

A ‘How To’ Guide To Simulating The Universe

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Public Astronomy Lecture

27-05-2022

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EAGLE: Evolution and Assembly of GaLaxies and their Environments

The evolution of intergalactic gas. Colour encodes temperature

$z = 19.8$
 $t = 0.2 \text{ Gyr}$
 $L = 25.0 \text{ cMpc}$

Simulation by the EAGLE collaboration
Visualisation by Jim Geach & Rob Crain



Image Credit: <https://www.epcc.ed.ac.uk>



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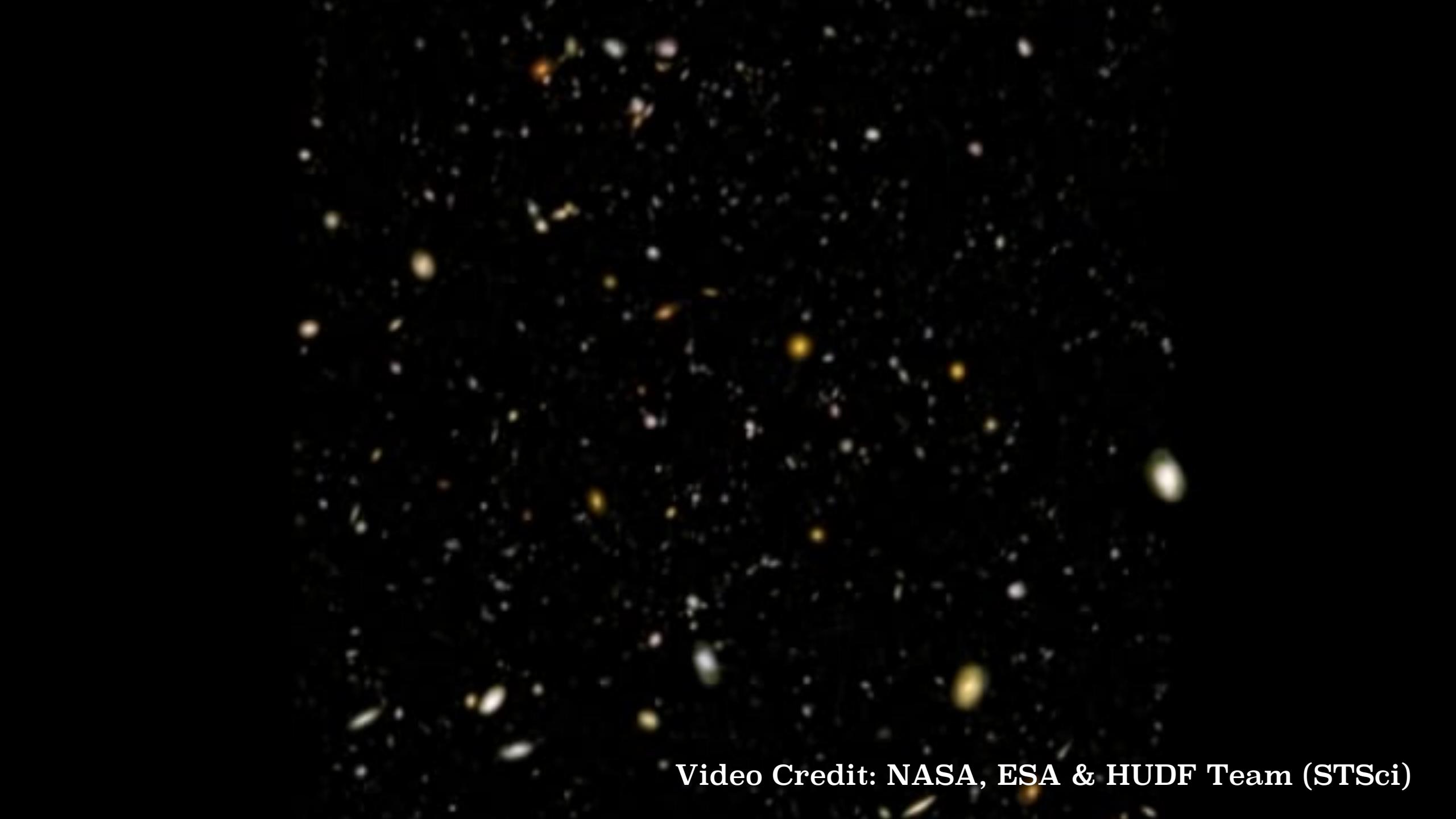
OzSTAR
Supercomputing

The necessary bits required to simulate the Universe

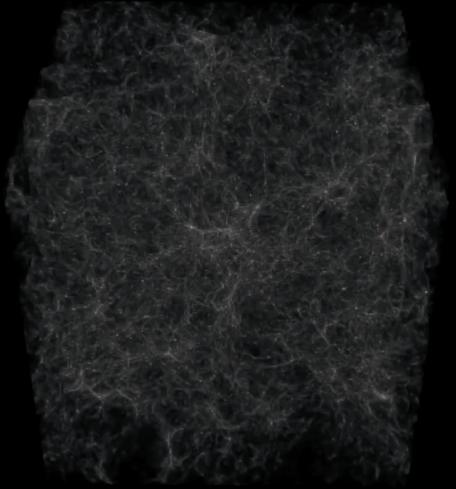
0. Gravity
1. Cosmic Microwave Background
2. Dark Matter
3. Atoms
4. Stars
5. Black Holes
6. Feedback



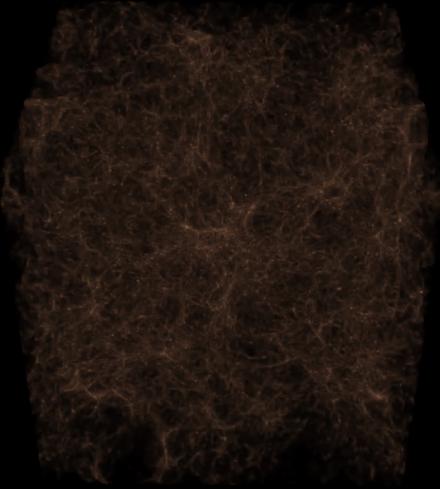
Image Credit: Ángel R. López-Sánchez (AAO-MQ)



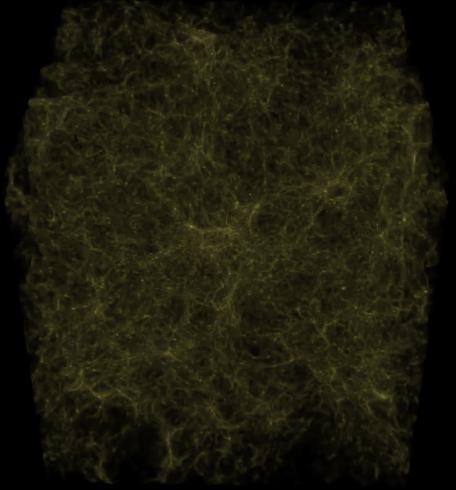
Video Credit: NASA, ESA & HUDF Team (STScI)



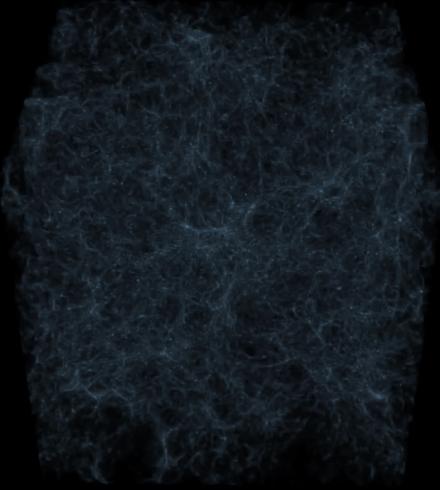
Normal Universe



No Supermassive Black Holes



Strong Supermassive Black Holes



Stronger Supernovae

Large volume (statistics)

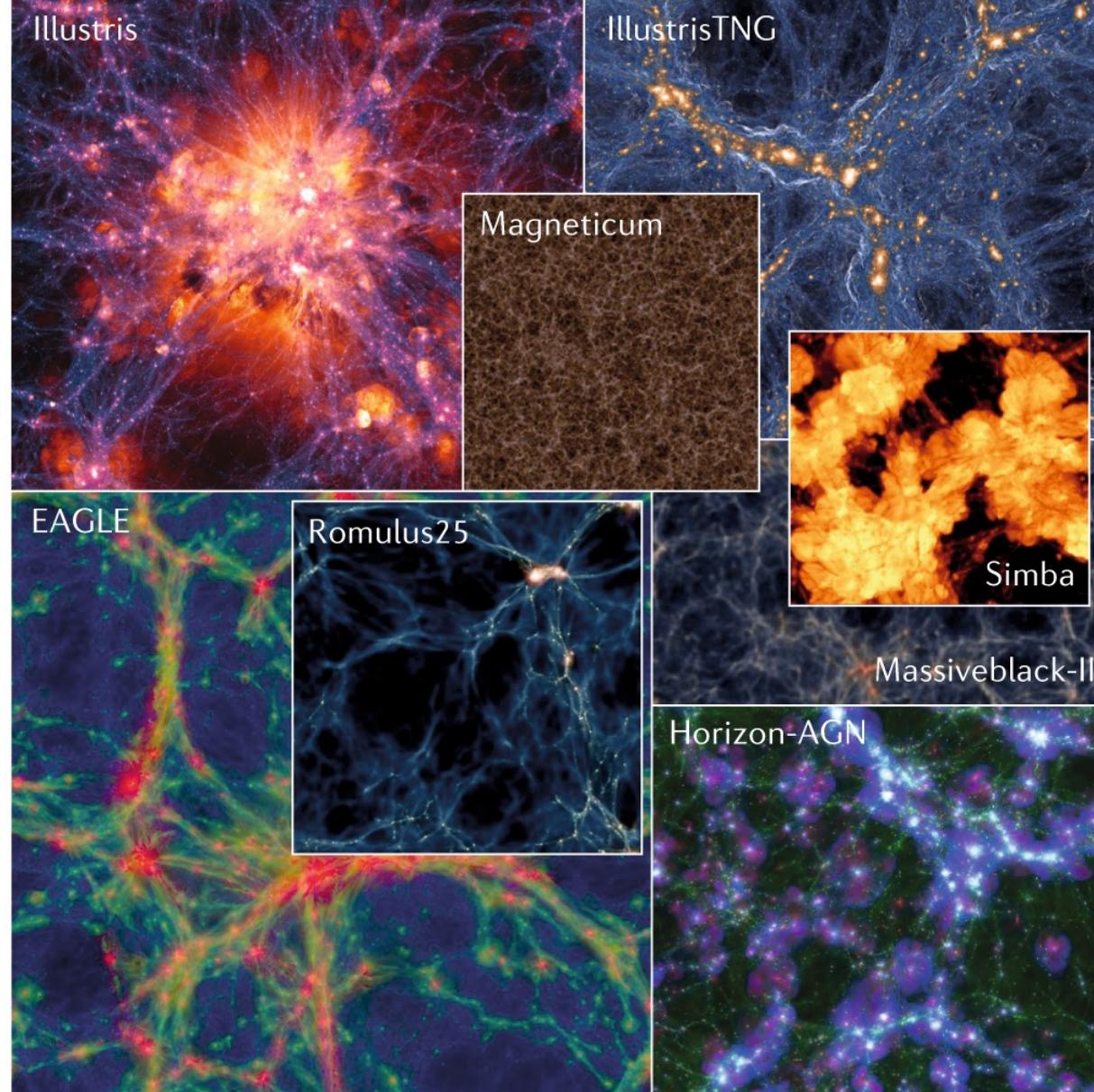
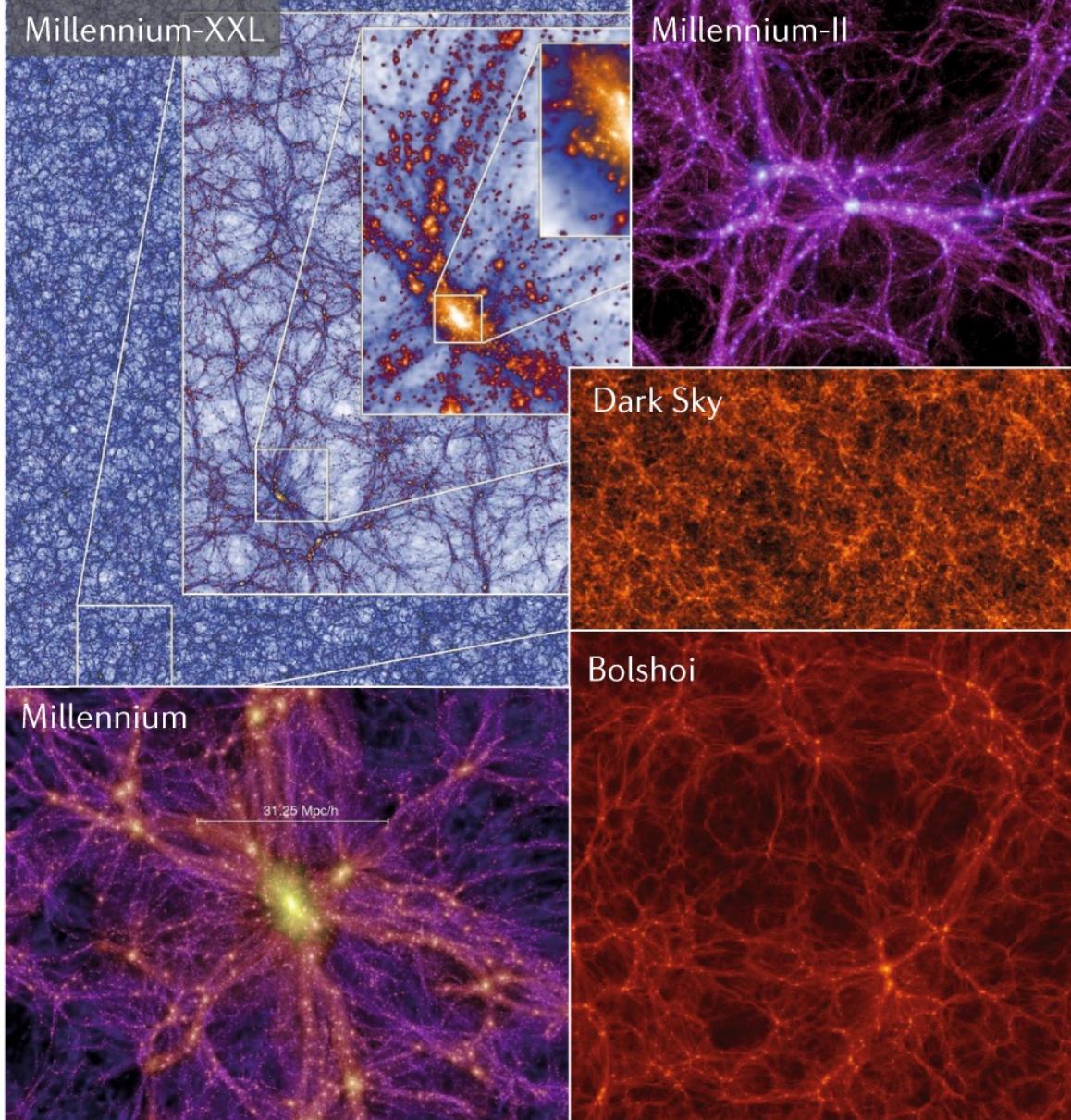


Image Credit: Vogelsberger et al. (2020)

0. Gravity

0. Gravity

General Relativity:

$$G_{\mu\nu} + \Lambda g_{\mu\nu} = \frac{8\pi G}{c^4} T_{\mu\nu}$$

0. Gravity

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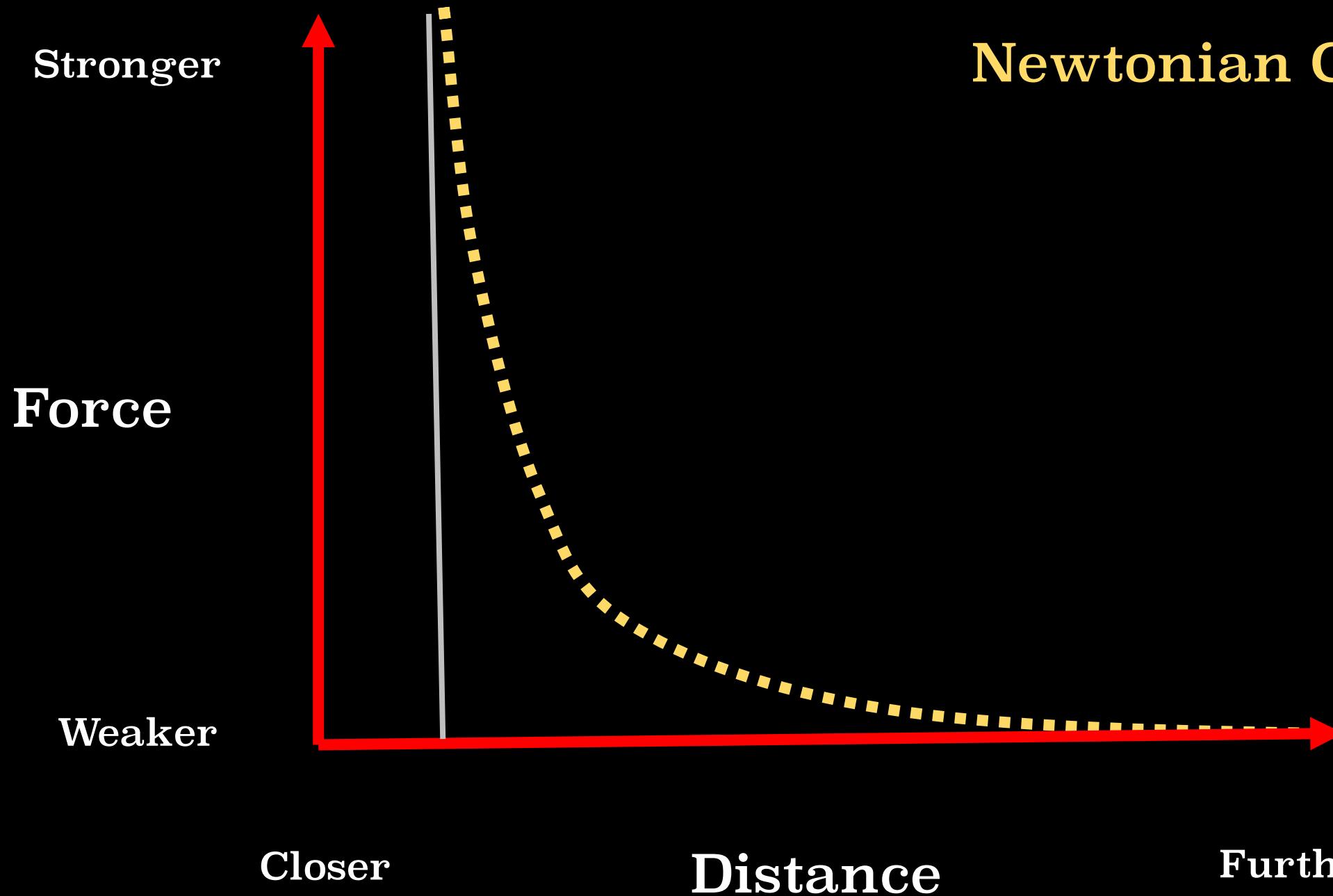
0. Gravity

General Relativity:

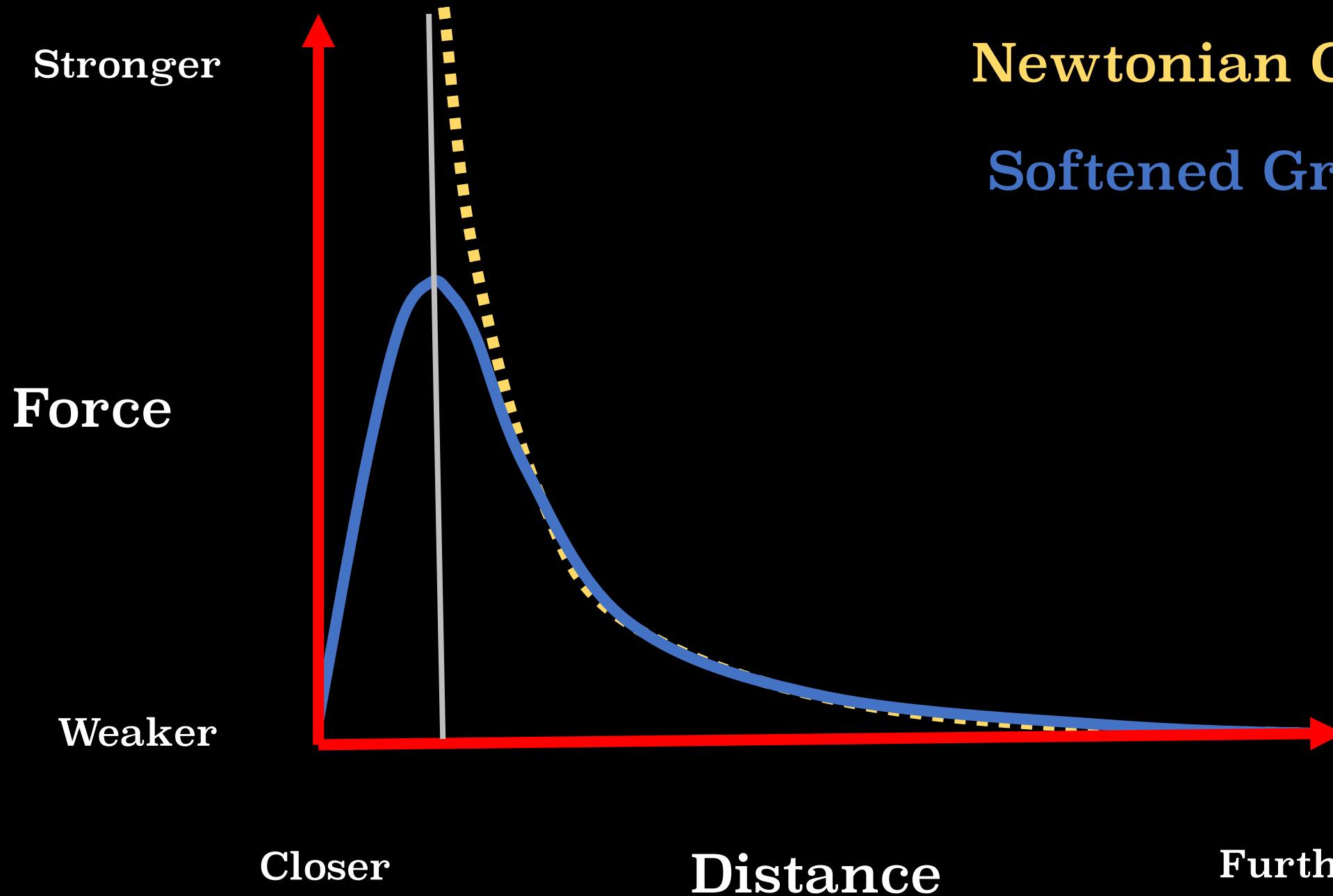
$$G_{\mu\nu} + \Lambda g_{\mu\nu} = \frac{8\pi G}{c^4} T_{\mu\nu}$$

