

# We are not the 99 percent: quantifying asphericity in the distribution of Local Group satellites

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## ABSTRACT

We use numerical simulations to build an explicit probability distribution for the asphericity in the satellite distribution around galaxies similar to the Local Group (LG) in the Lambda Cold Dark Matter (LCDM) paradigm. This allows us to estimate the atypicality of the satellite distributions in the LG even when the underlying simulations do not have enough systems that fully resemble the LG. We demonstrate the method using three different simulations: Illustris-1, Illustris-1-Dark and ELVIS. Detailed results differ greatly among the simulations, suggesting a strong influence of resolution, simulated physics and criteria to select LG pairs. However, there are three common trends in these simulations. First, at most of 1% the pairs are expected to have satellite distributions with the same asphericity as the Local Group; second, between 38% to 59% of the pairs have a halo with a satellite distribution as aspherical as in M31; and third, at most 2% of the pairs have a satellite distribution as planar as in the MW. The brightest satellites in M31 galaxy turn out to have a rather typical distribution in the LCDM context, while the MW is atypical in this respect. These results place the LG at the level of a  $3\sigma$  outlier in the LCDM paradigm. The approach presented here facilitates the comparison between asphericity results from different numerical setups and choices to select halo pairs in simulations, hopefully helping to clarify the reasons behind the LG atypicality in LCDM.

**Key words:** Galaxies: halos — Galaxies: statistics — Dark Matter — Methods: numerical

## 1 INTRODUCTION

The spatial distribution of satellite galaxies around our Milky Way and the M31 galaxy provides a stringent test for structure formation theories in an explicit cosmological context. The presence of a Vast Polar Structure of satellites around the Milky Way has been established in the last decade. This structure is present in the 11 brightest “classical” dwarf galaxies and has been confirmed as fainter satellites have been detected. In the case of M31, observational studies have found a planar satellite distribution in a subset of 15 satellites out of an homogeneous sample of 27 satellites. However, taking into account the 11 brightest satellites (as a comparison with the MW) there does not seem to be any planar structure.

The question whether this kind of planar configurations are atypical in the observable Universe cannot be properly addressed given the current limits in observations. So far it has only been possible to constraint that 5% of the Milky-

Way like galaxies in the Universe have two bright satellite galaxies as is the case of the Milky Way and its Large and Small Magellanic Clouds.

However, the typicality of satellite distributions can be addressed in simulations. Most of the studies have been performed in the currently preferred paradigm, which is based on a Cold Dark Matter cosmology in an expanding Universe described by General Relativity including a Cosmological Constant, the so-called Lambda Cold Dark Matter (LCDM). The broad consensus is that these highly structured satellite distributions are hard to find in simulations. There are still disagreements on the precise degree of atypicality and the conditions required to simulate satellite distributions closer

Some studies have focused on the study of individual dark matter halos that could host galaxies like the MW or M31, while some other studies have tried to simulate the formation of galaxy pairs resembling both the MW and M31. The transition from studying individual halos to pairs has been motivated by the hypothesis that the location of the Local Group in the cosmic web should play a determinant role

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in building up planar satellite distributions through preferential alignments and accretions histories of galaxies.

Studies of the Local Group in an explicit cosmological context have also renewed the interest on performing constrained simulations that could reproduce the observed large scale structure of the Universe. This has also motivated the study of the two dominant galaxies in the Local Group as a pair in a cosmological context finding that the LG itself has relatively uncommon kinematic configuration in the context of LCDM and a strong preference to lie along filaments.

The aim of our work is to put together some of these pieces in a simple framework to quantify how atypical is the satellite distribution found in the Local Group. Some of the explicit characteristics we want to keep in our treatment of the problem are the (i) cosmological context, (ii) the LG as a pair of galaxies and (iii) a simple and robust description for the satellites.

The first condition constrains the kind of simulations we use. We either results from fully cosmological simulations or resimulations of cosmological sub-volumes. The second condition constraints the kind of samples we want to use. We build explicit samples of pairs that resemble the kinematic structure of the MW and M31. The third condition motivates to only use the inertia tensor as a description for the satellites positions, while dropping any information that might come from 3D velocities (high uncertainties) or involved algorithms for plane fitting.

Finally, to compute the odds of finding the LG satellite distributions in LCDM simulations, we do not make a direct search for the exact observed values into simulations because, as we show in Section, the LG is itself rare and we do not have enough simulated systems to perform such a brute force solution. We also avoid this direct comparison for a second reason. We do not want to implicitly use either LCDM or observations as what is considered normal or standard and compute the deviations from that point. Our approach is different. We set as a point of reference an spherical satellite distribution to quantify how both simulations and observations deviate from the spherical distribution. Simulations help us to construct an explicit probability distribution for deviations from sphericity in LCDM. Then we compute the odds of finding  $n$  satellite distributions that deviates from sphericity as much as the LG does.

In Section we list the sources of the observational and simulated data to be used throughout the paper. Next, in Section we describe the methods we use to quantify and characterize the satellite distributions. In section we present the results. In the discussion section we quantify the correlations between the main plane properties as described by the simulations. We use this results to quantify the degree of atypicality of the LG and estimate the volume that has to be probed in simulations in order to find a pair with a satellite distribution as atypical as the LG. Finally, we summarize our conclusions in Section .

## 2 DATA SAMPLES

### 2.1 Observational Data

We use three dimensional positions as reported by. The appendix includes the tables with the values we use in our computations. For each central galaxy we use the 15 brightest

satellites within a distance of 300kpc to the central galaxy. The satellites for the MW are: The satellites for M31 are:

### 2.2 Data from the Illustris project

We use publicly available data from the Illustris Project (Vogelsberger et al. 2014). This suite of cosmological simulations, performed using the quasi-Lagrangian code AREPO (Springel 2010), followed the coupled evolution of dark matter and gas and includes parametrizations to account for the effects of gas cooling, photoionization, star formation, stellar feedback, black hole and super massive black hole feedback. The simulation volume is a cubic box of  $75 \text{ Mpc } h^{-1}$  on a side. The cosmological parameters correspond to a  $\Lambda$ CDM cosmology consistent with WMAP-9 measurements (Hinshaw et al. 2013).

We extract halo and galaxy information from the Illustris-1 and Illustris-1-Dark simulations, the former simulation includes hydrodynamics and star formation prescriptions while the latter only includes dark matter physics. These simulations have the highest resolution in the current release of the Illustris Project. Illustris-1 has  $1820^3$  dark matter particles and  $1820^3$  initial gas volumen elements, while Illustris-1-Dark has  $1820^3$  dark matter particles. This corresponds to a dark matter particle mass of  $6.3 \times 10^6 M_\odot$  and a minimum mass for the baryonic volume element of  $8.0 \times 10^7 M_\odot$  for Illustris-1 and a dark matter particle mass of  $7.6 \times 10^6 M_\odot$  for Illustris-1-Dark. In both simulations the dark matter gravitational softening is 1.4 kpc.

We build a sample of pairs that resemble the conditions in the LG. To construct this sample we select from Illustris-1 all galaxies with an stellar mass in the range  $1 \times 10^{10} M_\odot < M_\star < 1.5 \times 10^{11} M_\odot$ . Then we select the pairs with the following conditions.

