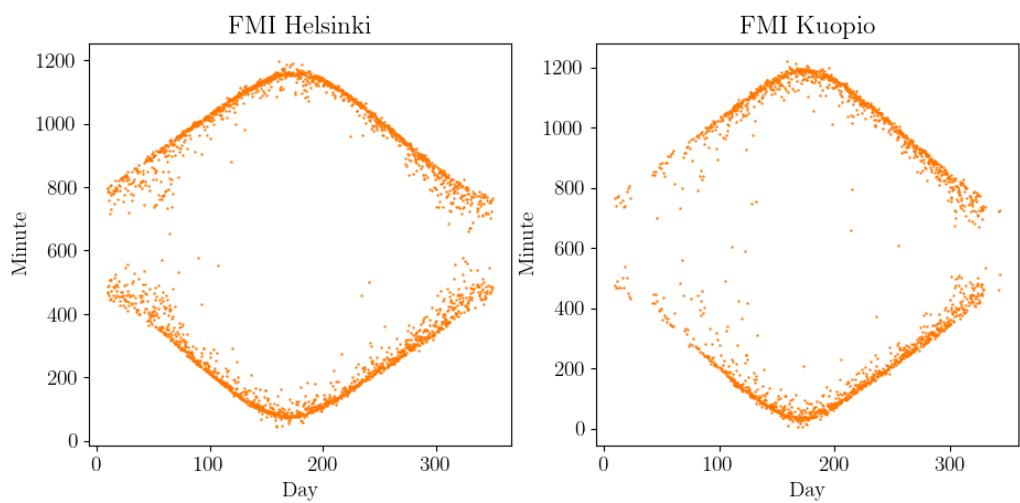


Figure 4.2: First and last non-zero power minutes of each day during years 2017 to 2021 from FMI Helsinki and Kuopio datasets.

4.1.1 Longitude estimation results

The longitude estimation algorithm was tested on a day range of 125th to 250th for each year in the FMI Helsinki dataset and the results are seen in figure 4.3. By taking the mean of the longitude predictions for each year in the dataset, the resulting longitude estimates varied between -0.804° to 0.0766° from known installation longitudes. Results for the Kuopio dataset shown in figure 4.4 were similar with mean delta range of -0.1042° to 0.9424° .

In kilometers the year with the highest mean deviation of -0.804° in the Helsinki dataset corresponds to an error of approximately 45 kilometers as per 4.2. With Kuopio results the highest delta of 0.9424° is 49 kilometers respectively.

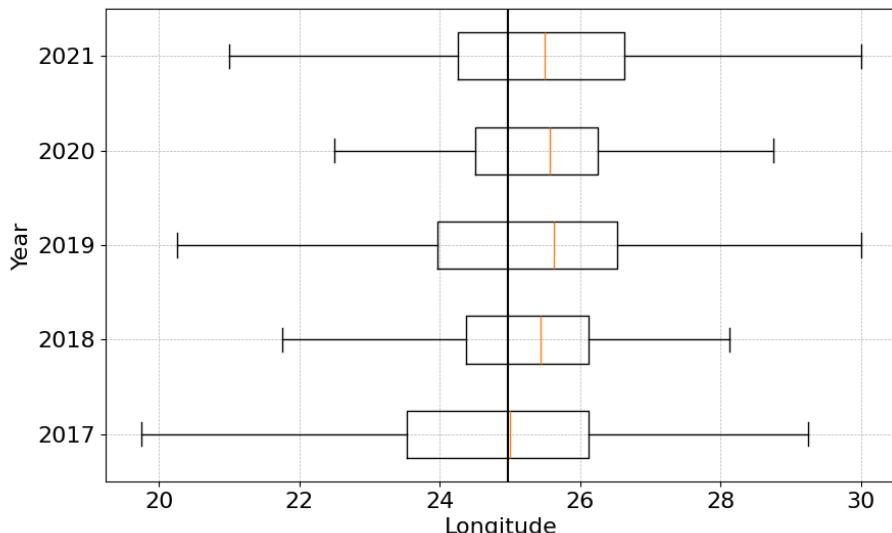


Figure 4.3: Box plot of longitude estimation using FMI Helsinki dataset. Boxes mark 25% to 75% quantile intervals and the orange middle lines are the medians. Correct longitude marked as a black line near 25 degrees.

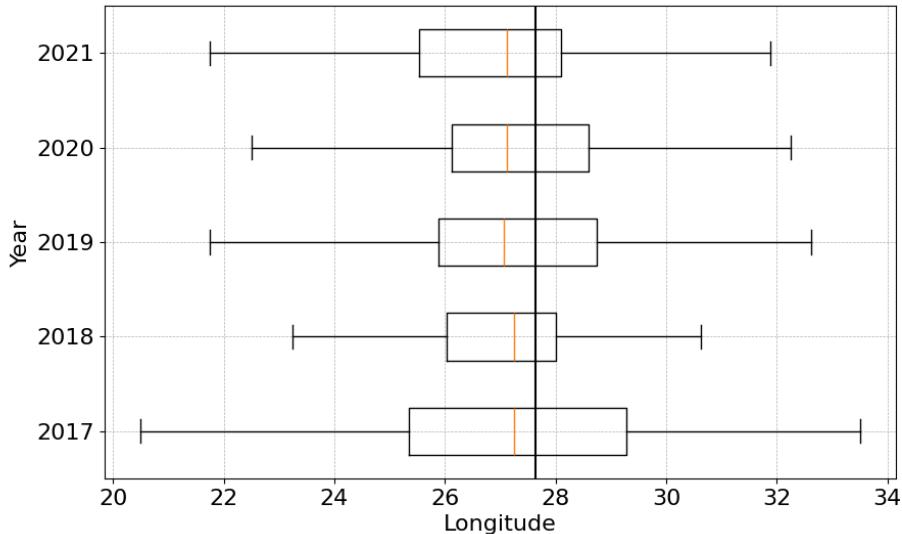


Figure 4.4: Box plot of longitude estimation using FMI Kuopio dataset. Boxes mark 25% to 75% quantile intervals and the orange middle lines are the medians. Correct longitude marked as a black line near 25 degrees.

4.1.2 Possible issues and further development ideas

PV output measuring systems may not be able to accurately measure when the power output is zero. The capacitance of electrical circuits, instrument accuracy, or other phenomena could lead to small positive or negative power readings even when the PV system is not producing any electricity. Small negative power output readings occur in the FMI datasets and the first and last non-zero power minutes are used instead of first and last greater than zero power minutes, this results in a bias of -3° in the longitude predictions. If similar positive false readings are present in datasets, this could complicate the estimation of the times when power production begins or stops.

While experimenting with the solar minute estimation functions, a curious trait was found. In Figure 4.5, the average of the first and last minute is approximately the same for each day at different latitudes as long as the latitude is below 50° . As the latitude is increased, the solar noon estimates begin to deviate significantly, becoming strongly skewed after 70° .

At first this behavior seems strange as astronomical solar noon should occur at the same time when longitude and the day are the same regardless of latitude. However as the solar noon estimates are calculated based on the first and last non-zero irradiance minute of the day, it would make sense that the estimations could be off by significant amount during equinoxes due to rapid changes in day lengths. Thus the day range was split into two sections and the color of days was chosen according to their distance from equinoxes. Line color was chosen based on distance from equinoxes where darker lines were the furthest away. This structure visible in the figure suggest that day length variation is a likely to be a factor in the phenomena.

In theory, if the same bias occurs in both the measurements and the modeled first and last positive power minutes, no corrections would be needed. The effect should be also lessened by using longer day ranges for predicting longitudes or by making sure that the intervals include an equal amount of days from both halves of the year as the direction in which this error occurs appears to be the opposite on both halves of the year.

Improvements in the algorithm accuracy could be achieved via by increasing the sampling interval of the irradiance simulations. PVlib POA simulations include a parameter for sampling frequency which is currently set to 1-per-minute in order to match the measuring frequency of FMI datasets. This could be increased to 1-per-second and the added resolution could help in determining more accurate estimates for solar noon times, resulting in possible gains in algorithm accuracy.

PVlib POA model was used instead of the more complex reflection and temperature aware model. This could be done as the geolocation is connected to first and last non-zero minute times which should be the same for both models. However even the POA model might be overly complicated as only two time values are needed. The usage of a much simpler model could increase computational speed significantly.

