

Fast lognormal realizations for multi-probe experiments

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1 Notation

χ is comoving radial density. We use units $c = 1$ throughout.

2 Algorithm

- Generate a Gaussian realization of the linear density field δ_G and the newtonian potential ϕ at $z = 0$. For this we use FFTW in a box able to hold a sufficiently large sphere. In what follows, let Δx be the comoving resolution of this Cartesian grid.

- Interpolate the Cartesian grid into spherical shells. These are generated with a width $\Delta r = f_r \Delta x$ (with $f_r = 1$ by default). The pixels are defined using a CAR scheme (equirectangular projection), with constant intervals in ϕ and θ .

Let N_θ be the number of divisions used in θ . Firstly, we will use $2 \times N_\theta$ divisions in φ . Secondly, the largest transverse scale covered by a pixel is given by $r_\perp^{\max} = \chi \pi / N_\theta$, where χ is the comoving distance to the shell. We determine N_θ as the smallest power of 2 that satisfies $r_\perp^{\max} \leq f_\theta \Delta x$ (with $f_\theta = 1$ by default).

- While doing the interpolation we compute the relevant quantities at $z = 0$ that will be used later. These are:
 - The radial velocity v_r , proportional to the radial gradient of ϕ . This is always computed, and is needed for RSDs.
 - The second derivatives of ϕ in the transverse directions. This is only computed if any lensing quantity (galaxy shear or convergence maps) are required.
 - The time derivative of ϕ , $\dot{\phi}$. This is only computed if ISW maps are required.

All radial and transverse derivatives are estimated by first computing the gradient or Hessian in Cartesian coordinates using finite differences in the box and then rotating into spherical coordinates.

- The interpolation is carried out in three steps:
 - First, each spherical voxel is sub-divided into $N_r^{\text{sub}} \times (N_\theta^{\text{sub}})^2$ sub-voxels, where N_r is the number of divisions taken in the radial direction and N_θ is the number of divisions in the two transverse directions.
 - Each sub-voxel is then assigned a value of the corresponding field (δ_G , v_r , $\nabla_\perp^2 \phi$ or $\dot{\phi}$) by using tri-linear interpolation on the sub-voxel centre. Specifically, let (x, y, z) be comoving Cartesian coordinates of the sub-voxel, and let (i, j, k) denote the cell in the Cartesian grid such that

$$x_i \leq x < x_{i+1}, \quad x_j \leq y < x_{j+1}, \quad x_k \leq z < x_{k+1}, \quad (1)$$

where $x_i \equiv i \Delta x$. Let $f_{i,j,k}$ be the value of the corresponding field in the Cartesian grid denoted by (i, j, k) , and let $h_x = (x - x_i) / \Delta x$ etc.. Then, the field value assigned to the

sub-voxel is given by:

$$f(x, y, z) = f_{i,j,k}(1-h_x)(1-h_y)(1-h_z) + f_{i,j,k+1}(1-h_x)(1-h_y)h_z + \quad (2)$$

$$f_{i,j+1,k}(1-h_x)h_y(1-h_z) + f_{i,j+1,k+1}(1-h_x)h_yh_z + \quad (3)$$

$$f_{i+1,j,k}h_x(1-h_y)(1-h_z) + f_{i+1,j,k+1}h_x(1-h_y)h_z + \quad (4)$$

$$f_{i+1,j+1,k}h_xh_y(1-h_z) + f_{i+1,j+1,k+1}h_xh_yh_z \quad (5)$$

– The field value assigned to the voxel is then computed as the average of the different sub-voxels.

- The different fields are evolved according to their linear perturbation growth factors and integrated along the line of sight in the case of lensing and ISW (see details below).
- The evolved field values are used to Poisson-sample sources, give them RSDs and shapes, produce convergence, ISW and intensity maps.

