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# Work Package Summaries and Reports

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| WORK PACKAGE 1 (WP1):  MANAGEMENT |
| WP1 ACTIVITY TYPE: MGT |
| WP1 DURATION: MONTHS 1 – 36 |
| WP1 LEAD BENEFICIARY: STFC (1) |
| WP1 LEADER: Professor Richard A Harrison |
| WP1 CONTRIBUTORS: n/a |
| WP1 OVERVIEW: The main tasks of WP1 are the initial establishment of the management structure for the HELCATS project, and its support and communication tools, the general operation of this structure, and its termination at the end of the project together with the delivery of the final reports. The work package consists of two tasks, namely, the HELCATS Project Management and the HELCATS website maintenance. The HELCATS consortium consists of 8 European groups (involving 7 EU counties), plus one third-party group from the USA. The project is a complex coordination of activities involving observations and cataloguing, techniques, the development of models and their application, and exercises in validation and assessment. The objective of WP1 is to provide the necessary management structure to implement the HELCATS project effectively, overseeing all administrative matters, assembling and submitting formal reports, overseeing finance auditing, and arranging meetings as appropriate. WP1 will oversee the scheduling and progress of all activities, and the production and management of the HELCATS website. |
| WP1 TASK 1.1: HELCATS PROJECT MANAGEMENT (TASK LEAD: STFC)  Task 1 covers the formal management roles of the HELCATS project. This includes administrative tasks, maintaining project infrastructure, coordinating inputs and reporting as required and handling a range of project issues.  At the onset of the project, the Executive Board established the formal communication with the Steering Committee, setting up monthly teleconferences and the kick off meeting. Regular monthly telecons have been held since, the latest one being September 2015, and minutes of the telecons are lodged on the HELCATS website wiki pages (see below) in the project private section. On rare occasions, monthly telecons have been cancelled due to lack of availability of people or to the coincidence of a HELCATS meeting, when the team is assembled in any case (e.g. the Annual meeting, and the first six monthly meeting).  The combined Second Bi-Annual HELCATS Project meeting and Annual Open HELCATS Workshop were held in tandem in Göttingen, Germany on 18-19 May 2015 and 19-22 May 2015, respectively. The former, project meeting included representatives from all beneficiaries and the EC Project Officer (Sabri Mekaoui) and Reviewer (Brigitte Schmeider). The Annual Report of the project was formally delivered at this meeting and comments received from the Reviewer. The formal response, which was very positive, was received by the Project Coordinator after the meeting. The presentations given at the meeting were placed on the project website, as was the Annual Report.  The Annual Open Workshop was well attended by members of the HELCATS Team plus representatives from many groups from Europe and further afield, including representatives from the key space weather groups. A report on the meeting, including links to the presentations can be found on the website. This was an extremely successful meeting both for showcasing the unique work being done within the HELCATS project but also for debating issues with a wider forum that are relevant to the detection, analysis and tracking of solar transients.  With regard to project deliverables, the project has uplinked the deliverable items to the EC website on the following, to date: D1.1 – HELCATS website launched (report on this action and the website structure were uplinked to the EC website), D1.2 – Minutes of the kick off meeting, D1.3 – Six monthly progress report, D1.4 – Annual Report. Also delivered to the project and uplinked is the report on the feasibility of automatic identification of CMEs (D2.2). D3.1 involves the delivery of time-elongation maps for the cataloguing and, whilst this is complete, it remains for us to uplink a formal statement to the EC website. With the completion of the 18 month meeting (see below), we will be submitting (D1.5) our 18 month report and this will be followed by the 18 month cost statements.  Our 18 month Bi-annual HELCATS Project Meeting was held in Helsinki on 3-4 November 2015. This was run in the same format as the 6 and 12 month Bi-annual Project Meetings with reports on all Work Packages. The 18 month report draft was tabled at the meeting. The meeting was attended by representatives from all beneficiaries, except TCD, plus the EC Project Officer and the Reviewer.  Through this Work Package, the project maintains contact with the EU through the Project Officer and deals with any formal project communications. We coordinate the production of required documentation and oversee meetings with support for logistical arrangements, information; we set the agendas for meetings and have responsibility for overseeing any formal aspects of the project, if required, such as any legal issues or changes to the project structure (e.g. due to staff changes). However, to date, the basic structure of the project and the progress of the Work Packages have been established and delivered as planned.  Finally, we note that the a presentation of the HELCATS project was made to the wider community (Harrison, Davies et al) at the European Space Weather Week, Oostende, Nov. 2015. |
| WP1 TASK 1.2: HELCATS WEBSITE MAINTENANCE (TASK LEAD: STFC)  The HELCATS website (<http://www.helcats-fp7.eu/>) continues to be a major tool for the project, especially with the exchange and posting of information. The homepage provides links to details of the project and a growing list of news items and products of the project workpackages as they develop (such as links to the CME catalogues). This is all publically available. The website has a wiki facility, both public and private (to the project). The private wiki pages contain the Steering Committee minutes, meeting reports, project newsletters, meeting information and documentation. Note that the Dissemination workpackage, WP8, has formal responsibility for placing reports, catalogues, publications etc. on the website. |
| WP1 SUMMARY/NEXT STEPS:  WP1 continues throughout the lifetime of the HELCATS project, overseeing all formal communications, meetings, and the project coordination. Thus, the Steering Committee telecons will continue on a regular basis, the Bi-Annual and Annual meetings will proceed as planned and scheduled throughout the project, and reports will be produced as required. Following the Helsinki meeting, the emphasis is on completing the cost statements of the beneficiaries, submitting the project financial statement, and submitting the final version of the 18 month report. Preparation for the second Annual Open Workshop and 24 month Bi-annual Project Meeting, both to be arranged by the Toulouse group (UPS) in May 2016, will start shortly. |