- For each galaxy  $A$  we find its closest galaxy  $B$ , if galaxy  $A$  is also the closest to halo  $B$ , the two are considered as a pair.
- With  $d_{AB}$  the distance between the two galaxies and  $M_{\star,min}$  the lowest stellar mass in the two galaxies, we discard pairs that have any other galaxy  $C$  with stellar mass  $M_\star > M_{\star,min}$  closer than  $3 \times d_{AB}$  from any of the pair's members.
- The distance  $d_{AB}$  is greater than 700 kpc.
- The relative radial velocity between the two galaxies, including the Hubble flow, is  $-120 \text{ km s}^{-1} < v_{AB,r} < 0 \text{ km s}^{-1}$ .

We find 27 pairs with these conditions. We then select the pairs where in both halos there are at least 15 detected subhalos, thus discarding pairs with halos with the lowest mass. We end up with a total of 20 pairs in Illustris-1. In Illustris-1-DM we use the center of mass position of the 27 pairs in Illustris-1 to find the matching halo pairs. After discarding the pairs with less than 15 detected subhalos in one of the halos we end up with a total of 24 pairs. This corresponds to a pair number density of  $\sim 2 \times 10^{-5}$  pairs  $\text{Mpc}^{-3}$ . Appendix A shows the physical properties (stellar masses, maximum circular velocities, radial velocities and separation) in those pairs.

Although Illustris-1 has stellar particles, we do not use their properties to select the satellite population because the

smallest galaxies are barely resolved in stellar mass at magnitudes of  $M_V = -9$ , close to the limit of the 11 “classical” MW satellites. We use instead the dark matter information as the smallest sub-halos we care about are sampled with at least 40 particles.

Both in Illustris-1 and Illustris-1-Dm we chose the satellite samples by ranking the subhalos in decreasing order of its current maximum circular velocity and select the first  $N_p$  halos in the list. The results presented here correspond to  $11 \leq N_p \leq 15$ .

### 2.3 Data from the ELVIS project

We use data from the public release of the Exploring the Local Universe In Simulations (ELVIS) project. For a detailed description of that project and its data we refer the reader to ?. Here we summarize the elements relevant to our discussion.

ELVIS data comes from resimulations of dark matter halo pairs selected in dark matter only cosmological simulations. The parent cosmological boxes have a cosmology consistent with the Wilkinson Microwave Anisotropy Probe 7 results.

The ELVIS project used the results from 50 simulation boxes of side length 70.4 Mpc to select pairs with kinematic characteristics similar to the LG. These selection criteria included the following

- The virial mass of each host must be in the range  $1 \times 10^{12} M_\odot < M_{vir} < 3 \times 10^{12} M_\odot$
- The total pair mass must be in the range  $2 \times 10^{12} M_\odot < M_{vir} < 5 \times 10^{12} M_\odot$
- The center of mass separation is in the range  $0.6 \leq d \leq 1$  Mpc.
- The relative radial velocity is negative.
- No halos more massive than the least massive halo within 2.8 Mpc and no halos with  $M_{vir} > 7 \times 10^{13}$  within 7 Mpc of the pairs’ center of mass.

This corresponds to a pair number density of  $\sim 8 \times 10^{-6}$  pairs  $\text{Mpc}^{-3}$ , this is a factor  $\sim 2.5$  lower than the pair number density we find in the Illustris data. There were a total of 146 pairs that met those criteria, but only 12 were chosen for resimulation. Additionally, the selected pairs for resimulation have a relative tangential velocity less than  $75 \text{ km s}^{-1}$ . The dark matter particle resolution in these resimulations is  $1.9 \times 10^5$ , a thousand times larger than Illustris-1. In this paper we only use the results from these 12 resimulated pairs.

Appendix A compares the physical properties (stellar masses, maximum circular velocities, radial velocities and separation) in those pairs against the results from the Illustris-1 simulations.

## 3 BUILDING, CHARACTERIZING AND COMPARING SATELLITES SPATIAL DISTRIBUTIONS

### 3.1 Building Satellite Samples

We compare the joint satellite distributions in the MW and M31 at fixed satellite number,  $N_s$ . This means that the magnitude cut corresponding to the faintest satellite included in

the sample is different in each case. We make this choice for two reasons. First, to be sure that there is a non-zero number of satellites in the simulations to make the computations. Second, to rule out the influence of satellite numbers in the statistics.

We compute the satellite statistics for 11 up to 15 satellites. The lowest bound corresponds to the number of classical Milky Way satellites. The upper limit corresponds to the maximum number of satellites that can be resolved in both halos for most of the isolated pairs in Illustris-1. In simulations we rank the subhalos by their maximum circular velocity, in observations we rank the satellites by its  $M_V$  magnitude.

We also use two kinds of satellite distributions. The first keeps the positions for the satellites fixed as provided in the observations/simulations; the second randomizes the angular positions of the satellites around the central galaxy while keeping its radial distance fixed. The randomization process is done 1000 times for each galaxy.

### 3.2 Describing Samples with the Inertia Tensor

We base all our results on the description provided by the inertia tensor defined by the satellites’s positions.

$$\bar{\mathbf{I}} = \sum_{k=1}^{N_s} [(\mathbf{r}_k - \mathbf{r}_0)^2 \cdot \mathbf{1} - (\mathbf{r}_k - \mathbf{r}_0) \cdot (\mathbf{r}_k - \mathbf{r}_0)^T], \quad (1)$$

where  $k$  indexes the set of satellites of interest  $\mathbf{r}_k$  are the satellites’ positions,  $\mathbf{r}_0$  is the location of the central galaxy  $\mathbf{1}$  is the unit matrix, and  $\mathbf{r}^T$  is the transposed vector  $\mathbf{r}$ . We use  $\mathbf{r}_0$  as the position of the central galaxy, and not the satellites’ geometrical center, to allow for a fair comparison once the angular positions of the satellites are randomized around this point.

From this tensor we compute its eigenvalues,  $\lambda_1 > \lambda_2 > \lambda_3$ , and corresponding eigenvectors,  $\hat{\mathbf{I}}_1, \hat{\mathbf{I}}_2, \hat{\mathbf{I}}_3$ . We define the size of the three ellipsoidal axis as  $a = \lambda_1$ ,  $b = \lambda_2$  and  $c = \lambda_3$ . We also define  $\hat{\mathbf{n}} \equiv \hat{\mathbf{I}}_1$  as the vector perpendicular to the planar satellite distribution. We also define the width,  $w$ , of the planar satellite distribution, as the standard deviation of all satellite distances to the plane defined by the vector  $\hat{\mathbf{n}}$ .

To summarize we characterize the satellite distribution by for quantities obtained from the inertia tensor:

- Plane width,  $w$ .
- $c/a$  axis ratio.
- $b/a$  axis ratio.

