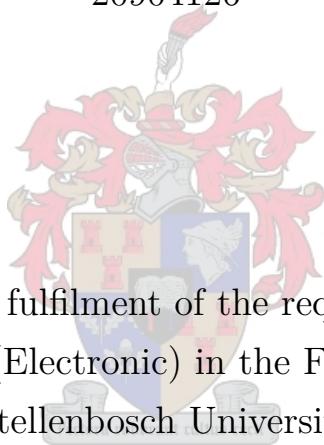


Diagnosis of Sensor Anomalies to Ensure Fault-Tolerant Control

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Abstract

English

The English abstract.

Afrikaans

Die Afrikaanse uittreksel.

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Nomenclature

Variables and functions

$p(x)$	Probability density function with respect to variable x .
$P(A)$	Probability of event A occurring.
ε	The Bayes error.
ε_u	The Bhattacharyya bound.
B	The Bhattacharyya distance.
s	An HMM state. A subscript is used to refer to a particular state, e.g. s_i refers to the i^{th} state of an HMM.
\mathbf{S}	A set of HMM states.
\mathbf{F}	A set of frames.
\mathbf{o}_f	Observation (feature) vector associated with frame f .
$\gamma_s(\mathbf{o}_f)$	A posteriori probability of the observation vector \mathbf{o}_f being generated by HMM state s .
μ	Statistical mean vector.
Σ	Statistical covariance matrix.
$L(\mathbf{S})$	Log likelihood of the set of HMM states \mathbf{S} generating the training set observation vectors assigned to the states in that set.
$\mathcal{N}(\mathbf{x} \mu, \Sigma)$	Multivariate Gaussian PDF with mean μ and covariance matrix Σ .
a_{ij}	The probability of a transition from HMM state s_i to state s_j .
N	Total number of frames or number of tokens, depending on the context.
D	Number of deletion errors.
I	Number of insertion errors.
S	Number of substitution errors.

Acronyms and abbreviations

AOCS	Attitude and Orbit Control System
ADCS	Attitude Determination and Control System
EKF	Extended Kalman Filter
FDIR	Fault Detection, Isolation and Recovery
EIC	Earth Inertial Coordinate
EFC	Earth Fixed Coordinate
GHA	Greenwich Hour Angle
ORC	Orbit-referenced Coordinate
SBC	Satellite Body Coordinate
DCM	Direct Cosine Matrix
TLE	Two-line Element
RAAN	Right Ascension of the Ascending Node
AP	Argument of Perigee
SGP	Simplified General Perturbations
IGRF	International Geomagnetic Reference Field
IAGA	International Association of Geomagnetism and Aeronomy
LEO	Low Earth Orbit
CoM	Centre of Mass
CoP	Centre of Pressure
RW	Reaction Wheel
FoV	Field of View
DMD	Dynamic Mode Decomposition
CART	Classification and Regression Trees
BST	Binary Search Tree
LOF	Local Outlier Factor
IRC	Inertial-referenced Coordinate

Chapter 1

Introduction

1.1. Background

All systems are prone to failure and one of an engineer's tasks is to ensure that crucial systems never fail and that continual maintenance is done on a system. For satellite's the problem is even more severe, and most failures are unrecoverable. Therefore, satellite systems must be tested thoroughly and be robust to any anomaly. According to [?] the attitude and orbit control system (AOCS) contributes to the largest percentage of satellite failures as shown in Figure 1.1. A study conducted by [?] on small satellite mission failures provide a deeper insight into the attitude control's role in satellite failures, since most missions are highly dependent on the attitude control for complete mission success. This is due to many mission specifications that rely on accurate control of the satellite to ensure that payloads such as cameras are able to operate as required.

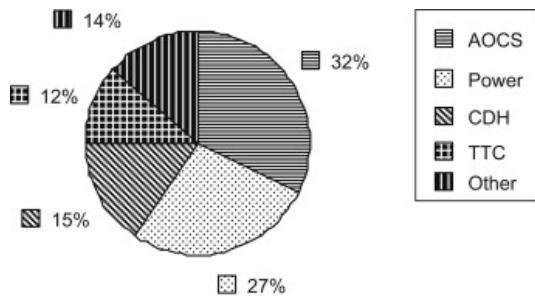


Figure 1.1: On-orbit failure per subsystem [?]

The database provided by [?] demonstrates the increase in the number of satellites launched every year. The drastic increase of launched satellites in the past few years even emphasizes this need of ensuring that satellites are robust to failures. Therefore, using traditional methods of tracking satellites with ground stations and manually checking for possible failures are not feasible, especially with satellite constellations. Consequently, most aspects of the satellite should operate autonomously, especially for attitude control. The focus of this thesis is on a specific aspect of the attitude control that is sensitive to anomalies, namely the attitude determination.

The extended Kalman filter (EKF) is implemented for attitude estimation of the satellite. Therefore, the objective of this thesis is to develop methods to avoid unstable

and inaccurate estimations from the EKF caused by sensor anomalies.

1.2. Problem Description

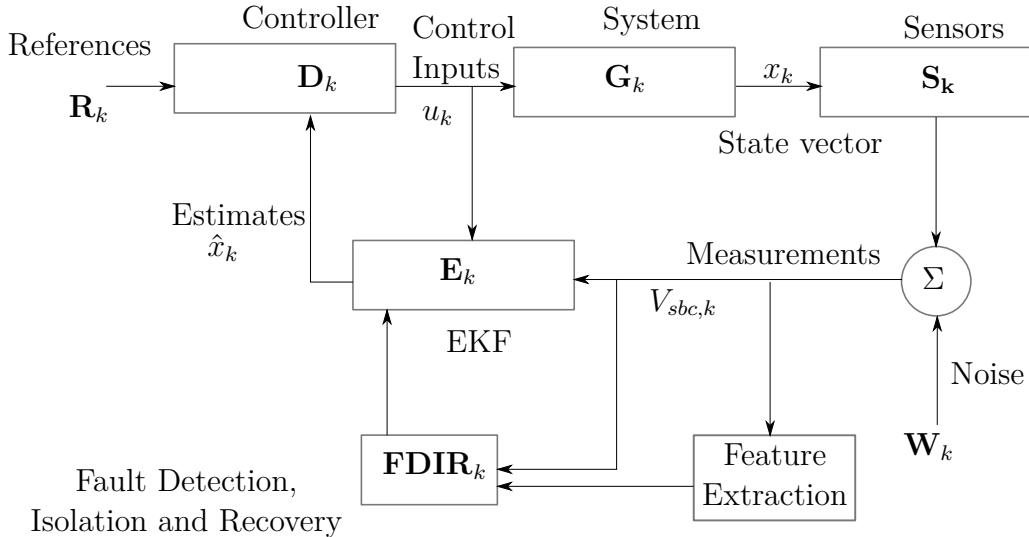
For many satellite missions, attitude determination and control system (ADCS) is of high importance. It is required to effectively control the attitude to fulfill the required mission requirements. The control performance is also limited by the estimation accuracy and performance. A classical satellite attitude requirement is to be earth-following during eclipse, point the payload to a target and otherwise point and track the sun for solar charging. Good attitude estimation throughout the entire orbit is required to best fulfill the requirements.

The attitude estimation is however highly influenced by the sensor readings. Different sensor measurements are fused together, normally with the use of an EKF, to produce a single attitude estimate. If an erroneous, or false measurement is present in the collection of sensors, it will deter and influence the outcome of the fusion algorithm. Depending on the number of sensors available and the size of the error the result can be either just reduced estimation accuracy or even divergence and instability. It is good practice to develop appropriate tests to protect the EKF against incorrect measurements.

Anomaly detection in satellite sensors have been investigated in many previous research. The current trend is to use generic sensor anomalies, such as bias drift, high noise, sudden failure or any drastic change in the behavior of the sensor. This is only a subset of possible errors and does not assist in diagnosing the anomaly, detecting intermittent errors, or coupled events between sensors. An example of a practical anomaly which can occur and which is difficult to detect using standard techniques are solar reflections from solar panels on a sun sensor. Majority of satellites, even with relatively low attitude requirements, have some form of sun sensor. The sun sensor also provides an accurate measurement during the periods of the orbit where targeting and solar tracking is most likely and where the attitude requirement is the highest. Thus it would be beneficial to have good interventions to ensure robust sun vector measurements for the EKF.

Since the EKF is also reliant on the mathematical model of the system the control inputs should also be accurate. Due to actuator failure on satellites, the command control input and the actual control input can differ significantly. This must also be detected and isolated to ensure robust estimation. The satellite must be able to detect, classify (isolate) and recover from the anomaly, autonomously to ensure safe operation during orbit. A block diagram for the fault detection, isolation and recovery (FDIR) of the EKF within the ADCS system is provided in figure 1.2.

The anomalies discussed and modelled in this thesis are specific to the design of the satellite. The attitude sensors and the anomalies for each sensor is a sun sensor with solar reflection from the solar panels, a infrared-nadir sensor with the moon on the earth horizon

**Figure 1.2:** System Diagram

and a magnetometer with magnetic disturbances caused by the magnetic induced dipole moment of the solar panels. The actuator failure is that of a reaction wheel not responding to control inputs. These anomalies should be accurately detected and classified to ensure a reliable EKF. The recovery methods after classification will be discussed in chapter 7.

1.3. Project definition

This project aims to develop a simulation model wherein practical anomalies are simulated during satellite orbit. The simulation must then be used to create a database of sensor measurements produced by different anomalies. This database provides labelled data for the training of binary and multi-class classification models for detection and isolation respectively. The trained models should be tested on the simulation environment and the estimation accuracy should be compared between different models and different recovery methods.

A comparison of model accuracy depending on training data from generic sensor anomalies vs practical anomalies should also be discussed. A thorough analysis of the methods and the results should be used to discuss the difference in estimation accuracy based on the model's training on generic or specific anomalies.

1.4. Thesis outline

Chapter 1 provides the background and motivation for this research as well as the project definition and thesis outline.

Chapter 2 discusses the relevant research that has been done on the subject of FDIR. Chapter 3 demonstrated the development and implementation of the simulation environment.

Chapter 4

Chapter 5 provides the mathematical models of the specific anomalies and the effects thereof on the satellite.

Chapter 6 describes the feature extraction methods used in this thesis to enhance the accuracy of the detection and classification models.

Chapter 7 provides various recovery methods and demonstrates the theoretical possibility of the methods based on perfect prediction accuracy.

Chapter 8 describes the different algorithms and methods used to detect an anomaly in the system.

Chapter 9 describes the different algorithms and methods used to classify an anomaly in the system.

Chapter 10 provides a summary of all the results for the combination of best methods as provided in chapter 7, chapter 8 and chapter 9.

Chapter 11 discusses the influence of modelling specific anomalies on the prediction accuracy and robustness of a EKF.

Chapter 2

Literature Study

The implementation of FDIR on satellites have multiple complications with regards to the type of data generated by a satellite and the methodologies that can be implemented within the time and memory constraint of a cube-sat processor.

TODO: Research within the larger field of FDIR diagram

2.1. Anomaly Detection on Satellites

Various methodologies have been tested on different component of satellites. Therefore a summary of these research articles are provided in this section.

2.1.1. Analysis and Prediction of Satellite Anomalies

2.1.2. Agent-based algorithm for fault detection and recovery of gyroscope's drift in small satellite missions

To ensure that the ADCS of satellites are autonomous every aspect of the control must be able to recover from faults. [?] developed an algorithm to evaluate the control of a gyroscope and detect whether drifting exists. If drifting is detected another algorithm is deployed to ensure the recovery of the gyroscope drift by updating the error state vector.

Multivariate Anomaly Detection in Discrete and Continuous Telemetry Signals Using a Sparse Decomposition in a Dictionary [?]

Fault isolation of reaction wheels onboard three-axis controlled in-orbit satellite using ensemble machine learning [?]

Fault tolerant control for satellites with four reaction wheels [?]

Innovative Fault Detection, Isolation and Recovery Strategies On-Board Spacecraft: State of the Art and Research Challenges [?]

Machine learning methods for spacecraft telemetry mining [?]

Machine learning techniques for satellite fault diagnosis [?]

Satellite fault diagnosis using a bank of interacting Kalman filters [?]

A scheme of satellite multi-sensor fault-tolerant attitude estimation [?] implements a fault tolerant federated Kalman filter with three sub-filters for multi-sensor

fault estimation.

Detection of satellite attitude sensor faults using the UKF [?] provides a fault detection method by using the residuals generated by an unscented Kalman filter to detect anomalies with a threshold based on a confidence level.

Sensor fault detection and recovery in satellite attitude control [?]

Sensor Failure Detection in Dynamical Systems by Kalman Filtering Methodology While methods for sensor failure detection in other dynamical systems has also been developed which includes kalman filter methodology [?],

Sensors Anomaly Detection of Industrial Internet of Things Based on Isolated Forest Algorithm and Data Compression isolation forests [?] and using LSTM on sensor data to detect anomalies on machines

LSTM-based Encoder-Decoder for Multi-sensor Anomaly Detection [?]

Sensor fault detection and isolation using adaptive extended Kalman filter [?]

2.2. Summary

Chapter 3

Simulation

To implement and research various FDIR systems on satellites a simulation of satellite dynamics and kinematics is developed. The focus of this thesis is on small satellites and more specifically CubeSats. For the simulation of the ADCS of the satellite [?, ?, ?] were referenced during the development of the satellite simulation. The simulation was developed in Python to simulate the dynamics and kinematics during a satellite orbit.

TODO: Decide whether satellite design should be within this chapter.

3.1. Attitude Determination and Control System

For the mission of the specific satellite in this document the main operational goal of the ADCS on this specific satellite mission is to control the payload to point towards the centre of the Earth during eclipse and point the solar panels towards the sun during the sunlit phase. To ensure this is accurately simulated the different coordinate frames dominating a satellite orbit, the attitude of the satellite as well as the satellite dynamics and kinematics is discussed in this section.

3.1.1. Coordinate Frames

The coordinate frames in aerospace is a fundamental part of the ADCS. To determine the orientation and position of an object, it should be relative to a fixed frame. Consequently, the Earth inertial coordinate (EIC) frame, $\mathcal{E}\{\bar{x}_\mathcal{E}, \bar{y}_\mathcal{E}, \bar{z}_\mathcal{E}\}$, is the fixed frame from which every other frame is relative to.

A coordinate frame consists of three orthogonal vectors which is commonly referred to as x, y, and z. The axis of the coordinate frame is appropriately named as the X-axis, Y-axis and Z-axis. A vector (\mathbf{r}) within the coordinate frame can thus be expressed as

$$\mathbf{r} = x\mathbf{i} + y\mathbf{j} + z\mathbf{k} \quad (3.1)$$

where the magnitude of \mathbf{r} is denoted as $\|\mathbf{r}\|$ and is equal to

$$\|\mathbf{r}\| = \sqrt{x^2 + y^2 + z^2}. \quad (3.2)$$

The Earth-centered coordinate frames are divided into two, namely the EIC and Earth fixed coordinate (EFC) frame, $\mathcal{F}\{\bar{\mathbf{x}}_{\mathcal{F}}, \bar{\mathbf{y}}_{\mathcal{F}}, \bar{\mathbf{z}}_{\mathcal{F}}\}$. EFC is fixed to the Earth and rotates with it, while EIC is inertial fixed.

The EIC is defined as the Z-axis pointing towards the north pole, the X-axis pointing towards the Vernal Equinox, Υ , and the Y-axis completing the orthogonal set. The EFC is a copy of the EIC, with the Z-axis being identical, however the EFC rotates with the Earth. The EFC in relation to the EIC can be expressed by a single angle of rotation, which is the Greenwich Hour Angle (GHA), α_G . With the elapsed time, t , since t_0 , the angular rate of the Earth, ω_E , and the GHA, $\alpha_{G,0}$, at $t = t_0$, α_G can be calculated as

$$\alpha_G = \omega_E t + \alpha_{G,0}. \quad (3.3)$$

To transform a vector from one coordinate frame to another, a transformation matrix, \mathbf{A} , is required. For example vector $\mathbf{r}_{\mathcal{F}}$ can be transformed to $\mathbf{r}_{\mathcal{E}}$ with

$$\mathbf{r}_{\mathcal{E}} = \mathbf{A}_{\mathcal{F}}^{\mathcal{E}} \mathbf{r}_{\mathcal{F}} \quad (3.4)$$

with $\mathbf{A}_{\mathcal{F}}^{\mathcal{E}}$ being the EFC-to-EIC transformation matrix. Due to the definition of both coordinate frames, $\mathbf{A}_{\mathcal{F}}^{\mathcal{E}}$ can be defined as

$$\mathbf{A}_{\mathcal{F}}^{\mathcal{E}} = \begin{bmatrix} \cos(\alpha_G) & -\sin(\alpha_G) & 0 \\ \sin(\alpha_G) & \cos(\alpha_G) & 0 \\ 0 & 0 & 1 \end{bmatrix}. \quad (3.5)$$

To determine the satellite position, satellite-centred coordinate frames must be used. Three satellite-centred coordinate frames are used, namely the inertial-reference coordinate frame (IRC), $\mathcal{I}\{\bar{\mathbf{x}}_{\mathcal{I}}, \bar{\mathbf{y}}_{\mathcal{I}}, \bar{\mathbf{z}}_{\mathcal{I}}\}$, which remains inertial fixed, the orbit-referenced coordinate (ORC) frame, $\mathcal{O}\{\bar{\mathbf{x}}_{\mathcal{O}}, \bar{\mathbf{y}}_{\mathcal{O}}, \bar{\mathbf{z}}_{\mathcal{O}}\}$ and the satellite body coordinate (SBC) frame, $\mathcal{B}\{\bar{\mathbf{x}}_{\mathcal{B}}, \bar{\mathbf{y}}_{\mathcal{B}}, \bar{\mathbf{z}}_{\mathcal{B}}\}$.