Newtonian Gravity:

$$F = -\frac{Gm_1m_2}{r^2}$$



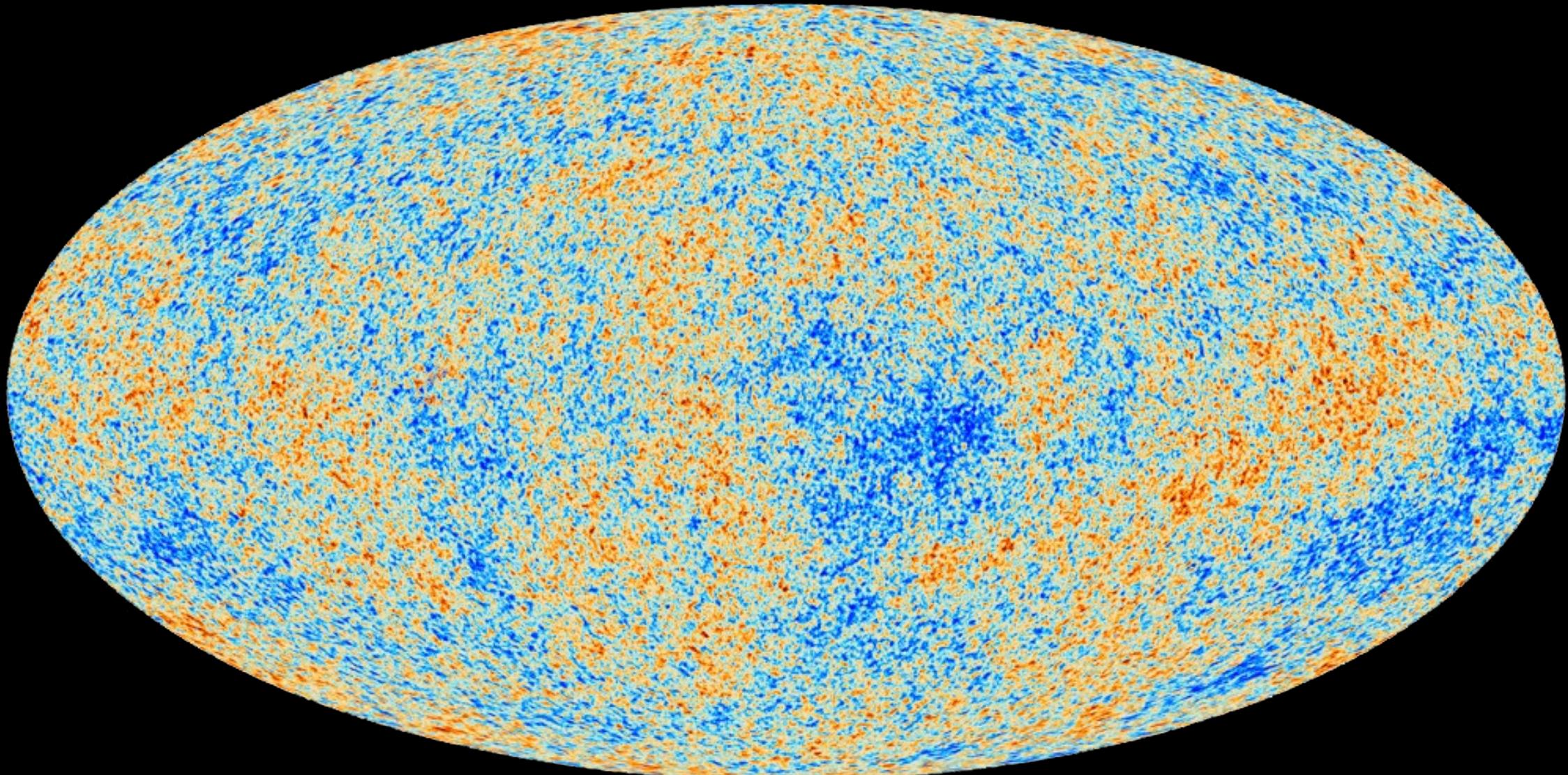
Newtonian Gravity



Where do we start
everything?

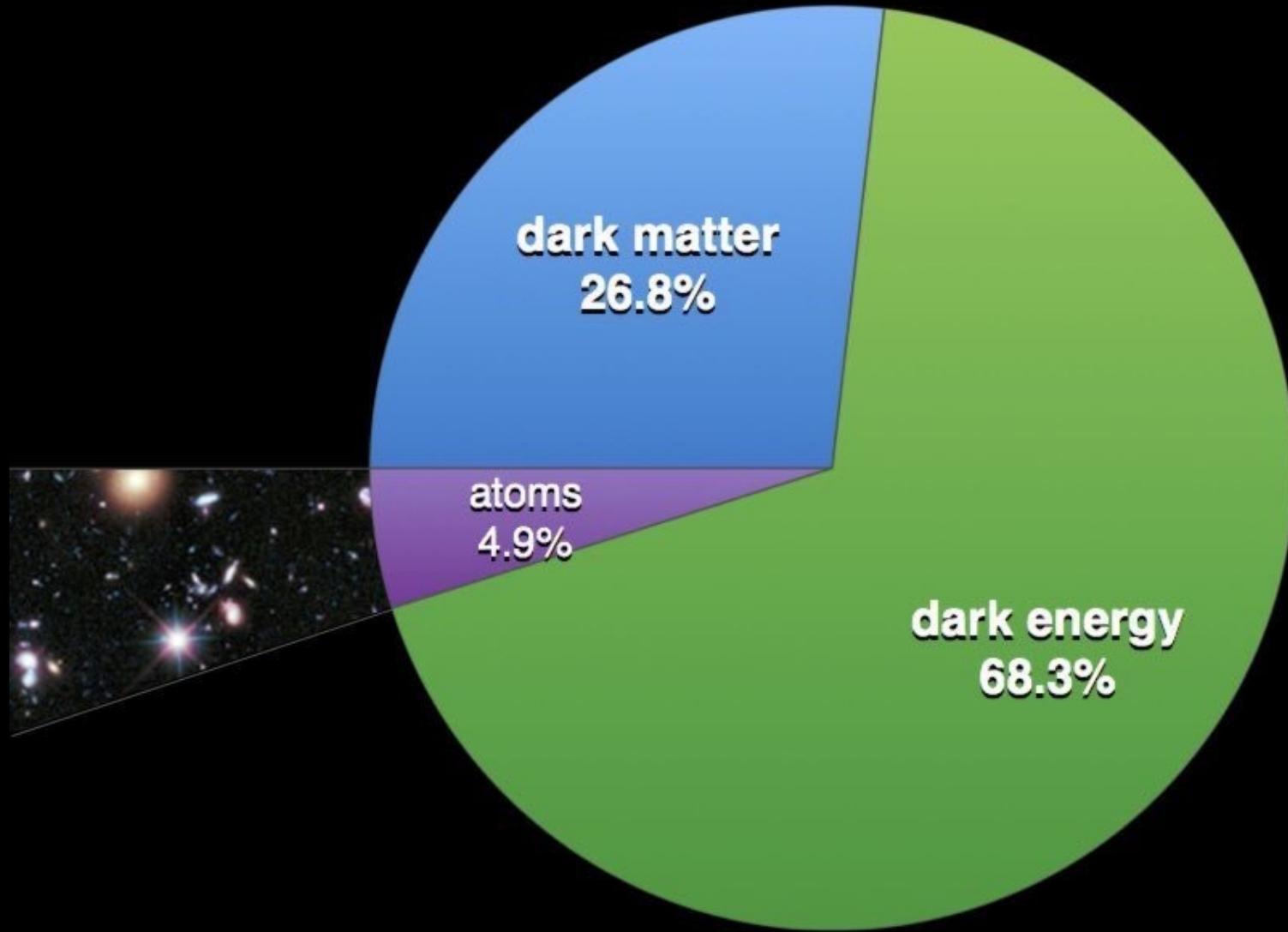
1. Cosmic Microwave Background

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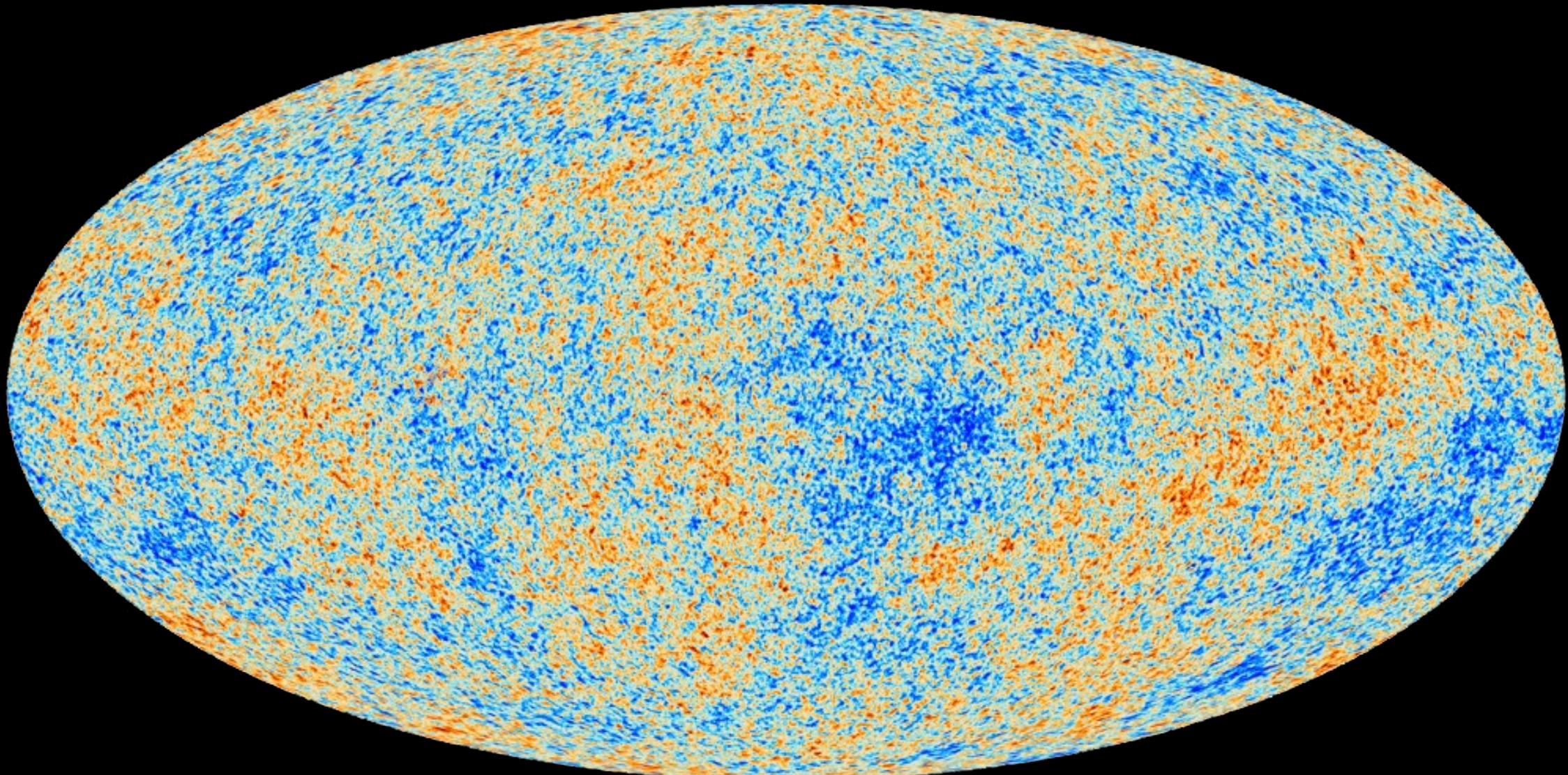


2. Dark Matter





1. Cosmic Microwave Background





planck CMB Simulator

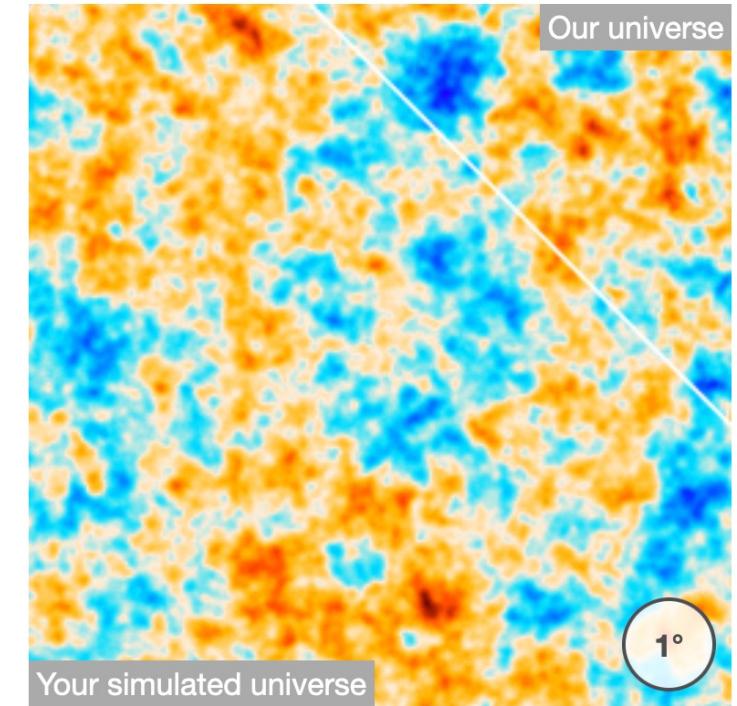


Normal Matter ($\Omega_b = 0.05$)

Dark Matter ($\Omega_c = 0.275$)

Dark Energy ($\Omega_\Lambda = 0.675$)

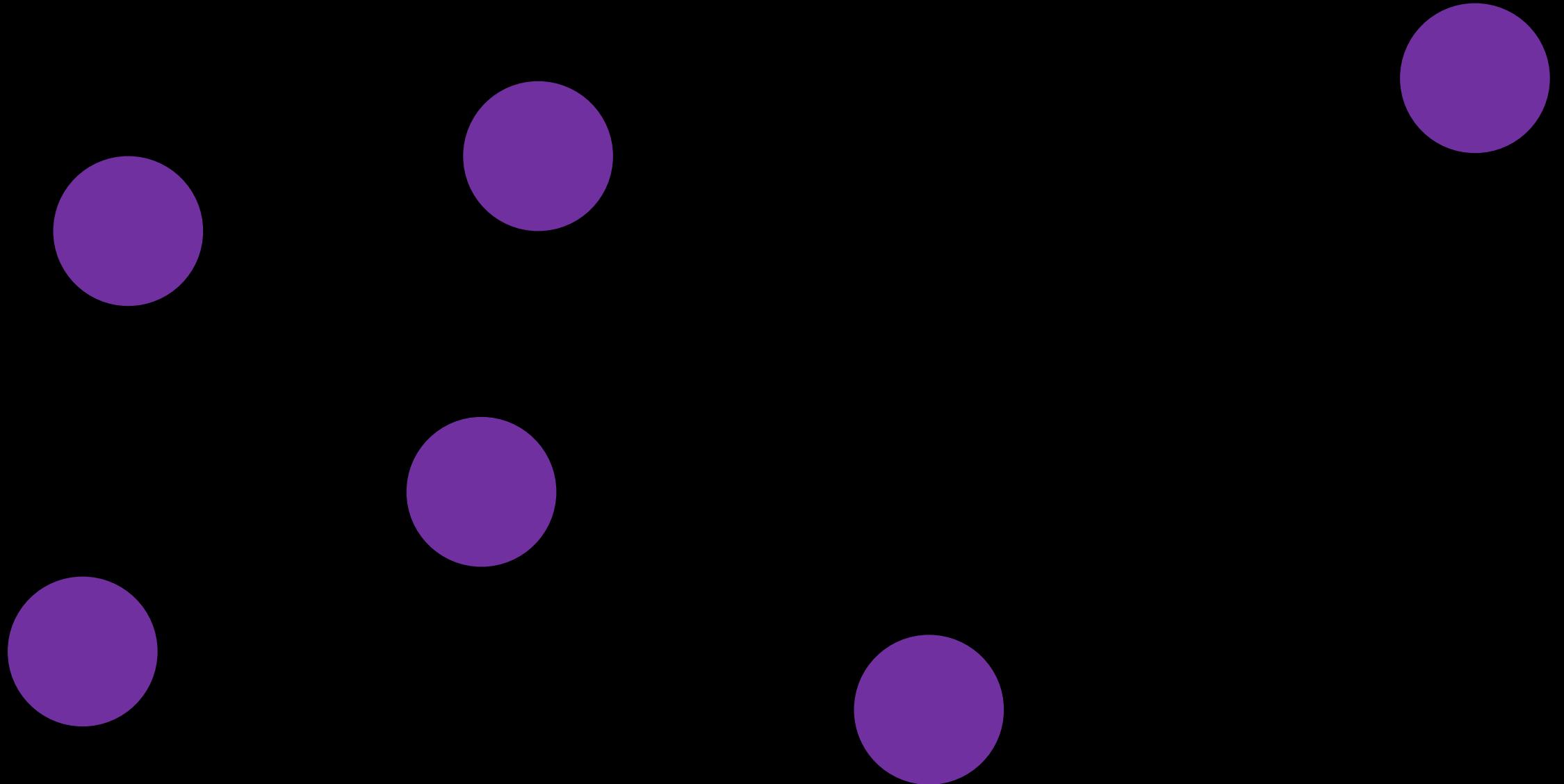
Normal matter only



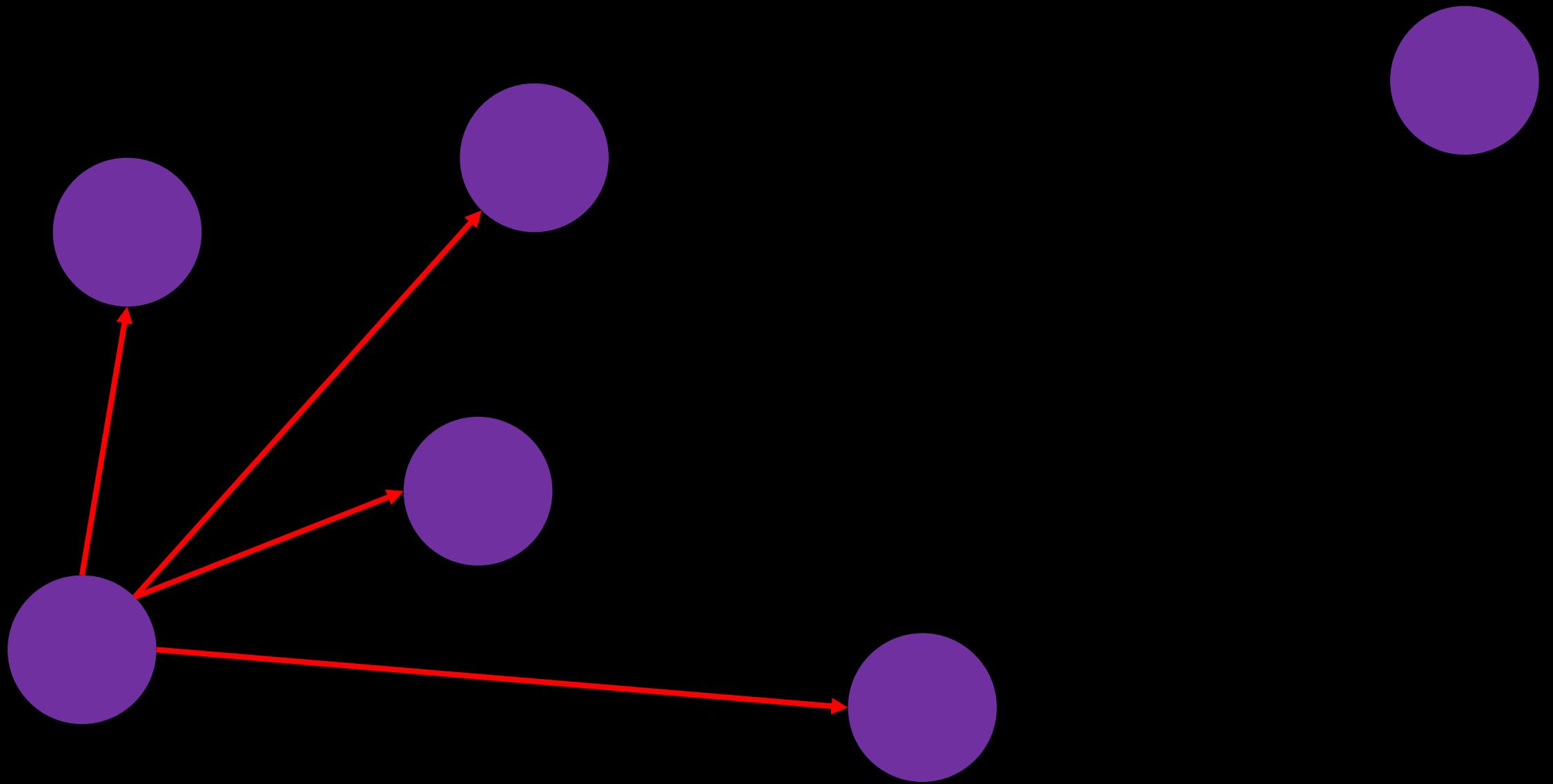
Your simulated universe

13.8 billion years old - just right
flat universe
Fundamental scale ~0.8°
Universe similarity **100%** - the same as our universe

