PHYS644 Problem Set 6

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Problem 1: Photons in Today's Universe

We are told that the CMB $T_0 = 2.725\,\mathrm{K}$ (Kelvin isn't a degree!). We are asked to calculate the number density of photons in the universe today AND the energy density of photons - $\Omega_{\gamma,0}H^2$.

From lecture notes we know for bosons (photons are bosons)

$$n_\gamma = \frac{g\zeta(3)T^3}{\pi^2} \tag{1}$$

In natural units of $k_b = c = \hbar = 1$, g is the number of internal degrees of freedom per particle in case $g = 2$.

To add back in units, we know that n should have units of 1/volume^3,

$$\mathrm{cm}^{-3} = T^3 \tag{2}$$

but the right hand side has units of temperature. k_bT gets us units of energy, and then we need to get energy to volume, recall that $\hbar c$ has units of energy * length.

$$\boxed{n_\gamma = \frac{g\zeta(3)}{\pi^2}(\frac{k_bT}{\hbar c})^3} \tag{3}$$

Now we can plug in our value of T.

$\boxed{n_\gamma = 411\,\mathrm{cm}^{-3}}$.

Problem 1b: $\Omega_{\gamma,0}h^2$

From lecture notes we know:

the energy density of photons is

$$\rho_{\gamma,0} = 4\sigma T^4 \tag{4}$$

where σ is the stefan-boltzman constant, but in natural units.

and $\rho_{\mathrm{crit},0}$ is given by:

$$\rho_{\mathrm{crit},0} = \frac{3H_0^2}{8\pi G} \tag{5}$$

This as units of kgR^{-3} but should have units of energy density - $\rho = KgR^{-1}t^{-2}$. We are off by c^2.

and $\Omega_{\gamma,0}$ is given by:

$$\Omega_{\gamma,0} = \frac{\rho_{\gamma,0}}{\rho_{\mathrm{crit},0}} \tag{6}$$

So $\Omega_{\gamma,0}h^2$ is given by:

$$\Omega_{\gamma,0}h^2 = \frac{8\pi^3 G}{45\,(100\ \mathrm{km\,s^{-1}\,Mpc^{-1}})^2}\frac{(k_BT_0)^4}{\hbar^3c^3} \tag{7}$$

To add the units back I remember from stat mech that there is a factor of c, and then its just ρ_{crit} which we found and then using the unit form of σ (from wiki).

$\boxed{\Omega_{\gamma,0}h^2 = 2 \times 10^{12}}$

Problem 2: Acceleration Redshift

z_{acc} happens when the universe expansion just starts (postively) accelerating IE when $\ddot{a} = 0$.

The second Friedmann equation becomes

$$\frac{\ddot{a}}{a} = 0 = -\frac{4\pi G}{3}(\rho + 3p) \tag{8}$$

Therefor our condition is when $\rho + 3p = 0$. We are absorbing Λ into its own effect ρ and p. Baryonic matter gives $p = 0$, and dark energy has $p = w\rho_{DE}$.

Now we have

$$0 = \rho_m + \rho_{DE}(1 + 3w) \Rightarrow \rho_m = -\rho_{DE}(1 + 3w) \tag{9}$$

recall that

$$\rho_{\text{crit},0} = \frac{3H_0^2}{8\pi G} \tag{10}$$

$$\Omega_{m,0} = \frac{\rho_{m,0}}{\rho_{\text{crit},0}} \tag{11}$$

$$\Omega_\Lambda = \Omega_{DE,0} = \frac{\rho_{DE,0}}{\rho_{\text{crit},0}} \tag{12}$$

We now insert these into our condition, we convert to values at z recall that $a = 1/(1 + z)$

$$\Omega_{m,0}\,\rho_{\text{crit},0}\,(1 + z)^3 = -(1 + 3w)\,\Omega_\Lambda\,\rho_{\text{crit},0}\,(1 + z)^{3(1+w)}. \tag{13}$$

Now we solve for z — this will be z_{acc}.

$$(1 + z)^{-3w} = -(1 + 3w)\,\frac{\Omega_\Lambda}{\Omega_{m,0}} \tag{14}$$

$$\boxed{z = [-(1 + 3w)\,\frac{\Omega_\Lambda}{\Omega_{m,0}}]^{3w} - 1} \tag{15}$$

We learned from class (and from $a(t) = 1/(1 + z)$ that $z = 0$ is now, and $z = -1$ is $t = \infty$. Aka $z > 0$ if we are looking backwards in time. We see that the first term on the left is always positive when $w < -\frac{1}{3}$.

