

## Bridging EUV and White-Light Observations to Inspect the Initiation Phase of a “Two-Stage” Solar Eruptive Event.

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**Abstract** The initiation phase of coronal mass ejections (CMEs) is a very important aspect of solar physics, as these phenomena ultimately drive space weather in the heliosphere. This phase is known to occur between the photosphere and low corona, where many models introduce an instability and/or magnetic reconnection that triggers a CME, often with associated flaring activity. To this end, it is important to obtain a variety of observations of the low corona in order to build as clear a picture as possible of the dynamics that occur therein. Here, we combine the EUV imagery of the *Sun Watcher using Active Pixel System Detector and Image Processing* (SWAP) instrument on board the *Project for Onboard Autonomy* (PROBA2) with the white-light imagery of the ground-based Mark-IV K-coronameter (Mk4) at Mauna Loa Solar Observatory (MLSO) in order to bridge the observational gap that exists between the disk imagery of the *Atmospheric Imaging Assembly* (AIA) on board the *Solar Dynamics Observatory* (SDO) and the coronal imagery of the *Large Angle Spectrometric Coronagraph* (LASCO) on board the *Solar and Heliospheric Observatory* (SOHO). Methods of multiscale image analysis were applied to the observations to better reveal the coronal signal while suppressing noise and other features. This allowed an investigation into the initiation phase of a CME that was driven by a rising flux rope structure from a “two-stage” flaring event underlying an extended helmet streamer. It was found that the initial outward motion of the erupting loop system in the EUV observations coincided with the first X-ray flare peak, and led to a plasma pile-up of the white-light CME core material. The characterized CME core then underwent a strong jerk in its motion, as the early acceleration increased abruptly, simultaneous with the second X-ray

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flare peak. The overall system expanded into the helmet streamer to become the larger CME structure observed in the LASCO coronagraph images, which later became concave-outward in shape. Theoretical models for the event are discussed in light of these unique observations, and it is concluded that the formation of either a kink-unstable or torus-unstable flux rope may be the likeliest scenario.

**Keywords:** Coronal Mass Ejections, Low Coronal Signatures, Initiation and Propagation

## 1. Introduction

Coronal mass ejections (CMEs) represent the largest, most dynamic phenomena that originate from the Sun (Webb and Howard, 2012). Propagating at speeds of hundreds up to thousands of kilometers per second (Yashiro *et al.*, 2004), the particle densities and energies involved can cause shocks and adverse space weather effects at Earth and elsewhere in the heliosphere (Prangé *et al.*, 2004; Schwenn *et al.*, 2005; Carley *et al.*, 2013; Liu *et al.*, 2014). They can lead to geomagnetic storms upon impacting our magnetosphere, damaging satellites, affecting communication and navigation systems, and increasing the radiation risk for astronauts (Lockwood and Hapgood, 2007). Given their potentially hazardous impact on Earth's geomagnetic environment, the physics governing their eruption and propagation needs to be understood.

Various theoretical CME models exist in the effort to reproduce the physical driver mechanisms responsible for their initiation and propagation (see the review by Chen, 2011). All are based upon some form of triggering mechanism, most likely as a result of a magnetic energy imbalance often described in the context of tether straining and release (Klimchuk, 2001). Considering the pre-eruption structure of the CME as a flux rope (*e.g.*, Chen, 1996) or strongly sheared arcade, possible causes for eruption may include flux injection and magnetic twisting (Krall *et al.*, 2001; Kliem and Török, 2006), reconnection beneath the flux rope (Moore and Labonte, 1980; Forbes and Priest, 1995; Amari *et al.*, 2003; Lin *et al.*, 2007), or reconnection between the overlying field and neighboring flux systems (Antiochos, DeVore, and Klimchuk, 1999; van der Holst, Jacobs, and Poedts, 2007; Lynch *et al.*, 2008). Such models may provide an interpretation of the observations, and thus allow some deeper understanding of the forces governing CMEs and their relationship with associated phenomena like flares (see the review by Priest and Forbes, 2002).