3 Relevant equations

Given the Cartesian realization of the Gaussian density field δ_G at $z = 0$, we compute the Newtonian potential by solving Poisson's equation in Fourier space:

$$\phi(\mathbf{k}, z = 0) = -\frac{3}{2}\Omega_M H_0^2 \frac{\delta_G(\mathbf{k}, z = 0)}{k^2}, \quad (6)$$

The radial velocity field is computed in terms of the gradient of the Newtonian potential as

$$v_r(z = 0) = -\frac{2f_0}{3H_0^2\Omega_M}(\hat{\mathbf{n}} \cdot \nabla)\phi(z = 0), \quad (7)$$

where $\hat{\mathbf{n}} \equiv (\sin\theta \cos\varphi, \sin\theta \sin\varphi, \cos\theta)$.

The transverse hessian of the gravitational potential is computed as

$$\hat{H}_\perp \phi = \hat{R}(\hat{H}\phi)\hat{R}^T \quad (8)$$

where \hat{H} is the Cartesian Hessian operator, and \hat{R} is a rotation matrix:

$$\hat{H} \equiv \begin{pmatrix} \partial_x \partial_x & \partial_x \partial_y & \partial_x \partial_z \\ \partial_x \partial_y & \partial_y \partial_y & \partial_y \partial_z \\ \partial_x \partial_z & \partial_y \partial_z & \partial_z \partial_z \end{pmatrix}, \quad \hat{R} \equiv \begin{pmatrix} \cos\theta \cos\phi & \cos\theta \sin\phi & -\sin\theta \\ -\sin\phi & \cos\phi & 0 \end{pmatrix}. \quad (9)$$

The evolution equations for the different quantities are:

$$\delta_G(z) = D(z)\delta(z = 0), \quad v_r(z) = \frac{D(z)f(z)H(z)}{f_0H_0}v_r(z = 0), \quad (10)$$

$$\phi(z) = (1+z)D(z)\phi(z = 0), \quad \dot{\phi}(z) = (1+z)D(z)H(z)[f(z) - 1]\phi(z = 0) \quad (11)$$

Finally, we obtain the shear (γ_1, γ_2) , convergence κ and ISW Δ^{ISW} fields after integrating along the line of sight:

$$\begin{pmatrix} \kappa + \gamma_1 & \gamma_2 \\ \gamma_2 & \kappa - \gamma_1 \end{pmatrix}(\chi) = 2 \int_0^\chi d\chi' \chi' \left(1 - \frac{\chi'}{\chi}\right) \hat{H}_\perp \phi(\chi') \quad (12)$$

$$\Delta^{\text{ISW}}(\chi) = 2 \int_0^\chi d\chi' a(\chi') \dot{\phi}(\chi') \quad (13)$$

4 Sources

At each voxel, we compute the physical galaxy density for each galaxy type a as

$$n_a(\chi, \hat{\mathbf{n}}) = \bar{n}_a(\chi) \exp \left[D(\chi) b_a(\chi) (\delta_G(\chi, \hat{\mathbf{n}}) - D(\chi) b_a(\chi) \sigma_G^2/2) \right], \quad (14)$$

where \bar{n}_a is the mean number density of galaxies, and σ_G^2 is the variance of the density field at $z = 0$. \bar{n} is related to $dn/(d\Omega dz)$ as $dn/(d\Omega dz) = \bar{n} \chi^2/H$.

Then, at each pixel we sample a number of galaxies from a Poisson distribution with mean $\lambda \equiv V_{\text{pix}} n_a$, where V_{pix} is the comoving volume of each spherical voxel. We then place the resulting number of particles inside each voxel at random within it. Each galaxy is given a cosmological redshift corresponding z_c corresponding to its comoving distance by inverting

$$\chi = \int_0^{z_c} \frac{dz}{H(z)}. \quad (15)$$

In addition, each galaxy is given a redshift distortion $\Delta z = v_r$ according to the value of the comoving velocity field in their corresponding voxel.

