

A Fast Algorithm for Finding Point Sources in the Fermi Data Stream: FermiFAST

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ABSTRACT

This paper presents a new and efficient algorithm for finding point sources in the photon event data stream from the Fermi Gamma-Ray Space Telescope. It can rapidly construct about most significant half of the Fermi Third Point Source catalogue (3FGL) with nearly 80% purity from the four years of data used to construct the catalogue. If a higher purity sample is desirable, one can achieve a sample that includes the most significant third of the Fermi 3FGL with only five percent of the sources unassociated with Fermi sources. Outside the galaxy plane, the contamination is essentially negligible. This software allows for rapid exploration of the Fermi data, simulation of the source detection to calculate the selection function of various sources and the errors in the obtained parameters of the sources detected.

Key words: methods: data analysis — methods: observational — techniques: image processing — astrometry

1 INTRODUCTION

The launch of the Fermi Gamma-Ray Telescope in July 2008 brought dramatically more sensitive instruments to bear on the study of the Universe in gamma rays. The observations of the Fermi LAT (e.g. Ackermann et al. 2012) have resulted in a series of catalogues of point sources both within the Galaxy and beyond (e.g. Abdo et al. 2010; Ackermann et al. 2011; Nolan et al. 2012; Abdo et al. 2013; Ackermann et al. 2013; Acero et al. 2015). Both the identification of sources and their characterisation rely on the calculation of a likelihood ratio for the observed data with and without a source and for different characteristics for the potential source (Cash 1979; Sutherland & Saunders 1992; Mattox et al. 1996). These telescopes as well as Cerenkov telescopes on the ground rely on the conversion of gamma-rays into electron-positron pairs as they enter the telescope (or the atmosphere). By tracing the motion of these pairs, one can reconstruct the momentum of the incoming gamma ray. For low gamma-ray energies the reconstruction is poorer than at higher energies. Because of this broad point-spread function of gamma-ray telescopes the contribution of several potential sources will overlap in a particular region of the sky making the generation of a catalogue even more difficult and time consuming.

Although both Fermi and EGRET before it used the likelihood ratio to assess the significance of potential point

sources, several alternatives have been proposed to search for point sources to account for the photons observed by these instruments. Massaro et al. (2009); Campana et al. (2008) developed an algorithm to build a minimal spanning tree from the arrival directions of the photons over a large portion of the sky. They then divide the tree into subclusters by remove branches larger than a particular threshold. If a subcluster contains more than a particular number of photons, it is considered to be a potential point soource. Campana et al. (2013) extended the technique to include a noisy background. Damiani et al. (1997a,b); Ciprini et al. (2007) propose to use wavelets to find point sources from photon counts for ROSAT, EGRET and Fermi data. Although these techniques are significantly faster than the traditional likelihood ratio, they require simulated data to assess the significance of the sources detected because they do not rely of statistics with well understood distributions given the null hypothesis of no point source at a particular location. In general the estimated significances are proportional to those found in the full likelihood analysis.

In this paper we propose a technique to calculate likelihood ratios much more quickly that the usual tehcniques by storing the photon data stream in a data structure optimized for finding photons within a given percentile of the point spread function of the instrument from a potential source. Furthermore, we will assume a simple expression for the likelihood that assumes that the photons from within a particular region of the sky come from the combination of a point source at the centre of the region and a flat back-

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ground. Because we assume that the point-spread function is circularly symmetric, a flat background or even a background with a uniform gradient do not affect the likelihood ratio. We do not assume a particular spectral model but only that the measured momenta of photons from a point source follow the measured distribution (Ackermann et al. 2013) from the telescope.

2 THE PHOTON DATABASE

The key to the speed of this algorithm is the database that contains the position of the observed photons on the sky. Each photon is stored in a four-dimensional $k-d$ tree (Bentley 1975). We use the particularly memory efficient implementation of Lang (2009) (used in *astrometry.net*). The coordinates are actually stored as shorts instead of floats to save additional memory. The typical coordinates range from -1 to $+1$, so using shorts yields an angular precision of about six arcseconds much finer than that of the Fermi point-spread function (PSF). This memory efficient implementation allows us to store all of the photons detected by Fermi above 100 MeV and within a zenith angle of 100 degrees in memory simultaneously.

The first three dimensions contain the position of the photon on the celestial sphere as shown in the upper portion of Fig. 1. Storing the direction of the photon momentum in this manner removes the coordinate singularity of the spherical coordinates. Additionally it makes integrating over the celestial sphere straightforward because $\int d\Omega = \int 2\pi \delta p d\delta p$ where $\delta \mathbf{p} = \mathbf{p}' - \mathbf{p}$ is the three dimensional vector between two points on the sphere. The fourth coordinate that we denote by w depends on the point-spread function for the photon in question. In particular $w = \pm \sqrt{R_{\max}^2 - R_i^2}$ where R_i is the radius of the ninety-fifth percentile at the energy, entrance angle and front or back conversion for the photon. For convenience we use positive values of w for front-converted photons and negative values of w for back-converted photons. Furthermore, R_{\max} is the ninety-fifth percentile for the photon with the poorest angular resolution. This is depicted in the lower panel of Fig. 1.

Once the $k-d$ tree is created, it is efficient to find all of the entries within the database within a given Cartesian distance of a particular point. In our case we query the database for all of the photons within a distance R_{\max} of a particular point on the celestial sphere and use $w = 0$ for the fourth coordinate. The particular choice of w for the observed photons means that the query will yield all of the photons that are within the ninety-fifth percentile of the PSF. In other words, if there is a point source located at that particular point the query will return on average 95% of the photons from that source. Of course, it will also return photons from the background and other nearby sources. Additionally, the region of interest for a particular prospective source is energy dependent. The form of the exposure map also is energy dependent as shown in Fig. 2. It peaks at 0.95 times the exposure time in the direction of the potential source and slowly drops to the ninety-fifth percentile of the PSF in radius and then drops according to the power-law of the tail component of the linear combinations of King function that are used to characterize the Fermi PSF.