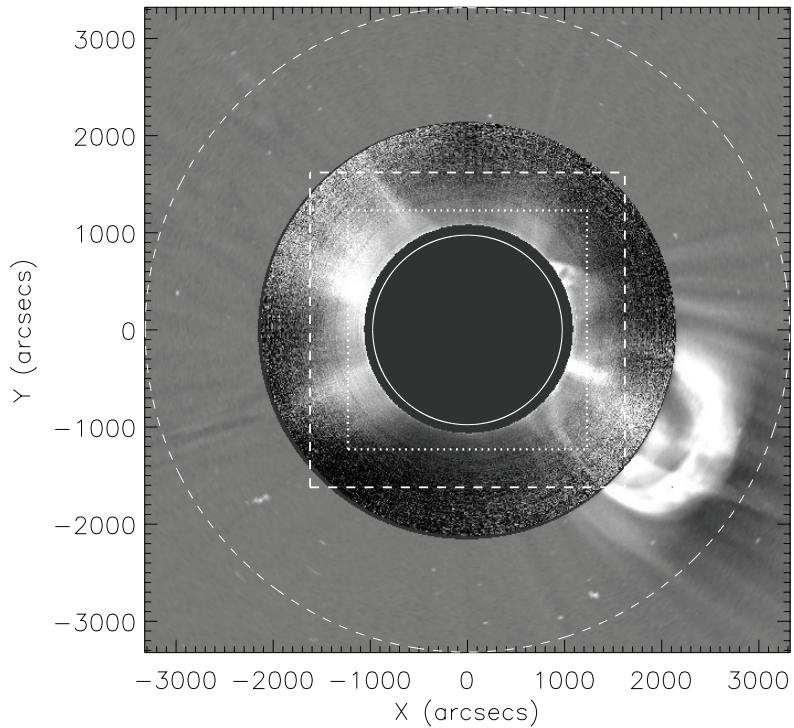
An important aspect of studying CME initiation, is the ability to resolve their low-coronal propagation and associated source regions on the disk: be it a flaring or non-flaring active region, a prominence/filament eruption or other rising loop system (Subramanian and Dere, 2001; Zhang and Wang, 2002), or else a “stealth CME” without any specifically detectable source (Howard and Harrison, 2013). Prominence lift-offs often become the core material of a CME (Gopalswamy *et al.*, 2003; Filippov and Koutchmy, 2008), and rising loops often form some part of the CME morphology (Cremades and Bothmer, 2004; Dauphin, Vilmer, and Krucker, 2006).

The low-coronal kinematics and morphology of CMEs provide insight into the early forces at play, and so a rigorous study of such phenomena is key to understanding the physics involved. However, a difficulty exists in studies of coronal structures that are prone to low signal-to-noise ratios in the observational data. Low-coronal white-light observations using a coronagraph are problematic due to the strong radial brightness gradient and issues with scattered light in the instrument, while extended EUV disk observations are problematic due to the strong drop-off in emission brightness with increasing coronal height. These common issues with solar observational data motivate the ongoing development and use of advanced image processing techniques to suppress noise and enhance structures in the image data (Stenborg and Cobelli, 2003; Morgan, Habbal, and Woo, 2006; Young and Gallagher, 2008; Pérez-Suárez *et al.*, 2011; Druckmüllerová, Morgan, and Habbal, 2011).

In this paper, a relatively unique “two-stage” solar eruptive event is studied with a combination of multiple, overlapping EUV disk observations and white-light coronal observations. This complements the study of Su *et al.* (2012), who reported this event as evidence for secondary heating during a flare and associated CME. In Section 2 we describe the observations and use of multiscale image processing methods. In Section 3 we describe the event, that occurred on 8 March 2011, and how the combination of observations and techniques can provide deeper insight into the initiation phase of CMEs. A discussion of the interpretation of this study is presented in Section 4, and final conclusions in Section 5.