Each source is also given two ellipticity components based on the value of the local shear field:

$$\epsilon_1 = 2\gamma_1, \quad \epsilon_2 = 2\gamma_2. \quad (16)$$

It is also possible compute the fully non-linear values of the ellipticities, given by

$$\epsilon_i = 2 \frac{1 - \kappa}{(1 - \kappa)^2 + \gamma_1^2 + \gamma_2^2} \gamma_i. \quad (17)$$

5 Intensity mapping

The brightness temperature for a line-emitting species a in a voxel i is

$$T_a(\nu, \hat{\mathbf{n}}) = \bar{T}_a(\nu) [1 + \Delta_i^a(\chi \hat{\mathbf{n}})]. \quad (18)$$

where the mean brightness temperature is

$$\bar{T}_a(z) = \frac{3\hbar A_{21} x_2 c^2}{32\pi G k_B m_a \nu_{21}^2} \frac{H_0^2 \Omega_a(z) (1+z)^2}{H(z)} \quad (19)$$

and Δ_i^a is the redshift-space overdensity of the line-emitting species smoothed over the voxel. Here x_2 is the fraction of atoms in the excited state, Ω_a is the fractional density of the species, ν_{21} is the rest-frame frequency of the transition and A_{21} is the corresponding Einstein coefficient for emission.

The procedure to generate intensity maps is:

- We cycle over each voxel in the spherical shells for which we have stored the value of the density and velocity fields.
- We compute the overdensity in the voxel using a log-normal model:

$$1 + \delta^a = \exp \left[D(\chi) b_a(\chi) (\delta_G(\chi, \hat{\mathbf{n}}) - D(\chi) b_a(\chi) \sigma_G^2/2) \right]. \quad (20)$$

- We sub-sample the voxel in N_{sub} random points. Each point is assigned a brightness temperature

$$T_{a,\text{sub}} = \bar{T}_a(\nu) (1 + \delta_i^a) \frac{v_{\text{vox}}}{N_{\text{sub}} v_{\text{IMAP}}}, \quad (21)$$

where v_{vox} is the comoving volume of the voxel and v_{IMAP} is the comoving volume of the intensity mapping pixel. Each point is also assigned a redshift displacement given by v_r , the value of the lightcone-evolved radial velocity field in the voxel.

- We compute the frequency channel corresponding to each point from their redshift (computed as the sum of its cosmological redshift and the RSD term), as well as the pixel index in that frequency channel corresponding to the angular coordinates of the point. We then add the brightness temperature of this point computed in the previous step to the total brightness temperature in the pixel.
- Each intensity mapping pixel is finally divided by the total comoving volume covered by the pixel.

6 Shear

We compute the shear tensor as

$$\hat{T} \equiv \begin{pmatrix} \phi_{xx} & \phi_{xy} & \phi_{xz} \\ \phi_{yx} & \phi_{yy} & \phi_{yz} \\ \phi_{zx} & \phi_{zy} & \phi_{zz} \end{pmatrix}, \quad \hat{R} \equiv \begin{pmatrix} \sin \theta \cos \phi & \sin \theta \sin \phi & \cos \theta \\ \cos \theta \cos \phi & \cos \theta \sin \phi & -\sin \theta \\ -\sin \phi & \cos \phi & 0 \end{pmatrix}, \quad (22)$$

$$\hat{t} \equiv \hat{R} \cdot \hat{T} \cdot \hat{R}^T, \quad \hat{\tau} \equiv \begin{pmatrix} t_{11} & t_{12} \\ t_{21} & t_{22} \end{pmatrix}, \quad (23)$$

$$\hat{\Gamma}(\chi, \hat{\mathbf{n}}) \equiv \int_0^\chi d\chi' \hat{\tau}(\chi' \hat{\mathbf{n}}) \chi' \left(1 - \frac{\chi'}{\chi}\right) = \int_0^\chi d\chi' \hat{\tau}(\chi' \hat{\mathbf{n}}) \chi' - \frac{1}{\chi} \int_0^\chi d\chi' \hat{\tau}(\chi' \hat{\mathbf{n}}) \chi'^2. \quad (24)$$

$$\hat{\Gamma}(\chi_i) = \hat{I}_i - \frac{1}{\chi_{i+1/2}} \hat{J}_i, \quad \hat{I}_i = \hat{I}_{i-1} + \hat{\tau}(\chi_i) \frac{\chi_{i+1/2}^2 - \chi_{i-1/2}^2}{2}, \quad \hat{J}_i = \hat{J}_{i-1} + \hat{\tau}(\chi_i) \frac{\chi_{i+1/2}^3 - \chi_{i-1/2}^3}{3} \quad (25)$$