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| WORK PACKAGE 2 (WP2):  PRODUCING A DEFINITIVE CATALOGUE OF CMES IMAGED BY STEREO/HI |
| WP2 ACTIVITY TYPE: RTD |
| WP2 DURATION: MONTHS 1 – 36 |
| WP2 LEAD BENEFICIARY: STFC (1) |
| WP2 LEADER: Dr Jackie Davies |
| WP2 CONTRIBUTORS: UGOE (4); ROB (5) |
| WP2 OVERVIEW: WP2 involves the production of a catalogue of CMEs in the heliosphere through visual inspection of white-light imagery from the HI Instruments on NASA’s twin-spacecraft STEREO mission. Autonomous cataloguing of STEREO/HI CMEs, based on use of the long-established CACTus software package, is also being investigated within this WP, and since it has proven to be successful, these catalogues are currently being compared with each other, as well as with other, pre-existing, coronal CME catalogues. WP2 is divided into four tasks. Task 2.1 (led by STFC) comprises the manual CME cataloguing of STEREO/HI CMEs, while the automatic cataloguing is performed under Task 2.2 (led by ROB); Task 2.3 (led by UGOE) covers the comparison of the resultant CME catalogues, both with each other and pre-existing coronal CME catalogues. Task 2.4 (led by STFC) covers the HELCATS scientific management. A summary of the progress in each of these tasks, over the first 18 months of the HELCATS project, is given below. |
| WP2 TASK 2.1: MANUAL CATALOGUING OF STEREO/HI CMES (TASK LEAD: STFC)  WP2 is core to the HELCATS project, in that it provides the fundamental manually-generated catalogue of coronal mass ejections (CME) in the heliosphere, which will become the benchmark for future studies of interplanetary CMEs. It provides the foundation on which the subsequent CME-oriented WPs, e.g. WP3, are based. This catalogue is generated by the visual inspection of heliospheric images from HI-1 observations, from the start of the mission science phase in April 2007. Through extensive discussion within the HELCATS consortium, the format of the manual CME catalogue (Task 2.1 of WP2) contains the following fields:  *•* A unique heliospheric CME identifier.  *•* The time of first observation of the CME in the HI-1 field of view (UTC).  *•* The observing spacecraft (A or B).  *•* The northernmost position angle extent of the CME (degrees).  *•* The southernmost position angle extent of the CME (degrees).  *•* A quality flag indicating whether the CME observation is considered to be *poor*, *fair* or *good*.  The unique heliospheric CME identifier is a string containing both the date on which the CME is first observed in HI-1 and the observing spacecraft, with an additional two-digit number to differentiate between multiple CMEs occurring on the same day. So, for example, a CME observed by STEREO-A on 31 Dec. 2007 will have the identifier “HCME A 20071231 01”. CMEs that exceed the position angle range of the field of view are indicated with a *greater than “>”* or *less than “<”* symbol in the appropriate field. It should be noted that in recognition of the numerous blob-like transients observed close to the Sun, for the unambiguous identification of CMEs, we demand that CMEs have position angle widths of 20o or more. The quality flag has been introduced to account for the ambiguity that results from using human observers to identify events. This field is used as a means to quantify the confidence of the observer, based on the following criteria. A *poor* event is any object (spanning at least 20*◦* in position angle, as noted), but which poorly resembles a CME. A *fair* event is one that resembles a CME, though not all observers may be convinced that this is the case (due to some limitation in the event observation such as a faint, disjointed or otherwise irregular CME structure). A *good* event is one that is unquestionably a CME.  At the time of the 18 month report this cataloguing activity has been completed for the full nominal STEREO mission science phase up to the operational interruption for superior conjunction, i.e. the catalogue includes the years 2007*−*2014, providing the observational properties of over 1000 events.  A paper (Harrison et al.) is in preparation on this unique and comprehensive catalogue of HI-observed CMEs. It outlines the statistical properties, making clear associations with related studies (such as coronal CME catalogues) and will provide the first stage of what we regard to be a benchmark for understanding CMEs in the heliosphere. This will be complemented by a second paper covering the kinematic properties (see WP3).    *Figure 2.1: The WP2 HI catalogue facility on the HELCATS website*  Figure 2.1 shows a snapshot of the HI catalogue facility that resides on the HELCATS website (www.helcats-fp7.eu), under the ‘products’ tab. The catalogue is updated as events are identified and inspected. The facility includes a capability to select date ranges and specific position angle location and width. |
| WP2 TASK 2.2: AUTOMATIC CATALOGUING OF STEREO/HI CMES (TASK LEAD: ROB)  In addition to manual CME identification, task 2.2 comprised an assessment of the potential of autonomous detection of CMEs in the STEREO/HI imagery using ROB’s CACTus software application (http://sidc.oma.be/cactus/). CACTus, a long-established algorithm for automatic CME detection and tracking, has, in the past, been applied exclusively to coronagraph images, from SOHO/LASCO and STEREO/COR. Indeed, we can report that CACTus has been successfully adapted to work with STEREO-HI images and we note that this was concluded in the Annual Report (which was uplinked to the ECAS website as action D1.4) and in a dedicated report (uplinked as action D2.2). We plan to publish the results of this activity (Rodriguez et al.) to make the findings available to the wider community.  CACTus applies the Hough transform to distance–time maps, which are then converted to a time vs angle map which contains all information about the CME (see Robbrecht and Berghmans, 2004, Astron. Astrophys. 425, 1097 for more details). This information includes the starting time of the CME, its position angle, angular width and speed. The catalogue now covers the full STEREO mission; a sample is shown in the following figure:  *Figure 2.2: A sample of the CACTus catalogue for the HI data.*  In the left-hand column each CME is marked with a number and the following columns describe the CME parameters, measured automatically. On the right-hand side, a graphic version of the detections is shown. The user can click on each CME to obtain further details of the event, including movies.  In order to validate the catalogue, we have been making comparisons with the manual catalogue of task 2.1. We have identified the events that are common to both catalogues, using an algorithm that compares CME start times, propagation direction and angular widths. We are currently comparing their characteristics, in order to identify differences between manual and automatic classification. Furthermore, we have started to do an eye inspection of all the events and a manual matching between both catalogues. The outcome of this comprehensive comparison and analysis will help to fine-tune the parameters in the automatic detection algorithm, and also provide physical insights into the properties of CMEs detected in HI data. |
| WP2 TASK 2.3: COMPARISON OF CME CATALOGUES (TASK LEAD: UGOE)  UGOE established a STEREO/SECCHI/COR2 coronal CME database under the framework of the DLR project STEREO/Corona, and the EU FP7 projects SOTERIA and AFFECTS. This database, available online at http://www.affects-fp7.eu/cme-database/, has been updated, as planned (as part of WP3), to the end of 2011 (see below), to enable a comparison of the SOTERIA/AFFECTS database with the HELCATS catalogues. As discussed in the annual report, 241 of the CMEs (referred to as the STEREO COR2 “Best-of-list”) were selected for further analyses, i.e. the data were such that a Graduated Cylindrical Shell (GCS) model could be applied. One such event from the UGOE database is illustrated in Fig. 2.1. As part of WP2, the COR2 list is currently being compared with the STEREO/SECCHI/HI1A and B lists established under WP2 task 2.1. Preliminary comparison of the HI lists with the COR 2 “Best-of” list has yielded a set of 109 CME events that could be uniquely associated in both the HI1A or B and COR2 A or B fields of view without duplicate events. The comparison lists will be made available at the HELCATS website.  Bildschirmfoto 2014-11-03 um 16.04.54.png  *Figure 2.1: Sample CME observed in the STEREO-B/SECCHI/COR2 field of view on March 12, 2012.* |
| WP2 TASK 2.4: SCIENTIFIC MANAGEMENT (TASK LEAD: STFC)  Whereas WP1 covers the administrative management of the HELCATS project, WP2 includes an element for the scientific and technical management. Thus, whereas WP1 includes aspects such as the oversight of the meeting and teleconference structure, the WP2 effort coordinates the scientific and technical discussion between the WP leaders and the wider team. Throughout the project activities to date, this has focused on the teleconference discussions, the kick-off, six monthly and annual meetings, and the technical details of the report writing. It also covers the regular technical communication with the WP teams on a range of issues as the project progresses. Thus, this WP runs in parallel with the logistical management of the project and continues throughout the HELCATS project. It is coordinated by both the Scientific Manager (Jackie Davies) and the newly appointed Technical Manager (Jason Byrne). |
| WP2 SUMMARY/NEXT STEPS:  The activities of WP2 are well underway, with active cataloguing of HI events, both by manual inspection and by automated processes.   * The HI heliospheric CME (manual) catalogue now covers the period 2007 to 2014 and is a unique resource forming the backbone of the HELCATS project. * The CACTus automated CME identification method, which has proven to be valid for HI heliospheric CME detection, has been used to produce an event list that is being compared to the manual catalogue, to assess performance. * Comparisons of the HI manual heliospheric CME catalogue to previously compiled coronal CME event lists are under way.   Many of the WP2 activities continue throughout the HELCATS project; the final catalogue deliverables are actually required by month 36 (deliverables D2.1, D2.5). This is a natural consequence of the fact that the STEREO project is still in operation and the event lists will grow throughout the mission. However, there are three WP2 deliverables that have either been finalised so far, or are up to date, pending further inputs from the extended STEREO mission: Deliverable D2.1 which is the observational cataloguing of heliospheric CMEs in the HI images, is now complete for 2007 to late 2014; Deliverable D2.2 which is the report on the feasibility of automated CME identification and this has indeed been demonstrated and delivered; and Deliverable D2.3 which is the report on the comparison of the manual and automated catalogues is now underway.  Publications are in preparation on the basic statistical analysis of the heliospheric CME catalogue (Harrison et al.), as a new benchmark on heliospheric CME observation, and on the successful CACTus results in the application to HI data (Rodriguez et al.).  In summary, the cataloguing activities of WP2 are progressing as planned. |