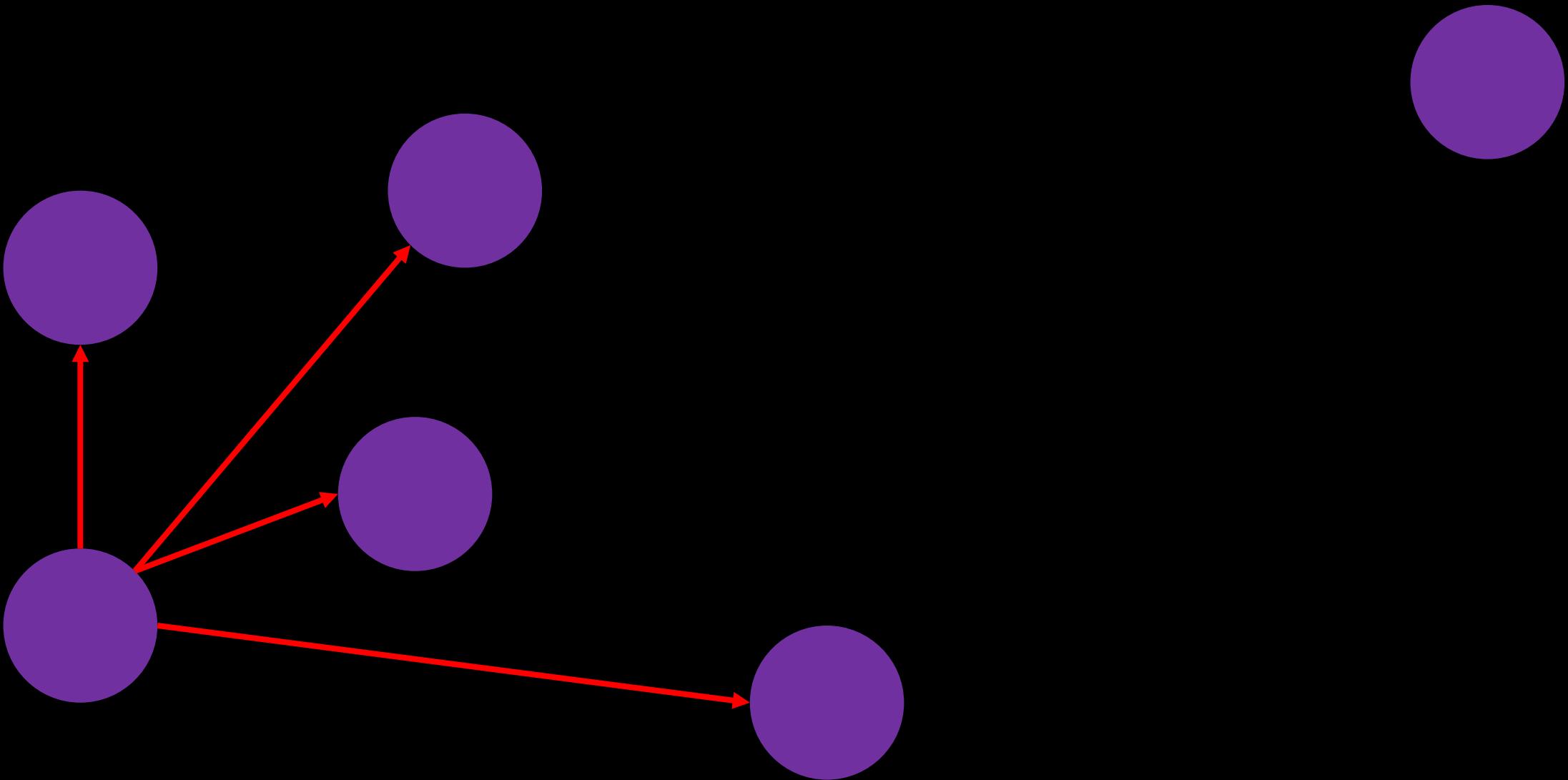
<https://plancksatellite.org.uk/cmb-sim/>



