FYS3150 - Project 3

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Abstract om du vil

I. INTRODUCTION

In this report we will study the solar system in three dimensions.

II. THEORY

A. Newton's laws and the motion of planets

The well-known Newtons second law reads

$$\sum \mathbf{F} = m\frac{\mathrm{d}^2\mathbf{r}}{\mathrm{d}t^2}, \tag{1}$$

where $\sum \mathbf{F}$ is the sum of forces acting on an object, m is its mass and $\mathbf{r}$ its position.

Newton's law of gravity says that the total gravitational force acting on an object i from all the other objects in a system is

$$F_{G,i} = -Gm_i \sum_{j\neq i} m_j \frac{\mathbf{r}_i - \mathbf{r}_j}{|\mathbf{r}_i - \mathbf{r}_j|^3} \tag{2}$$

where $\mathbf{r}_i$ are the positions and m_i are the masses of the objects. G is the gravitational constant.

For the motions of planets, on which the gravitational force is the only force acting, the two equations 1 and 2 combined give us the differential equation that describes the motion of an object:

$$\frac{\mathrm{d}^2\mathbf{r}_i}{\mathrm{d}t^2} = \frac{1}{m_i}F_{G,i} = -G\sum_{j\neq i} m_j \frac{\mathbf{r}_i - \mathbf{r}_j}{|\mathbf{r}_i - \mathbf{r}_j|^3} \tag{3}$$

B. Velocity Verlet

Classical problem of Newtonian mechanics often involve solvinga set of two coupled first order differential equations, namely

$$\frac{\mathrm{d}x}{\mathrm{d}t} = v$$
$$\frac{\mathrm{d}v}{\mathrm{d}t} = a$$

where x is position, v is velocity and a is acceleration. Doing a Taylor expansion of x around a point in time $t = t_0$ gives

$$\begin{aligned} x(t) &= x(t_0) + (t-t_0)\frac{\mathrm{d}x}{\mathrm{d}t}(t_0) + \frac{1}{2}(t-t_0)^2\frac{\mathrm{d}^2x}{\mathrm{d}t^2}(t_0) + O(h^3) \\ &= x(t_0) + (t-t_0)v(t_0) + \frac{1}{2}(t-t_0)^2 a + O(h^3) \end{aligned}$$

An expression can then be found for position at a time $t = t_0 \pm h$:

$$x(t_0 \pm h) = x(t_0) \pm hv(t_0) + \frac{1}{2}h^2 a \pm O(h^3) \tag{4}$$

Discretizing equation 4 and letting $x_i = x(t)$, $x_{i+1} = x(t+h)$, $v_i = v(t)$ and $v_{i+1} = v(t+h)$, we get

$$\begin{aligned} x_{i+1} &\approx x_i + hv_i + \frac{h^2}{2}a_i \\ v_{i+1} &\approx v_i + ha_i + \frac{h^2}{2}\dot{a}_i, \end{aligned}$$

where $\dot{a}_i$ is the derivative of the acceleration with respect to time. Similarly to the case of Forward Euler, doing a Taylor expansion of a gives after some manipulation

$$\dot{a}_i \approx \frac{a_{i+1} - a_i}{h} \tag{5}$$

Insering this into the expressions we have for x_{i+1} and v_{i+1} gives us the equations that describe the Velocity Verlet method:

$$\begin{aligned} x_{i+1} &\approx x_i + hv_i + \frac{h^2}{2}a_i \\ v_{i+1} &\approx v_i + \frac{h}{2}(a_{i+1} + a_i) \end{aligned}$$

From equation 4 we see that the mathematical error in this approximation goes like $O(h^3)$.

C. Forward Euler

From equation 4 we see that by including only the two first terms in the Taylor expansion of $x(t)$ and $v(t)$ we get

$$\begin{aligned} x_{i+1} &\approx x_i + hv_i \\ v_{i+1} &\approx v_i + ha_i \end{aligned}$$

This is referred to as the Forward Euler method of solving differential equations. The mathematical error in this approximation goes like $O(h^2)$.

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D. Conservation of angular momentum

Kepler's second law states that if you draw a line from the Sun to a planet orbiting it, then that line would sweep out the same area in equal periods of time. We will use this law to derive the conservation of angular momentum.

For short periods of time dt the area swept out by the line from the Sun to the planet is approximately a triangle with area

$$A = \frac{1}{2} r v_\theta \mathrm{d}t$$

where r is the distance from the planet to the sun and v_θ is the tangential velocity of the planet. Kepler's second law states that this area is constant for all intervalls of time of the same length. For this to be true, we must require

$$r v_\theta = \text{constant}$$

The angular momentum $\mathbf{L}$ of the planet around the Sun is

$$\begin{aligned} \mathbf{L} &= \mathbf{r} \times \mathbf{p} \\ &= r\hat{\mathbf{i}}_r \times m(v_\theta \hat{\mathbf{i}}_\theta + v_r \hat{\mathbf{i}}_r) \\ &= m r v_\theta \hat{\mathbf{k}}, \quad \text{as } \hat{\mathbf{i}}_r \times \hat{\mathbf{i}}_\theta = \hat{\mathbf{k}} \text{ and } \hat{\mathbf{i}}_r \times \hat{\mathbf{i}}_r = 0 \\ &= \text{constant} \end{aligned}$$