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| WORK PACKAGE 3 (WP3):  Deriving/cataloguing the kinematic properties of STEREO/HI CMEs based on geometrical and forward modelling |
| WP3 ACTIVITY TYPE: RTD |
| WP3 DURATION: MONTHS 1 – 36 |
| WP3 LEAD BENEFICIARY: UGOE (4) |
| WP3 LEADER: Dr. Volker Bothmer |
| WP3 CONTRIBUTORS: UGOE (4); TCD (9); STFC (1); UNIGRAZ (2) |
| WP3 OVERVIEW: The key objectives of WP3 are deriving and cataloguing the kinematic properties of the STEREO/HI CMEs identified and catalogued in WP2, based on geometrical and forward modelling. For these purposes the recently established geometrical, forward (and prototype) inverse CME-modelling methods are applied. The CME kinematic parameters are added to the WP2 catalogue, including back- and forward-projections to ‘predict’ CME launch and arrivals at various solar system locations. The WP includes comparisons between the parameters yielded by the different models. |
| WP3 TASK 3.1: Geometrical modelling of STEREO/HI CMEs (TASK LEAD: STFC)  The STFC contribution to WP3 is principally through task 3.1, the geometrical modelling of the STEREO/HI heliospheric CMEs. This task involves the derivation of the kinematic properties of those CMEs visually identified in WP2.1. This is achieved by manually identifying the track made by each CME in a time-elongation map (J-map) and applying the assumptions about its geometry and dynamics summarized in Davies et al. (2012, Astrophys. J., 750, 23). This makes use of three established geometrical approaches, assuming that the CME can be modelled (i) as a point-like feature propagating outwards, or Fixed-Phi (FP) method, (ii) as an expanding circle whose diameter is defined by the CME apex and Sun-centre, the Harmonic Mean (HM) method, and (iii) as a self similarly expanding circle of 30° half-width, propagating outward between two fixed position angles, the Self Similar Expansion (SSE) method.  The CMEs are assumed to travel at a constant speed in a fixed direction away from the Sun. The three different fitting methods are applied to each CME (see Davies et al., for details). For a given CME, the path of its (apparent) leading edge through a J-map is manually tracked at a position angle close to its apex and each of the three fitting procedures is applied to estimate its 3D speed and propagation direction. These values are, in turn, used to derive launch times for each event, which are then applicable to WP4.1, and to generate arrival times at various locations in the heliosphere for in-situ comparison, to be made in WP4.2. Thus, the definitive product of this task is a unique and comprehensive list of heliospheric CMEs with fitted kinematic parameters, notably including longitudinal direction.  To date, events occurring between the beginning of the STEREO science phase, in April 2007, and the beginning of reduced science operations in the lead up to superior conjunction of STEREO, from August 2014, for both STEREO-A and -B have been analysed in this way. Events identified as poor in WP2.1 are excluded, as are those which are directed at position angles far from the ecliptic, due to the limited number of HI frames in which they appear. A number of small CMEs occur, which quickly become subsumed by subsequent, larger events; these are also excluded. This process has been completed for a total of 1210 events; 635 for STEREO-A and 575 for STEREO-B. A comparison between the speeds and propagation angles resulting from each of the three fitting methods is shown in Fig. 3.1, with a screenshot of the catalogue output shown in Fig. 3.2.    *Figure 3.1: A comparison between the speeds and propagation directions obtained from each of the three fitting methods. The scatter plots in the left column show the difference in velocity between the FP and SSE models (top), FP and HM models (middle) and SSE and HM models. The right hand column shows the differences between the propagation angles,* φ*, for the same methods. The colour of each data point represents the maximum elongation angle to which the CME was tracked. Filled and open dots correspond to CMEs observed by STEREO-A and B, respectively.*  A paper is being prepared on this unique and comprehensive catalogue of geometrically-modelled CMEs observed by HI that will outline the statistics of their kinematics during the nominal STEREO mission science phase of 2007-2014. This builds on the publication being prepared for WP2 and will provide a fundamental analysis of CME kinematic properties in the heliosphere as a reference for the wider community.    *Figure 3.2: Example of the CME catalogue available on the HELCATS website. The kinematic properties (from SSE fitting) of the first ten CMEs are displayed.*  At UNIGRAZ, the routines for calculating planetary and spacecraft arrivals of CMEs observed by HI have been further automated and enhanced. Based on the fitting results (FPF, HMF, SSEF) of the launch time, speed and direction, they produce the arrival times and speeds at each spacecraft location consistent with different CME geometries (see examples plotted in Figs. 3.3 and 3.4). Now that the first HI catalogue with the fit results has been established by RAL (STFC), a few adaptations to our codes concerning the data formats are necessary and are currently underway. The final results of CME impacts will be derived for Earth, Mars, Venus, Mercury, MESSENGER, Saturn, STEREO-A, STEREO-B and Ulysses and will be used in WP4 for establishing the linked catalogue of CME events observed from the Sun to in-situ observing spacecraft.    *Figure 3.3: Distributions of speeds of the modelled CMEs, separated by year, spacecraft and fitting technique (Fixed Phi (FPF), Self-Similar Expansion (SSEF), and Harmonic Mean (HMF) fits).*  This model also allows one to estimate whether, and to what degree, each CME will ‘hit’ any in-situ location, and this is provided as a parameter in the catalogue, along with speeds and arrival times. A user friendly package in the IDL programming language has been created to read in the geometrically modelled CME catalogue files and output the predicted arrival time files (Fig. 3.5) for any of the in situ locations. The initial catalogue contains over 400 predicted arrivals at Earth, and over 100 for Mercury-, Venus-, Mars- and Saturn- directed CMEs. The arrival times will be compared to other model results and in situ data in WP4.    *Figure 3.4: SSEF ICME speed vs. Transit time to different planetary locations.*    *Figure 3.5: Sample predicted arrival time file.* |
| WP3 TASK 3.2: Forward modelling of STEREO/HI CMEs (TASK LEAD: UGOE)  The HI-COR2 event list established in WP2 (see above) contains 109 events. Out of the total number of 109 events, 96 have been modelled with the Graduated Cylindrical Shell (GCS) method, which assumes a flux-tube style geometry, which is then used to match to the observed intensity (density) distribution of the HI images. These results are available online as STEREO/SECCHI/COR2 database on the AFFECTS homepage at http://www.affects-fp7.eu/cme-database/. This is work that combines the output of the AFFECTS and SOTERIA projects with a comparison as part of the HELCATS project, in effect, comparing coronal and heliospheric CME events. The content of the AFFECTS/SOTERIA database has been provided to RAL (STFC) for inclusion as a standardised list on the HELCATS website.      *Figure 3.6: Sample observed mass and velocity determinations from a GCS modelled CME observed on 19 April 2010. Then initial mass curves (~4-5 Rs) reflect the emergence of mass from below the inner edge of the FOV. Forward Modelling fits are shown superimposed on the STEREO-A and B COR images (top).*  In order to improve the accuracy of CME speed and mass determinations, the GCS modelling procedure applied to the events of the COR2 CME database has been extended by making additional use of SOHO LASCO C2/C3 observations. The extended GCS-method will now be applied for the determination of the CME directional propagation velocities and tentative source region positions and also to derive their total masses. The data analysis methods were developed in collaboration with the STEREO/SECCHI PI institution at the Naval Research Laboratory (A. Vourlidas, R. Howard, A Thernisien, N. Savani). The first version of this so called HELCATS “KINematic CATalogue” is now available online at the AFFECTS-FP7 homepage (<http://www.affects-fp7.eu/helcats-database/database.php>), including explanations of the parameters. The database is ready for implementation on the HELCATS website. A sample CME speed and mass determination is shown in Fig. 3.6.  First results of the statistical analysis of COR2-HI comparisons were presented at the first biannual HELCATS workshop held in May 2015 at Göttingen University (http://www.affects-fp7.eu/helcats-meeting/).  For proper comparison of the derived speeds of different CME events, the initial speeds provided in the database were derived for a constant CME apex distance from the Sun. The individual height-time-profiles of the modelled KINCAT events will be established as the next step and can then be used for further research on CME kinematics, e.g. to study the influence of drag forces, and for comparisons with other data sets, e.g. onset of radio waves or energetic particle fluxes. In order to enhance the precision of the calculated total CME masses, methods to improve the accuracy in subtracting non CME background brightness caused by ambient coronal structures and objects like comets, stars or planets are currently under investigation. |
| WP3 TASK 3.3: Inverse modelling of STEREO/HI CMEs (TASK LEAD: UGOE)  UGOE has established an “Inverse Modelling List“; a sample is shown in Fig. 3.7 for illustration. The list is being used by TCD to inspect related solar phenomena to provide a “Low Coronal Event Catalogue“ available at <http://data.rosseobservatory.ie/HELCATS/hi_catalogue/> (Fig.3.8). It contains information on the back-projected CME from HI to COR2 and properties of the source region.    *Figure 3.7: Sample of the UGOE Inverse Modelling List.*  The results of WP3 enabled the identification of the source regions and coronal activity associated with interplanetary CMEs in the HELCATS catalogue. This allows for the analysis of low coronal events including CMEs, EUV waves and type II radio bursts. An essential aspect of analysing type II radio bursts is determine the underlying shock speed and starting height using coronal density models (this is being done as part of WP3). To do aid this analysis, TCD produced density and Alfvén speed maps in a number of events in the HELCATS catalogue. These maps are produced from SDO AIA and LASCO C2 observations so are primarily used for limb events as seen from Earth. These Alfvèn and density maps are available at http://data.rosseobservatory.ie/HELCATS/ (Figure 3.9).  ../../../Desktop/Screen%20Shot%202015-10-04%20at%2023.09.33.png  Figure 3.8: Catalogue of 1673 events available at http://data.rosseobservatory.ie/HELCATS/hi\_catalogue/  However, there is also the possibility of producing such density maps from STEREO COR1 observations, allowing for the analysis of limb events as seen from either STEREO A or B. Furthermore, COR1 produces polarised brightness observations of the corona routinely, allowing for an hourly estimate of coronal density in the COR1 field of view (as opposed to just once per day with LASCO C2).  ../../Seminar_talk_lesia/Presentation1/Slide1.png  Figure 3.9: Example of Density and Alfven speed maps available at http://data.rosseobservatory.ie/HELCATS/.  This is particularly pertinent for analysing type II radio bursts for Earth directed CMEs. In order to carry out this, the density inversion procedures in IDL which currently operate only on LASCO C2 data will need to be re-written for the format of COR1 data. This will benefit type II and shock analysis for events observed by the STEREO spacecraft and avoid having to use a density model of the corona, which can often lead to spurious results. The production of density inversion procedures for the COR1 data sets will be done over the next 18 months. |
| WP3 TASK 3.4: Comparison of modelling results (TASK LEAD: UGOE)  The results of the COR2 CME speed determinations derived under WP3 task 3.2 are being compared with the speeds obtained from HI observations derived from geometrical modelling under WP3 task 3.1. A sample comparison is shown in Figure 3.8. Further comparisons will be established for the full set of events of the established KINCAT database as this WP progresses. This will also include comparisons of the determined source region positions.  HI-speeds:  456 (FPF)  470  481  *Figure 3.9: Sample comparison of modelling results* |
| WP3 SUMMARY/NEXT STEPS:  WP3 has seen considerable progress.   * The application of geometrical fits to the events catalogued in WP2 are providing the first comprehensive kinematic and geometrical catalogues of heliospheric CMEs, which is also allowing a thorough coomparison of the geometrical models being exploited. * The kinematic analyses are now being used to project arrival times at key locations in the Solar System (Earth, other planets and spacecraft). * Forward Modelling techniques used in previous FP7 programmes are being extended and adapted for use with the HI data, and these are now beginning to allow comparisons of model results. * Inverse Modelling is being used to project back to solar source regions, allowing the compilation and analysis of associated solar activity and structures.   These activities are progressing as planned and promise significant advances and results that will provide valuable resources for the wider community in the near-future. Various publications are planned including a statistical analysis of the HI heliospheric CME kinematic results. |