7 CMB lensing

For $\kappa \equiv \text{Tr}(\hat{\Gamma})/2$ we compute, from the simulation, $\kappa(\chi_{\text{max}})$, where χ_{max} is the maximum radial distance that fits in the box. In order to get CMB lensing we need $\kappa(\chi_{\text{LSS}}) = \tilde{\kappa}_{\text{max}} + \Delta\kappa$, where

$$\tilde{\kappa}_{\text{max}} \equiv \int_0^{\chi_{\text{max}}} d\chi \delta(\chi) \frac{\chi}{2} \left(1 - \frac{\chi}{\chi_{\text{LSS}}}\right) \quad (26)$$

$$\Delta\kappa \equiv \int_{\chi_{\text{max}}}^{\chi_{\text{LSS}}} d\chi \delta(\chi) \frac{\chi}{2} \left(1 - \frac{\chi}{\chi_{\text{LSS}}}\right). \quad (27)$$

The strategy to compute these is:

- We compute $\tilde{\kappa}_{\text{max}}$ from $\kappa(\chi_{\text{max}})$ using the same strategy used for shear, but taking care to divide by χ_{LSS} instead of χ_{max} .
- We compute $\Delta\kappa$ as a Gaussian realization constrained to have the right correlation with $\tilde{\kappa}_{\text{max}}$. Do do this we start by rewriting the previous equation as

$$\tilde{\kappa}_{\text{max}} \equiv \int d\chi w_1(\chi) \delta(\chi \hat{\mathbf{n}}), \quad \Delta\kappa \equiv \int d\chi w_2(\chi) \delta(\chi \hat{\mathbf{n}}), \quad (28)$$

$$w_1(\chi) \equiv \frac{\chi}{2} \left(1 - \frac{\chi}{\chi_{\text{LSS}}}\right) \Theta(\chi, 0, \chi_{\text{max}}), \quad w_2(\chi) \equiv \frac{\chi}{2} \left(1 - \frac{\chi}{\chi_{\text{LSS}}}\right) \Theta(\chi, \chi_{\text{max}}, \chi_{\text{LSS}}). \quad (29)$$

where $\Theta(x, x_0, x_f)$ is a top-hat function.

The covariance matrix of the two terms is therefore

$$\langle |\tilde{\kappa}_{\text{max}, \ell m}|^2 \rangle = \frac{2}{\pi} \int_0^\infty dk k^2 P(k) w_{1, \ell}^2(k), \quad \langle |\Delta\kappa_{\ell m}|^2 \rangle = \frac{2}{\pi} \int_0^\infty dk k^2 P(k) w_{2, \ell}^2(k) \quad (30)$$

$$\langle \text{Re}(\tilde{\kappa}_{\text{max}, \ell m} \Delta\kappa_{\ell m}) \rangle = \frac{2}{\pi} \int_0^\infty dk k^2 P(k) w_{1, \ell}(k) w_{2, \ell}(k), \quad w_{i, \ell}(k) \equiv \int_0^\infty d\chi w_i(\chi) j_\ell(k\chi) \quad (31)$$

Thus we generate a realization of $\Delta\kappa$ at each multipole order as a Gaussian number with distribution $\mathcal{N}(\mu, \sigma)$, where

$$\mu = \frac{\langle \text{Re}(\tilde{\kappa}_{\text{max}, \ell m} \Delta\kappa_{\ell m}) \rangle}{\langle |\tilde{\kappa}_{\text{max}, \ell m}|^2 \rangle} \tilde{\kappa}_{\text{max}, \ell m}, \quad \sigma = \langle |\Delta\kappa_{\ell m}|^2 \rangle - \frac{\langle \text{Re}(\tilde{\kappa}_{\text{max}, \ell m} \Delta\kappa_{\ell m}) \rangle^2}{\langle |\tilde{\kappa}_{\text{max}, \ell m}|^2 \rangle} \quad (32)$$