### 3.3 Comparing Satellite Samples

We compare every satellite distribution against its own spherically randomized distribution. We keep fixed the radial position of every satellite with respect to the central galaxy and then randomizing its angular position. We repeat this procedure 1000 times for each satellite distribution and proceed to measure the quantities mentioned in the last section:  $w$ ,  $c/a$  and  $b/a$ . For each quantity we compute an average and standard deviation from the 1000 random samples. This allows us to build a normalized version of all

	Observations		Randomized Obs.	
	M31	MW	M31	MW
Plane width (kpc)	$59 \pm 3$	$22 \pm 2$	$64 \pm 12$	$45 \pm 8$
$c/a$ ratio	$0.45 \pm 0.04$	$0.28 \pm 0.03$	$0.55 \pm 0.10$	$0.53 \pm 0.10$
$b/a$ ratio	$0.82 \pm 0.06$	$0.78 \pm 0.02$	$0.82 \pm 0.07$	$0.81 \pm 0.08$

**Table 1.** Results from observations. Mean values and standard deviations for the different quantities describing satellite distribution: plane width,  $c/a$  ratio and  $b/a$  ratio. The first column refers to the results from observational data, the second column uses the spherically randomized version of the observational data.

	Illustris-1-Dark		Illustris-1		ELVIS	
	M31	MW	M31	MW	M31	MW
Plane width (kpc)	$62 \pm 3$	$61 \pm 3$	$70 \pm 4$	$67 \pm 2$	$70 \pm 2$	$68 \pm 4$
$c/a$ ratio	$0.50 \pm 0.01$	$0.50 \pm 0.02$	$0.52 \pm 0.01$	$0.53 \pm 0.01$	$0.54 \pm 0.01$	$0.49 \pm 0.02$
$b/a$ ratio	$0.79 \pm 0.01$	$0.81 \pm 0.02$	$0.80 \pm 0.01$	$0.80 \pm 0.02$	$0.80 \pm 0.01$	$0.81 \pm 0.01$

**Table 2.** Results from simulations. Mean values and standard deviations for the different quantities describing satellite distributions: plane width,  $c/a$  ratio and  $b/a$  ratio. The first column summarize the results from the 24 pairs in Illustris-1-Dark, the second column from the 20 pairs in Illustris-1 and the third column from the 12 pairs in the ELVIS project.

quantities of interest by subtracting the mean and dividing between the standard deviation of the randomized samples.

This allows us to make a first comparison. The observed/simulated distribution against its randomized version. The second intermediate comparison we make are observations against simulations. The final comparison is between the normalized quantities, both observed and simulated.

The comparison between the normalized quantities is the one that carries the important information about the deviations from sphericity. We do not want to directly compare how the results observations deviate from simulations but to compare the deviations from asphericity in observations and simulations.

### 3.4 Describing joint satellite distributions

After building the normalized variables with the simulated data we perform a Kolmogorov-Smirnov test with the null hypothesis of belonging to a normal distribution with mean and standard deviation computed from the mean and standard deviation from the data itself. We find that the distributions for the normalized  $w$ ,  $c/a$  and  $b/a$  are consistent with gaussian distributions.

Based on this result we build a multivariate normal distribution for the joint distributions of the normalized  $w$ ,  $c/a$  and  $b/a$ :

$$p(X; \mu, \Sigma) = \frac{1}{(2\pi)^{3/2} |\Sigma|^{1/2}} \exp \left( -\frac{1}{2} (X - \mu)^T \Sigma^{-1} (X - \mu) \right), \quad (2)$$

where  $X = [w, c/a, b/a]^T$  is a vector variable with the normalized quantities,  $\mu$  is the vector mean and the  $\Sigma$  is the covariance matrix.

We compute the preferred covariance matrix and the mean distribution values with a jackknife technique. That is, out of the  $n$  pairs we have, we perform  $n - 1$  pairs. The reported covariance and mean values correspond to the average of all measurements, the corresponding standard deviation also help us to estimate the uncertainty on every reported coefficient. This compact description allows us to generate samples of size  $N$  that are consistent by construction with their parent simulation.

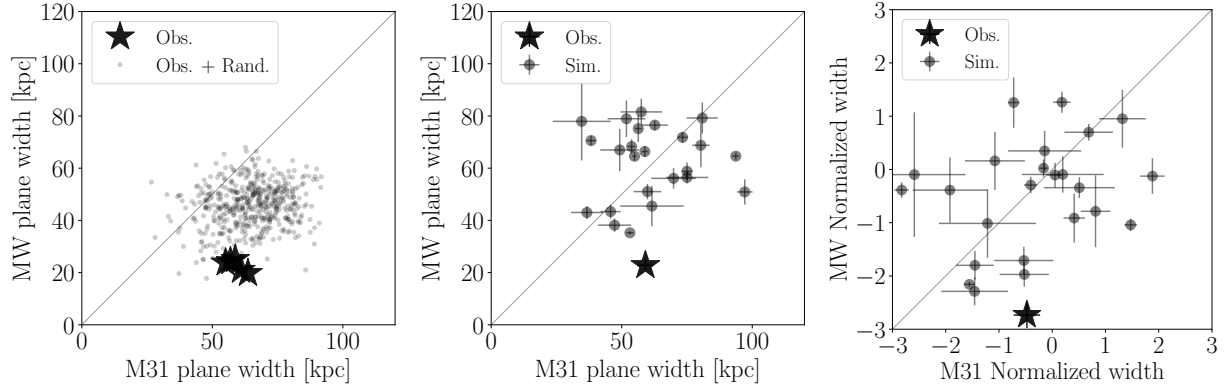
Finally, we use the generated samples to estimate how common are the deviations from sphericity that we measure in the observational data. We use a double-tailed test in this comparison, meaning that we always measure the fraction of points with absolute values larger than the threshold absolute observed value.

## 4 RESULTS

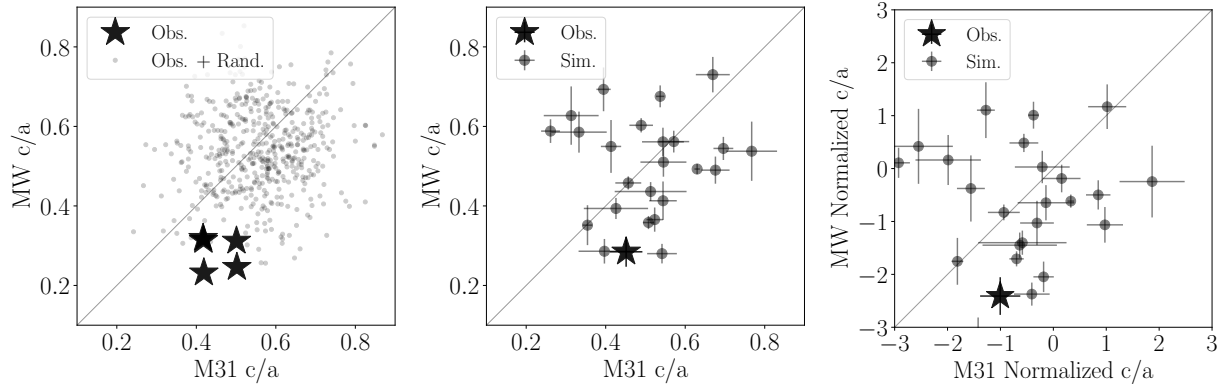
Table 1 and Table 2 summarize the mean values and uncertainties for the plane  $w$ ,  $c/a$  ratio and  $b/a$  ratio in the observations and simulations, respectively. The uncertainty in the observations is computed over the results with different number of satellites, while in the simulations it corresponds to the standard deviation over different halos.