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| WORK PACKAGE 4 (WP4):  Verifying the kinematic properties of STEREO/HI CMEs against in-situ CME observations and coronal sources |
| WP4 ACTIVITY TYPE: RTD |
| WP4 DURATION: MONTHS 10 – 36 |
| WP4 LEAD BENEFICIARY: UNIGRAZ (2) |
| WP4 LEADER: Dr Christian Möstl |
| WP4 CONTRIBUTORS: UNIGRAZ (2), UPS (3), UGOE (4), ROB (5), IMPERIAL (6), UH (7) |
| WP4 OVERVIEW: The primary goal of WP4 is to provide researchers with a catalogue that links CMEs observed in the low corona, in coronagraph and heliospheric imager fields of view, with their *in situ* counterparts. To this end, individual lists of CMEs and their parameters in the low corona (EUV), outer corona (coronagraph observations) and in the heliosphere (heliospheric imaging (HI)), and in situ data, are underway. The activities relating to the solar, coronal and heliospheric cataloguing are detailed above. Most of these data and lists are now available, as summarized in this report. Windows in time and space for which associations between different phenomena in different datasets are being established now, such that the lists can be linked – forming the linked catalogue “LINKCAT“. A first release of LINKCAT on the HELCATS/products website is scheduled for month 24 of the HELCATS project (April 2016). LINKCAT will provide the basis for many investigations on CME physics and kinematics, of which the first examples are shown here. A summary of the progress in each of the tasks is given below. |
| WP4 TASK 4.1: Comparing to coronal sources (TASK LEAD: UGOE)  For establishing the LINKCAT, the coronal sources of a list of CMEs that will be established by UNIGRAZ/UH will be compiled by UGOE. For this, the solar launch time from the heliospheric CME catalogue from WP3 and the CME direction form proxies for temporal and spatial windows to look for coronal signatures of the CME. An example for this process is shown in Figure 4.1. A CME on 14 April 2013 at Earth is unambiguously connected to a CME in the heliospheric (HI) CME catalogue (WP3) (panel 1c) that left the Sun on 11 April 2013, associated with a flare near disk center. Details on the validity of spatial and temporal windows for defining “unambiguous“ connections will be worked out iteratively once the first event list by UNIGRAZ/UH is made available to UGOE.  UGOE will use the Low Coronal Signature catalogue (LOWCAT, see below) and KINCAT lists for CME events that happened until end of 2011, and the Heliophysics Event Knowledgebase for events after 2011. The databases mentioned can be found as follows:  UGOE LOWCAT (see “source region results“): http://www.affects-fp7.eu/cme-database/database.php  UGOE KINCAT: http://www.affects-fp7.eu/helcats-database/database.php  Heliophysics Event Knowledgebase: https://www.lmsal.com/isolsearch.    FIGURE 4.1: Linking CME observations from the Sun to 1 AU - LINKCAT (C. Möstl, P. Boakes, UNIGRAZ; P. Vemareddy, IIA). (a) SDO/AIA image of the source location of a flare on April 11 2013. (b) STEREO-B COR2 observation of the CME, which was Earth directed. (c) Screenshot of a visualization of the CME fronts (blue = from HI on Behind, red = from Ahead) at 2 timesteps, derived from the preliminary WP3 catalogue, with CME constant directions and speeds taken from the SSEF30 results. The two fronts indicate the CME as it propagates away from the Sun towards the Earth, with the two fronts at slightly different directions as seen by HI on each spacecraft. (d) In situ magnetic field observation of the ICME near Earth, arriving on April 14, consistent with the results of HI modeling. (e) Visualization of the CME MFR configuration at 1 AU from an application of Grad-Shafranov reconstruction (e.g. Möstl et al. 2009, Solar Phys. 256, 427). |
| WP4 TASK 4.2: Comparing to in-situ measurements (TASK LEAD: UH)  The goal of task 4.2 is to construct a comprehensive catalogue of heliospheric or interplanetary CMEs (ICME) measured *in situ* and to list their key parameters and the results of modelling based on *in situ* data. Since the annual report, the in situ data (DATACAT) and ICME catalogues (ICMECAT) have been revised. As stated in the work plan, all data shall be made available until end of 2015. However, the following datasets will be updated during this WP until all data are available, which is expected to extend at least into 2016. Before this, the catalogues will be internally released between all project partners, and then publicly released once completed.  Summary of currently available in situ data and lists relevant to WP4  DATACAT: the catalogue of in situ data contains plasma and magnetic fields at various spacecraft, starting in January 2007. We summarize what is currently available for each spacecraft as stated in the work plan, noting which HELCATS group is associated with the data.  At UNIGRAZ (C. Möstl, P. Boakes):  For Venus Express (VEX), the available data of magnetic fields (VSO coordinates) covers 1 January 2007 to the end of the mission on 25 November 2014, thus it is considered as complete. Two datasets are available, one with all data and one with an exclusion of intervals that VEX spent inside the Venus bow shock, based on a general model that does not include bow shock modifications during CME impacts. Both datasets have a time resolution of 1-minute and the IDL.sav file for each version is 263 MB in size.  For MESSENGER at Mercury, with magnetic fields in MSO coordinates, there are also 2 datasets available, one with and the other without the Mercury magnetosphere, covering Mercury flybys in January 2008, October 2008, October 2009, and including continuous data after Mercury orbit insertion on 18 March 2011, currently up to 18 March 2014. The data have 1-minute time resolution, and each IDL.sav file so far is 101 MB in size. The mission ended on 30 April 2015, thus about 1 year of data at Mercury is still left for processing, and about 4 years during the cruise phase (in RTN coordinates).  Ulysses: Plasma and magnetic field data at 1-hour resolution have been processed for 1 January 2007 - July 2009, when the mission ended, thus our dataset is complete. Useful for HELCATS is the last pass of Ulysses through the ecliptic in August 2007.  At UH (A. Isavnin, E. K. J. Kilpua):  The following data all have magnetic fields in HEEQ coordinates and plasma data available (speed, density, temperature):  STEREO-A data are available from 1 January 2007 to 31 August 2014. Due to the superior conjunction of STEREO with the Sun, there is a complete data gap from 20 March to 9 July 2015. Thus the data need to be updated until the end of 2015 when available.  STEREO-B data are available from 1 January 2007 until 31 August 2014, which is all data but one month; loss of contact with the spacecraft was on 27 September 2014. Thus the data are almost complete. However, NASA is still attempting to regain contact with STEREO-B; potentially more data could be available in the future.  Wind data from the L1 point in the near Earth solar wind are available from 1 January 2007 to 31 December 2014. These data will also be updated until the end of 2015.  ICMECAT: This will include clear ICME events, which satisfies the two criteria of elevated total field strength compared the background wind and the presence of a magnetic obstacle, which contains a smooth magnetic field, either rotating (which defines a magnetic flux rope, MFR) or with constant direction, or a complex ejecta. Due to the different data availability of VEX and MESSENGER, which deliver magnetic fields, whereas Wind and STEREO data contains magnetic field and plasma parameters, the criteria for the lists at each spacecraft are slightly different. How much this introduces a bias will be assessed with the studies on the HI to in situ connections in WP 4.3, and the ICMECAT will be adapted if necessary (e.g. clean the STEREO ICME lists for events that do not contain magnetic obstacles).  The lists from various sources for each spacecraft have been further brought into standardised time formats as described extensively in the annual report. A paper on the VEX/MESSENGER Catalogue by S. Good et al. (2015) has been submitted to Solar Physics. The following table provides an overview of the event numbers and time-ranges we have so far.  However, we first bullet the specific updates since the first annual report:   * the new list by Winslow et al. (2015, JGR (Space Phys) 120, 6101) has been added which consists of 61 ICMEs observed at Mercury after orbit insertion, with 10 events overlapping between the Winslow and Good lists, thus 84 unique ICMEs observed by MESSENGER remain. * a shock list was provided by V. Krupar (IMPERIAL) for Wind. From Jan 2007 to Jul 2015, the list contains shock times, shock angles and before / after magnetic field and solar wind speed for 383 events. * another shock database also available at UH http://ipshocks.fi. Both are very useful for comparison to WP3 results if no ICME is present. * A collaboration has been started by UNIGRAZ with the University of New Hampshire (R. Winslow, N. Lugaz, C. J. Farrugia) to identify spacecraft line-ups, i.e. ICMEs that have been subsequently detected at the different spacecraft, relevant for understanding the interplanetary CME evolution.  |  |  |  |  |  | | --- | --- | --- | --- | --- | | Spacecraft | Provider | Number of ICMEs | Start time | End time | | Wind | T. Nieves-Chinchilla | 135 | Jan 2007 | Dec 2013 | | STEREO-A | L. Jian | 167 | Jan 2007 | Jun 2014 | | STEREO-B | L. Jian | 133 | Jan 2007 | Dec 2013 | | VEX | S. Good | 81 | Jan 2007 | Dec 2013 | | MESSENGER | S. Good  R. Winslow | 33  61 | Jan 2007  May 2011 | May 2012  Sep 2014 |   In summary, 435 ICMEs were observed close to 1 AU (Wind, STEREO-A, STEREO-B), plus 165 events in the inner heliosphere by VEX and MESSENGER. This number contains only events after January 2007 so they overlap with the HI catalogue from WP2/3. In total, 600 events are currently in the ICMECAT. Once the full data are available until end of 2015, this number can be expected to rise to approximately 650 events. We note that IMPERIAL (V. Krupar, J.P. Eastwood) is also working on connecttions between CMEs in radio and in situ data, which provides a complementary study, in WP7, with WP4. |
| WP4 TASK 4.3 Assessing the validity of the HI modelling (TASK LEAD: UNIGRAZ)  In this task we validate the modelling methods used in WP3 to extract CME parameters from HI with multipoint *in situ* data. Since the annual report, progress has been made on CME case studies using the preliminary HICAT and ICMECATs:  Source position vs. CME propagation direction:  The paper Möstl et al. (2015, Nature Comm.) has been published about a major false alarm for a CME on 7 January 2014. A synthesis of data from 7 spacecraft (as used in this WP) has shown that a CME may erupt in a direction that is about 40° different in heliospheric longitude compared to the position of the source region. A follow up paper by Mays et al. (2015, ApJ, in press) has further confirmed this with ensemble numerical simulations, and showed that the 3D orientation of the CME was also favourable in missing the Earth, even though the CME source on the Sun was perfectly aimed towards us.  Multipoint studies of ICMEs with spacecraft line-ups:  Figure 4.2 shows an example of an ICME detection that is contained in ICMECAT, taken from Good et al. (2015, ApJ). A triple line-up of a CME flux rope observation has been studied, from MESSENGER to VEX to STEREO-B, in November 2011. Few such radial observations of MFRs have been previously reported, and they provide direct evidence on the evolution of CMEs from the Sun to Earth. The results confirmed previous results on CME flux rope expansion, but showed that the magnetic flux was conserved between 0.44 and 1.09 AU, suggesting that the rope underwent no significant erosion through magnetic reconnection between MESSENGER and STEREO-B. Derived from force-free fits, the rope axis rotated by 30° between the spacecraft to lie closer to the solar equatorial plane at STEREO-B. Such a rotation, if it is a common feature of coronal mass ejection propagation, would have important implications for space weather forecasting.  Macintosh HD:Users:chris:Desktop:figure_aus_good_2015.JPG  Figure 4.2: Example for an ICME in ICMECAT, observed by Venus Express, taken from Good et al. (2015, IMPERIAL). From top to bottom: Magnetic field magnitude, field components, inclination and azimuthal angle, in RTN coordinates. VEX observed an interplanetary shock, entered the induced magnetosphere of Venus as visible by a sharp peak in the total magnetic field, entered the solar wind again and observed the magnetic flux rope part of the ICME (delimited by two dashed vertical lines). The solid lines show model results of a force-free magnetic field configuration fit. The same event was studied as part of a multipoint lineup that implies a significant rotation of the CME in the heliosphere.  On-going studies:  Establishing the links from the observations of CMEs unambiguously from Sun to Earth is a largely unsolved problem in heliophysics. However, it is crucial e.g. for case studies concerning MFR production on the Sun, which may make possible to give accurate predictions of the internal magnetic field, and hence geoeffectiveness, of CMEs. This it is extremely important for space weather forecasting because CME MFR orientations can affect dramatically their ability to drive geomagnetic disturbances. With the combined catalogues from HELCATS, this problem can be solved.  In-situ magnetic structures compared to magnetic field of their photospheric source regions:  An example of how we tackle this problem is shown in Figure 4.1. A flare and CME erupted on 11 April 2013 and impacted the Earth on 14 April 2013. The movies we have produced with the preliminary catalogue (screenshots in panel c, contained in the previous annual report) allow a straightforward connection of the solar to the in situ data. This event will be further studied with P. Vemareddy (Indian Institute of Astrophysics). The event features a very clear example of a rotation of the EUV wave, and its flux rope can be modelled with reconstruction methods. (CME id: HCME\_A\_\_20130411\_01 and HCME\_B\_\_20130411\_01)  Interaction of CMEs:  A study on an event on June 5-7 2011 seems of high interest concerning the interaction of 2 CMEs, with a line-up of MESSENGER, VEX, STEREO-A.  Direct comparison of GCS results to in situ data:  The COR2/Graduate Cylindrical Shell modelling results from WP3 will be compared to the detections in the ICMECAT, to identify the predictive capability of the GCS model. |
| WP4 SUMMARY/NEXT STEPS:  Building on the core heliospheric CME catalogue of WP3, WP4 builds links to catalogues of in situ and coronal phenomena and events, and this involves the inspection and analysis of a range of databases. Careful management of the developing databases to exploit the links and enable the best physical interpretations is at the heart of WP4. The following bullets highlight the next steps in this WP:  4.1: UGOE will search for the coronal sources of the CMEs in the selected event list (from task 4.3, provided by UNIGRAZ/UH), and send back to UNIGRAZ/UH until March 2016.  4.2:   * UPS will look at if MAVEN solar wind plasma and magnetic field data are available, and process them to into our normalized formats. * UNIGRAZ: Magnetic field cruise data with RTN coordinates for MESSENGER for 2007-2011 will be processed. Update MESSENGER data at Mercury until mission end April 2015 when available. * IMPERIAL and UNIGRAZ: Append events from the Winslow et al. (2015) MESSENGER catalogue to the Good et al. (2015, submitted) catalogue. Add “ICME\_start\_time“ to VEX and MESSENGER catalogue from Simon Good (currently contains “magnetic obstacle“ start and end times). Update VEX ICME catalogue until the end of mission. * UH updates L1 Wind data to end of 2015, STEREO-B to September 27 2014, and STEREO-A as available. * UH starts to determine ICME parameters from the ICMECAT and DATACAT. MFR modeling (Minimum Variance Analysis and Grad-Shafranov reconstruction when applicable) will be applied for suitable events in the ICMECAT using the DATACAT.   4.3:   * UNIGRAZ/UH: a list of CME events in ICMECAT that can be traced back to ARRCAT and HICAT will be established and given to UGOE, to be done until end of January 2016. Stealth CMEs (no coronal signatures) and very slow CMEs need particular attention. * UNIGRAZ/UH: With the coronal sources provided by UGOE, the first LINKCAT will be established, and its descriptiom will be put into a report and released on the HELCATS products website until April 2016, fulfilling the deliverable D4.1. * ROB: directly compare GCS results from WP3 (UGOE) to ICMECAT (similar to Rodriguez et al. 2011, Solar Phys. 270, 561)   Papers related to WP4 with HELCATS in acknowledgements:  Good, S. W., Forsyth, R. J., Raines, J. M., Gershman, D. J., Slavin, J. A., & Zurbuchen, T. H., Radial Evolution of a Magnetic Cloud: MESSENGER, STEREO, and Venus Express Observations, The Astrophysical Journal 807, 177 (2015). doi:10.1088/0004-637X/807/2/177  Good et al., Interplanetary Coronal Mass Ejections observed by MESSENGER and Venus Express, Solar Physics, submitted, 2015.  Möstl, C., et al., Strong coronal channelling and interplanetary evolution of a solar storm up to Earth and Mars, Nature Communications, 6, 7135 (2015). doi:10.1038/ncomms8135 open access.  Mays, M. L., Thompson, B. J., Jian, L. K., Colaninno, R. C., Odstrcil, D., Möstl, C., Temmer, M., Savani, N. P., Taktakishvili, A., MacNeice, P. J., & Zheng, Y., Propagation of the 7 January 2014 CME and Resulting Geomagnetic Non-Event, The Astrophysical Journal,  812, 2, 145, 15 pp, 2015. <http://fr.arxiv.org/abs/1509.06477> |