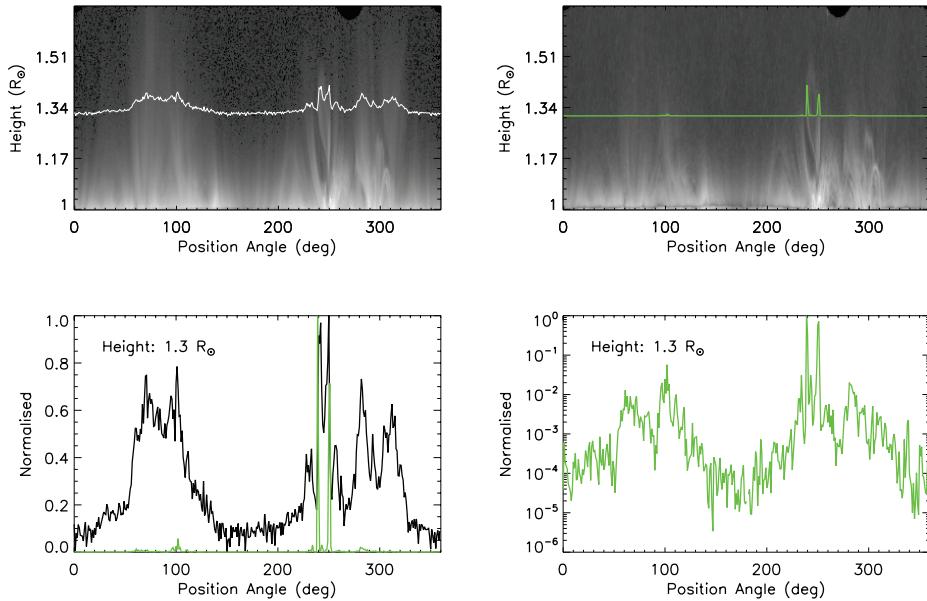
## 2. Observations and Techniques

In order to connect CMEs to their source regions, data from disk imagers such as the *Sun Watcher using Active Pixel System Detector and Image Processing* (SWAP: Seaton *et al.*, 2013, Halain *et al.*, 2013) on board the second *Project for Onboard Autonomy* (PROBA2: Santandrea *et al.*, 2013) and the *Atmospheric Imaging Assembly* (AIA: Lemen *et al.*, 2012) on board the *Solar Dynamics Observatory* (SDO: Pesnell, Thompson, and Chamberlin, 2012), may be used in tandem with coronagraph observations. However, difficulties in the interpretation of the observed features arise due to the varying instrument specifications, particularly the limitations on their fields of view, image passbands, and cadences of observations. In order to bridge the gap between the EUV observations of the low corona and the white-light images of the upper corona, the SWAP imager was used in conjunction with the ground-based Mauna Loa Solar Observatory (MLSO) Mark-IV K-coronameter (Mk4: Elmore *et al.*, 2003) to directly compare the observations of CMEs as they erupt through the overlapping fields of view (Figure 1). SWAP has a spectral bandpass centered on 174 Å, with 3.2 arcsec pixels over a 54×54 arcmin field-of-view, and a cadence of ≈2 minutes. Mk4 is a rotating coronameter that produces white-light images of the polarization brightness of the corona from 1.1–2.8 R<sub>⊕</sub> at a cadence of ≈3 minutes over a five-hour observing day. This allows a direct correspondence of features in the EUV images with those in the white-light images, and can therefore provide new insight into the initial phases of CME eruption and acceleration.



**Figure 1.** A zoomed-in LASCO/C2 image of the 8 March 2011 CME, with an MLSO/Mk4 image overlaid in the range  $1.1\text{--}2.2 R_{\odot}$ , at times 20:24 and 20:22 UT respectively. The C2 image has been processed *via* the CORIMP techniques of normalizing radial graded filter (NRGF) and quiescent background subtraction. It has been trimmed to a half-width of  $3.4 R_{\odot}$ , which is the upper limit of the PROBA2/SWAP field-of-views as indicated by the dashed circle. The SWAP field-of-view during nominal operations is indicated by the outer dashed box, and the SDO/AIA field-of-view is indicated by the inner dotted box. The limb of the Sun behind the occulter is indicated by the solid white circle. A CME is observed off the Southwest limb as a bright loop structure with some inner core material, as seen here in the Thomson-scattered white-light coronagraph images. It is clear how the fields of view of SWAP and AIA can be used to bridge CME observations to the low corona and solar disk, for gaining insight into their initiation phase.

Methods of multiscale image processing have been developed in recent years for use on coronagraph images to enhance the underlying structure (Stenborg, Vourlidas, and Howard, 2008; Young and Gallagher, 2008; Gallagher *et al.*, 2011). The fundamental idea behind these methods is to highlight details apparent on different scales within the data. Therefore, multiscale techniques provide an ability to remove small-scale features in images, essentially suppressing the noise such that the structures of interest can be revealed in greater detail. By applying them to coronagraph images, the morphology of CMEs as they propagate through the corona in a sequence of observations can be determined with better accuracy than previously possible, and can allow a characterization of



**Figure 2.** The top two panels show polar-unwrapped images of the solar corona across the PROBA2/SWAP field-of-view on 8 March 2011 at 19:53:59 UT; left being the level-1 data, right being the multiscale enhanced data (shown at log-scale image intensity). Across each image, at a constant height of  $1.3 R_{\odot}$ , a normalized intensity profile is plotted in white continuous and green continuous lines, respectively, to demonstrate how the background coronal structure is suppressed by the multiscale techniques, to highlight only the complex structure of the prominence. The bottom left plot shows a direct comparison of the two intensity profiles (drawn in black continuous and green continuous lines), where the prominence is located between  $230 - 260^{\circ}$ . The bottom right plot shows, with a green continuous line, a log scale of the normalized intensity profile across the enhanced image to demonstrate that the rest of the coronal structure is still present, just strongly suppressed relative to the prominence material. (The green continuous line in all cases corresponds to the intensity slice of the multiscale enhanced image at top right.)

the erupting structure to determine various properties in their evolution (Byrne *et al.*, 2009, 2010, 2012).