The observed widths are always smaller than its randomized version. For M31 there is barely a ratio of 0.92 between observed and randomized, while for the MW this factor goes down to less than half 0.48. The observed  $c/a$  ratio follows the same extreme trend for the MW compared to a M31 distribution closer to spherical, The  $b/a$  ratio is statistically the same between observations and the spherical randomization.

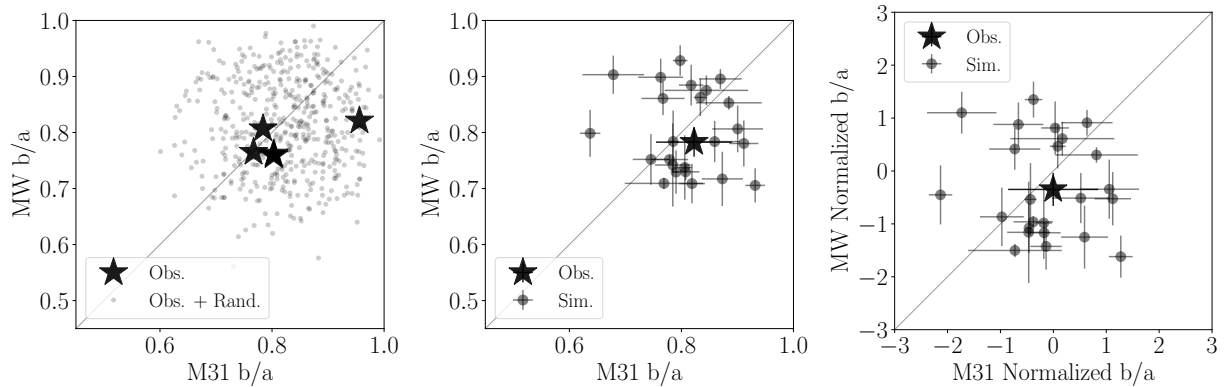
In the following subsections we describe in detail the results on the distributions for  $w$ ,  $c/a$  and  $b/a$ . The plots



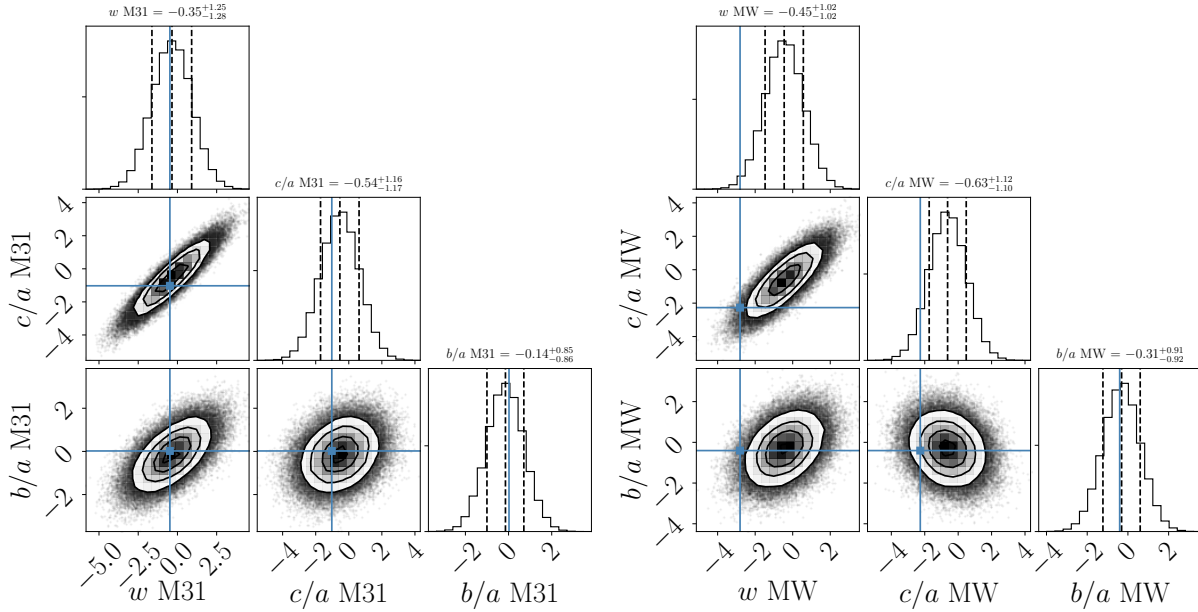
**Figure 1.** Plane width characterization in observations and simulations (Illustris1-Dark in this case.) In all panels the horizontal/vertical axis corresponds to the M31/MW or the most/less massive halo in the pair. Left. Plane width in physical units comparing the results from observations (stars) against the result of spherically randomizing the satellite positions (circles). Middle. Average from the observations (star) and the average from each pair in the simulation (circles). Right. Same as the middle panel only that this time each point has been normalized (median subtracted and normalized by the standard deviation) to the results of its randomization. The main message of these panels is that the MW has a significantly thinner plane both compared to the result of its own satellite spherical randomization (left panel) and the expectation from simulations (middle and right panel). This low value is  $2\sigma$  away from what is expected in a spherical distribution. In M31 the satellite distribution is in agreement with the expectations both from an spherical distribution and the results from the simulations.



**Figure 2.** Same layout as in Figure 1. This time for the  $c/a$  axis ratio. The same message holds in this case; namely, that the MW is has a significantly low  $c/a$  value than the expectation from the spherical distribution and the simulations. This low value is also  $2\sigma$  away from the expectations for an spherical distribution. On the other hand, M31 is consistent both with an spherical distribution and the results from simulations.



**Figure 3.** Same layout as in Figure 1. This time for the  $b/a$  axis ratio. In this case both the MW and M31 are consistent with the results of a spherical distribution and simulations.



**Figure 4.** Results from multivariate gaussian model fitted on the normalized values for the plane width  $w$ ,  $c/a$  ratio and  $b/a$  ratio from the Illustris1-Dark results. Left/right panel correspond to M31/MW. The cross indicates the observed values. The contour levels in the 2D histograms correspond to the  $1\sigma$ ,  $2\sigma$  and  $3\sigma$  contours in two dimensions. The dashed vertical lines in the histograms along the diagonal correspond to the  $1\sigma$  boundaries in one dimension. The results for the gaussian model are built from  $10^6$  point realizations in the three-dimensional space spanned by the variables of interest for each halo. This plot clearly shows how the M31 results are well within the expectations from simulations while the MW has an unusual low value for the plane width and the  $c/a$  axis ratio.

in the main body of the paper correspond to the Illustris1-Dark simulation; the results from the other simulations are included in the Appendix.

#### 4.1 Plane Width

Figure 1 summarizes the results for the width measurements. The left panel compares the results for the MW and M31 observations (stars) against its spherically randomized satellites (circles). The most interesting outcome is that the MW plane width is smaller than  $\approx 98\%$  of the planes computed from the randomized distribution, while the M31 plane width is only slightly smaller than the average of the distribution.

The middle panel in Figure 1 compares the observational result (star) against the measurements of all pairs from the Illustris1-Dark simulation (circles). In this case we have a similar result as before. The observed MW width is smaller than all the results in the simulation, there is not a single halo with similar values. On the other hand, the results for the M31 are entirely consistent with observations. Most of the halos in the simulation show a width value similar to M31.

