Evolution of galaxy dynamics over the last 10 Gyrs with MUSE/VLT

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Galaxy evolution through time

Why studying kinematics ?

Explain evolution of morphology with kinematics

- ▷ main processes responsible for disc formation
- ▷ impact of merging, inflows, outflows can be measured
- ▷ kinematics → rotation curve → dark matter distribution

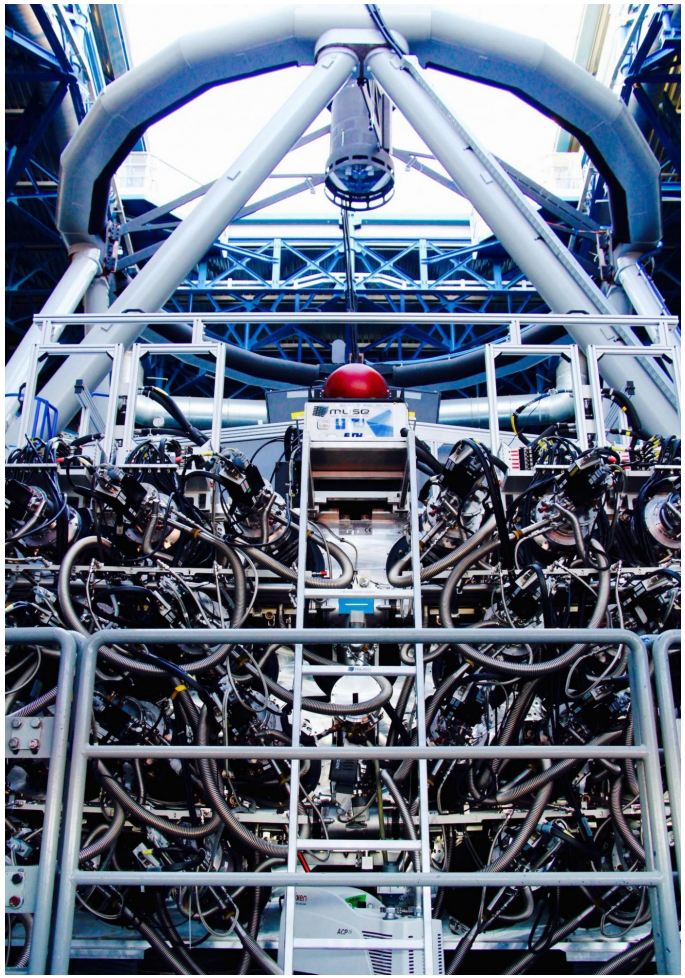
Integral Field Spectroscopy & MUSE

IFS:

- ▷ 3D cubes (2D spatial + 1D spectral)
- ▷ photometry + kinematics

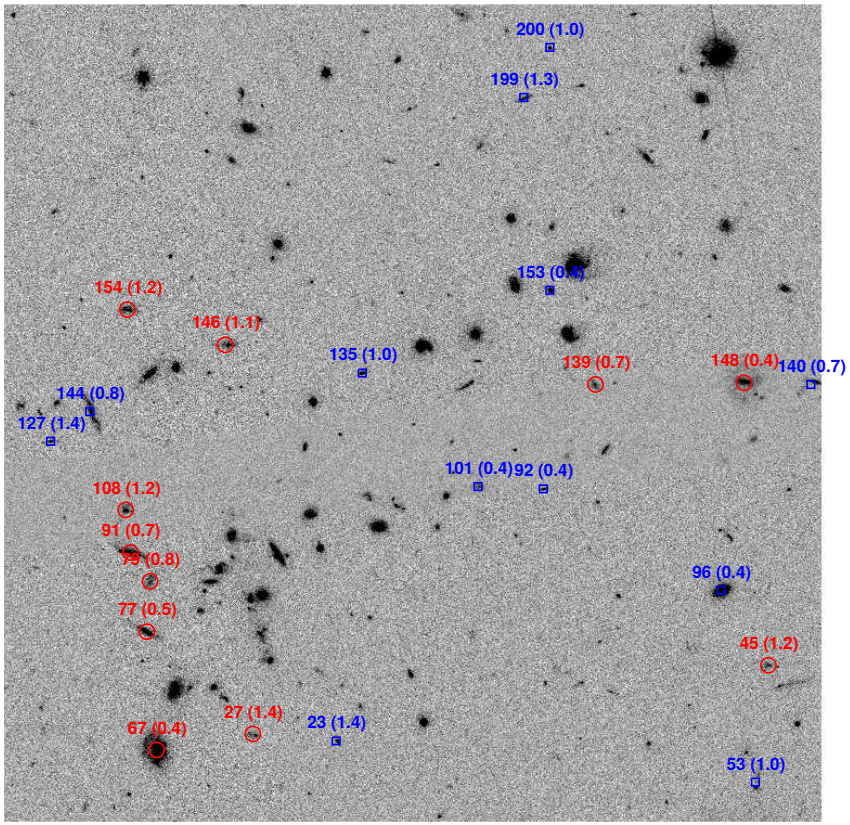
MUSE:

- ▷ $1 \times 1\,\mathrm{arcmin}^2$ FoV
- ▷ 0.2 arcsec spatial sampling
- ▷ spectral range $[4650\,\text{Å}, 9300\,\text{Å}]$
- ▷ seeing-limited or AO observations



MUSE instrument. Credit: Contini Thierry (IRAP)

Our sample



HST image of COSMOS group CGr30

- ▷ 16 MUSE fields in COSMOS area
- ▷ exposures from 1 to 10 hr
- ▷ seeing-limited (FWHM $\lesssim 0.7$”) or AO (FWHM $\lesssim 0.5$”)
  - · *deep* and *best_seeing* observations
- ▷ ~ 500 field galaxies with [OII] detection
  - · HST-ACS counterparts
  - · $0.4 \leq z \leq 1.4$

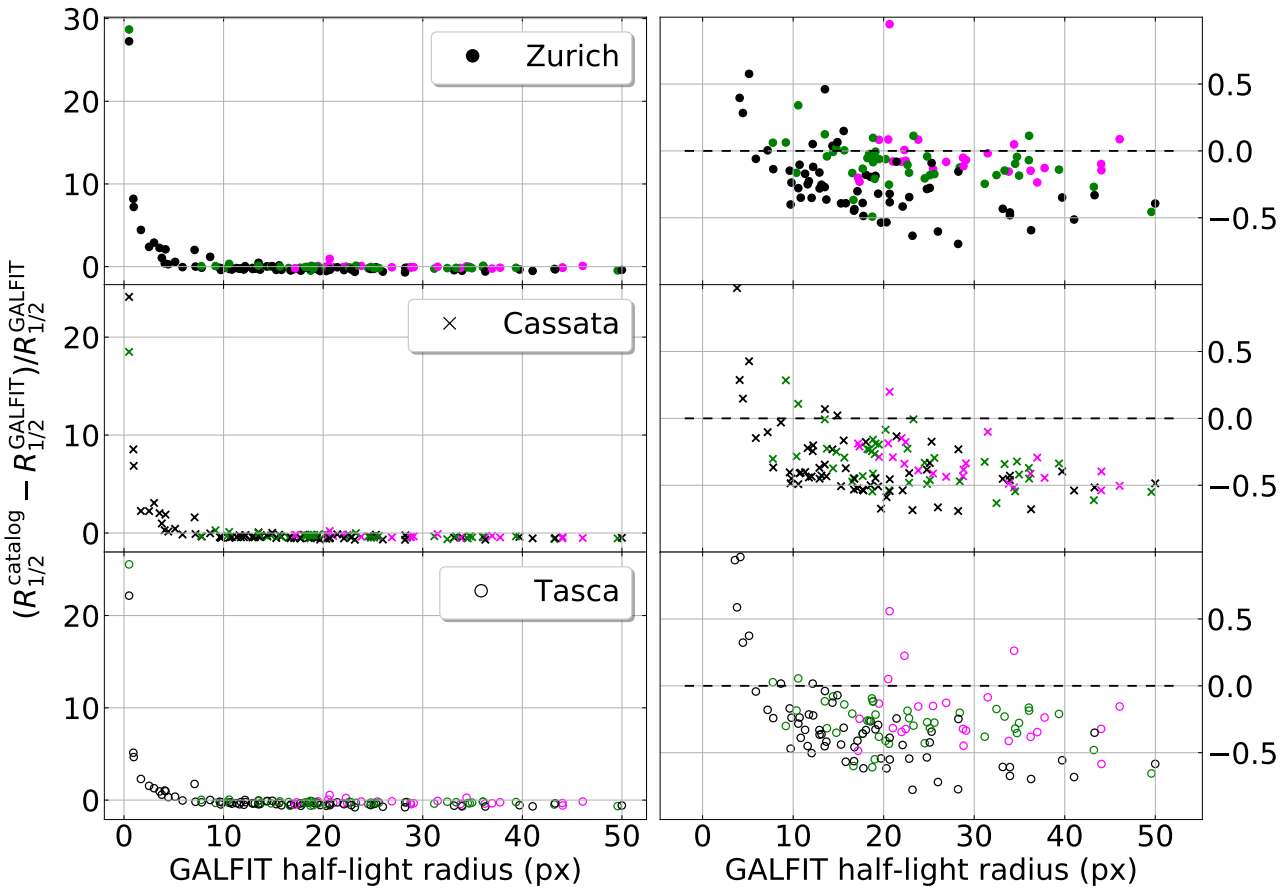
Methodology

I. Morphological information
   - ▷ **half-light radius** $R_{1/2}$ to select resolved galaxies
     - · Cassata, Tasca and Zurich catalogues
   - ▷ **ellipticitiy** to compute the inclination
     - · fixed input for the kinematical model

