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Index number: 431289

Search for dormant Black Holes in OGLE survey database.

Bachelor degree in the field of Astronomy

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Warsaw, NAN 2022 r.

Abstract

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Keywords

Dormant Black Holes,

Thesis title in Polish language

Poszukiwanie uśpionych czarnych dziur w bazach danych przeglądu OGLE.

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Chapter 1

Introduction

According to current knowledge most of the massive stars end their evolution as black holes, further investigation suggest existence of $\mathcal{O}(10^8)$ stellar mass BH's (Brown & Bethe, 1994) in our galaxy. Despite huge theoretical expectations only few tens of such objects have been found up to day. Most of them either emit X-rays through accretion mechanism or can distinguish themselves with exotic astrometry/radial velocity. Search for X-ray emission nor exotic movement are viable method to search BH search. In the first case binary system is typically composed of either low mass star which overflows it's Roche lobe (BH-LMXB) or high mass star which is accreting mass onto companion BH via stellar wind. In first case X-ray emission mostly have transient character (Belloni & Motta, 2016) making it hard to observe them due to the accretion disc instabilities (Lasota, 2001). What is crucial to emphasize is relative small period of the Black Hole binaries which can be uncovered from X-ray emission as longest one reported today for confirmed BH binary is 33 d (GRS1915+105), wider orbits decrease relative X-ray brightness decreasing chances to detect such objects. In recent years new stellar mass BH's candidates are being found due to the new methods not based on X-ray emission. Noticeable cases include spectroscopic observations (Liu et al., 2019; Jayasinghe et al., 2021), microlensing (Sahu et al., 2022) and astrometric observations (El-Badry et al., 2022). Claims about BH nature of some of those objects have been contested nevertheless current spectroscopic, photometric and astrometric surveys provide new way to search for those exotic objects, which can accelerate more in the next years due to better observational data.

In recent years new approach was proposed aiming to find objects which do not emit X-rays but have filled theirs Roche lobe. In the case of such binaries one can expect to observe ellipsoidal modulation caused by surface disruption. This approach have great advantage due to the fact that photometric data is very vast in contrast to spectroscopic one. Relative potential of this approach was shown by (Masuda & Hotokezaka, 2019), it's estimated that light curve modulation due to ellipsoidal effect can be enough to track down at least few of them in TESS data. Other noticeable case include series of publications (Gomel et al., 2021b,a,c) on which this works is based. Publications presented new way to search for this type of objects using publicly available list of eclipsing variables in the direction of galactic bulge from OGLE database (Soszyński et al., 2016). This approach itself provides a rather good and robust way to look for the systems with high mass ratios. In the chapter 2 theoretical introduction to ellipsoidal modulation will be given and in the chapters it following method is employed to search for BH candidates in OGLE database.

Chapter 2

Theoretical considerations

2.1. Morris & Naftilan expression

Let's consider binary system with major axis of orbit a , primary star mass M_1 and radius R_1 which is tidal disrupted by second object with mass M_2 . Traditionally one can define mass ratio $q = \frac{M_2}{M_1}$, in our case we want this value to be as high as possible which indicate that our primary star is less massive then (hopefully unseen) companion contrary to normal convention where we define q to be strictly smaller then 1 (assuming primary star is more massive). Up to day many publications tackled problem of ellipsoidal modulation including (Kopal, 1959) leading to formula proposed by (Morris & Naftilan, 1993) which take form

$$\frac{\Delta L}{\bar{L}} = \frac{\alpha_1}{\bar{L}/L_0} \left(\frac{R_1}{a} \right)^4 q (4 \sin i - 5 \sin^3 i) \cos \phi - \quad (2.1)$$

$$\frac{1}{\bar{L}/L_0} \left(\alpha_2 \left(\frac{R_1}{a} \right)^3 + \beta_2 \left(\frac{R_1}{a} \right)^5 q (6 \sin^2 i - 7 \sin^4 i) \right) \cos 2\varphi \quad (2.2)$$

$$- \frac{5}{3} \frac{\alpha_1}{\bar{L}/L_0} \left(\frac{R_1}{a} \right)^4 q \cos 3\varphi = \quad (2.3)$$

$$A_1 \cos \varphi + A_2 \cos 2\varphi + A_3 \cos 3\varphi \quad (2.4)$$

where φ is relative phase, L_0 is luminosity of star *without tidal disruption* while \bar{L} stands for mean luminosity of our primary star. From this point onward equation 2.1 will be referred to as MN93. Coefficients α_1 , α_2 and β_2 are connected to limb darkening coefficient u and gravity darkening coefficient τ via equations

$$\alpha_1 = \frac{15u(2 + \tau)}{32(3 - u)} \quad (2.5)$$

$$\alpha_2 = \frac{3(15 + u)(1 + \tau)}{20(3 - u)} \quad (2.6)$$

$$\beta_2 = \frac{15(1 - u)(3 + \tau)}{64(3 - u)} \quad (2.7)$$

. Both values \bar{L} and L_0 are also related and can be found via following equation

$$\bar{L} = L_0 \left(1 + \frac{1}{9} \left(\frac{R_1}{a} \right)^3 (2 + 5q)(2 - 3 \sin^2 i) \right) \quad (2.8)$$

Equations can be rewritten in more suitable form when we replace $\frac{R_1}{a}$ as

$$\frac{R_1}{a} = \frac{R_1}{R_{Roche}} \frac{R_{Roche}}{a} = E(q)f$$

where R_{Roche} is radius of Roche lobe, f stands for roche lobe fillout and $E(q)$ is ratio of Roche lobe radius to semi major axis which can be described by classic Eggleton formula (Eggleton, 1983)

$$E(q) = \frac{0.49q^{-\frac{2}{3}}}{0.6q^{-\frac{2}{3}} + \log(1 + q^{-\frac{1}{3}})} \quad (2.9)$$

. MN93 equation in original form is valid only for small values of f, q making it unsuitable in our case as one would like to find objects with high values of q .

2.2. Amplitude correction

There are few approaches one can take to extend MN93 formula, in this work original formulation from (Gomel et al., 2021b) was adopted. In original work each coefficient was multiplied by correction $C_i(q, f, i)$ computed using PHOEBE simulation software. For each fourier coefficient relevant model was fitted to PHOEBE data and then compared allowing to find such function $C(q, f, i)$ that

$$\left| C(q, f, i) - \frac{A_{Phoebe}}{A_{MN93}} \right| \quad (2.10)$$

will be minimized. In Gomel et al. (2021b) following functions were chosen for second and third correction coefficients

$$C_2(q, f) = 1 + \left(0.0379 + \frac{0.005}{0.0446 + q} \right) \left(\frac{f}{1.0909 - f} \right) \quad (2.11)$$

$$C_3(q, f) = 1 + \frac{(1 + 0.0698q \sin^2 i) f^6 + 0.2075 \cdot f^2}{(2.0223 + 0.3880 \ln q) f + \sin^4 i} \quad (2.12)$$

as there had closed analytical form and small relative error.