The ORC frame changes location as the satellite moves, however the Z-axis is always pointing towards the centre of the Earth, with the Y-axis being the orbit anti-normal and the X-axis completing the orthogonal set. To transform a vector from the EIC frame to the ORC frame the unit position vector, \mathbf{r}_{sat} and the unit velocity vector, \mathbf{v}_{sat} in EIC is required [?]. The EIC to ORC transformation matrix, $\mathbf{A}_{\mathcal{E}}^{\mathcal{O}}$, is calculate as

$$\begin{aligned} \mathbf{A}_{\mathcal{E}}^{\mathcal{O}} &= [\mathbf{u} \ \mathbf{v} \ \mathbf{w}]^T \\ \text{where } \mathbf{w} &= -\frac{\mathbf{r}_{sat}}{\|\mathbf{r}_{sat}\|} \\ \mathbf{v} &= -\frac{\mathbf{r}_{sat} \times \mathbf{v}_{sat}}{\|\mathbf{r}_{sat} \times \mathbf{v}_{sat}\|} \\ \mathbf{u} &= \mathbf{v} \times \mathbf{w}. \end{aligned} \quad (3.6)$$

The SBC frame is the frame fixed to the satellite and it is the relative rotation of the

satellite in relation to the ORC. Thus for the mission of this satellite it is required that the SBC and ORC frames coincide during eclipse. For the transformation of a vector from the ORC to SBC frame, the direct cosine matrix (DCM) also referred to as \mathbf{A} or $\mathbf{A}_{\mathcal{O}}^{\mathcal{B}}$ is used. For the remainder of the document the DCM will be referred to as $\mathbf{A}_{\mathcal{O}}^{\mathcal{B}}$ to avoid any confusion. The calculation of this transformation matrix is discussed in §3.1.2 and implemented with Eq 3.10.

3.1.2. Attitude

To determine the attitude of an object, a model must be used to determine the rotation of an object in three dimensions. For this the visual and intuitive example of the Euler angles exist. Euler angles are the rotation of an object around three orthogonal axis, that change orientation with the rotation of the object. The three axes, denoted by x , y and z rotate with the object as depicted in Figure 3.1.

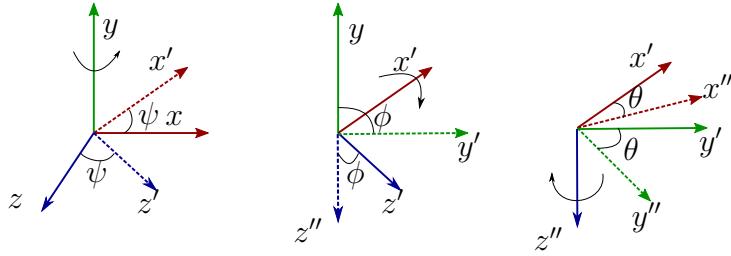


Figure 3.1: Euler angles

$\mathbf{A}_{\mathcal{O}}^{\mathcal{B}}$ can be used to calculate the attitude transformation from given Euler angle rotations. This is done by multiplying the transformation matrices representing each individual Euler angle rotation. The $\mathbf{A}_{\mathcal{O}}^{\mathcal{B}}$ can therefore be calculated as

$$\begin{aligned} \mathbf{A}_{\mathcal{O}}^{\mathcal{B}} &= \mathbf{A}_\psi \mathbf{A}_\phi \mathbf{A}_\theta \\ &= \begin{bmatrix} \cos\psi & \sin\psi & 0 \\ -\sin\psi & \cos\psi & 0 \\ 0 & 0 & 1 \end{bmatrix} \begin{bmatrix} 1 & 0 & 0 \\ 0 & \cos\phi & \sin\phi \\ 0 & -\sin\phi & \cos\phi \end{bmatrix} \begin{bmatrix} \cos\theta & 0 & -\sin\theta \\ 0 & 1 & 0 \\ \sin\theta & 0 & \cos\theta \end{bmatrix}. \end{aligned} \quad (3.7)$$

Euler angles however are not always a suitable method in determining the attitude of a satellite. This is because of singularities that can occur such as the gimbal-lock effect. Where two rotational axis, coincide to form a single rotational axis. Consequently, not all 3D rotations can be described with Euler angles, because with gimbal-lock only two effective rotations can occur instead of three [?]. Therefore the method of describing 3D rotation with quaternions is more often used and more convenient.

A quaternion, \mathbf{q} , has four components that are dependent on one another and constrained by

$$q_1^2 + q_2^2 + q_3^2 + q_4^2 = 1. \quad (3.8)$$

The attitude quaternion is also related to the Euler angles in that if the Euler rotational axis from ORC to SBC is defined as a unit vector $\mathbf{e} = [e_1 \ e_2 \ e_3]^T$ and the angle of the Euler rotation is Φ then \mathbf{q} can be expressed as

$$\mathbf{q} = \begin{bmatrix} e_1 \sin\left(\frac{\Phi}{2}\right) \\ e_2 \sin\left(\frac{\Phi}{2}\right) \\ e_3 \sin\left(\frac{\Phi}{2}\right) \\ \cos\left(\frac{\Phi}{2}\right) \end{bmatrix} \quad (3.9)$$

It is difficult to visualize a quaternion, however the most simplistic method of understanding it is shown in Figure 3.2. A quaternion is consequently a unit vector protruding from the centre point of an object as well as the angle of rotation of that object around that unit vector. As seen in Figure 3.1 the angle θ is the angle of rotation around the z -axis, for quaternions the angle of rotation is the same principle, however the axis around which the object is rotating, is the unit vector. Therefore, q_4 provides the angle of rotation while, q_{1-3} represents the unit vector, however with the condition of Eq 3.8.

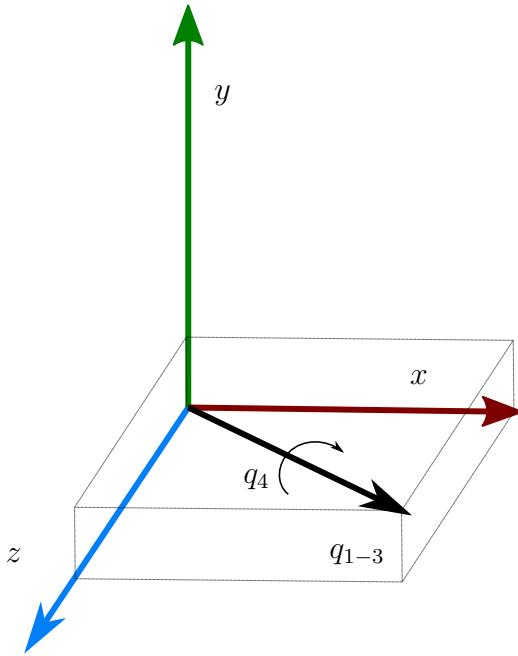


Figure 3.2: Graphical quaternion representation

\mathbf{A}_O^B can also be transformed as a function of \mathbf{q} [?] through

$$\mathbf{A} = \begin{bmatrix} q_1^2 - q_2^2 - q_3^2 + q_4^2 & 2(q_1 q_2 + q_3 q_4) & 2(q_1 q_3 - q_2 q_4) \\ 2(q_1 q_2 - q_3 q_4) & -q_1^2 + q_2^2 - q_3^2 + q_4^2 & 2(q_2 q_3 + q_1 q_4) \\ 2(q_1 q_3 + q_2 q_4) & 2(q_2 q_3 - q_1 q_4) & -q_1^2 - q_2^2 + q_3^2 + q_4^2 \end{bmatrix}. \quad (3.10)$$

The quaternion is used for attitude determination and therefore also for the attitude control. A error between the commanded quaternion, \mathbf{q}_c and the current quaternion, \mathbf{q} , is required for proportional control. This is discussed in section 3.6.1.

3.1.3. Satellite Kinematics and Dynamics

The conservation of momentum dominates the dynamics of a satellite. This consists of the torques applied to the satellite, and are mainly control torques, \mathbf{N}_c , or disturbance torques, \mathbf{N}_d , as well as the moment of inertia of the satellite, \mathbf{I} , multiplied by the inertial-referenced angular acceleration of the satellite, $\dot{\omega}_B^I$. The control torques used in this design are only reaction wheel torques, \mathbf{N}_w , and magnetorquer torques, \mathbf{N}_m . The disturbance torques are discussed in detail in section 3.4, therefore it can only be mentioned that the disturbance torques are the gravity gradient torque, \mathbf{N}_{gg} , the wheel imbalance torque, \mathbf{N}_{rw} , the gyroscopic coupling torque, \mathbf{N}_{gyro} , and the aerodynamic disturbance torque, \mathbf{N}_{aero} . The Euler dynamic equation can therefore be given as

$$\mathbf{J}\dot{\omega}_B^I = \mathbf{N}_c + \mathbf{N}_d,$$

where $\mathbf{N}_d \approx \mathbf{N}_{aero} - \mathbf{N}_{gyro} + \mathbf{N}_{gg} + \mathbf{N}_{rw}$,

and $\mathbf{N}_c = \mathbf{N}_m - \mathbf{N}_w$.

(3.11)

This is the overarching equation that will be used to determine the control torque as well as the model update of the extended Kalman Filter (EKF). The integration method used in the simulation is the 4th order Runge-Kutta method to solve the differential equations. This is demonstrated with Algorithm 3.1.

Algorithm 3.1: 4th order Runge-Kutta

```

1: Definitions: Ts - Timestamp;
2:  $h = Ts/I$ 
3: for  $n := 1$  to  $I$  do
4:    $k_1 = hf(x_n, y_n)$ 
5:    $k_2 = hf(x_n + \frac{h}{2}, y_n + \frac{k_1}{2})$ 
6:    $k_3 = hf(x_n + \frac{h}{2}, y_n + \frac{k_2}{2})$ 
7:    $k_4 = hf(x_n + h, y_n + k_3)$ 
8:    $y_{n+1} = y_n + \frac{k_1}{6} + \frac{k_2}{3} + \frac{k_3}{3} + \frac{k_4}{6}$ 
9: end for
```

Where h is the step size, which is set to $T_s/10$, and the time step, T_s , is equal to one second. $f(x_n, y_n)$ is the Euler dynamic function and Algorithm 3.1 is used to calculate ω_B^I . With this procedure the dynamics and kinematics of the satellite can be simulated after each time step.

3.2. Environment

To ensure an accurate simulation environment certain aerospace phenomena must be simulated to ensure that all anomalies can be accurately modelled as well as creating an accurate representation of the satellite orbit. Therefore, the position of the satellite with

respect to the Earth (orbit propagation) is required to determine most of the elementary principles of the satellite mission. The orbit propagation is further used to determine the moon on the Earth horizon anomaly. The sun position is required for eclipse as well as simulating the sun reflecting from the solar panels unto the sun sensor. The magnetic field is required to simulate \mathbf{N}_m of the magnetorquer as well as the solar panel dipole anomaly and disturbance torque. **TODO: Consider a section on CubeSat design.**

3.2.1. Orbit Propagation

The satellite position, \mathbf{r}_{sat} and velocity \mathbf{v}_{sat} at a given time step is required to determine the multiple different variables required for the simulation environment. Therefore the refined version and fourth generation of the simplified general perturbations (SGP) model, namely SGP4, is used as orbit propagator of the satellite after each time step [?].

To determine \mathbf{r}_{sat_k} and \mathbf{v}_{sat_k} at time step, k , the two-line element, (TLE), set of the satellite is required. The TLE set is an encoding of the specified satellite orbit, that requires parameters such as the semimajor axis, a , right ascension of the ascending node (RAAN), Ω , argument of perigee (AP), ω , inclination, i , eccentricity, e , and the time at the beginning of the orbit as a Julian date, J_t . With these parameters and the elapsed time since J_t , both \mathbf{r}_{sat_k} and \mathbf{v}_{sat_k} can be determined from the World Geodetic System 72 constants that is implemented through the SGP4 model. An example of the satellite orbit propagated by the SGP4 model is illustrated in Figure 3.3.

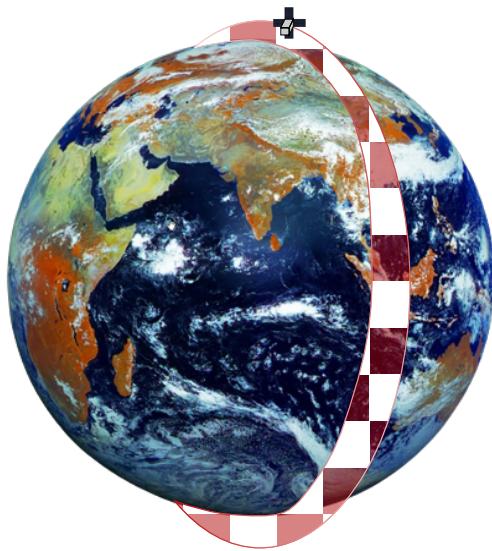


Figure 3.3: SGP4 orbit propagation

The SGP4 outputs the \mathbf{r}_{sat_k} and \mathbf{v}_{sat_k} in the EIC reference frame. The SGP4 is implemented with the SGP4 python package [?]. Therefore, with \mathbf{r}_{sat_k} and \mathbf{v}_{sat_k} known, $\mathbf{A}_{\mathcal{E}}^O$ can now be calculated according to Eq 3.6.

3.2.2. Sun

For the mission to be successful it is critical to determine the position of the sun relative to the satellite. This is because the satellite must determine whether it is in the eclipse to determine the control operation. Therefore, the model from [?] is implemented to determine the position of the sun in the EIC frame.

From this model the vector from the centre of the Earth to the centre of the sun, \mathbf{r}_{sun} , is provided in the EIC frame. This model requires various calculations as given in Eq 3.13. For this calculation the difference between the current Julian date, J_t , and the J_{2000} epoch is required. Where $J_{2000} = 2\,451\,545$ and the difference is thereafter converted to the amount of Julian centuries (365.25 days). The time difference in Julian centuries, T_{JC} can therefore be calculated as

$$T_{JC} = \frac{J_t - 2\,452\,545}{36\,525}. \quad (3.12)$$

With T_{JC} known, \mathbf{r}_{sun} can then be calculated with

$$\mathbf{r}_{sun} = r_{\oplus} \begin{bmatrix} \cos(\lambda_e) \\ \cos(\epsilon)\sin(\lambda_e) \\ \sin(\epsilon)\sin(\lambda_e) \end{bmatrix},$$

where $r_{\oplus} = 1.000\,140\,612 - 0.016\,708\,617 \cos(M_{\oplus}) - 0.001\,395\,89 \cos(2M_{\oplus})$,

$$M_{\oplus} = 357.527\,723\,300^{\circ} + 35\,999.050\,340 T_{JC}, \quad (3.13)$$

$$\lambda_e = \lambda_{M_{\oplus}} + 1.914\,666\,471 \sin(M_{\oplus}) + 0.019\,994\,643 \sin(2M_{\oplus}),$$

$$\lambda_{M_{\oplus}} = 280.460\,618\,400^{\circ} + 36\,000.770\,053\,610 T_{JC},$$

$$\epsilon = 23.439\,291^{\circ} - 0.013\,004\,200 T_{JC}$$

$$\text{and } T_{JC} = \frac{J_t - 2\,451\,545}{36\,525}.$$

The definitions of the parameters used in the calculation and the description thereof is tabulated in Table 3.1. After determining the sun position, it is crucial to calculate whether the satellite is in the eclipse or not. This can be done with basic geometry after calculating the position of the sun relative to the satellite through

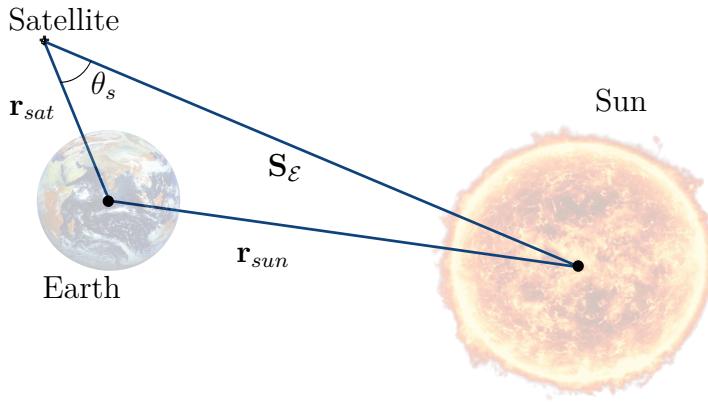
$$\mathbf{S}_{\mathcal{E}} = \mathbf{r}_{sun} - \mathbf{r}_{sat}. \quad (3.14)$$

The assumption is made that whenever the satellite is not able to view the centre of the sun it is in the eclipse. This a valid assumption given the very small angle required to change the satellite from a partial eclipse to a full eclipse, due to the comparative distances of the sun to satellite and satellite to Earth. Therefore, the eclipse is defined as the period

Table 3.1: Description and definition of Earth orbit parameters

Symbol	Definition	Description
r_{\oplus}	Sun position magnitude	The absolute distance of the Earth to the sun
λ_e	Ecliptic longitude	The angle between the primary direction 0° of the plane in which the Earth is orbiting and the current Earth position.
M_{\oplus}	Mean anomaly	The fraction of the orbit's period after the Earth has passed the furthest position from the sun
ϵ	Obliquity	The inclination of the plane of orbit to the celestial equator
$\lambda_{M_{\oplus}}$	Sun's mean longitude	The average angle subtended at the Earth between the vernal equinox and the sun. [?]

during which θ_s is smaller than θ_E . Where $\theta_E = \sin(\frac{R_E}{\|r_{sat}\|})$ and $\theta_s = \mathbf{r}_{sat} \cdot \mathbf{S}_{\mathcal{E}}$ as shown in Figure 3.4. R_E is the radius of the Earth.