II. Kinematical modelling
   - ▷ recover $V_{\rm max}$ and $\sigma_{\rm v}$

III. Tully-Fisher relation

Checking morphological parameters
Half-light radius

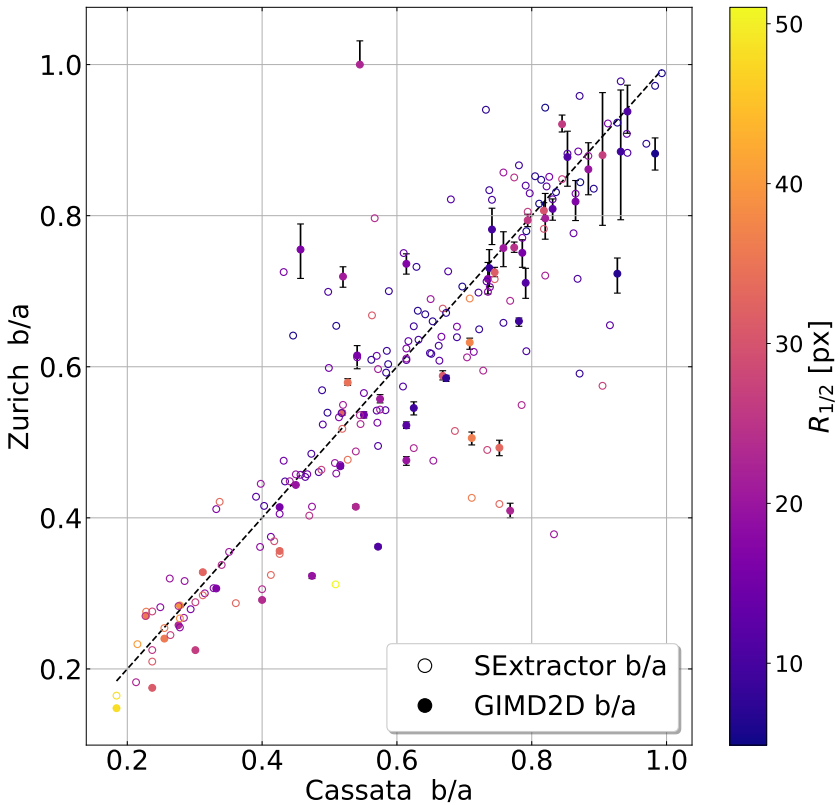