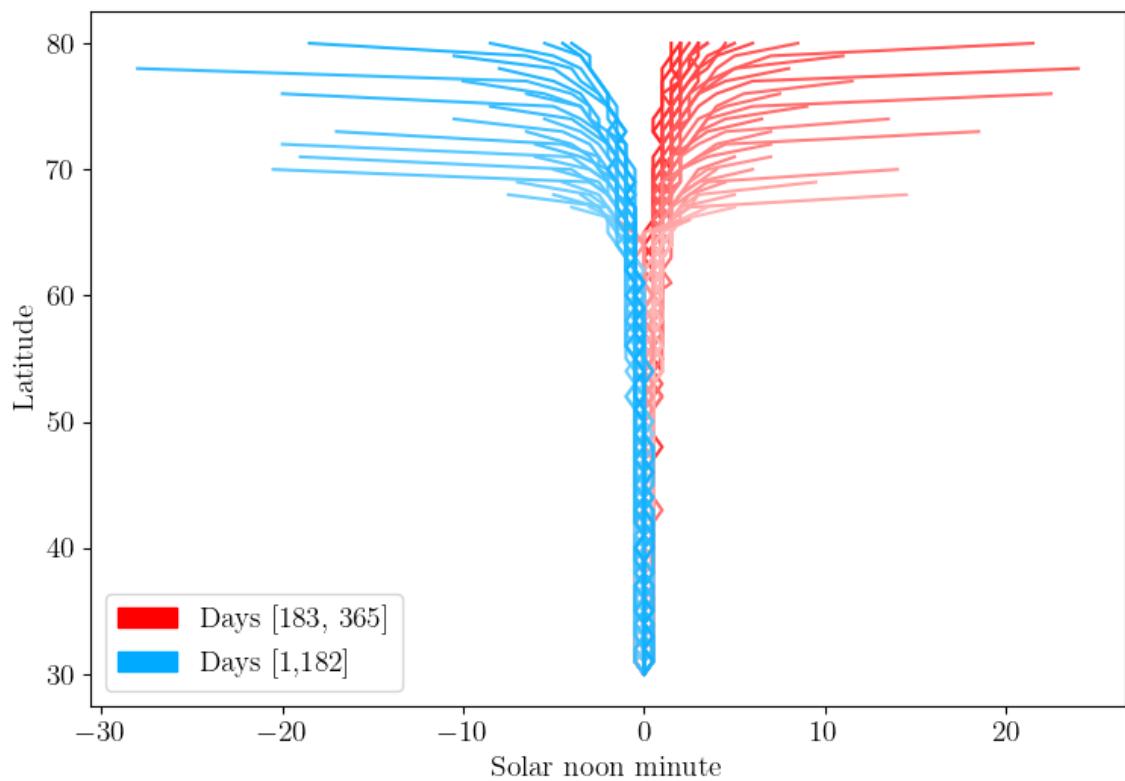


Figure 4.5: Each line represents a set of solar noon estimates based on first and last simulated non-zero power minute at different latitudes. Days are split based on whether they belong to the first or second half of the year. Darker line color marks days which are further away from equinoxes.

4.2 Estimating geographic latitude

Similarly to the longitude, the latitude of an installation is strongly connected to the timing of the first and last non-zero minutes of the day. This means that PVlib POA simulations can be used instead of the more complex PV model.

In earlier Figure 3.6, the simulated first and last minutes can be seen to change day by day at varying rates based on the latitude. In mathematical terms it could be said that the slope of the day-to-first-minute function is determined by the latitude of the installation. And for the days around equinoxes, and at higher latitudes of 50° to 70° , this relationship would seem to be bijective as per earlier Figure 3.6. The following algorithm is based on the former observations.

4.2.1 Latitude algorithm

1. Simulate first non-zero minutes for a day range A to B at a given latitude.
2. Fit a line to the simulated day to first minute pairs from step 1. Save the line slopes and the corresponding latitude.
3. Repeat steps 1 and 2 for multiple latitudes.
4. Create a slope to latitude graph from the earlier steps.
5. Fit an n-degree polynomial equation to the graph from step 4.
6. Select the day range A to B from a solar PV dataset and fit a line to the first non-zero power minutes. Feed the slope of this line to the polynomial from step 5. Value of this polynomial is an approximation of the latitude of the PV system.

Notes: Due to seasonal differences in data quality, polar winters and the midnight sun, the range of days chosen for the algorithm is important. If the range is short, individual outliers in measurements can result in large errors. Whereas if the range is too long, it will be harder to choose the range while avoiding low data quality sections. In the algorithm visualization Figure 4.6, the range of 250th to 300th seems to result in a good linear fit.

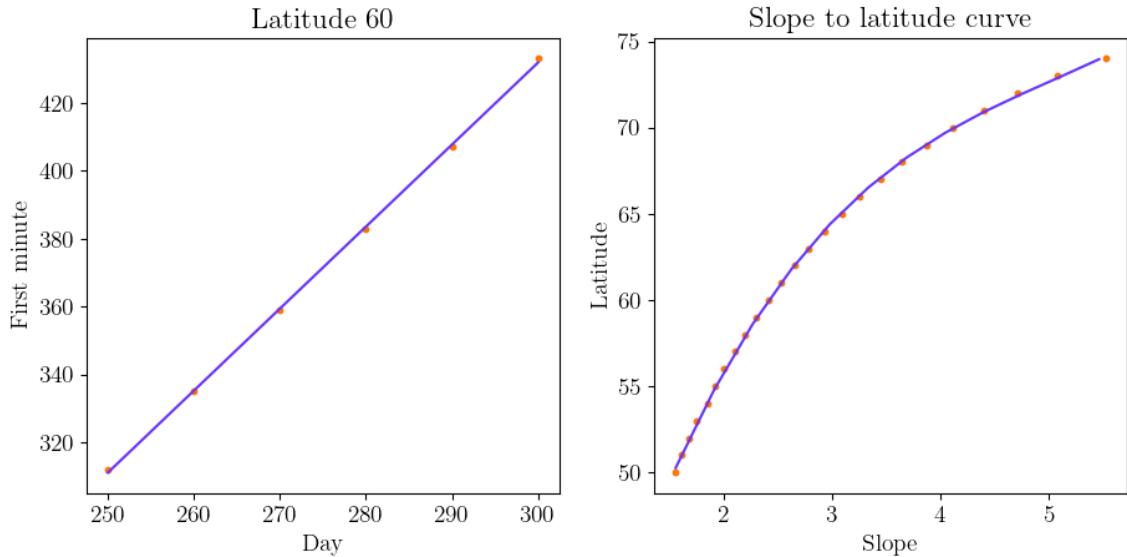


Figure 4.6: Left shows the linear relationship between day and simulated first minutes. Right shows the relationship between the slope angle and latitude.

FMI Kumpula[250-300]		
Year	Predicted latitude	Error
2021	61.365°	1.161°
2020	64.493°	4.289°
2019	63.121°	2.917°
2018	61.190°	0.986°
2017	57.515°	-2.789°

Table 4.7: Results from estimating the latitude of FMI Kumpula PV installation with the preceeding algorithm. Day range of 250th to 300th was used.

4.2.2 Improving latitude prediction algorithm

The results of the algorithm shown in Table 4.7 are somewhere in the correct region, but the delta of over 4° in the 2020 estimate is significant and much higher than the error of the longitude estimation algorithm. The first step in improving the algorithm would be the use of last non-zero minutes as well as the first non-zero minutes. This doubles the amount of outputs from the algorithm and while doubling the amount

of outputs does not directly increase the accuracy of the algorithm, it can provide additional insights into the performance of the algorithm.

FMI Kumpula[250-300]				
Year	First min. p.	Error	Last min. p.	Error
2021	61.365°	1.161°	63.685°	3.481°
2020	64.493°	4.289°	64.288°	4.084°
2019	63.121°	2.917°	66.762°	6.558°
2018	61.190°	0.986°	60.230°	0.026°
2017	57.515°	-2.789°	62.256°	2.052°

Table 4.8: Latitude algorithm with added output for last minutes based prediction.

The second step in improving the algorithm is choosing the best possible day range for latitude estimation. One way of choosing the day ranges would be by testing multiple day ranges and choosing the range which results in the lowest average absolute error from the known latitude, but this is problematic as the correct latitude should not be assumed to be known. However if there were multiple datasets with complete metadata, this could be used in order to find universally well-behaving day ranges.