**Figure 3.4:** Geometry for satellite eclipse

3.2.3. Geomagnetic field

The Earth generates a magnetic field through electric currents due to motion within the molten core of the Earth, which is commonly referred to as the geomagnetic field. The magnetorquers interact with the geomagnetic field for momentum dumping and the magnetometers measure the geomagnetic field for attitude estimation. Therefore, the modelling of the geomagnetic field is required for an accurate simulation environment.

The geomagnetic field is modelled with the time-varying International Geomagnetic Reference Field (IGRF) model released by the International Association of Geomagnetism and Aeronomy (IAGA). This model is used for increased ADCS accuracy and the 13th generation of the model is implemented [?]. The scalar potential function,

$$V(r_s, \theta, \phi, t) = R_E \sum_{n=1}^N \left(\frac{R_E}{r_s} \right)^{n+1} \sum_{m=0}^n (g_n^m(t) \cos(m\phi) + h_n^m(t) \sin(m\phi)) P_n^m(\cos\theta), \quad (3.15)$$

is used to calculate the geomagnetic field, \mathbf{B} with

$$\mathbf{B} = -\nabla V. \quad (3.16)$$

Therefore, the geomagnetic field is the gradient of the scalar potential function given in Eq 3.15. Where R_E is the mean Earth radius of 6371.2km, r_s is the radial distance from the centre of the Earth, θ is the latitude and ϕ is the longitude. $g_n^m(t)$ and $h_n^m(t)$ is known as the Gauss coefficients that slowly change with time and consequently the IGRF-13 provide values for these coefficients at 5-year epoch intervals. The $P_n^m(\cos\theta)$ is the Legendre functions of the degree n and m [?].

The magnitude of the geomagnetic field is visually demonstrated in Figure 3.5. The measuring of the magnetic field with the magnetometer is further discussed in section ??.

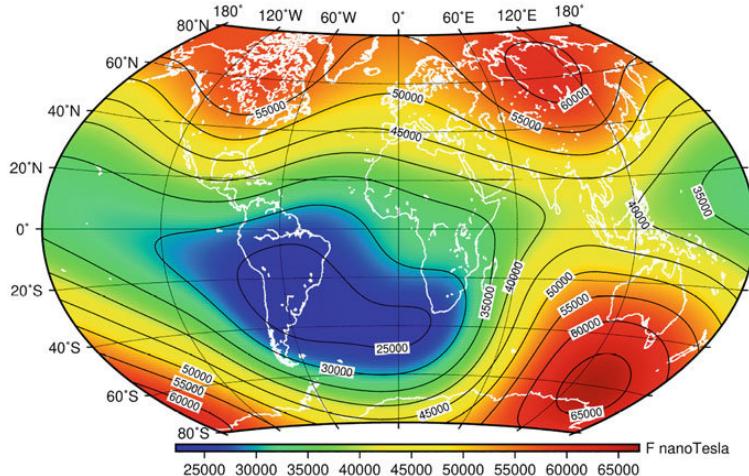


Figure 3.5: The magnitude of geomagnetic field according to the 13th generation of the IGRF model [?].

3.3. Sensor models

The positioning of the sensors on the satellite is necessary to meet mission requirements. The exact position of the sensors also impact the modelling of the anomalies on the sensors. Therefore, each sensor's position on the satellite is provided as well as the measured vector of each sensor. It is further assumed that each sensor has a zero-mean Gaussian noise and consequently, the low frequency noise such as drift is negligible. The sensor measurement

in the SBC frame, \mathbf{v}_B , can be calculated as

$$\mathbf{v}_B = A_O^\varepsilon \mathbf{v}_I + \mathbf{m}_v, \quad (3.17)$$

where \mathbf{m}_v is the measurement noise of the current sensor and \mathbf{v}_I is the reference ORC vector. The measured unit vectors of the sun as an example of the sensor measurements is shown in Figure 3.6 where the grey background sections of the graphs are the eclipse periods, while the sections with the white background is the sunlit phase of the orbit.

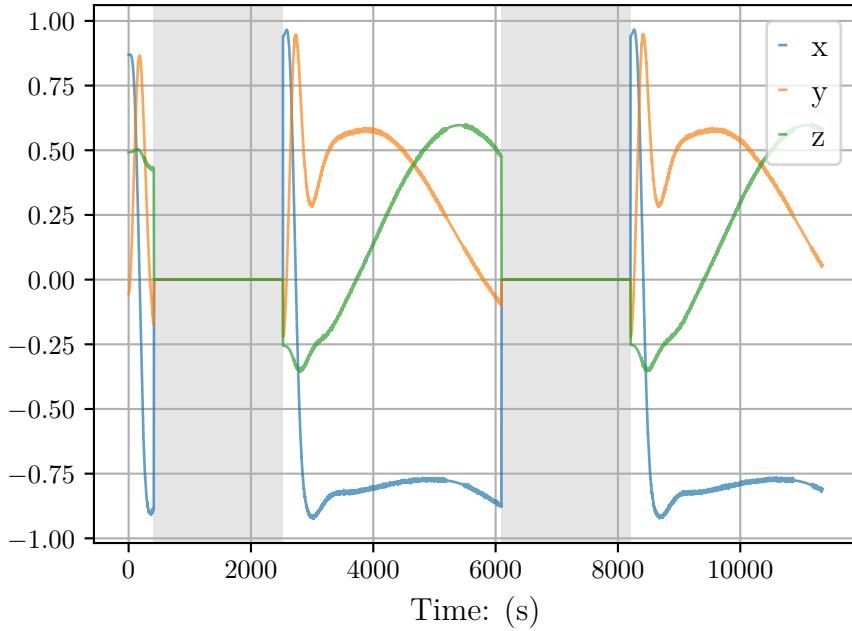


Figure 3.6: Sun vector in SBC

3.4. Disturbance models

During orbit a satellite is exposed to various disturbance torques. It is these torques that cause the modelled attitude to differ from the actual attitude. Therefore, these torques are modelled and are assumed to influence the attitude of the satellite continuously. Other disturbances that occur only with anomalies are discussed in Chapter 5. Some disturbances are excluded from the simulation environment and only the major sources of disturbance torques are modelled and discussed.

The first disturbance torque is that of the gyroscopic coupling which can be calculated with

$$\mathbf{N}_{gyro} = \dot{\boldsymbol{\omega}}_B^I \times (\mathbf{I}\dot{\boldsymbol{\omega}}_B^I + \mathbf{h}_w), \quad (3.18)$$

where \mathbf{h}_w is the angular momentum of the reaction wheels. The other disturbance torques are discussed in more detail below.

3.4.1. Gravity Gradient

The gravity gradient is caused by both the centrifugal force on the satellite due to the orbit around the Earth as well as the gravitational force. The part of the satellite nearest to the Earth will experience the largest gravitational force and the smallest centrifugal force of the satellite. While the part of the satellite furthest from the Earth will experience the smallest gravitational force and the largest centrifugal force. According to [?] the gravity gradient disturbance torque, \mathbf{N}_{gg} can be calculated as

$$\mathbf{N}_{gg} = 3\omega_o^2(\mathbf{z}_B \times \mathbf{I}\mathbf{z}_B). \quad (3.19)$$

The orbit nadir vector is calculated as,

$$\mathbf{z}_B = \mathbf{A} \begin{bmatrix} 0 & 0 & 1 \end{bmatrix}^T. \quad (3.20)$$

Due to the design of the satellite and the positions of the solar panels, the equation for \mathbf{N}_{gg} can not be simplified. The gravity gradient torque is the only torque that can be accurately modelled on-board the satellite, and therefore is also included in the model update of the EKF. \mathbf{N}_{gg} in SBC is shown in Figure 3.7.

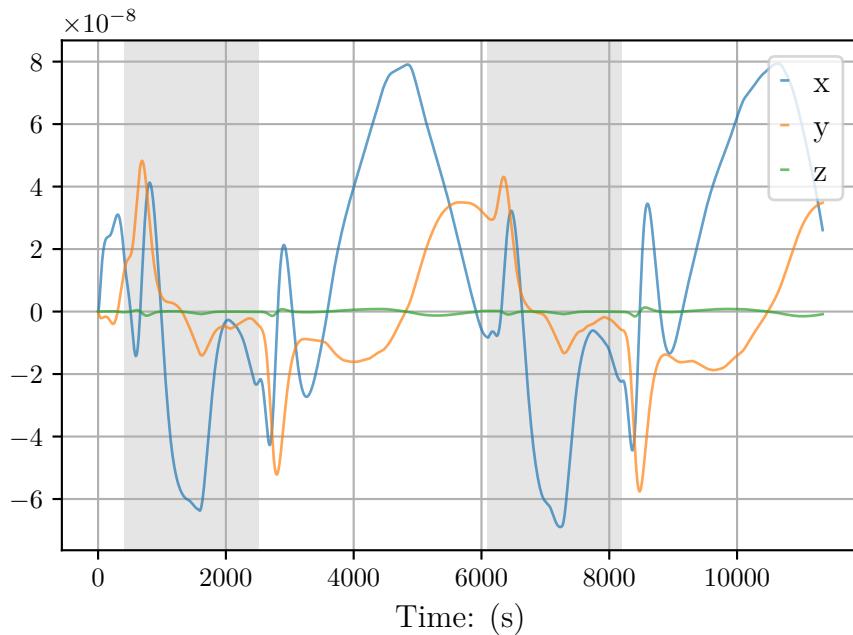


Figure 3.7: Gravity Gradient Torques in SBC

3.4.2. Aerodynamic Disturbance

The aerodynamic disturbance torques are cause by air density in the atmosphere creating a force on each individual segment of the satellite [?]. This is significant due to the low

Earth orbit (LEO) of the satellite, where the atmospheric density is higher. Therefore, the aerodynamic disturbance torque, \mathbf{N}_{aero} is a summation of all the torques created by the air force on each segment area, A_i . \mathbf{N}_{aero} can therefore be calculated as

$$\mathbf{N}_{aero} = \sum_{i=1}^n \left(\rho \|\mathbf{v}_B\|^2 A_i H \{\cos(\alpha_i)\} \cos(\alpha_i) (\sigma_t (\mathbf{r}_{pi} \times \bar{\mathbf{v}}_B) + [\sigma_n S + (2 - \sigma_n - \sigma_t) \cos(\alpha_i)] (\mathbf{r}_{pi} \times \bar{\mathbf{n}}_i)) \right), \quad (3.21)$$

where n is the number of segments of the satellite. The factors that influence the aerodynamic disturbance torque is the atmospheric velocity in SBC, \mathbf{v}_B , the atmospheric density, ρ , each segment's surface area, A_i and the offset vector between the segment's centre of mass (CoM) and the centre of pressure (CoP), \mathbf{r}_p . $H\{x\}$ is the Heaviside function which is equal to 0 when x is smaller than 0 and equal 1 otherwise. α_i is the incidence angle of \mathbf{v}_B on segment i , while σ_t is the tangential accommodation coefficient and σ_n is the normal accommodation coefficient. S is the ratio of molecular exit velocity to \mathbf{v}_B and $\bar{\mathbf{n}}_i$ is the unit inward normal vector of segment i [?].

The atmospheric density is a function of the distance from the Earth surface. The density model provided by [?] is given as

$$\rho = \rho_o e^{-\frac{h(t)-h_o}{H}}, \quad (3.22)$$

where ρ_o is the reference density at the reference altitude, h_o , and $h(t)$ is the satellite's altitude as a function of time and H is the scale height. According to [?] the atmospheric density is $\frac{1}{2}\rho$ during during eclipse. Furthermore \mathbf{v}_B is calculated as

$$\mathbf{v}_B = \mathbf{A}_O^\mathcal{B} \mathbf{A}_E^\mathcal{O} \mathbf{v}_E$$

where $\mathbf{v}_E = \begin{bmatrix} 0 \\ 0 \\ \omega_E \end{bmatrix} \times \mathbf{r}_{sat} - \mathbf{v}_{sat}$

$$(3.23)$$

σ_n and σ_t are both assumed to be equal to 0.8, while S is 0.8 [?]. From these equations the aerodynamic disturbance can be calculated and is shown in Figure 3.8.

3.4.3. Wheel Imbalance

The reaction wheel imbalance is considered to be the most significant disturbance on the reaction wheel and consequently, only this disturbance torque is modelled for this simulation [?]. Although reaction wheels are manufactured with low tolerances the reaction wheel will have a slight imbalance, since the mass of the reaction wheel will not be uniform and evenly displaced.

The static imbalance of the reaction wheel is the caused by CoM offset from the

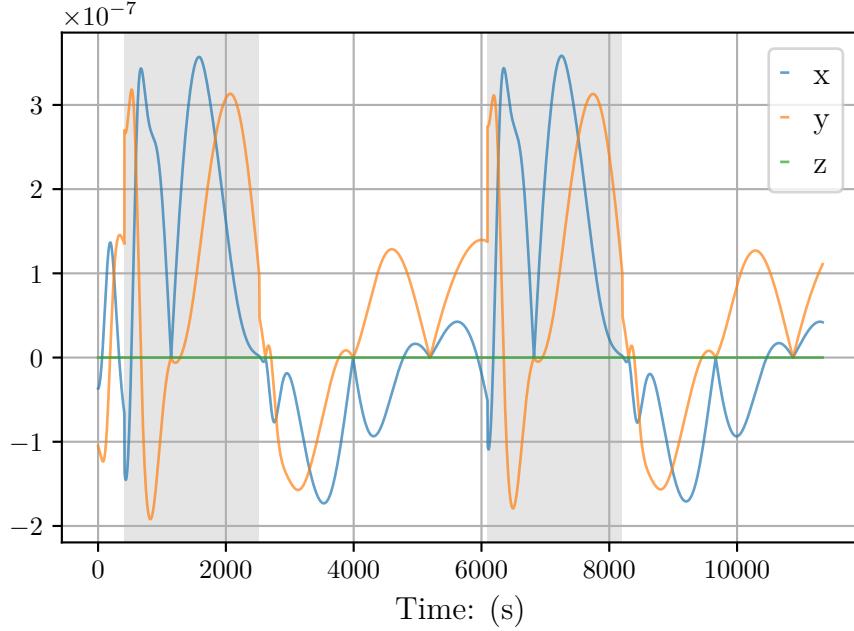


Figure 3.8: Aerodynamic Torques in SBC

rotational axis. Therefore, to model the static imbalance of the reaction wheels it is assumed that the unevenly distributed mass of the reaction wheel can be simplified to a point mass, m , a distance, r from the rotational axis as shown in Figure 3.9. The static imbalance, U_s is equal to mr and this value is given by the reaction wheel manufacturers. To determine the resulting torque from the wheel imbalance the torque generated by each