GALFIT run by V. Abril-Melgajero (LAM) on structure galaxies

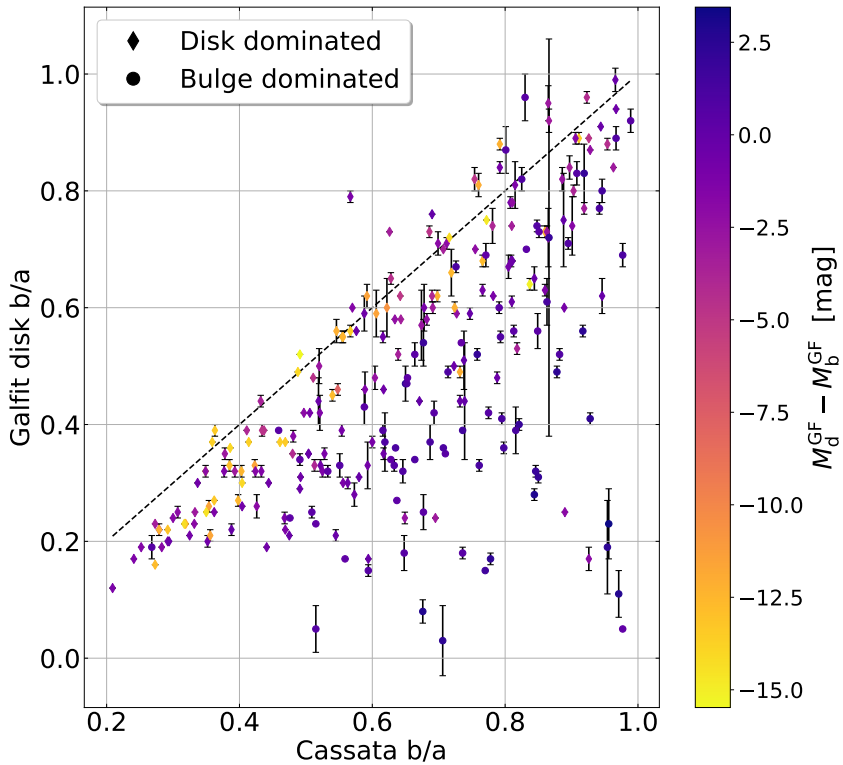
- ▷ GALFIT radius used as a reference

spheroidal disk-like irregulars

Checking morphological parameters