Although the construction of the tree is not done in par-

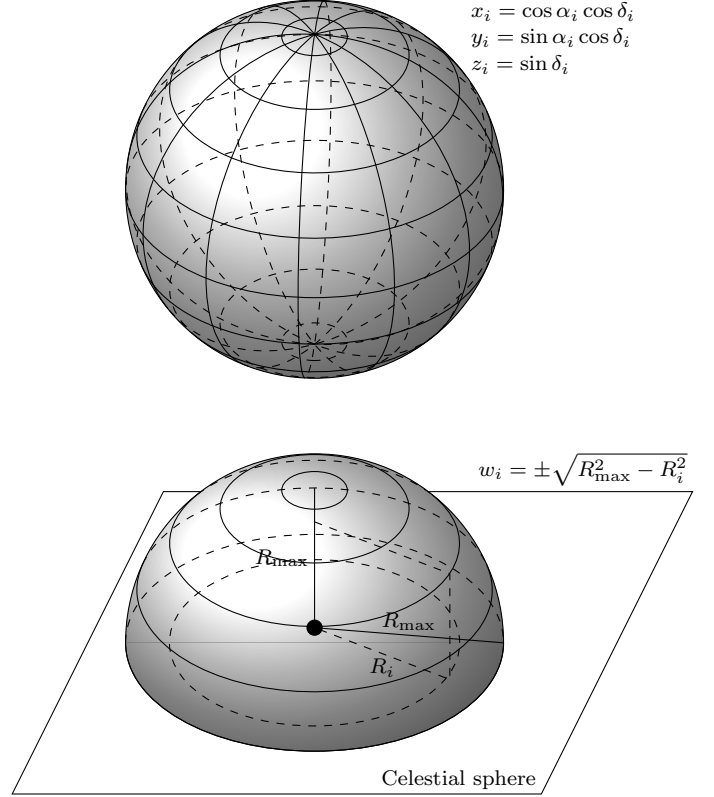


Figure 1. The location of a given photon event on the celestial sphere and in the additional dimension. R_i is the ninety-fifth percentile radius for the photon in question and R_{\max} is the largest ninety-fifth percentile radius for the photons in the sample.

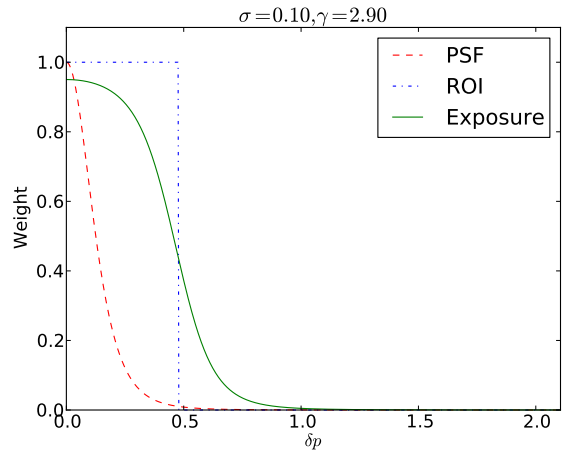


Figure 2. The Exposure Map, Region of Interest (ROI) and Point-Spread Function (PSF)

allel, the queries are performed in parallel using the tree held in shared memory. For example if one uses all the photons above 100 MeV from weeks 9 through 216 with a standard zenith angle cut of less than 100 degrees (89,684,009 photons) requires one gigabyte to store the tree and about ten minutes to construct. The 200,000 location queries and likelihood calculations require 140,000 seconds (700 ms each), so the speed-up through parallelisation can be dramatic. On

the other hand, if one restricts to photons above 1 GeV (13,193,171), it only takes 90 seconds to construct the tree. At higher energies, it makes sense to make more location queries because the PSF is smaller. In this example 786,426 queries require 5,600 seconds of CPU time (7 ms each) or only six minutes on sixteen cores.

3 SOURCE LIKELIHOOD

Using the $k-d$ tree to determine the photons that would like within the 95% enclosure region of the point-spread function, we calculate several statistics of the observed photons to assess the likelihood of a source being at a particular position on the sky. We first determine two statistics whose distributions are known. First, if all of the photons within the region of interest (all photons that lie within the 95th-percentile cone of the potential source) indeed come from a uniform background, the ratio of the solid angle enclosed in a circle centred on the potential source and running through the observed to the total solid angle within the region of interest for that particular photon should be uniformly distributed between zero and one. We denote the mean of this ratio over the observed photons \bar{r}^2 . This is a Bates distribution with mean of $1/2$ and variance of $1/(12N_{\text{photons}})$. Second, if all of the photons within the region of interest indeed come from a point source at the centre of the region of interest, the ratio of the percentile of a given photon within the cumulative PSF distribution for that particular photon to 0.95 should be uniformly distributed between zero and one. We denote this statistic by \bar{f} .

If we assume that the observed photons originate from a linear combination of these two possibilities, we can determine the ratio of the two contributions from these statistics. In particular the fraction of photons that come from the point source would be

$$A_f = \frac{\frac{1}{2} - \bar{r}^2}{\bar{f} - \bar{r}^2} \quad (1)$$

and we can estimate the significance of the value of A_f by

$$S(r^2) = \left(\bar{r}^2 - \frac{1}{2}\right) \sqrt{12N_{\text{photons}}} \quad (2)$$

and the probability of getting a value of $S(r^2)$ larger than x by chance is

$$P[S(r^2) > x] = \frac{1}{2} \text{erfc}\left(\frac{x}{\sqrt{2}}\right) \approx \exp\left(-\frac{x^2}{2}\right) \quad (3)$$

if we take the limit of many photons in the region of interest where the Bates distribution tends to the normal distribution.

These basic statistics are summarized in Tab. 1, and Tab. 2 lists these statistics for the ten most significant sources detected. These basic statistics really just compare two numbers about the distribution of the photons within the region of interest. We can use the detailed knowledge of the point spread function to develop a more comprehensive test of the distribution of photons. In particular we define the unbinned likelihood (this is very similar to the expression used by Braun et al. 2008, for neutrino telescopes)

$$\log L = \sum_{\text{photons}} \log \left[A_{\text{PSF}} \frac{\text{PSF}_i \Omega_{\text{max},i}}{0.95} + (1 - A_{\text{PSF}}) \right] \quad (4)$$

where we have dropped N_{pred} from the usual definition because we have defined the model in such a way that $N_{\text{pred}} = N_{\text{photons}}$ automatically and $dN_{\text{pred}}/dA_{\text{PSF}} = 0$. Furthermore, for $A_{\text{PSF}} = 0$, $\log L = 0$ and because we are fitting a single variable, $\log L$ is distributed as a chi-squared distribution with a single degree of freedom and the probability of getting a value of $\log L$ larger than x by chance is

$$P(\log L > x) = \sqrt{\pi} \text{erfc}\left(\sqrt{\frac{x}{2}}\right) \approx \exp\left(-\frac{x}{2}\right). \quad (5)$$

From Tab. 2 we can see that the values of A_f are similar to those of A_{PSF} at least for highly significant sources.