Standard deviation minimization is another option for day range selection. As there are two estimated latitude values per year, datasets with n years of data would provide $n * 2$ estimated latitude values. Standard deviation of these values could be expected to be small if the day interval does not contain days with bad data quality and this means that the interval selection can be automated. Following Figure 4.9 shows a heatmap for FMI Helsinki latitude estimation. Lowest found standard deviation occurred with the day interval 180th to 310th with standard deviation of 0.3788°.

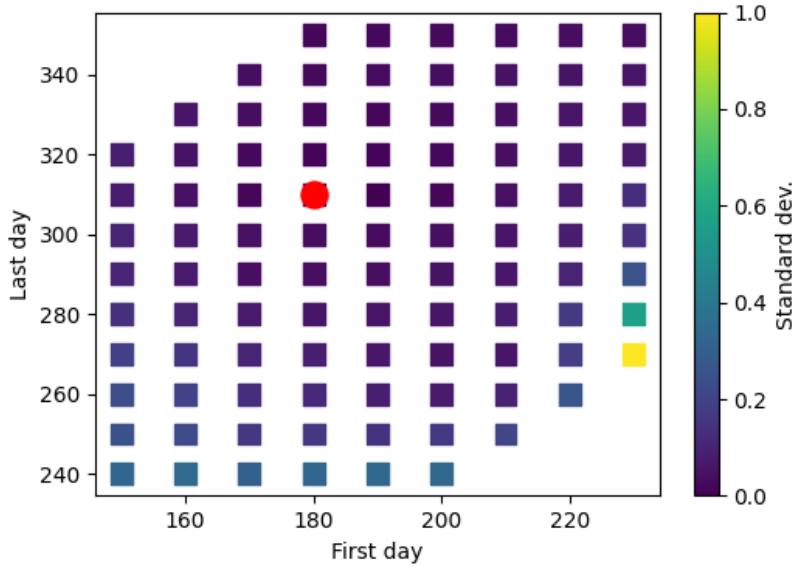


Figure 4.9: Standard deviation map for FMI Helsinki latitude estimation algorithm. Tightest grouping of latitude results was achieved with day interval 180th to 310th.

4.2.3 Latitude estimation results

Latitude estimation appears to be more challenging to optimize than longitude estimation. With the Helsinki dataset the found low standard deviation interval 180-310 predictions shown in Table 4.10 do not seem to be better than those achieved by an arbitrarily chosen interval 190-250 shown in Table 4.11. This observation is supported by the flatness of the standard deviation space in Figure 4.9 which also suggest that large sections in the day interval space have a similarly low standard deviation. Results for the Kuopio installation have higher variance but the scatter appears to otherwise be similar as seen on Table 4.12.

FMI Helsinki latitude estimation results(180-310)				
Year	First min. p.	Error	Last min. p.	Error
2021	60.7453°	0.5409°	61.107°	0.9026°
2020	60.4969°	0.2925°	60.82°	0.6156°
2019	60.5981°	0.3937°	60.4222°	0.2178°
2018	60.6537°	0.4493°	60.5745°	0.3701°
2017	61.416°	1.2116°	61.2206°	1.0162°

Table 4.10: Estimated latitudes for FMI Helsinki Kumpula dataset with day range of 180th to 310th

FMI Helsinki latitude estimation results(190-250)				
Year	First min. p.	Error	Last min. p.	Error
2021	59.792°	-0.677°	60.186°	-0.334°
2020	59.792°	-0.412°	60.186°	-0.018°
2019	59.896°	-0.308°	59.558°	-0.646°
2018	59.945°	-0.259°	59.463°	-0.741°
2017	60.577°	0.373°	60.008°	-0.196°

Table 4.11: Estimated latitudes for FMI Helsinki Kumpula dataset with day range of 190th to 250th

FMI Kuopio latitude estimation results(190-280)				
Year	First min. p.	Error	Last min. p.	Error
2021	62.626°	-0.266°	63.197°	0.305°
2020	62.259°	-0.633°	61.895°	-0.997°
2019	62.983°	0.091°	62.708°	-0.184°
2018	62.722°	-0.170°	62.874°	-0.018°
2017	61.669°	-1.223°	61.152°	-1.740°

Table 4.12: Estimated latitudes for FMI Kuopio Kumpula dataset with day range of 190th to 280th.

4.2.4 Possible issues and further development ideas

PVlib POA based first and last minute estimations are slower to compute than necessary as only two timestamps are needed. The use of a simpler sunrise and sunset equation would increase the speed significantly, allowing for the use of brute force day range selection algorithms.

Different methods could also be used. In Hagdadi 2017 [1] latitude estimations are done by fitting solar irradiance models with 3 unknown parameters to power generation measurement data. The latitude deltas of 1.65° to 3.42° in the 2017 article are higher than those achieved in this thesis, however as the datasets, geographical regions and algorithms are different, direct comparison can not be made.

In earlier Figure 4.6 the slope to latitude fitting is not perfect. The use of higher degree polynomials may result in a closer fit which in turn may result in a marginal improvement in latitude prediction accuracy. Alternatively a piecewise linear interpolation based fitting could be used. Both methods can in theory result in higher estimation accuracy at the risk of overfitting.

The two day intervals 180 to 310 and 190 to 250 may be biased. For the interval 180 to 310 all of the intervals result in positive error values whereas with 190 to 250 these errors mostly negative. This suggests that there could be biases in the data which do not influence standard deviation of the results. This will be left as a possible topic of study for future studies.

4.3 Combined latitude and longitude estimations

As it is unlikely that the longitude and latitude estimation algorithms are used in isolation from one another, their results should be examined together. This can be done by plotting the estimated locations on a map. Here the two installations in Helsinki and Kuopio and their predicted locations per year are plotted side by side. Plotted latitudes are from earlier tables 4.11, 4.12 and longitudes are from the longitude prediction algorithm using an arbitrarily chosen day range of 190th to 280th.

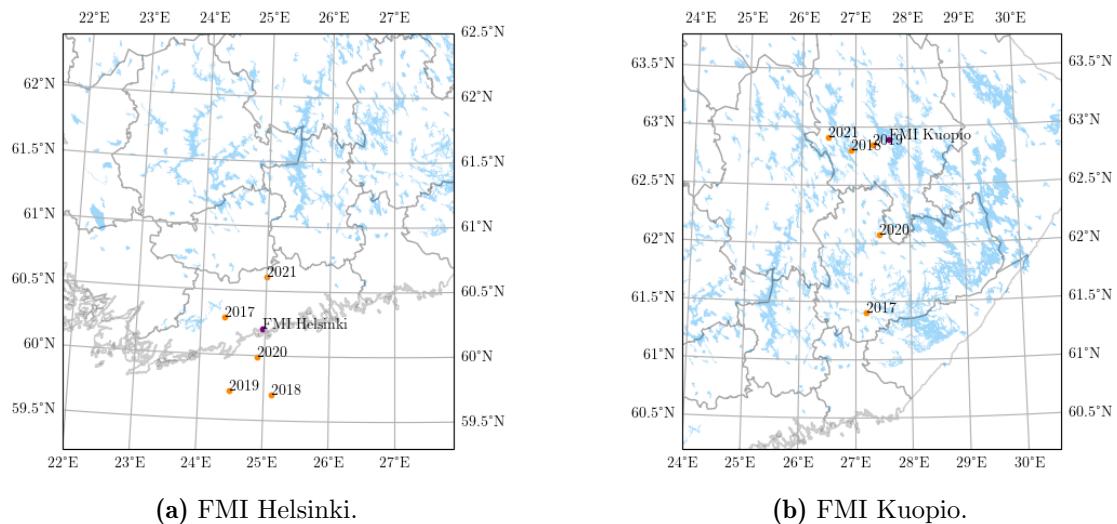


Figure 4.13: Geolocation estimations for FMI datasets.

In the Helsinki predictions figure, the estimated geolocations are scattered around the known installation location, showing very little bias and some random noise. Similar behavior can be seen in Kuopio predictions where two outliers 2017 and 2020 deviate more significantly. One degree on the latitude axis is approximately 110 km regardless of latitude and longitude, one degree of longitude is 56km at 60° N and 50 km at 63° N. As the variance is strongest on the latitude axis, it is likely that the latitude prediction algorithm is more sensitive to variations in the data and further development should be focused on more accurate latitude prediction and day range selection.