Ellipticity

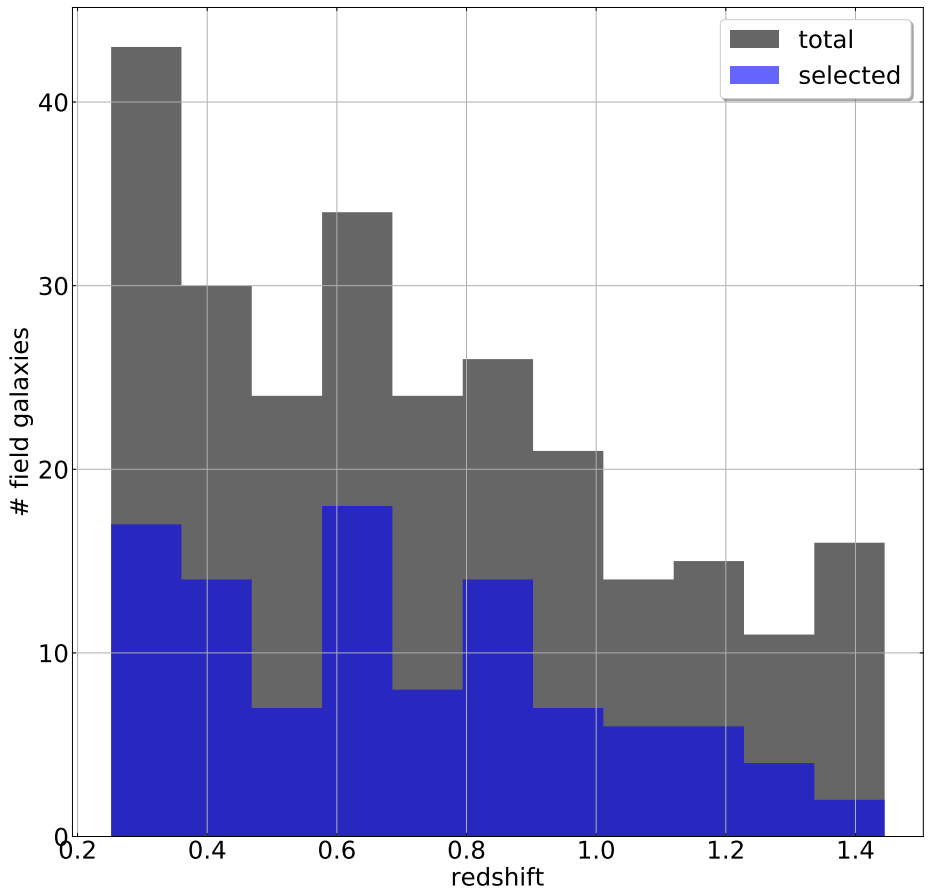


▷ values are consistent between catalogues



▷ scatter is due to bulge dominated (spherically symmetric) systems

Characteristics of our sample