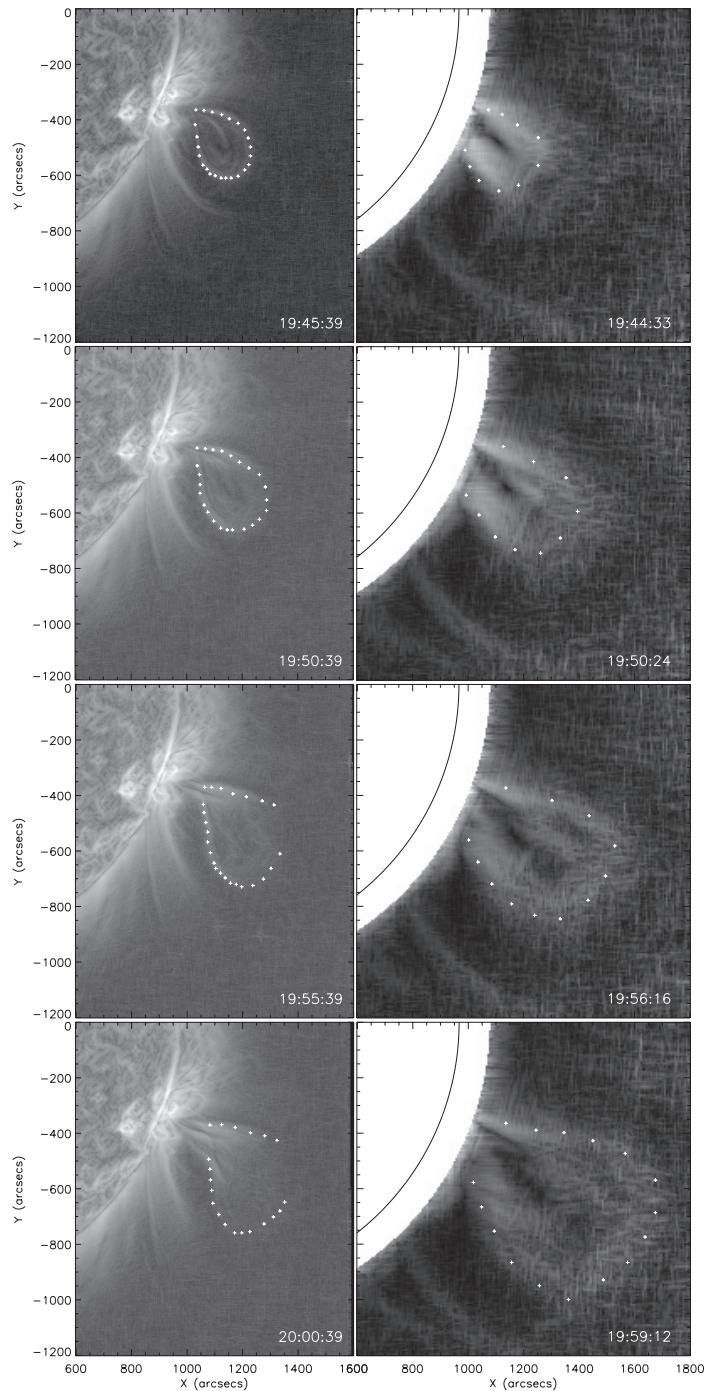
Here, multiscale methods are used on the Mk4 coronameter and SWAP EUV imager, to provide insight into the low-coronal morphology of erupting structures that form CMEs. Details on the fundamental techniques are outlined in Young and Gallagher (2008) wherein the magnitude of the multiscale gradient is used to show the relative strength of the detected edges in the image structure at a particular scale of the multiscale decomposition (*i.e.*, the strongest edges appear brightest). To further increase the signal-to-noise ratio of the edge detections, the magnitude information from the scales most relevant to the coronal structures of interest may be multiplied together, neglecting the largest scales that smooth out the coronal signal, and the smallest scales that reveal the finer structure and noise (see Byrne *et al.*, 2012, for details). Thus the magnitude of the multiscale gradient across the dominant edges of coronal loops and CMEs is further enhanced for subsequent characterization of their morphology.

Figure 2 shows the effectiveness of the multiscale techniques for detecting the structure of an ejection observed by SWAP at 19:53:59 UT on 8 March 2011. The top left image shows the level-1 processed data, polar-unwrapped about Sun-center at coronal heights of  $1 - 1.7 R_{\odot}$ . The top right image shows the result of the multiscale filtering technique applied in such a manner as to enhance the edges of the detected structure in the data. The bottom left plot shows the comparison of a normalized intensity profile at a height of  $1.3 R_{\odot}$  in each, revealing how the multiscale techniques best characterize the complex structure of the erupting prominence material located between  $230 - 260^{\circ}$ . The bottom right plot shows a log scale of the normalized intensity profile across the multiscale filtered image to reveal the suppressed signal of the background features.