The results can be compared to the data resolution. One minute delta in measurements corresponds to a longitudinal shift of 14 kilometers in longitudinal axis at

60° . As the point cloud width is approximately 1° or 50km, the estimates can be thought to have a longitudinal range of 50km, 1° or 3.5 minutes with nearly the same values for the Kuopio installation. As the temporal resolution of measurements is 1 minute, the algorithm should not be limited by temporal resolution.

CHAPTER V

Estimating panel angles

Solar panel installation angles are a large factor in deciding the energy output of a PV system. If panel angles can be freely chosen during planning and installation phases, it can make sense to either optimize for total power generation or power generation during peak consumption hours. This means that even if installation angles could be freely chosen, installation angles are unlikely to be the same for every system in the same geographical region. Panel angles may also be restricted by installation sites and mounting types.

One reason for lacking or faulty metadata is that panel angles can be difficult to measure accurately. The tilt angle of the panels or the angle between the panel normal and zenith can easily be measured with an angle ruler and a bubble level, but the azimuth angle of the panels is much harder to measure with the same degree of accuracy. If an accurate compass is used and the difference between the magnetic north and the geographic north is taken into account, metal structures and electrical systems nearby can still distort local magnetic fields enough to cause errors in measurements. The challenges in taking accurate measurements are not insurmountable, but they may contribute to the inaccuracies and the lack of available information in PV installation parameter metadata.

The space of possible panel installation angles can be thought as a hemisphere in a spherical coordinate system where each point on the surface represents a direction to which the normal of the solar panels could be directed towards. A visualization of parameter space in 3D and 2D is shown in Subfigures 5.1a and 5.1b. The 3 dots in the Subfigure 5.1b mark the zenith for which azimuth is not well defined(red), the installation angles of FMI Helsinki installation azimuth 135° tilt 15° (blue) and a

close to power generation maximized installation with directly south facing panels with the tilt of 45°(black).

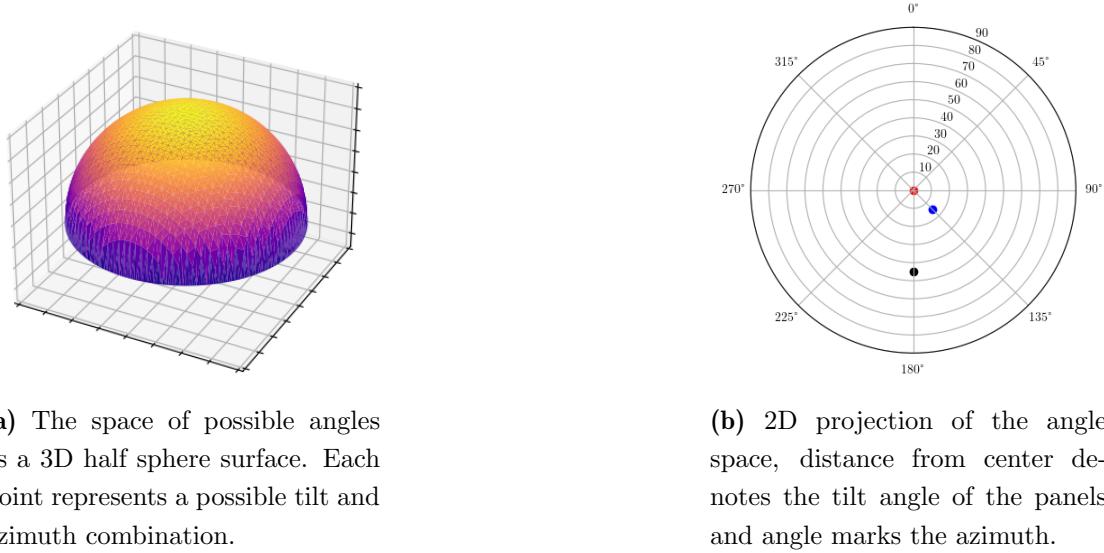


Figure 5.1: Angle space visualizations.

Estimating panel installation angles requires the use of multiple functions, each of which can be defined in multiple ways. These functions are defined in the following sections.

- Prediction error function for quantifying how good a prediction was when the correct panel parameters are known.
- Model fitness function for measuring the difference between simulated power values and measured power values.
- Angle space discretization function for discretizing the angle space into n discrete points which can then be tested with model fitness function.

5.1 Prediction error function

In this thesis, the proposed error estimation method combines the tilt and azimuth delta values into one error angle value, the angular distance between two points on a spherical surface. The goal is then to develop a panel angle estimation function which achieves the lowest angle error value with the available datasets.

Alternative approaches can also be chosen as the function or functions for measuring the distance between two points in angle space can be defined in multiple ways. The simplest way is to use the delta of known tilt and azimuth angles as two separate error values without normalizing in any way. This method was used in Hagdadi's 2017 [1] article but such values are not directly comparable between installations as the significance of azimuth delta depends on tilt angle.

Deriving angle space distance equation

Let $v = [v_1, v_2]$ and $k = [k_1, k_2]$ be two component angle-space vectors so that $v_1, k_1 \in [0, 90]$ and $v_2, k_2 \in [0, 360]$. These vectors represent points on the surface of a unit sphere and their components are the angles of spherical coordinate system. The cartesian coordinates of these points are:

$$x_v = \sin(v_1)\cos(v_2) \quad (5.1)$$

$$y_v = \sin(v_1)\sin(v_2) \quad (5.2)$$

$$z_v = \cos(v_1) \quad (5.3)$$

And

$$x_k = \sin(v_1)\cos(v_2) \quad (5.4)$$

$$y_k = \sin(v_1)\sin(v_2) \quad (5.5)$$

$$z_k = \cos(v_1) \quad (5.6)$$

And the cartesian distance between these two points can be calculated with the following equation:

$$d = \sqrt{(x_v - x_k)^2 + (y_v - y_k)^2 + (z_v - z_k)^2} \quad (5.7)$$

The two points and the origin form an isosceles triangle with the sides from the origin to the vector end points having the length of 1 while the distance between the vector end points is the same as d.

As the lengths of three sides are known, the angles of the triangle can be calculated with the cosine rule.

$$a^2 = b^2 + c^2 - 2bc \cos(A) \quad (5.8)$$

Where

a = Side opposing the angle A, same as earlier value d

b = Side opposing angle B, value is 1

c = Side opposing angle C, value is 1

Substituting known values into the cosine equation.

$$a^2 = b^2 + c^2 - 2bc \cos(A) \quad (5.9)$$

$$d^2 = 1^2 + 1^2 - 2 \cos(A) \quad (5.10)$$

$$d^2 = 2 - 2 \cos(A) \quad (5.11)$$

Solving for angle A

$$d^2 = 2 - 2 \cos(A) \quad (5.12)$$

$$2 \cos(A) = 2 - d^2 \quad (5.13)$$

$$\cos(A) = \frac{2 - d^2}{2} \quad (5.14)$$

$$A = \cos^{-1}\left(\frac{2 - d^2}{2}\right) \quad (5.15)$$

Renaming A as *Error*.

$$Error = \cos^{-1}\left(\frac{2 - d^2}{2}\right) \quad (5.16)$$

By first calculating the distance between the vectors using equations 5.1-5.7 and then substituting the distance into equation 5.16, the resulting angle can then be used as an error value between two panel angle measurements.

5.2 Simulation fitness function

The earlier model delta function defined as Equation 3.13 can be re-used here as the simulation fitness Equation 5.17. By then computing multiple simulations with different panel angles and choosing the simulation with the best fitness, meaning lowest delta value, the panel angle values can be estimated.

Simulation fitness function

$$Delta_{model} = \sum_{t=1}^{1440} |P_{measured}(t) - P_{simulated}(t)| / 1440 \quad (5.17)$$

As the function normalizes delta values by division with 1440, the delta for shorter days is lower than the delta for comparable long days. This should not matter for parameter estimation as the algorithm works by minimizing delta for each day independently.

Visualization of the function is shown in Figure 5.2 where simulation with tilt of 90° and azimuth of 135° is tested against known power measurements from a day in FMI Helsinki dataset, resulting in an average delta of 2959.5 W delta per minute.