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| WORK PACKAGE 5 (WP5):  PRODUCING A DEFINITIVE CATALOGUE OF CIRs IMAGED BY STEREO/HI THAT INCLUDES VERIFIED MODEL-DERIVED KINEMATIC PROPERTIES |
| WP5 ACTIVITY TYPE: UPS |
| WP5 DURATION: MONTHS 1 – 36 |
| WP5 LEAD BENEFICIARY: STFC (1) |
| WP5 LEADER: Dr Alexis Rouillard |
| WP5 CONTRIBUTORS: UH (7) |
| WP5 OVERVIEW: The primary goal of WP5 is to provide a catalogue of the spatial and temporal evolution of Stream Interaction Regions (SIRs) (and their substructures) observed by HI in 3-D, following their complete formation process using different observations (mainly imaging but also in-situ) from the Sun out to 1 AU. The output of the solar wind stream advanced catalogue will be optimized to help the space physics community in the search for clues on the origin, propagation, 3D morphology, and the planetary effects of CIRs and the slow solar wind. |
| WP5 Tasks 5.1 to 5.4  Here, we group the WP5 descriptions together to aid the report on progress, first summarising the status of identified tasks.  WP5 Task 5.1: CATALOGUING THE OCCURRENCE OF CIRs (TASK LEAD: UPS)   * Using J-maps and optimized running-difference images, we will list the times of observation of each CIR in HI images **[Task completed – i.e. includes data up to superior conjunction]** * From the CIR fitted trajectories, we will provide a catalogue of the arrival times of CIRs at Mercury, Venus, Mars, Earth, Saturn, thereby providing support to European-funded space missions around these planets **[Task completed for STEREO A – i.e. includes data up to superior conjunction]**   WP5 Task 5.2: DERIVING/CATALOGUING THE KINETIC VARIATION OF CIRs (TASK LEAD: UPS)   * We will fit the leading edge of each CIR in HI J-maps derived at all available latitudes to obtain the spatial/temporal evolution of each CIR over ~180o longitude **[Task completed for STEREO A – i.e. includes data up to superior conjunction]** and ~900 latitude **(Task in progress]**   WP5 Task 5.3: COMPARING BACK-PROJECTED CIR TRACKS WITH CORONAL SOURCES (TASK LEAD: UPS)   * Using the derived trajectories and kinematic properties of CIRs and their small-scale transients, we will determine for each CIR observed in white-light images if there is an associated coronal hole observed in EUV **[Task completed for STEREO A – i.e. includes data up to superior conjunction]**   WP5 Task 5.4: COMPARING FORWARD-PROJECTED CIR TRACKS WITH IN-SITU MEASUREMENTS (TASK LEAD: UPS)   * We will track small-scale transients to 1 AU and make a list of predicted impacts at points in the heliosphere where in-situ measurements are taken **[Task completed for STEREO A and B – i.e. includes data up to superior conjunction, and for ACE and Wind]**   *More detailed summary for Task 5.1, Task 5.2, Task 5.3 and Task 5.4:*  Using J-maps derived from running-difference images, Illya Plotnikov (recruited in May 2014) has listed the times of passage of each Co-rotating Density Structures (CDS) in the field of view of HI images obtained from STEREO-A (Goal 1 in T5.1). We then used our technique to determine the location of each CDS in the ecliptic plane as a function of time. In order to determine the time-dependent evolution of the 3-D trajectory of each SIR, the characteristic tracks left by these SIRs in the J-maps were fitted manually by clicking on J-maps produced with HI-1 and HI-2 images from STEREO-A (and in the future STEREO-B). For each SIR, we fitted first the clearest tracks associated with individual density irregularities entrained by the SIR. We then used the fact that a SIR co-rotates during its passage in the heliospheric imagers, thereby leaving a characteristic pattern in the J-map. This pattern was then computed theoretically and superposed onto the real J-map. The algorithm to trace this SIR pattern folds in a correction for the orbital motion of the spacecraft. This orbital motion changes the location of the viewpoint of the probe (STEREO-A) and can have an important effect when the SIR pattern is considered over a 180o of co-rotation in the HI field of view lasting up to 18 days. To validate a fitting we use two criteria (1) we require that the correspondence between the simulated and measured SIR patterns are satisfactory, (2) we require that the fitting technique (‘fixed-point technique’) of the trajectory of individual small-scale transients entrained by the SIR gives small errors in both speed (<40 km/s) and direction (<20o). Both of these criteria usually require that the SIR be well observed by the heliospheric imagers over a large range of elongations (>40o). CDSs that are poorly observed because (1) the white-light signature is too weak, (2) too many CMEs passed in the field of view are not fitted.  Once a SIR is fitted, we assign it a time that corresponds to the launch time of the small-scale transient entrained by the SIR that is best observed during the event. We adopted this definition of the onset time because it guarantees that if a future user wishes to find this SIR in the heliospheric imagers, the single feature will be most easily spotted as a bright signature in the images (especially with the ‘propagation tool’: c.f. WP8 for more details). The Carrington longitude and latitude of the source region of the SIR was also recorded and was systemically compared with EUV images. For each event, we recorded whether a coronal hole was present less than 20o of longitudinal separation to the East of the SIR source region and recorded its location in Carrington coordinates (Goals 1 and 2 of T5.3). The format of the catalogue is given in the first annual report.  Since the annual report, Illya Plotnikov has completed the entire catalogue of CDSs from STEREO-A (2007-May 2014). He has compared the catalogue with various sets of in-situ measurements taken by STEREO-A, ACE/Wind and STEREO-B at the predicted times of impact of the CDS. The comparison shows that the fitted CDSs have speeds close to the slow solar wind speed and not the speed of the stream interface measured in situ. He has also demonstrated that CDSs are mostly S/CIRs and that all S/CIRs are detected by HI during the solar minimum period but at solar maximum the large number of CMEs passing in the field of view prevents the identification of half of the passing S/CIRs. He has also demonstrated that CDSs occur mostly at sector boundaries in situ and at the neutral line at the Sun (see Figure 5.1) and therefore the close relation between CDSs and small transients expelled in the plasma sheet. All these findings are included in a paper that is in preparation (*Plotnikov, I., Rouillard, A.P., Davies, J.A. et al., Long-term tracking of corotating density structures using heliospheric imaging, to be submitted to Solar Physics, 2015).*  Illya Plotnikov has also investigated the possibility of fitting CDS in STEREO-B heliospheric images, however this is very difficult because the pattern left by the passage of CDSs in STEREO-B images is less clear. He has therefore developed a new scheme that is based on the STEREO-A fitting. He predicts the pattern of CDSs from STEREO-A to STEREO-B by letting the structure co-rotate in HI1-/2B, if the pattern generated in STEREO-B is reproduced, then a CDS is validated in that field of view. So we are deriving a new CDS catalogue for STEREO-B based on the STEREO-A catalogue.  Macintosh HD:Users:ALEXIS:Desktop:Capture d’écran 2015-11-03 à 09.44.35.png  *Figure 5.1: A series of stack plots derived from measurements of the near-Earth plasma (panels a-c) and of the solar magnetic field using a PFSS model (panels e-g). All Stack plots are presented in the same way, with Carrington longitudes running along the X-axis and Carrington Rotation numbers along the Y-axis. The red diamonds are the estimated Carrington rotation number versus longitude of each CDS fitted in our catalogue. All measurements refer to magneto-plasma parameters (described in detail in the body of the article). Time runs from high to low values of Carrington longitudes (right to left along the X-axis) and from bottom to top along the Y-axis [Taken from Plotnikov et al. 2015]* |
| WP5 SUMMARY/NEXT STEPS:  We have reported significant progress in the identification, cataloguing and fitting of CIR events recorded in the HI data and their comparisons to solar and in—situ data. The task summary, above outlines what has been accomplished and what is yet to be done. However, we touch also, here, on two points of interest and future work:  Illya Plotnikov is now investigating the effect of spacecraft motion which was recently demonstrated by Conlon et al (2015) to affect fitted speeds by a few tens of km/s. An updated CDS catalogue will be produced including spacecraft motion. Future deliverables (due month 24) include the fitting of the leading edge of CIRs and determination of their latitudinal extend.  Eduardo Sanchez-Diaz (recruited in Nov. 2014); has investigated the speed of small transients in HI and has found that they have very slow speeds, that these small-scale transients retain very slow speeds up to 0.3-0.4AU and investigated whether HELIOS-measured such low speeds in solar cycle 22. These speeds are indeed measured in situ up to 0.6 AU, this provided additional confidence that our speed calculations are correct. A paper was submitted to Geophysical Research Letters but after a mitigated referee report mainly due to suitability of the paper for the GRL format, we decided to re-submit to JGR. Eduardo has identified times when blobs emitted from the corona observed in STEREO-A images are propagating towards STEREO-B and L1, he is currently analysing these data.  We note two papers related with the HELCATS work in this WP:  *Sanchez-Diaz, E., Segura, K., Rouillard, A.P., Lavraud, B., Chihiro T., The very slow solar wind: origin, variability and properties, to be submitted to JGR, 2015.*  *Plotnikov, I., Rouillard, A.P., Davies, J.A. et al., Long-term tracking of corotating density structures using heliospheric imaging, to be submitted shortly to Solar Physics, 2015* |