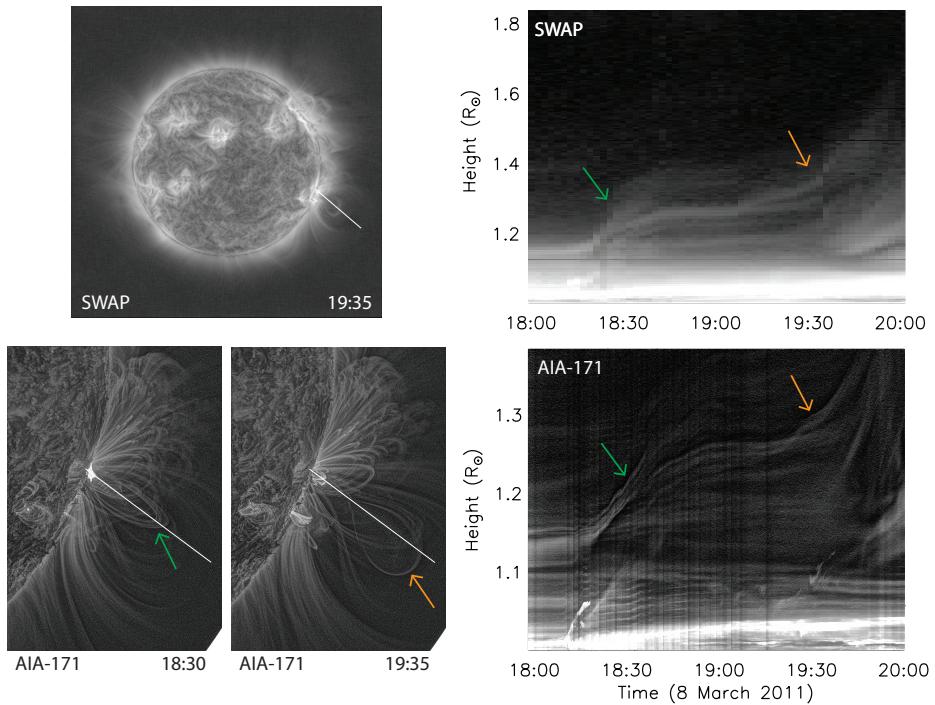
The extended corona from  $\approx 2 - 30 R_{\odot}$  is often observed with the *Large Angle Spectrometric Coronagraph* (LASCO: Brueckner *et al.*, 1995) on board the *Solar and Heliospheric Observatory* (SOHO: Domingo, Fleck, and Poland, 1995) which orbits the L1 point. Coronal structures, and specifically CMEs, have been studied in the white-light image data from these instruments through the use of a number of steps outlined in the Coronal Image Processing package (CORIMP: Morgan, Byrne, and Habbal, 2012; Byrne *et al.*, 2012). Here, these techniques are extended for use on the MLSO/Mk4 coronameter data. The Mk4 data is prepared *via* an instrumental vignetting function that maximizes the image contrast by offsetting the radial brightness gradient in order to best reveal structures such as CMEs and streamers. The multiscale methods are then applied to the data in order to produce magnitude images of the relative edge strengths in the images, and highlight the detected structure. Figure 3 shows four multiscale enhanced images from each of the SWAP and Mk4 observations of the 8 March 2011 event, allowing a point-and-click characterization of the erupting structure of interest. This allows, for example, an ellipse fit to the outward propagating fronts, to obtain kinematical and morphological information as described in Byrne *et al.* (2009).

### 3. “Two-Stage” Solar Eruptive Event

A CME erupted from the Southwest active region NOAA 11165, first visible in LASCO/C2 at 20:12 UT on 8 March 2011. The active region caused numerous flares during its evolution across the disk, among them an M4.4 flare at GOES start-time 18:08 UT (peaking at 18:28 UT) associated with the rising loop system that later erupted to form the core material of the CME. Of particular interest for this event, is the “two-stage” X-ray flare profile seen in the GOES flux, identified as two individual M-class flares separated by almost 2 hours. Su *et al.* (2012) present this as evidence for a secondary heating phase. The loop system evolution was visible up to  $\approx 1.3 R_{\odot}$  in AIA images, at which height a set of loops that were most strongly observed in AIA 171 Å images began to erupt, coinciding with the time of the secondary M1.4 flare  $\approx 20:00$  UT. These were observable to a height of  $\approx 1.6 R_{\odot}$  in the larger field-of-view of the SWAP 174 Å imager. The observations of the source active region are exhaustively reported by Su *et al.* (2012), though it is noted that the system of loops they track (in the stack



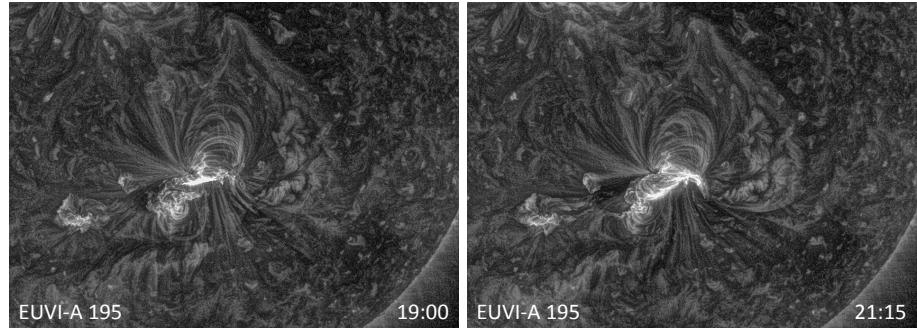
**Figure 3.** SWAP (left) and Mk4 (right) observations of the erupting loop system that formed the core of the CME on 8 March 2011. The images have been processed using a multiscale filter, and have intensities that represent the relative magnitudes of the edges in each image. The images are zoomed to the same scale, with Mk4 extending further in the x-direction. The erupting structure was traced manually by a point-and-click procedure, as shown.