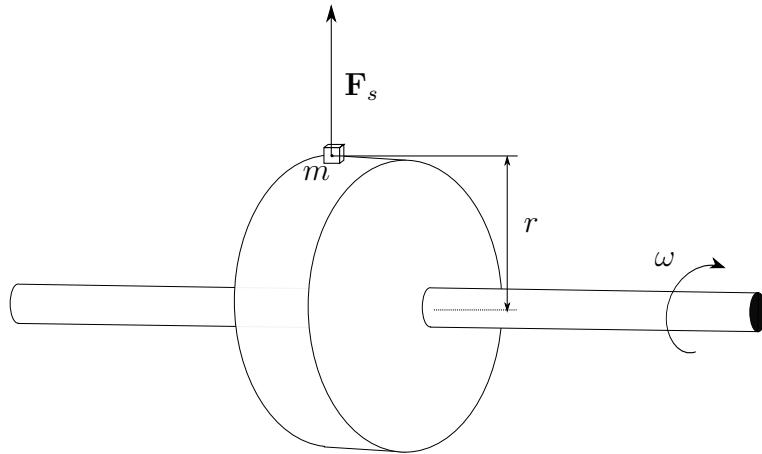


Figure 3.9: Static Imbalance

wheel is individually calculated. Consequently, for the reaction wheel in direction x_B , RW_{x_B} , the force, F_{sx_B} , generated by U_s is dependent on the the angular rate, ω , of the

reaction wheel as well as the current position of m . Therefore, F_{sx_B} can be expressed as

$$\mathbf{F}_{sx_B} = U_s \omega^2 \begin{bmatrix} 0 \\ \sin(\omega t + \phi_s) \\ \cos(\omega t + \phi_s) \end{bmatrix}, \quad (3.24)$$

where the current angle of m is defined as $\omega t + \phi_s$, with ϕ_s as an arbitrary phase and for the sake of simplification is set to 0. With F_{sx_B} exerted on RW_{x_B} known, the torque on the satellite can be calculated with the known position vector, w_{x_B} , of F_{sx_B} to the satellite CoM. Therefore, N_{sx_B} can be calculated as

$$\mathbf{N}_{sx_B} = w_{x_B} \times F_{sx_B}. \quad (3.25)$$

This is calculated for each reaction wheel to determine the resulting static imbalance disturbance torque on the satellite. Another aspect of the reaction wheel imbalance is also modelled, namely the dynamic imbalance. The dynamic imbalance is caused by the principal inertia of the reaction wheel being misaligned with the rotational axis. This can be simplified to two equal point masses, m , with an axial displacement, d , and distance, r , from the rotational axis. These two masses are 180° apart with respect to the rotational axis and consequently create two equal forces in opposite directions. The dynamic imbalance is graphically represented in Figure 3.10. The dynamic wheel imbalance

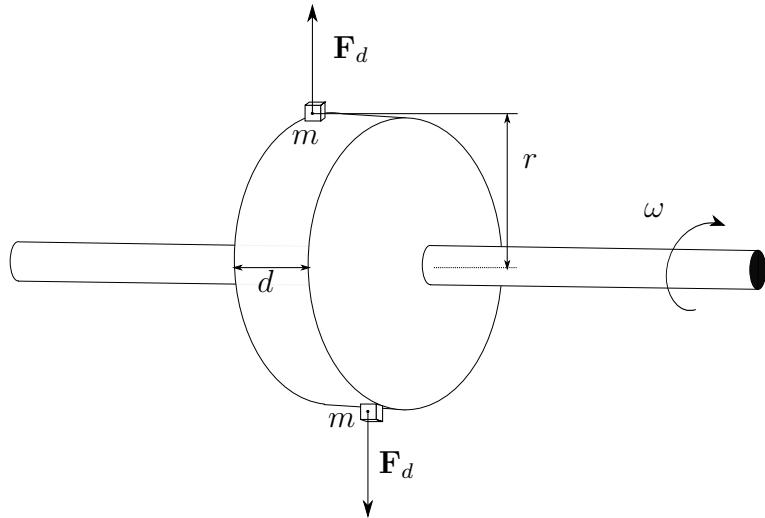


Figure 3.10: Dynamic Imbalance

torque, \mathbf{N}_{dx_B} , for RW_{x_B} with the U_d known can be calculated as

$$\mathbf{N}_{dx_B} = U_d \omega^2 \begin{bmatrix} 0 \\ \sin(\omega t + \phi_d) \\ \cos(\omega t + \phi_d) \end{bmatrix}. \quad (3.26)$$

where $U_d = mrd$ is the dynamic imbalance. Both U_d and U_s are provided by the manufacturer and based on the reaction wheel, RW-0.06 from Sinclair Interplanetary, $U_d = 2.08e^{-9}$ and $U_s = 2.08e^{-7}$. The net wheel imbalance torque from both the static and dynamic wheel imbalance is provided in Figure 3.11.

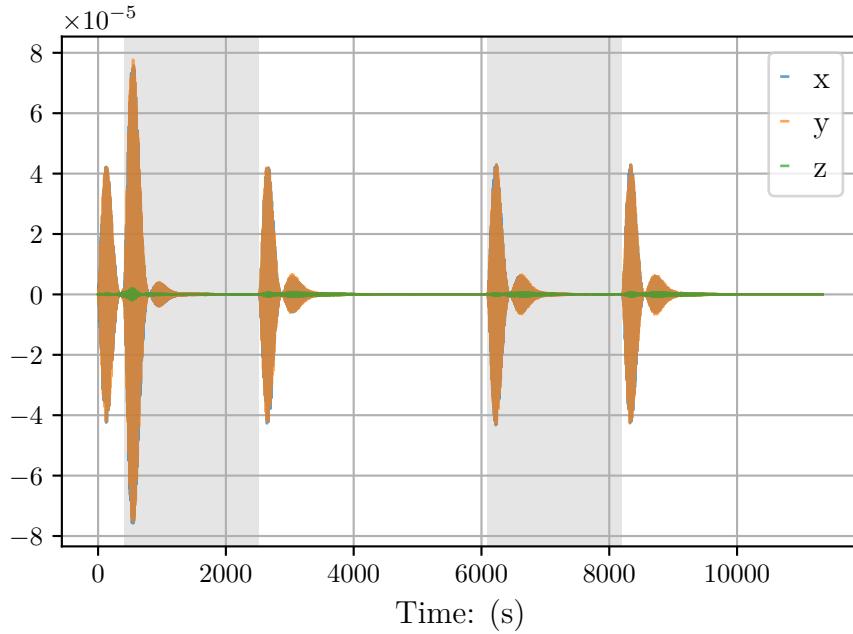


Figure 3.11: Wheel disturbance torques in SBC

3.5. Attitude Determination

This section discusses the estimation algorithm for attitude estimation of the satellite. This is done with the EKF, which utilizes the sensor measurements as well as modelled vectors according to underlying physics to estimate the current attitude. The EKF is highly sensitive to sensor anomalies and actuator failures and this section provides the background knowledge to the fundamentals of the EKF.

3.5.1. Extended Kalman Filter

The implementation of the estimated kalman filter *EKF* is for estimation of the current satellite attitude with sensor fusion of the magnetometer, star tracker, sun sensor and nadir sensor to accurately estimate the attitude and rotation rate of the satellite. The EKF will be used due to the non-linear nature of the system. The EKF consists of two fundamental parts, the model update and the measurement update. The general form for a system model can be expressed as

$$\dot{x}_t = \mathbf{f}(x_t) + s_t \quad (3.27)$$

where $\mathbf{f}(x_t)$ is a non-linear function of x_t . The state vector, x , for the full 7-state EKF consists of the quaternion vector, q and the inertial-referenced angular velocity, ω_B^T .

$$x = [q, \omega_B^T]^T \quad (3.28)$$

The estimated state vector x will be denoted as \hat{x} and the estimated vector before and after the measurement update will be indicated with a superscript ' $-$ ' and ' $+$ ' respectively. To calculate the model update the dynamics and kinematics of the system model is used to calculate both ω_B^T and q . The integration method used in the simulation is the 4th order Runge-Kutta method to solve the differential equations. The integration method is shown in Algorithm 3.2 where $(\hat{\omega}_B^I)_k^-$ is calculated for the first step of the model update.

Algorithm 3.2: Runge-Kutta 4th order Algorithm at time k

```

1: Satellite Body Inertia  $\mathbf{J} = \begin{bmatrix} 0.4 & 0 & 0 \\ 0 & 0.45 & 0 \\ 0 & 0 & 0.3 \end{bmatrix}$ 
2: Timestep ( $T_s$ ) = 1s
3: Number of iterations ( $I$ ) = 10
4: Step size  $h = \frac{T_s}{h}$ 
5: Disturbance torques  $N_d = N_{gg} - N_{gyro}$ 
6: Control torques  $N_c = N_m - N_w$ 
7:  $\mathbf{N} = N_c + N_d$ 
8: for  $n \in I$  do
9:    $k_1 = h(\mathbf{J}^{-1}\mathbf{N})$ 
10:   $k_2 = h(\mathbf{J}^{-1}\mathbf{N} + \frac{k_1}{2})$ 
11:   $k_3 = h(\mathbf{J}^{-1}\mathbf{N} + \frac{k_2}{2})$ 
12:   $k_4 = h(\mathbf{J}^{-1}\mathbf{N} + k_3)$ 
13:   $\omega_{n+1} = \omega_n + \frac{k_1}{6} + \frac{k_2}{3} + \frac{k_3}{3} + \frac{k_4}{6}$ 
14: end for
15:  $(\hat{\omega}_B^I)_k^- = \omega_{n+1}$ 
16: return  $(\hat{\omega}_B^I)_k^-$ 

```

where N_{gg} , gravity gradient disturbance torque, N_{aero} , aerodynamic disturbance torque, N_{rw} , reaction wheel disturbance torque, N_{gyro} , gyroscopic torque, N_m , magnetic control torque and N_w , reaction wheel control torque, are modelled according to [?].

The calculation for the \hat{q}_k^- is done with Eq 3.29 [?]

$$\hat{q}_k^- = \left[\cos(k_q) \mathbf{I}_{4 \times 4} + \frac{1}{\|(\hat{\omega}_B^O)_k^-\|} \sin(k_q) \boldsymbol{\Omega}_k^- \right] \hat{q}_{k-1}^+$$

where $k_q = \frac{T_s}{2} \|(\hat{\omega}_B^O)_k^-\|$

$$\begin{aligned} (\hat{\omega}_B^O)_k^- &= (\hat{\omega}_B^I)_k^- - \hat{\mathbf{A}} \begin{bmatrix} 0 & -(\omega_k) & 0 \end{bmatrix}^T \\ &= [\hat{\omega}_{ox} \quad \hat{\omega}_{oy} \quad \hat{\omega}_{oz}]^T \end{aligned} \quad (3.29)$$

$$\|(\hat{\omega}_B^O)_k^-\| = \sqrt{\hat{\omega}_{ox}^2 + \hat{\omega}_{oy}^2 + \hat{\omega}_{oz}^2}$$

and $\boldsymbol{\Omega}_k^- = \begin{bmatrix} 0 & \hat{\omega}_{oz} & -\hat{\omega}_{oy} & \hat{\omega}_{ox} \\ -\hat{\omega}_{oz} & 0 & \hat{\omega}_{ox} & \hat{\omega}_{oy} \\ \hat{\omega}_{oy} & -\hat{\omega}_{ox} & 0 & \hat{\omega}_{oz} \\ -\hat{\omega}_{ox} & -\hat{\omega}_{oy} & -\hat{\omega}_{oz} & 0 \end{bmatrix}$

The estimated state vector, \hat{x}_k^- can now be expressed as

$$\hat{x}_k^- = [(\hat{\omega}_B^I)_k^- \quad \hat{q}_k^-] \quad (3.30)$$

Q_k is the covariance matrix representing the discrete system noise and is assumed to be zero-mean and Guassian. Φ_k is the discrete system perturbation model. H_k is the discrete measurement perturbation Jacobian Matrix. \mathbf{R}_k is the measurement noise covariance matrix. The state covariance matrix P_k can be propagated with Eq 3.31.

$$\mathbf{P}_k^- = \Phi_k \mathbf{P}_{k-1}^+ \Phi_k^T \quad (3.31)$$

The \mathbf{e}_k calculated with Eq 3.32

$$\mathbf{e}_k = v_{meas,k} - \hat{\mathbf{A}}_k^- v_{model,k} \quad (3.32)$$

where $v_{meas,k}$ is the measured vector in \mathcal{B} and $v_{model,k}$ is the modelled \mathcal{O} vector. The gain matrix \mathbf{K}_k is used to determine the influence of \mathbf{e}_k on updated state vector, \hat{x}_k^+ . \mathbf{K}_k can be calculated as

$$\mathbf{K}_k = \mathbf{P}_k^- (\mathbf{H}_k^-)^T \left[\mathbf{H}_k^- \mathbf{P}_k^- (\mathbf{H}_k^-)^T + \mathbf{R}_k \right]^{-1} \quad (3.33)$$

after which the updated state vector can be calculated with Eq 3.34

$$\hat{x}_k^+ = \hat{x}_k^- + \mathbf{K}_k \mathbf{e}_k \quad (3.34)$$

The state covariance matrix can then be updated as

$$\mathbf{P}_k^+ = [\mathbf{I}_{7 \times 7} - \mathbf{K}_k \mathbf{H}_k^+] \mathbf{P}_k [\mathbf{I}_{7 \times 7} - \mathbf{K}_k \mathbf{H}_k^+] + \mathbf{K}_k \mathbf{R}_k \mathbf{K}_k^T \quad (3.35)$$

During the measurement update of \hat{x}_k^- with the error of $v_{meas,k}$ and $v_{model,k}$, the \hat{x}_k^+ is largely affected by anomalous behaviour in the sensor measurements. The sensitivity of the Kalman filter to this behaviour as well as related work will be discussed in section ??.

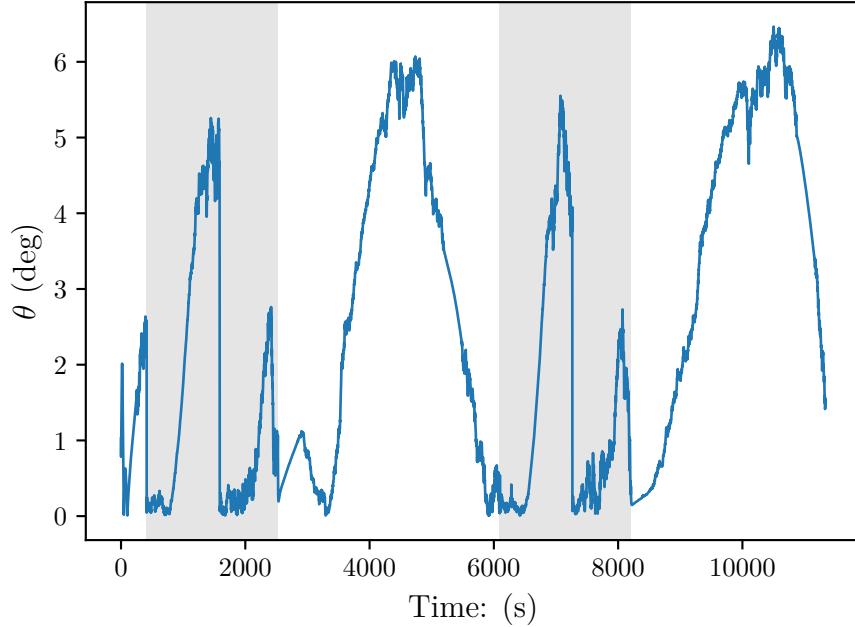


Figure 3.12: Estimation Metric

3.6. Attitude Control

To ensure that the satellite is able to satisfy the mission requirements, control of the satellite attitude is required. Therefore, the satellite's payload must be in the direction of the Earth during eclipse and the solar panels should be pointing in the direction of the sun during the sunlit phase. For this control a quaternion feedback controller of the reaction wheels is implemented and a momentum dumping with the magnetorquers is implemented to ensure that the wheel disturbance remains within reasonable boundaries.

3.6.1. Quaternion Feedback Controller

To ensure that the satellite is in the desired orientation the reaction wheels are used. To ensure stable control in all three axes [?] the quaternion feedback reaction wheel controller is implemented. The controller is designed from the estimated orientation and angular velocity, while ignoring disturbance torques. Eq ?? without the disturbance torques

becomes

$$\mathbf{J}\dot{\boldsymbol{\omega}}_B^I = -\mathbf{N}_w - \mathbf{N}_{gyro} \quad (3.36)$$

To calculate the required torque, $\mathbf{N}_{w,req}$, the definition according to [?] for all cases at time step, k , is given as

$$\mathbf{N}_w(k) = K_{PI}\mathbf{J}\mathbf{q}_{err}(k) + K_{DI}\mathbf{I}\hat{\boldsymbol{\omega}}_B^O - \hat{\boldsymbol{\omega}}_B^I(k) \times [\mathbf{J}\hat{\boldsymbol{\omega}}_B^I(k) + \mathbf{h}_w(k)] \quad (3.37)$$

where $\mathbf{h}_w(k)$ is the measured angular momentum of the wheels, the control gains are defined as

$$\begin{aligned} K_{PI} &= 2\omega_n^2 \\ K_{DI} &= 2\zeta\omega_n \end{aligned} \quad (3.38)$$

the quaternion error is calculated with the quaternion difference operator, Θ , as

$$\begin{aligned} \mathbf{q}_{err}(k) &= \mathbf{q}_c(k)\Theta\hat{\mathbf{q}}(k) \\ \begin{bmatrix} q_{1e} \\ q_{2e} \\ q_{3e} \\ q_{4e} \end{bmatrix} &= \begin{bmatrix} q_{4c} & q_{3c} & -q_{4c} & -q_{4c} \\ -q_{3c} & q_{4c} & q_{1c} & -q_{2c} \\ q_{2c} & -q_{1c} & q_{4c} & -q_{3c} \\ q_{1c} & q_{2c} & q_{3c} & q_{4c} \end{bmatrix} \begin{bmatrix} \hat{q}_1 \\ \hat{q}_2 \\ \hat{q}_3 \\ \hat{q}_4 \end{bmatrix} \end{aligned} \quad (3.39)$$

where $\hat{\mathbf{q}}$ is the current estimated quaternion and q_c is the command quaternion which is $\{0 \ 0 \ 0 \ 1\}^T$ during eclipse and during the sun following phase, the attitude command according to [?] can be calculated as

$$\mathbf{q}_c = \begin{bmatrix} \mathbf{u}_c \sin(\frac{\delta}{2}) \\ \cos(\frac{\delta}{2}) \end{bmatrix}, \quad (3.40)$$

where

$$\mathbf{u}_c = \frac{\mathbf{u}_{sp}^B \times \mathbf{s}_o}{\|\mathbf{u}_{sp}^B \times \mathbf{s}_o\|}. \quad (3.41)$$

\mathbf{s}_o is the measured unit sun vector in ORC, and the main solar panel's position is denoted as a unit vector, \mathbf{u}_{sp}^B . The angle between \mathbf{u}_{sp}^B and \mathbf{s}_o , δ , can be calculated with the vector dot-product. This can then be used as the reference for the control. The reference $\boldsymbol{\omega}_b^I$ is always $[0, 0, 0]$.