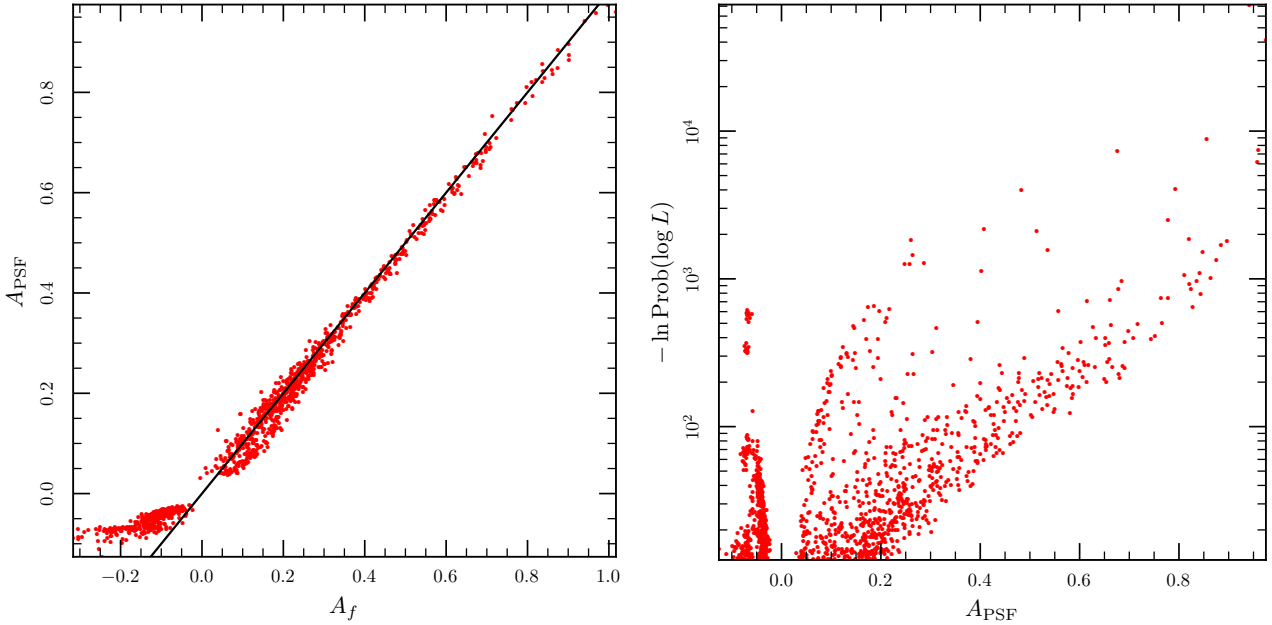
The first pass is to determine the value of $\log L$ on a HEALPix grid (Górski et al. 2005) of potential sources. For example for the photons above 1 GeV we used $\text{NSIDE} = 256$ or 786,432 grid points. Of these 786,432 points, 18,425 have $P(\log L > x) < e^{-12.5} \approx 4 \times 10^{-6}$. Next we take this list of potentially significant sources and find the local maxima of $\log L$; this reduces the number to 1,226 unique potential sources. However, 426 have $A_{\text{PSF}} < 0$, indicating not a source but a point-like deficit, leaving 800 sources. Each of these potential source positions is used more precisely to determine the local maxima of $\log L$ and get a more precise position for each source. For the lower-energy cutoff, we start with $\text{NSIDE} = 128$ or 196,608 grid points. Of these 196,608 points, 96,066 have $P(\log L > x) < e^{-12.5} \approx 4 \times 10^{-6}$ and 967 unique peaks of which 310 have $A_{\text{PSF}} > 0$.

Before discussing the results (*i.e.* the catalogue of sources and how they compare with the third Fermi catalogue, we will examine the performance of the technique and focus on the sources detected from photons with energies exceeding 1 GeV. The left-panel of Fig. 3 depicts the results of the initial set of peaks with $P(\log L) < e^{-12.5}$ before position refinement. We see that for $A_f > 0$ there is a strong linear correlation between the value of A_f and A_{PSF} . Whereas for negative values of A_f where there is a hole in the photon distribution, the value of A_{PSF} saturates at -0.1 . The right panel shows the values of A_{PSF} against the significance. There are both significant sources and holes in the photon distribution. The fractional deficit in the “hole” regions is limited to one tenth. Because of the focus of this paper is to look for gamma-ray sources, we won’t discuss these “holes” further, but they would be an interesting focus of further investigation.

For the sources in the preliminary catalogue (*i.e.* those with $A_{\text{PSF}} > 0$ we further refine the source estimate of hte source. The left panel of Fig. 4 depicts the change in $\log L$ during the optimization. Sometimes the value of $\log L$ actually decreases, but usually it increases modestly by say ten percent. The right panel shows the change in the position of the source. The size of each HEALPix region is about 0.2 degrees on a side, so an optimization within a given HEALPix cell would result in a typical movement of about one tenth of a degree, and the results bear this output. Occasionally the estimated source position moves by a large distance indicating that the optimization has found another nearby peak. We use the optimized position only if the value of $\log L$ has actually increases, and the estimated position of the source has shifted by less than 0.5 degrees.

Table 1. Basic statistics calculated for the photon distribution around a potential source

Statistic	Symbol	Abbreviation	Definition
Number of photons	N_{photons}	N	Number of photons that lie within the 95% percentile
Mean Solid Angle Ratio	\bar{r}^2	MEANR2	The mean of the ratio of solid angle enclosed between observed photon position and source location and the solid angle enclosed with the 95% percentile of the PSF
Mean Percentile Ratio	\bar{f}	MEANFRAC	The mean of the ratio of PSF percentile to 95%
Significance of MEANR2	$S(r^2)$	SIGR2	How many standard deviations is MEANR2 away from 0.5
Significance of MEANFRAC	$S(f)$	SIGFRAC	How many standard deviations is MEANFRAC away from 0.5
Fraction of photons from point source	A_f	AFRAC	If one assumes that the photons come from the sum of a uniform background and a point source, what fraction come from the point source? $A_f = (0.5 - \bar{r}^2)/(\bar{f} - \bar{r}^2)$

**Figure 3.** Left: the correlation of the value A_f , determined from the means of the photon distributions, to the value of A_{PSF} , determined from the maximum likelihood technique. Right: The values of A_{PSF} against the significance.

4 RESULTS

To test the algorithm we used essentially the same data set as used to construct the Fermi Large Area Telescope Third Source Catalog (3FGL Acero et al. 2015). We used weeks 9 through 216. That is, photons detected between 2008 August 4 (15:45:36 UTC) and 2012 July 26 (01:07:25 UTC), a span of nearly four years. This is about five days shorter than the span of the 3FGL observations. We used the good-time intervals (GTI) as defined by the Fermi team and the Pass 7 response function (P7REP_SOURCE_V15) and the Pass 7 reprocessed data. From the form of the likelihood function, Eq. 4, it is apparent that we do not use a model for the background, and the key ingredient of the instrument response is the estimate of the point-spread function (Ackermann et al. 2012).