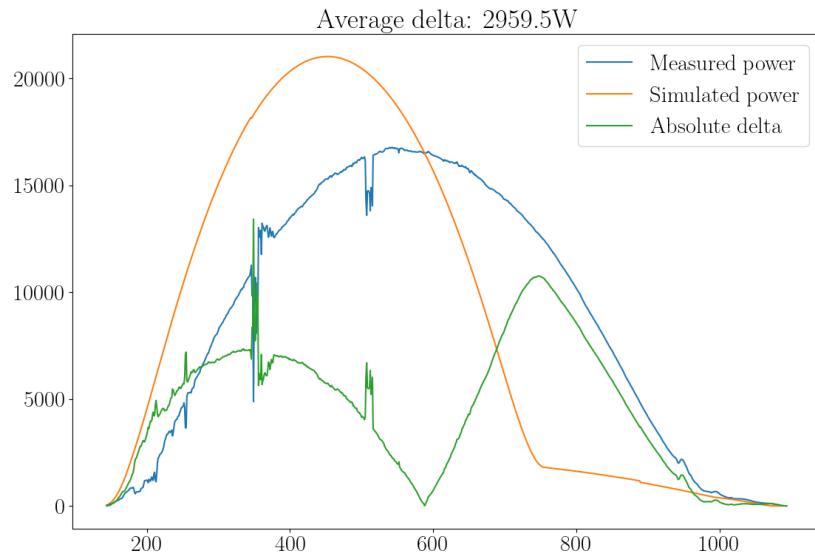


Figure 5.2: Visualization of fitness function.

5.3 Angle space discretization

The next step is angle space discretization. The panel angles are denoted with a doublet of tilt and azimuth values, ranging from 0 to 90 and 0 to 360 respectively. If the tilt and azimuth axes are discretized individually in steps of 5 so that tilt is [0, 5, 10, 15... 90] and azimuth [0, 5, 10, 15... 355], the permutations of these tilt and azimuth values create an even grid in the euclidean projection of angle space where $x = \text{tilt}$, $y = \text{azimuth}$. However as the physical phenomena represented by the angle values is not a point on a flat plane but a point on a half-sphere surface, this results in an uneven discretization seen in Figure 5.3a.

Sphere discretization problems are relevant for 3D graphics and real world problems involving geometry and so there are pre-existing methods available for discretization. One of the mathematically more elegant methods is the Fibonacci lattice which was used in a similar fashion in González 2009 [12]. The mathematical formulation of similar lattices is an older process and an earlier example is found in Vogel 1979 [13]. The following mathematical notation for the lattice is based on a code sample included in a blog post by Vagner Seibert [14].

Fibonacci lattice point n of k equation

$$s = n + 0.5 \quad (5.18)$$

$$\phi = a \cos(1 - 2s/k) \quad (5.19)$$

$$\theta = \pi s(1 + \sqrt{5}) \quad (5.20)$$

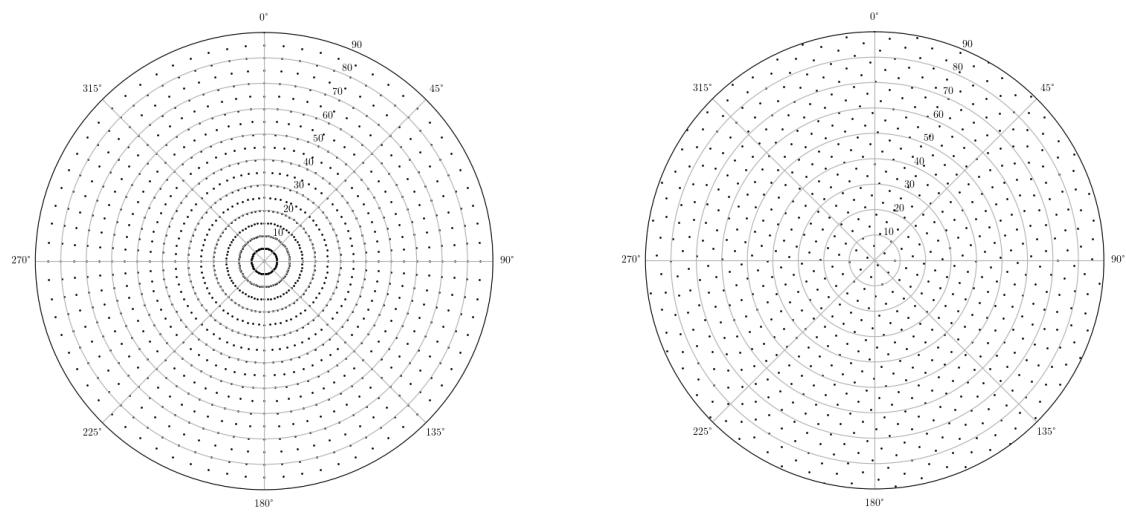
Where n is the point number, k is the amount of points, ϕ is the panel tilt angle and θ is the azimuth angle.

$$x = \cos(\theta) \sin(\phi) \quad (5.21)$$

$$y = \sin(\theta) \sin(\phi) \quad (5.22)$$

$$z = \cos(\phi) \quad (5.23)$$

x , y and z are the corresponding cartesian coordinates.



(a) In steps of 5 discretization with 1296 points

(b) Fibonacci lattice-based discretization with 756 points.

Figure 5.3: Comparison of two different discretization patterns. Fibonacci lattice based discretization on right shows a more even distribution of points than the latitude-longitude lattice. The minimum density is approximately the same in both graphs despite the difference in point counts.

5.3.1 Importance of lattice density

Using the correct lattice density is important for using exhaustive search algorithms for panel angle estimation. Regardless of the lattice point count, a discrete lattice is unlikely to ever include the best fit in the whole \mathbb{R}^2 parameter subspace. This means that lattices of a sufficient density should be used in order to guarantee that a lattice point exists near to the optimal fit. However as increasing the lattice density increases the computational cost of angle estimation, choosing a good density becomes an optimization problem.

Lattice densities behave differently in different latticing patterns. With *in-steps-of-n* lattices, the main benefit is easy readability. If the lattice is given one point for latitudinal and longitudinal degree, resulting in $360 * 90 = 32400$ points, then the lattice can be aligned so that each tilt and azimuth degree pair where angles are integers is tested. This discretization makes the results easily understandable as if the known installation angles are given as integers, a point representing the exact known installation angles already exists on the lattice.

With Fibonacci lattices and other lattice patterns in which lattice points may appear to be at seemingly random coordinates, evaluation of search algorithms is more difficult without visual aid. For example, in a 1000 point Fibonacci lattice the closest point to an arbitrary angle space coordinate is not obvious nor is the center angle distance between datapoints in the lattice as intuitive as with in-steps-of lattices.

5.4 Solving panel angles

Now that the geographic location and the multiplier values of installations are known to be solvable and fitness functions have been defined, the next step is solving the panel angles. The simple method is evaluating all points on a sufficiently dense lattice and choosing the point with the lowest delta value.

Figure 5.4 contains 10 Fibonacci lattice points and their normalized delta values. The best fit was at tilt 31.79° , azimuth 153.79° and it had a delta value of 369W. The lattice density leaves large gaps between lattice points and this found best fit is the closest point to the known installation angles of 15° and 135° . Center angle error as per Equation 5.16 is 18.17° .

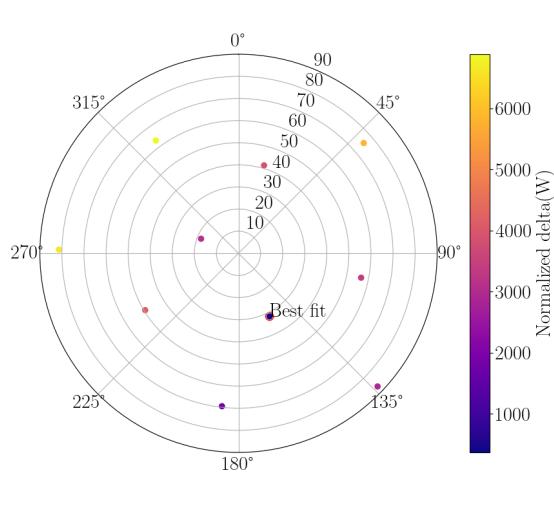


Figure 5.4: Polar plot of test points for a single day of data from FMI Kumpula dataset.

Table 5.5: Tilt, azimuth and error table for single day.

The fit achieved with the 10 point lattice is not very good and better results can be achieved by increasing the lattice density. The trivial method is to use a Fibonacci lattice with a higher point count, for example the fit achieved with a 10 000 point lattice show in Figure 5.6 found the best fit at tilt 14.79° , azimuth 136.2° , delta value of 136.2W and center angle error of 0.3659 degrees $^\circ$. This is a much better fit but increasing the lattice density comes with a higher computational cost. With moderately well written python code lacking multicore optimizations, evaluating a full 10 000 point lattice for a single day can take up to several hours on a modern

laptop.

