

H α and Ca II Infrared Triplet Variations During a Transit of the 23 Myr Planet V1298 Tau c

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ABSTRACT

Young transiting exoplanets (< 100 Myr) provide crucial insight into atmospheric evolution via photoevaporation. However, transmission spectroscopy measurements to determine atmospheric composition and mass loss are challenging due to the activity and prominent stellar disk inhomogeneities present on young stars. We observed a full transit of V1298 Tau c, a 23 Myr, 5.59 R $_{\oplus}$ planet orbiting a young K0-K1.5 solar analogue with GRACES on Gemini-North. We were able to measure the Doppler tomographic signal of V1298 Tau c using the Ca II infrared triplet (IRT) and find a projected obliquity of $\lambda = 5^\circ \pm 15^\circ$. The tomographic signal is only seen in the chromospherically driven core of the Ca II IRT, which may be the result of star-planet interactions. Additionally, we find that excess absorption of the H α line decreases smoothly during the transit. While this could be a tentative detection of hot gas escaping the planet, we find this variation is consistent with similar timescale observations of other young stars that lack transiting planets over similar timescales. We show this variation can also be explained by the presence of starspots with surrounding facular regions. More observations both in- and out-of the transits of V1298 Tau c are required to determine the nature of the Ca II IRT and H α line variations.

Keywords: Stellar activity, exoplanet astronomy, exoplanet atmospheres, transmission spectroscopy

Features: Affiliated Python scripts/Jupyter notebooks 

1. INTRODUCTION

Young planets provide valuable insights into the early stages of planet formation and evolution. Through transit missions like *K2* (Howell et al. 2014) and *TESS* (Transiting Exoplanet Survey Satellite; Ricker et al. 2014) the population of young transiting planets (here defined as < 100 Myr) has grown to approximately a dozen (David et al. 2016; Mann et al. 2016; Benatti et al. 2019; Newton et al. 2019; Mann et al. 2020; Plavchan et al. 2020; Rizzuto et al. 2020). These planets are key to understanding planetary migration, atmospheric evolu-

tion, and measuring atmospheric mass-loss rate due to photoevaporation.

1.1. Obliquity Measurements

Close-in transiting exoplanets have brought into question the mechanisms in which they get there. Measuring the obliquity, λ , or the relative angle between the spin of the star and orbit of the planet (Lai 2014), can yield insights into different migratory paths. A low-obliquity hints at a smooth disk migration history (Goldreich & Tremaine 1979; Ford 2014), while binary companions or nearby stars in stellar birth clusters can torque planets to a high obliquity over millions of years (Fabrycky & Tremaine 2007).

Such measurements have only just begun for young systems. The recently discovered 20 Myr planet

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AU Mic b (Plavchan et al. 2020) has been the subject of significant follow-up. All independent studies have measured a low obliquity (Addison et al. 2020; Hirano et al. 2020; Martioli et al. 2020; Palle et al. 2020), suggesting the planet formed beyond the ice-line within the protoplanetary disk and experienced a smooth migration inward.

Similar studies of spin-orbit alignment have been completed for DS Tuc Ab, a planet in a slightly older (35–45 Myr) system, in a known binary (Newton et al. 2019; Montet et al. 2020; Zhou et al. 2020). The results from both studies concluded the planet has a low obliquity and that the migration was not significantly torqued through Kozai-Lidov oscillations by DS Tuc B. This is additionally supported by the fact that the age of the system is younger than the timescale required for Kozai-Lidov interactions (Montet et al. 2020). V1298 Tau has no known companions and thus any spin-orbit misalignment would hint towards potential torquing of the disk due to neighboring stars in the birth cluster.

1.2. Photoevaporation & Young Stellar Activity

Photoevaporation of young close-in planets is one of the leading explanations for the “radius gap” of transiting planets (Owen & Wu 2017; Fulton et al. 2017; Van Eylen et al. 2018; Bean et al. 2021). This gap is between a bimodal distribution of small planets, centered at 1.3 and $2.4 R_{\oplus}$ with a distinct lack of planets in between. However, early studies did not account for the ages of the systems; Berger et al. (2020) presented the first analysis of the planet-radius distribution while accounting for such ages. They found tentative evidence of an age dependence of the planet radius distribution, demonstrating that the number of super-Earths ($R/R_{\oplus} < 1.8$) increases with age, assuming a fixed radius gap location. This could be attributed to a core-powered mass-loss mechanism (Gupta & Schlichting 2019), which may continue for the first 1 Gyr (Rogers & Owen 2021). However, the general contributions of atmospheric removal from photoevaporation and core-powered mass-loss is currently unknown. While photoevaporation rate is a function of flux, these young transiting planets have short orbital periods and allow us to probe the most dramatic environments for atmospheric removal.

Atmospheric photoevaporation is driven by the high energy irradiation from young stars (Lammer et al. 2003; Owen 2019). The effects are particularly dramatic in the first 100 Myr where stellar XUV emission is higher and planets are still contracting, which makes their atmospheres more vulnerable to mass loss (Preibisch et al. 2005; Feigelson et al. 2007). Using *GALEX* observa-

tions, Shkolnik & Barman (2014) found that early M stars remain at high, saturated levels of NUV and FUV flux for the first \sim 650 Myr. The same relation was found of K stars (Richey-Yowell et al. 2019). The activity lifetime for low-mass stars is longer than that of solar-type stars. However, the higher levels of XUV flux makes photoevaporation more efficient for planets around low-mass stars (Rogers & Owen 2021), thus the evolution of planetary radii is faster (Owen & Wu 2013).

An increase in magnetic activity leads to elevated UV/X-Ray radiation in young stars, manifests itself in larger spot coverage, and results in increased photometric variability (Feigelson & Montmerle 1999). Several detailed studies of T Tauri stars have estimated that spots cover 29–41% of the less spotted hemisphere and 61–67% of the more spotted hemisphere (Grankin 1999), with the extreme case of LkCa 4 with of $\sim 80\%$ spot coverage (Gully-Santiago et al. 2017). Larger studies of young stars have also seen a lack of phase dependence for flares, another proxy for magnetic activity, indicating similar levels of high spot coverage across both hemispheres (Feinstein et al. 2020b). For detailed spectroscopic studies of young planets (e.g. studying their spin-orbit alignment or atmospheric characterization), it is crucial to understand the underlying starspot and active region coverage.

1.3. V1298 Tau

V1298 Tau is a 23 Myr pre-main sequence K star ($M_{\star}/M_{\odot} = 1.101$; $R_{\star}/R_{\odot} = 1.345$) hosting four known transiting planets (David et al. 2019a,b) observed in *K2* Campaign 4 (Howell et al. 2014). V1298 Tau is the first known young transiting multi-planet system. V1298 Tau was first identified as a young star via strong X-ray emission observed with *ROSAT* and strong Li I 670.8 nm absorption (Wichmann et al. 1996, 2000). V1298 Tau belongs to a young association in the foreground of the Taurus-Auriga star-forming complex (Oh et al. 2017) with an estimated age of 20–30 Myr based on comparison to empirical and theoretical isochrones (Luhman 2018; David et al. 2019b). V1298 Tau c ($R \sim 5.59 R_{\oplus}$) and d ($R \sim 6.41 R_{\oplus}$) are predicted to evolve to sub-Neptunes ($< 4 R_{\oplus}$) or super-Earths ($< 2 R_{\oplus}$), depending on the XUV evolution of the host star, after 5 Gyr (David et al. 2019a; Poppenhaeger et al. 2021). The light curve for V1298 Tau also shows strong starspot variability, with a peak-to-trough amplitude of $\sim 6\%$ (David et al. 2019b). Therefore, any in-transit spectra need to be carefully searched for the signatures of contaminating ARs and spots.

V1298 Tau b is seen to be well aligned through both a clear Rossiter-McLaughlin signal and Doppler tomographic analyses (Johnson et al. in prep). Additionally, David et al. (in prep) reports detections of excess H α absorption during transits of planets V1298 Tau b, hinting at highly extended atmospheres, which is expected for young planets.

Particularly, the excess absorption dominates in the red wing of the line and enhanced absorption at transit ingress of V1298 Tau b, which is confined to the core of the line. As the transit continues, there is an excess in absorption in the wings of the line and more so on the red side. The excess absorption during ingress could be the planet occulting a bright stellar surface or limb feature. Several observations during transits of V1298 Tau c and V1298 Tau d suggest tentative evidence for excess H α absorption (David et al. in prep). Observations do not span full transits, and therefore it is challenging to attribute the excess to the planets alone.

Here, we present a spectroscopic analysis during a full transit of V1298 Tau c, the innermost known planet in the system. We measure the spin-orbit alignment of the planet via Doppler Tomography and explore potential signs of atmospheric escape through changes in H α . The paper is presented as follows: Section 2 describes our observations and data analysis through both Doppler tomography and detailed line work. We present our results in Section 3. Section 4 discusses potential sources of the H α variations observed including center-to-limb variations, underlying starspots and faculae, and an extended planetary atmosphere from V1298 Tau c, and the source of the Ca II Infrared Triplet (IRT). We present our interpretation and future work in Section 5. We conclude by putting V1298 Tau c in context with other young planets and a summary of our key findings in Section 6.

2. OBSERVATIONS & METHODOLOGY

Due to the youth and high levels of stellar activity for V1298 Tau, we took several approaches to interpreting our observations. We observed other young active stars to compare spectral line variations to those seen in V1298 Tau. An updated ephemeris for V1298 Tau c was obtained from recent *Spitzer* observations (Livingston priv. comm.).