We create two catalogues: one using photons above 1 GeV and the second with all photons above 100 MeV. The ten most significant sources in the 1 GeV map are given in Tab. 2 with the refined positions and the original detection significances. We see a general trend that the significances in sigma-units of the 3FGL are about three times that of

the FermiFAST technique. The one exception in the table is the Crab pulsar which in the Fermi analysis is split among three sources the pulsar itself, the synchrotron and inverse Compton components of the pulsar wind nebulae. The FermiFAST source most closely coincides with the pulsar, and the 3FGL significance for this component alone is quoted in the table.

All of the sources detected by FermiFAST above five sigma in the 1 GeV and the 100 MeV sample are depicted in Fig. 5. One can see that nearly all of the FermiFAST sources correspond to 3FGL sources and in fact a few may correspond to several sources. On the other hand, there are many 3FGL sources that do not have counterparts in the FermiFAST catalogue. We can examine the statistics of the correspondances more carefully by examining the distance across the sky between the nearest neighbours in each of the two catalogues: FermiFAST and 3FGL. These are plotted as the red curves in Fig. 6 and the cyan curve is a multi-Rayleigh distribution fit to the red curves. We find that the nearest neighbour in the 3FGL of each almost source in the FermiFAST catalogue lies within about 1-2 arcminutes. We

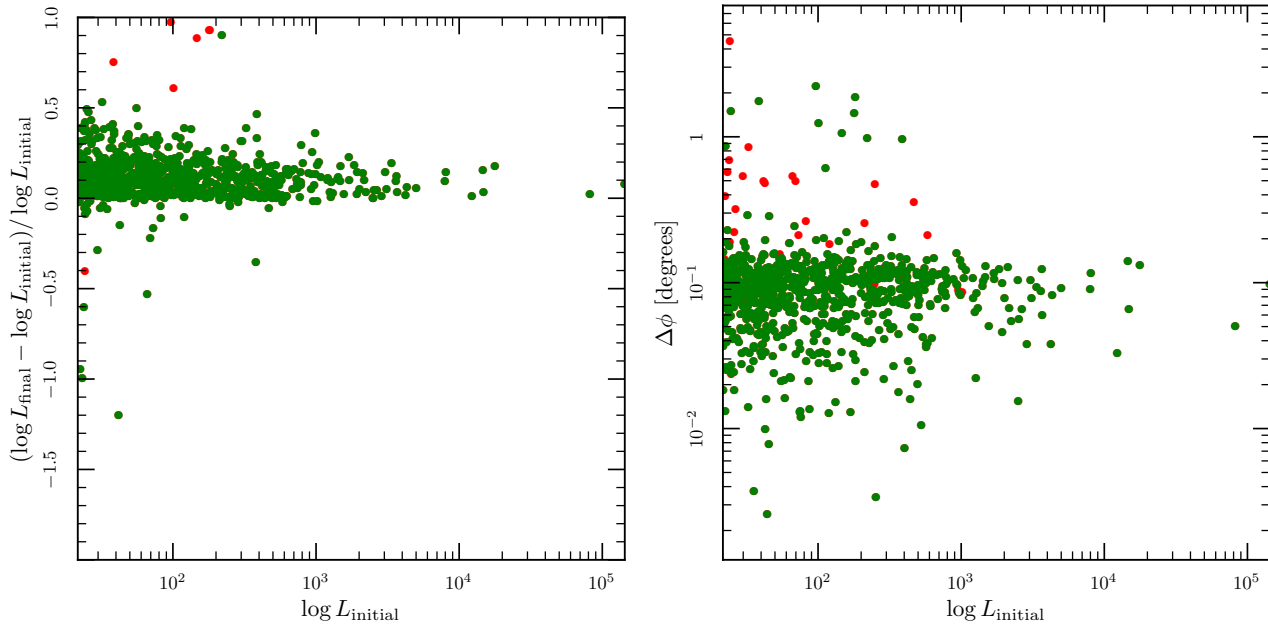


Figure 4. Left: the relative change in $\log L$ during the position refinement. Green is for sources whose positions moved less than one degree. Red is for greater movement. Right: the relative change in position during the position refinement. Green is for sources whose positions likelihood increased. Red is for those whose likelihood decreased.

Table 2. The results for the ten most significant peaks in the 1 GeV map

Source	RA	Dec	N_{photons}	\bar{r}^2	\bar{f}	$S(r^2)$	$S(f)$	A_f	A_{PSF}	$S(\text{FF})$	$S(3\text{FGL})$
PSR J0835-4510 (Vela)	128.84	-45.18	171821	0.174	0.521	-467.5	29.5	0.94	0.94	378.92	1048.96
PSR J0633+1746 (Geminga)	98.48	17.77	90467	0.164	0.501	-349.9	0.6	1.00	0.97	286.73	1012.14
PSR J0534+2200 (Crab)	83.64	22.02	26443	0.204	0.558	-166.7	32.6	0.84	0.86	133.34	30.67
LAT PSR J1836+5925	279.06	59.43	16595	0.167	0.494	-148.8	-2.5	1.02	0.96	121.85	438.12
PSR J1709-4429	257.42	-44.48	32526	0.265	0.614	-146.9	71.3	0.67	0.68	121.02	360.82
3C 454.3	343.50	16.15	14021	0.170	0.511	-135.4	4.4	0.97	0.96	110.98	480.74
LAT PSR J0007+7303	1.76	73.05	13558	0.224	0.564	-111.5	25.6	0.81	0.79	89.89	288.75
LAT PSR J2021+4026	305.39	40.45	31304	0.331	0.669	-103.8	103.9	0.50	0.48	89.31	237.05
PSR J1057-5226	164.49	-52.46	8194	0.230	0.570	-84.6	21.8	0.80	0.78	70.87	200.06
PSR J2021+3651	305.26	36.85	22683	0.355	0.693	-75.5	100.5	0.43	0.41	65.78	145.14

must assess whether these are chance coincidences. If the sources in the two catalogues are not correlated with each other (*i.e.* there are no real counterparts), then one would expect that the cumulative distribution of nearest neighbour distances to grow as a $1 - \exp(-\lambda\Omega)$ where λ is the density of sources on the sky and Ω is solid angle enclosed by a circle centred on the object and passing through the nearest neighbour. Given that there are 3029 Fermi 3FGL sources on the sky, the cumulative distribution in this case would approximately be $1 - \exp(-r^2/2\sigma^2)$ where $\sigma \approx 1.5^\circ$, so it is unlikely that these counterparts at a typical distance of 1 arcminutes are by chance.