Redshift distribution

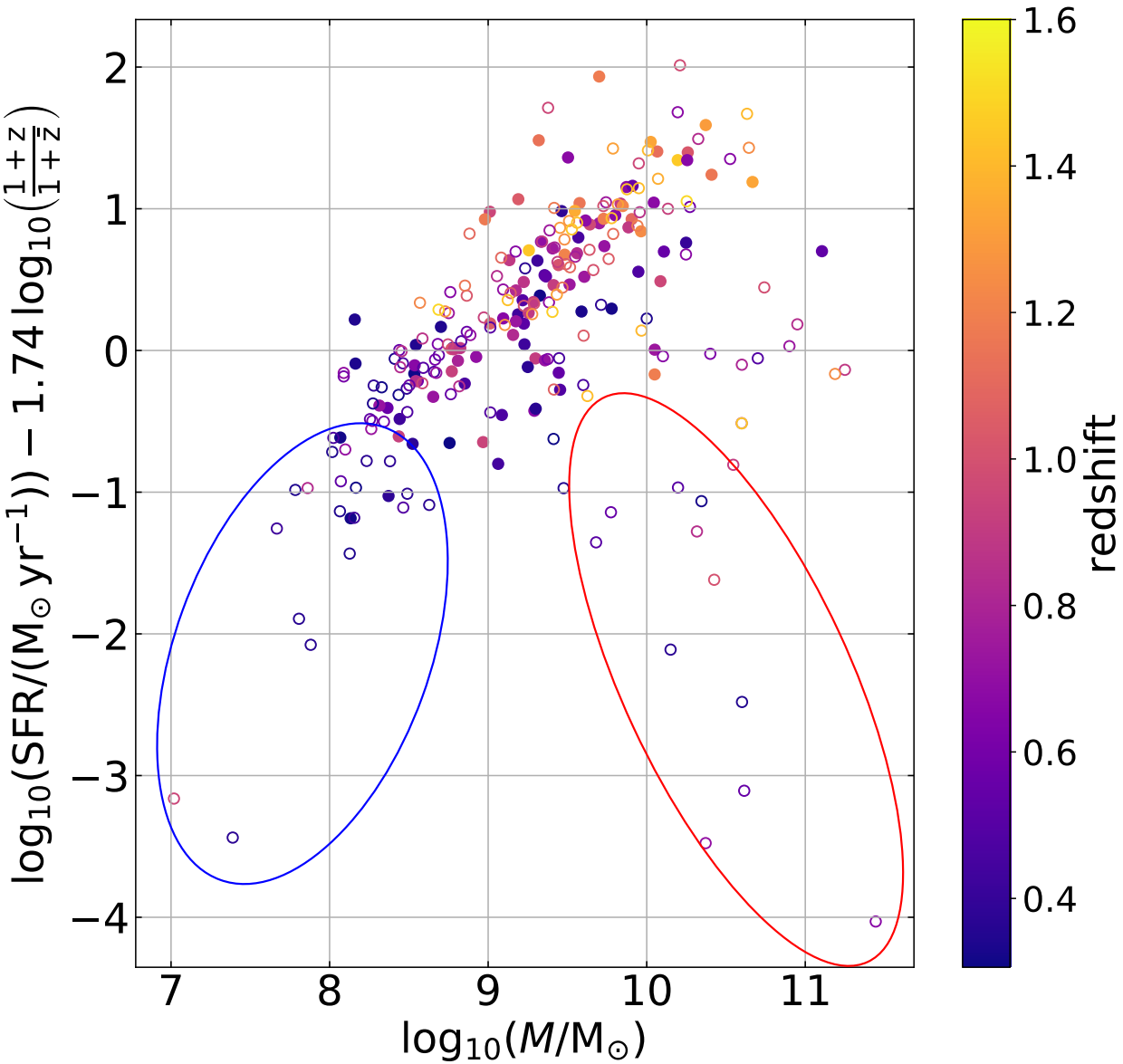


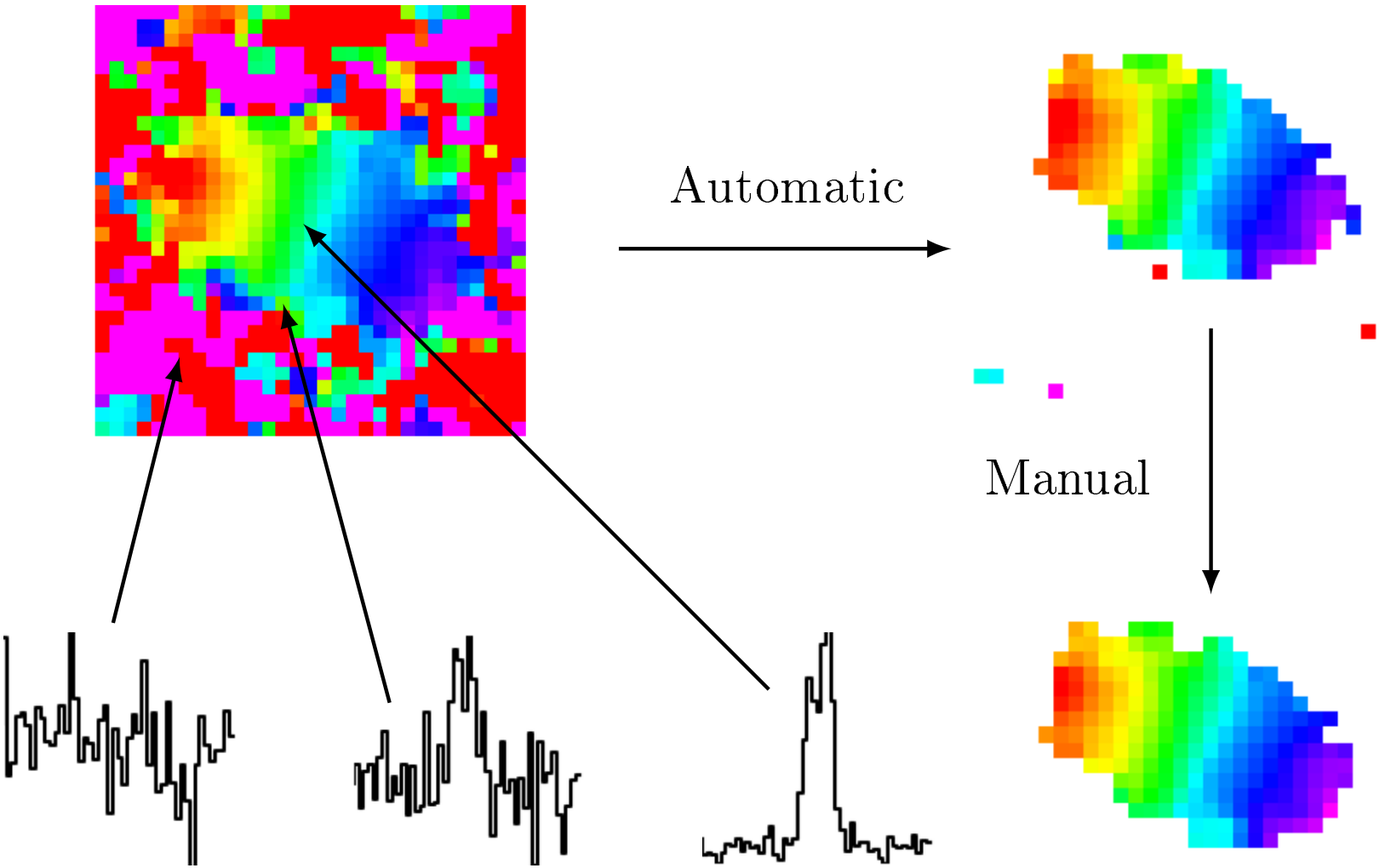
The total number corresponds to galaxies with photometric data in Cassata and/or Zurich catalogues.