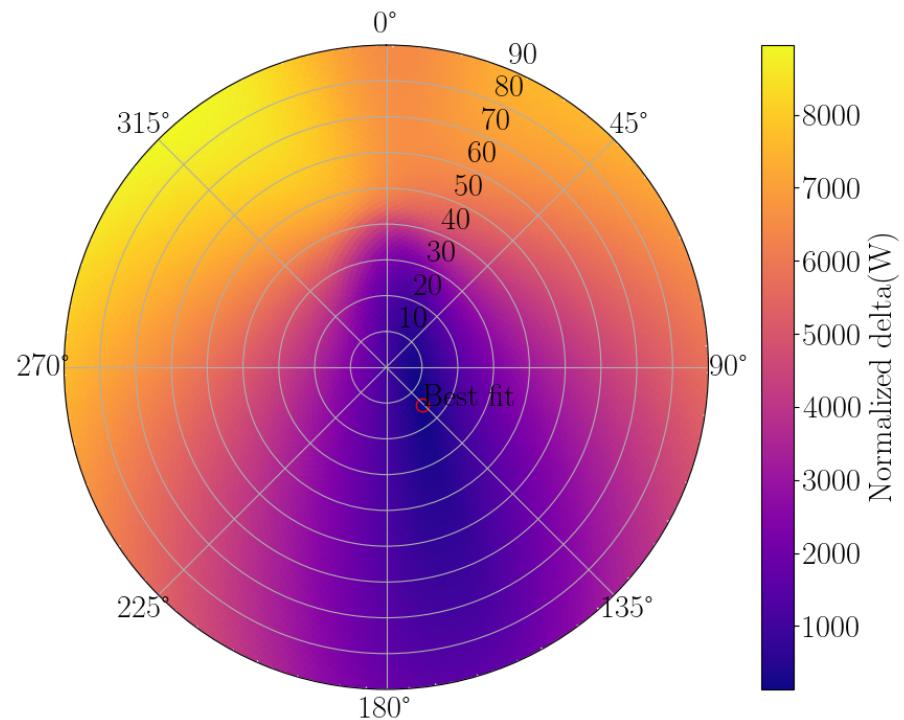


Figure 5.6: Results of a 10 000 datapoint lattice fitting against a single day from FMI Helsinki dataset. Center angle error between the found best fit and known installation angles is 0.3659° .

The results of the exhaustive search algorithm are excellent for the Helsinki dataset. Angle delta values as low as the achieved 0.37° are low enough to be within the assumed 1° precision of known installation angles.

The next verification step is using the algorithm on multiple days from the same dataset and checking the spread pattern. This was done on sets of days with both Helsinki and Kuopio datasets. Figures 5.7 and 5.8 display the tight scatter patterns in the Helsinki predictions. For the tight grouping in the 2018 figure the average tilt of 13.89° and azimuth of 130.81° are close reported angles of 15° tilt and 135° azimuth. And with year 2019 where the scatter pattern is wider, the average of the results is 14.17° tilt, 132.62° azimuth which are both even closer to the known installation angles. Fitness values or the per minute delta values are for the multi day 10 000 point evaluations were in the range of 76W to 117W for the Helsinki dataset.

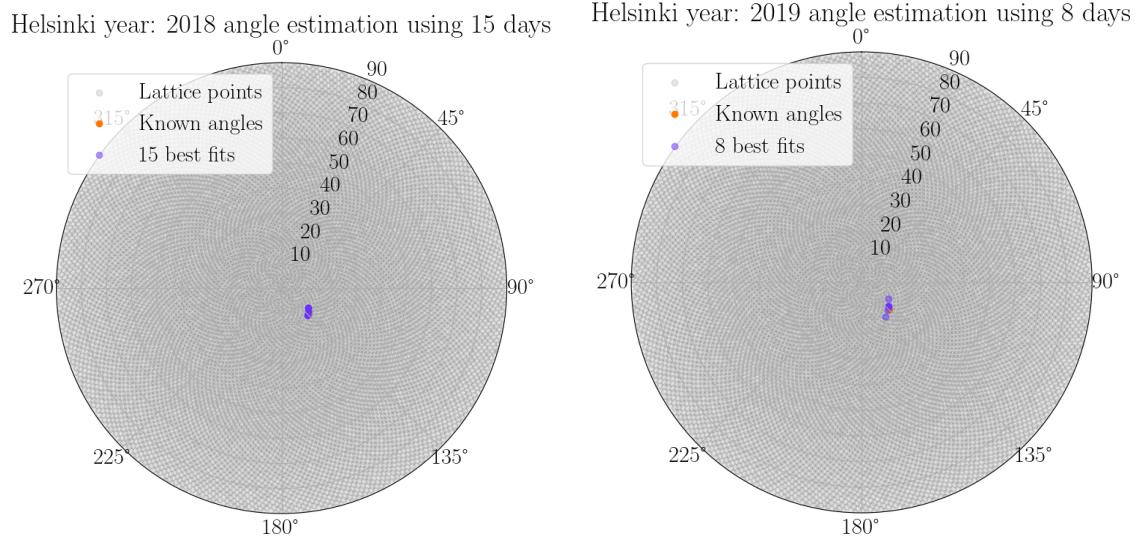


Figure 5.7: 10 000 point exhaustive search on 15 cloud free days from FMI Helsinki dataset using year 2018.

Figure 5.8: 10 000 point exhaustive search on 8 cloud free days from FMI Helsinki dataset using year 2019.

Results for the Kuopio dataset are not nearly as good as the Helsinki predictions. Figures 5.9 and 5.10 clearly indicate that estimations are converging approximately 13° off from the known installation angles. The reason for this significant delta is not

known and there could be multiple contributing factors. One possible explanation is panel shading which is more likely to occur when the Sun is near the horizon. This causes a proportionally higher energy output loss during first and last hours of the day, resulting in a more narrow power generation plot shape. This *sharpness* is also influenced by the tilt angle of an installation as was previously seen in Figure 3.3 and thus the error in the estimated tilt angle could be partially caused by either panel self shadowing or shadowing caused by other structures near the PV panels.

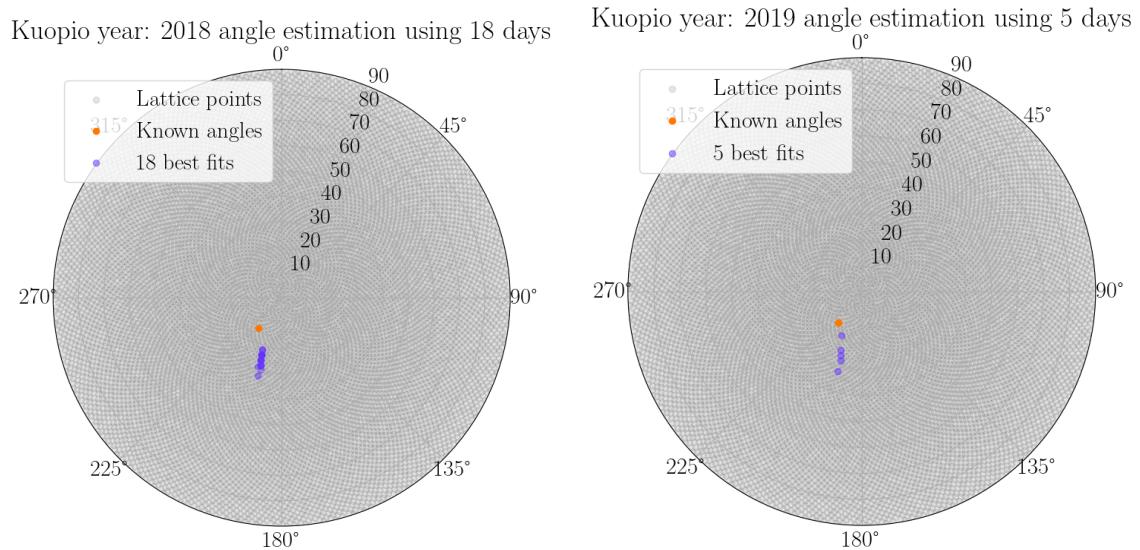


Figure 5.9: 10 000 point exhaustive search on 18 cloud free days from FMI Kuopio dataset using year 2018.