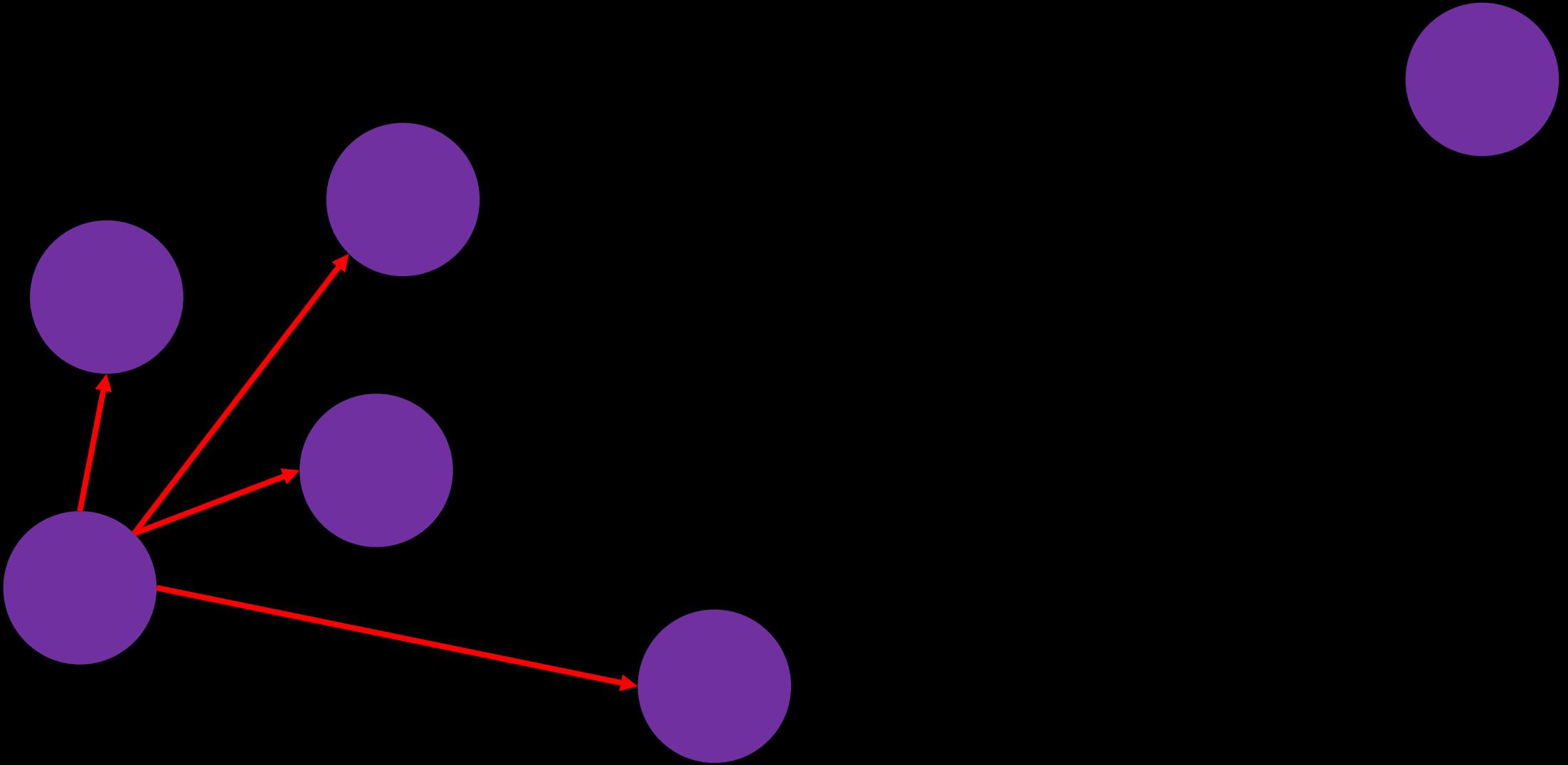
N-Body Problem



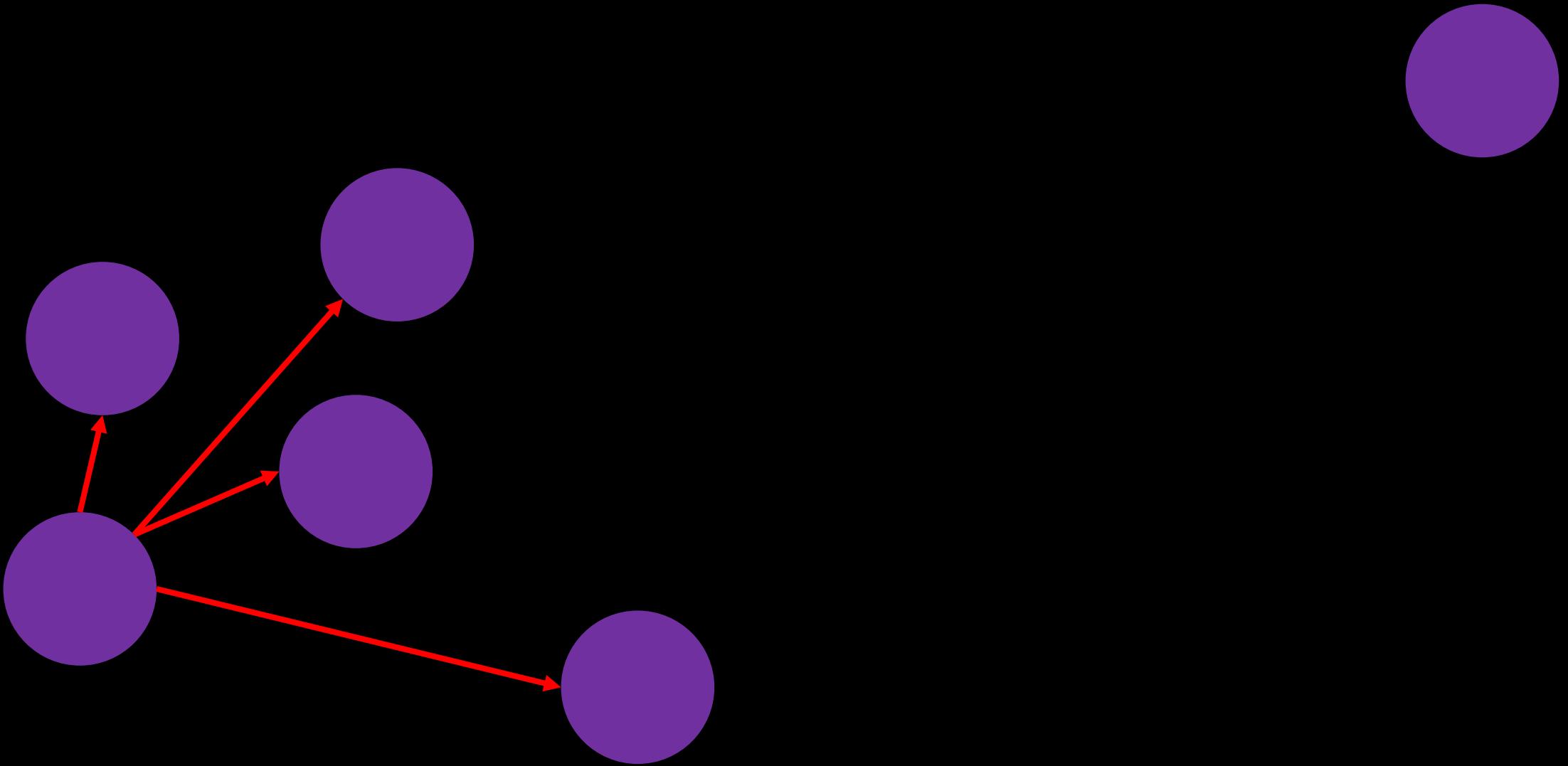
N-Body Problem



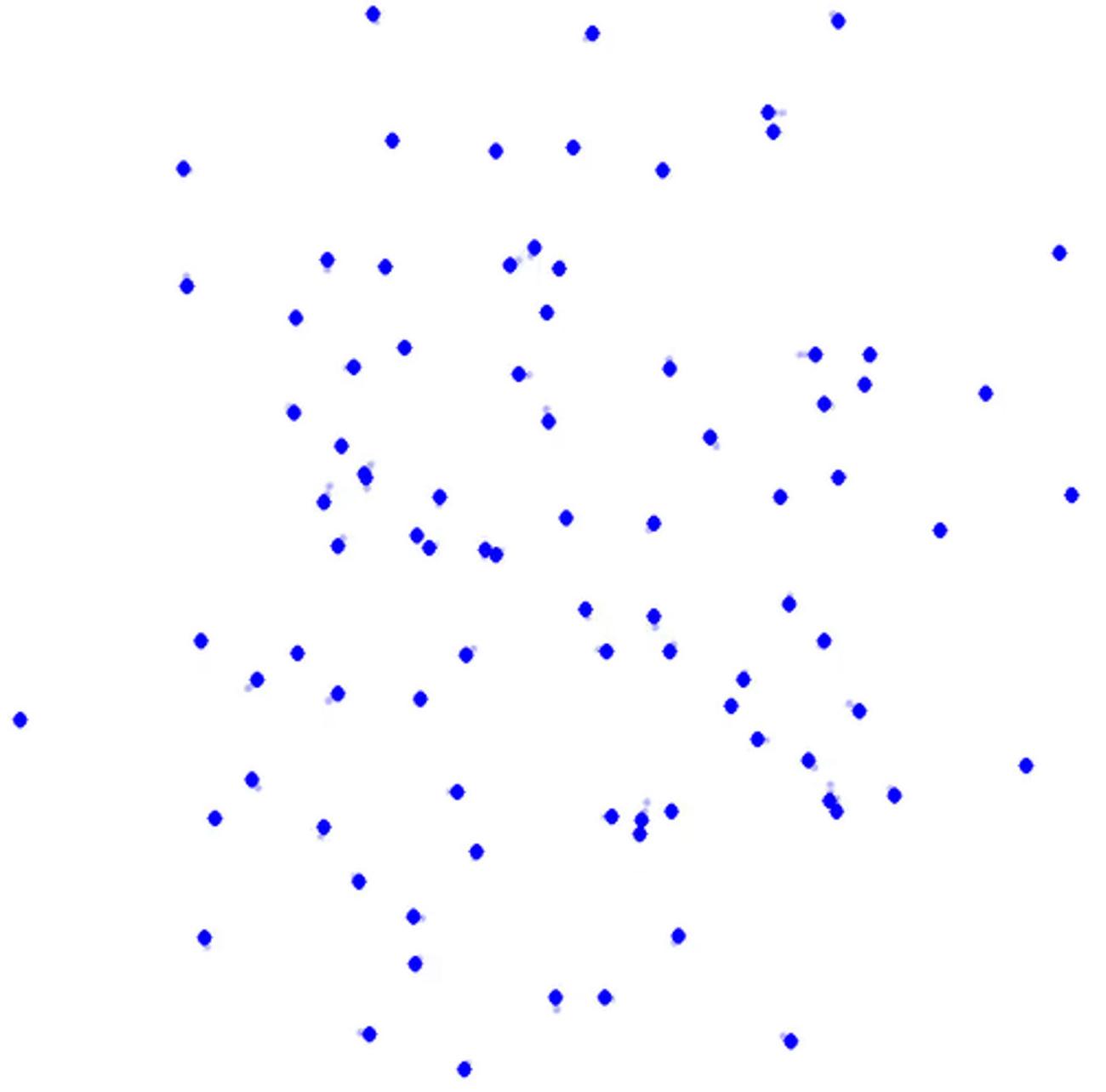
N-Body Problem



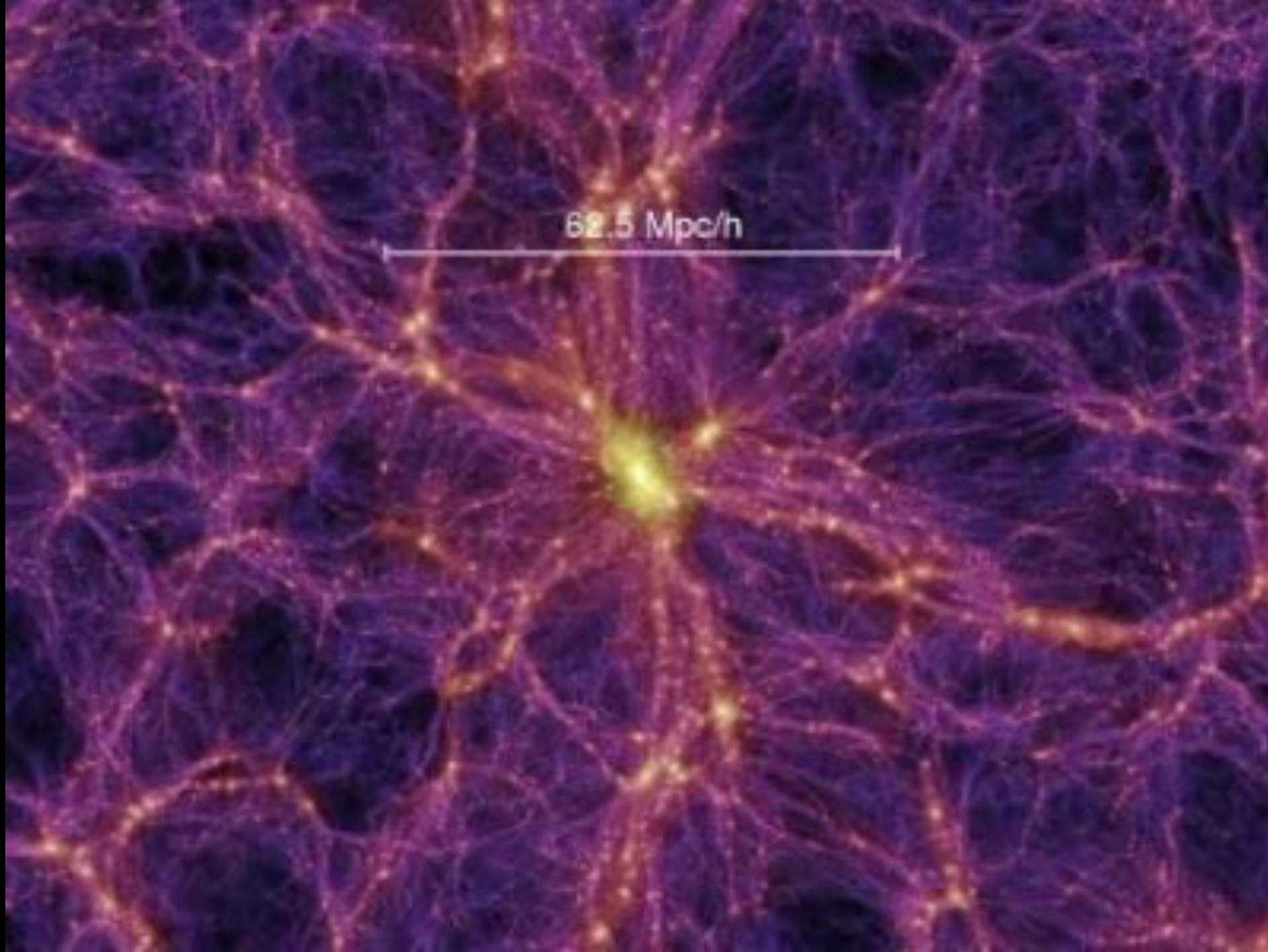
N-Body Problem



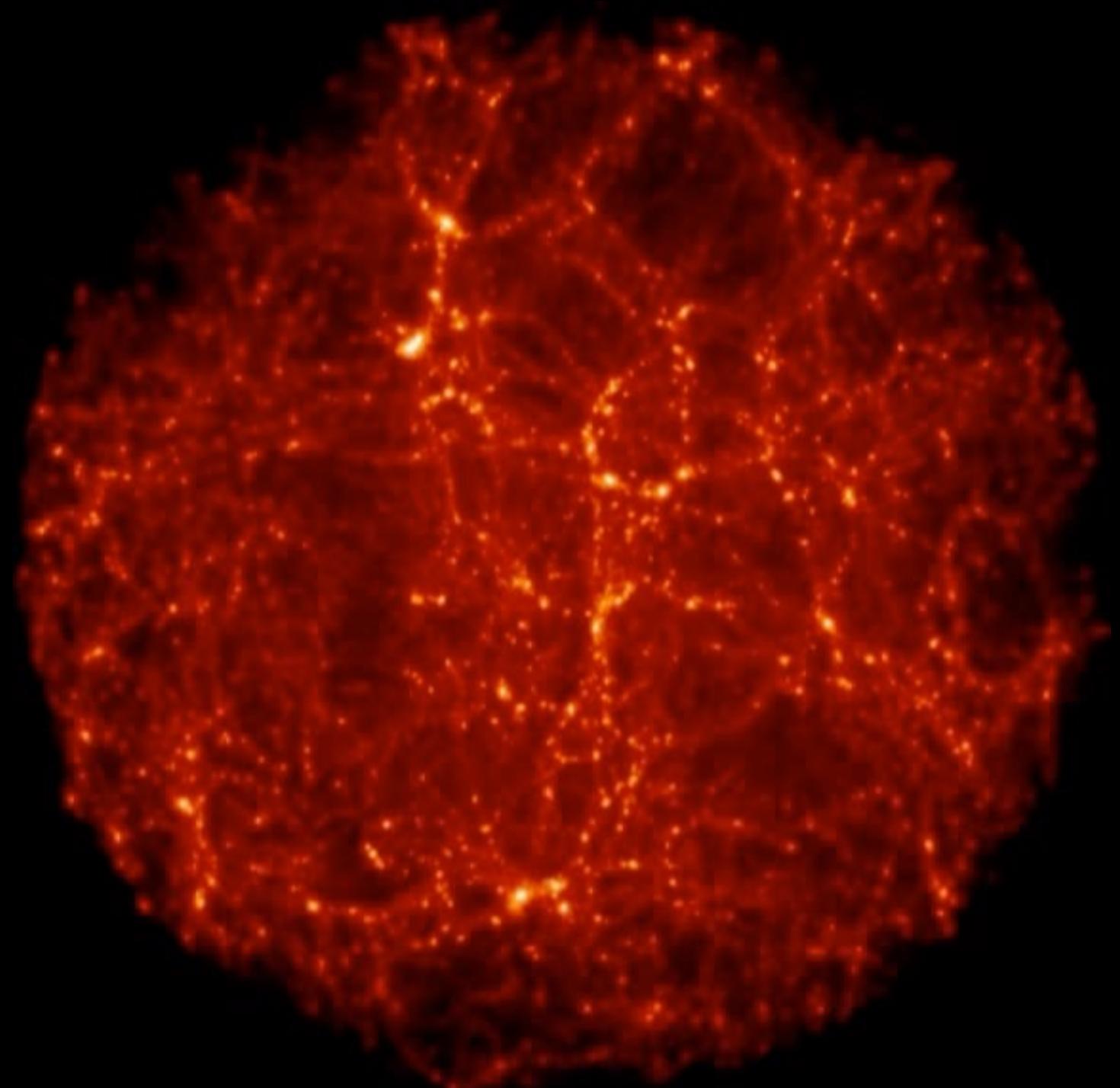
N-Body Problem



**N-Body Simulations are
actually relatively easy!**



62.5 Mpc/h



3. Atoms

3. Atoms (Gas)

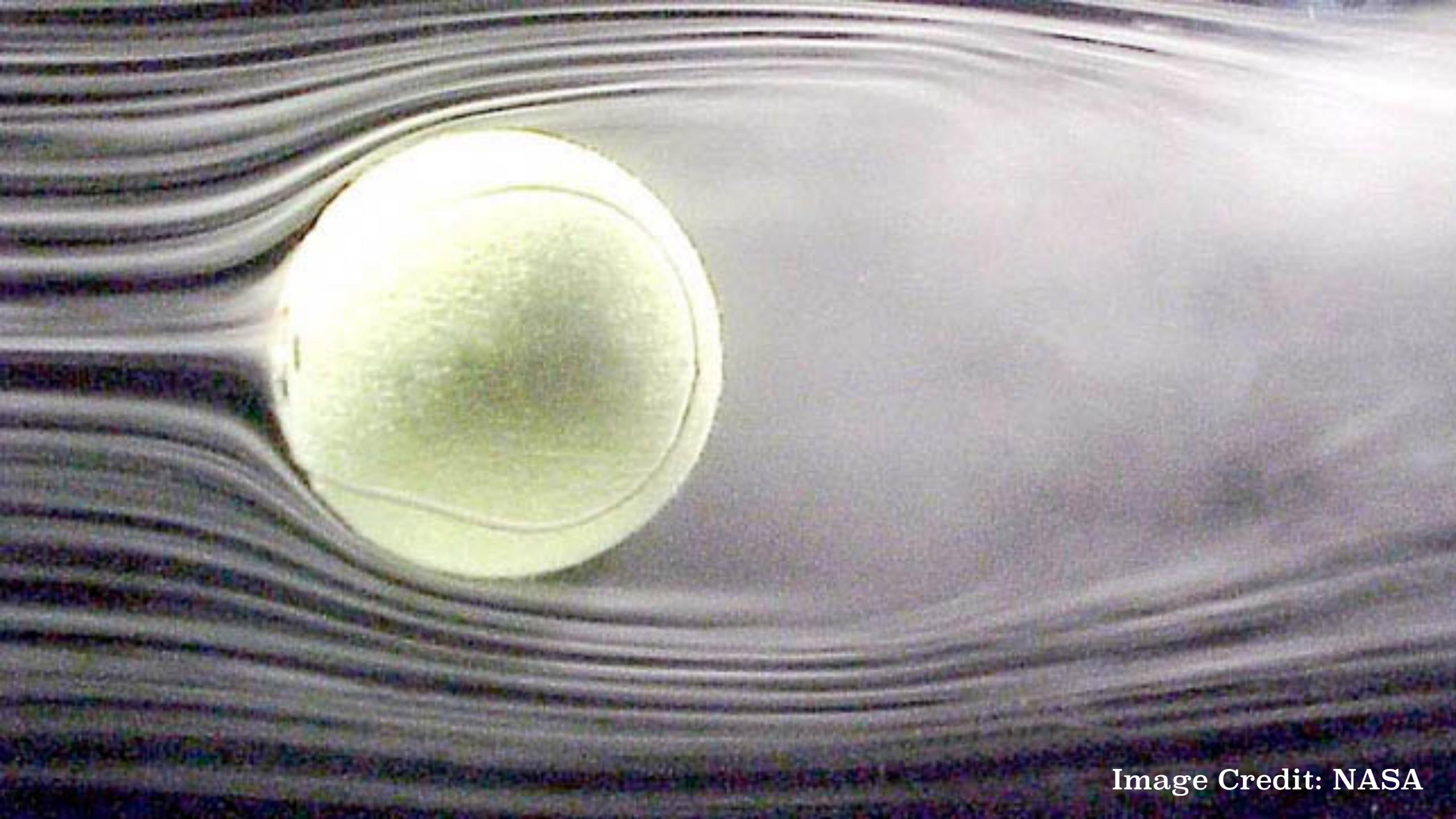
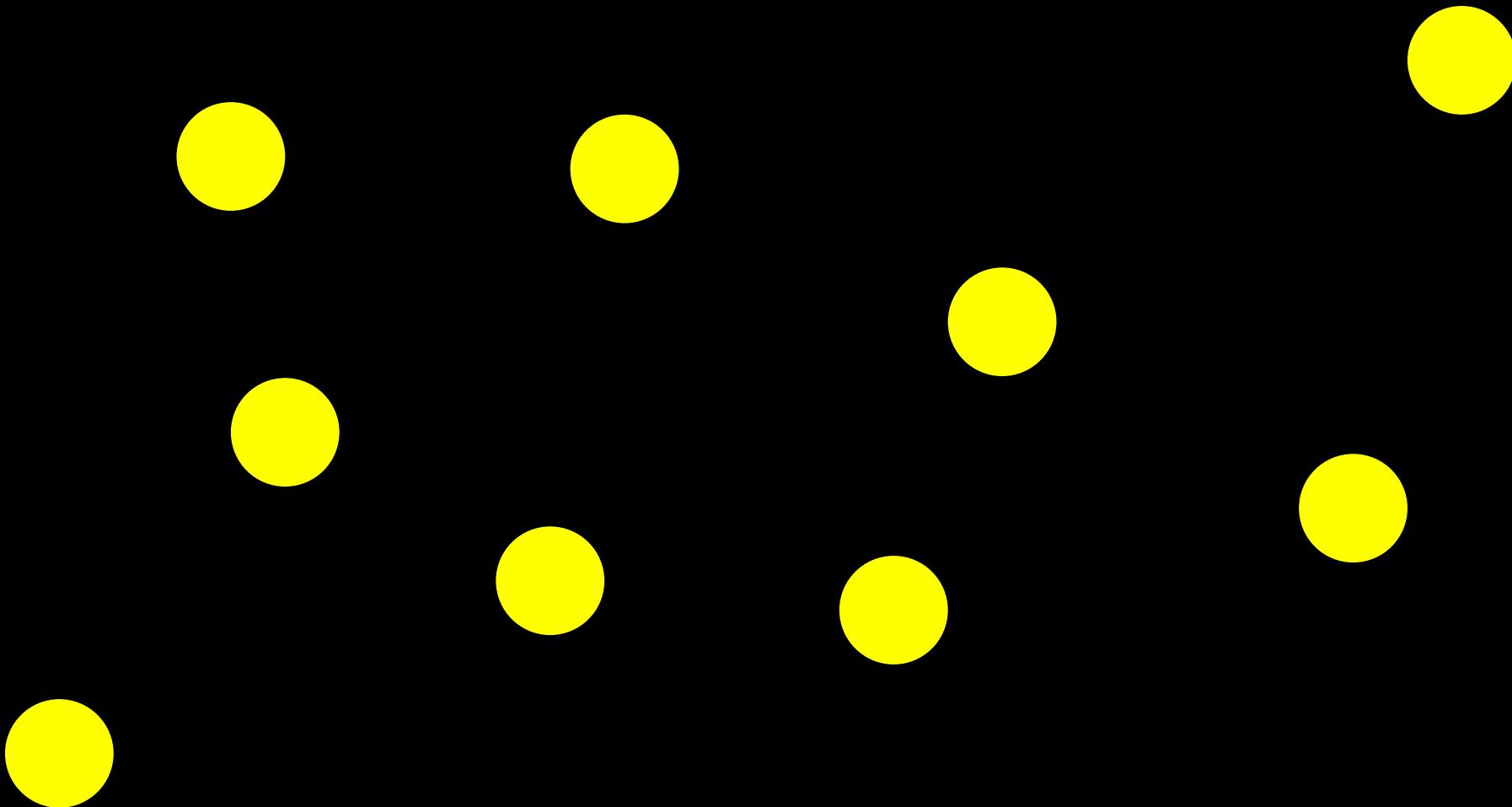
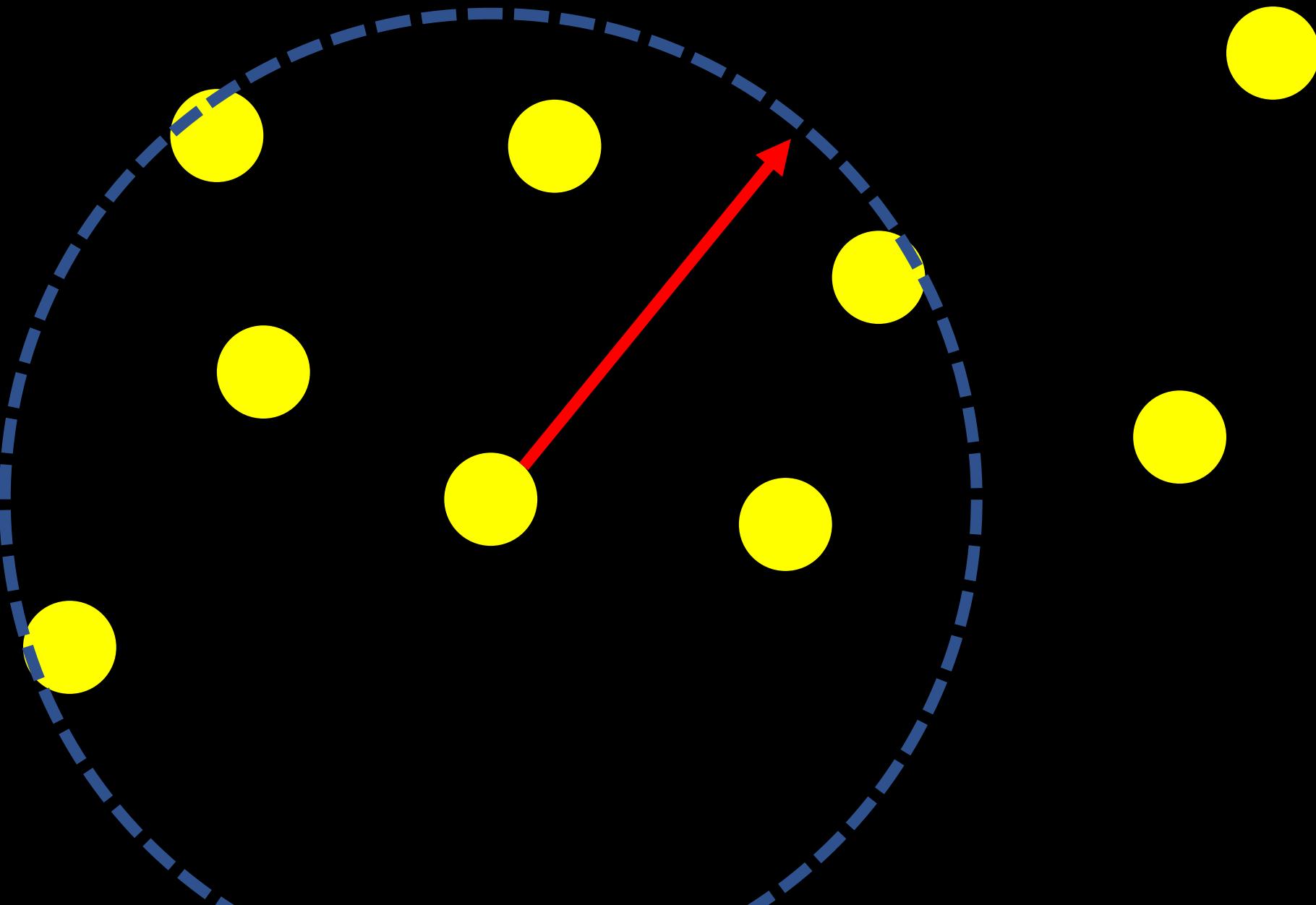
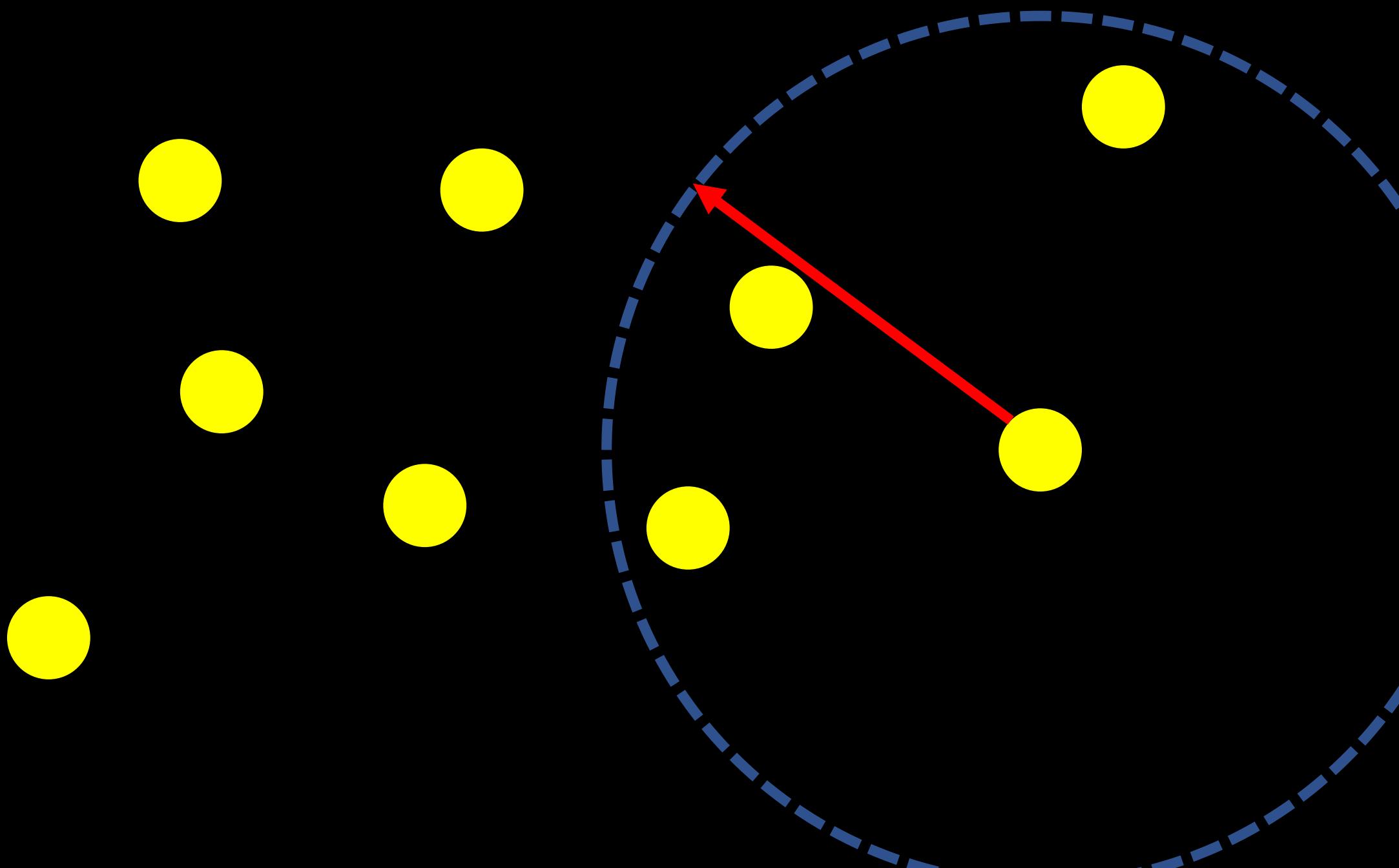
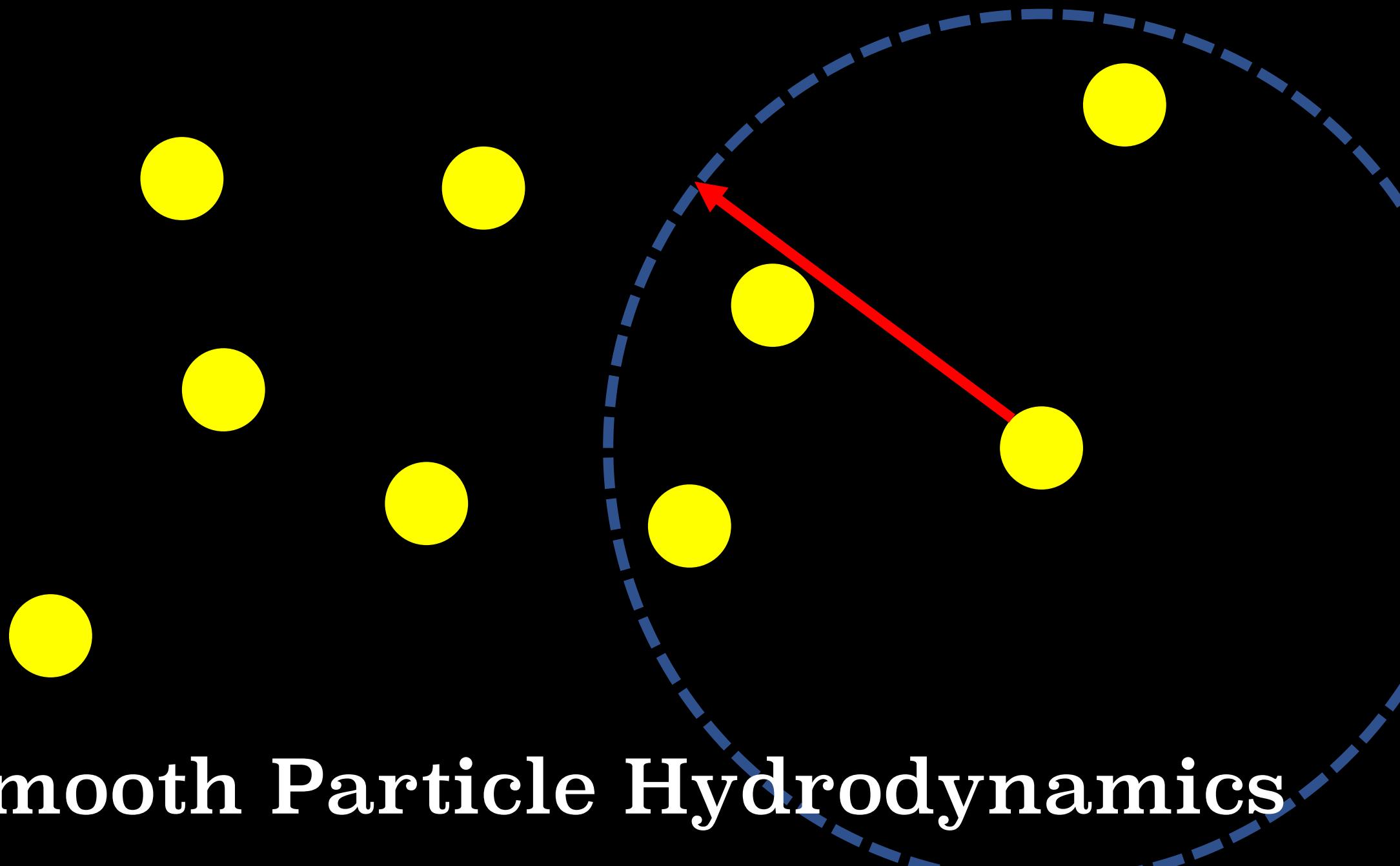