2.3. Limitations

In previous section new corrected formula was introduced allowing to predict amplitudes of fourier series based on physical parameters q, f, i . After correction our formula have relatively good accuracy but what need to be emphasised are limitations due to other physical phenomena. To begin with MN93 formula is valid only for circular orbits. In order to lift that limitation one can reproduce results using other type of expansion such as TODO ENGEL. Moreover if the binary will be compact we need to take into account relativistic beaming effect. Following (Loeb & Gaudi, 2003)

$$F_\lambda = F_{\lambda,0} \left(1 - B \frac{v_r}{c} \right) \quad (2.13)$$

where F_λ is spectral energy density, v_r radial velocity and B is coefficient defined as $B = 5 + \frac{\partial \ln F_\lambda}{\partial \ln \lambda}$. One need to emphasize that relativistic beaming effect does not affect second or third harmonic coefficient. This is purely due to the fact that beaming is caused by motion while our main tidal effect is in second harmonic and can be interpreted as effect connected

purely to position. This important fact makes second harmonic best for any search when the goal is to predict physical parameters based on amplitude of our harmonics. There are also few other effects one can consider such as reflection effect, one can find detail analysis in (Gomel et al., 2021b).

2.4. Modified mass ratio

After derivation of expression for second harmonic coefficient we can introduce *modified mass ratio* following (Gomel et al., 2021a). As one can verify $A_2(q, f, i, \alpha)$ is increasing function of both $\sin i$ and f . This allows us to introduce q_{mmr} such that

$$A_2(q_{mmr}, 1, \pi/2, 1.2) = \tilde{A}_2 \quad (2.14)$$

where \tilde{A}_2 is measured value of second harmonic amplitude. This definition assures that $q \leq q_{mmr}$ is lower bound for true mass ratio for any values of i and f . This property is crucial and provide a simple way to search for objects with potentially compact companion. Second thing one need to emphasize is value of α_2 . Following original work $\alpha_2 = 1.2$ was set for observation in I band and also standard deviation $\Delta\alpha_2 = 0.1$ was assumed. This value solely represents effects of temperature, gravity and chemical composition making it very important as it's encapsulating all parameters we don't know. Exact value of parameter greatly influence modified mass ratio due to which inference of mass without exact knowledge of α_2 is impossible. This fact lead authors of method to introduce second parameter which was originally defined as 16th percentile of our estimator q_{mmr} . Solely role of this parameter is to take into account uncertainty of α acting as lower bound of q_{mmr} . In this work other parameter was introduced denoted as $q\tilde{q}_{mmr}$ and defined as

$$A_2(\tilde{q}_{mmr}, 1, \pi/2, 1.1) = \tilde{A}_2 - \Delta A_2$$

where ΔA_2 is uncertainty of estimate of A_2 coefficient. This approach not only allow to filter out those stars which can have large amplitude due to surprisingly high value of α_2 which is taken into account by decreasing α_2 in definition 2.4 to 1.2 but also allows to filter out stars that have high amplitude but also relatively high uncertainty ΔA_2 .

2.5. Selection process

In order to provide reliable and robust way to search for compact companions following procedure using aforementioned estimators of q_{mmr} was used. For each object two coefficients q_{mmr} and \tilde{q}_{mmr} were calculated and each object for which $\tilde{q}_{mmr} > 1$ was qualified as candidate for compact companion. Secondly each object with amplitude $\tilde{A}_2 > 0.24$ was removed from the list. One can find, that using 2.14 one need mass ratio of order $q_{mmr} \approx 1000$ to obtain configuration with such great amplitude suggesting that modulation is not due to ellipsoidal effect but rather caused by other type of star variability. Exact threshold of \tilde{A}_2 can be a matter of debate due to previously mentioned dependence on α_2 which is unknown.

Chapter 3

Observational data

3.1. Description of data and preprocessing

From OGLE database two object samples were analysed using described method. First sample consisted of 6620 object from Small Magellanic Cloud (SMC), Large Magellanic Cloud (LMC) and Magellanic Bridge (MBR) while second sample consisted of stars from Galactic Disc (GD) and from the direction towards Galactic Bulge (BLG) (1895 objects). While first sample consisted of any variable objects found in the direction of Magellanic Clouds second sample was constructed by cross matching OGLE data with list of objects from Gaia DR 3 (Collaboration et al., 2022) with values of velocity amplitude (*rv_amplitude_robust*) higher then 100 km/s. Light curves used in the analysis were collected during 4 th season of OGLE (Udalski et al., 2015), (Udalski et al., 1992) and were collected in *I* band. Before preprocessing data objects with brightness lower then 17 mag were removed as faint objects would be unsuitable for any kind of spectroscopic observations.

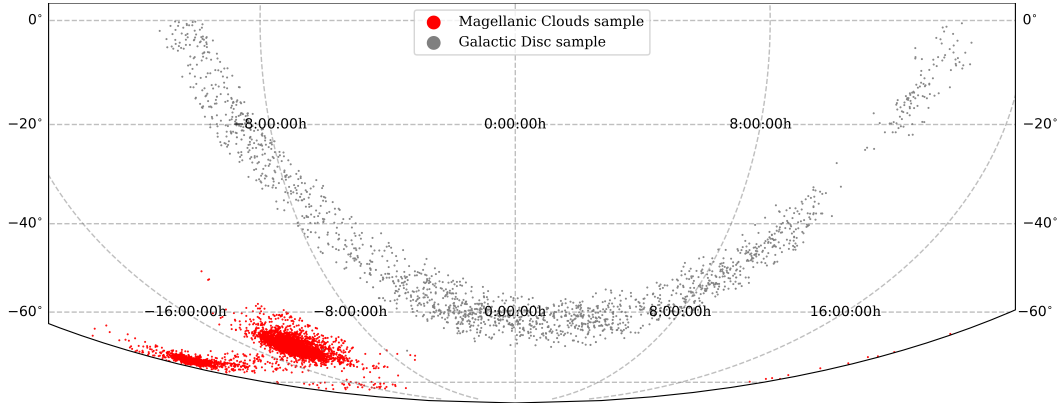


Figure 3.1: Distribution of objects in each sample.

Each object was analysed using ANOVA (analysis of variance) method (Schwarzenberg-Czerny, 1989) in order to determine period of light curve. Then each light curve was fitted with 4th degree harmonic model

$$I(t) = A_0 + \sum_i^4 A_{1i} \sin\left(2\pi i \frac{t}{P}\right) + A_{2i} \cos\left(2\pi i \frac{t}{P}\right) \quad (3.1)$$