Figure 5.10: 10 000 point exhaustive search on 10 cloud free days from FMI Kuopio dataset using year 2020.

5.5 Solving panel angles iteratively

The exhaustive search used in earlier section suggest that the method is capable of estimating panel installation angles accurately. However evaluating 10 000 point lattices is somewhat inelegant and this can be avoided by using multiple less dense lattices iterarively if the fitness space satisfies some requirements.

The first requirement is that the fitness space should be smooth. This smoothness does not have to be perfect, fine patterns and details in the fitness space surface do not cause issues with iterative search algorithms if the noise pattern is not observable in the iterative lattices. Based on an earlier Figure 5.6 generated by exhaustive search, this requirement would appear to be met.

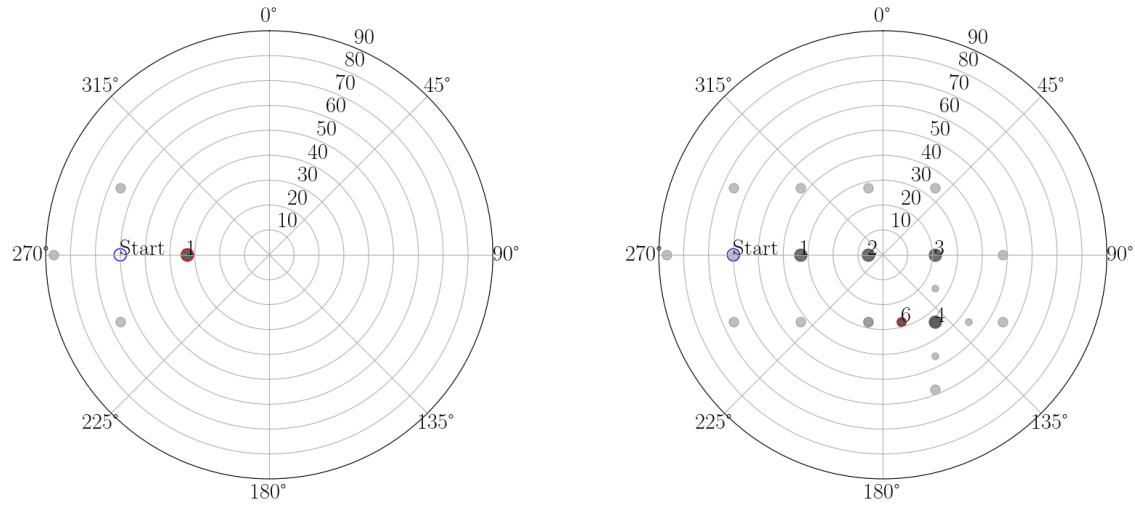
Second requirement is that the fitness space should contain as few convergence points and their respective basins as possible. These basins are regions in fitness space which are defined by a convergence point and their surrounding regions where the slope of the space leads to the respective convergence point. The region around best found fit in Figure 5.6 forms a large basin but there would also appear to be a second converge point somewhere around azimuth 0° tilt 90°.

Iterative panel angle estimation algorithm

1. Choose a cloud free day from the dataset for evaluation.
2. Choose a starting or "center" point from the angle space. This can be either the best result from a low density lattice or a fixed point such as tilt 45° azimuth 180°.
3. Evaluate the fitness at the center point and store that as the center point fitness.
4. Choose a few points near this center point within a given distance and evaluate their fitness. If any of the neighboring points results in a better fit than the center point, this point will then be chosen as the new center point.
5. Repeat steps 3 and 4 until step 4 does not find a better fit. When this happens, decrease the distance used for the local lattice in step 4.
6. Repeat steps 3,4 and 5 for a set number of times. Last center point is the best iteratively found fit.

5.5.1 Local lattice generation

The step 4 in the iterative panel angle solving algorithm detailed above is the first non-trivial step in the algorithm as it requires the generation of local lattices. One option would be to use a subsection of a Fibonacci lattice, but generating large Fibonacci lattices and utilizing only a small subsection would require unnecessary computation.



(a) Iterative best fit search after 1 cross pattern search.

(b) Iterative best fit search after 6 cross pattern searches.

Figure 5.11: Visualization of a + pattern iterative best fit search algorithm. Numbers next to markers indicate that a point was the best found during n:th cross pattern search.

The method shown in Subfigure 5.11a is based on generating a + pattern by first transferring the starting point from angle space to a unit disk with Equation 5.24. On this unit disk 4 points are generated, each of which is s units away from the start point either on x or y -axis using Equations 5.25. When these 4 points are transformed back into angle space with 5.26, they can be used as the local lattice surrounding the chosen center point.

Angle space to unit disk equations

$$\begin{aligned}
 d &= \text{Tilt}/90^\circ \\
 x &= \cos(\text{Azimuth})d \\
 y &= \sin(\text{Azimuth})d
 \end{aligned} \tag{5.24}$$

Unit disk points

$$\begin{aligned} p_1 &= (x + s, y) \\ p_2 &= (x - s, y) \\ p_3 &= (x, y + s) \\ p_4 &= (x, y - s) \end{aligned} \tag{5.25}$$

Unit disk to angle space equations

$$\begin{aligned} d &= x^2 + y^2 \\ Tilt &= \sqrt{d} * 90^\circ \\ Azimuth &= atan2(x/d, y/d) \end{aligned} \tag{5.26}$$

5.5.2 Results

By using 30 iterations with initial search distance of 0.3 on a set of cloud free days from the FMI Helsinki and Kuopio dataset, scatter patterns shown in Figures 5.12a and 5.12b are generated. The center angle distances between known angles and cluster means for the Helsinki and Kuopio scatter patterns are 1.97° and 12.47° respectively. The Helsinki scatter pattern is marginally worse than the pattern achieved by exhaustive search algorithm shown in Figure 5.8. This may be due to the algorithm getting stuck on a local minimum. Whereas Kuopio results are similar fits found with the exhaustive search.

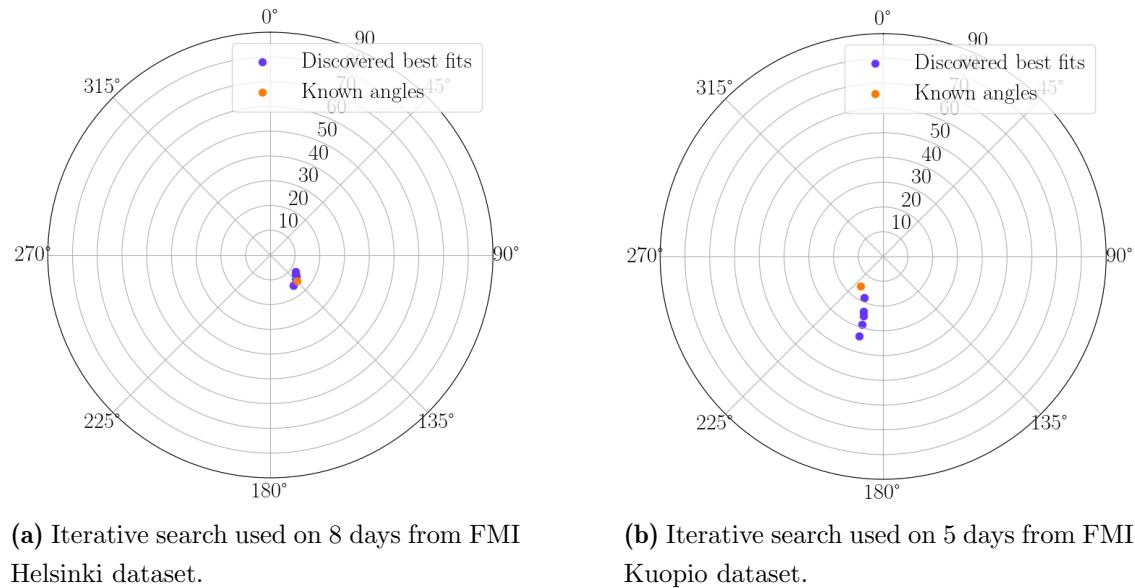


Figure 5.12: Scatter patterns for iterative panel estimation algorithm on multiple days.

One method for comparing the iterative method against the exhaustive is to compare the fitness values achieved by both methods. If for example the resulting fitness values are better for the exhaustive search, then there could be multiple convergence points in the basin of the fitness space and the iterative function could be converging on the wrong point. Whereas if the iterative method achieves worse center angle distances and better fitness values than the exhaustive, the seemingly worse results of the iterative method would be a result of a better fit existing in the angle space which the Fibonacci lattice points did not contain.