2.1. GRACES Observations

We obtained data with GRACES (Gemini Remote Access to CFHT ESPaDOnS Spectrograph; Chene et al. 2014) mounted on Gemini North at the Gemini Observatory on UT 2020 January 22 through the Fast Turnaround program. GRACES has a wavelength cov-

erage of 400-1000 nm and a resolving power of $R \sim 81,000$. To obtain an optimal signal-to-noise of approximately 300 at 600 nm, we observed V1298 Tau at 360 second exposures. V1298 Tau was monitored for approximately 6 hours in total, resulting in 53 exposures for this analysis. Observations were taken under photometric conditions. Our observations cover the entire transit of V1298 Tau c. Five exposures at the beginning and six at the end are out-of-transit (OOT). We note that a transit of V1298 Tau d ended approximately 12 minutes before our observations began.

2.1.1. Data Reduction

Raw and reduced spectra were provided by Gemini-N. The spectra were reduced using the OPERA¹ pipeline (Martoli et al. 2012; Teeple 2014). The OPERA pipeline completes several pre-reduction stages including creating master calibration images, a bad-pixel mask, and calibrating the instrument profile, aperture, and wavelength. The pipeline then extracts spectra from the original CCD images, and normalizes and telluric-corrects the spectra. However, we found a custom reduction of the data provides cleaner spectra for our analysis, via visual inspection.

We extracted each spectrum through the box-extraction method. We dark-subtracted and flat-corrected the spectra using master frames created by taking a median of all available dark and flat exposures. Cosmic rays were identified and removed using `ccdproc` (Craig et al. 2017), an extension of `astropy` (Astropy Collaboration et al. 2013; Price-Whelan et al. 2018) that tracks bad pixels in images.

To extract individual orders, we used five equally spaced columns across the frame and identified local minima across each column. The minima corresponded to regions between the orders. A second-degree polynomial was then fit to each of the five points in a given row to trace and extract the order. We visually inspected that each model corresponded to a single order, and did not cross over into a different order. We discretized the models and ensured the box was the same width across the orders. We then box-extracted 23 orders in each exposure.

To correct for the blaze, we removed the first and last 350 data points and created a 95th-percentile filter using `scipy` with a window length of 150 data points. For stars with broad lines, such as V1298 Tau, this is a standard procedure for identifying the blaze when there is no distinct continuum available (Pineda et al. 2013; Mon-

¹ Open source Pipeline for ESPaDOnS Reduction and Analysis

tet et al. 2015). An 8th-degree polynomial was fit to this filter and divided out of the spectra, resulting in our final corrected and flattened spectra (Figure 1). The number of cut data points and window length were visually found to result in the cleanest fully reduced spectra. Barycentric-corrected wavelengths, errors, and order numbers were taken from the OPERA-reduced data (Teeple 2014) and mapped onto our extracted spectra. Spectra of the same orders from different exposures were interpolated onto the same wavelength grid with three times the original wavelength resolution.

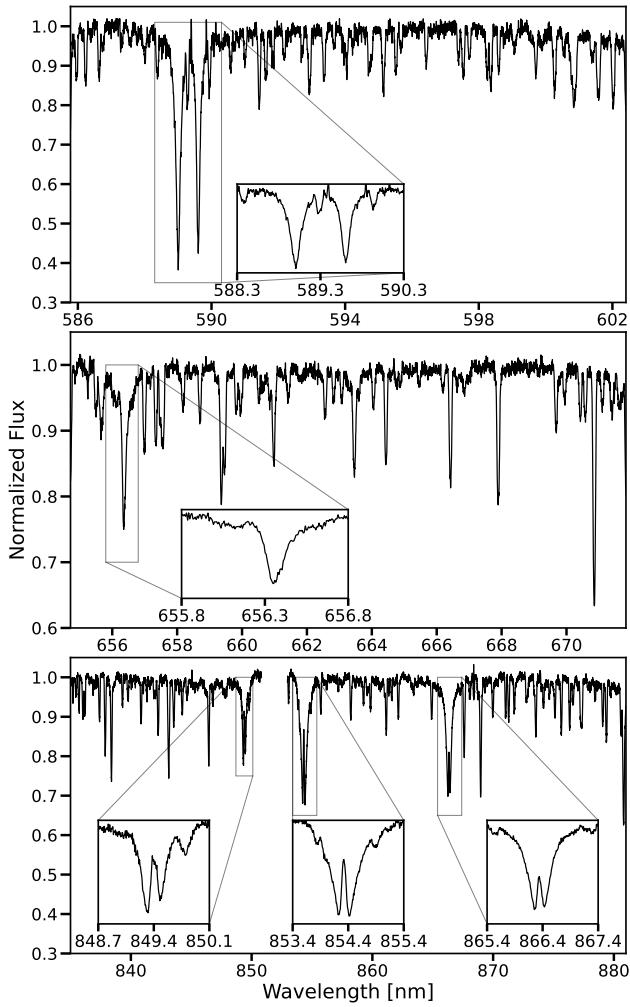


Figure 1. Example extracted spectra for V1298 Tau obtained with GRACES on Gemini-North. The bottom panel contains spectra from two different orders to highlight the Ca II IRT at 849.8, 854.2, and 866.2 nm. Inset panels highlight relevant lines used in this analysis including the Na I doublet (top), H α (middle), and Ca II IRT (bottom).

2.1.2. Doppler Tomographic Analysis

Doppler tomography relies on a transiting planet creating perturbations in rotationally broadened stellar line profiles that trace the orbit of the planet. Unlike typical radial velocity observations, which rely on tracing changes in the photocenter of lines, the Doppler tomography method allows us to resolve this perturbation in high resolution spectra. Here, we used the Doppler tomographic code MISTTBORN² (Johnson et al. 2014, 2017). In summary, this method compares line profiles between in-transit (IT) and OOT generated models. The initial guess for the line profiles are created by convolving a line list from the Vienna Atomic Line Database (Kupka et al. 2000) with a Gaussian line profile and are rotationally broadened given the $v \sin i$ of the system. MISTTBORN is flexible, such that after given an initial guess it fits for the line profile and allows it to take on an arbitrary best-fit shape. An average line profile is extracted from each order of each observation and is weighted by the signal-to-noise. The tomographic signal is determined by subtracting the IT from the average line profile across all OOT observations.

In addition to following the methods presented in Johnson et al. (2014), we focused on detailed analysis of the Fraunhofer lines (Fraunhofer 1817), which trace photospheric activity, and tried to extract the tomographic signal from just these lines. Each feature was compared to a median template created from the OOT observations. Of all the lines, H α and the Ca II IRT at 849.8, 854.2, and 866.2 nm were the only lines that showed variation throughout the night. All other lines appeared stable (see examples in Figure 4).

2.2. Veloce-Rosso Observations

As part of our analysis, we place the H α variability of V1298 Tau in context of other young stars of similar activity levels. During 7 nights between UT 2020-11-08 to 2020-11-26, we observed five young M dwarfs ($T_{\text{eff}} = 3323 - 3584$ K)³ with the Veloce-Rosso Spectrograph (Gilbert et al. 2018) mounted on the Anglo-Australian Telescope. The targets were originally identified as high-probability ($> 94\%$) young moving group candidates using BANYAN- Σ (Gagné et al. 2018), which uses *Gaia* kinematics to assign membership, and were presented in Feinstein et al. (2020b). Three of the targets (TIC 333680372, 246897668, and 178947176) are high-probability kinematic members of the AB Doradus Moving Group ($t_{\text{age}} \approx 150$ Myr) and the other two

² <https://github.com/captain-exoplanet/misttborn>

³ T_{eff} taken from the TESS Input Catalog v8 (TIC; Stassun et al. 2018).