E. Escape velocity

The inverse square law of gravity (Newton's law of gravity) is a conservative force with a potential

$$U_G(r) = -G\frac{m_1 m_2}{r} \tag{6}$$

The sum of this potential energy and the kinetic energy will be conserved for celestial objects. In order to escape the gravitational pull of the Sun, a planet in orbit must have a total energy $E = U_G + K \geq 0$ as the potential energy goes to zero infinitely far away:

$$\begin{aligned} E = U_G + K &\geq 0 \\ K &\geq -U_G \\ \frac{1}{2} M v^2 &\geq G\frac{M M_\odot}{r} \\ v &\geq \sqrt{G\frac{2M_\odot}{r}} \end{aligned}$$

A planet which begins at a distance 1 AU from the Sun, will therefore need an initial velocity of

$$\begin{aligned} v_0 &= \sqrt{G\frac{2M_\odot}{1\ \text{AU}}} \\ &= \sqrt{8\pi^2 \frac{\text{AU}^2}{\text{yr}^2}} \\ &= 2\sqrt{2}\pi \frac{\text{AU}}{\text{yr}} \end{aligned} \tag{7}$$

in order to be able to escape the gravitational pull of the Sun.

F. Energy and angular momentum

In this project we are going to be working with conservative forces. Therefore both energy and spin for our system should be conserved. When we are working with non-relativistic sizes, the total energy is given by the sum of kinetic and gravitational potential energy. Kinetic energy is given by

$$K = \frac{1}{2} m v^2, \tag{8}$$

where K is the kinetic energy, m is the mass and v is the velocity. The potential energy of an object in a gravitational field is given by

$$U = G \sum_i \frac{M m_i}{r_i} \tag{9}$$

Where U is the potential energy, G the gravitational constant, r_i the relative distance between the objects, m_i mass of an object creating the gravitational field, and M mass of the object we want to find the energy of. Notice that the zero-point of potential energy is set to infinity.

Another conserved size is angular momentum, and has the equation

$$\mathbf{L} = \mathbf{r} \times \mathbf{v}. \tag{10}$$

$\mathbf{L}$ is the angular momentum and $\mathbf{r}$ the position of an object with velocity $\mathbf{v}$. We can choose from what point we calculate the angular momentum ($\mathbf{r}$ depends on it), but in order for them to be conserved we have to be consistent.

G. Adjusting speed and position of the center of mass

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III. METHODS

A. Discretization of the equations of motion

Equation 3 is the equation of motion that describes the system of planets and stars in motion. We want to solve this equation numerically for all the planets in our solar system. In order to do this we must discretize the equations. We can rewrite the equation as a set of first order differential equations using the velocity:

$$\begin{aligned} \frac{\mathrm{d}\mathbf{r}}{\mathrm{d}t} &= \mathbf{v} \\ \frac{\mathrm{d}\mathbf{v}}{\mathrm{d}t} &= -G \sum_{j \neq i} m_j \frac{\mathbf{r}_i - \mathbf{r}_j}{|\mathbf{r}_i - \mathbf{r}_j|^3} \end{aligned}$$

The Verlet is not a self-starting algorithm. Mainly applicable when we are not interested in the velocity Velocity Verlet is energy conserving, Forward Euler is not

B. Units of measurement

Measuring distance in meters and time in seconds leads to very large numbers when doing calculations on the solar system. In order to do our calculations with numbers of magnitudes closer to 10^0, we will measure time in years and distance in astronomical units AU, which is the mean distance between the Sun and the Earth. For an object in circular motion with radius R and period T, the acceleration is given by the sentripetal acceleration

$$a = \frac{v^2}{r} = \frac{(2\pi R/T)^2}{R} = \frac{4\pi^2 R}{T^2}$$

With a small error, we can model the orbit of the Earth around the Sun as circular. To further simplify the expressions, we will measure masses in solar masses $M_\odot$. For the Earth's orbit we have $R = 1$ AU and $T = 1$ yr. This means that the force acting on the Earth from the Sun is

$$\begin{aligned} F_{G,\mathrm{Earth}} &= G\frac{M_{\mathrm{Earth}}M_\odot}{\mathrm{AU}^2} \\ &= M_{\mathrm{Earth}}a_{\mathrm{Earth}} \\ &\approx M_{\mathrm{Earth}}\frac{4\pi^2\mathrm{AU}}{\mathrm{yr}^2} \\ G &\approx 4\pi^2\frac{\mathrm{AU}^3}{M_\odot \mathrm{yr}^2} \end{aligned}$$

In units of AU, $M_\odot$ and yr the gravitational constant has a value of approximately 4π.

C. Testing the stability of the algorithms

In order to make sure our algorithms runs correctly, we want to test it. Our first test will be to simulate the simplest possible system, namely earth orbiting the sun. We only look at the gravitational force from the sun on earth. We place the sun in origin and with zero velocity. The earth has a radii of $r = 1$AU, and orbits the earth every year. This gives it an initial velocity of $v = 2\pi$ AU/yr. Testing for different time steps Δt we can look at the stability of velocity Verlet and Euler's forward method. In this system earth's distance from the sun and velocity should be constant, meaning both potential and kinetic energy ideally is conserved. Therefore we also test for conservation of energy.
In the Theory section, from Kepler's second law, we showed that angular momentum is conserved. Therefore we also tested whether that is the case for our Earth-Sun system. We did the tests for our ideal orbit over, and for the actual values collected from NASA (see [1]).

IV. RESULTS

V. DISCUSSION

discuss eventual differences between the verlet algorithm and the euler algorithm. Consider also the numver of flops involved and perform a timing of the two algorithm for equal final times.

VI. CONCLUSION

[1] ref til nasa dataen