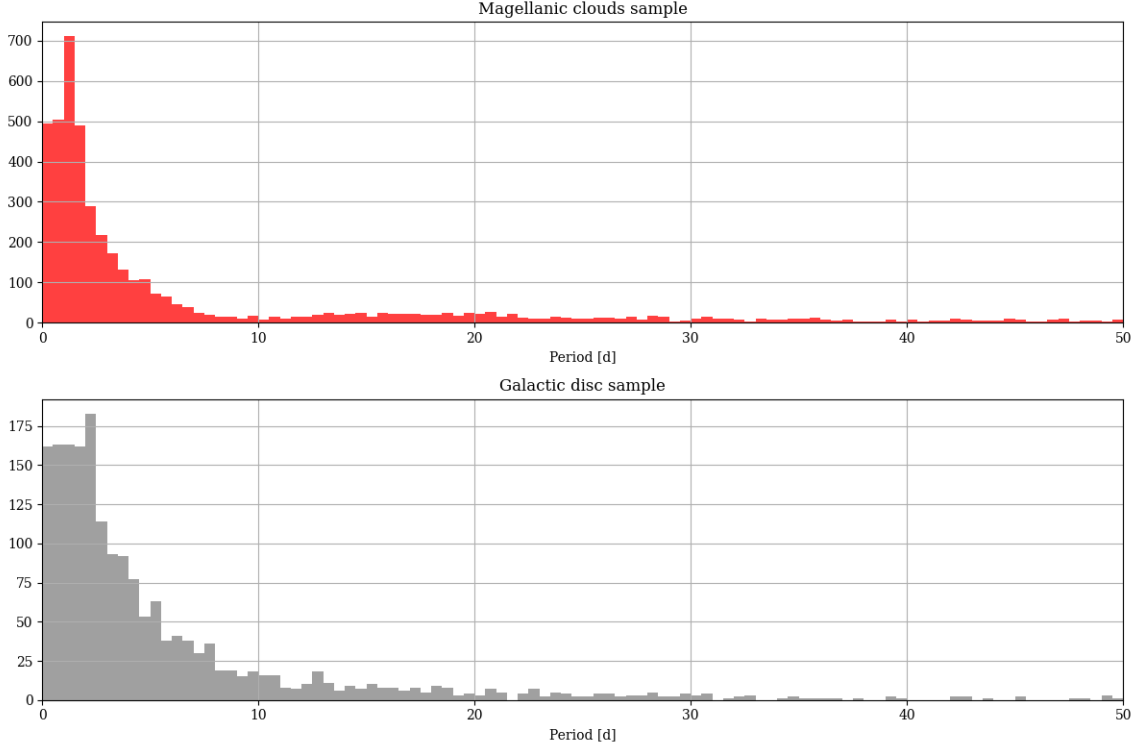


Figure 3.2: Distribution of periods from the analysed samples.

with sigma clipping threshold set at 3σ allowing to determine fourier coefficients from the relationship $A_i = \sqrt{A_{1i}^2 + A_{2i}^2}$. Subsequently minimum mass ratio q_{mmr} together with it's lower bound \tilde{q}_{mmr} were calculated and previously described procedure of selecting candidates was carried out. Objects with light curves indicating other type of variability then ellipsoidal one were removed together with those with period higher then 50 d. As we assume objects should be composed of star nearly filling Roche lobe suggesting rather small period. Exact value of threshold can be debated, here it's used as mean to remove pulsating stars that pollute our sample. After this part of preprocessing only 41 objects from first sample left together with 22 objects from second sample.

3.2. Spectral Energy Distribution

Each object was cross matched with Gaia DR3 catalogue obtaining parallax (denoted from here as π_0 together with uncertainty denoted as σ_π) estimates of objects. If parallax was statistically significant ($\pi > 3\sigma_\pi$) it was used to derive distance to object. For objects towards Magellanic Clouds many entries lacked statistically significant parallax indicating that they really reside in Magellanic Clouds and do not line up accidentally. In this case distance estimate d_0 and distance uncertainty σ_d were based on the work (Jacyszyn-Dobrzniecka et al., 2016) assuming that presented distribution of cepheid reflects underlying distribution of stars in sample. In the case of LMC and SMC assumed distance estimates were $d_{LMC} = 49.93 \pm 1.79$ kpc and $d_{SMC} = 64.62 \pm 4.95$ kpc respectively.

To further investigate nature of objects spectral energy distribution analysis was performed with two models: single star model and double star model. Main goal of this type

of analysis is to use photometry from various parts of spectra to reconstruct whole spectral characteristic and hence obtain parameters of our object temperature. Single star model depends on three free parameters: logarithm of temperature $\log_{10} T$ (measured in kelvins), logarithm of luminosity $\log_{10} L$ (measured in L_{\odot}) while third parameter was either parallax (if one was measured) or distance if it wasn't available. In the case of double star model there were two additional parameters describing second star $\log_{10} T_2$ and $\log_{10} L_2$. In the case of the first model star was assumed to be main sequence object (MS) with $\log g = 4$. In double star model more massive star was assumed to have $\log g_1 = 4$ as in the first case while second object was assumed to be a giant with $\log g_2 = 2$. Objects with statistically significant parallax were assumed to have solar like metallicity ($Z = 0.013$) as they reside inside galactic disc while objects in Magellanic Clouds were assumed to have metallicity equal to $Z = 0.010$ in the case of LMC/MBR and $Z = 0.005$ in the case of SMC. BaSeL library of stellar spectra (Lejeune et al., 1998) was used to find spectrum for any given $\log_{10} T$, $\log_{10} L$, d by means of interpolation as implemented in python library `pystellib`¹. Theoretically calculated stellar spectrum is then processed using `pyphot`² Python library yielding observed magnitude in required filter.

In order to find set of best fitting parameters Bayesian approach was adopted. Let's denote set of observed magnitudes as \tilde{m}_i , theoretically predicted magnitudes as m_i while errors of magnitudes as σ_i . Let's denote $\mathcal{U}(a, b)$ to be uniform distribution with support in form of interval $[a, b]$ and $\mathcal{N}(\mu, \sigma^2)$ as normal distribution with mean μ and variance σ_i^2 . In order to prepare Bayesian model one need to specify probability distributions that allow to evaluate likelihood of data. Each of the observed magnitudes \tilde{m}_i is expected to be normally distributed with mean m_i and variance σ_i^2 while prior distributions on each of the three parameters is either uniform in case of $\log_{10} T$ and $\log_{10} L$ or normal in case of distance or parallax. Detailed model can be written as:

$$\log_{10} T \sim \mathcal{U}(3.31, 4.6) \quad (3.2)$$

$$\log_{10} L \sim \mathcal{U}(-3, 5) \quad (3.3)$$

$$d \sim \mathcal{N}(d_0, \sigma_d^2) \text{ or} \quad (3.4)$$

$$\pi \sim \mathcal{N}(\pi, \sigma_{\pi}^2) \quad (3.5)$$

$$\tilde{m}_i \sim \mathcal{N}(m_i(\log_{10} T, \log_{10} L, d), \sigma_i^2) \text{ or} \quad (3.6)$$

$$\tilde{m}_i \sim \mathcal{N}(m_i(\log_{10} T, \log_{10} L, \pi), \sigma_i^2) \quad (3.7)$$

where $m_i(\log_{10} T, \log_{10} L, d)$ is written to indicate that predicted magnitude is function of parameters. Similarly we can write down our second model together with priors for parameters

$$\log_{10} T_1 \sim \mathcal{U}(3.31, 4.6) \quad (3.8)$$

$$\log_{10} L_1 \sim \mathcal{U}(-3, 5) \quad (3.9)$$

$$\log_{10} T_2 \sim \mathcal{U}(3.444, 4.21) \quad (3.10)$$

$$\log_{10} L_2 \sim \mathcal{U}(-3, 5) \quad (3.11)$$

$$d \sim \mathcal{N}(d_0, \sigma_d^2) \text{ or} \quad (3.12)$$

$$\pi \sim \mathcal{N}(\pi, \sigma_{\pi}^2) \quad (3.13)$$

$$\tilde{m}_i \sim \mathcal{N}(m_i, \sigma_i^2) \quad (3.14)$$

where we hidden explicit dependence of m_i on parameters for clarity. In both cases normal distribution is truncated to positive numbers as we do not permit negative parallax/distance.