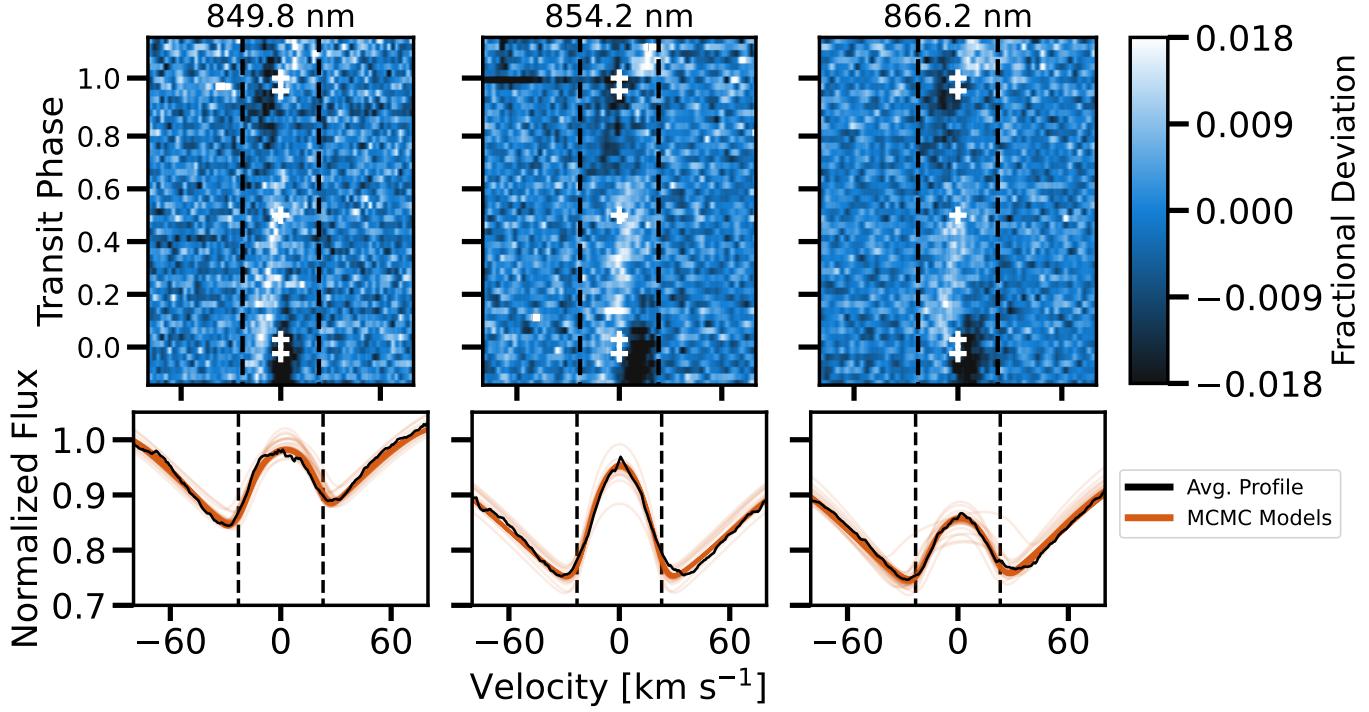


Figure 2. Top: Tomographic signal for each line in the Ca II IRT. Signals are plotted in the rest frame of the star. Excess absorption is shown in white and traces the transit. White pluses represent the four contact points of the transit and transit midpoint. Vertical dashed black lines represent $\pm v \sin i$. Bottom: 200 randomly selected MCMC fits (orange) used to derive the projected obliquity, λ , of V1298 Tau c, compared to the average line profile (black). Our MCMC best-fit parameters and associated priors are presented in Table 1.

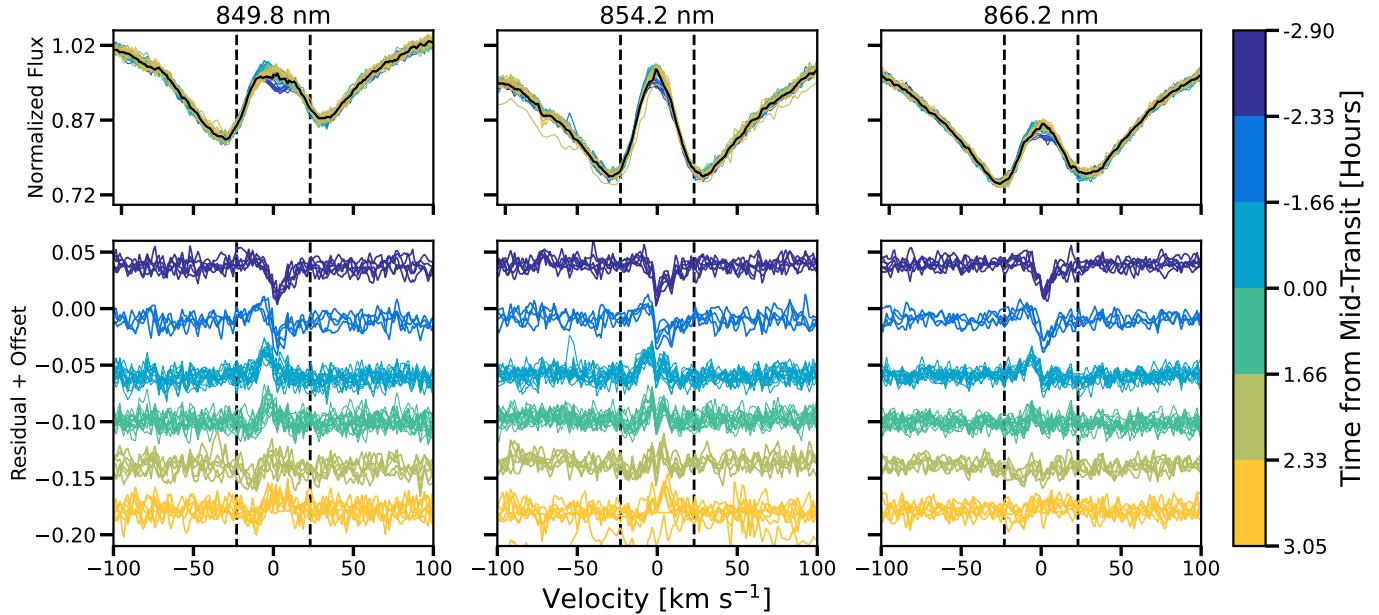


Figure 3. Zoom-in of the core and wings of each of the calcium triplet lines. Top: Normalized spectra. The OOT template is plotted as a solid black line. Bottom: The calcium feature with the OOT template subtracted out. There is a deficit in the red side of the core at the beginning of the night (purple), which quickly disappears. An excess in the blue can be seen during the beginning of the transit (dark blue) and slowly disappears through the remainder of the observations (teal to yellow). Lines are binned by time to mid-transit, where the transit duration is 4.66 hours; purple and yellow lines are OOT observations. Vertical dashed black lines represent $\pm v \sin i$, centered at the core of each calcium feature.

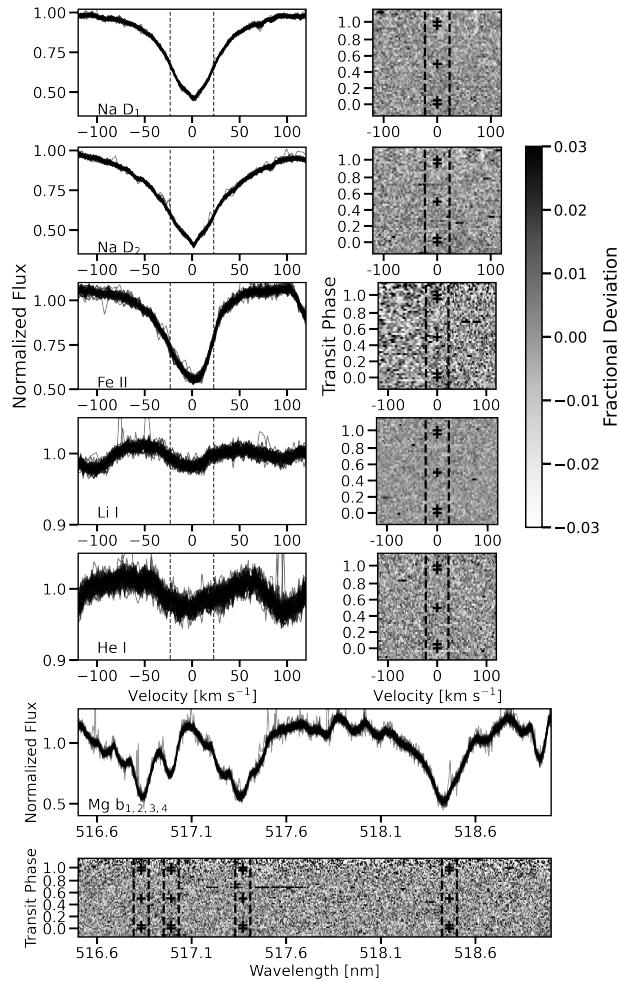


Figure 4. Spectral features (line plots) and associated waterfall plots for additional lines where we do not see a Doppler tomographic signal. Each line plot is labeled with the feature and the affiliated waterfall plot is located directly below. We present the Fe II line at 546.6 nm in the third row, Li I in the fourth row, and He I in the fifth row. The last two rows trace over several nanometers to cover the Mg I $b_{1,2,3,4}$ lines. Lines are colored by time of observation similarly to Figure 3. Waterfall plots represent deviations from a median OOT template from our observations. White pluses represent the four contact points of the transit and transit midpoint. Vertical dashed black lines represent $\pm v \sin i$.

(TICs 146522418 and 427346731) are high-probability kinematic members of the β Pictoris Moving Group ($t_{age} \approx 25$ Myr; Bell et al. 2015).

BANYAN- Σ probabilities are strictly based on kinematic arguments and do not guarantee these stars are members of assigned known groups. However, the rapid rotation seen in the *TESS* photometry for these stars is consistent with young ages. Spectroscopic youth indicators, such as strong H α emission are present in all of our observations, providing further evidence these stars are likely

young and potentially members of the AB Doradus and β Pictoris moving groups.

Targets were observed for a total of 70-120 minutes per night of observation, with exposure times ranging from 300-600 seconds. Data were taken simultaneously with observations of a laser comb, used for precise wavelength calibration (Murphy et al. 2007). The laser comb was additionally used to trace individual orders for box extraction. The observed targets have no known transiting planets. We reduced all spectra using methods similar to those described in Section 2.1.1. The wavelengths were mapped using a pixel-to-wavelength reference made available by the Veloce-Rosso team.⁴

Targets were observed simultaneously with *TESS*. Due to the youth of these stars, it was possible some observations were taken during flare events. The *TESS* 2-minute light curves were searched for flares using the convolutional neural networks developed by Feinstein et al. (2020b). These neural network models assign a probability that a given event is a flare or not. We considered events “true flares” when they were assigned a probability ≥ 0.9 . Times of flares were cross-matched with our observation; spectra that were taken during identified flare events were removed from this analysis.

3. RESULTS

Because we found our spectra to show little variation over the night, we could not use all spectral features to extract a Rossiter-McLaughlin or Doppler tomographic signal. Of the Fraunhofer lines checked, including the Mg I triplet at 517 nm, and Na D₁ and D₂ (Figure 4), the tomographic signal was only seen in the Ca II IRT at 849.8, 854.2, and 866.2 nm.