3.6.2. Momentum Dumping

Momentum dumping is crucial to ensure that the wheel disturbance does not cause the system to become unstable. Momentum dumping is implemented during eclipse after the satellite is in a stable nadir-pointing attitude. The momentum dumping is implemented with magnetic torquers based on a Cross-Product controller.

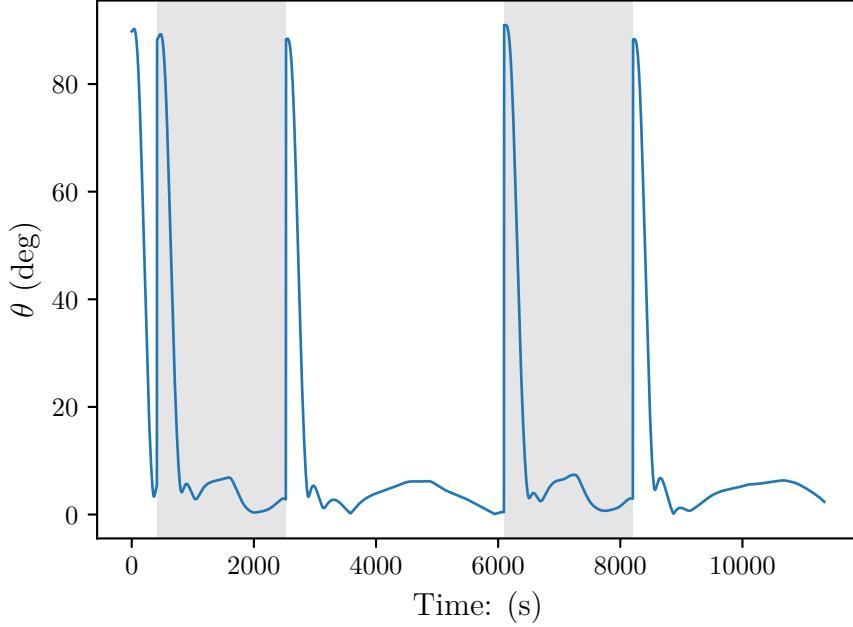


Figure 3.13: Pointing Metric

The magnetic dipole moment \mathbf{M} is calculated as

$$\mathbf{M} = \frac{\mathbf{e} \times \mathbf{B}}{\|\mathbf{B}_b\|^2} \quad (3.42)$$

where \mathbf{B}_b is the geomagnetic field and the error vector, \mathbf{e} can be calculated as

$$\mathbf{e} = -K_w(\mathbf{h}_w - \mathbf{h}_{w,\text{ref}}) \quad (3.43)$$

where K_w is a positive gain. This momentum dumping is implemented 200s after sun-following phase is implemented, to ensure stable control and reduce the momentum in the reaction wheels. The magnetorquers torques are shown in Figure 3.14 and it is evident that when the satellite control changes from eclipse to sunlit and from sunlit to eclipse the magnetorquers torques increase to compensate for the increase in reaction wheel torques and to minimise the dynamic reaction wheel disturbance.

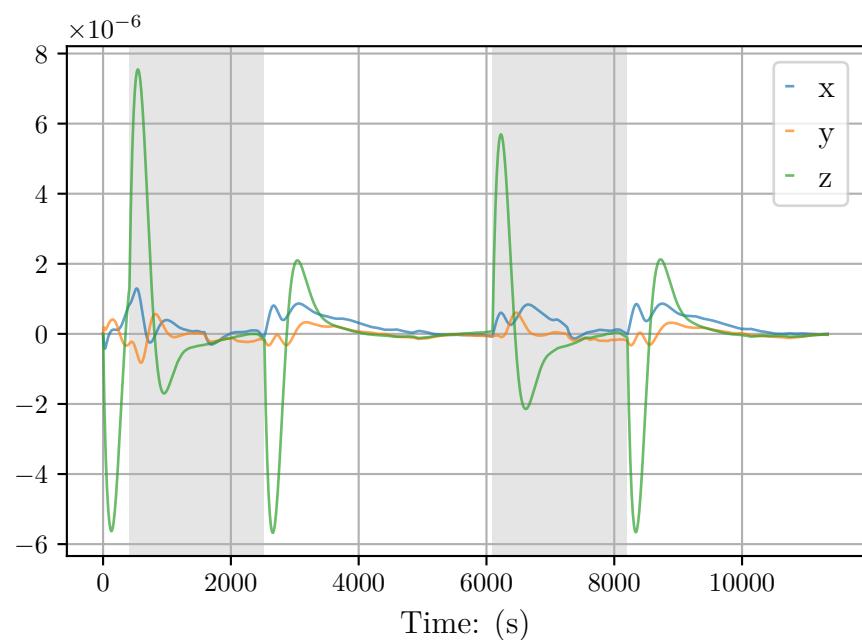


Figure 3.14: Magnetic Control Torques

Chapter 4

Implementation of Methods on Actual Satellite

4.1. Fault Detection, Isolation and Recovery

4.2. Feature Extraction

4.3. Detection

4.4. Isolation

4.5. Recovery

4.6. Summary

Chapter 5

Anomalies

To ensure that the prediction and classification of anomalies are not based on generalised sensor failures a few anomalies are simulated. These anomalies are either chosen to show the significant effect of these anomalies on the ADCS or are modelled based on research that label the anomaly as a possible influence on the ADCS. There is a anomaly for each sensor, that will create inaccuracies for that specific sensor measurement. An anomaly for the reaction wheels is also implemented to show the resulting estimation failure based on an inaccurate model update, since the control torque and the torque implemented is not the same. This anomaly will also be predicted based on the sensor readings, since the effect will be evident on all the sensors.

5.1. Reflection of Solar Panels on Sun Sensor

Sun reflection from solar panels unto sun sensors is a very probable anomaly, if the sun sensor is not placed in a position where the reflection will not influence the sensor. Reflection on the sun sensor is modelled to determine whether the reflection has a significant effect on the ADCS. The reflection anomaly is modelled for the specific shape and design of the CubeSat as shown in Figure 5.1. The CubeSat will also be the same design for all of the modelled anomalies.

The reflection anomaly is modelled for the specific shape and design of the satellite as shown in Figure 5.1. The modelling takes place within the SBC frame.

The assumption is that the solar panel can be modeled as a geometric plane. Therefore, light from the solar panel will reflect similarly to a perfectly smooth mirror. This model also assumes, that if the sun sensor detects any reflection from the solar panel, the measured sun vector will default to the reflection ray instead of the direct sun vector. In practice, this is a function of the exact detection algorithm within the sensor, and some reflections might be ignored. This assumption will produce the worst-case behaviour. The intensity of the light vector is also disregarded.

The solar panel $ABCD$ -plane can be represented in the SBC by a point and normal

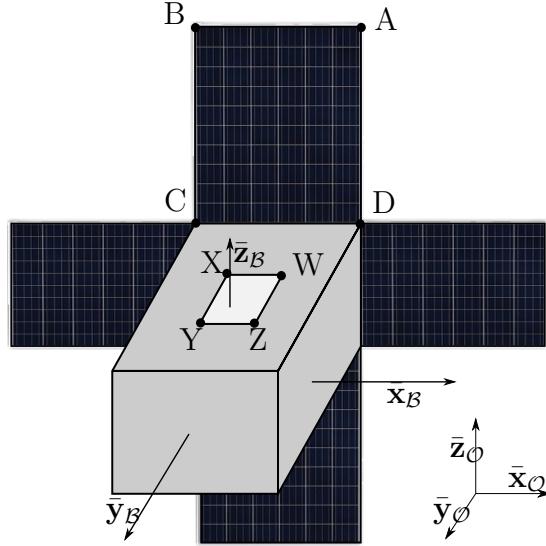


Figure 5.1: The modelled satellite, with solar panels and sun sensor within the SBC frame.

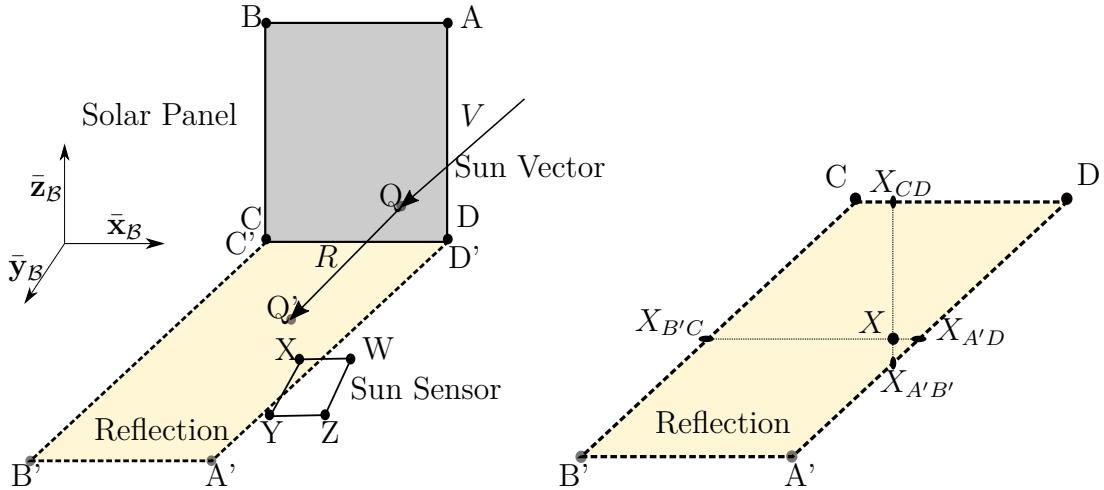


Figure 5.2: Definition of solar reflection point from $ABCD$ -plane.

vector defined as

$$\mathbf{p}_{ABCD} = [p_x, p_y, p_z]^\top, \text{ and} \\ \bar{\mathbf{n}}_{ABCD} = [n_p x, n_p y, n_p z]^\top, \quad (5.1)$$

respectively. Similarly, the sun sensor $WXYZ$ -plane is represented by the point, \mathbf{p}_{WXYZ} , and normal vector, $\bar{\mathbf{n}}_{WXYZ}$.

The reflected sun vector, \mathbf{r}_{ref} , can be calculated by

$$\mathbf{r}_{\text{ref}} = \mathbf{r}_{\text{sun}} - 2\bar{\mathbf{n}}_{ABCD}^\top (\mathbf{r}_{\text{sun}} \cdot \bar{\mathbf{n}}_{ABCD}), \quad (5.2)$$

where \mathbf{r}_{sun} is the incoming sun vector. To calculate the intersection of the reflected vector with the $WXYZ$ -plane of the sun sensor, the equation of the $WXYZ$ -plane, the reflected

vector, and the point of origin is required. The reflection of the sun vector is illustrated in Figure 5.2. The reflection from Q to Q' can thus be calculated as a projection of \mathbf{r}_{ref} unto the $WXYZ$ -plane.

To model reflection from the solar panels to the sun sensor, only two corners of the solar panel and two corners of the sun sensor are to be considered. From Figure 5.2 it is evident that if the solar panel reflects on Y that the reflection will also cover X . The same is true for corner Z and W . Since C' will be at the same position as C , which is valid for D' and D , the calculation can be omitted. Therefore it is only necessary to calculate the reflected positions A' and B' . This simplifies the reflection model significantly.

The reflected position A' can be calculated as the intersection of the reflected vector R with plane $WXYZ$. We also know the position of A , based on the satellite design, and can therefore calculate A' . The same applies to B and B' . To determine whether Y or X is within the reflection region, we assume that the plane $WXYZ$ is a 2D plane, and we omit the third dimension. Therefore, the axis changes from x, y, z to only x, y . We calculate whether x is between the lines of $A'D$ and $B'C$ and between the lines CD and $A'B'$. By determining the line equation between reflected points in the form

$$y_{A'B'} = mx_{A'B'} + c, \quad (5.3)$$

from the coordinates of A' and B' , the corresponding $X_{A'B',y}$ can be calculated by substituting X_x into Eq 5.3. With the same method the coordinates of $X_{B'C}$, $X_{A'D}$, $X_{A'B'}$ and X_{CD} can be determined. After that, with logical if statements, it can be determined whether X is in the reflection zone. If X_x is to the right of $X_{B'C,x}$ and to the left of $X_{A'D,x}$, as well as X_y is above $X_{A'B',y}$ and below $X_{CD,y}$ then X is within the reflection zone.

The results for the sun vector with and without reflection are shown in Figure 5.3. There is a clear difference between the true sun vector and the possible measurement influenced by the reflection. This false reflection vector may affect the estimation and, thus also the attitude control of the satellite.

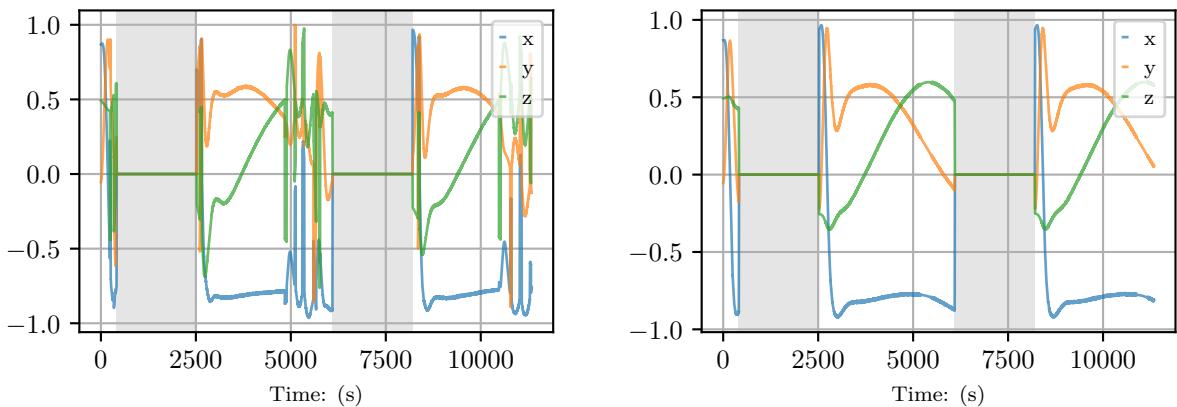


Figure 5.3: Comparison of Sun Vector with and without Reflection

5.1.1. Influence of anomaly on estimation

To determine whether the reflection on the sun sensor has a influence on the ADCS, the estimation metric is shown in Figure 5.4. The estimation metric, is the angle difference between the actual attitude and the estimated attitude. It is evident that the reflection has a large influence on the estimation. It is also clear that during the eclipse the estimation returns to a more accurate estimation. This is due to the fact that all sensors are ignored if the measured vector is 0.

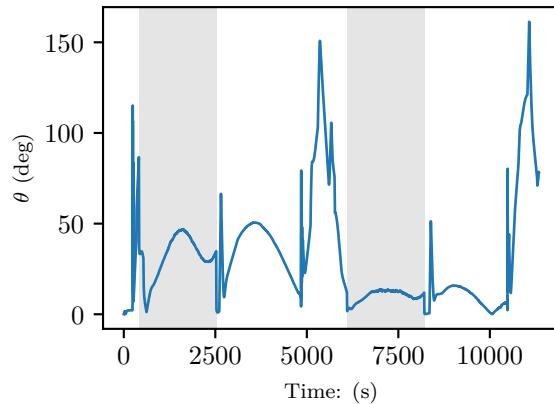


Figure 5.4: Estimation Metric with reflection on sun sensor

5.2. Moon and Sun in Field of View of Nadir Sensor

An anomaly that can be experienced by an infra-red (IR) nadir sensor is the moon overlapping the horizon of the Earth in die nadir sensor's field of view (FoV). This is shown in Figure 5.7. This influences the edge detection and circular fit algorithm [?, ?] and consequently the calculated centre of the Earth. Firstly, it is required to simulate the image seen by the nadir sensor, thereafter the algorithm for detecting the centre of the Earth can be implemented.