**Figure 4.** Stack plots were generated for a radial intensity profile through the enhanced SWAP and AIA-171 observations. The profile location (white straight line) is shown in a full field-of-view SWAP image at 19:35 UT (top left), and at two stages of the eruption in cropped AIA-171 images at 18:30 and 19:35 UT on 8 March 2011 (bottom left). The stack plots (right) clearly show the two-stage eruption, with green and orange arrows to indicate the different loop structures of interest; the latter of which becomes the CME core that is tracked through the SWAP, Mk4 and LASCO observations. An online movie (20110308\_AIA.mov) accompanies this figure.

plot of their Figure 4) is not the same as the specific flux rope loops observed and tracked here, which erupted as the core of the CME. Figure 4 shows these different sets of loops that correspond to the two stages of the eruption. The CME core was then observed in the white-light Mk4 coronameter images to a height of  $\approx 2.2 R_{\odot}$ , entrained within a faint CME bubble that started to become visible at these heights before being clearly observed to propagate outwards in the extended LASCO coronagraph images. This type of cavity-CME morphology is often seen in low coronal observations, studied for example by Gibson *et al.* (2006) for long-lasting structures seen in Mk4 data, which erupt to become the typical three-part CME structure with a bright front, darker cavity and bright core (Illing and Hundhausen, 1986).

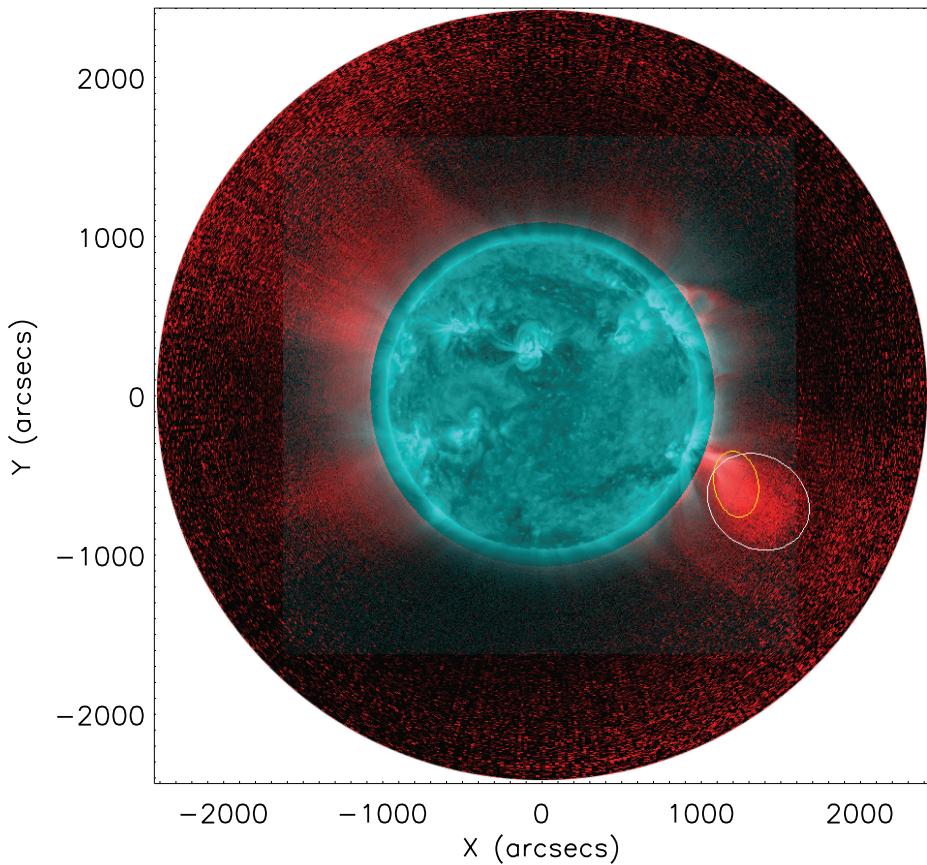
The *Solar Terrestrial Relations Observatory* (STEREO: Kaiser *et al.*, 2008) *Ahead* spacecraft was at a separation of almost  $88^{\circ}$  from Earth, providing direct observations of the active region with the *Sun-Earth Connection Coronal and Heliospheric Investigation* (SECCHI: Howard *et al.*, 2008) *Extreme Ultraviolet Imager* (EUVI: Wuelser *et al.*, 2004) at wavelengths of  $171 \text{ \AA}$ ,  $195 \text{ \AA}$ ,  $284 \text{ \AA}$  and  $304 \text{ \AA}$ . These images, particularly the  $195 \text{ \AA}$ , reveal a top-down view of the loop