3.1. CaII Infrared Triplet

3.1.1. Doppler Tomography

The Doppler tomographic analysis of the Ca II IRT is shown in Figure 2. The fractional deviation compares the difference between the Ca II IRT lines of each observation with a median profile from all OOT data. Bright areas represent regions of excess absorption, here due to the planet crossing the surface of the star, tracing from the bottom left to top right and approximately from the beginning to end of the transit. The transit signal is seen independently in all three lines. We fit for the tomographic signal in the core of the Ca II IRT lines using MISTTBORN (Johnson et al. 2014). The line profile model in MISTTBORN accounts for the distortion of the

⁴ <https://newt.phys.unsw.edu.au/~cgt/Veloce/Veloce.html>

line shape during the planet's transit and is a function of the parameters presented in Table 1. We do not account for differential rotation or macroturbulence.

MISTTBORN is typically used to model individual lines and not spectral features with cores. Here, we combined the line core with an upside-down Gaussian to model the regions around the core, presented in Figure 2. For Ca II at 854.2 nm, we set the central velocity of the Gaussian equal to the fitted line center; the other two Gaussian line centers were allowed to vary due to noticeable asymmetry in the lines. We used emcee (Goodman & Weare 2010; Foreman-Mackey et al. 2013) to fit each line. We ran our MCMC fit with 350 walkers and 750 steps. After visual inspection, we removed 150 burn-in steps. We verify our chains have converged via visual inspection and following the method of Geweke (1992). The parameters, fit results, and priors are given in Table 1. The median Ca II IRT and 200 randomly selected MCMC examples are shown in the bottom row of Figure 3.

From this, we measure a projected obliquity of $\lambda = 5^\circ \pm 15^\circ$. We find the Doppler tomographic signal is driven by changes in the core of the Ca II IRT, formed in the stellar chromosphere (Figure 3). The individual lines (top row) and residuals with a median OOT template (bottom row) are colored by transit phase of V1298 Tau c. The changes in the lines are visible by eye, and are more pronounced in the residuals. There is very little difference between each observation in the wings of the lines (regions outside of $\pm v \sin i$).

3.1.2. Weighted Mean “Light Curves”

Comparing the weighted contributions of spectral features of IT to OOT observations has been used to infer the presence of planetary atmospheres, in particular the Na I D lines (Charbonneau et al. 2002; Wyttenbach et al. 2015). For this, we follow a similar approach to (Barnes et al. 2016). Per each spectral feature, we calculate:

$$x_i(F_j, \sigma_j^2) = \sum_{j=0}^m (F_j / \sigma_j^2) / \sum_{j=0}^m (1 / \sigma_j^2) \quad (1)$$

where x_i refers to a specific spectral feature, F_j is the flux array and σ_j^2 is the variance array for j observations. The total weighted mean “light curve” across multiple spectral features is then given by

$$\bar{x}(F_j, \sigma_j^2) = \sum_{i=0}^n x_i \quad (2)$$

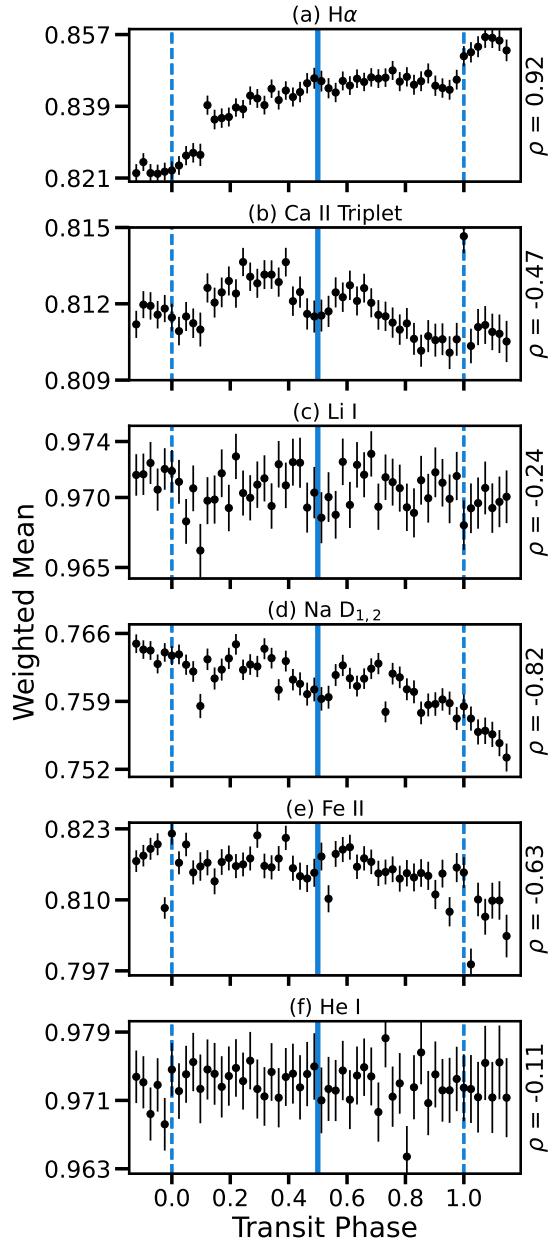


Figure 5. Weighted mean “light curves”, \bar{x} , and affiliated errors, $\sigma_{\bar{x}}$, of the (a) H α , (b) Ca II IRT, (c) Li I at 670.7 nm, (d) the Na I doublet at 589 nm, (e) Fe II at 546.6 nm, and (f) He I at 587.8 nm. The Spearman’s correlation value, ρ , for each feature is presented on the right hand side. Note the scale for each subplot was chosen to optimize the entire region and error bars are comparable. The solid line marks t_{mid} and the dashed lines represent t_1 and t_4 . An increase in weighted mean corresponds to excess absorption in the spectral feature. There is a visible increase in Ca II with transit ingress and egress, while no such trend is seen in the Li I. This provides additional confidence the deviations in Ca II is planetary in nature. The spectra for additional lines are presented in Figure 4.

Table 1. Doppler Tomographic Modeling Results for the Ca II IRT

Planet Parameters	849.8 nm	854.2 nm	866.2 nm	Combined	Prior
λ [°]	$13.6^{+16.9}_{-8.8}$	$-0.5^{+9.3}_{-22.7}$	$4.2^{+5.8}_{-13.8}$	$4.9^{+15.0}_{-15.1}$	$\mathcal{U}[-180, 180]$
T ₀ [MJD]	$58846.09709^{+0.00008}_{-0.00019}$	$58846.09720^{+0.00018}_{-0.00018}$	$58846.09716^{+0.00017}_{-0.00015}$	$58846.09715^{+0.00016}_{-0.00019}$	$\mathcal{G}(58846.097156, 10^{-5})$
Period [Days]	$8.24964^{+0.00019}_{-0.00012}$	$8.24962^{+0.00021}_{-0.00011}$	$8.24953^{+0.00013}_{-0.00020}$	$8.24959^{+0.00022}_{-0.00018}$	$\mathcal{G}(8.24958, 10^{-5})$
R _p /R _*	$0.03923^{+0.00034}_{-0.00013}$	$0.03926^{+0.00063}_{-0.00012}$	$0.03869^{+0.00052}_{-0.00088}$	$0.03921^{+0.00033}_{-0.00062}$	$\mathcal{G}(0.039208, 10^{-3})$
v sin i [km s ⁻¹]	$23.59^{+0.97}_{-0.73}$	$23.92^{+0.61}_{-1.25}$	$23.45^{+0.99}_{-1.71}$	$23.65^{+0.87}_{-1.17}$	$\mathcal{G}(23.0, 1.0)$
a/R _*	$13.188^{+0.15}_{-0.136}$	$13.187^{+0.003}_{-0.005}$	$13.218^{+0.185}_{-0.088}$	$13.189^{+0.133}_{-0.047}$	$\mathcal{G}(13.19, 0.01)$
impact parameter, b	$0.20^{+0.11}_{-0.12}$	$0.21^{+0.12}_{-0.09}$	$0.22^{+0.10}_{-0.09}$	$0.21^{+0.11}_{-0.10}$	$\mathcal{G}(0.20, 0.05)$
eccentricity, e	$0.11^{+0.03}_{-0.01}$	$0.09^{+0.02}_{-0.05}$	$0.09^{+0.03}_{-0.04}$	$0.10^{+0.03}_{-0.04}$	$\mathcal{U}[0, 0.5]$
periastron, ω [°]	$90.18^{+59.52}_{-65.7}$	$87.27^{+59.61}_{-63.2}$	$86.4^{+51.07}_{-56.23}$	$87.84^{+57.05}_{-61.78}$	$\mathcal{U}(-180, 180)$
limb darkening, u ₁	$0.41^{+0.21}_{-0.10}$	$0.86^{+0.13}_{-0.49}$	$0.63^{+0.25}_{-0.26}$	$0.55^{+0.38}_{-0.19}$	$\mathcal{U}[0, 1]$
limb darkening, u ₂	$0.21^{+0.12}_{-0.14}$	$0.30^{+0.21}_{-0.15}$	$0.35^{+0.20}_{-0.09}$	$0.30^{+0.19}_{-0.17}$	$\mathcal{U}[0, 1]$
line center [km s ⁻¹]	$-1.31^{+0.02}_{-0.68}$	$-7.29^{+0.19}_{-0.02}$	$-14.72^{+0.67}_{-0.11}$	—	$\mathcal{U}[-20, 20]$
Gaussian Parameters					
mean, μ , [km s ⁻¹]	$4.99^{+0.19}_{-0.26}$	—	$-12.57^{+1.55}_{-0.99}$	—	$\mathcal{U}[-20, 20]$
width, σ	$39.36^{+3.55}_{-0.72}$	$57.66^{+1.43}_{-2.37}$	$53.30^{+1.74}_{-1.75}$	—	$\mathcal{U}[35, 60]$
scaling factor	$23.56^{+1.62}_{-0.38}$	$41.53^{+1.92}_{-2.71}$	$36.99^{+2.28}_{-1.68}$	—	$\mathcal{U}[10, 50]$
core scaling factor	$0.176^{+0.005}_{-0.012}$	$0.24^{+0.003}_{-0.003}$	$0.145^{+0.003}_{-0.004}$	—	$\mathcal{U}(0, 1)$
model y-offset	$1.04^{+0.01}_{-0.0}$	$1.00^{+0.01}_{-0.01}$	$0.99^{+0.01}_{-0.01}$	—	$\mathcal{U}[0.5, 1.5]$

NOTE—Ca II at 852.4 nm was symmetric about the core so the mean of the underlying Gaussian, μ , was not fit separately from the line center.