5.2.1. Simulating Nadir Sensor Infra-red Image

Firstly the vectors of both the satellite to Earth and the moon the Earth is required. The moon position is determined with the Julian date, since the propagation of the moon position relative to the centre of the Earth has already been done. These vectors are shown in Figure 5.5. From the vector, \mathbf{R}_{SE} , and the position of the centre of the Earth, P_{Earth} , a 3D plane normal to \mathbf{R}_{SE} and at P_{Earth} can be calculated. Where both P_{Earth} and \mathbf{R}_{SE} are defined as

$$P_{Earth} = [x_0, y_0, z_0] \quad (5.4)$$

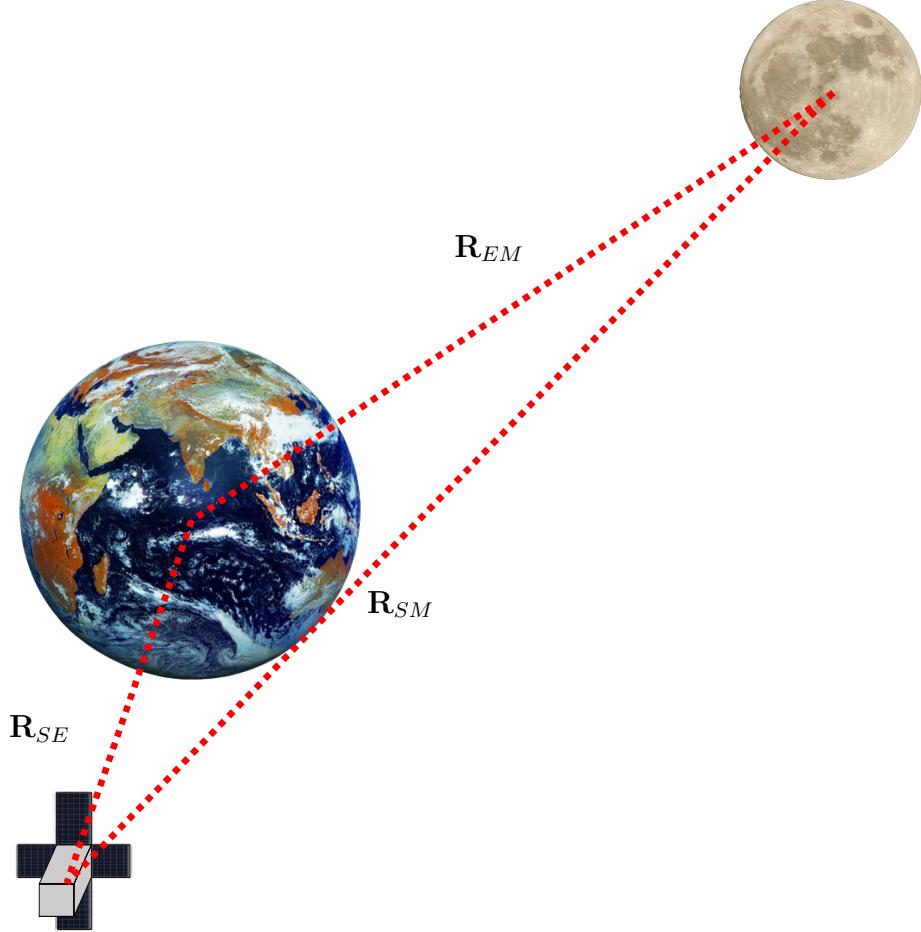


Figure 5.5: Earth to Moon and Earth to Satellite Vectors

and

$$\mathbf{R}_{SE} = [n_x, n_y, n_z]. \quad (5.5)$$

Therefore with the equation for the 3D plane defined as

$$Ax + By + Cy = D \quad (5.6)$$

the parameters A, B, C, D can be calculated as

$$\begin{bmatrix} A \\ B \\ C \\ D \end{bmatrix} = \begin{bmatrix} n_x \\ n_y \\ n_z \\ n_x x_0 + n_y y_0 + n_z z_0 \end{bmatrix} \quad (5.7)$$

This 3D plane slicing the Earth in half as seen from the satellite is shown in Figure 5.6. The moon and Earth can both be projected unto the 3D plane to determine the image seen by the nadir sensor. Therefore the nadir vector must also be projected unto the 3D

plane. A circle can be drawn for the Earth, moon and the nadir sensor FoV. The radius, of the moon as projected on the 3D plane can be calculated as

$$R_{moon} = \|R_{SE}\| \frac{r_{moon}}{\|R_{SM}\|} \quad (5.8)$$

the radius of the nadir sensor FoV, R_{FoV} can calculated as

$$R_{FoV} = \|R_{SE}\| \tan \theta \quad (5.9)$$

With these variables defined and calculated, the edges of the moon and Earth within the nadir FoV can be determined.

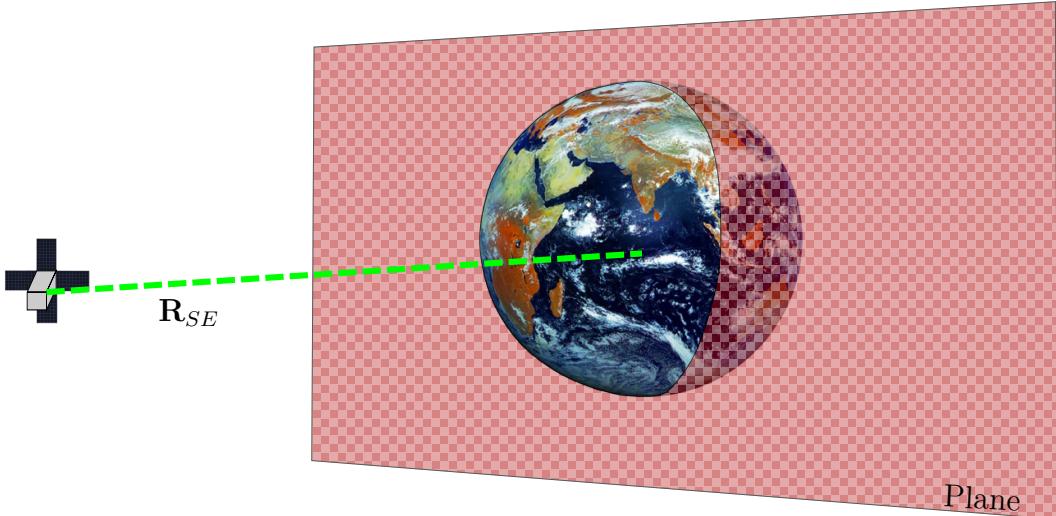


Figure 5.6: Plane perpendicular to \mathbf{R}_{SE} and at center of Earth

Firstly the edges of the moon and Earth are discretely determined based on a fixed number of points for the Earth, N , and the number of discrete points on the moon is determined based on the ratio of R_{moon} to R_{Earth} . The projected Earth and moon unto the 3D plane is shown in Figure 5.6. The discrete edges of the Earth that is within the FoV and will be used for the algorithm discussed in section 5.2.2 are determined with the following logical statements.

1. Distance between point and centre of nadir Sensor FoV must be smaller than R_{FoV} .
2. Distance between point and centre of moon must be larger than R_{moon}

The discrete edges of the moon used for the algorithm is must satisfy the following conditions

1. Distance between any discrete point and centre of Earth must be smaller than R_{Earth} for the moon to overlap the horizon
2. Distance between point and centre of nadir Sensor FoV must be smaller than R_{FoV} .

3. Distance between point and centre of Earth must be larger than R_{Earth}

This creates the array of points that will be used in the algorithm used to calculate the centre of the Earth.

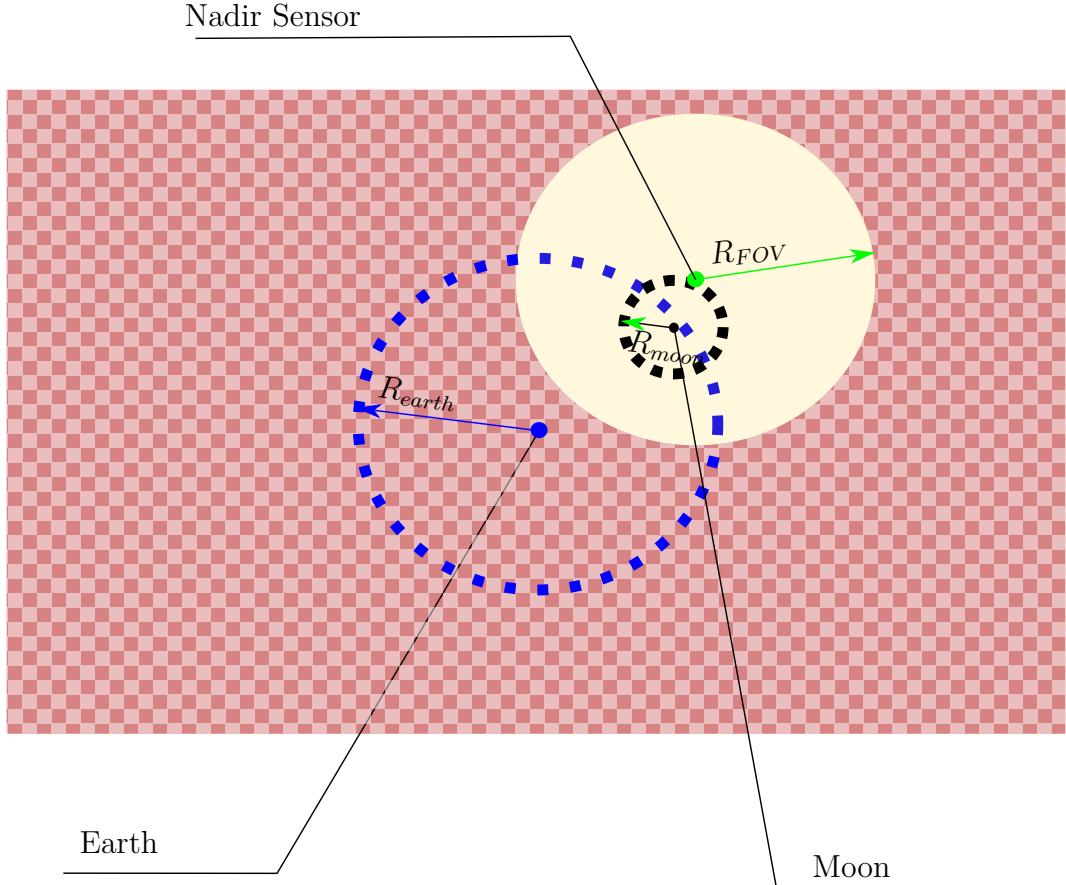


Figure 5.7: Projection of moon and Earth on plane

5.2.2. Calculating the Centre of the Earth

The edges of the Earth are detected based on a gradient between the lowest temperature and the highest temperature within the IR nadir sensor's FoV. This will not be implemented in our case, since we can determine discrete points of both the Earth and the moon from the simulation environment. Furthermore the visible phases of the moon will not be accounted for. The reasoning for this is due to the coldest side of the moon being $140K$ and the warmest part, $400K$. The temperature of space is $2.7K$ and the coldest part on the Earth is $180K$. Consequently, the IR horizon sensor must be calibrated to always use the minimum value for edge detection as $180K$ or it must use the smallest value in the image, which will most likely be $2.7K$. Therefore, it can be assumed that the moon will not have any detectable phases for the IR horizon sensor and it will always be seen as a full moon, due to its lowest temperature being warmer than that of space.

With this assumption the circular fit algorithm as shown in Figure 5.8 can now be used to determine the centre of the Earth on the plane [?]. For this calculation the 3D plane is transformed to a 2D plane and all the coordinates is also transformed.

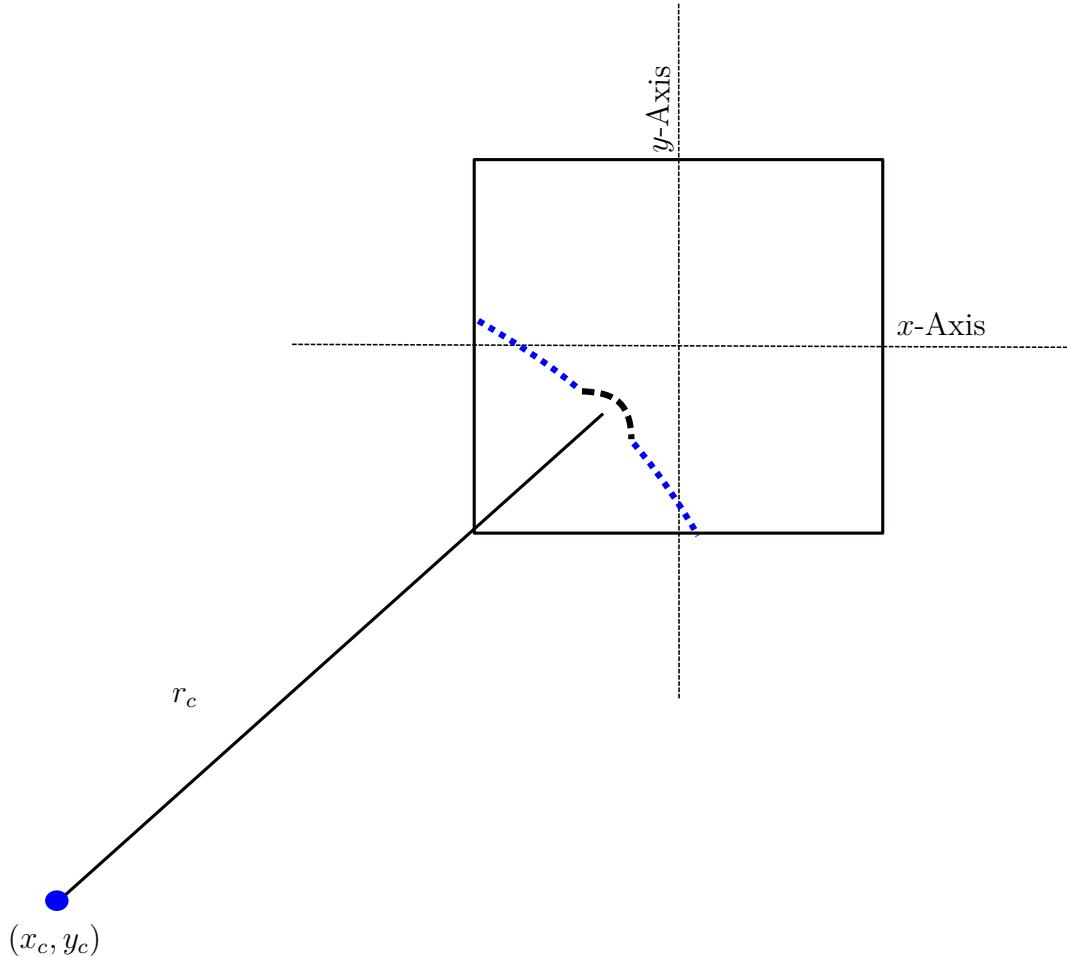


Figure 5.8: Circular Fit Algorithm

Firstly the curvature is described as

$$ax + bx + c = x^2 + y^2 \quad (5.10)$$

where

$$\begin{aligned} a &= 2x_c \\ b &= 2y_c \\ c &= r_c^2 - \sqrt{x_c^2 + y_c^2} \end{aligned} \quad (5.11)$$

Therefore using all the discrete edges within the image a , b and c can be calculated as

$$\begin{bmatrix} x_0 & y_0 & 1 \\ x_1 & y_1 & 1 \\ \vdots & \vdots & \vdots \\ x_n & y_n & 1 \end{bmatrix} \begin{bmatrix} a \\ b \\ c \end{bmatrix} = \begin{bmatrix} x_0^2 + y_0^2 \\ x_1^2 + y_1^2 \\ \vdots \\ x_n^2 + y_n^2 \end{bmatrix} \quad (5.12)$$

where (x_0, y_0) to (x_n, y_n) are the coordinates of the discrete edges. It is thus evident that when the moon overlaps the horizon of the Earth from the nadir sensor's perspective the centre of the Earth will be incorrectly calculated and this anomaly must be dealt with. The other anomaly in this section where the sun is in the FoV of the nadir sensor will not provide a measurement, since the sun will saturate the Infra-red nadir sensor [?].

5.2.3. Influence of anomaly on estimation

To determine the effect of the moon in the circular fit algorithm the measured Earth vector with and without the moon on horizon anomaly is shown in Figure 5.9. It is clear from Figure 5.9 that the anomaly has no visible effect on the Earth vector. It is also evident in Figure 5.10 that the estimation metric is also not influenced negatively by this anomaly. Therefore, this anomaly is not included in the FDIR development, since there is no evident difference due to the anomaly.