Problem 4: Distance Measures

Problem 4A:

We are told that $\Omega_{m,0} = 0.3$, $\Omega_{r,0} = 8.5 \times 10^{-5}$, $\Omega_\Lambda = 1 - \Omega_{r,0} - \Omega_{m,0}$

We know that the physical distance R is given by

$$R(t) = a(t)\chi \tag{16}$$

and that $a(t)$ is the scale factor and χ is the commoving distance which is fixed if objects are just drifting due to $a(t)$. χ

Co-moving distance χ is also written as D_c, We know that the co-moving distance is

$$D_c(z) = c\int_0^z H(z)^{-1}\,dz \tag{17}$$

and we can write $H(z)$ as

$$H(z) = H_0\sqrt{\Omega_m(1+z)^3 + \Omega_r(1+z)^4 + \Omega_k(1+z)^2 + \Omega_\Lambda} \tag{18}$$

This is the corrected for z version! In our case $\Omega_k = 0$ because like the Earth[1] - the universe is flat. We now have:

$$D_c(z) = c\int_0^z \frac{1}{H_0\sqrt{\Omega_m(1+z)^3 + \Omega_r(1+z)^4 + \Omega_\Lambda}}\,dz \tag{19}$$

While I could solve this analytically, we are asked to plot so why bother?

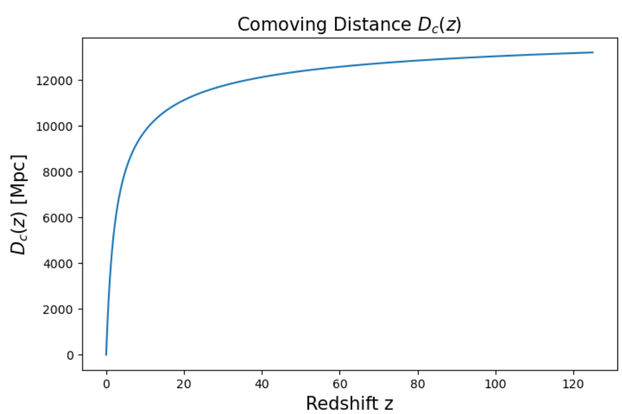

Figure 1: I'm not writing a caption

Here is our plot figure 1! Most of the evolution happens at "low" redshift and then it looks like we start to approach an asymptote as we go on.

Problem 4B:

We do the same thing now until the egde of the observable universe! This is $z = \infty$

Here is our plot figure 2. We can see we run into a numerical issue at high z, but that D_c asymptotically approaches roughly a constant value.

[1]This is a meme.

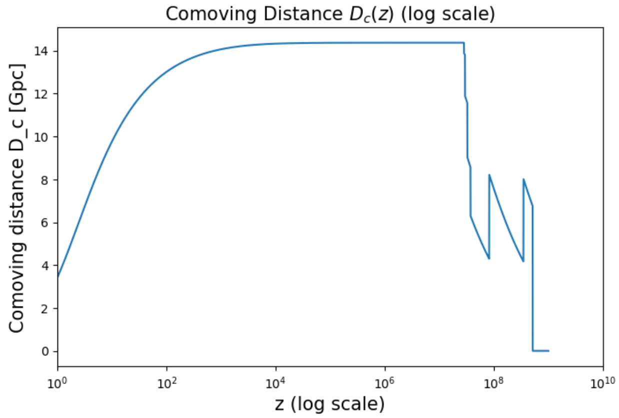

Figure 2: no caption!

Problem 4C

From figure 2 we can read the asymtotically value as $\sim 14\,\mathrm{Gpc}$. So half of this value is $\sim 7\,\mathrm{Gpc}$ — by my eye that is around $z = 10$ ish. (Its actually easier to see on 4A!)

Problem 5: Chemical Potential of Electrons and Positrons

This reminds me of the good times back in advanced statistical mechanics! If only I remembered more from the class :(.

Problem 5A

We are asked to calculate the difference between the number densities of electrons and positrons, $n_{-} - n_{+}$, in the relativistic limit $m_e << T$.

From phase-space density we know:

$$f(P) = \frac{1}{\exp((E-\mu)/T) \pm 1} \tag{20}$$

Where $+$ is for Fermions, and $-$ is for Bosons. We use the $+$ as electrons are Fermions.

The number density is

$$n = \frac{g}{(2\pi)^3} \int f(P) dP^3 \tag{21}$$

These are also given in the lecture notes for lecture 10.

Switching to spherical coordinates and into momentum space, $dp^3 = p^2 dp d\Omega$, Ω is the solid angle! so the integral over it $\int \Omega = 4\pi$ as always.

we can write:

$$n = \frac{g}{(2\pi)^3} 4\pi \int_0^\infty p^2 f(p) dp = \frac{g}{2\pi^2} \int_0^\infty p^2 f(p) dp \tag{22}$$

In our natural units $E = \sqrt{P^2 + m^2}$, and in the relativistic limit this becomes $E \approx |p|$.

Let's define the electron $\mu_e = \mu$, and positrons as $-\mu_e = -\mu$. We can then write:

$$n_{\mp} = \frac{g}{2\pi^2} \int_0^\infty \frac{p^2}{\exp((p \mp \mu)/T) \pm 1} dp \tag{23}$$