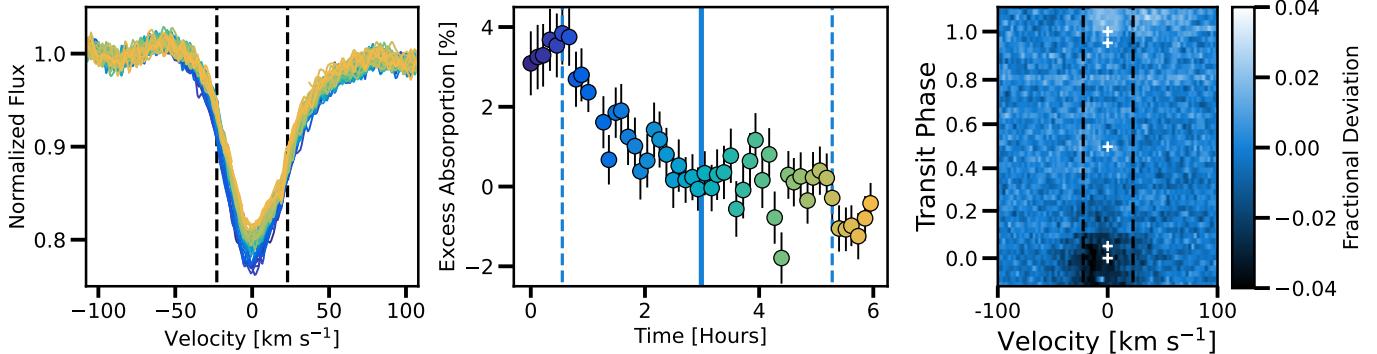


Figure 6. Demonstration of variation in H α observed in our observations. Left: Normalized spectra colored by observation time. Purple and yellow represent the beginning and end of the night, respectively. The smooth variation in depth of the H α feature is clear. Middle: Measurement of excess absorption seen in H α as a function of time. The color of points correspond to color of the spectra in the left plot. The solid line corresponds to t_{mid} . The dashed lines correspond to transit contact points t_1 and t_4 . Right: Tomography of H α colored by fractional deviation from the OOT template. White pluses represent the four contact points of the transit and transit midpoint. Vertical dashed black lines in the left and right panels represent $\pm v \sin i$.

where n is the total number of combined features (e.g. for the Ca II IRT, $n = 3$). In a similar manner, such that the associated standard deviation for a given feature is:

$$\sigma_i(\sigma_j^2) = \sum_{j=0}^m \frac{1}{\sigma_j^2} \quad (3)$$

where σ_i refers to errors for a specific spectral line and σ_j^2 is the variance array for j observations. The total

weighted standard deviation across i spectral features is calculated as:

$$\sigma_{\bar{x}}(\sigma_j^2) = \left(\sum_{i=0}^n \sigma_i \right)^{-1/2} \quad (4)$$

We compare a weighted mean of the the Ca II IRT lines to various other lines and activity indicators, such as Li I at 670.7 nm, the Na I doublet at 589 nm, and the He I

line at 587.8 nm; the results are shown in Figure 5. An increase in weighted mean can be interpreted as excess absorption in the spectral feature.

Li I has been seen to remain relatively stable in the presence of high levels of activity and photometric spot variations (Pallavicini et al. 1992, 1993; Soderblom et al. 1993), with a potential abundance increase in the presence of the more-spotted hemisphere (Flores Soriano et al. 2015). We measured the Li I equivalent width, and found the average to be 140 mÅ, which is consistent with other K0-K1.5 stars of this age (Wichmann et al. 2000). As a comparison, we use Li I as a control to what is intrinsically occurring in the stellar chromosphere. We test for correlations between transit phase and the light curves for Li I and Ca II IRT in Figure 5 by calculating the Spearman’s rank correlation coefficient (Spearman 1907). This test assess how well two data sets can be described using a monotonic function. A correlation coefficient of ± 1 implies an exact monotonic relationship, while 0 implies no correlation. We find correlation values of -0.24 and -0.65 for Li I and Ca II IRT respectively, implying Ca II IRT is slightly more correlated with transit phase than Li I . Li I could be enhanced from flare events; however, we see no evidence of such event in other spectroscopic features affected by flares (e.g. variations in the blue wing of $\text{H}\alpha$; Maehara et al. 2021).

3.1.3. Correlations Between Spectral Features

We evaluate other known activity indicators, such as Na I doublet, Fe II , and He I (subplots d - f in Figure 5). We calculated the Spearman’s correlation value for each to be -0.80, -0.63, and -0.11, respectively. High Spearman’s correlation values could be due to long term stellar activity trends seen in the data. This could be particularly true for the Na I doublet. Although trends in the weighted mean is similar to Ca II IRT, there is no evidence of a Doppler tomographic signal in the lines (Figure 4). For completeness, we also evaluate the Spearman’s correlation value for $\text{H}\alpha$ (subplot a in Figure 5) and find it to be 0.93, indicating a strong correlation with transit phase. We further inspect the $\text{H}\alpha$ signal below.

We additionally evaluate the Spearman’s correlation value between each of the spectral features in Figure 5. In order to look for correlations in the features, rather than long term trends, we fit a first degree polynomial to each weighted mean light curve. From the trend-removed light curve, we found the correlation between each line. The correlation values are shown in Figure 7. The median correlation value between all the lines is 0.1, which indicates there is very little to no correlation. The

lines that are most strongly correlated are $\text{H}\alpha$ and the Ca II IRT. Once the long term trend is removed from the $\text{H}\alpha$, there is a clear upwards “bump”, similar to what is seen in Figure 6. Due to the strong correlation between that “bump” and the increase that is seen in the Ca II IRT, which is planetary in nature, this could indicate that the trend in $\text{H}\alpha$ is also originating from V1298 Tau c.

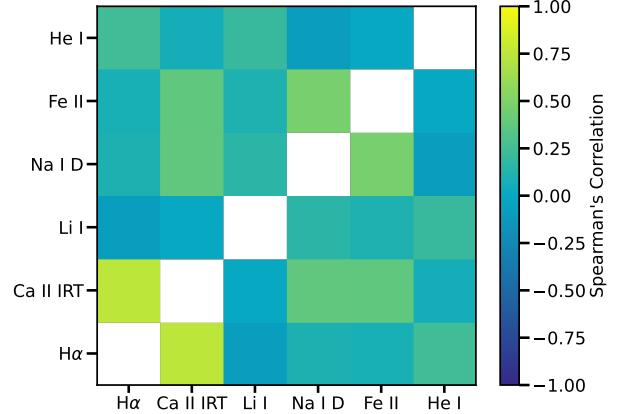


Figure 7. The Spearman’s correlation value for spectral features presented in Figure 5. The average Spearman’s value across all spectral features is 0.1, indicating that there is no correlation. The strongest correlation exists between $\text{H}\alpha$ and the Ca II IRT.

3.2. $\text{H}\alpha$ Variations

The $\text{H}\alpha$ variations are shown in Figure 6. The leftmost panel is the $\text{H}\alpha$ line and surrounding continuum used in our analysis. By eye, there is a clear trend in decreasing absorption throughout the night. To quantify this trend, we calculate the equivalent absorption (EA) of each line, defined as:

$$\text{EA} = - \sum_{i=0}^n \frac{F_i}{F_{out}} - 1 \quad (5)$$

where F_i is the line flux for one observation and F_{out} is the median template of the OOT observations. The equivalent absorption for $\text{H}\alpha$ is shown in the middle panel of Figure 6. There is a trend of decreasing excess absorption throughout the night, with potentially sharp features after V1298 Tau c is fully visible across the stellar surface. This could be the presence of an extended hydrogen atmosphere. However, given the long term trend and minimal OOT observations, it is challenging to disentangle this signature from general stellar activity.

The Doppler tomographic analysis is shown in the rightmost panel of Figure 6. Dark and light regions at the beginning and end of the night are the result of over-subtraction from the template, which is constructed from these specific observations. There is no clear tomographic signal seen between $\pm v \sin i$, as was seen in Figure 2. This is likely due to stellar activity dominating the long-term trend in H α seen over the night. The excess absorption seen in the middle panel is therefore most likely stellar activity.