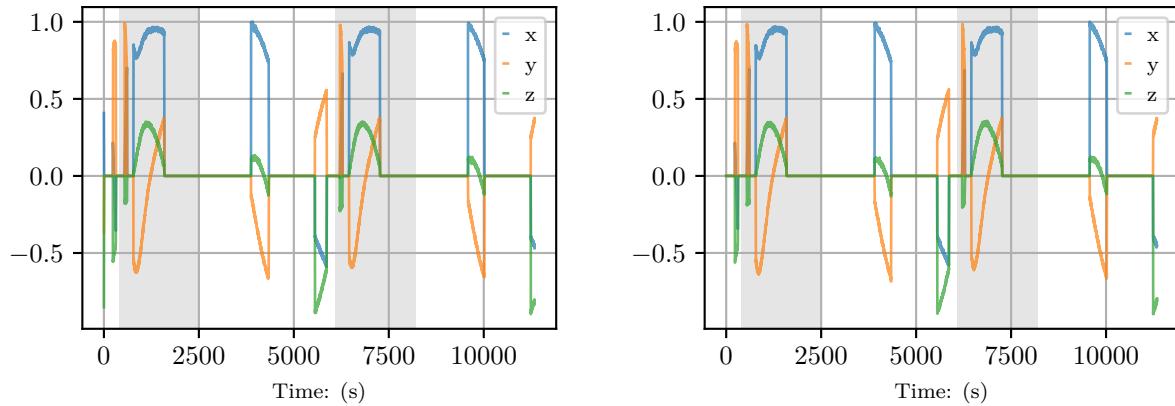
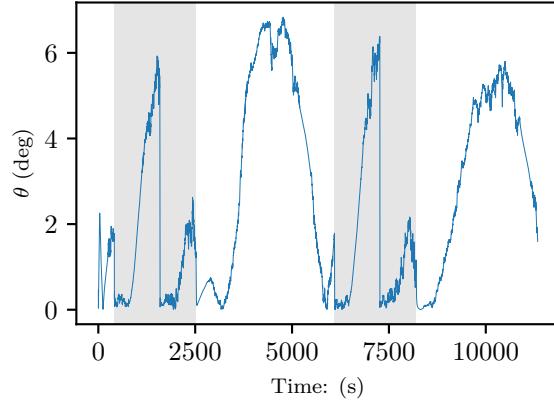


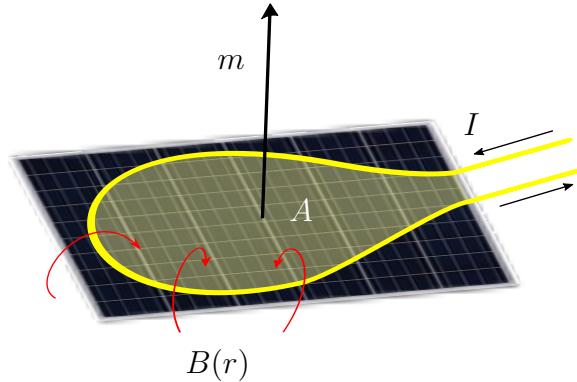
Figure 5.9: Comparison of Earth Vector with and without moon on the horizon

5.3. Magnetic Moment Disturbance from Satellite Bus

Magnetic moments produced by a coil in solar panels on a CubeSat can create a disturbance torque and influence the magnetometer measurements due to the induced magnetic field in the coil of the solar panel [?, ?]. According to [?] the current, I , in each individual cell of the solar panel can be modelled as a cumulative current for the entire solar panel, since the normal vector to each cell and the solar panel is the same. This magnetic moment is

**Figure 5.10:** Earth Values

modelled for the specific size of the CubeSat model in Figure 5.1. The coil in the solar panel and the resulting magnetic field, $B(r)$, as well as the resulting dipole moment, m , is shown in Figure 5.11. The inner area of the coil is assumed to be the same as the surface area of the solar panel.

**Figure 5.11:** Dipole Moment from circular loop in solar panel

The dipole moment is calculated as

$$m = AI \quad (5.13)$$

where A is the area within the loop and I is the current within the coil. The current is generated by the solar panel and thus depends on the incoming sun vector as well as the area on the solar panel illuminated by the sun as demonstrated in Figure 5.12. Therefore the current, I , can be calculated with

$$I = I_{max} \frac{A_{total}}{A_{illuminated}} \cos(\theta) \quad (5.14)$$

where θ is the angle between the normal vector to the solar panel and the incoming sun vector and I_{max} depends on the solar panel. The dipole moment in term produces a

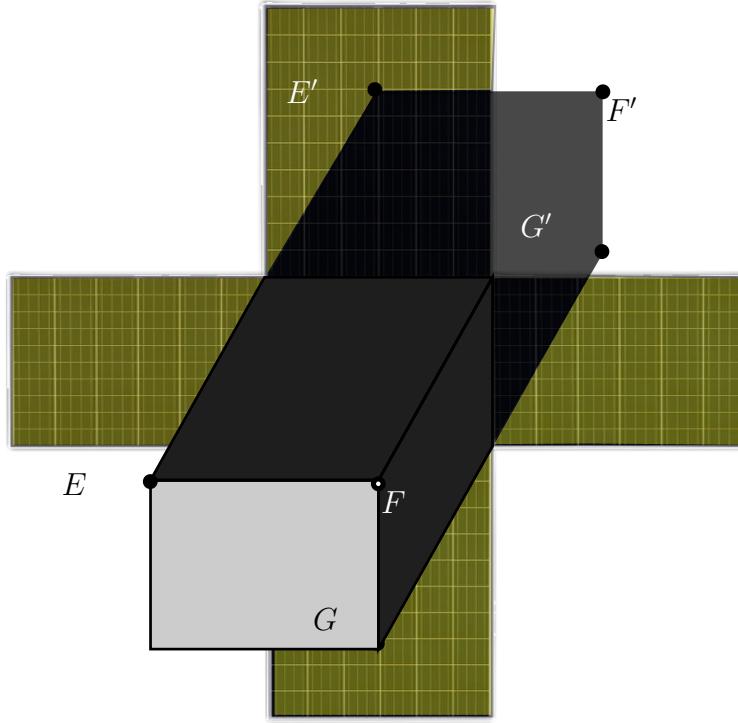


Figure 5.12: Shadow created by CubeSat body on Solar Panels

disturbance torque on the CubeSat. With the resulting torque expressed as

$$\tau = m \times B \quad (5.15)$$

where B is the magnetic field of the Earth. Since the only external magnetic field that creates a resulting torque is that of the Earth. The resulting torque for five orbits are shown in Figure 5.13.

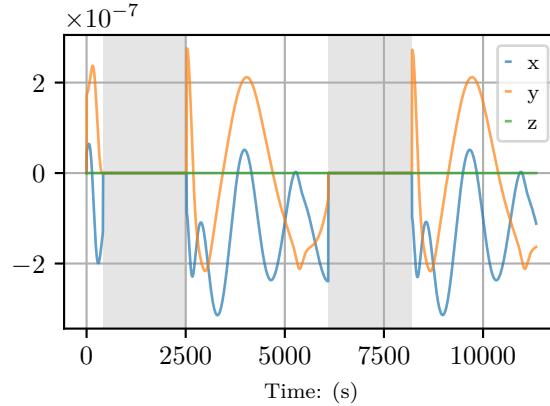


Figure 5.13: Solar Panel Disturbance Torques

The magnetometer measurement influenced by the magnetic field produced by the coil

in the solar panel. This magnetic field experienced at the magnetometer can be calculated with

$$B(r) = \frac{\mu_0}{4\pi} \frac{3\hat{r}(\hat{r} \cdot m) - m}{\|r\|^3} \quad (5.16)$$

5.3.1. Influence of anomaly on estimation

The vector r between the position of the magnetometer and the solar panel influences the magnetic field significantly. The experienced magnetic field by the magnetometer will be different for each solar panel. The resulting measured vector by the magnetometer is the summation of the Earth's magnetic field and the magnetic field produced by the coils in the solar panels. The resulting magnetometer measurement with and without the induce moment is shown in Figure 5.14. From Figure 5.15 is is evident that this anomaly has a significant effect on the estimation, but not as large as the reflection on the sun sensor.

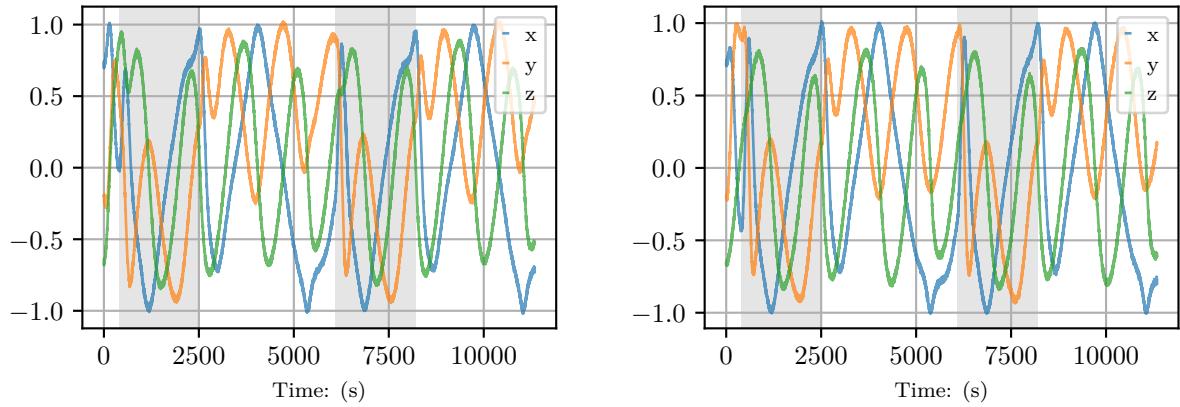


Figure 5.14: Comparison of Magnetic field Vector with and without solar Panel magnetic field

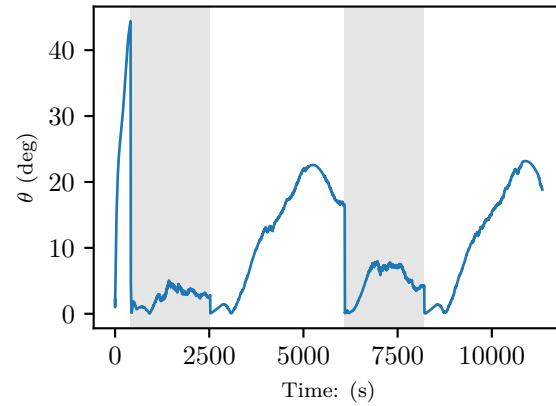


Figure 5.15: Estimation Metric with induced dipole moment

5.4. Reaction wheels

When an actuator fails it influences all the sensor measurements. The anomaly will be modelled as a sudden failure in the actuator when it does not react to inputs. This will influence the EKF, since the model update will be inaccurate. Therefore, this anomaly is included even though it is not a sensor anomaly, because it often occurs and a lot of research is done in this. The recovery of this will however will not be within the specification of this thesis, but the model update for the estimation will be modified based on a modified torque vector. The resulting estimation metric for this anomaly is shown in Figure 5.16 and it is evident that this anomaly has a large negative effect on the EKF.

maybe use the angular moment sensors to update the torque.

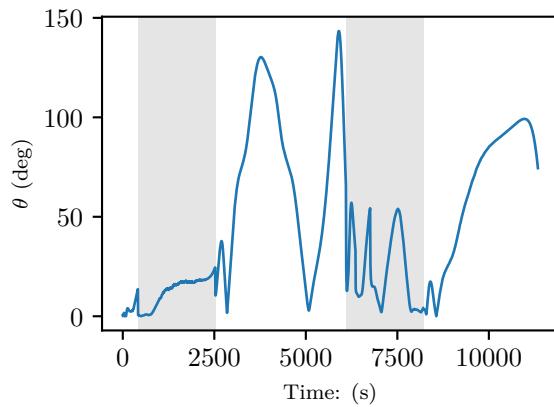


Figure 5.16: Estimation Metric with failure of Reaction Wheels

5.5. Summary

Based on the influence of the modelled anomalies on the estimation metric. Therefore the reflection of the solar panels on the sun sensor, the induced magnetic dipole and the failed reaction wheel are used for the FDIR development.

Chapter 6

Feature Extraction

6.1. Binary Feature Extraction

The classes hyperparameter will be set to 2. The only problem with this, is a very evident 2 classes is eclipse and non-eclipse. Consequently each feature extraction model will be trained twice. Once for the eclipse and once for the non-eclipse period. The features extracted will be used with the detection algorithms.

6.1.1. Local Outlier Factor

6.1.2. Dynamic Mode Decomposition

The proposed method by [?] uses Dynamic Mode Decomposition (DMD), which was initially developed by [?] and further expanded to include control by [?], to provide an estimation of a sensor vector based on the previous measurement of the sensor as well as the measurements of the other sensors in the system. DMD was first developed in the fluids community and constructed a matrix \mathbf{A} to relate the state vector x with the following time step of the state vector, x_{k+1} . The state vector, in our case, will be the measurement vector of the specific sensor that we want to monitor.

$$\mathbf{x}_{k+1} = \mathbf{A}\mathbf{x}_k \quad (6.1)$$

Where \mathbf{x}_k and \mathbf{x}_{k+1} during a specified number of time steps, will be denoted as \mathbf{X} and \mathbf{X}' respectively.

The method of DMD, however, is useful for high order systems where the calculation of \mathbf{A} is computationally intensive. This is not the case for our system, and using DMD is not justifiable and consequently, a linear regression model is implemented. Therefore with the pseudo-inverse of \mathbf{X} , denoted as \mathbf{X}^\dagger , we calculate \mathbf{A} as

$$\mathbf{A} = \mathbf{X}\mathbf{X}^\dagger \quad (6.2)$$

This necessitates data of the state vector over time. The article by [?] however includes \mathbf{B} to relate the vector measurements of the other sensors to adjust the predicted state, X_{k+1}

of the monitored sensor.

$$\mathbf{X}_{k+1} = \mathbf{AX}_k + \mathbf{BY}_k \quad (6.3)$$

Where \mathbf{Y}_k is the other sensor measurements, this is adjusted for our use case, where \mathbf{Y}_k is the control torques for the magnetorquers and reaction wheels, while \mathbf{X}_k is all of the sensor measurements. Consequently, the model of Eq 6.3 denotes the prediction of the sensor measurements at time step $k + 1$ based on the current sensor measurements and control inputs. Thereafter, as implemented by [?] the model is adjusted with a Kalman Filter. From \mathbf{A} and \mathbf{B} the Kalman filter can be implemented to predict $\hat{\mathbf{X}}_{k+1}$

$$\hat{\mathbf{X}}_{k+1} = \mathbf{AX}_k + \mathbf{BY}_k + K(\mathbf{X}_k - \hat{\mathbf{X}}_k) \quad (6.4)$$

where $K = 0.001$. After the calculation of $\hat{\mathbf{X}}_{k+1}$ [?] proposes a moving average of the innovation covariance

$$\mathbf{V}_k = \frac{1}{N} \sum_{i=k-N}^k (\mathbf{X}_i - \hat{\mathbf{X}}_i)(\mathbf{X}_i - \hat{\mathbf{X}}_i)^T \quad (6.5)$$

where N is the number of timesteps to account for. The moving average is used as an additional input parameter for the classification of anomalies based on \mathbf{X} .

6.2. Multiple Feature Extraction

The classes hyperparameter will be set to the number of anomalies (plus 1). The features will be used in the isolation algorithms. The only problem with this, is a very evident 2 classes is eclipse and non-eclipse. Consequently each feature extraction model will be trained twice. Once for the eclipse and once for the non-eclipse period.

6.2.1. Linear Regression

6.2.2. Local Outlier Factor

6.2.3. K-means Clustering

6.2.4. Prony's Method

6.2.5. Partial Least Square

6.2.6. t-Distributed Stochastic Neighbor Embedding

6.3. Summary

Chapter 7

Recovery

7.1. Analysis

7.2. EKF-ignore

7.3. EKF-combination

7.4. EKF-reset

7.5. EKF-top2

Chapter 8

Detection

8.1. Analysis

8.2. Supervised Learning

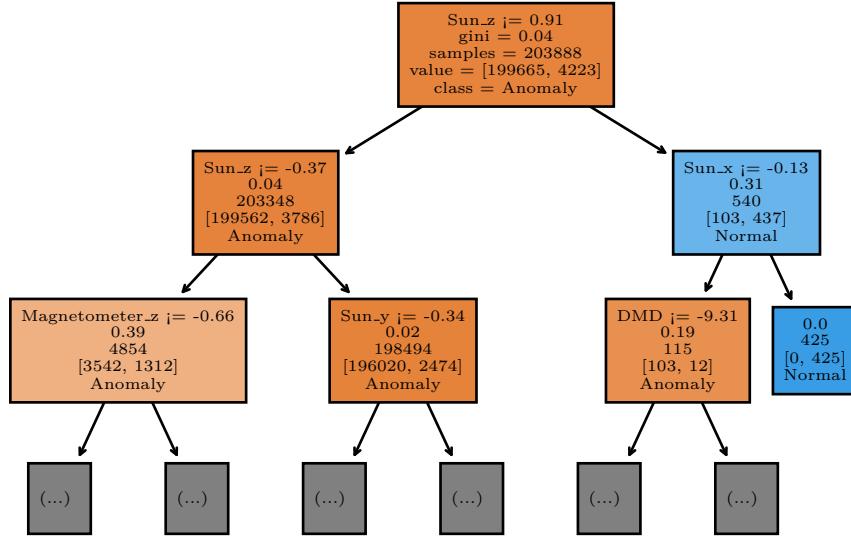
8.2.1. Random Forests

8.2.2. Decision Trees

However, to split the data for the anomalies, we need to decide which input parameter will be used to make the first split, the root node. The Gini index measures the probability of a data sample being wrongly classified at a given node. This can be calculated with Eq 8.1.

$$GI = 1 - \sum_{i=1}^n (P_i)^2 \quad (8.1)$$

The operator split that produces the lowest Gini index provides the purest split and will be used as the root node. For our use case, the CART algorithm will be used to optimise the decision tree, which also considers the most prominent information gain to construct the decision tree. Figure 8.1 is a graphical representation of the decision tree developed to classify anomalies. The depth of a decision tree determines how many splits occur from the root node to the leaf node the furthest from the first split. If the depth is unspecified, the decision tree will split until all the data samples are perfectly split into anomalous and normal data samples. However, the larger the depth, the more biased the decision tree is to the training data. This depth can be altered to optimise the efficiency and accuracy of the decision tree.