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| WORK PACKAGE 6 (WP6):  INITIALISING ADVANCED NUMERICAL MODELS BASED ON THE KINETIC PROPERTIES OF STEREO/HI CMEs AND CIRs |
| WP6 ACTIVITY TYPE: UPS |
| WP6 DURATION: MONTHS 7 – 36 |
| WP6 LEAD BENEFICIARY: UPS (1) |
| WP6 LEADER: Dr Alexis Rouillard |
| WP6 CONTRUBUTORS: GMU (Third-party) |
| WP6 OVERVIEW: The primary goal of WP6 is to transform the catalogues of CMEs and CIRs observed by HI, accomplished in WP2/3 and WP5, into more advanced catalogues of simulations results of CIRs and CMEs. This advanced database will provide to the space community a set of simulation results optimised by assimilating direct images of the solar wind into ENLIL simulations. The delivery of these advanced catalogues will enhance forefront research on the ‘background’ solar wind (fast and slow solar wind) and on the spatial and temporal evolution of CIRs and CME shocks, and will provide unique material to study and interpret particle radiation measurements in the inner heliosphere. |
| We have addressed a subset of the goals of the four tasks defined for WP6, namely:  WP6 Task 6.1 - Assimilating HI images to model the background solar wind [Months: 7-36], UPS  The combination of the catalogues of CIRs derived in WP5.1 and WP5.2, J-maps derived from HI images and movies has been compared with preliminary synthetic J-maps of CIRs derived from numerical simulations of the background solar wind (ENLIL) through the propagation tool. We are currently dividing events in two classes: Class 1 for which a good correspondence is immediately obtained between simulated and observed height-time maps and Class 2 for which J-maps differ significantly.  Since the annual report, Rui Pinto has enabled his solar wind code to use different magnetogram sources (Wilcox Solar Observatory, Mount Wilson and Solar Dynamics Observatory) to produce potential field source surface (PFSS) maps of the coronal magnetic field. The use of other methods of coronal field reconstruction (e.g, non-linear force-free methods) is also possible. His model calculates an ensemble of solar wind profiles on a large sample of open magnetic flux-tubes covering all solar latitudes and longitudes. The wind solutions cover all the heights ranging from the surface up to about 30 solar radii (including the chromosphere and the transition region).  He has pursued his numerical effort to produce Carrington Maps of solar wind parameters at 21.5Rs to initialise ENLIL. He has run his solar wind model for several consecutive Carrington rotations (so far covering the year 2009), in addition to a later case present in the catalogue of Corotating Density Structures (S/CIRs) from WP5 (Carrington rotation 2137, May-June 2013).  He is currently using these solutions to fine-tune and and validate the model against in-situ data from OMNI at 1 AU (especially the case from 2013, which displays a highly warped heliospheric current sheet and an intricate mixture of fast and slow wind flows). The model already produces good quality predictions for the wind speed and temperature, and the fine-tuning in progress is leading to improvements on the predictions of the density of the wind flows (which is a common difficulty for all current day solar wind models). Additional sources of calibration come from detection of Alfvén wave propagation within the wind flows at the lowest-lying layers of the corona, as described in: *Morton, Tomczyk, Pinto, Nature Communications, Volume 6, 7813 (2015).*      *Figure 6.1: A sample of magnetic field-lines (in blue) extrapolated from an SDO surface magnetogram (grey-scale map over the solar surface) corresponding to the Carrington rotation 2137 used for fine-tuning the solar wind model and already present on the S/CIR catalogue (WP5).*  The developments undertaken under the scope of this task also led to the two following papers (to be submitted soon):  *Pinto, R., Brun, S., Rouillard, A.P., Wang, Y.-M., The slow and fast wind during an activity cycle, to be submitted to Astrophys. Journal, 2015.*  *Pinto, R., Rouillard, A.P., Wang, Y.-M., Grappin, R. The acceleration of the solar wind along magnetic flux tubes of varying geometries, to be submitted to Astrophys. J., 2015.*  We have furthermore started comparing our own PFSS reconstructions based on Wilcox Solar Observatory data with those using SDO data (obtained via SolarSoft), and the first results suggest that one of the key parameters for PFSS extrapolations (the height of the source-surface) should be re-evaluated for more realistic results. Fig. 6.1 shows a series of magnetic field-lines calculated via PFSS for one of the Carrington rotations tested.  The outcome of the model comes in the form of a series of maps of wind speed, temperature and density for all Carrington latitudes and longitudes at 21.5 solar radii (see example in Fig. 6.2). Rui Pinto defined a longitude-latitude grid with variable angular resolution (δθ = 2.5°, 5° and 10°, respectively on the equatorial, mid-latitude and high-latitude regions) that corresponds roughly to an uniform spatial coverage over the sphere. The set of maps can be completed at the highest resolution, if necessary.  *Figure 6.2: Maps of calculated wind speed (blue panel) and temperature (red panel) at 32 Rs for Carrington rotation 2137 (represented in Fig 6.1). The X-axis represents solar longitude and the Y-axis solar latitude.*  WP6 Task 6.2 - Assessing the use of HI to initialize ENLIL [Months: 7-36] , UPS  Since the annual report, we have read the catalogues produced by WP3 to locate CMEs relative to CIRs/SIRs by projecting their trajectories on the source surface. We are also comparing the 3-D trajectory of CME derived from WP3 with the trajectory of CMEs located by the NOAA alert system and available through synthetic J-maps derived by the ENLIL model. We are also evaluating the latitudinal extent that will have to be considered for the injection of CMEs in the MHD model. The synthetic J-maps are produced by computing the white-light heliospheric images that STEREO-A and STEREO-B should see from their vantage points. These images are derived by applying the equations of Thomson scattering from the simulated 3-D volume of electronic densities. The code computes running-difference images and like ordinary J-maps, a band of pixels is extracted along the ecliptic plane. These J-maps have so far been produced for 2012 and 2013. We integrated them in our software tool and can switch from these maps to real J-maps. In the tool, we can now superpose the catalogue of J-maps derived in WP5. This allows us to compare the trajectories and patterns left by CIRs derived from observations with the trajectories of CIRs derived from numerical simulations.  WP6 Task 6.3 - Continual assimilation of HI data in ENLIL and comparison with standard implementation techniques [Months: 7-36]  We will start this task in Month 24.  Other papers in preparation:  Rouillard, A.P., Plotnikov, I., R. Pinto,… ,The expansion of coronal pressure waves in 3-D during Coronal Mass Ejections, to be submitted to Astrophysical Journal, 2015. |