We also can determine the distribution of unassociated pairs from the data itself by inverting the coordinates of the sources in one catalogue and looking for the nearest neighbours again. To be precise we change the sign of the Galactic latitude and longitude of each source in the 3FGL and repeat the nearest neighbour search. Here all of the correspondances will be by chance. These results are given by the green curves in the various panels and let us assess that

nearly all of the sources in the FermiFAST have counterparts in the 3FGL. In the all-sky map (left panel) in the FermiFAST, we can assess the number of FermiFAST sources that are not in the 3FGL by fitting the cumulative distribution with several Rayleigh distributions that quantify the positional error for the true counterparts and the chance of a false counterpart determined by fitting the cumulative distribution of false counterparts with a Rayleigh distribution as well. We find that about 50 of the 800 sources do not have a counterpart in 3FGL. If we now focus on the right panel of Fig. 6 we find that out of the 580 high-latitude sources in FermiFAST at most only a few lack a counterpart in 3FGL.

One thing that is glaringly obvious from Fig. 5 is that most sources in the 3FGL lack a counterpart in the FermiFAST catalogue. Both the 1 GeV and the 100 MeV version. Given that the algorithm outlined here is much less comprehensive than the techniques used to generate the Fermi catalogues (there is no modelling of the background, multiple sources or spectra), this is not surprising. However, it is useful to figure out what sources are missing from the Fer-

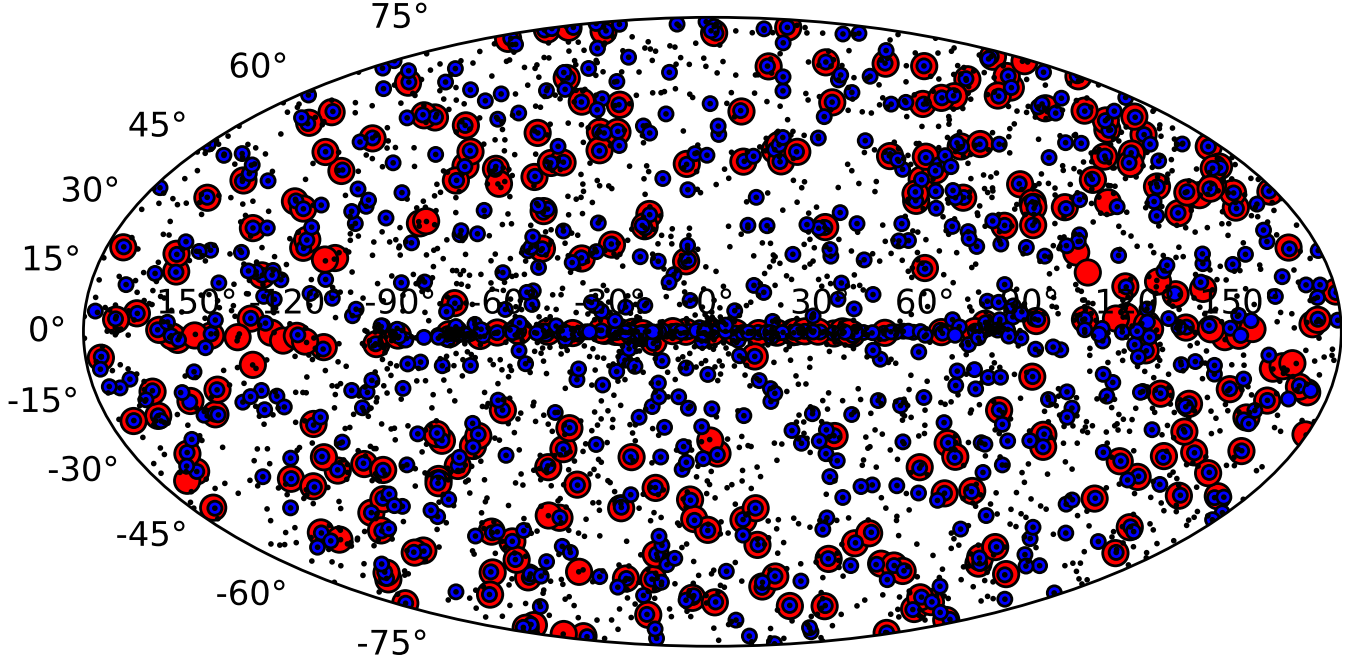


Figure 5. Sources with $\ln P(\log L) < -12.5$ (i.e. five sigma) in the 1 GeV catalogue (small blue circles), in the 100 MeV catalogue (big red circles) and the 3FGL (black dots).