¹<https://github.com/mfouesneau/pystellib>

²<https://github.com/mfouesneau/pyphot>

Range on the logarithm of temperature is limited due to boundaries of library while limit of the logarithm of luminosity is set in the boundaries in order to eliminate possible unphysical solutions. Under those assumptions log-likelihood function can be written as

$$\mathcal{L} = - \sum_i \frac{(\tilde{m}_i - m_i)^2}{2\sigma_i^2} - \frac{(d - d_0)^2}{2\sigma_d^2} \text{ or} \quad (3.15)$$

$$\mathcal{L} = - \sum_i \frac{(\tilde{m}_i - m_i)^2}{2\sigma_i^2} - \frac{(\pi - \pi_0)^2}{2\sigma_\pi^2} \quad (3.16)$$

depending whether distance or parallax was used.

Following catalogues were used to assemble SEDs:

1. Catalogues shared by both samples of objects:

- 2MASS survey (Skrutskie et al., 2006)
- Gaia DR2 (Gaia Collaboration et al., 2018)
- VISTA Hemisphere Survey DR5 (McMahon et al., 2021)
- ALLWISE/WISE survey (Wright et al., (2010) Cutri et al., (2021))
- GALEX Survey (Bianchi et al., 2011)
- Spitzer IRAC data (Meixner et al., 2006)
- Denis survey (Denis, 2005)
- SkyMapper DR2 (Wolf et al., (2018) Onken et al., (2019))
- XMM Optical Monitor data (Page et al., 2012)

2. Catalogues exclusive to the Magellanic Clouds:

- VISTA Magellanic Cloud survey DR4 (Cioni et al., 2017)
- Denis catalog of objects in Magellanic Clouds (Cioni et al., 2000)
- Magellanic Clouds Photometric Survey: SMC and LMC (Zaritsky et al., 2002) (Zaritsky et al., 2004)

3. Catalogues exclusive to the Galactic Disc:

- Bochum Galactic disc survey (Hackstein et al., 2015)
- AAVSO Photometric All Sky Survey DR9 (Henden et al., 2015)
- VISTA Variable in Via Lactea Survey DR2 (Minniti et al., 2017)

. In the case of magellanic clouds extinction estimates were based on the map (Skowron et al., 2021) while in the case of Galactic disc extinction was obtained using `mw dust` (Bovy et al., 2016) Python library with 3D dust map being combination of (Green et al., 2019), (Greiner et al., 2001), (Drimmel et al., 2003). In the calculations Cardeli extinction law was assumed (Cardelli et al., 1989) and python implementation from `extinction`³ was used. Using described setup MCMC python based library `emcee`⁴ (Foreman-Mackey et al., 2013) was used to construct compact library⁵ of routines used to sample from the posterior of our

³<https://extinction.readthedocs.io/en/latest/index.html>

⁴<https://emcee.readthedocs.io/en/stable/>

⁵<https://github.com/Wesenheit/Iris>

models allowing to estimates parameters together with associated uncertainties. In order to help to decide between models BIC score was used

$$BIC = k \log n - 2\mathcal{L} \quad (3.17)$$

where k is number of estimated parameters and n is number of data points. From all objects 13 were selected based on high quality fit with single star model while one of the objects was selected due to observed X-ray emission that can indicate existence of accretion disc. Complete list of objects together with key parameters is listed in table 4.1.

Name	Period [d]	q_{mmr}	\tilde{q}_{mmr}	RA [deg]	DEC [deg]	A_0	A_1	A_2	A_3	A_4	π_0 [mas]
BLG986.08.7	0.5132	1.464	1.029	260.451241	-43.019464	12.192	0.0028	0.0985	0.0039	0.0093	1.20
GD2246.03.18414	0.4270	2.299	1.548	179.327372	-57.091968	12.0764	0.0045	0.1090	0.0030	0.0106	1.41
GD1097.20.23000	0.4570	9.739	5.178	252.108419	-44.137004	11.782	0.0008	0.1414	0.0029	0.0200	1.78
GD1448.27.17	1.2410	1.777	1.233	139.969949	-45.759178	11.875	0.0115	0.1035	0.0030	0.0056	0.77
BLG931.27.36745	0.7050	2.437	1.628	261.74489	-40.31792	11.633	0.0033	0.1113	0.0026	0.0105	1.46
GD1070.18.22288	45.1467	7.409	3.726	257.007546	-41.048747	11.369	0.0091	0.1362	0.0043	0.0049	1.03
LMC574.11.3407	0.2551	1.479	1.026	80.1415	-63.185667	16.919	0.0064	0.0989	0.0013	0.0092	0.45
LMC606.30.48	0.2698	1.668	1.113	92.27575	-63.376167	15.906	0.0141	0.1019	0.0038	0.0122	0.52
LMC751.15.2886	0.3753	1.703	1.161	98.082958	-66.307861	16.958	0.0090	0.1024	0.0012	0.0121	0.18
MBR108.18.3	0.2947	1.487	1.024	33.257292	-72.993167	13.376	0.0063	0.0989	0.0033	0.0128	1.10
MBR236.09.433	0.4398	1.511	1.070	50.642958	-80.540667	14.585	0.0024	0.0994	0.0028	0.0089	0.47
SMC711.22.1068	0.4466	2.346	1.555	9.833875	-70.378028	13.226	0.0029	0.1104	0.0053	0.0154	0.66
SMC720.28.40576	0.5674	4.290	2.555	11.938042	-73.13625	16.367	0.0017	0.1246	0.0037	0.0035	—
SMC742.26.330	0.3453	1.676	1.118	350.85925	-77.530417	13.864	0.0042	0.1020	0.0049	0.0074	1.07

Table 3.1: Selected objects together with period, estimated mass ratios, coordinates, fourier coefficients (A_i) and parallax from Gaia DR3

What should be emphasised here is meaning of BIC score. As we assume little about our objects especially we do not investigate best fitting $\log g$ which can be crucial to estimate stellar parameters. Due to this effect it's unreasonable to calculate values of χ^2 and compare them across objects as few of them can be better fitted with given metallicity/ $\log g$.

Chapter 4

Results

4.1. Physical parameters of objects

Each of 14 final object was further investigated using available data in order to determine physical properties of objects. Out of all 14 objects 12 of are of spectral types G or F, one object is of spectral type O while last one is composed of two objects. One of those objects is spectral type G while second is spectral type M. 13 stars from the list have measured parallax from Gaia DR3 while one of the objects is located in SMC. 11 objects were selected based on good fit with single star model while one was selected due to counterpart emission in X-ray.

For each entry mass of object was estimated using PARSEC (Costa et al., 2019a),(Costa et al., 2019b) evolutionary tracks. Assuming solar like metallicity simple approximate fits were obtained by choosing best fitting entry from PARSEC track using temperature and luminosity estimates from SED fits.