3.2.1. Comparison to Veloce-Rosso Targets

The EA of H α was additionally calculated for the young stars in our Veloce-Rosso sample using Equation 5. We note that the activity of M and K stars are comparable at ages ≤ 1 Gyr (Schneider & Shkolnik 2018; Richey-Yowell et al. 2019). Thus, this comparison will help provide a stellar context for the H α variations. The results are shown in Figure 8. We note that the H α for targets observed with Veloce-Rosso is in emission, while V1298 Tau clearly shows H α in absorption without a significant core component.

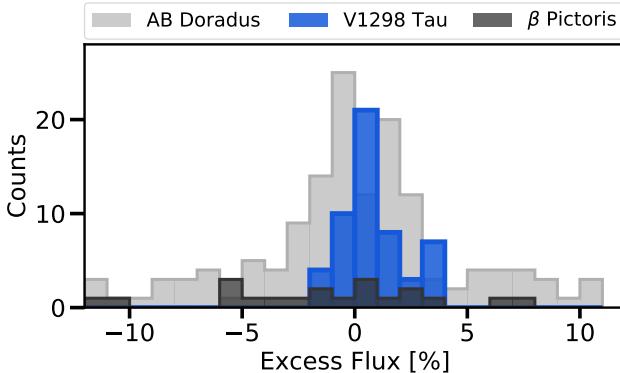


Figure 8. Comparison of H α variations observed in V1298 Tau (purple) with five young M stars observed with Veloce-Rosso. The levels of H α emission in these young stars varies on the order of a few percent over a few hours. The H α for V1298 Tau is correlated with time, hinting this may be the signature of an escaping hydrogen atmosphere. However, observations of other young active stars show similar trends over a single night. The H α excess variations of V1298 Tau trace the distribution of AB Doradus members. This is likely because the relative activity of a 23 Myr K star is comparable to a 150 Myr M star. The light gray histogram are members of AB Doradus moving group ($t_{age} = 150$ Myr); dark the gray histogram dark gray points are members of the β Pictoris moving group ($t_{age} = 25$ Myr). Veloce-Rosso observations during flare events were removed. Each bin represents 1% variation in excess flux.

A median template of H α was created per star per night of observation. The excess flux variation is calculated by comparing each exposure to the median template. This is similar to the way excess H α flux for V1298 Tau was calculated. There is clear variation for all young stars, with excess variation extending out to $\pm 10\%$. The predominant amount of excess flux varies between $\pm 3\%$, which follows the trend seen for V1298 Tau. We additionally found stars in our sample exhibit smooth increasing or decreasing H α variation over a single night, in a similar way to V1298 Tau (Figure 6, middle panel). This demonstrates the behavior of H α in young stars can change dramatically over a night.

The average strength H α emission for young stars is comparable to what is seen in our Veloce-Rosso sample. AB Doradus and other confirmed young moving group members have H α EWs ranging from -10 to 0 , where negative values indicate emission (Riedel et al. 2017). Confirmed β Pictoris moving group members were seen to have H α EWs ranging from ~ -15 to 0 for M0-M9 stars (Shkolnik et al. 2017). Schneider et al. (2019) measured H α EWs in 336 young low-mass stars (K5-M9; $t_{age} \sim 22 - 200$ Myr) with values ranging from ~ -20 to 2 , where earlier spectral types tend towards positive EWs. The values for our H α EA measurements fall well within the ranges for other young stars with ages between $t_{age} \sim 23 - 150$ Myr. In the context of other young stars, the variations seen in V1298 Tau can be solely attributed to youth and general stellar activity.

4. DISCUSSION

We speculate on several different explanations for the observed H α variations, such as limb darkening, the presence of star spots, and the presence of an extended H α atmosphere from V1298 Tau c or V1298 Tau d.

4.1. Center-to-Limb Variations

The underlying wavelength dependence of limb darkening can create signals masquerading as those expected from planet atmospheres (Czesla et al. 2015). We model center-to-limb variations (CLVs) to address this possibility. A synthetic spectrum was obtained from the Castelli & Kurucz (1994) grid of stellar atmospheric models generated with ATLAS9 (Kurucz 1993). Based on stellar parameters presented in David et al. (2019b), we chose a model with $T_{eff} = 5000$ K and $\log(g) = 4.00$ with solar metallicity. Spectra with different limb darkening values were generated using `spectrum` (Gray & Corbally 1994) with a wavelength step size of $\Delta\lambda = 0.002\text{\AA}$. The spectra were additionally rotationally broadened ($v \sin i \sim 23 \text{ km s}^{-1}$) following the formalism of Gray

(2008). A total of 50 synthetic spectra were generated from $u = 0 - 1$, where $u = 0$ is the edge of the stellar disk and $u = 1$ is the center. We sampled limb angles at $\Delta u = 0.02$; $u = 0.0$ was replaced with $u = 0.001$ (Czesla et al. 2015).

We created a discretized stellar surface with $R_{pixels} = 20001$, filled with concentric circles for each value of u as prescribed by Vogt et al. (1987). An example of the modeled stellar disk and accompanying spectral feature for one of the Ca II IRT lines at 849.8 nm is shown in Figure 9. Variations in line depth as a function of u are highlighted in the inset panel. We model the transit of V1298 Tau c with a radius of 19 pixels, maintaining an $R_p/R_\star = 0.03$ and at an impact parameter of $b = 0.2$ (David et al. 2019a). The spectrum at each step is the summation of all unocculted elements.

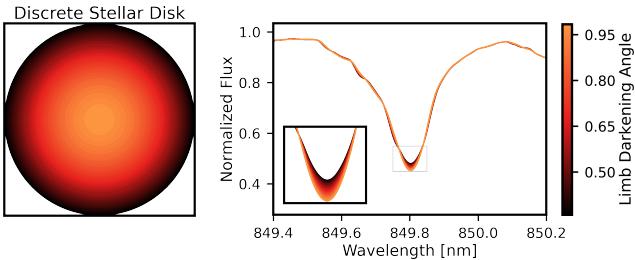


Figure 9. Modeled stellar disk and accompanying spectra for V1298 Tau at different limb darkening angles. The original spectrum was taken from the grid of stellar atmospheric models from Kurucz (1979). Spectra were then generated at different limb darkening angles using `spectrum`. We accounted for rotational broadening due to high $v \sin i$ of V1298 Tau. The averaged effect of limb darkening over the entire stellar surface is negligible with respect to the H α variation seen during our observations.

We find that CLVs for this system induce changes in the spectral features on the order of 10^{-5} . These effects are negligible compared to the observed 8 ppt signal (Figure 5). The CLV testing described above does not account for non-local thermodynamic equilibrium effects. Although it would be more rigorous to include such calculations, the magnitude of the variations is negligible compared to the contribution of ARs (Cauley et al. 2018a). Therefore, the behavior seen here is not attributable to CLV.

Next, we account for potential limb-brightening in the core of the Ca II IRT lines, following the methods presented in Czesla et al. (2015). First, we create a differ-

ence light curve (DC) with a modified version of Equation 9 to account for three, instead of two, lines:

$$DC_{\text{core}}(t_i) = \frac{1}{3} \sum_{j=0}^3 LC_j(t_i) - \frac{1}{3} \sum_{j=0}^3 LC_{C_j}(t_i) \quad (6)$$

where t_i is an observation, LC_j is the mean weighted light curve of each of the three individual Ca II IRT line, j , and LC_{C_j} is the mean weighted light curve for the continuum around each Ca II IRT line. Due to there being many lines near the Ca II IRT, we select the continuum region as 0.5 nm on either side of the lines. We find the effect of limb brightening to be on a similar scale to that in Czesla et al. (2015), but does not exhibit the same trend. While limb brightening could be responsible for this signal, it is near-impossible to distinguish between variability from the planet and CLV-induced changes, and therefore we rule out limb-brightening as the source of this variability.

4.2. Stellar Nature of the Variability

Due to H α being a tracer of stellar activity, we must explore the possibility that the observed excess absorption variations are strictly stellar in nature. The smooth variation over our observations was seen in similar observations for targets in our Veloce-Rosso sample. Additionally, young stars are believed to have significant photospheric inhomogeneities, seen in the original light curve of V1298 Tau. Similar night-long trends in activity have been seen for other young planet hosts (Montet et al. 2020) and the trend in the excess H α absorption during our observations can be explained by stellar activity alone.

4.2.1. Spot & Faculae Modeling

We use `starry` (Luger et al. 2019), an analytic solution to time series of planet transits and stellar surfaces based on applications of spherical harmonics, in combination with `spectrum` to recreate the H α absorption with starspots and surrounding facular regions. We follow the methods discussed in Section 4.1 and add in the presence of two starspots with surrounding facular regions. We take a stellar spectrum model of the same log(g) and metallicity as V1298 Tau at 5500 K and 4500 K for the faculae and spots respectively. Although these T_{eff} may not accurately represent all starspot/facular regions, they provide a reasonable estimation for determining changes in line depths/profiles in the presence of such features (Catalano et al. 2002).

We created `starry` models as a function of rotation that are 2500×2500 pixels in size. Then we discretized

the surface to indicate values of either surface (orange), starspot (red), or facula (yellow); examples are shown in the top row of Figure 10. The approximate area of spot to facula is 60%, similar to that in Cauley et al. (2017) and the associated modeled photometric variability (Figure 10, middle) corresponds to variability seen in the original *K2* light curve (David et al. 2019b). Forward modeling of *K2/TESS* light curves also suggests the variability is driven more strongly by starspots than faculae (Johnson et al. 2021). We weighted the spectra by the area of each feature (spot, facula, or surface) and accounted for limb-darkening.