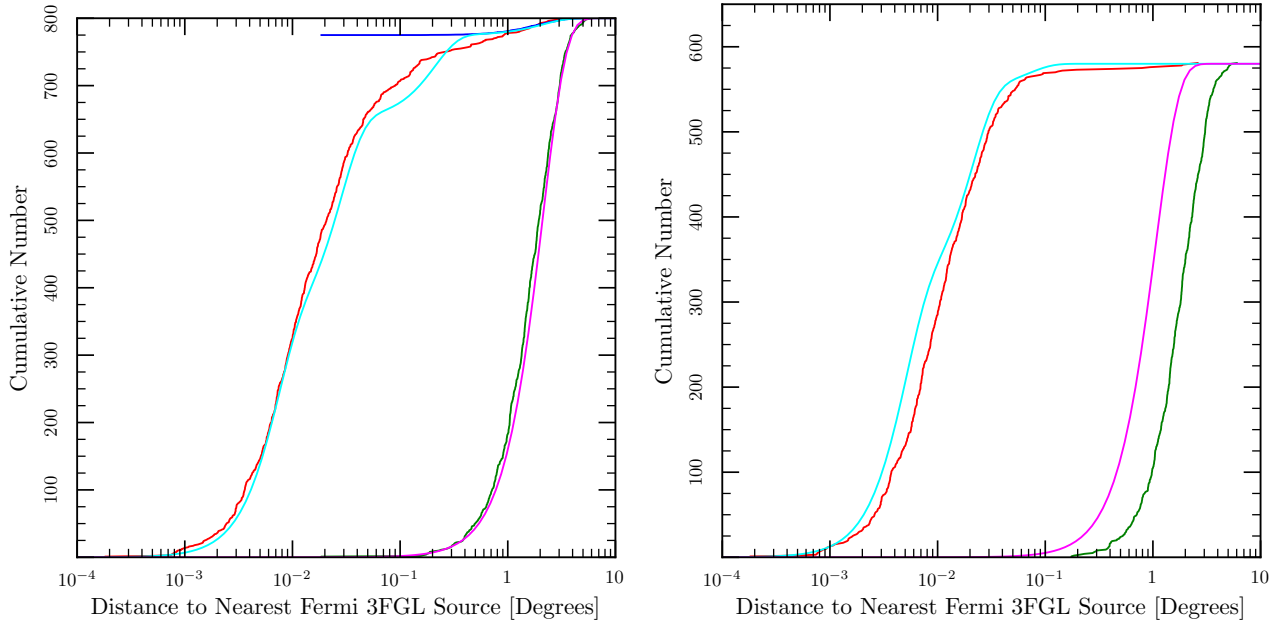


Figure 6. The left panel gives the results for all sources, and the right panel has $|b| > 10^\circ$. The distance from a Fermi FAST source to the nearest Fermi 3FGL source. This demonstrates that ninety percent of the Fermi FAST sources are associated with sources in the FERMI 3FGL (across the whole sky) and nearly all at Galactic latitudes greater than ten degrees. The red curve give the observed cumulative distribution of nearest distances. The green curves give the cumulative distribution that one would expect if there were no associated between the Fermi FAST and Fermi 3FGL sources. This is calculated by performing the same analysis as the red curves but with the Galactic coordinates inverted. The blue curve yields the false positive rate. The cyan and magenta curves are Rayleigh distributions that are fit to the observed distributions. The typical positional error between associated 3FGL and FAST sources is 1-2 arcminutes.

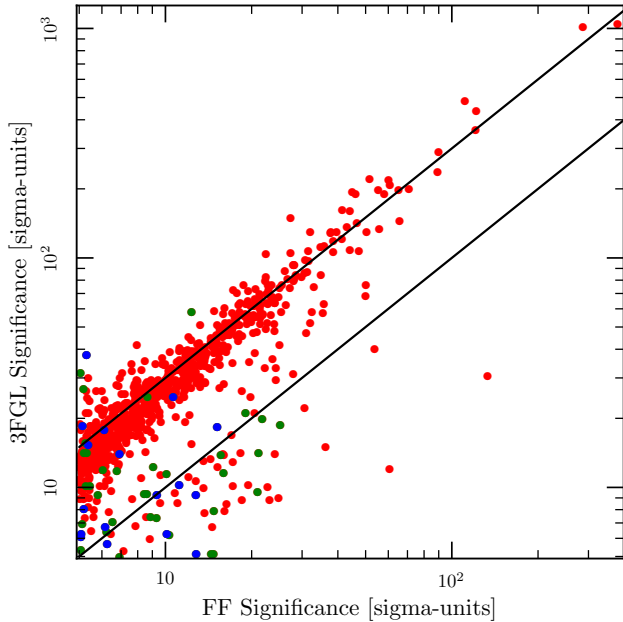


Figure 7. 3FGL Significances vs Fermi FAST. The red points are where the counterparts lie within 0.2° of each other, the green the counterparts lies between 0.2° and 1° , and the blue have counterparts further than 1° away.

miFAST catalogue and why. Fig. 7 shows the relationship between the significances of a source assigned by FermiFAST and by the 3FGL. In particular the 3FGL significances are three times larger (the upper line) than FermiFAST typically, so if both apply a threshold of five sigma, FermiFAST will find fewer sources. However, there is a large population of sources for which the two likelihoods are similar or the FermiFAST likelihood is larger. Since FermiFAST does not perform any spectral modelling, perhaps these sources have poor spectral fits in the 3FGL, yielding smaller likelihoods. Furthermore, this indicates a discovery space for FermiFAST to find point sources where we do not have a good prior notion of the spectral model. The outlying point in the lower-right is the Crab pulsar whose 3FGL likelihood is artificially too low due to the way it is fit within the catalogue.

To understand further whether FermiFAST is simply missing less significant 3FGL sources we can look at the cumulative distribution of 3FGL significances of sources that appear in the FermiFAST 1 GeV catalogue and all of the 3FGL sources in Fig. 8. We see that FermiFAST catalogue is essentially complete for all 3FGL sources above about 18-sigma, so FermiFAST can quickly (in six minutes) generate a sample from the Fermi data stream of the upper quartile of sources that would appear a Fermi catalogue using the full likelihood technique to construct the catalogue.

The question arises: is it possible to do better? We can reduce the significance threshold for find sources in the 1-GeV FermiFAST catalogue. We expect that the number of sources that do not have firm associations in the 3FGL to increase but also for the completeness to increase as well. We can use the empirical distance distribution for the false matches depicted in yellow in Fig. 9 to split the distribution of the matches between Fermi FAST and the 3FGL statist-

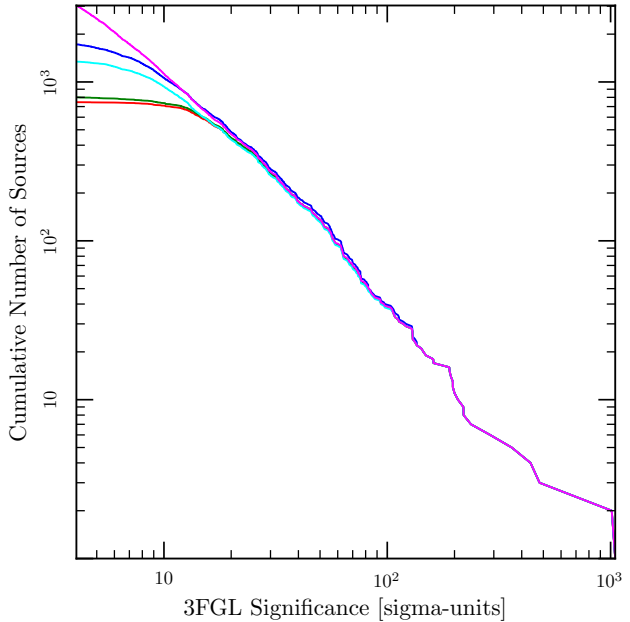


Figure 8. Cumulative distribution of 3FGL significances for the entire 3FGL (blue) and the entire Fermi FAST 1-GeV catalogue with a five-sigma cut (green) and Fermi FAST where the 3FGL counterpart lies within 0.2° (red) again with a five-sigma cut. The blue and cyan give the same as green and red but with a three-sigma cut.

Table 3. Completeness and Purity

Threshold	Sources	True	False	Purity	Comp
$3 - \sigma$	1727	1374	352	79.8%	48.4%
$3.5 - \sigma$	1312	1178	134	89.8%	41.8%
$4 - \sigma$	1076	1009	67	93.8%	36.2%
$4.5 - \sigma$	923	877	46	95.1%	31.3%
$5 - \sigma$	800	765	35	95.8%	27.0%

cally into the true matches and the false matches to measure the completeness and purity of the samples. To do this we assume that all of the true matches are closer than any of the false matches and use a ranked-sum-test to determine the fraction of true matches from the observed distribution of matches (see Appendix for details). For the 1-GeV catalogue this appears to be a good assumption. The contribution of the false matches to the cumulative distributions is depicted in the same colour as the measured distributions but dashed. The distribution of the close matches here is somewhat different than in Fig. 6 because we have not used the improved localizations to find the matches. The localization improvement tends to fail more often for sources of low significance (see Fig. 4). Tab. 3 gives the results of this trial. The key result is that the five-sigma sample is very pure; only about five percent of the sources lack firm associations in the 3FGL. On the other hand, if one is willing to sacrifice the purity of the sample, one can achieve nearly 50% completeness relative to the 3FGL by reducing the significance threshold to $3 - \sigma$. Fig. 8 shows by accepting a lower purity one can get a complete 3FGL sample about about ten-sigma, comprising nearly half of the 3FGL sample.