Name	T [K]	R [R_{\odot}]	$E(B - V)$	d [kpc]	BIC	M_{PARSEC} [M_{\odot}]
BLG986.08.7	6118^{+19}_{-20}	2.22 ± 0.01	0.226	0.83 ± 0.01	1.3	
GD2246.03.18414	7246 ± 41	1.63 ± 0.02	0.273	0.71 ± 0.01	91.4	0.8
GD1097.20.23000	6068 ± 13	1.90 ± 0.01	0.221	0.56 ± 0.008	386	1.2
GD1448.27.17	8065 ± 39	4.13 ± 0.04	0.611	1.30 ± 0.02	162.8	2.3
BLG931.27.36745	6434 ± 10	2.25 ± 0.01	0.199	$0.69^{+0.04}_{-0.03}$	193.3	1.4
LMC574.11.3407	4680 ± 10	1.01 ± 0.01	0.034	$2.20^{+0.46}_{-0.30}$	48.5	—
LMC606.30.48	5280 ± 11	1.08 ± 0.01	0.047	$1.93^{+0.15}_{-0.12}$	101.7	0.8
LMC751.15.2886	5768 ± 6	1.30 ± 0.01	—	0.91 ± 0.01	3318.9	3.2
MBR108.18.3	5943^{+23}_{-20}	1.56 ± 0.01	0.000	$5.70^{2.98}_{-1.49}$	75.1	1.1
MBR236.09.433	6285 ± 8	1.88 ± 0.01	—	2.12 ± 0.07	950.5	1.2
SMC711.22.1068	6717 ± 13	1.99 ± 0.01	0.021	1.52 ± 0.03	1541.2	1.4
SMC720.28.40576	34183^{+985}_{-655}	$4.41^{+0.17}_{-0.24}$	0.095	—	263.7	16
SMC742.26.330	5615 ± 5	1.57 ± 0.01	—	0.93 ± 0.01	4736.4	1

Table 4.1: Estimated physical parameters of objects using single star model together with PARSEC mass estimates and extinction estimates.

4.2. Radial velocity semi-amplitude estimation from Gaia DR3 data

Half of the objects from final list have available high-quality radial velocity information that is normally available for bright stars from Gaia DR3 catalogue. While estimate of radial velocity is based on median of observed radial velocity, error of this estimate is based on the epoch standard deviation of observed values. According to (ref) radial velocity error δv is calculated via equation

$$\delta v = \sqrt{\sigma_{med}^2 + 0.11^2} \quad (4.1)$$

$$\sigma_{med} = \sqrt{\frac{\pi}{2}} \frac{\sigma(V_j)}{\sqrt{N}} \quad (4.2)$$

where $\sigma(V_j)$ is standard deviation of radial velocity while N stands for number of transits used to compute radial velocity. This simple formula allows to obtain variance of radial velocity measurements as

$$\sigma^2 = \frac{2N}{\pi} ((\delta v)^2 - 0.11^2) \quad (4.3)$$

. It can be proven (for details see appendix A) that if we assume error is dominated by sinusoidal radial movement with amplitude K variance of velocity will be equal to

$$\sigma^2 = \frac{K^2}{2} \quad (4.4)$$

. Using this fact one can estimate semi-amplitude of velocity using gaia measurements (that will be denoted from here as K_{Gaia}) as

$$K_{Gaia} = \sqrt{\sigma^2} \sqrt{2} = 2 \sqrt{\frac{N}{\pi} ((\delta v)^2 - 0.11^2)} \quad (4.5)$$

This radial velocity amplitude then is used to calculate binary mass function for each of objects using definition

$$f(M_1, M_2, i) = \frac{M_1^3 \sin^3 i}{(M_1 + M_2)^2} = \frac{K^3 P}{2\pi G} \quad (4.6)$$

where P stands for period while G is gravitational constant. For each object with RV estimates based on aforementioned method mass functions were obtained and together with information about radial velocity estimates are presented in table 4.2.

Name	Period [d]	RV [km/s]	N	δv [km/s]	K_{Gaia} [km/s]	$f(M_1, M_2, i)$
GD1549.19.348	2.8224	36.72	21	8.64	44.67	0.026
GD2246.03.18414	0.4270	27.68	25	23.03	129.93	0.097
GD1097.20.23000	0.4570	36.5	10	21.03	75.03	0.020
BLG986.08.7	0.5132	-1.69	15	18.11	79.14	0.026
BLG931.27.36745	0.7050	-4.35	8	14.81	47.26	0.0077
GD1448.27.17	1.2410	33.08	22	7.69	40.69	0.0086
GD1070.18.22288	45.1467	38.71	20	7.66	38.65	0.27

Table 4.2: Estimated mass function using Gaia radial velocity error.

4.3. Detailed analysis of objects

In this section detailed characteristic of individual objects will be provided. It will be divided into 3 subsections where first two will quickly cover information available for objects GD1070.18.22288 and GD1070.18.22288 while third one will be dedicated to rest.

4.3.1. GD1070.18.22288

GD1070.18.22280 has longest period from selected objects with $P \approx 45$ d and was listed as plausible candidate only due to the reported X-ray emission. Due to excess amount of observations in many parts of spectra it was possible to obtain high quality spectral energy distribution fit that revealed two sources with different temperatures, presented distribution can be seen on plot 4.1.

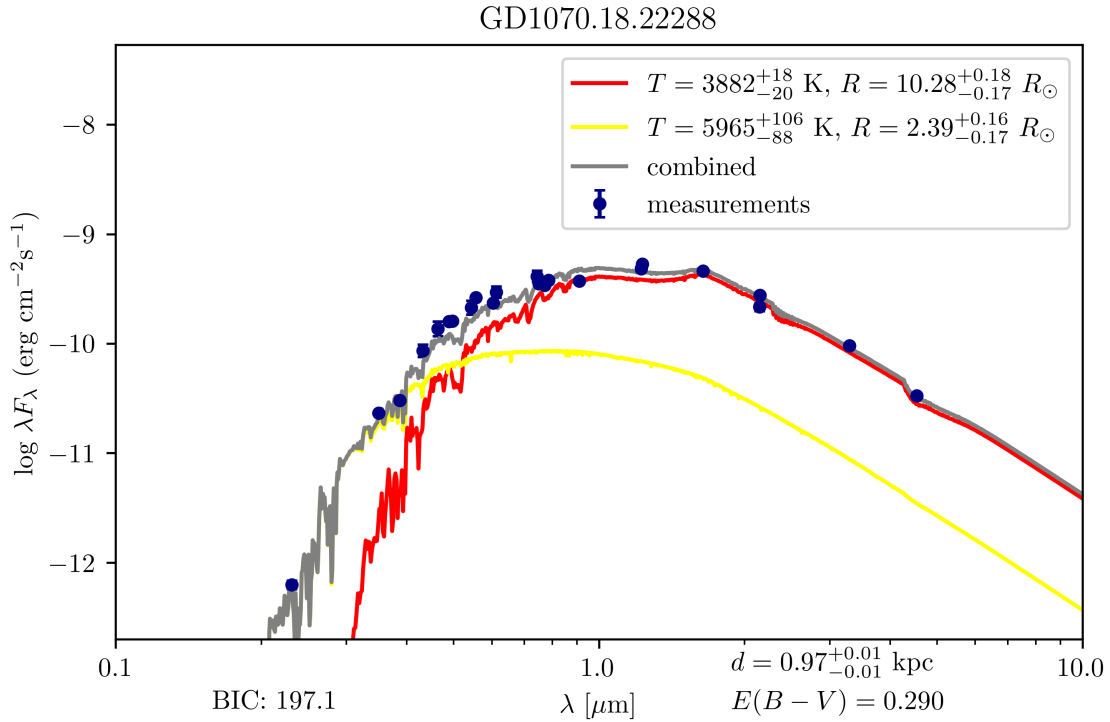


Figure 4.1: Spectral Energy Distribution fit for GD1070.18.22280 using model with two components.