Image Credit: NASA



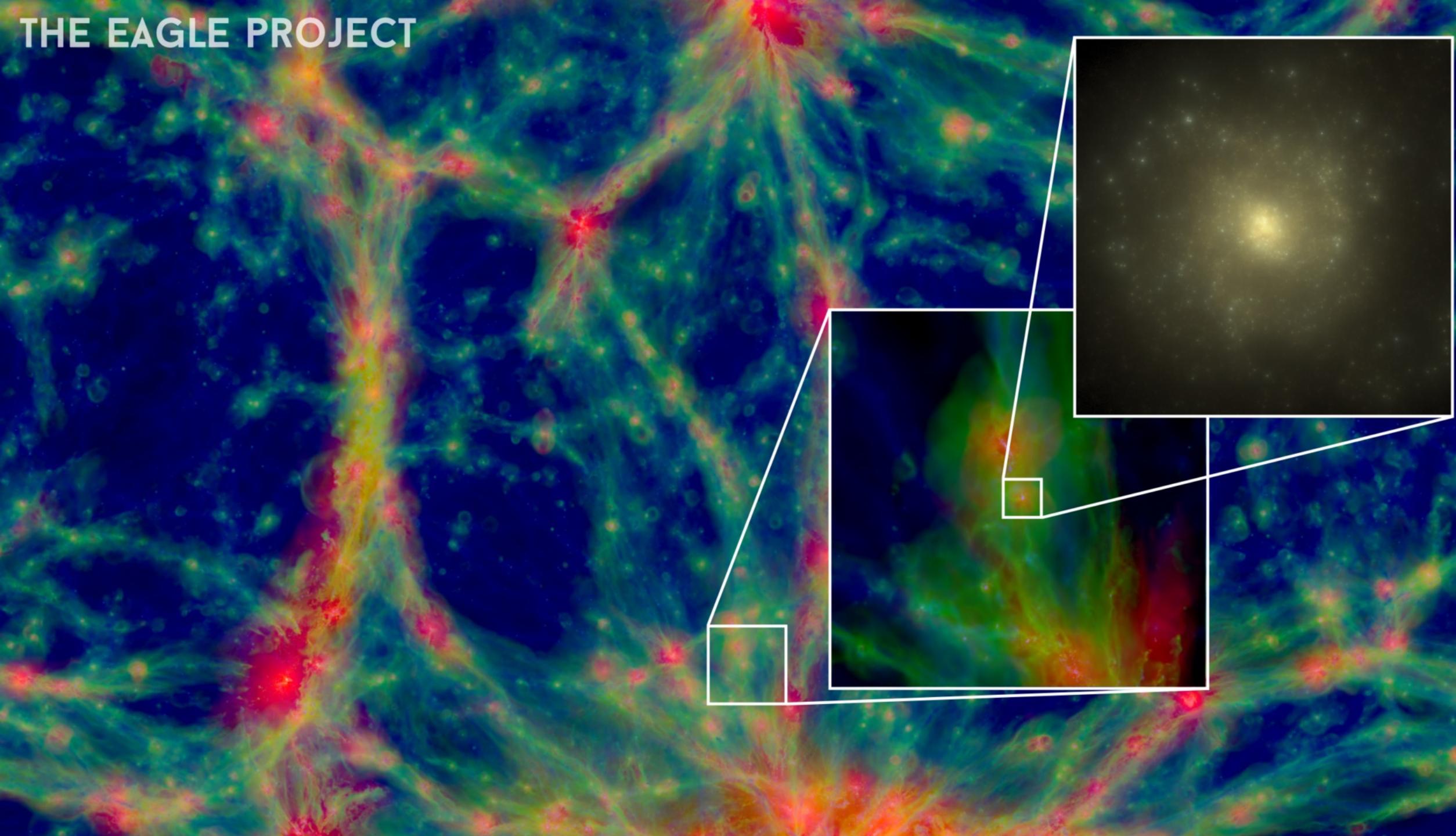




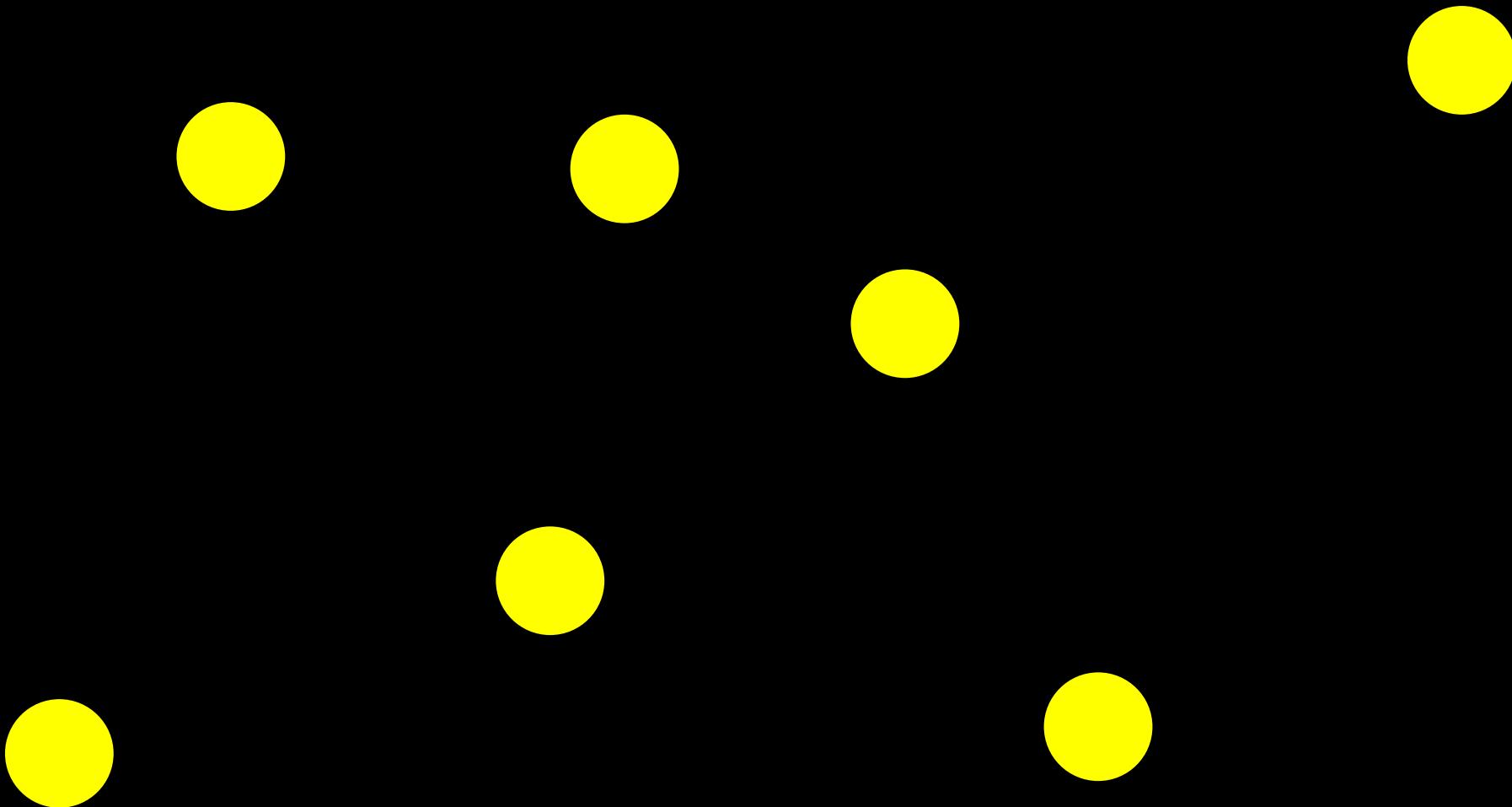


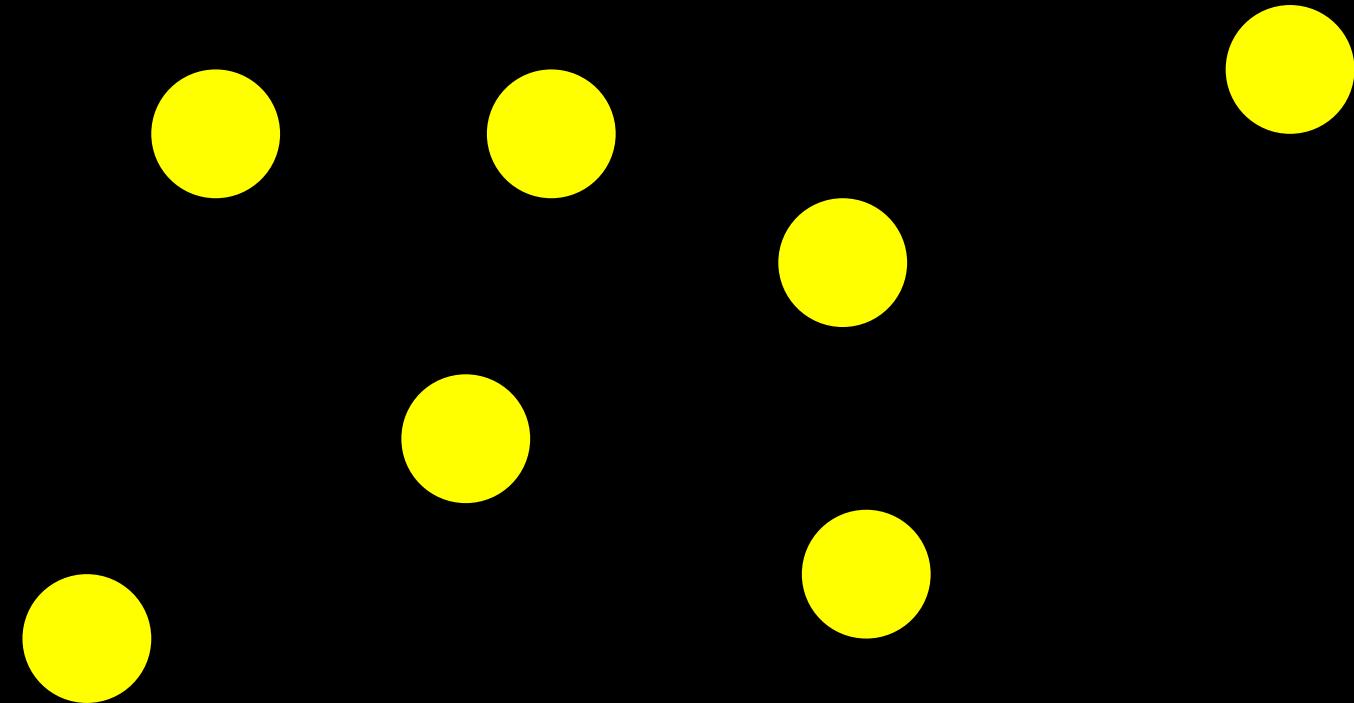
Smooth Particle Hydrodynamics

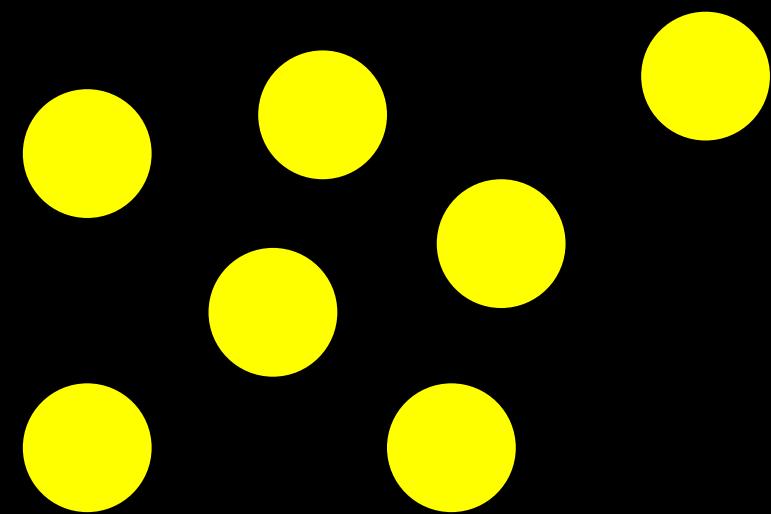
THE EAGLE PROJECT

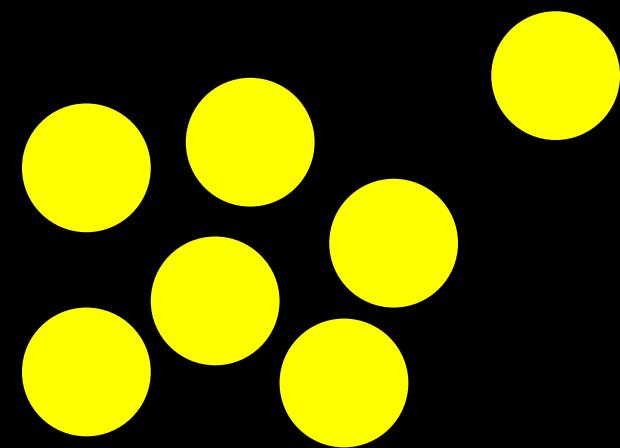


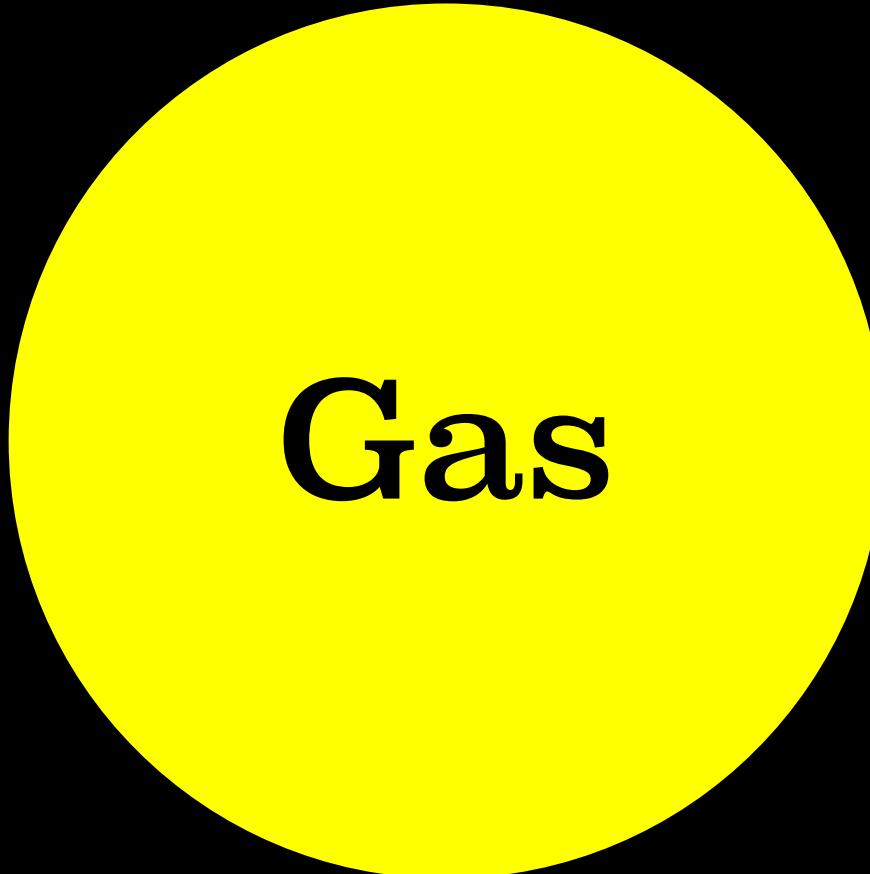
4. Stars



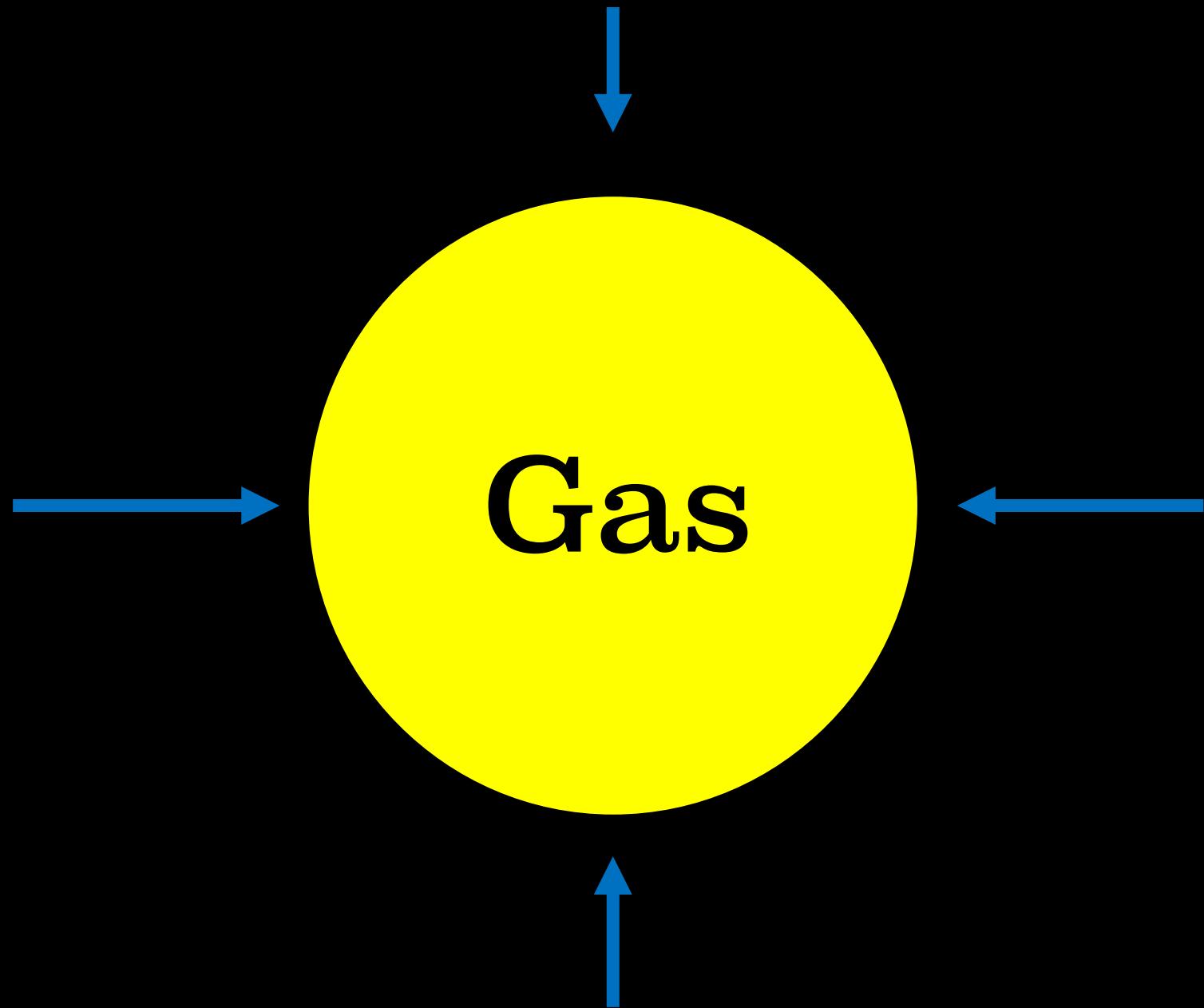


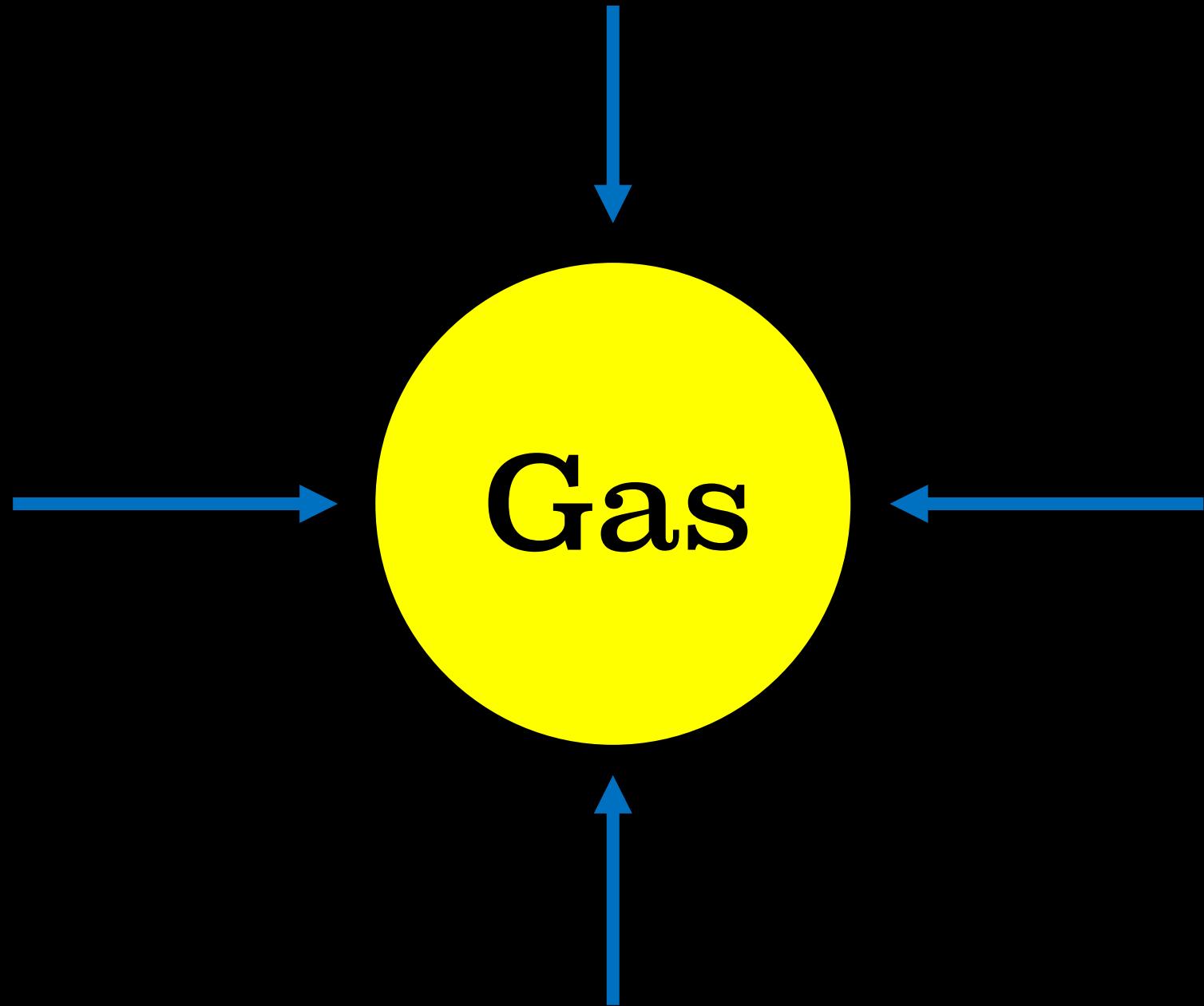




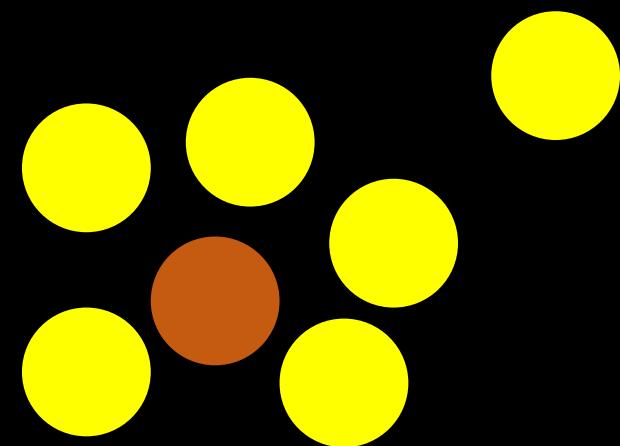


Gas



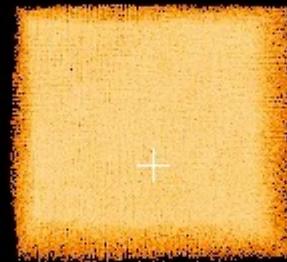


Star



Gas Evolution

Redshift: 18.978



Centre: 52.8056, 52.5858, 50.2844

Stars Evolution

Redshift: 18.978



Centre: 52.8056, 52.5858, 50.2844

Videos Credit: Adam Ussing (Swinburne)

5. Black Holes



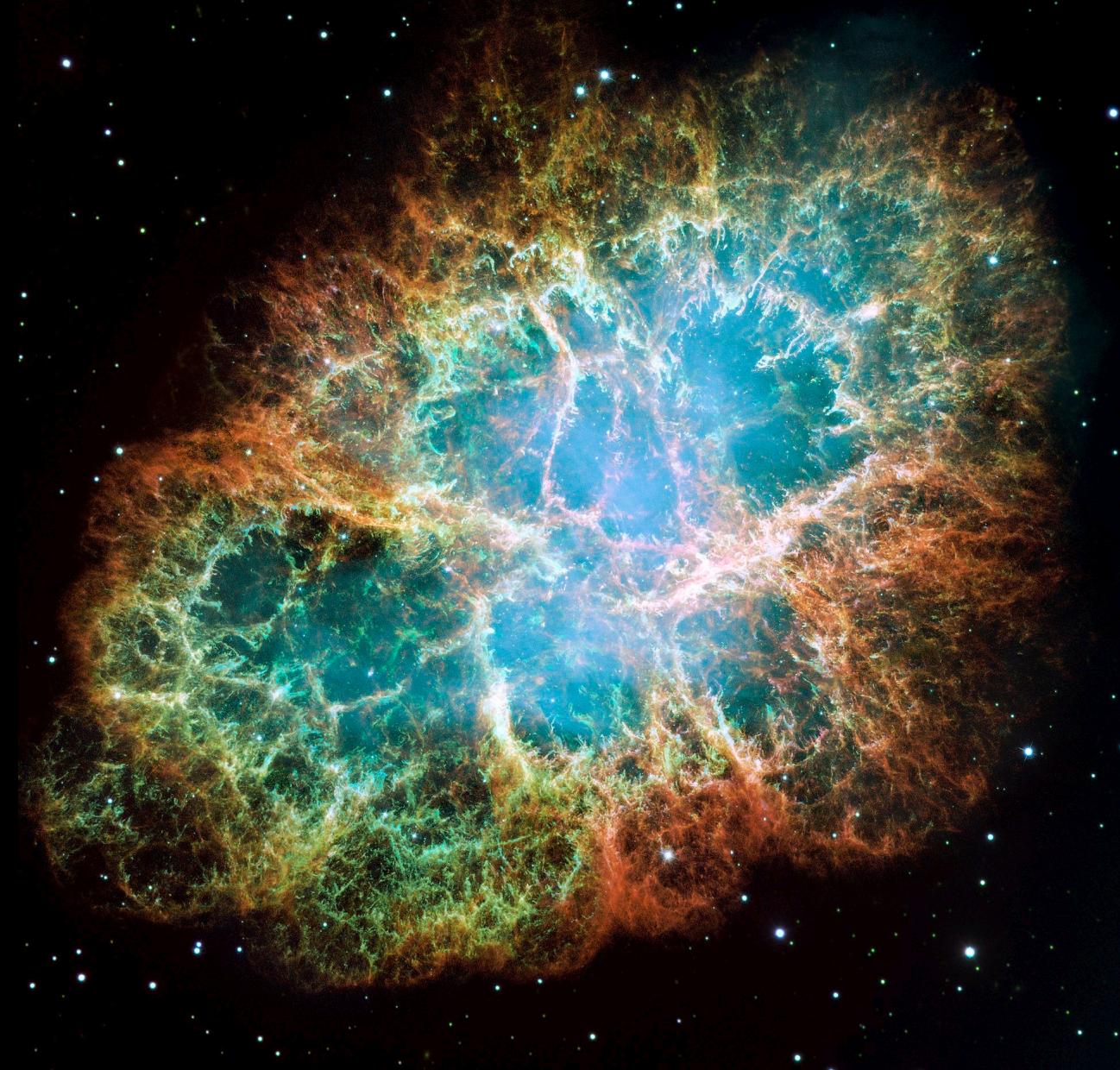
Supermassive black holes are added to the centre of a galaxy when it grows to 10 billion solar masses