**Figure 5.** EUVI-A 195 Å observations of the active region NOAA 11165 at times 19:00 and 21:15 UT on 8 March 2011. The left image shows the post-flare loops of the first stage of the eruption, following the initial X-ray flare peak at  $\approx$ 18:28 UT; while the right image shows the post-flare loops of the second stage of the eruption, following the secondary X-ray flare peak at  $\approx$ 20:15 UT. An online movie (20110308\_EUVI\_195.mov) accompanies this figure.

structures involved in the two-stage flaring eruption. Initial post-flare loops were observed following the M-class flare X-ray peak from about 19:00 UT, proceeded by relatively higher post-flare loops following the secondary X-ray peak from about 21:15 UT (Figure 5). The outward propagation of the erupting loop structure is also very faintly visible, with a slow and steady rise before the secondary eruption, and appearing to have undergone a slight clockwise untwisting motion and/or possible asymmetric expansion (see the movie 20110308\_AIA.mov). Thus it may be that the erupting loops, oriented North-South above the East-West flare ribbons seen in EUVI-A, were impeded by overlying field lines oriented in a more east-west direction as seen in the AIA images.

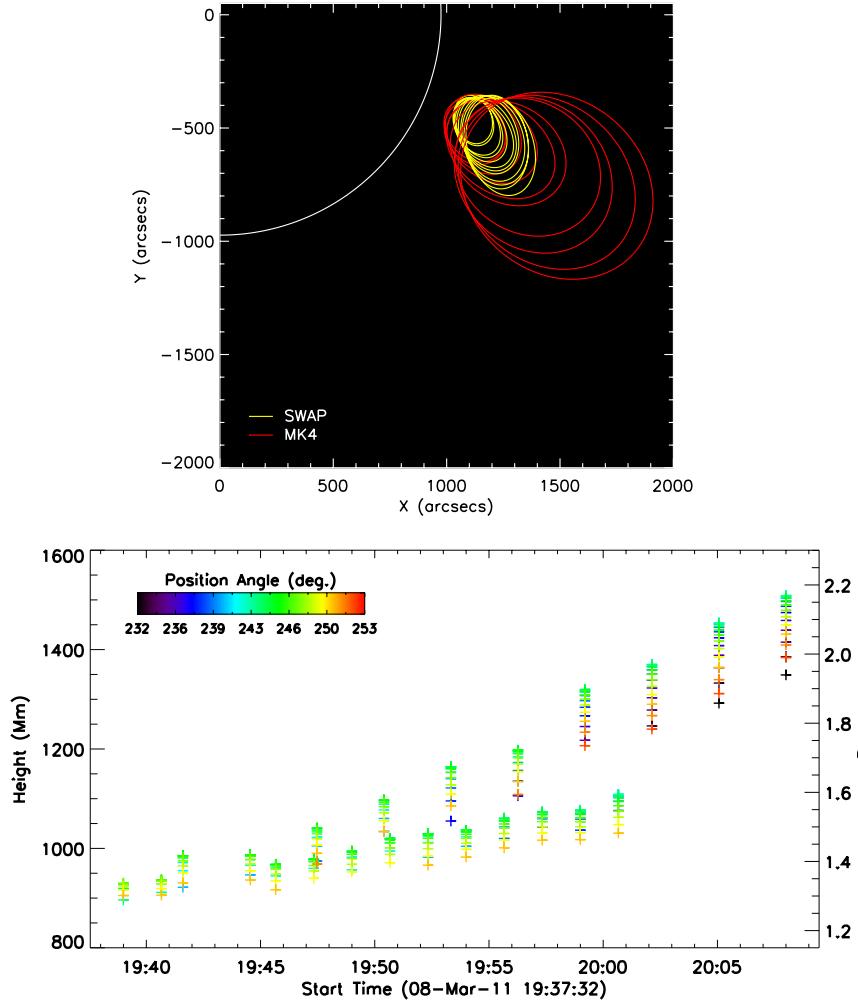
Morgan, Jeska, and Leonard (2013) reported on the expansion of the loops in NOAA 11165 into the extended solar corona in the few days leading up to this CME. The region lay beneath a helmet streamer structure that appeared to contain the observed coronal loops, with a number of faint brightenings due to small outward-propagating plasma blobs. These, and the pointed shape of the rising loops, are postulated to be indicators of helmet streamer interchange reconnection at the apex of the closed field (Wang *et al.*, 2012). A subsequent brightening and expanding of the loops, accompanied by a swelling of the helmet streamer, preceded the CME from this region. This is evidence for an energy input to the system that leads to an explosive release. Such a process manifested as the two-stage solar eruptive event outlined in this article and in the observations of Su *et al.* (2012). The AIA images of the active region underlying this system were processed with a multiscale Gaussian normalization technique (Morgan and Druckmüller, 2014) to enhance the observations of the coronal loop evolution over the course of the event. The right panels of Figure 4 show stack plots obtained from a radial intensity profile through both the enhanced SWAP and AIA-171 observations. The stack plots clearly show the two stages of the eruption, with green and orange arrows on the AIA images indicating the different loop structures that each intensity track corresponds to; the latter being the flux rope structure that erupts as the CME core. The loop structures