One of the objects with rather low temperature of $T \approx 3800$ K has radius $R \approx 10 M_{\odot}$ and was quite distant from any evolutionary track in PARSEC database as it's probably giant stripped by it's companion. Object can be also found in ASAS-SN database (Jayasinghe et al., 2019) under id J170801.81-410255.6 where it's classified as rotational variable star with half of the period found in this work. It was observed by ASAS-SN in V and SDSS g band allowing to compare light curves with OGLE. Moreover it was observed in SDSS i band by Bochum Disc Survey (Hackstein et al., 2015). Comparison between all light curves can be find on the plot ??

GD1070.18.22288

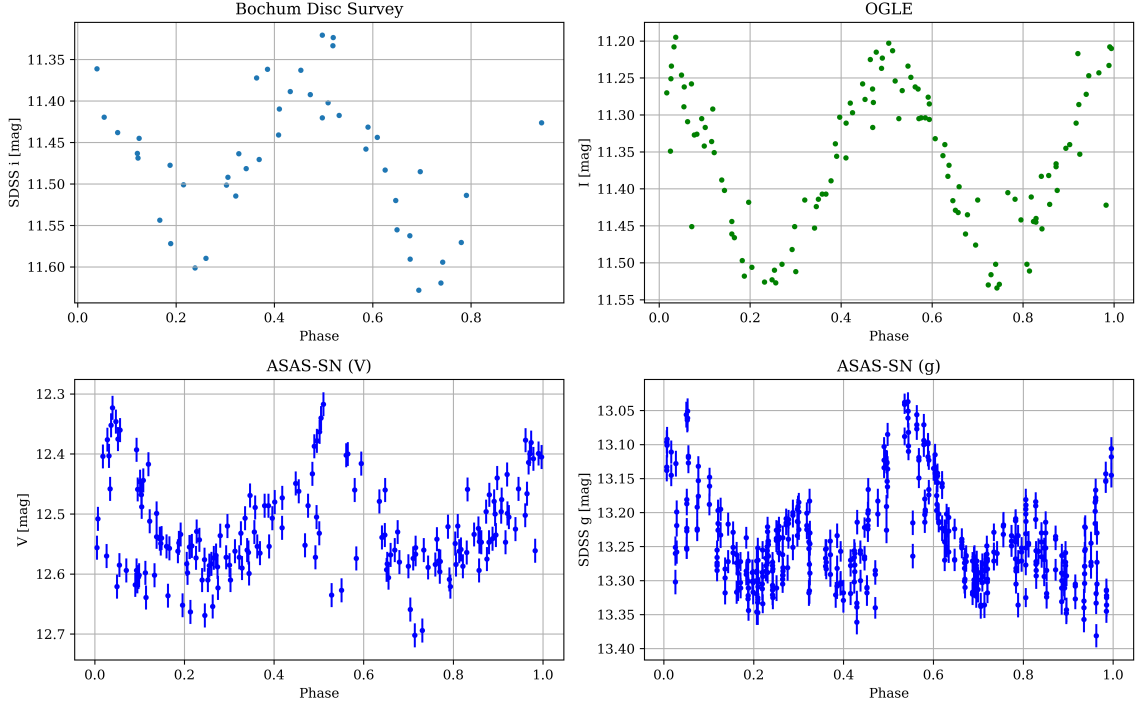


Figure 4.2: Light curves of GD1070.18.22288 in 4 different bands

ASAS observation in g band can be traced back to 2016 and allow to get better insight into evolution of light curve over time, change of observed brightness is presented on the plot 4.3. Object have also greatest value of mass function from whole sample that is equal to $\sim 0.25M_{\odot}$.

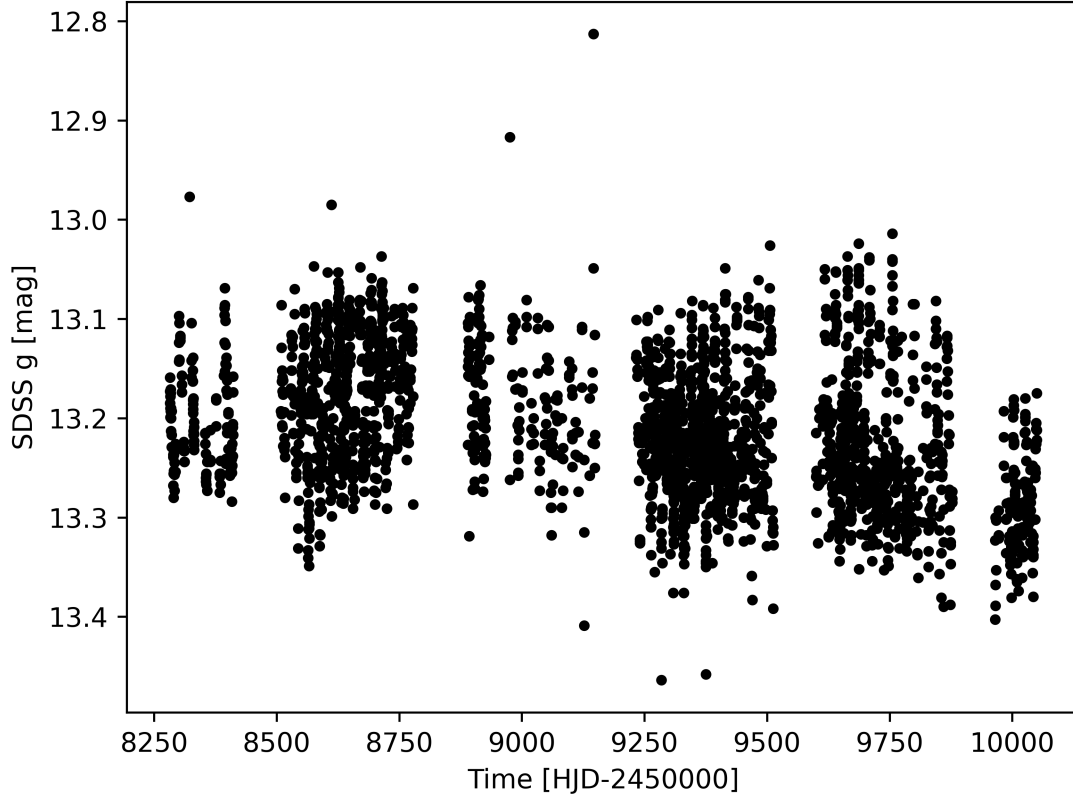


Figure 4.3: Change of brightness in SDSS g band over time

Object was observed in X-ray by Swift and XMM missions and published in serendipitous sources catalogs ((Evans et al., 2020) and (Traulsen et al., 2020)) with id names J170801.8-410254 and J170801.8-410255 respectively. Angular distance between position reported by XMM DR11 catalogs (Traulsen et al., 2020) and OGLE position is 105 mas and has reported statistical/positional error of order 365 mas. XMM object also came to attention due to gamma ray source source HESS J1708-410 (Aharonian et al., 2008) located 4.212 arcminutes from estimated source position. Results from elongated 2-D Gaussian model fit to gamma excess yielded standard deviations 3.6 arcminutes and 4.8 arcminutes along principal axis. Up to date no plausible explanation was found to the origin of aforementioned source despite multiple efforts and search in various parts of optical spectra. One particular work (Van Etten et al., 2009) observed close neighbourhood of gamma ray source and found emission from point source labeled in the publication as nr 1 (hereinafter referred to as [VFH2009] 1). This particular point source although quite faint coincides with position of GD1070.18.22280 with accuracy of 382 mas. In the publication it was found that X-ray emission is best fitted with absorbed power law by hydrogen density $n_H = 2.0 \times 10^{21} \text{ cm}^{-2}$. No conclusion was reached although publication explored few possibilities of the potential origin of gamma emission including Pulsar Wind Nebulae kind scenario. Interestingly assumed distance of 3 kpc based on radio observations is nearly 3 times higher then one obtained using parallax from Gaia DR3. If emission really originates from the system based on the Gaia parallax luminosity in X band would be around $L_X = 1.32 \cdot 10^{31} \text{ erg/s}$.