The right panel in Figure 1 shows the result for the normalized width. This panel tells the same story as the middle panel. The M31 values are typical while the MW is an outlier. The added value of the data in this panel is that it is the normalized which is consistent with normal distributions. This is the data used to build the mean values vector and covariance matrix described in Equation

#### 4.2 $c/a$ axis ratio

Figure 2 shows the results for the minor to major axis ratio. The layout is the same as in Figure 1. The results for the  $c/a$  ratio follow the same trends as for the width  $w$ .

The left panel in Figure 2 shows how the MW the  $c/a$  ratio is significantly lower as the measured values for spherical distributions, it is smaller than  $\approx 98\%$  of the randomized distributions. On the other hand the ratio for M31 is lower than the mean of the spherical values but still well within its variance.

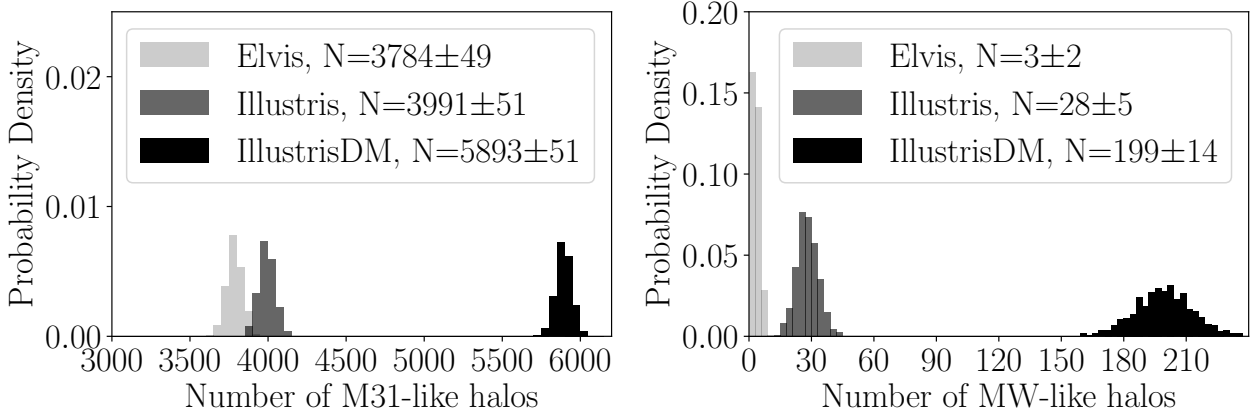
The middle panel in the same Figure shows the LG compared against the results in the simulations. In this case we find a similar trend as before. The MW is atypical and M31 is within the variance from the simulation data. This time, however, there are two MW-like halos out of the total of 24 that show an  $c/a$  as small as the MW.

The right panel shows the normalized results. The MW shows a low  $c/a$  ratio close to between two and three standard deviations away from the mean value of the spherical distribution; this contrasts the results for M31 which are close to 1 standard deviation away.

#### 4.3 $b/a$ axis ratio

Figure 3 shows the results for the minor to major axis ratio with the same layout as Figure ?? In all cases of comparison (against randomized distribution and simulations the results for both the MW and M31 are typical.





**Figure 5.** Probability distribution for the expected number of M31 and MW halos showing the same degree of atypicality as the Local Group if drawn from a sample of 10000 isolated pairs. The distributions correspond to results derived from Illustris-DM, Illustris, and ELVIS data. In the most optimistic case (Illustris-DM), 1% of the pairs have two halos with properties similar to the LG. In the case of Illustris this percentage drops to 0.1% and 0.01% for ELVIS.

#### 4.4 Fit to a Multivariate Gaussian Distributions

Figure 4 illustrates the results from compute the covariance matrix and mean vector in Eq.2 from the normalized quantities computed from the Illustris1-Dark simulation. The distributions in the Figure are computed from  $10^6$  with the multivariate gaussian. Similar plots for Illustris1 and ELVIS are in the Appendix. The values for all the covariance matrices and mean vectors for all the simulations are listed in the Appendix.

This results nicely summarizes the results we had in the previous sections. The left triangular plot shows how M31 falls into the middle of all 2D distributions always close to the peak and within the  $1 - \sigma$  range. The right plot clearly places the MW observations outside the  $3\sigma$  range in the joint distributions that involve the width  $w$ .

In both cases the strongest positive correlation is present for the width and the  $c/a$  axis ratio. A weaker but still positive correlation is present for the width and the  $b/a$  axis ratio.

#### 4.5 Number of Expected LG Systems

We use the fits to the multivariate gaussian distributions to compute the expected number of pairs with characteristics similar to the LG. To do this we generate  $10^3$  samples, each sample contains  $10^4$  pairs, where each pair member is drawn from the corresponding multivariate gaussian distributions.

We consider that a sampled system is similar to the M31/MW galaxy if the distance of each of its characteristics ( $w$ ,  $c/a$ ,  $b/a$ ) to the sample mean is equal or larger than the distance of the observational values to the sample mean. That is, we perform a double-tailed test using the observational values as a threshold.

Figure 5 summarizes the results from this experiment. The left panel shows the probability density for the number of M31 systems in a parent sample of  $10^4$  pairs, the right panel shows the results for the MW. For M31, between 37% to 58% of the pairs have a satellite distribution as aspherical as M31 in observations; this fraction drops dramatically for

the MW where only 0.03% to 2% of the satellites are expected to have extreme aspherical distributions as the MW.

Considering the joint distribution of M31 and MW we find that at most 1% of the pairs are expected to be similar to the LG. In a three dimensional gaussian distribution, having a  $1\sigma$ ,  $2\sigma$  and  $3\sigma$  interval corresponds to have respectively 19%, 73% and 97% of the total of points in the distribution. With this result in mind the LG has the same degree of atypicality as an  $3\sigma$  outlier.

Among the three simulations, the results inferred from ELVIS data show the lowest fraction of M31 and MW systems; for Illustris1-Dark we have the highest fraction of M31/MW systems. The results from Illustris1 are in between these two, but closer to ELVIS.

The most probable reason for these trends is the different median mass for the MW/M31 halos in the pairs from these simulations. For instance for the MW halo the median maximum circular velocity is  $\sim 160 \text{ km s}^{-1}$ ,  $\sim 150 \text{ km s}^{-1}$  and  $\sim 120 \text{ km s}^{-1}$  in the ELVIS, Illustris1 and Illustris1-Dark simulations, respectively. For the M31 halo this median velocity is  $\sim 200 \text{ km s}^{-1}$  both for the ELVIS and Illustris1 simulations, while for the Illustris1-Dark it is  $\sim 160 \text{ km s}^{-1}$ .

## 5 CONCLUSIONS

In this paper we develop and demonstrate a method to quantify the asphericity of the satellite distribution in the Local Group. We build up on the suggestions made in ?. We do not look for planes or coherent kinematical structures in the interest of keeping the method straightforward and robust.

The method uses as a reference point the spherically randomized data of the system under study. We first measure the width and axis ratios for the satellite distribution of interest. Then, we measure the same quantities for spherically randomized data. Finally, renormalize the initial results to the mean value and standard deviation from the randomized samples.