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| WORK PACKAGE 7 (WP7):  ASSESSING THE COMPLEMENTARY NATURE OF RADIO MEASUREMENTS OF SOLAR WIND TRANSIENTS |
| WP7 ACTIVITY TYPE: RTD |
| WP7 DURATION: MONTHS 10 – 36 |
| WP7 LEAD BENEFICIARY: IMPERIAL (6) |
| WP7 LEADER: Dr Jonathan Eastwood |
| WP7 CONTRIBUTORS: STFC (1); ROB (5) |
| WP7 OVERVIEW: The main goal of HELCATS WP 7 is to assess the potential for combining white-light imaging of the inner heliosphere with both ground- and space-based radio data, in particular Interplanetary Scintillation (IPS) and Type II radio bursts. WP7 is divided into two activities.  Task 7.1 (led by STFC) comprises the identification and analysis of potentially geoeffective solar wind events that are observed both by the Heliospheric Imager (HI) and IPS.  Task 7.2 (led by IMPERIAL) comprises the identification, analysis and cataloguing of solar wind transients that are observed both by HI and in Type II radio burst emission, principally detected by the STEREO/WAVES instruments.  This work package is now well underway. In particular, since the 12 month report a postdoctoral research associate has been hired at IMPERIAL to work exclusively on the HELCATS project and WP7 in particular (Vratislav Krupar, start date 1 July 2015, has considerable prior experience in the analysis of radio data). Regarding Task 7.1 the work effort has focused on the cataloguing and development of new and improved software to assist analysis. Regarding Task 7.2 the analysis of a detailed case study has already been completed and is being prepared for publication. This has established the foundations of the approach to cataloguing the data, and this process is now well underway. |
| WP7 TASK 7.1: IDENTIFYING AND ANALYSING POTENTIALLY GEOEFFECTIVE SOLAR WIND EVENTS THAT ARE OBSERVED BY BOTH HI AND IPS (TASK LEAD: STFC)  Work on this task has been performed by STFC (Mario Bisi and David Barnes). Note, since IPS data are not uniformly available (the radio-telescope systems used here run on a campaign basis only for observations of IPS) it is first necessary to establish the data availability working from the catalogues of CMEs and CIRs/SIRs provided by WP2 and WP5, respectively.  Cataloguing:   1. As a first accomplishment, all the EISCAT IPS data for the STEREO era have now been sorted. Bad and/or problematic observations have been sifted and removed. 2. Approximately half of the STEREO era data has been analysed with the UCSD IPS tomography. However, this has made use of an older version of the software. It is intended to run it for the entirety of the STEREO era with the most up-to-date versions of the CAT and the visualisation routines (see below). 3. The IPS analysis programme enables all of the available data to be analysed using an automated mode that finds the CME events. This is in progress. Planning is underway for how to best address the SIR aspects of Task 7.1 (using ideas and approaches based on those used in Bisi et al., Solar Physics, 2010).   With a data catalogue now coming online, work is focusing on the development of new software that can be used to batch-process the event list and thus make efficient progress in pursuit of the overall project goals.  Software development:   1. The IDL scripts for directly plotting the IPS P-Point into the STEREO HI field of view have been rewritten and now run much more simply via the updated geometrical routines available in SSW-IDL. This is superior to various ad hoc work-around scripting that has previously been used by e.g. Dorrian et al. and Hardwick et al. 2. More generally, the porting of the IPS data-analyses cross-correlation package to new processing machines at STFC has in fact proved more problematic than originally anticipated. These issues were of a technical and/or computational nature. Most have now been resolved with final testing currently being performed. 3. Enhancing software is also being developed. A script for finding the CME (negative-lobe) signatures in a systematic way from the analysed cross-correlation EISCAT data is being developed. 4. In collaboration with B. Jackson (UCSD), the latest IDL visualisation routines for tomography are being implemented at STFC, supplemented by discussions and meetings arranged in conjunction with the Third Remote Sensing of the Inner Heliosphere & Space Weather Applications Workshop in Mexico, Oct. 2015.   *It must be noted that the WP7 IPS work is an assessment study, as presented in the HELCATS proposal. Thus, we are not taking IPS data from all available worldwide sources and cataloguing events for a comprehensive comparison to the HI catalogues. The aim is to identify sufficient events for comparison from European sources and expertise to provide an assessment of the IPS-related work in comparison to the events seen using HI observations.*  Presentations at international meetings and conferences:  Third Remote Sensing of the Inner Heliosphere & Space Weather Applications Workshop, Morelia, Mexico, 20-24 October 2015 (Mario Bisi) |
| WP7 TASK 7.2: IDENTIFYING AND ANALYSING SOLAR WIND TRANSIENTS THAT ARE OBSERVED BY BOTH HI AND IN TYPE II RADIO BURST EMISSION (TASK LEAD: IMPERIAL; ADDITIONAL PARTICIPANT: ROB)  As previously described, the goals of this task are to develop a joint catalogue of CMEs observed in HI and S/WAVES and Wind/WAVES data, extending the catalogue with ground-based radio observations to more closely examine the source region of each CME. Height-time statistics will then be constructed and the usefulness of radio data in constraining modelling of CME lift-off will be systematically studied. Interacting CME events will be explored in detail to examine how radio data can be used to decipher event kinematics and improve forecasting.  The initial effort has focused on two parallel activities: a detailed case study of a specific event that is very well observed, and construction of the foundations of the radio catalogue.  Case study: To fully understand the complementary nature of radio observations we have focused on a specific CME in the time interval 29 November to 1 December 2013. This event produced strong radio emission that was detected by both STEREO spacecraft, allowing direction-finding techniques to be deployed (Krupar et al., 2012). Furthermore, the CME was imaged in HI and was directed such that it passed over the MESSENGER spacecraft at Mercury before being intercepted by STEREO-A. This event therefore affords comprehensive insight into CME physics, with remote optical, radio and in situ plasma data available at multiple points. It cuts across essentially all of the HELCATS work packages. This event analysis has been performed by IMPERIAL (Krupar/Eastwood) with input from ROB (Magdalenic) and has called on other partners working in particular on WP2 and WP3 as described in more detail below.  The kinematic properties of the CME have been studied in four independent ways:   1. Radio data have been used to construct height-time profiles based on density modelling and also triangulation using the radio data from both STEREO spacecraft together. 2. HI data and height/time information contained in the WP2 catalogue for this event (with input from STFC; Davies/Barnes/Byrne). 3. In situ measurements of CME shock arrival time at STEREO-A and MESSENGER (with input from WP4 IMPERIAL; Forsyth/Good). 4. Coronagraph images have been used to determine height/time and therefore speed profiles: we have examined output of the automated SEEDS algorithm and also compared to the Graduated Cylindrical Shell model (WP3 input from GOTTINGEN; Bothmer/Pluta/Mrotzek).   Several important conclusions have been revealed. The first is the very good overall agreement between the GCS modelling, the radio data (based on density model), the HI data and the in situ detection. This is illustrated in Fig. 7.1 (Krupar et al., 2015, *in prep.*). By combining coronagraph, HI, radio and in situ, this work demonstrates in new detail how radio data can be used to accurately profile the height-time behaviour of a CME, and therefore how radio data could be used in the absence of other data to estimate CME properties and constrain models of CME lift-off. The disagreement between the direction finding and the density model radio height-time profiles is most probably due to refraction of signal, which requires further investigation. The good agreement between the radio and GCS, but disagreement with SEEDS, shows that radio data analysis may provide rapid and accurate insight to the dynamics of the CME at lift-off. This event also illustrates an important link to WP4 with in situ observations. The results of this case study are now being prepared for publication and are to be submitted to Astrophysical Journal Letters.  The results of the case study have informed the approach to constructing the radio catalogue, as we now describe.  Catalogue: As a first step, the WP2 catalogue was used to establish an appropriate event time list. Associated summary plots of STEREO and Wind radio data were then generated. Each summary plot corresponds to a 24 hour interval of data centred on the HI event time. However, as shown in Figure 1 the radio data frequency range corresponds to heights that are typically inside the HI field of view. It was therefore very difficult to precisely determine which signatures in the radio data should be associated with the HI data, especially if there was significant Type III emission from flaring.  This led to a more direct comparison with coronagraph imaging. As illustrated by the case study, it is clear that coronagraph data provides contemporaneous optical information about the CME when compared to the radio data. We are using the output of WP3 and, in particular, the CME catalogue that was made available by UGOE <http://www.affects-fp7.eu/helcats-database/database.php> to combine this data with the radio data overview plots, and establish more precisely the association of different features with different events. This task has proved somewhat complex but progress has been aided by the assistance of those working on WP2 and WP3.  It is worth noting that the WP3 catalogue only contains a fraction of all the CMEs observed by STEREO as it focuses on those events that are sufficiently bright so as to be analysed with the GCS modelling technique. Most recently we have started to use the entire CME list, e.g., <http://sidc.oma.be/cactus/>, which is automatically generated in order to gain some insight into the radio properties of the dataset associated with the entire STEREO mission. This work is underway. Other continuing work tasks include establishing the parameters that should be included in the radio data catalogue, and the technical implementation of the catalogue on the HELCATS website.  Presentations at international meetings and conferences:  National Astronomy Meeting, Llandudno, Wales, 5-9 July 2015  Solar Orbiter/Solar Probe Plus Meeting, Florence, Italy, 2-4 September 2015    *Figure 7.1: Results of the height-time analysis for a CME observed between 29 Nov. – 1 Dec. 2013 (Krupar et al., 2015, in prep.). The height-time profile is determined using a variety of independent techniques. The green squares show the result of GCS modelling applied to coronagraph images. The purple triangles show the results derived from radio data applying a density model. The cyan diamonds show the output of the HI data from WP2. The yellow and red squares correspond to the observations made in situ by MESSENGER and STEREO-A. The blue crosses show the results of the radio direction finding. It can be seen that the observations made by multiple spacecraft using remote and in situ measurements were able to track the CME all the way from the corona to 1 AU. In particular, the radio data bridges the gap between the corona and interplanetary space (as observed by HI). Therefore in the absence of one or more of the optical datasets it would still be possible to predict arrival times at 1 AU.* |
| WP7 SUMMARY/NEXT STEPS:  7.1: It is expected that good progress will be made towards a complete and consistent analysed data set of the available EISCAT/ESR and LOFAR IPS data throughout the STEREO mission period. Issues pertaining to software reliability are expected to be rapidly resolved and the development of enhancing software is underway. By the end of the calendar year it is anticipated that the IPS-ENLIL code will be updated whereby the current small errors in solar rotation will be corrected as well as enhanced visualisation routines. More generally, effort is expected to focus on cataloguing in the first instance through the first half of 2016. Two case studies have been identified that may also bear significant insight and these will also be pursued as a secondary objective. Complex interacting CME events will also be identified as targets for future study in the final part of the project.  7.2: The results of the first case study are now being prepared for publication. It is expected that this will be submitted by the end of the calendar year and published in early 2016. An abstract to present the work at the Fall AGU meeting in San Francisco, USA, has been accepted and this will take place in December 2015. The work to complete the catalogue is on-going. It is expected that the first version of the catalogue will be available in early 2016. At this point, it is anticipated that initial top-level statistics concerning occurrence, duration, brightness *etc*. will be available for review by month 24 of the project. This will follow on to an assessment of which CMEs are most radio-loud, and initiate an exploration in more detail of why this is the case. It will also enable the selection of case studies consisting of interacting CME events. |