We saw in Fig. 7 that in general the FermiFAST sig-

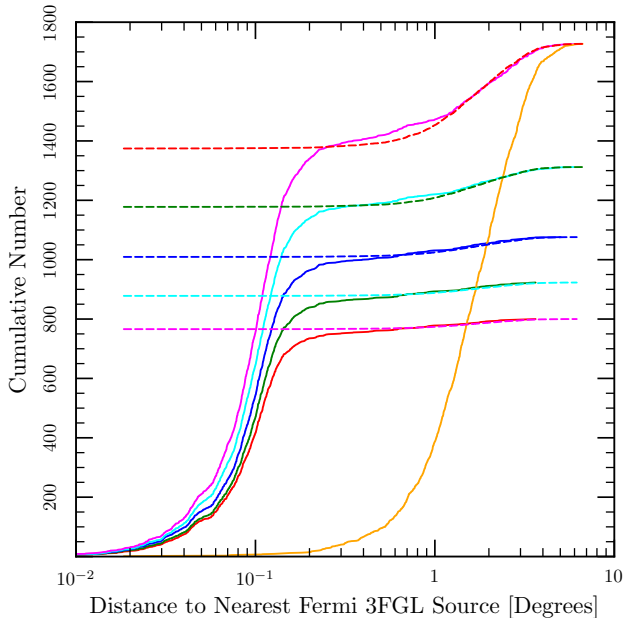


Figure 9. The cumulative distribution of distances between sources in Fermi FAST and 3FGL for various significance thresholds from top to bottom of 3-sigma, 3.5-sigma, 4-sigma, 4.5-sigma and 5-sigma. The dashed curves that start above zero and join each cumulative distribution at large radii trace the cumulative distribution of distances for the false matches in each sample.

nificances and the 3FGL significances are well correlated with However, for about one fifth of the sources $S(3FGL) < 2S(FF)$ and for about one tenth $S(3FGL) < S(FF)$. This demonstrates that at least for a fraction of the 3FGL the FermiFAST technique is more sensitive, so the sources that do lie in the FermiFAST catalogue but not the 3FGL may just be a class of sources to which the 3FGL is less sensitive, perhaps sources for which the suite of spectral models employed in the construction of the 3FGL to do apply. On the other hand, these unassociated sources could simply be regions where the diffuse background is especially bright and clumpy. Because the FermiFAST technique assumes that the background is smooth on the scale of the PSF, it would find a source where the background is locally strong. We possibly see the opposite of this effect in Fig. 3 where FermiFAST finds negative sources that are most likely places where the background is locally weak.

It is apparent from Fig. 5 that there are fewer 100-MeV sources and that most of these sources coincide with 3FGL and 1-GeV FAST sources. In Fig. 10 the localization is slightly poorer than Fig. 9 because we have used a coarser HEALPix grid to sample the likelihoods. For the 100-MeV catalogue we have not optimized the likelihood to determine more accurate positions, so Fig. 9 is the best comparison not Fig. 6. Again we have used a rank-sum test to determine the sources that are likely to be shared between the catalogues and also to determine those sources that appear in the 100-MeV catalogue but not the others. The 100-MeV catalogue contains 312 sources of which 23 are in the 3FGL but not the 1-GeV catalogue, 16 are in the 1-GeV catalogue but not the 3FGL and 24 are in neither of the other catalogues (249 are in all three). Relative to the 3FGL catalogue the purity

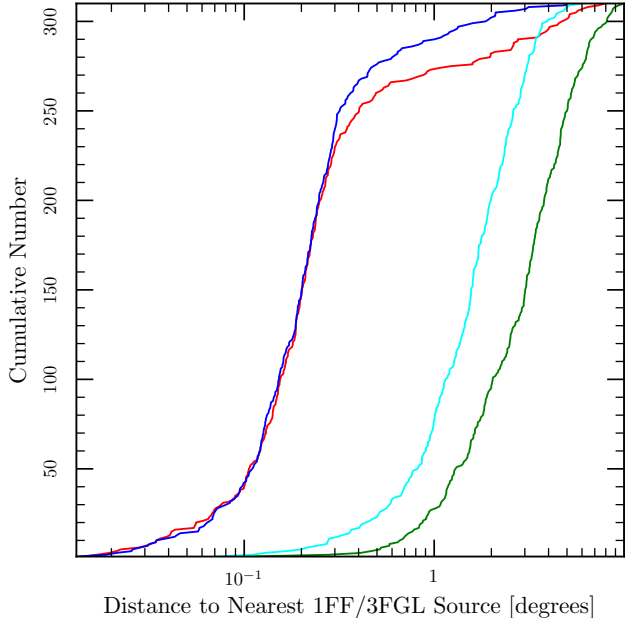


Figure 10. The cumulative distribution distances between sources in the 100-MeV catalogue and the 3FGL (blue, leftmost and higher) and the 1-GeV catalogue (red, second from left) and for the false matches between the 100-MeV and 3FGL (cyan, second from right) and 100-MeV and 1-GeV catalogue (green, rightmost).

is 88.1% and the completeness is 12.1%. If one takes the union of the five-sigma 1-GeV and the five-sigma 100-MeV FermiFAST catalogues, there is a total of 788 sources that correspond to 836 3FGL sources (*i.e.* one FermiFAST source can correspond to several 3FGL sources) for a slightly higher completeness fraction of 27.6% than that of the five-sigma 1-GeV FermiFAST catalogue.

5 DISCUSSION

We have outlined a technique that can rapidly find point sources from the photon data from the Fermi telescope. If one focusses on the photons above 1 GeV, one can reproduce about half of the 3FGL using the 3FGL data set in about six minutes with a false association rate of less than twenty percent. There are several immediate applications of this technique that come to mind as well as several avenues of potential improvements to increase the sensitivity and specificity and the data that we learn about each source. First, we will discuss some potential applications.