6. Feedback

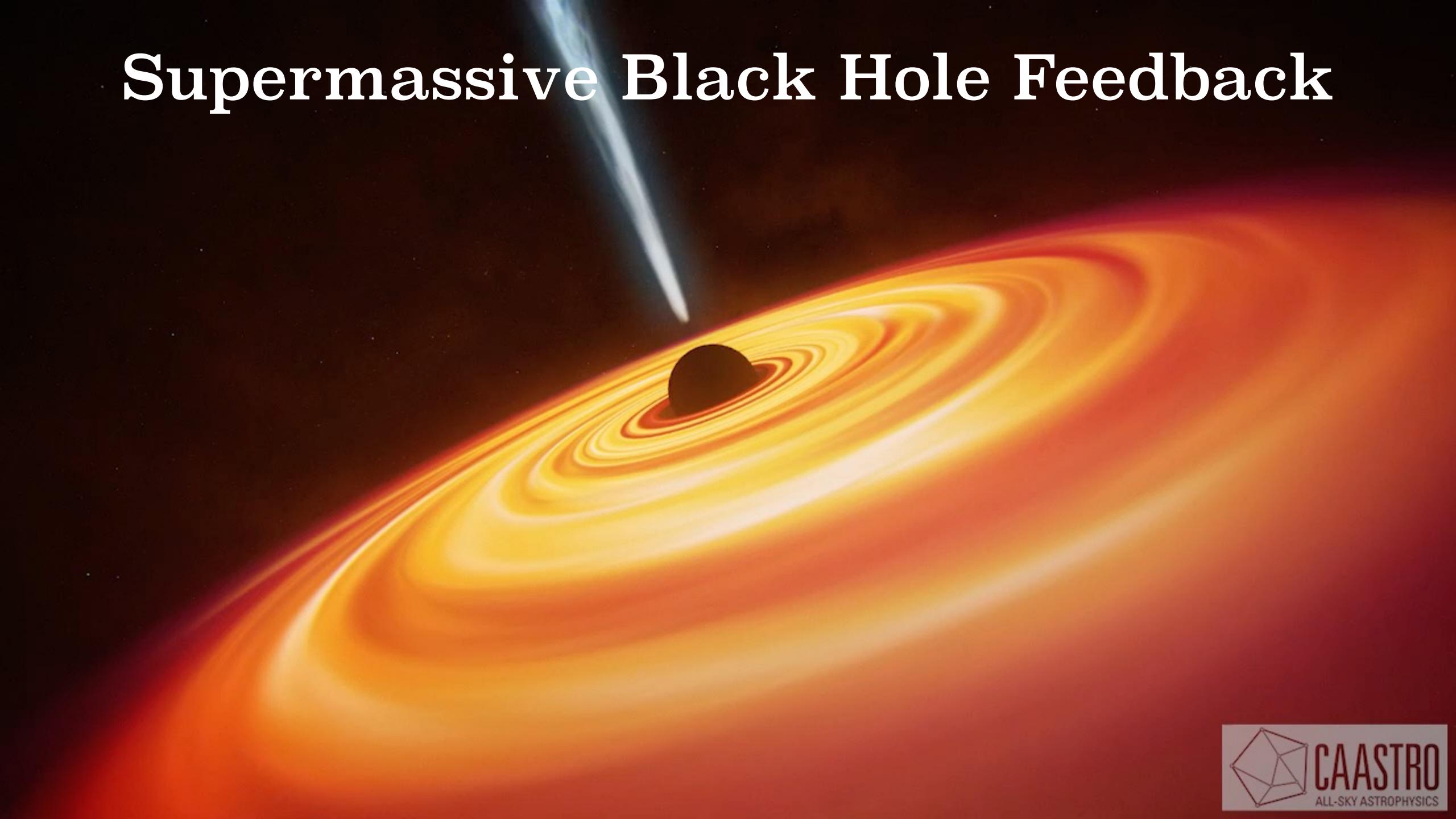
Star Formation Feedback

Supermassive Black Hole Feedback

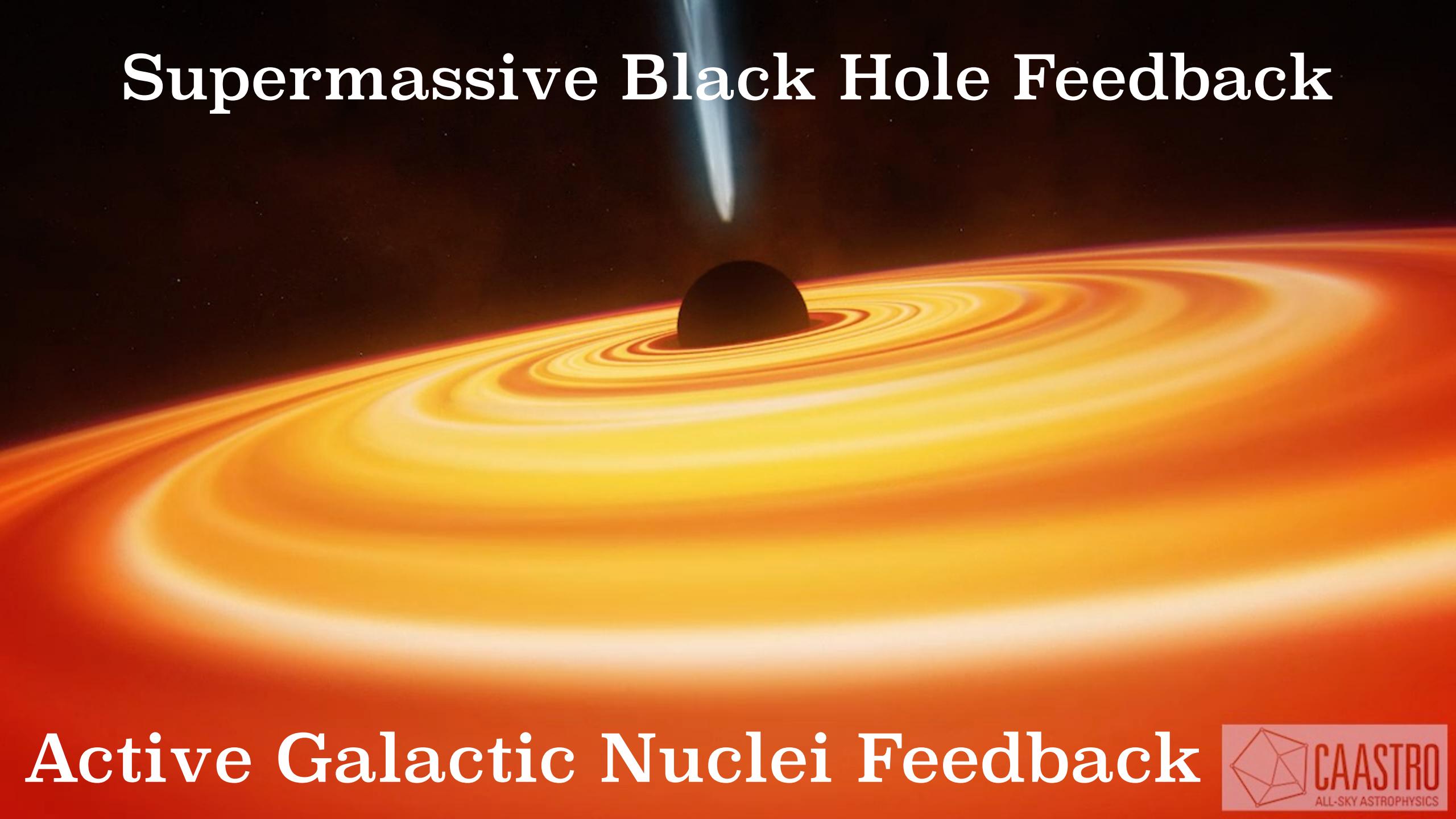
Star Formation Feedback



Supermassive Black Hole Feedback



Supermassive Black Hole Feedback



Active Galactic Nuclei Feedback

Mass of Stars / Total Mass

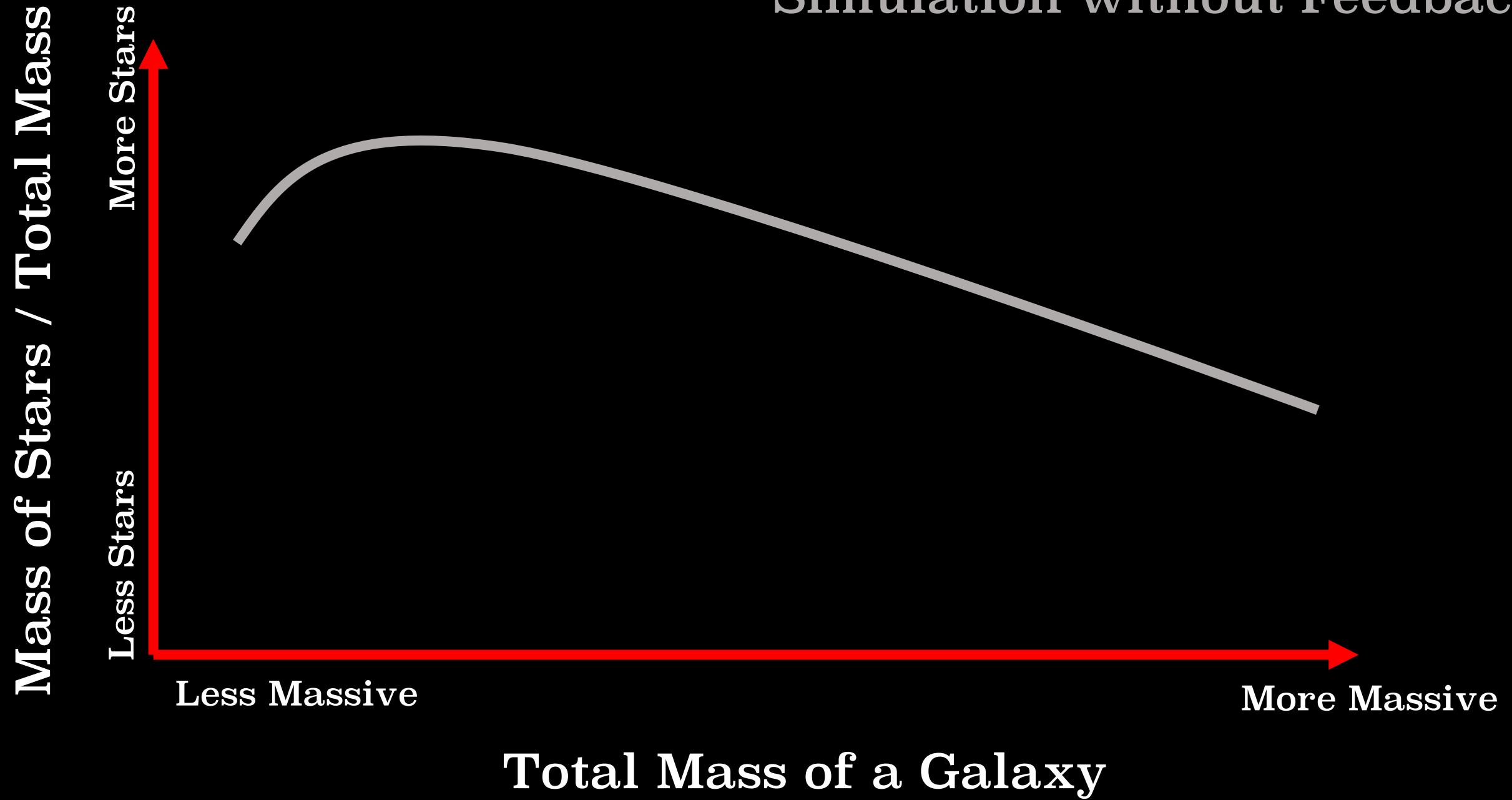
More Stars
Less Stars

Less Massive

More Massive

Total Mass of a Galaxy

Simulation without Feedback



Mass of Stars / Total Mass

More Stars
Less Stars

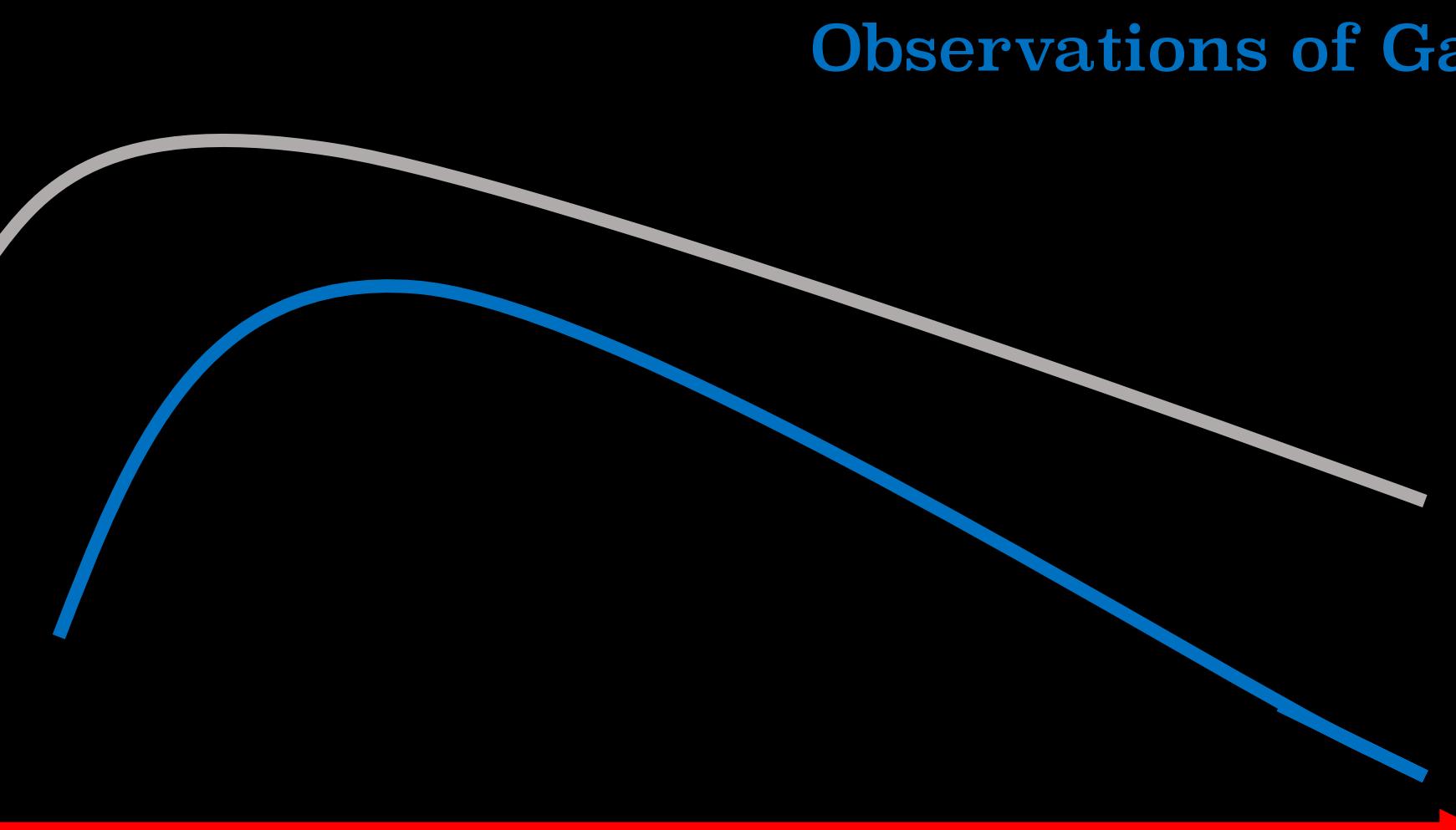
Less Massive

More Massive

Total Mass of a Galaxy

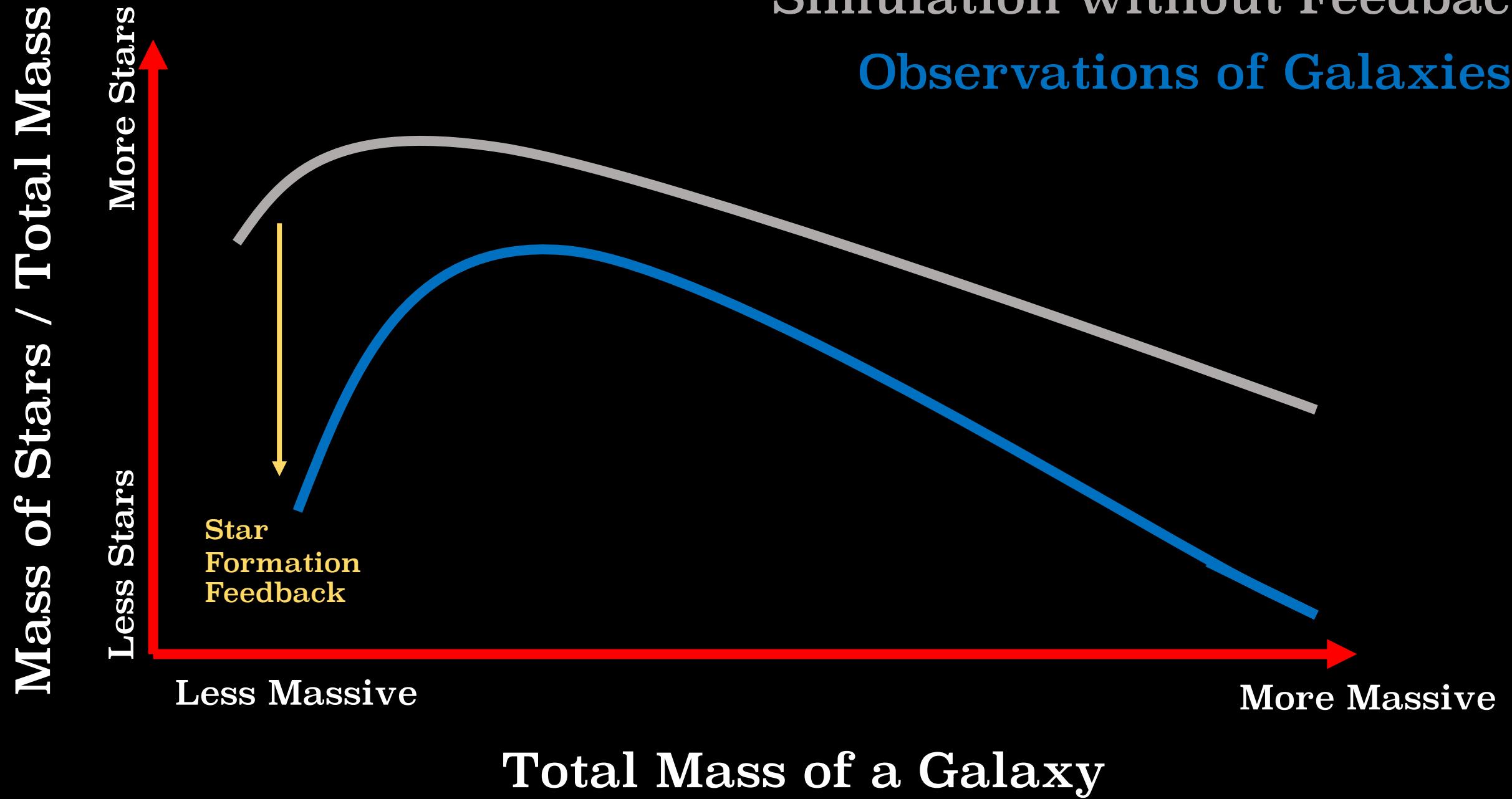
Simulation without Feedback

Observations of Galaxies



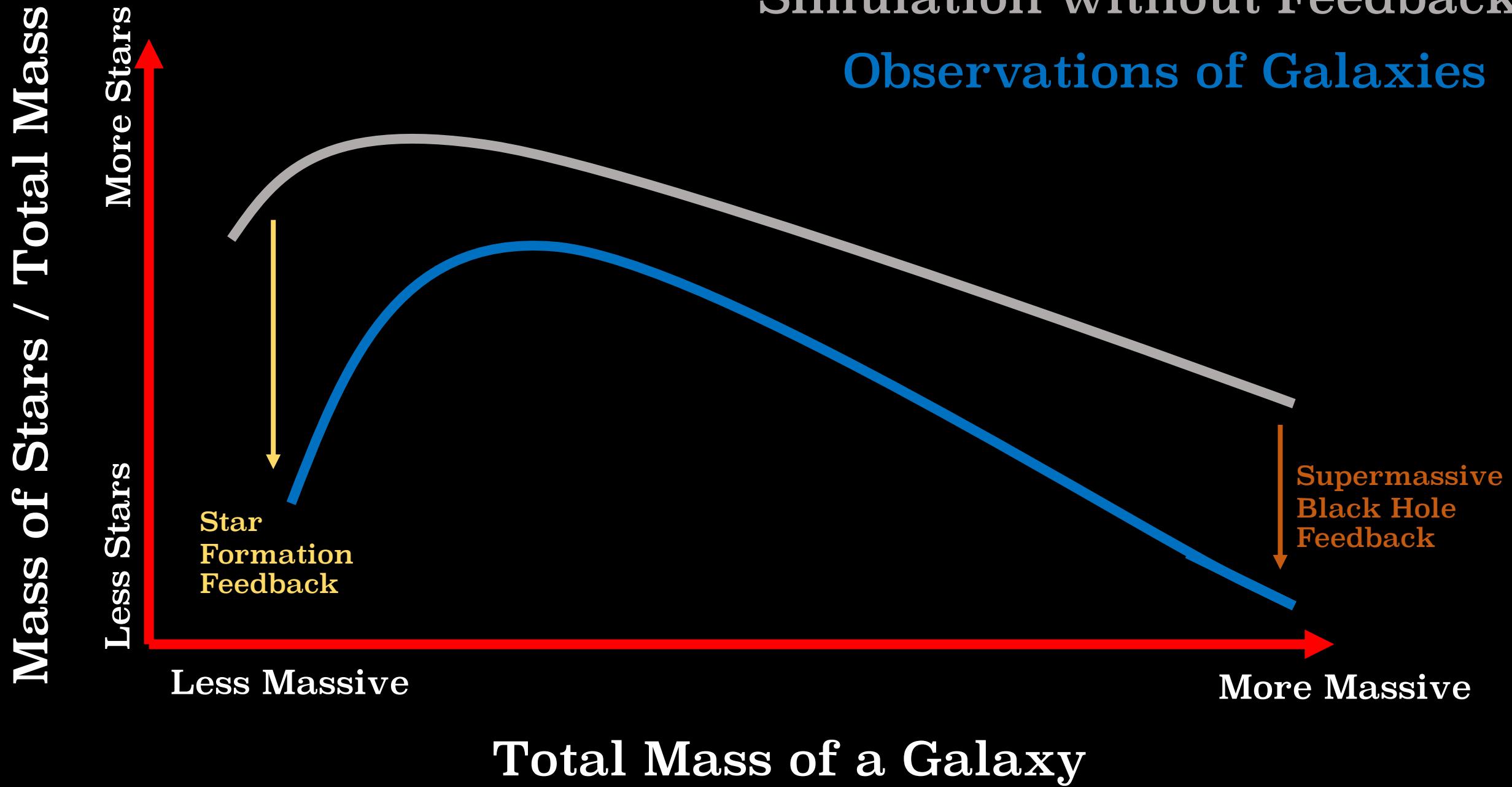
Simulation without Feedback

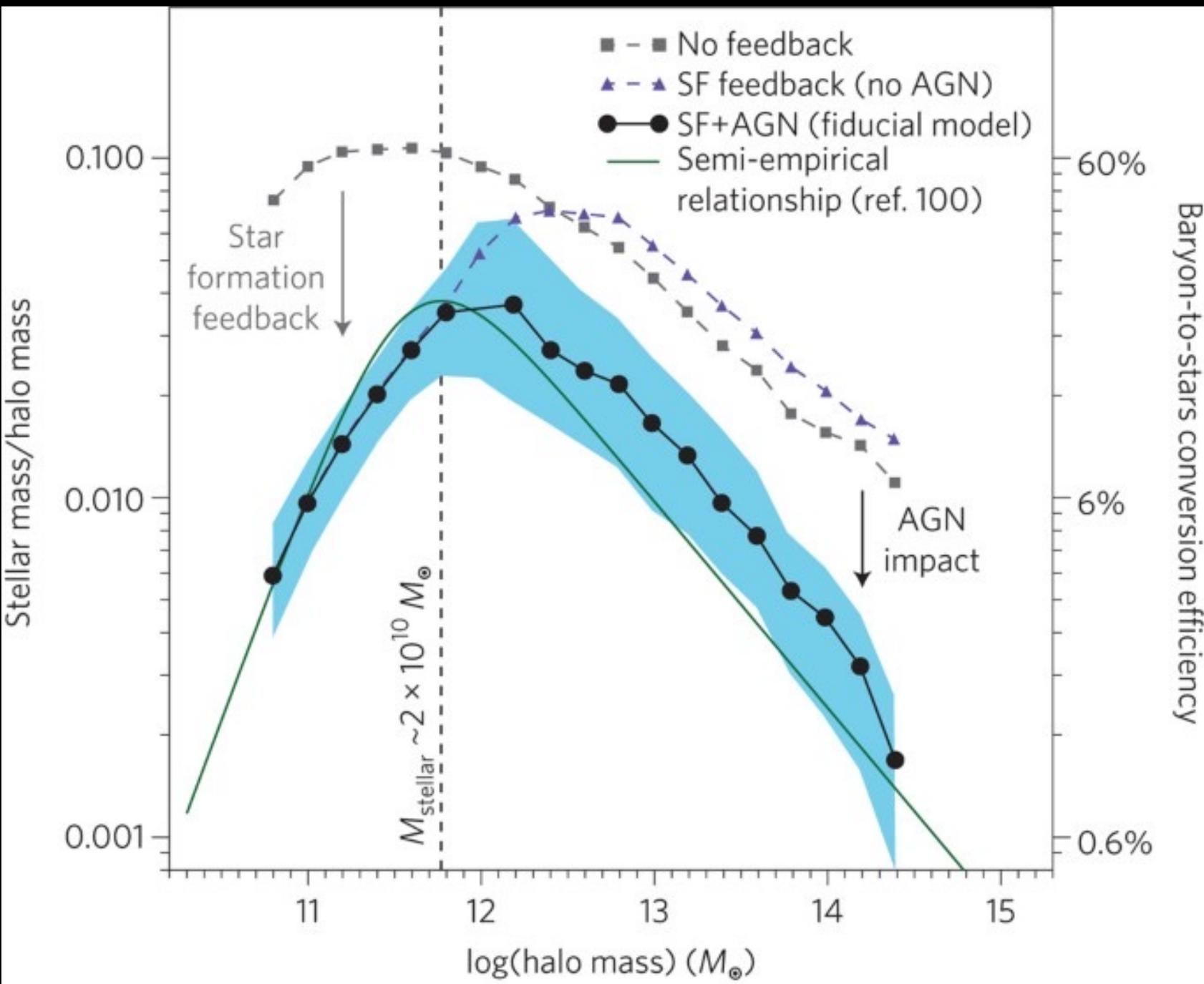
Observations of Galaxies



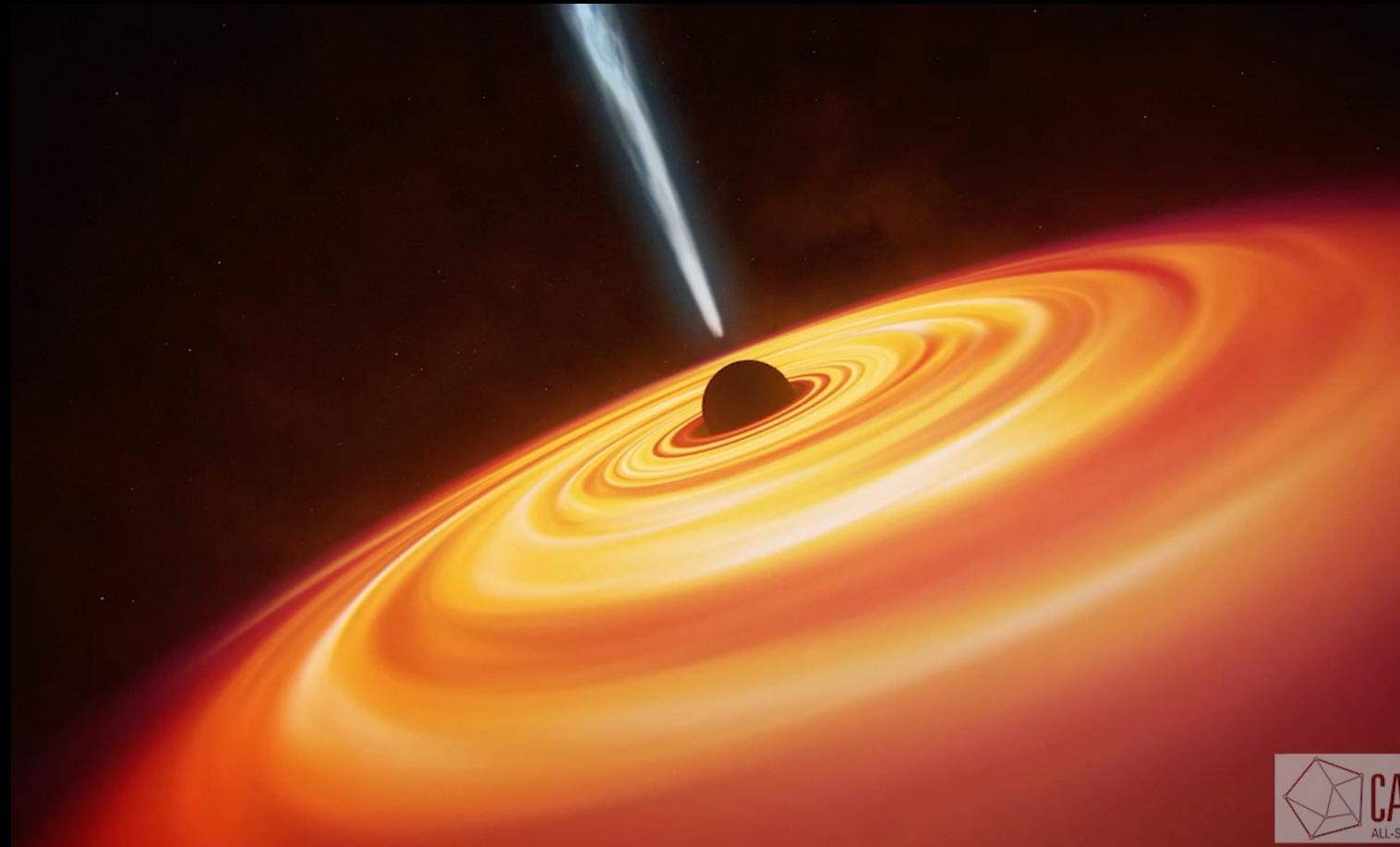
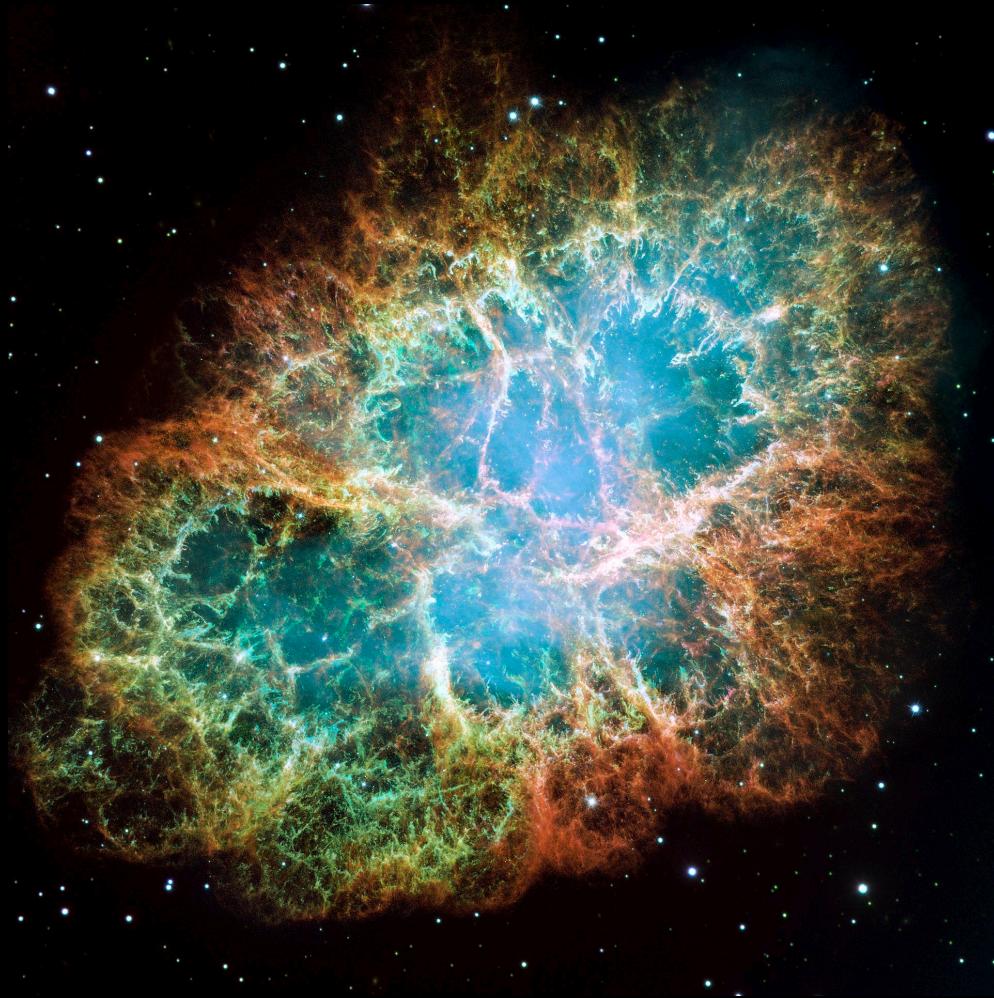
Simulation without Feedback