By testing out individual days and sets of days, the difference between achieved

fitness values with the different estimation methods do not seem to vary significantly. For the Helsinki dataset the fitness values are somewhere in the 75W to 120W range for both algorithms depending on chosen parameters, with even performance with both algorithms. This is also visible in Figures 5.13 and 5.14 where the curves generated by the results of both search methods are indistinguishable.

5.5.3 Choosing between exhaustive and iterative angle estimation methods

Both methods are capable of achieving good to results depending on data quality and algorithm parameters. As there are multiple algorithm parameters and as the exhaustive is computationally expensive, it is difficult to prove that one method is more accurate than the other. Each day in the datasets has an unique fitness space and parameters such as initial search distance, starting point and lattice point counts have a small and difficult to measure effects on the results of the algorithms.

That being stated, the computational speed of the iterative method is much higher than that of exhaustive search algorithms. If 30 steps are used with a 4 point local lattice, estimating the panel angles for a single day require 121 simulations and comparisons. This means that approximately 83 days can be iteratively evaluated in the time it takes to exhaustively evaluate one day with a 10 000 point lattice. As the 30 step iterative search algorithm reaches search distances as low as $9e - 6$ units on the unit disk, the iterative method is using finer precision than would practical and the speed of the iterative algorithm can be increased further by setting a search distance limit.

Instead of choosing one method, the two search methods can be combined and used together. An iterative method with 5 to 10 steps is enough to find the general location of the best fit and a local subsection of a dense Fibonacci lattice could be used in order to search a subsection of the angle space. This would however increase the algorithm complexity for small potential gains.

Figure 5.13 shows a comparison of fitting results for a single day. There the difference between pover curves is not significant regardless of the estimation method used and thus iterative method may be preferable. The difference between exhaustive and iterative search is likewise imperceivable in the example from FMI Kuopio dataset seen in Figure 5.14 despite the higher center angle distance achieved with both exhaustive and iterative methods in all earlier examples.

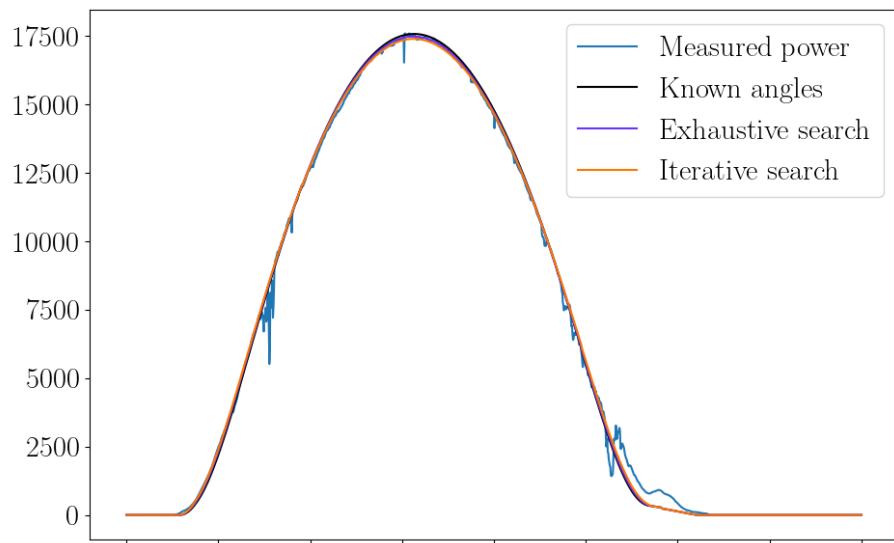


Figure 5.13: Comparison of power curves from measurements, simulation with known installation angles and estimated panel angles with exhaustive and iterative method. Using day 150 from year 2019 in Helsinki dataset.

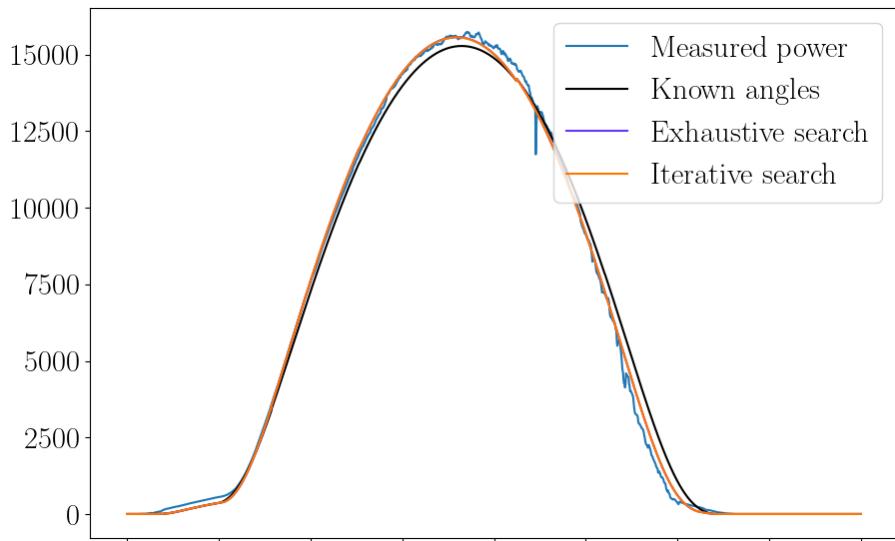


Figure 5.14: Comparison of power curves from measurements, simulation with known installation angles and estimated panel angles with exhaustive and iterative method. Using day 167 from year 2019 in Kuopio dataset.

5.5.4 Further development ideas

Both algorithms perform well enough to make some of the assumptions made on the properties of fitness spaces practically irrelevant. Thus concepts of basins were not explored thoroughly but they and the field of dynamic systems which studies attractors and basins could offer additional insights into similar optimization problems.

The problem of model fitting for angle estimation may also have connections to convolutions. Understanding the mathematical properties of signal impulses and their convolution shapes could help in proving how likely the formation of multiple basins are. This could be beneficial in starting point selection with iterative fitting algorithms.

The estimation algorithms should be evaluated with multiple different datasets in order to determine whether results from the Kuopio data are an anomaly or the algorithms are as noise sensitive as the Kuopio data suggests. Similarly data from different sources and with different temporal resolutions would help in verifying how well the algorithms perform with available data.

CHAPTER VI
Conclusion

PV system parameter estimation results for the FMI Helsinki dataset are very good to excellent. Angle estimation results with a center angle delta of less than 1° were achievable with both iterative and exhaustive algorithms. Geolocation estimation proved to be more difficult with the scatter pattern from multiple years having small amount of bias and a fair amount of noise. Scatter formation is spread around the FMI Helsinki installation and is approximately 50km by 200km in dimensions. Due to small sample size of 5 datapoints this algorithm is harder to evaluate.

Results for FMI Kuopio dataset were noticeably worse. Center angle delta with panel installation angles was approximately 13° regardless of estimation method used. Similarly the scatter pattern in geolocation estimation resulted in a 50km by 300km region with outliers.

Noise types and the influence of noise in PV data should be studied further. Yet unpublished research by Juha Karhu studies methods for detecting shadowing on solar PV panels and the algorithms developed by Karhu are showing signs of shadow induced output losses with the Kuopio installation. This shadowing could be the primary cause for the lower accuracy of parameter estimation with the Kuopio dataset.

.1 Source code

In total, the project source consists of 3000 lines of source code divided into multiple files and directories for easier usage. Including the full source in the appendix would not be feasible but the source is available online at github [15]. These project files can be downloaded and tested in python environments which meet the recommended python and python package versions detailed in the included Readme.md -file. This same file also includes further description on project structure and usage.

.2 Data files

The solar PV datasets used in this thesis are private but available upon request. PV datasets are planned to be released along with a data paper at a later point.

The program expects FMI Helsinki and Kuopio csv -files to have names "fmi-helsinki-2021.csv" and "fmi-kuopio-2021.csv" respectively. The exact filenames are not important and files can either be renamed to match these patterns or the data loading code in "solar_power_data_loader2.py" can be modified to match the used .csv files.

Datasets from other sources can also be used but their usage will require some modifications to the program code. If the temporal resolution or csv structure are different from the structure detailed in table 2.2, these modifications may be significant.

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