- ▷ sample of **103 galaxies** with $R_{1/2} > 0.35$" and SNR > 5
- ▷ we loose galaxies at $z \approx 1.4$
- ▷ redshift distribution is not drastically changed

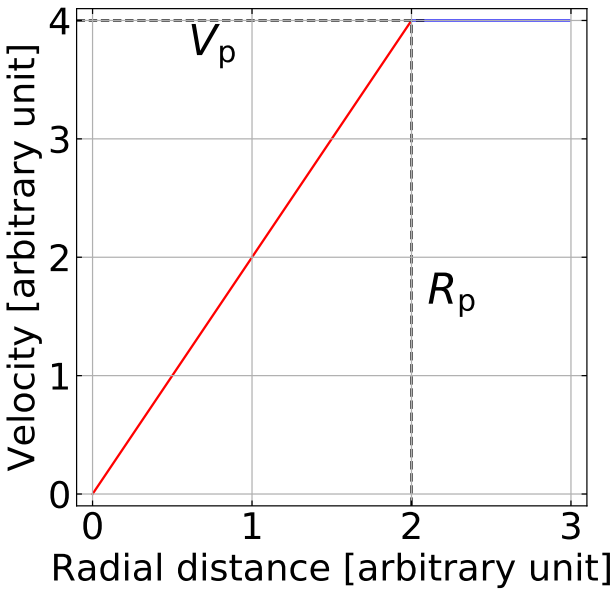
Characteristics of our sample
Mass-SFR relation