Observations of Galaxies





6. Feedback



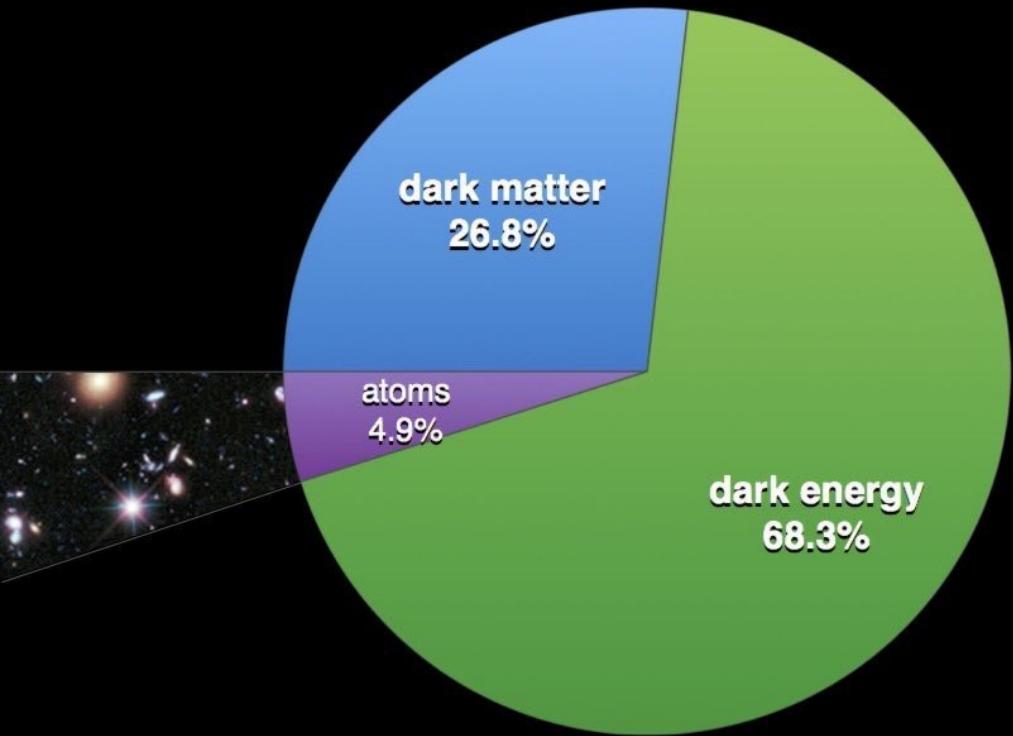
EAGLE: Evolution and Assembly of GaLaxies and their Environments

The evolution of intergalactic gas. Colour encodes temperature

$z = 19.8$
 $t = 0.2 \text{ Gyr}$
 $L = 25.0 \text{ cMpc}$

Simulation by the EAGLE collaboration
Visualisation by Jim Geach & Rob Crain

Dark Energy?



 **planck** CMB Simulator

Normal Matter ($\Omega_b = 0.05$)

Dark Matter ($\Omega_c = 0.275$)

Dark Energy ($\Omega_\Lambda = 0.675$)

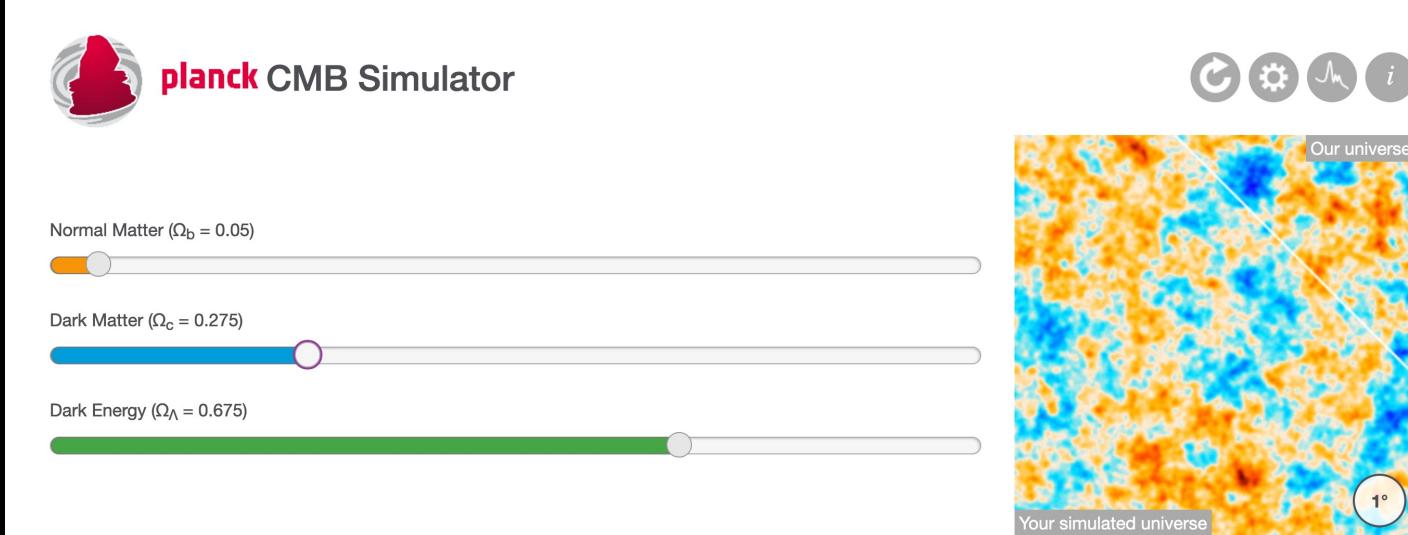
Normal matter only

Our universe

Your simulated universe

1°

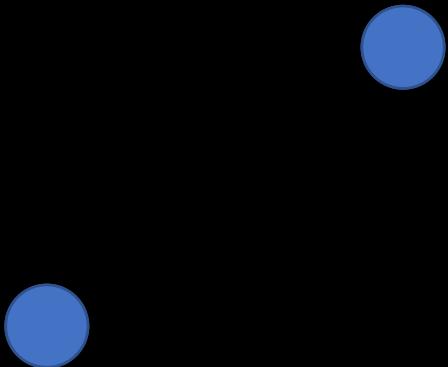
13.8 billion years old - just right
flat universe
Fundamental scale ~0.8°
Universe similarity **100%** - the same as our universe



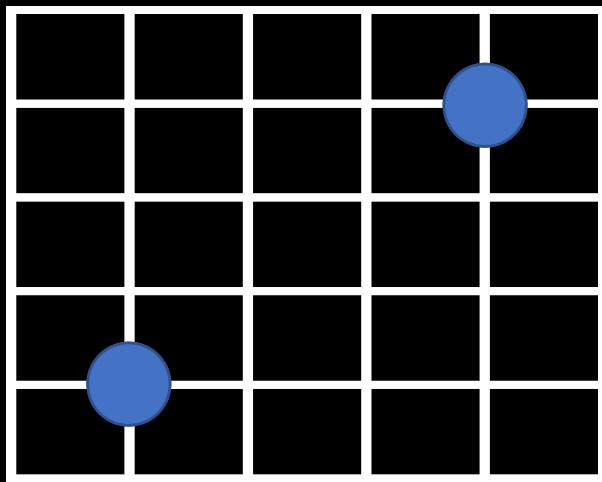
@AstroKatie/Planck13



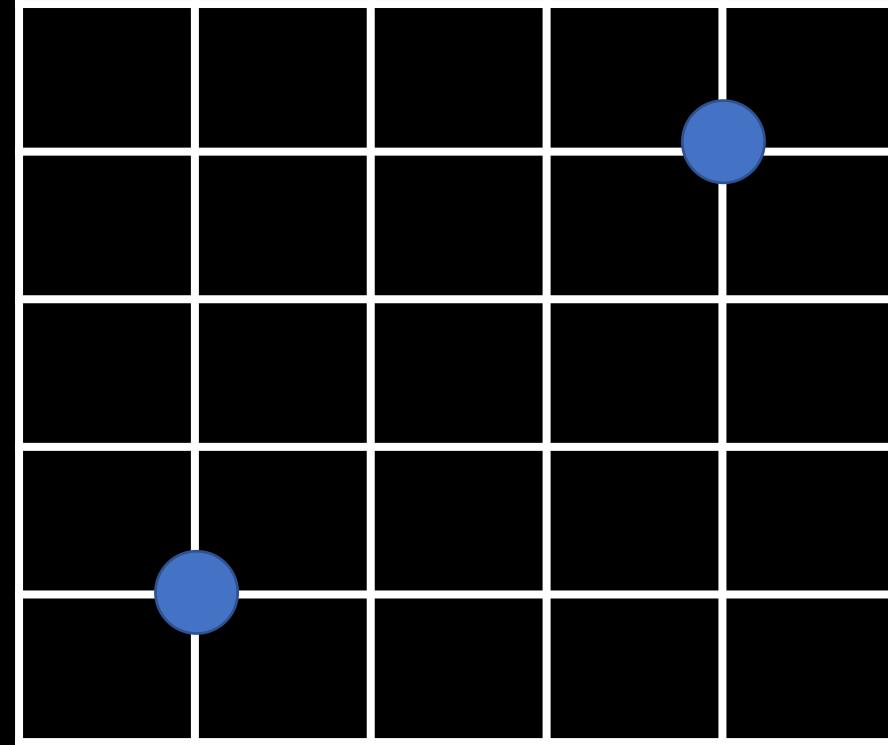
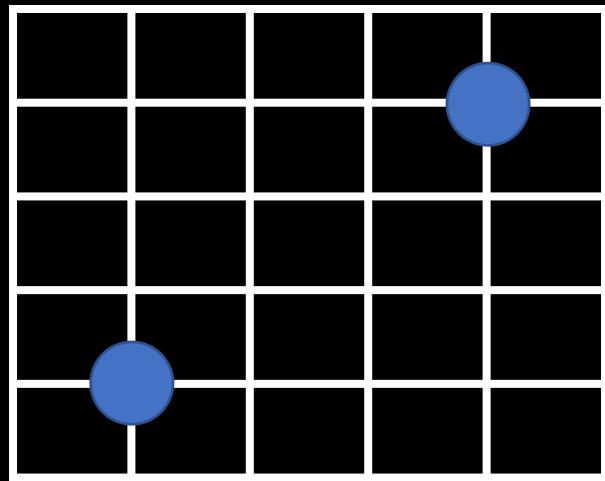
Comoving Coordinate System