The speed of the technique opens several possible exploratory research paths. One can search fractions of the data stream for potential sources that we active for a portion of the complete 3FGL observation time and so did not reach the threshold for detectability over the entire observation but were sufficiently bright to be detected over a shorter window. Given the large variability of blazars which form the bulk of the sources outside the plane of the Galaxy, this is an interesting avenue. In fact although the 3FGL catalogue contains many more sources than the 2FGL, many sources are absent in the 3FGL that were in the 2FGL. Some of

these may have been spurious, but others could have simply been variable, so the search for variable sources provides an exciting avenue. Furthermore, the sources in the FermiFAST catalogue that are not in the 3FGL would be great seeds for a full characterization using the Fermi tools. Are these potential sources simply knots in the gamma-ray background or do they remain significant even with the full analysis?

A second important direction is to use the FermiFAST algorithm to generate a catalogue of sources, most likely not as deep as the 3FGL and then use the catalogue for further statistical analysis of the various populations. Now what are the advantages of this when one considers that the 3FGL already exists? Because the FermiFAST catalogue only requires a few minutes to generate one can explore the statistical biases, selection and uncertainties in the catalogue through simulated data much more easily than for the full 3FGL catalogue. A first step would be to insert artificial sources of known fluxes and spectra into the data stream. This could be done in parallel by again inserting the sources on a HEALPix grid so that the regions of interest for each source (Fig. 2) do not overlap over the energy range considered. Now we can measure the relationship between the flux and significance throughout the sky to obtain the selection function for the survey. This is crucial ingredient to determine the luminosity function of the underlying population. Because the underlying sources may be variable, we could also insert variable sources into the data stream as well and combine this with the time slicing analysis described in the previous paragraph and develop for example a catalogue of blazars as a function of mean luminosity and variability that is both statistically well characterized and possibly contains new objects that were not discovered in the previous Fermi catalogues. Of course, many of the sources will have already been detected by Fermi and possibly followed up with further measurements, the key element here would be that the sample would be statistically well characterized, so clear conclusions about the population could be obtained.

Using the existing data stream without inserting sources, one could characterize whether the measurement significances of the individual sources are reasonable and whether the variability measurement determined say in an analysis as outlined earlier is significant. The first step would be to resample the photon data stream through a bootstrap — this would generate a new list of photons by selecting the same number of photons from the original list but with replacement, so that the same photon could appear many times in the new sample and of course some photons won't appear at all. Each new list of photons would generate new catalogue to determine what properties of the sources and the catalogue as a whole are robust against the Poisson fluctuations in the data stream. This would verify the underlying assumptions of the significances of the sources as well as determine how sources near the detection threshold enter and leave the sample. Because the sample is dominated by the faintest sources near the detection threshold (see Fig. 7), understanding this effect is again crucial to understand the sample.

To detect point sources from the Fermi photon data stream, we only used the data stream and the estimates of the point-spread function. There is additional data available that could help to improve the sensitivity of the technique, reduce the number of false positives and better characterize

the sources detected. First, the likelihood function, Eq. 4, assumes that the gamma-ray background is smooth everywhere. However, we know from independent measurements, for example the Galactic synchrotron emission, that the gamma-ray background is likely to be clumpy and the Fermi team provides estimates of this background (e.g. Acero et al. 2016). By combining this background estimate along with the integral of the effective area and time that a particular direction and energy was observed, we can develop an improved likelihood function that includes the estimated background. This could reduce both the number of candidate sources detected with $A_{\text{PSF}} < 0$ and the number of candidate sources without counterparts in the 3FGL. The addition of this information would not dramatically increase the time to construct the catalogue (we have already included this functionality in the software but it is not yet well characterized).

One could also take an alternative route. Looking again at Eq. 4, we can use the relative size of the two terms in the brackets to assign a fraction of each photon to the source and to the background. Again if we include information on the effective area and the exposure time, we obtain an estimate of the spectrum from each source and from the background. Because a given photon could be assigned to several sources as well as the background, we would also have to develop a good statistical model to split the photon among the sources. Of course, a first estimate could be obtained by ignoring the fact that several potential sources could claim a given photon and resolve the gamma-ray data into sources and background (as done by Selig et al. 2015). This would provide an independent confirmation of the structure of the gamma-ray background, but perhaps more importantly it would yield an estimate of the spectrum of each source. One could include a fit to the source spectrum as in the standard Fermi analysis either in the initial likelihood calculation or after the sources have been detected. In the first case this might improve both the sensitivity as we could find fainter sources against the background if they follow one of the spectral models and the specificity by excluding potential sources that do not follow the models.

FermiFAST provides the infrastructure both for exploratory analysis of the Fermi data to find new sources and new variable sources and also to create a sample of sources for population studies that can be characterized comprehensively by Monte Carlo simulation of the detection technique. The latter may allow new understanding of the population of Galactic and extragalactic gamma-ray sources and their evolution. In this paper we have purposefully used the Pass 7 response functions and just four years of data to compare with the 3FGL. Nearly an additional four years of data are now available with the Pass 8 reductions (Atwood et al. 2013), so a sneak preview of the 4FGL would be a very natural next step.

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APPENDIX

We use a ranked-sum test to estimate the fraction of coincidences between members of two catalogues that are likely to be true counterparts rather than chance coincidences. We have the distribution of the nearest-neighbour separations between the two catalogues and also an estimate of the distribution of the nearest-neighbour separations for the chance coincidences. We furthermore assume that the true matches all have smaller separations than the chance coincidences.

The basic idea is to compare the list of observed separations to the list of separations of chance coincidences in pair-by-pair match ups where the winner is the member of the pair that is smaller. If there were actually no true counterparts in the observed list, then the observed list and the list of chance coincidences would have the same distribution and one would expect the half of pairwise matches to be won by members of the observed list and half to be won by members of the chance list. On the other hand, if the entire observed list consists of true matches, then the members of the observed list should win all of the time, so we have

$$f_{\text{True Matches}} = 2f_{\text{Observed Wins}} - 1.$$

This formula holds even if the two lists have differing numbers of elements.

To calculate the fraction of wins, we do not actually perform a pair-by-pair comparison which would require $\mathcal{O}(N^2)$ comparisons but rather sort the two lists together, $\mathcal{O}(N \ln N)$, and keep track of which elements come from which list as in the following Python snippet:

```
rall=np.concatenate((r1,r2))
n1all=np.concatenate((np.ones(len(r1)),
                        np.zeros(len(r2))))
n2all=np.concatenate((np.zeros(len(r1)),
                        np.ones(len(r2))))
index=np.argsort(rall)
n1all=n1all[index]
n2all=n2all[index]
c1all=np.cumsum(n1all)
obswins=np.sum(n2all*c1all)
frac=2*obswins/(len(r1)*len(r2))-1
```

where `r1` contains the list of observed separations of the nearest neighbours in the two catalogues and `r2` contains the list of separations of the chance nearest neighbours.