As one can see on figure 4.2 there is discrepancy between curves in I band and curves from ASAS-SN. This is due to the different time light curves were collected. OGLE light curve is composed from observations collected from 2456467 to 2458734 HJD but most samples were collected before 2457201 HJD. On the other hand observations from ASAS in g band were collected after 2459797 HJD. Using archival data from ASAS there is visible point around 2459100 when period of the variable changes from ~ 45 d to nearly half the value around ~ 22.4 d. This observation together with evolution of amplitude of the variability as presented on the plot 4.3 clearly suggest that object represents class of rotational variables. This can be also partially supported by the X ray emission from the system, it's widely known that many rotational variables like RS Canum Venaticorum can emit X rays with luminosity around $\sim 10^{31}$ erg/s (Walter et al., 1980) so value obtained from XMM is consisted with emission from this type of system. It's hard to determine if OGLE curve exhibit similar changes in brightness as observations cover only relatively short period of time. It's highly unlikely that variability in I band is caused by ellipsoidal modulation if period would be equal to 45 d as inferred radii are to small for the system to fill their Roche Lobes (unless system would have very small total mas). After investing all of the observational data it's hard to pinpoint true nature of the system. Most plausible explanation states that the system is similar to RS Canum Venaticorum variable which underwent mass transfer in the past. Colder star is covered in spots that emerge due to high magnetic fields, such stars can develop powerful coronal heating responsible for X-ray emission. Although one could investigate spectrum of object to find whether it's consisted with emission from hot plasma but as the object isn't ellipsoidal variable this line of investigation is dropped as it's beyond scope of the work.

4.3.2. SMC720.28.40576

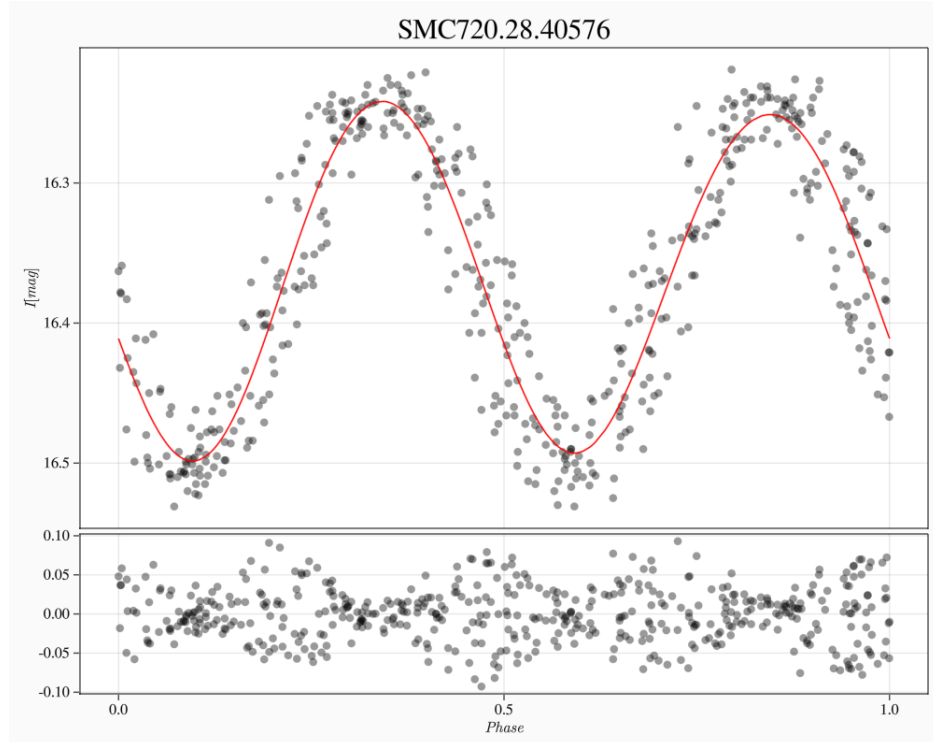


Figure 4.4: Light curve of SMC720.28.40576 collected by project OGLE.

SMC720.28.40576 is the only object in the sample that reside in Magellanic Clouds. It's very luminous and hot, estimated parameters suggest it's blue giant with spectral type O according to single model SED fit 4.5. Estimated period of the variable is nearly half of day, such short period paired together with such high mass suggest very compact setup of stars. PARSEC evolutionary tracks estimate mass of object around $16 M_{\odot}$. Star was published already by (Pawlak et al., 2016) and classified as contact binary. Although no companion is visible in SED distribution it might hide itself in the light of primary star. Unfortunately only IRAC data is available, lack of WISE magnitudes can be problematic when one would like to look for possible companion star with low temperature. Although NEOWISE-R (ref) reports one particular measurement from close neighbourhood (< 1 arcsec) it lacks values from *W3* or *W4* bands which are crucial if one would like to search for colder companion. No radial velocity from Gaia is available so mass function estimate remains unknown.