The normalized quantities for different systems are well reproduced by a multivariate gaussian distribution. We fit the parameters of this gaussian to the results from LG sam-

ples coming from three different simulations to finally compare the observational results against the distributions derived from the simulations.

We find that in the best case (Illustris1-DarkMatter) the deviation from asphericity in the observed LG is only expected in  $3 \pm 2$  pairs out of a sample of  $10^4$  isolated pairs. This places the LG as a  $3\sigma$  outlier. The weight to explain this atypical result is not distributed equally between the MW and M31. While M31 presents a fully typical asphericity in the expectations from LCDM, the MW shows aspherical deviations in plane width and the major-to-minor axis ratio highly atypical in the framework of LCDM. We estimate that with the M31 between 37% and 58% of the pairs pairs show normalized characteristics larger than M31, while this fraction drops to less than 1% for the MW. These fractions are robust to changes in the numerical simulations and to the criteria used to define the pairs.

The highly aspherical distribution of the brightest MW satellites is in stark contrast to the more spherical M31 distribution. Furthermore, the atypicality of the MW distribution in the LCDM context adds up to the also atypical presence of two bright satellite galaxies.

The focus of our approach is building explicit probability distributions for the observables of interest, instead of trying to find simulated objects that fulfill different observational criteria. This approach is particularly useful in the case of atypical observables. Building the probability distribution allows for an atypicality quantification without needing to build explicit samples of objects that are already scarce and difficult to find in simulations.

An extension of this framework to outliers in higher order deviations (i.e. a *subset* of satellites that are in a plane or coherent *velocity* structures) should also be possible, similar to the work presented by XX on the satellite velocity anisotropy.

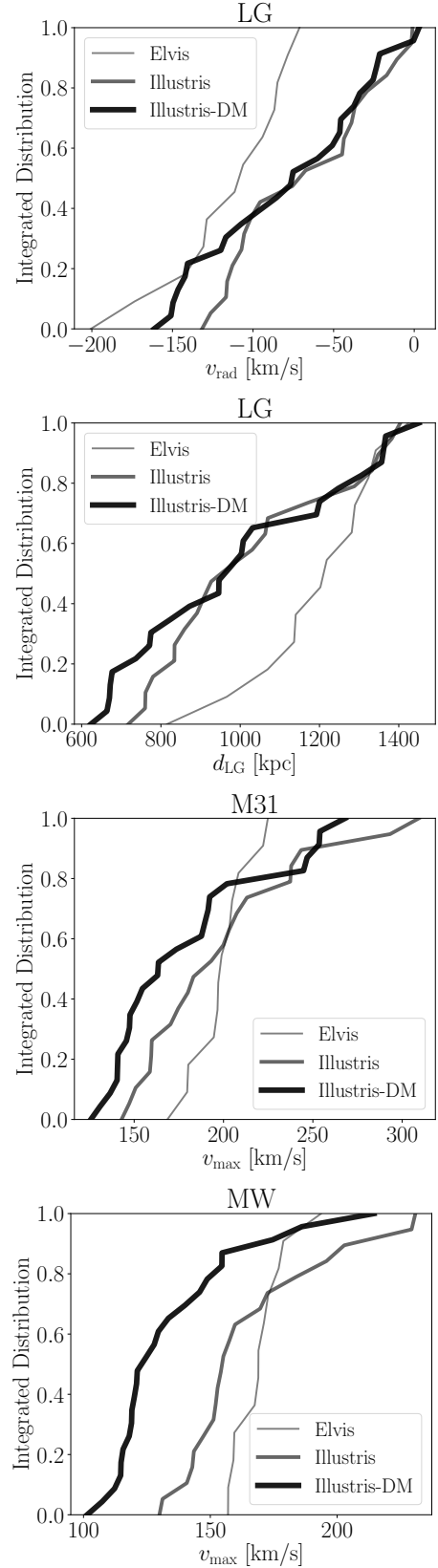
From the results we derive in this paper we also estimate that a cosmological volume of at least  $(XX)^3$  is required to find a halo pair that shows similar asphericity characteristics as the LG. This atypical satellite distribution should be seen as an opportunity to constrain in high detail the initial conditions and environment that allowed such a pattern to emerge.

## REFERENCES

- Hinshaw G., et al., 2013, [ApJS](#), **208**, 19  
 Springel V., 2010, [MNRAS](#), **401**, 791  
 Vogelsberger M., et al., 2014, [MNRAS](#), **444**, 1518