The configuration explored in our toy model is just one of many potential solutions and was not fit to the original *K2* light curve, but rather captures a general scale of spot variability. Therefore, it should not be taken as the ground-truth configuration (Luger et al. 2021). A six-hour window is highlighted in the photometric variability in red (Figure 10, middle), which corresponds to the H α excess absorption seen in the bottom row of the same figure. Our observations (gray) show a decrease in excess H α absorption of 4.29% over six hours; this model (red) has a similar trend, decreasing by 4.37% over the simulation.

4.3. Planetary Nature of the Variability

We speculate on several different planetary configurations that may be the source of the H α and Ca II IRT variations seen over our observations. Each scenario is evaluated separately, although the ground truth could be the result of some combination from each.

4.3.1. Extended H α Atmosphere of V1298 Tau c

Given the age of the system and current X-ray irradiation, V1298 Tau c is predicted to evolve significantly from $5.59R_{\oplus}$ to a final radius of $1 - 5.55R_{\oplus}$ given the different mass and X-ray activity level assumptions (Poppenhaeger et al. 2021). It would therefore not be surprising to detect a highly extended atmosphere from our target. We note that slight “bumps” in the measured H α equivalent width at transit ingress and egress for V1298 Tau c (Figure 6; middle panel) could be attributed to the planet, while the overall slope is dominated by stellar activity. We fit a line between the equivalent widths for ingress and egress to represent the activity longer-term trend and measured a depth in the excess absorption of 1.85%. Assuming spherical symmetry, this would correspond to an atmosphere with a thickness of $1.28 R_J$.

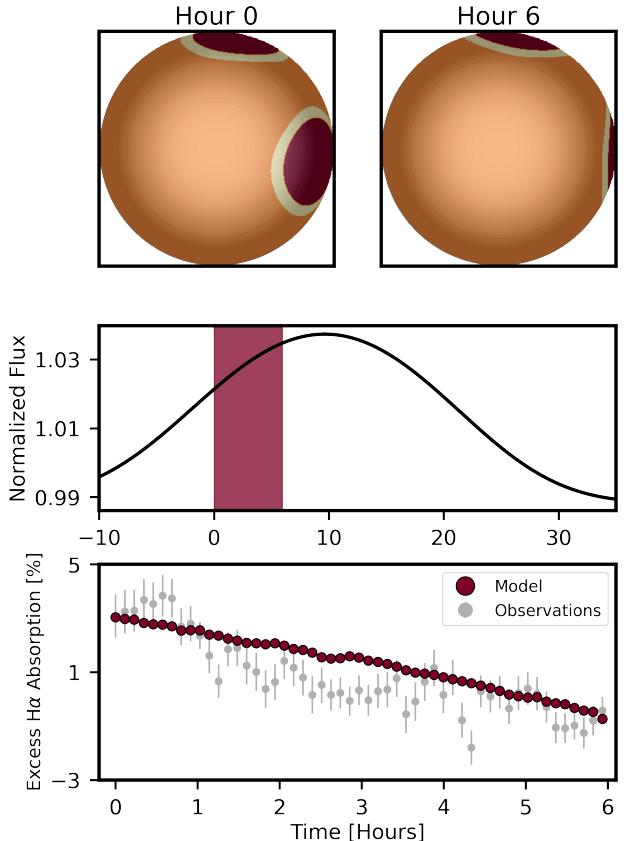


Figure 10. Toy spot and facula model created using `starry` of a potential configuration on V1298 Tau that produces similar levels of H α variation to that seen in our observations. Top row: Surface maps at the beginning and after 6 hours of rotation. Dark red regions represent starspots; bright yellow represent surrounding facular regions; orange is the surface. Maps are limb-darkened. The starspots and faculae comprise of $\sim 20\%$ of the visible surface at the start. Middle row: Normalized light curve of this starspot configuration. Variation amplitudes match those that are present in the original *K2* observations (David et al. 2019b). Shaded red region represents the 6 hour window for the above rotation configurations. Bottom row: Measured H α excess absorption at each rotation time step. Identifying individual starspots is a degenerate issue; this is one configuration that is able to reproduce a similar trend in H α excess absorption over the timescale of our observations.

4.3.2. H α Tail from V1298 Tau d

With updated *Spitzer* ephemerides, we also find that the transit of V1298 Tau d ended 12 ± 15 minutes before our observations (Livingston priv. comm.). V1298 Tau d is a $6.41R_{\oplus}$ planet on a 12.4 day period and is predicted to evolve to $1.5 - 6.4R_{\oplus}$ in 5 Gyr (Poppenhaeger et al. 2021). If an atmospheric tail is present from V1298 Tau d, it could be contaminating our observations at the beginning of the night. The first seven observa-

tions ($t_{\text{MJD}} = 58870.725 - 58870.753$) show a relatively stable H α excess absorption of 5% before decreasing. This could very well mean that our initial observations are catching the tail end of an inflated hydrogen atmosphere around planet d.

In this scenario, we can estimate the extent of the hydrogen atmosphere. If we assume the transit of the extended atmosphere continues for the first 2.5 hours, up until the excess H α absorption temporarily flattens out, we see the signal change $\Delta 5\%$. A 5% signal would correspond to an atmosphere thickness of $2.36 R_J$. Assuming the atmospheric signal continues for 2.5 hours, this would correspond to a tail length of $5.85 R_J$. The remaining observations would be affected by the transit of V1298 Tau c, with a potential active region crossing resulting in increased noise in our H α measured EAs between $t_{\text{MJD}} = 58870.871 - 58870.908$ (corresponding to 3.504-4.392 hours in the middle panel of Figure 6).

4.3.3. Ca II IRT from V1298 Tau c

We investigate if the Ca II is an extended ionized atmosphere of the V1298 Tau c. We evaluate the behavior of the Ca II IRT transmission signal in both the star and planet's rest frames (Figure 11) and evaluate the transmission signal for H α (rightmost column). The variations in Ca II IRT are seen to vary with planet orbital phase in both rest frames. We can rule-out that the Ca II IRT is part of an extended atmosphere from V1298 Tau c because it is seen in emission and is not at rest in the planet's rest frame. The H α line changes also are not at rest in the planet's rest frame and highlight the $\pm v \sin i$ change in absorption over the night.

Close-in planets interact with their host star's magnetic field lines, which have been seen to manifest as changes in flux from the cores of chromospherically driven lines. Such changes have previously shown variations with the planet's orbital phase (Cauley et al. 2019). Shkolnik et al. (2008) showed that the equivalent width of Ca II H & K varied consistently with the orbital phase of HD 179949 b over 5 years of observations, which is indicative of a magnetically-driven star-planet interaction. They also found a strong correlation between Ca II H & K and the mean emission from the chromospheric core of the Ca II IRT. Similar studies have been conducted for HD 189733 b, where no clear evidence of star-planet interactions were seen via a similar analysis of Ca II H & K (Fares et al. 2010). Following studies using 10 years of spectral observations of HD 189733 b found Ca II H & K variations corresponding to transits, but with a leading phase offset of $\Delta\phi = 40 - 53^\circ$ (Lanza 2012; Cauley et al. 2018b).

It is possible the Ca II IRT variations in our observations are driven by star-planet interactions. Given the increased magnetic activity of young stars and the close proximity of young planets, these interactions may be expected. However, with limited orbital phase coverage it is difficult to say if this is the case. Multiple transits of V1298 Tau c are required to determine if the Ca II IRT varies regularly with orbital phase. Simultaneous observations Ca II H & K would also prove useful here.

5. INTERPRETATION & FUTURE WORK

5.1. Interpretation

In Section 4.1, we explored the possibility of our observed H α trend to be the result of CLV. We can confidently rule this out as being the source of the equivalent width variations due to the small effect we find ($\approx 10^{-5}$) compared to our observed signal (8 ppt).

There appears to be slight indications of a transit in H α (Figure 6) with an underlying long term stellar activity trend. However, there is no visible Doppler tomographic signal. The change in excess H α absorption is also consistent with that seen for other young stars (Figure 8). We also find no evidence of H α changes in the planet's frame of reference (Figure 11, rightmost column). We also show such H α variations can be explained by a stellar surface with 20% coverage by starspots and faculae (Figure 10). A 20% inhomogeneity coverage, like that at the start of our toy model, is possible for young stars.

With no previous constraints on an extended atmosphere from V1298 Tau c or V1298 Tau d, it is difficult to address how much this could be contaminating our OOT observations at the beginning of the night. However, since the H α excess absorption variations are seen in other young stars and they can be explained solely by the presence of stellar inhomogeneities, we speculate the H α is most likely stellar in nature.

We only find the Doppler tomographic signal in Ca II IRT. If this is driven by star-planet interactions, then it would be expected to be seen in other lines that originate in the chromosphere. Since our observations do not contain any other such lines (e.g. Ca II H & K), we cannot definitively say this is the case. It is still unclear as to why we do not see variability in other strong lines, specifically the Na I doublet (Figure 4).