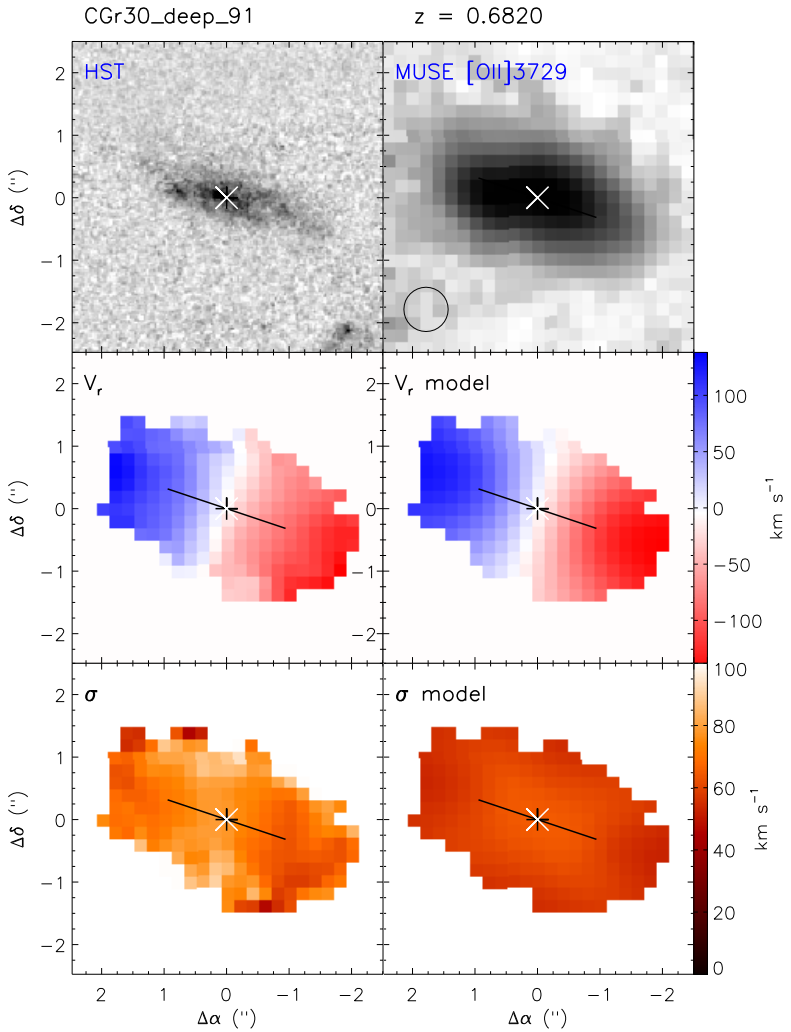
- ▷ Most of our sample galaxies are on the main sequence
- ▷ massive quiescent (low [OII]) and very low mass galaxies are lost

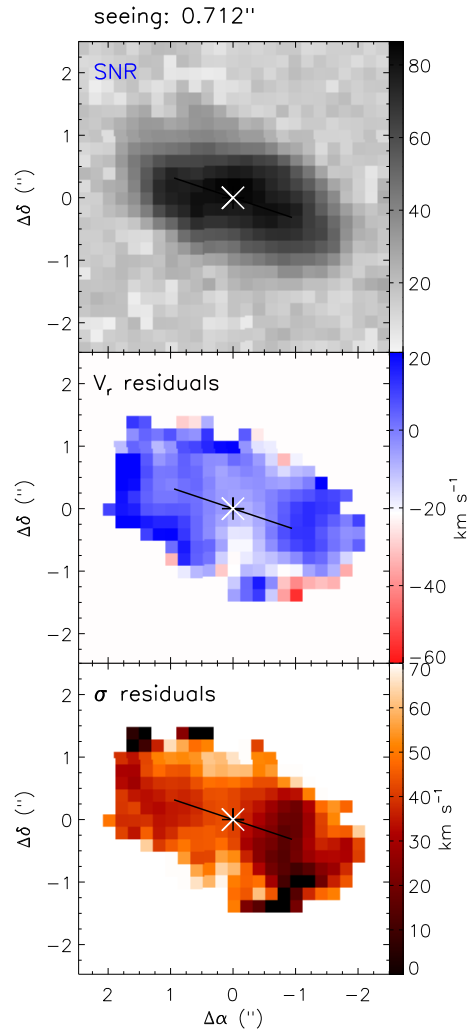



Automatic
Manual

Use a ramp model with a linear internal variation and a plateau.







First results
$V_{\mathrm{max}}/\sigma_{\mathrm{v}}$ distribution

Perspectives

Short term:

- ▷ Tully-Fisher evolution with z
- ▷ Variation between morphological and kinematical PA ?

Long term:

- ▷ Morphological modelling
- ▷ Angular momentum evolution ?
- ▷ Dark matter vs. luminous mass ?
- ▷ **better selection + larger sample**