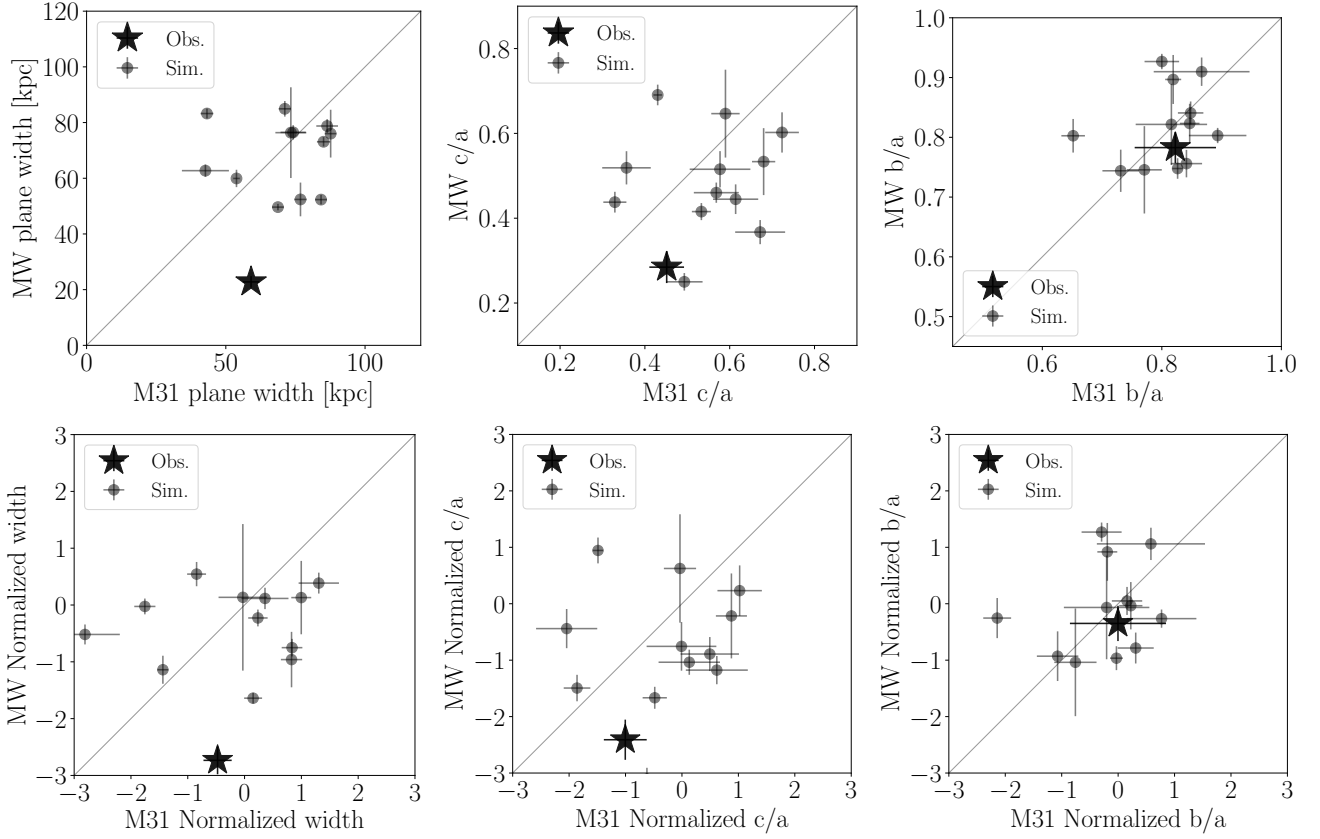
## APPENDIX A: PHYSICAL CHARACTERISTICS OF THE ISOLATED PAIRS SAMPLES

### APPENDIX B: RESULTS FROM ELVIS AND ILLUSTRIS1

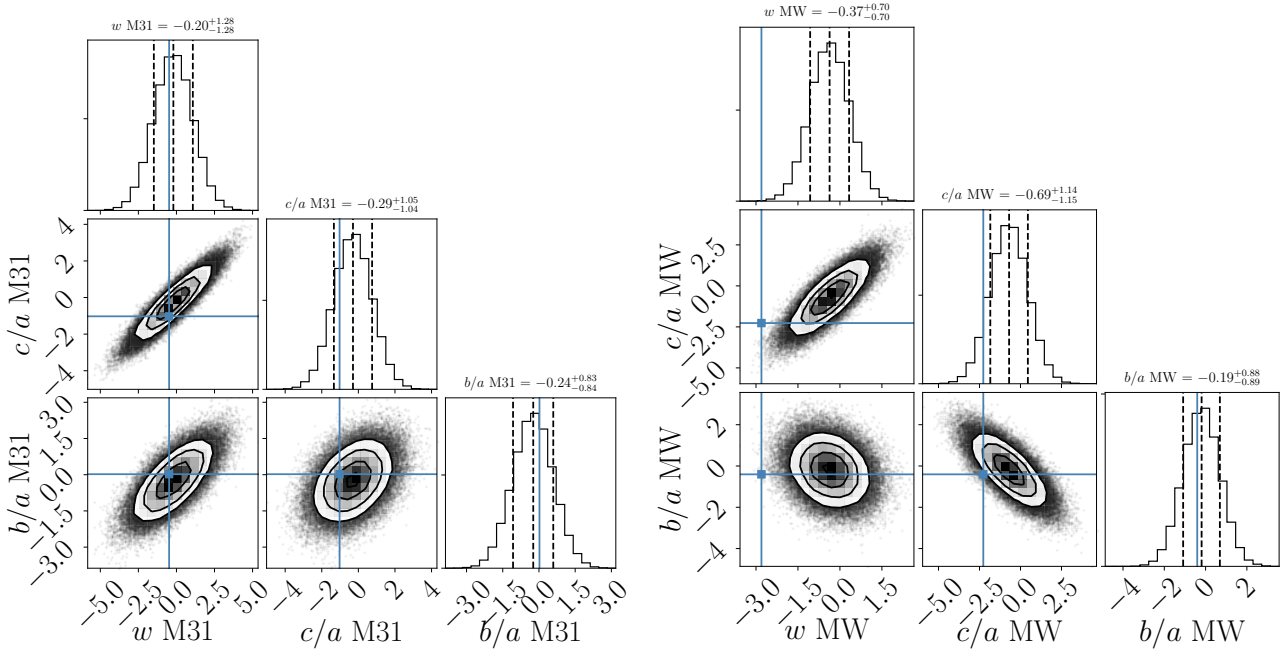


**Figure A1.** Physical characteristics of the LG pairs selected in the simulations. All lots show the integrated distributions. The physical properties are the radial comoving velocity between the MW and M31, the radial separation between the MW and M31 and the maximum circular velocity for the M31 and MW dark matter halo.