5.2. Future Work

While we speculate on several independent scenarios, the ground truth could very possibly be some combination of all presented interpretations. The observations presented here alone are insufficient to distinguish

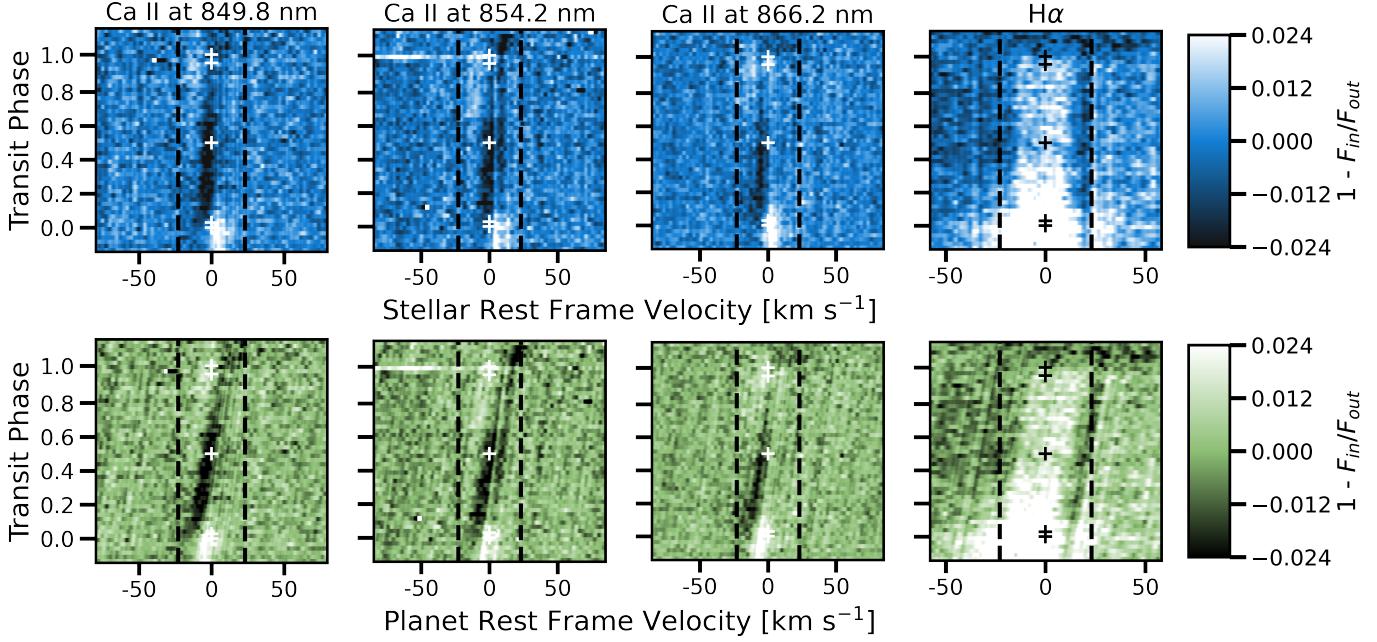


Figure 11. Comparison of transmission signal for both Ca II IRT and H α in the stellar rest frame (top; blue) and V1298 Tau c’s rest frame (bottom; green). Here, white represents absorption and black represents emission, which is not physical in transmission spectra. The excess in Ca II IRT stretches more so from $\pm v \sin i$ (vertical dashed black lines) in the planet’s rest frame velocity, indicating the signal is stellar in origin. Pluses represent the four contact points of the transit and transit midpoint.

between them. Multiple transits observed with both Ca II H & K and Ca II IRT could help answer why the Doppler tomographic signal is only seen in this feature of our observations. If it is the result of star-planet interactions, there will be periodic equivalent width variations corresponding to the period of V1298 Tau c (e.g. HD 179949 b; Shkolnik et al. 2008).

Since H α is known to contaminate transmission spectra (Rackham et al. 2018) via the presence of stellar inhomogeneities and young stars are extreme cases of these surfaces (e.g. Gully-Santiago et al. 2017), a detection of an extended H α atmosphere is very challenging. Multiple observations of V1298 Tau c transiting over different stellar surface phases need to be observed to try and recover this signal. Additional OOT observations may be required to understand the general activity of H α for this active young star. If the 5% variations of H α from our observations are stellar in nature, which we believe they are, then more OOT observations will help build a better spectral template to compare IT observations to.

5.2.1. HeI at 1083.3 nm

Starspot and AR contamination has a much less significant effect ($\approx 0.1\%$) for HeI at 1083.3 nm (Oklopčić & Hirata 2018). Several detections of atmospheres have been made through observations of HeI (e.g. Mansfield

et al. 2018; Spake et al. 2018; Allart et al. 2019; Paragas et al. 2021). Guilluy et al. (2020) conclude that planets around active stars are good candidates for searching for HeI atmospheric signatures and signs of atmospheric escape. This would also be ideal for young planets, not still embedded in debris disks (Hirano et al. 2020), as stellar inhomogeneities are lower in contrast at longer wavelengths (Strassmeier 2009). Additionally, simultaneous observations of V1298 Tau c in H α and HeI will allow us to further understand if the trend in our observations originates from the planet or is dominated by underlying stellar activity.

Ground-based follow-up of V1298 Tau c and other transiting planets in this system using NASA Keck + NIRSPEC could allow for the first detection of an extended young atmosphere in HeI. If the greatest variations we will see are driven by chromospheric lines, then transits of V1298 Tau c in the UV may prove more fruitful. Observations using the Hubble Space Telescope + Wide-Field Camera 3 would allow for a full characterization of planetary atmospheres and the high-energy flux from the young active host star. Future James Webb Space Telescope observations of transits of V1298 Tau c as well as V1298 Tau b and d would offer unprecedented detail of the composition as well as characterizing the extension of the young atmospheres. Observations of all three planets would allow for a detailed comparison

of atmospheric composition and mass loss for planets of different radii in the same harsh stellar environment.

5.2.2. V1298 Tau and *TESS*

Additionally, *TESS* will observe V1298 Tau during its extended mission from September - November 2021. Simultaneous spectroscopy with this guaranteed photometry would allow for a confident removal of stellar activity and first potential detection of an extended H α or He I atmosphere. The resulting *TESS* light curves will show much lower variability amplitudes by virtue of spot contrast that would not be obtainable by ground-based photometry at the *TESS* wavelength coverage. Additionally, *TESS* overlaps in wavelength coverage (600-1000 nm; Ricker et al. 2014) with H α , Ca II IRT, and He I, making it easier to compare the spectral response in simultaneous data sets.

X-ray observations of HD 189733 provided evidence of an intense flaring event 3 ks after the eclipse of HD 189733 b (Wolk et al. 2011). These observations lead to speculation that the flare was the result of magnetic interactions between the star and planet. There is evidence of flares from V1298 Tau in the *K2* light curve (David et al. 2019a) and future *TESS* light curves at a higher cadence may display similar flare trends. Any phase dependence of flares with planet transits could be evidence of star-planet interactions and are unlikely dependent on stellar surface spot coverage (Feinstein et al. 2020b).

6. CONCLUSIONS

We have presented a Doppler tomographic and spectral analysis of V1298 Tau, a 23 Myr K star, during a transit of its innermost known planet, V1298 Tau c.

1. We measure an alignment of V1298 Tau c of $\lambda = 5^\circ \pm 15^\circ$ from the Ca II IRT at 849-866 nm. This is the only spectral feature with an obvious transit signal, hinting that measuring the spin-orbit alignment of young planets may be more feasible at near-infrared/infrared wavelengths.
2. V1298 Tau c adds to a growing list of young planets with low obliquities, including V1298 Tau b (Johnson et al. in prep), indicating these planets undergo smooth migration shortly after they formed. Multi-planet systems tend towards being well-aligned at all ages (Albrecht et al. 2013) and we have demonstrated that the V1298 Tau system is no exception.
3. Variations in Ca II IRT could be an indication of star-planet interactions. More transit observations

of V1298 Tau c in Ca II H & K and Ca II IRT may be able to firmly assess whether this is the case.

4. The excess absorption in H α corresponds well to variations seen in other stars of younger and slightly older ages. The variations can also be replicated solely via the presence of starspots and faculae, leading us to believe stellar activity is the origin of the signal and it is not from an extended young atmosphere.
5. These observations alone are insufficient to characterize any potential extended atmosphere from V1298 Tau c. More transits are required to disentangle stellar from planetary signals. Future observations in the UV, for more chromospheric lines, and NIR/IR, where starspots are at lower contrast, may prove more optimistic for characterizing young planetary systems.
6. There is still much work to be done to understand the nature of the variability seen in our spectra. The data⁵ used in this analysis can be found on the Gemini Archive. In the interests of open science, our analysis code is also publicly available⁶ and can be found via the  icons throughout the paper.

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⁵ Gemini Program GN-2019B-FT-215 on the Gemini Data Archive: <https://archive.gemini.edu/searchform/>

⁶ https://github.com/afeinstein20/doppler_tomography

Council (Canada), Agencia Nacional de Investigación y Desarrollo (Chile), Ministerio de Ciencia, Tecnología e Innovación (Argentina), Ministério da Ciência, Tecnologia, Inovações e Comunicações (Brazil), and Korea Astronomy and Space Science Institute (Republic of Korea). Data were obtained under the program GN-2019B-FT-215.

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Software: numpy ([Van Der Walt et al. 2011](#)), matplotlib ([Hunter et al. 2007](#)), scipy ([Jones et al. 2001–](#)), astropy ([Astropy Collaboration et al. 2013; Price-Whelan et al. 2018](#)), ccdproc ([Craige et al. 2017](#)), OPERA pipeline ([Martíoliet al. 2012; Teeple 2014](#)), starry ([Luger et al. 2019](#)), batman ([Kreidberg 2015](#)), stella ([Feinstein et al. 2020a](#)), BANYAN- Σ ([Gagné et al. 2018](#)), MISTTBORN ([Johnson et al. 2017, 2018](#)), emcee ([Foreman-Mackey et al. 2013](#))

Facility: Gemini-North: GRACES ([Chene et al. 2014](#)), Anglo-Australian Telescope: Veloce-Rosso Spectrograph ([Gilbert et al. 2018](#))

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