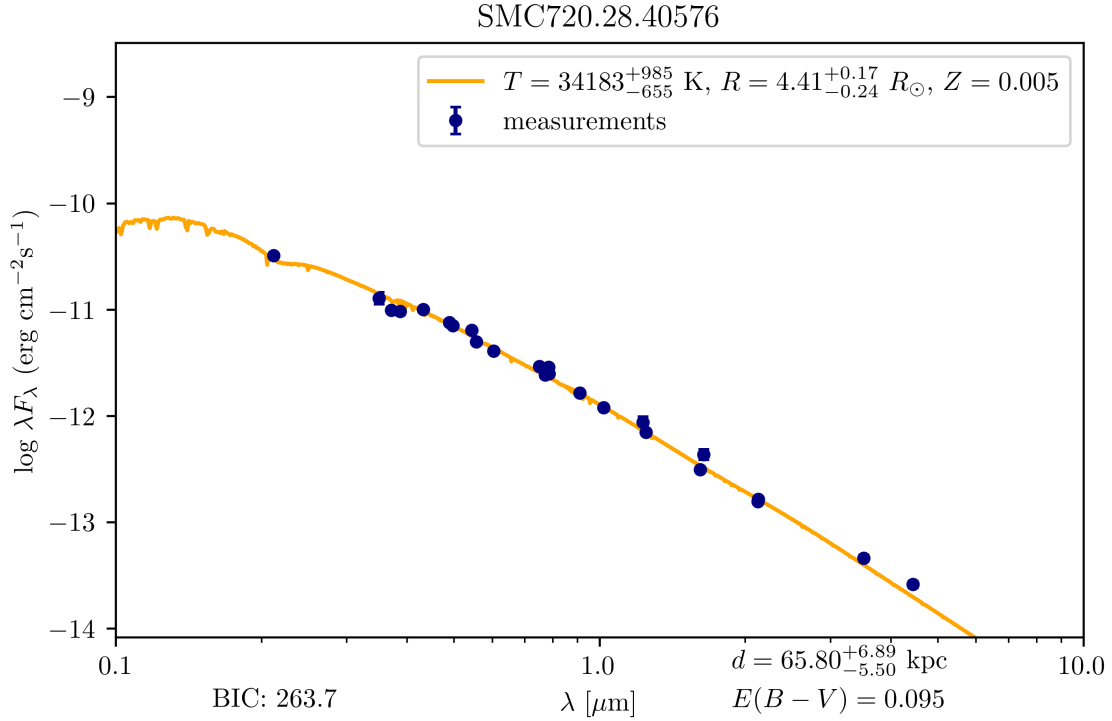


Figure 4.5: Single model SED fit for SMC720.28.40576.

Available data paired with MCMC allow to put some constraints on the parameters of objects. As primary star is supergiant second SED fit is performed assuming $\log g_1 = 4$ and $\log g_2 = 4$ in order to determine upper limit on companions luminosity. When MCMC chain isn't able to determine significance of some parameters it's reconstructing prior distribution which results in uniformly distributed region of samples. Results of MCMC sampling was plotted over samples from HYG database¹ and are presented on figure 4.6

¹<https://www.astronexus.com/hyg>

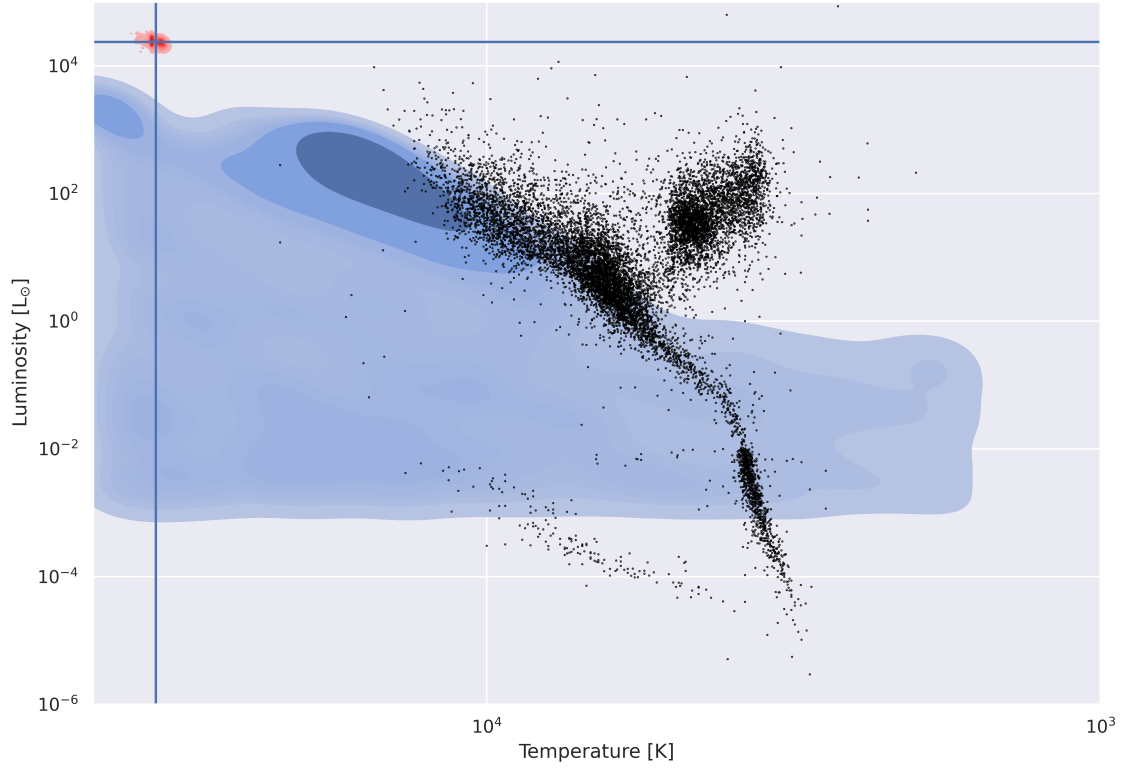


Figure 4.6: Posterior estimate for parameters of SMC720.28.40576 highlighted in red (primary component) and in blue (secondary component).

Although we can rule out any objects on giant branch many regions on HR diagram can host potential secondary companion. Most importantly few regions of Main Sequence stars cannot be ruled out giving plausible candidate for secondary companion.

4.3.3. Rest of objects

. TO DO

Chapter 5

Summary

TO DO

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Appendices

Appendix A

Semi-amplitude estimate from Gaia DR3

Let's now assume, that this radial velocity can be factored into two separate movements: center of mass movement (with velocity v_0) and circular motion with semi-amplitude K . Under following assumptions we can find, that samples of radial velocity can be written as

$$V_i = v_0 + K \cos(2\pi X_i) \quad (\text{A.1})$$

where X_i is sample from $\mathcal{U}(0, 1)$. This allows to determine, that if we write down our statistics of σ^2 we get

$$\sigma^2 = \frac{1}{n} \sum_i (\bar{V} - V_i)^2 \approx \frac{K^2}{N} \sum \cos^2(2\pi X_i) = \frac{K^2}{N} \sum_i \frac{\cos(4\pi X_i) + 1}{2} \quad (\text{A.2})$$

where we used the simplification that $\bar{V} \approx v_0$. Now we would like to obtain some statistical properties of this variance estimator such as expected value and variance. We can clearly see that $\mathbb{E} \cos(4\pi X_i) = 0$ while $\widehat{\text{Var}}(\cos 4\pi X_i) = \frac{1}{8}$. While first value is straightforward as mean of cosine is equal to zero second one is more complicated but follows from the fact that distribution of $\cos(4\pi X_i)$ is arcsine distributed and it's variance it's well known. Hence we see, that

$$\mathbb{E}\sigma^2 = \frac{K^2}{2N} \cdot N = \frac{K^2}{2} \quad (\text{A.3})$$

$$\widehat{\text{Var}}(\sigma^2) = \frac{K^4}{32N^2} \cdot N = \frac{K^4}{32N} \quad (\text{A.4})$$

Those two values allow us to determine quality of our estimator, using Central Limit Theory one can also approximate that

$$\sigma^2 \sim \mathcal{N}\left(\frac{K^2}{2}, \frac{K^4}{32N}\right) \quad (\text{A.5})$$

which can be quite useful to estimate K from obtained data (for most of the objects $N > 17$ so Central Limit Theorem gives good approximation). There are many simplification made here which are mainly based on the fact it's not entirely know how Gaia obtains RV measurements and what is measured after all. If we assume, that true velocity is contaminated by gaussian noise with standard deviation σ_{noise} estimated variance will be shifted and

$$\mathbb{E}\sigma^2 = \frac{K^2}{2} + \sigma_{noise}^2 \quad (\text{A.6})$$

so we will overestimate K .