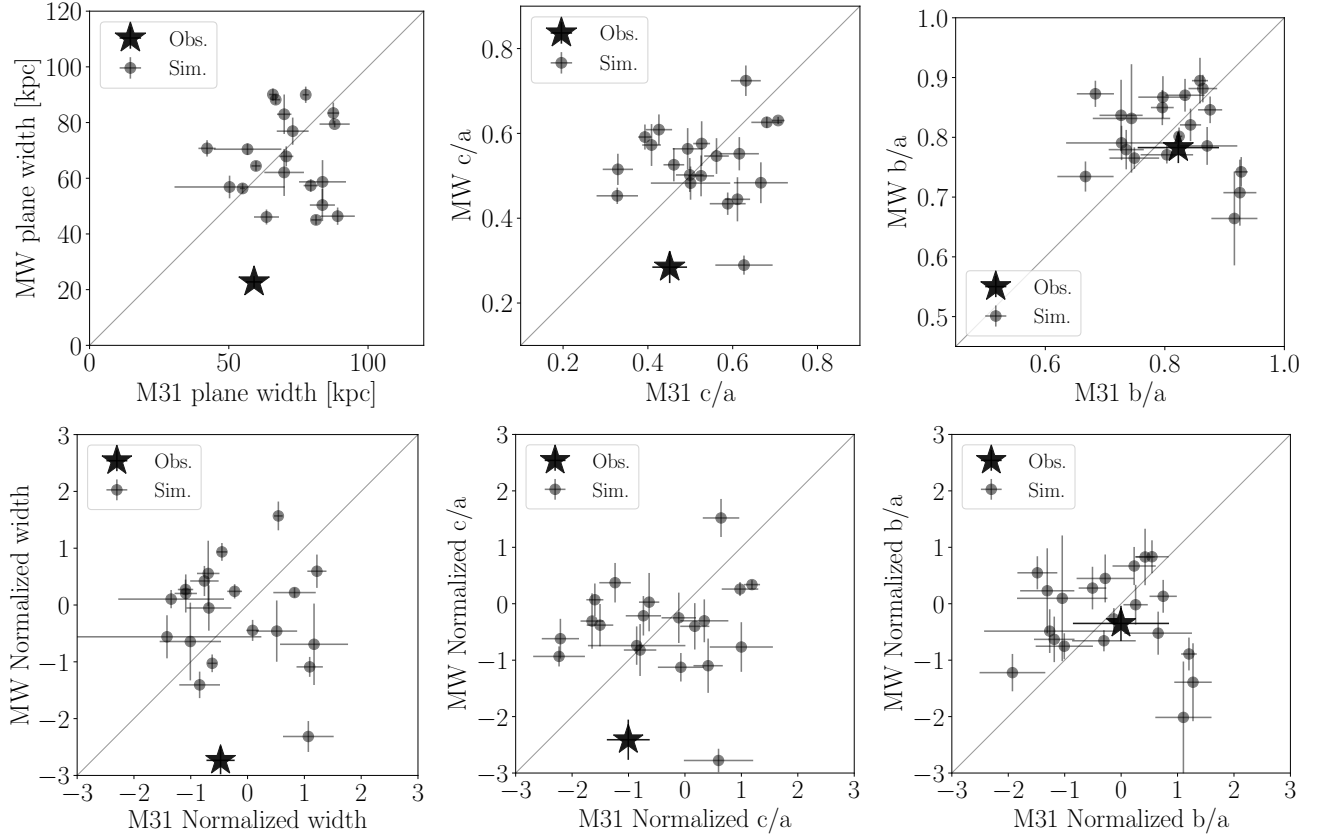




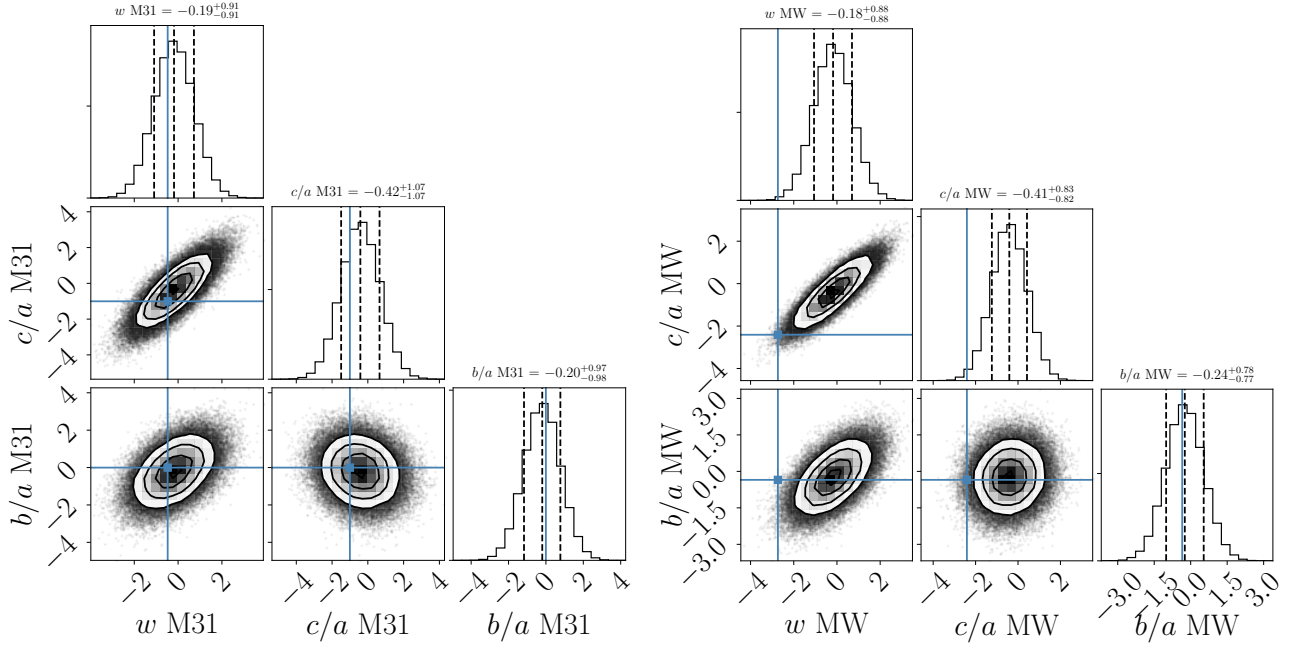
**Figure B1.** ELVIs results for the quantities presented for the Illustris1-Dark simulation in Figures 1, 2, 3. Upper row corresponds to the raw values from observations and simulated pairs, while the second row normalizes the same values to the mean and standard deviation on its spherically randomized counterparts.



**Figure B2.** Same layout as Figure 4, this time computed from the ELVIS data.



**Figure B3.** Same layout as Figure B1, this time computed from the Illustris1 data.



**Figure B4.** Same layout as Figure 4, this time computed from the Illustris1 data.

## APPENDIX C: COVARIANCE MATRICES AND MEAN VALUE VECTORS

### C1 Illustris-1

#### C1.1 M31

$$\Sigma = \begin{bmatrix} 0.83 \pm 0.03 & 0.78 \pm 0.04 & 0.41 \pm 0.04 \\ 0.78 \pm 0.04 & 1.16 \pm 0.05 & -0.15 \pm 0.05 \\ 0.41 \pm 0.04 & -0.15 \pm 0.05 & 0.96 \pm 0.04 \end{bmatrix}$$

$$\mu = [-0.18 \pm 0.04 \quad -0.41 \pm 0.05 \quad -0.20 \pm 0.05]$$

#### C1.2 MW

$$\Sigma = \begin{bmatrix} 0.78 \pm 0.06 & 0.63 \pm 0.07 & 0.40 \pm 0.03 \\ 0.63 \pm 0.07 & 0.69 \pm 0.08 & 0.06 \pm 0.02 \\ 0.40 \pm 0.03 & 0.06 \pm 0.02 & 0.61 \pm 0.04 \end{bmatrix}$$

$$\mu = [-0.17 \pm 0.04 \quad -0.40 \pm 0.04 \quad -0.23 \pm 0.04]$$

### C2 Illustris-1-Dark

#### C2.1 M31

$$\Sigma = \begin{bmatrix} 1.50 \pm 0.08 & 1.27 \pm 0.08 & 0.64 \pm 0.04 \\ 1.27 \pm 0.08 & 1.31 \pm 0.08 & 0.28 \pm 0.04 \\ 0.64 \pm 0.04 & 0.28 \pm 0.04 & 0.70 \pm 0.04 \end{bmatrix}$$

$$\mu = [-0.37 \pm 0.05 \quad -0.55 \pm 0.05 \quad -0.13 \pm 0.03]$$

#### C2.2 MW

$$\Sigma = \begin{bmatrix} 1.03 \pm 0.05 & 0.94 \pm 0.05 & 0.36 \pm 0.03 \\ 0.94 \pm 0.05 & 1.24 \pm 0.07 & -0.18 \pm 0.04 \\ 0.36 \pm 0.03 & -0.18 \pm 0.04 & 0.86 \pm 0.03 \end{bmatrix}$$

$$\mu = [-0.45 \pm 0.04 \quad -0.61 \pm 0.05 \quad -0.31 \pm 0.04]$$

### C3 ELVIS

#### C3.1 M31

$$\Sigma = \begin{bmatrix} 1.60 \pm 0.19 & 1.22 \pm 0.14 & 0.70 \pm 0.09 \\ 1.22 \pm 0.14 & 1.08 \pm 0.10 & 0.35 \pm 0.05 \\ 0.70 \pm 0.09 & 0.35 \pm 0.05 & 0.63 \pm 0.10 \end{bmatrix}$$

$$\mu = [-0.18 \pm 0.08 \quad -0.28 \pm 0.07 \quad -0.22 \pm 0.05]$$

#### C3.2 MW

$$\Sigma = \begin{bmatrix} 0.45 \pm 0.04 & 0.61 \pm 0.09 & -0.11 \pm 0.09 \\ 0.61 \pm 0.09 & 1.21 \pm 0.17 & -0.63 \pm 0.09 \\ -0.11 \pm 0.05 & -0.63 \pm 0.09 & 0.64 \pm 0.06 \end{bmatrix}$$

$$\mu = [-0.32 \pm 0.04 \quad -0.74 \pm 0.07 \quad -0.08 \pm 0.05]$$