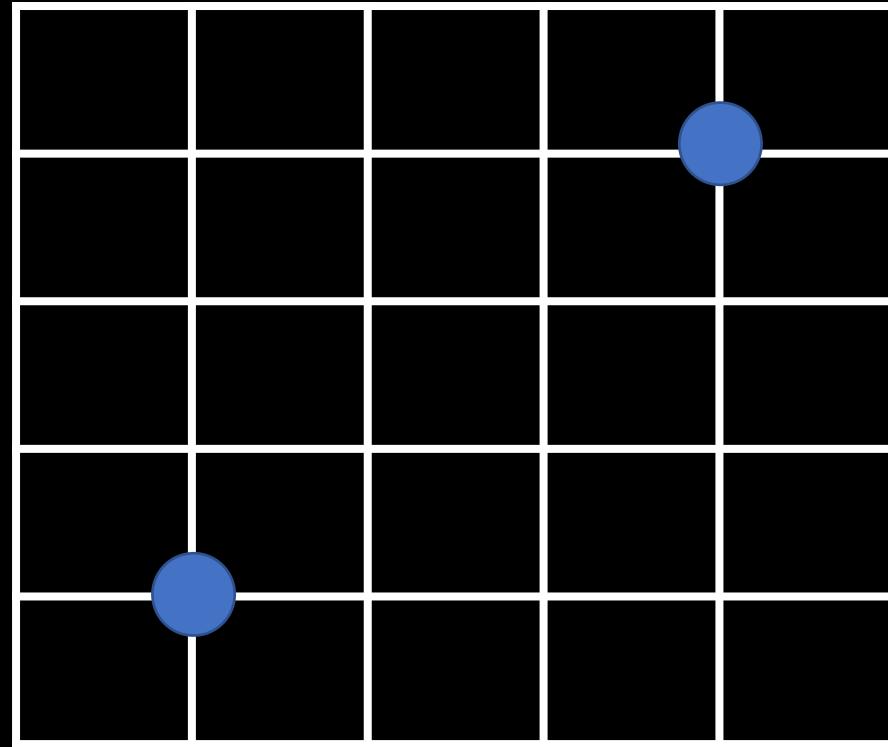
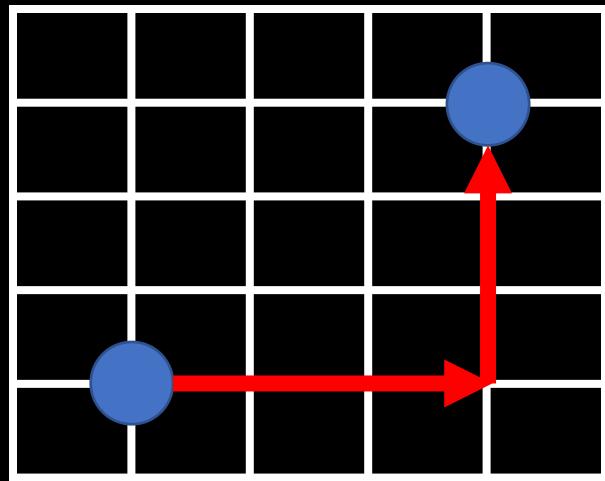
Comoving Coordinate System



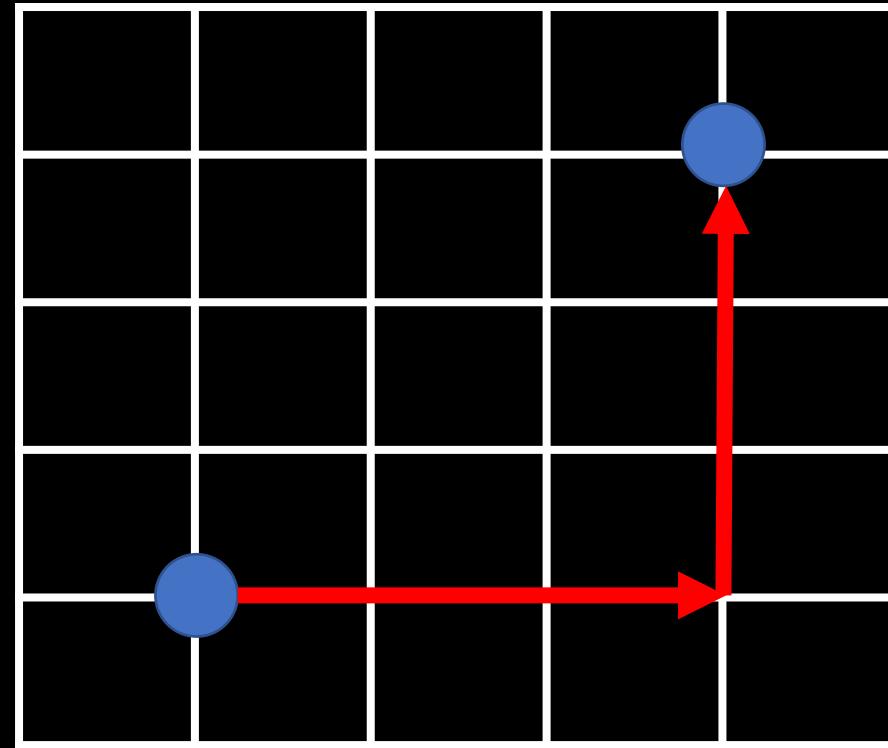
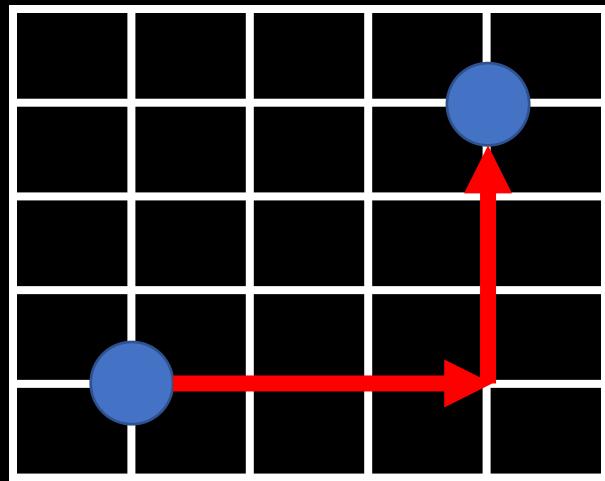
Comoving Coordinate System



Comoving Coordinate System



Comoving Coordinate System



What have we
learnt from
simulations?

Black holes can drive galaxy evolution

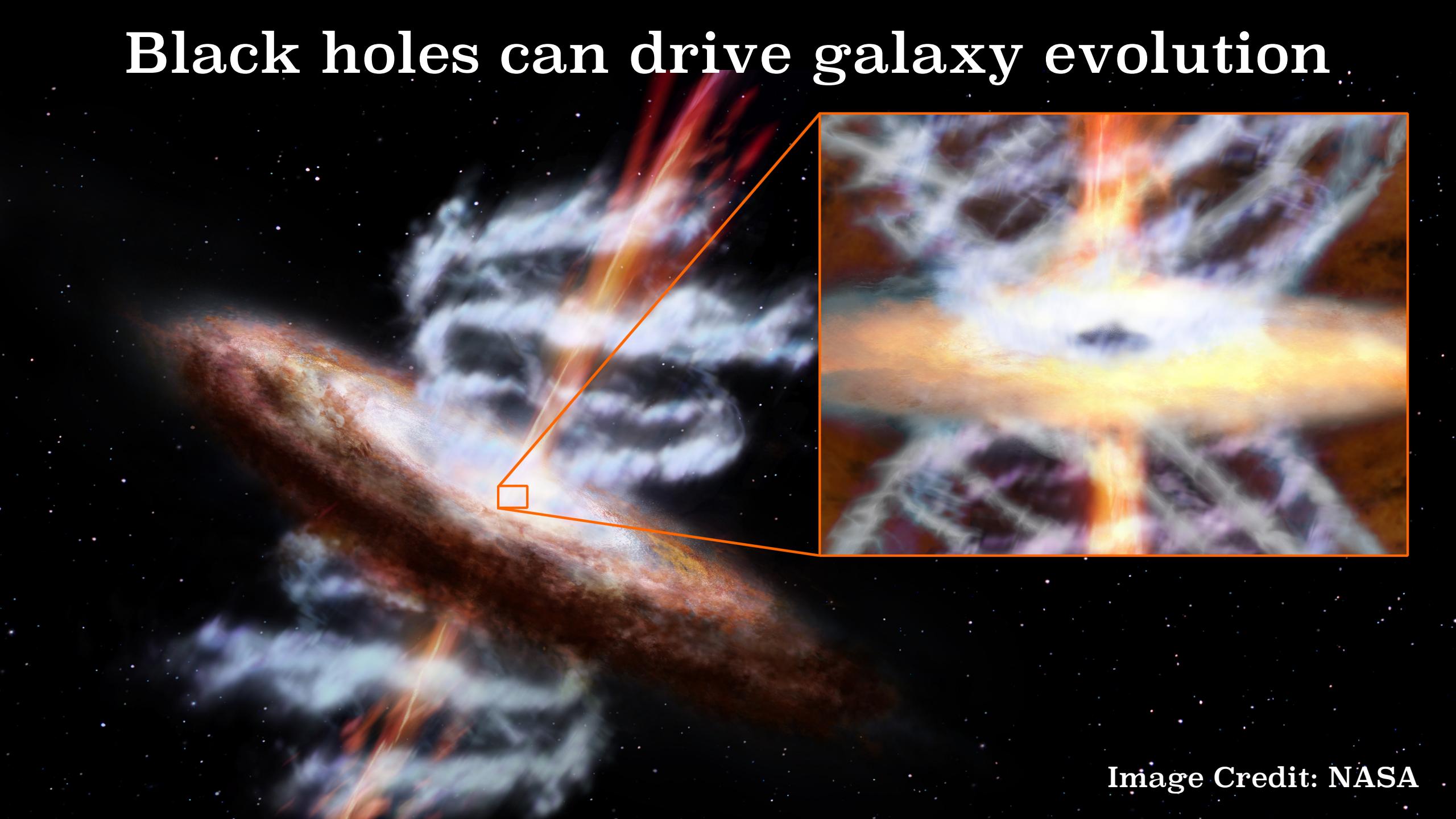


Image Credit: NASA

CDM

$z = 15.00$

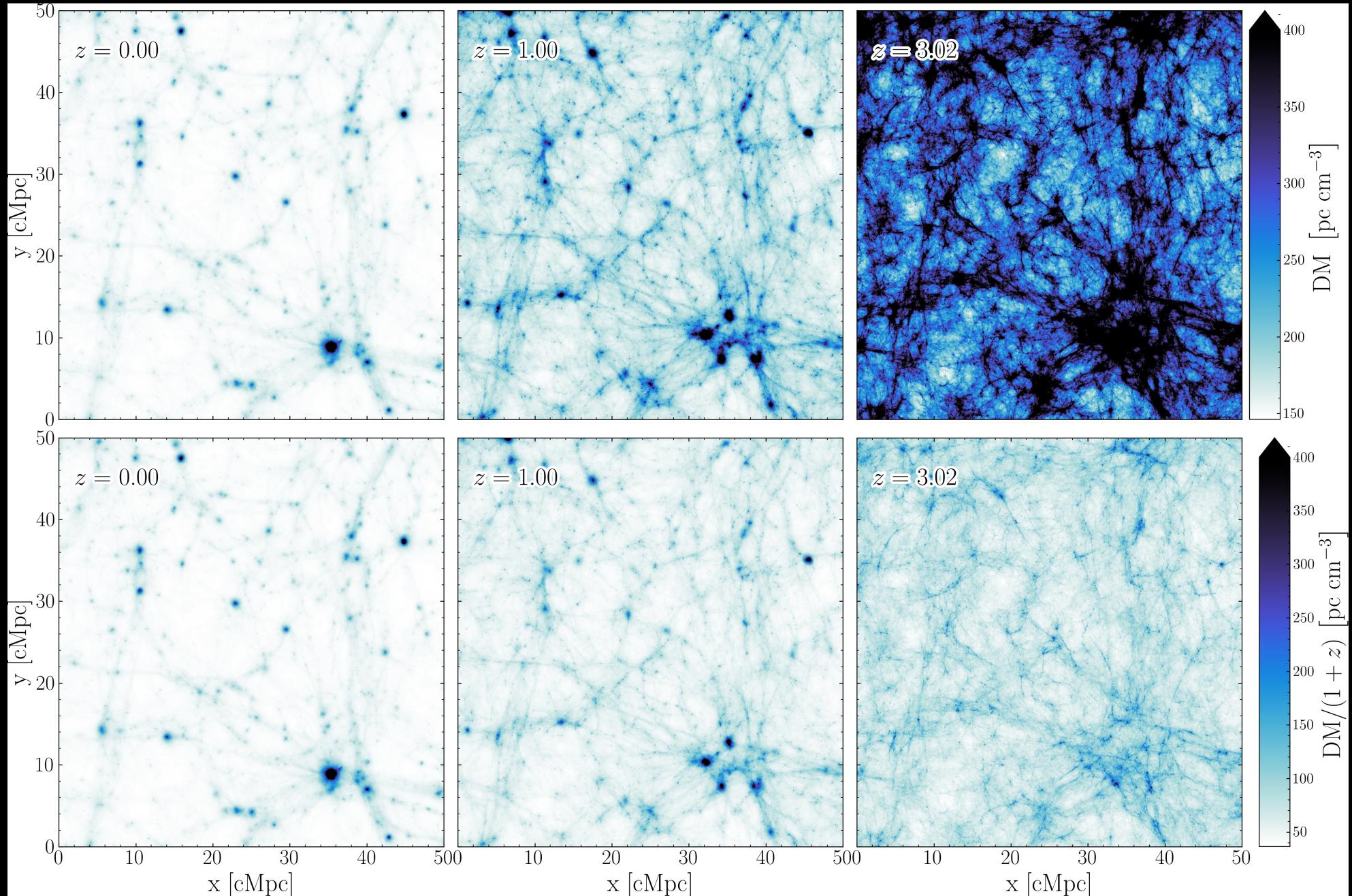
WDM

a

b

What are the properties of dark matter?

Image Credit: Gao & Theuns (2007)



Helping to reveal the location of the missing baryons.

Summary



Twitter: @adamjbatten

We use simulations to test our understanding of the laws of the Universe, and make predictions based on what we already know.

0. Gravity
1. Cosmic Microwave Background
2. Dark Matter
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4. Stars
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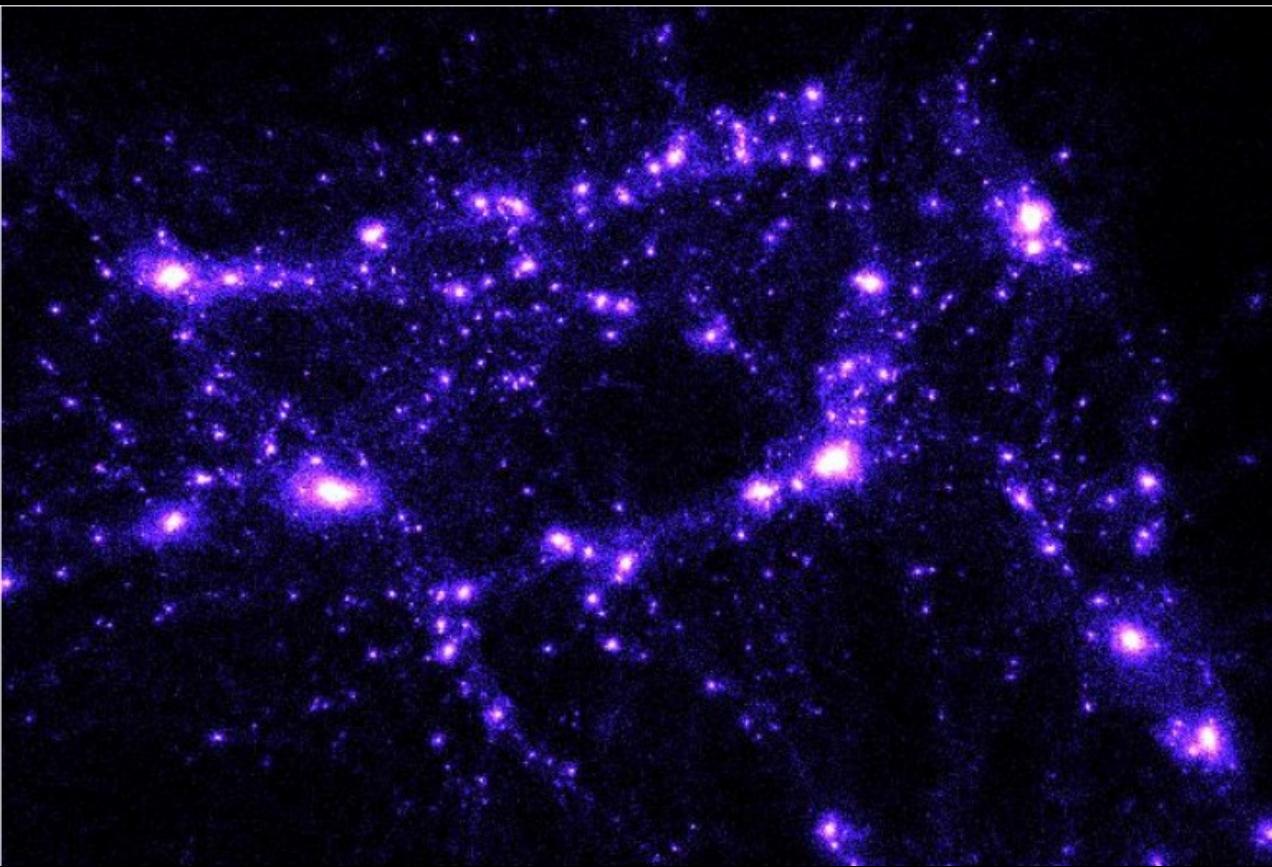
Slides:
<https://adambatten.com/talks/>

$$\frac{1}{2}g^{\alpha\beta}\partial_{\alpha}\partial_{\mu}g_{\beta\nu}+\frac{1}{2}g^{\alpha\beta}\partial_{\alpha}\partial_{\nu}g_{\mu\beta}-\frac{1}{2}g^{\alpha\beta}\partial_{\alpha}\partial_{\beta}g_{\mu\nu}-\frac{3}{2}g^{\alpha\beta}\partial_{\mu}\partial_{\nu}g_{\alpha\beta}$$

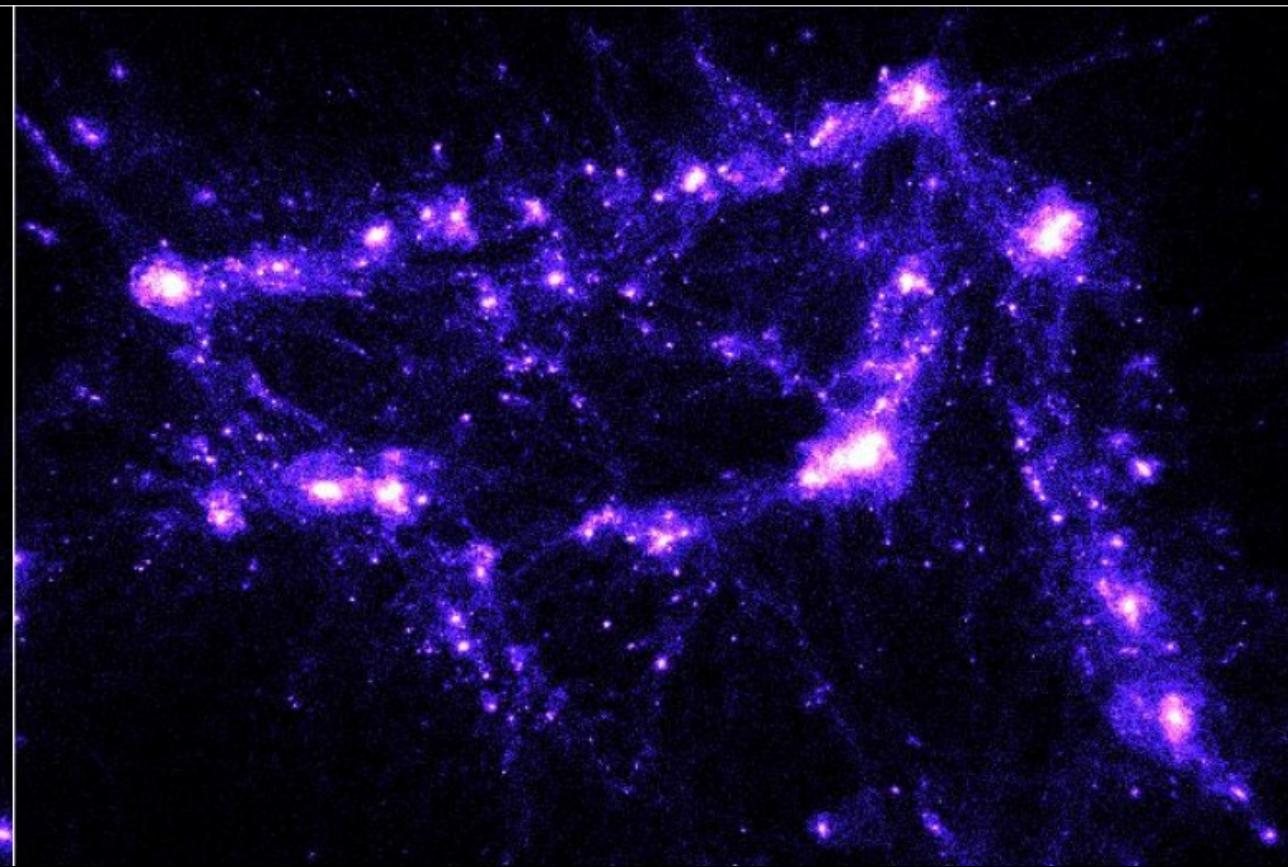
$$-\frac{1}{2}g^{\beta\lambda}g^{\alpha\rho}\partial_{\alpha}g_{\rho\lambda}\partial_{\mu}g_{\beta\nu}-\frac{1}{2}g^{\beta\lambda}g^{\alpha\rho}\partial_{\alpha}g_{\rho\lambda}\partial_{\nu}g_{\mu\beta}$$

$$+\frac{1}{4}g^{\beta\lambda}g^{\alpha\rho}\partial_{\nu}g_{\alpha\lambda}\partial_{\mu}g_{\rho\beta}+\frac{1}{4|g|}g^{\alpha\beta}\partial_{\beta}|g|\partial_{\nu}g_{\mu\alpha}$$

$$-\frac{1}{4|g|}g^{\alpha\beta}\partial_{\beta}|g|\partial_{\alpha}g_{\mu\nu}-\frac{1}{4|g|}g^{\alpha\beta}\partial_{\beta}|g|\partial_{\mu}g_{\alpha\nu}+\Lambda g_{\mu\nu}=\frac{8\pi G}{c^4}T_{\mu\nu}$$



Dark Matter + Dark Energy



Modified Newtonian Dynamics