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| WORK PACKAGE 8 (WP8):  DISSEMINATION |
| WP8 ACTIVITY TYPE: OTH |
| WP8 DURATION: MONTHS 1 – 36 |
| WP8 LEAD BENEFITIARY: STFC (1) |
| WP8 LEADER: Dr Chris Perry |
| WP8 CONTRUBUTORS: UPS (3), IRAP |
| WP8 OVERVIEW: WP8 oversees the logistics of the dissemination of data and information to the project, the Commission and the wider community, with the website as the principal tool ([www.helcats-fp7.eu](http://www.helcats-fp7.eu)). Although the website is formally the responsibility of WP1, the posting of information through material posted on the website, including the basic format and information on the website, posted documentation, links etc… is the job of WP8 and this has been done, to allow efficient running of the project and the distribution of information as required. This includes the project (private) wiki page. Meeting minutes and reports are posted as required. WP8 also oversees documentation standards and formatting, including the catalogue structure. Information distributed outside the webpage such as press releases are also managed through WP8. Dissemination is also concerned with the running of the HELCATS meetings such as the six monthly project meeting in ROB in November 2014 and the Annual Workshop and project meeting in UGOE in May 2015. The WP also includes the ingestion of products into the AMDA data-mining IRAP (UPS) propagation tool, and integration with projects such as HELIO. This WP coordinates the exploitation of the project outputs, such that they feed into numerous research activities and future space weather applications. |
| WP8 TASK 8.1- PUBLICATION OF RESULTS AND CONCLUSIONS (TASK LEAD: STFC)  This WP Task focuses on the publication of documents relevant to the project, including posting on the website. In the reporting period, this includes meeting reports and minutes, teleconference minutes, the catalogues that are being constructed as the project progresses. We also note that some scientific papers have been published and more are in preparation and will be submitted soon. A list of publications by workpackage is provided on the main website. These publications are identified in the sections of this report, by WP. |
| WP8 TASK 8.2 – ANNUAL OPEN MEETINGS WITH THE SCIENCE COMMUNITY AND PRESENTATION AT MAJOR SCIENCE MEETINGS (TASK LEAD: STFC)  The first Annual Open Workshop of HELCATS took place in May 2015 hosted by Georg-August-Universitaet Goettingen. It was attended by more than 50 participants and was a key opportunity for the project to engage with the HELCATS community and the wider scientific community including detailed scientific discussion on project-related science. The Open Workshop included key presentations from the HELCATS team and from the community (the presentations are available on the meeting website which can be accessed from the main HELCATS web pages) and was a major tool for advertising HELCATS progress and activities. A full programme of talks has been scheduled and a healthy attendance from within the project and beyond guarantees that this will be major event for heliospheric physics.  Whilst a number of project-related publications are listed in the WP descriptions above, some HELCATS participants have presented HELCATS results at major meetings, e.g. Eduardo Sanchez-Diaz has presented his results on the very slow solar wind at EGU in April 2015 and at the L ‘Aquila summer school in September 2015. Rui Pinto showed the results of work on solar wind acceleration at the AIC meeting in Tampa, Florida in April 2015 and as an invited talk at the Solar Wind 14 meeting in Weihai. Alexis Rouillard presented his work in an invited talk at the IUGG in Prague and the 1st Solar Orbiter/Solar Probe+ meeting in Florence. Mario Bisi presented both general and WP7 specific presentations at the Third remote-sensing of the inner heliosphere workshop in Mexico, October 2015. |
| WP8 TASK 8.3 – INSTALLATION OF RELEVANT DOCUMENTS, CATALOGUES, PUBLICATIONS ON THE PROJECT WEBSITE (TASK LEAD: STFC)  The website has been actively utilised from the start of the HELCATS project with basic information, but also with posted meeting information, documents, reports and minutes of meetings and teleconferences. This includes the catalogues themselves. Most of the site is public, but there is a private project area where some telecon minutes are lodged.  A number of catalogues are currently available via the HELCATS site including from WP2, WP3 and WP4. The catalogues are provided in several common formats (ASCII, VOTable and JSON) to maximise their use in different applications. Detailed release notes are provided to support the catalogue use and once a catalogue is consider to be stable it is published on Figshare providing a DOI: and archive reference that will be accessible over the long term and addressing the issue of sustainability of the project results.  The catalogue web pages provide some simple filtering capabilities and options to easily copy & paste or drag and drop into tools for further filtering or visualisation such as Excel or Topcat. Examples of the catalogue web pages have been provided in some of the individual catalogue reports above. |
| WP8 TASK 8.4 - INTEGRATE WITH RELEVANT, ESTABLISHED COMMUNITY FACILITIES AND WEBSITES (TASK LEAD: STFC)  The WP descriptions, above, have discussed the progress of the catalogues as they develop, and this includes the use and links to existing catalogues and websites for comparisons. The comparison of HI and coronagraph CME catalogues is a case in point, as is the integration of the catalogue facilities at RAL (STFC) with the UK Solar System Data Centre. Another facility that is key to the HELCATS work, for WP5 and 6, is the AMDA facility at UPS. The ENLIL work is also linking to an established MHD model and we have integrated ENLIL simulations to the propagationtool.cdpp.eu so that the community can also have access to the synthetic J-maps. Thus, the project has many ‘integration’ activities in progress and we refer the reader to the relevant WP sections.  The IRAP catalogues are being produced in both .xml format (VOtables) and .ascii readable by most computing languages. The CIR catalogue developed in WP5 has been integrated in the ‘propagation tool’ (propagationtool.cdpp.eu). This has required some modification of the tool that was supervised by Alexis Rouillard. We are currently integrating the CME catalogue produce by WP2 for both the fixed point and the harmonic mean techniques. This should be operational in the coming weeks. The next step will be to create a web-service between our tools and the source archive of all these catalogues at STFC.  IRAP have also integrated ENLIL simulations to the propagationtool.cdpp.eu so that the community can also have access to these synthetic J-maps. The CME catalogues produced in WP3 with the fixed phi and harmonic mean techniques have been integrated in the propagationtool.cdpp.eu and this is used in WP6. |
| WP8 TASK 8.5 – DISSEMINATION OF INFORMATION TO THE PUBLIC AND POLICY MAKERS (TASK LEAD: STFC)  The website is openly available to the public with information about the project and access to most of the project documentation and activities. The website includes news items, events, tweets and information on meetings. There is much background information for the non-professional, interested reader. The HELCATS team members are very active in ‘live’ space projects (e.g. STEREO, Cluster, Hinode) and in the increasing space weather activities, such as ESA’s Space Situational Awareness (SSA) space weather programme. The HELCATS project has been widely discussed, especially in space weather fora, because it is seen as pioneering and spearheading the assessment of HI observations along with associated data and modelling methods in a way that can have a major influence on future space weather strategic applications. Many members of the Steering Committee, including the Coordinator and the Scientific and Technical Managers, are actively involved in SSA projects. In the UK this has involved close involvement with the Met Office, which is now an established space weather forecasting facility for the UK Government. |
| WP8 SUMMARY/NEXT STEPS:  The WP8 descriptions, above, are clearly defining an on-going process in terms of the varying aspects of dissemination. The activities, as described will continue throughout the project. The basic activities that are described relate to deliverables D8.1 to D8.8, all of which run through the entire project and are formally ‘delivered’ by month 36. |

# Project Summary

The WP reports, above, describe a unique, multifaceted project with clear aims, focusing on a comprehensive understanding of transients (CMEs and CIRs) in the heliosphere, exploiting novel instrumentation, assessing and validating models to study the heliospheric transients and including the study of solar sources and impacts at locations in the Solar System. The comprehensive, coordinated cataloguing that is at the heart of this is a resource that will provide the benchmark for many years in this field, and much progress has been made in the cataloguing activities.

All of this work is being done with a keen eye on the research aspects (e.g. in terms of understanding CME sources or CME propagation) and on space weather applications (e.g. the validation of prediction methods).

The reports, above, stress that the wealth of work being performed in the WPs is delivering as planned in this important feld.

Indeed, this report shows that, at the end of the first 18 months, there are numerous elements to the work underway with clear paths to bringing aspects together as the project progresses. In terms of the project deliverables HELCATS is on track and we are seeing impressive results in bringing together a unique set of catalogues and the assessment of a range of key methods. Thus, the project is functioning well and we anticipate, already, a solid legacy for the future.

The WP descriptions provide specific summaries and pointers to the future but we bring out one point here that is mentioned in the WP8 summary. The growth of the ESA Space Situational Awareness space weather programme is something that has taken place in parallel with the conception of HELCATS through to its implementation. It is abundantly clear that HELCATS is a thoroughly unique, global assessment of the added value of HI observations, of related data and of modelling methods, and the results will undoubtedly be fundamental to the application of HI observations in any future HI programme within a European space weather programme.

Having said that, the careful coordination of HI and associated data with the key modelling codes means that the project is well set to provide a truly unprecedented view of the heliosphere for the scientific community.