**Figure 8.1:** Decision Tree

8.2.3. Support Vector Machines

8.3. Unsupervised Learning

8.3.1. Isolation Forests

This unsupervised learning method is based on the principle of isolating data points by slicing the data with random conditions [?]. The data is randomly split into specified sample sizes with a randomly selected dimension and a randomly selected cut-off value. For each sample size the data must be split until each data point within the sample is isolated from all other data points. Training of a single tree is completed when all the data points are isolated and this training must be repeated for all the data samples, however many are predefined.

The distance measured from the first split the *tree top* to the isolated data point is used to determine whether a data point is anomalous or not [?]. The logical reasoning for support of this algorithm is that data points which are non-anomalous will be more closely related and hence have more splits to separate the data points until isolation is achieved. Therefore, the distance from the tree top for non-anomalous data points will be longer than anomalous data points which will have a shorter distance from the tree top. Therefore non-anomalous data points are closer to the *root*.

Figure 8.3 demonstrates the splitting of the data points until isolated. Each split or

branch only splits the data into two groups. After training multiple trees, a single data point is "sent through the forest" and the distance from the tree top for each tree is calculated and the average of all the trees are used to calculate the average distance for the data point. Using a threshold for the distance, the data point is classified as anomalous or not.

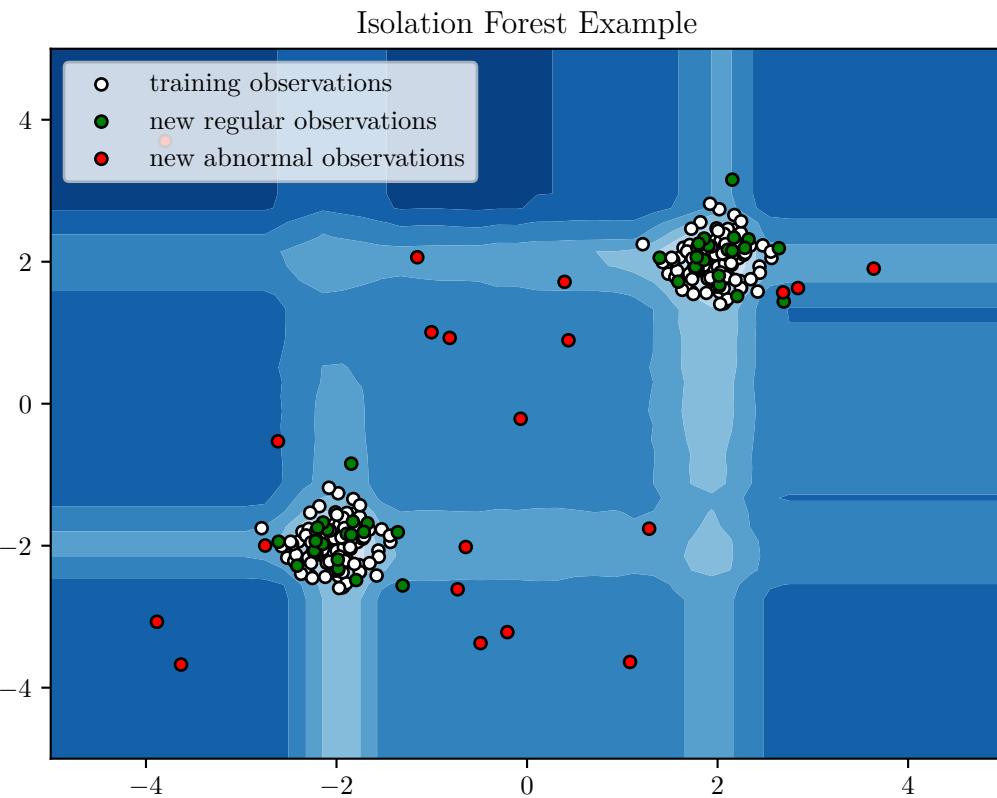


Figure 8.2: Isolation Forest

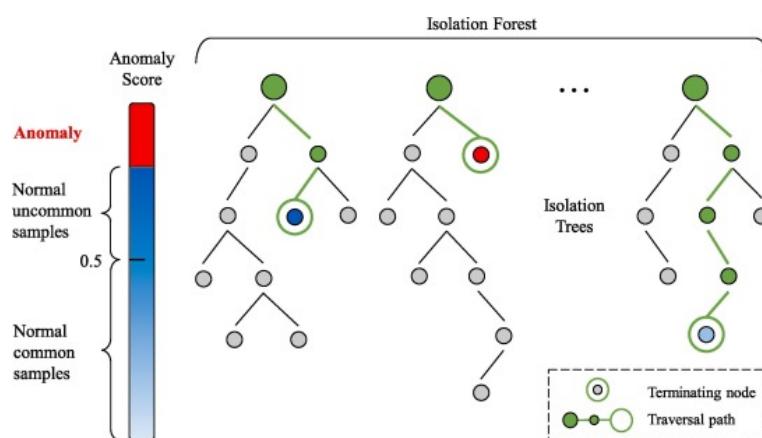


Figure 8.3: Isolation Forests [?]

The anomaly score is calculated with Eq 8.2

$$s(x, n) = 2^{-E(h(x))/c(n)} \quad (8.2)$$

where $E(h(x))$ is the average value of the distance measured from the tree top for a single data point in all the trees [?] and n is the size of a data sample used to train a single tree. For the distance to be normalized, $c(n)$ — the mean distance from the tree top in an unsuccessful search in a *Binary Search Tree* (BST) — is used and is calculated as

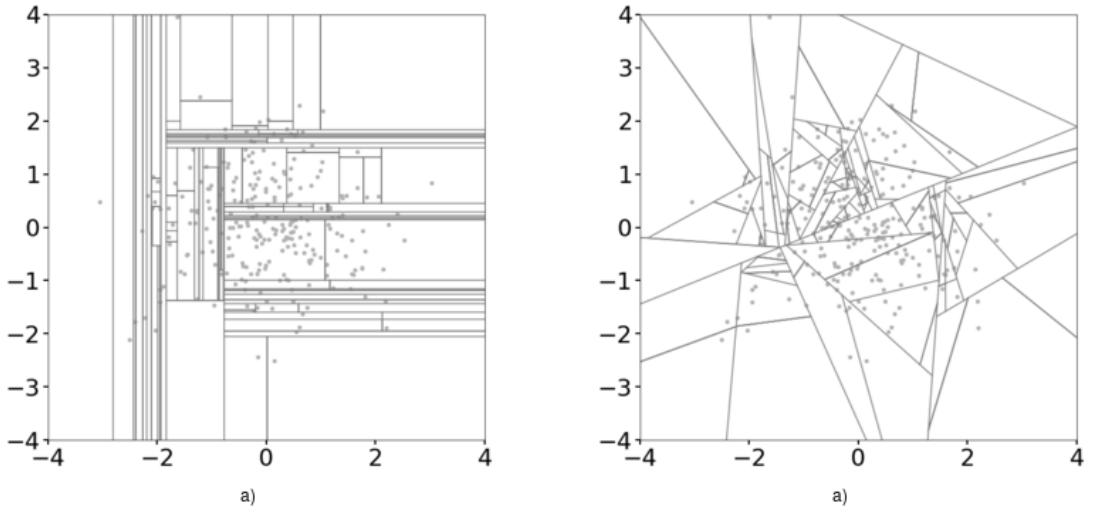
$$c(n) = 2H(n - 1) - \frac{2(n - 1)}{n}. \quad (8.3)$$

$H(i)$ in Eq 8.3 is the harmonic number and is estimated with Euler's constant as

$$H(i) \approx \ln(i) + 0.5772156649. \quad (8.4)$$

Isolation Forests, however have multiple issues, since it splits data in rectangles as seen in Figure a). This is due to the slicing algorithm selecting a feature, x and a cut-off value, v . Consequently, the data is either split vertically or horizontally — if seen as a two dimensional dataset. This split method is unable to categorise complex data structures. These issues however are addressed by [?] and led to the *Extended Isolation Forest* algorithm.

The extended isolation forest algorithm generalises the isolation forest algorithm by applying a slope to each slice. Data points are therefore divided into two groups depending on the "side" of the plane or slice as seen in Figure b).



(a) Isolation Forest Slicing example

(b) Extended Isolation Forest Slicing example

Figure 8.4: The slicing of Isolation Forest vs Extended Isolation Forest

It is evident that applying an angle of 0° to all the slices the general algorithm of the

extended isolation forest produces the standard isolation forest algorithm where planes or slices are perpendicular to the axis of the randomly selected feature, x .

8.3.2. Local Outlier Factor

Most algorithms for anomaly detection are based on a metric which accounts for the entire dataset [?]. However, many anomalies are identifiable in relation to the local neighbourhood of data points and not the overall dataset. Therefore, [?] developed the local outlier factor *LOF* algorithm that provides a measure of a data point's "outlierness". This implies that a data point is not classified as an anomaly or not, but a local outlier factor is calculated to determine how much a data point is distantiated from it's k -nearest neighbours. This is clearly demonstrated in Figure ?? where the data points which are clustered together have smaller LOF's than data points which are removed from the highly dense areas.

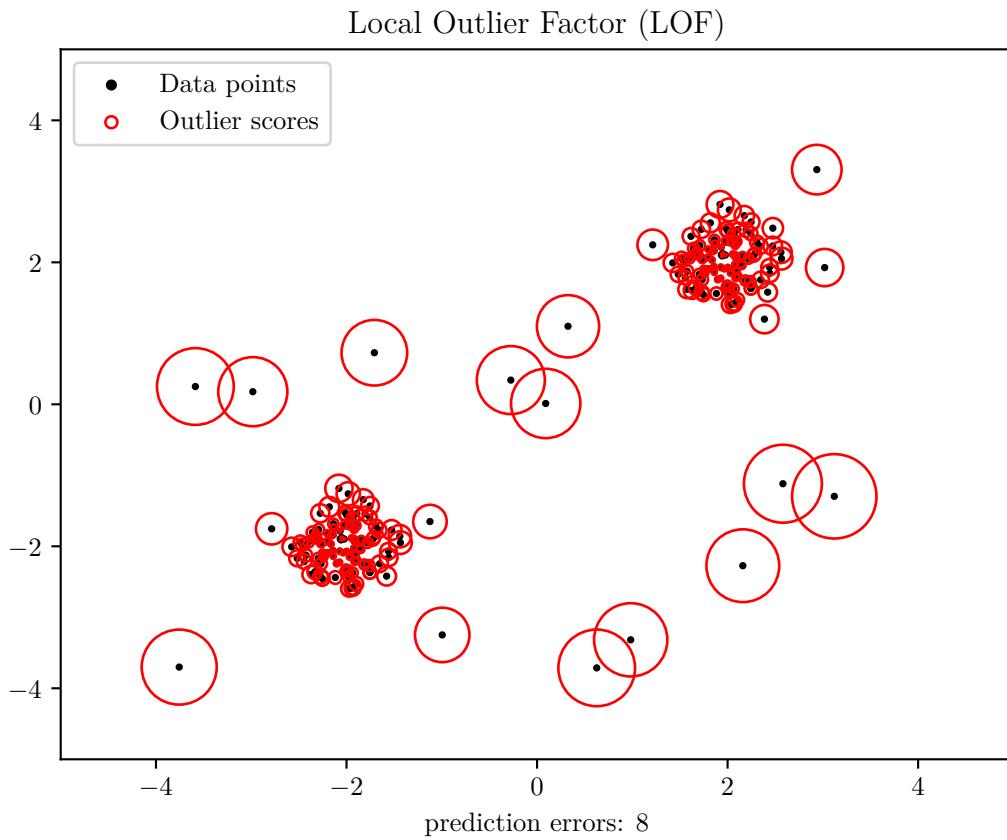


Figure 8.5: Plane perpendicular to \mathbf{R}_{SE} and at center of earth

To calculate the LOF, the k -distance must be calculated and also the local reachability density lrd . The k -distance, is the k^{th} ranked $distance(o, p_i)$. Where $distance(o, p_i)$ is the distance between data point o and any data point p_i , with $i \in N$, where N is the number of data points within the dataset with a minimum value of $MinPts$. To reduce

fluctuations in the $distance(o, p_i)$ the distance between o and p_i is replaced with

$$\max\{distance(o, p_i), k\text{-distance}\} \quad (8.5)$$

and will henceforth be referred to as the reachability distance [?]. The lrd of a data point, p , is calculated as

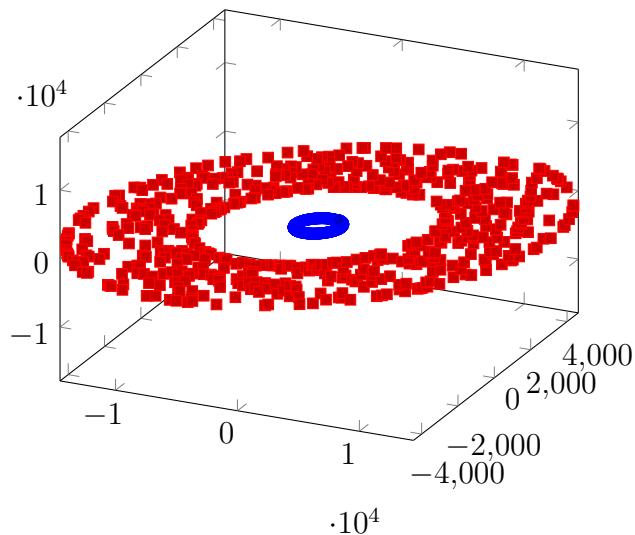
$$lrd_{MinPts}(p) = 1 / \left(\frac{\sum_{o \in N_{MinPts}(p)} reachdist_{MinPts}(p, o)}{|N_{MinPts}(p)|} \right) \quad (8.6)$$

and denotes "the inverse of the average reachability distance based on the $MinPts$ -nearest neighbours of the p " — [?]. Eq 8.6 enables the calculation for the LOF of point p as shown in Eq 8.7

$$LOF_{MinPts}(p) = \frac{\sum_{o \in N_{MinPts}(p)} \frac{lrd_{MinPts}(o)}{lrd_{MinPts}(p)}}{|N_{MinPts}(p)|} \quad (8.7)$$

The rule of thumb for detecting an outlier is that when the LOF is larger than 1, then the point is considered an outlier with respect to its neighbourhood. This however is not fixed and the threshold can be changed depending on the application. This method is aimed at producing a measure of the "outlierness" of a data point within a local neighbourhood and not for all the data points. This method will thus be implemented for the satellite anomaly detection, since it will detect anomalies within the two neighbourhoods produced by the eclipse during orbit. This method will also be able to detect measurements of earth sensors, sun sensors and magnetometers that drastically change from the previous orbital data. For example in Fig ?? it is evident that the LOF will be comparatively larger for the red data points, which are anomalies, to the blue data points that are the normal orbit of the satellite.

Earth Sensor During Multiple Orbits



8.4. Summary

Chapter 9

Isolation

9.1. Analysis

9.2. Random Forests

9.3. Decision Trees

9.4. Support Vector Machines

9.5. Summary

Chapter 10

Results

10.1. Perfect Extended Kalman Filter

For all these results the best recovery method will be demonstrated

10.2. Unsupervised Detection and Supervised Isolation

Use the best unsupervised learning algorithm from chapter Detection and the best supervised learning algorithm from chapter Isolation

10.3. Supervised Detection and Isolation

Use the best unsupervised learning algorithm from chapter Detection and the best supervised learning algorithm from chapter Isolation

10.4. Summary

Chapter 11

Conclusion

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11.1. Project/thesis/dissertation summary

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11.2. Appraisal of project/thesis/dissertation contributions

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11.3. Suggestions for future work

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11.4. What the student has learnt during this project

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Appendix A

Project Planning Schedule

This is an appendix.

Appendix B

Outcomes Compliance

This is another appendix.