**Figure 6.** A merged SWAP (blue) and Mk4 (red) image with the ellipse fits to the characterized CME core material as observed by each instrument at 19:59 UT on 8 March 2011. An online movie (20110308\_combined.mov) accompanies this figure.

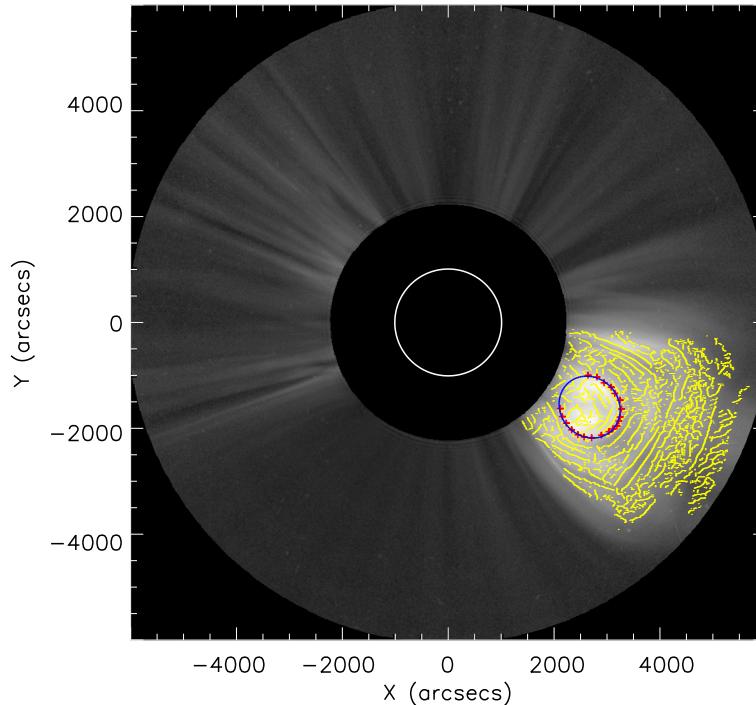
in the AIA-171 images also exhibit a kink-unstable topology, with a single twist observed most clearly after the first stage of the eruption. This is especially highlighted by the path of the upflowing material into the southern portion of the loops, tracing the kinked loop structure (see the movie 20110308\_AIA.mov).

In order to best reveal the eruption material when comparing the low signal-to-noise SWAP and Mk4 images, multiscale methods of noise suppression and edge enhancement were employed, as discussed above. This allowed a robust point-and-click characterization of the CME core material, which was the brightest structure to be tracked through the different imagers when the CME front was not yet fully formed (Figure 3). The rising loop system observed with SWAP and the erupting CME core material observed with Mk4 coincided both temporally and spatially, at least initially, and each was characterized by ellipse fits to the detected front edges of the core flux rope structure. Figure 6 shows an overlay of SWAP and Mk4 images during the eruption at times 19:58:59 and 19:59:12 UT respectively, with the ellipse fits to the erupting fronts. Figure 7 shows the



**Figure 7.** Top: The SWAP and Mk4 ellipse fits to the characterized CME core material over the course of the eruption. Bottom: The height-time profile of the characterized eruption observed simultaneously with the SWAP imager and Mk4 coronameter, where the plus symbols represent the span of heights of each of the ellipse fits, as indicated by the colorbar. The spatial and temporal offsets are due to the different speeds of the eruptions and different cadences of the instruments. The erupting EUV loops observed with SWAP moved at a speed of  $\approx 100 \text{ km s}^{-1}$ , while the associated core of the CME observed with Mk4 reached a speed of  $\approx 400 \text{ km s}^{-1}$ .

progression of the ellipse fits over the course of the eruption, indicating how the white-light material observed with Mk4 propagated away from the source quicker than the EUV material observed with SWAP. The different height-time profiles show the material in the Mk4 images attained a speed of  $\approx 400 \text{ km s}^{-1}$ , while the associated erupting loop structures in the SWAP images moved at only  $\approx 100 \text{ km s}^{-1}$ . This may be, at least partly, due to the increased intensity drop-off

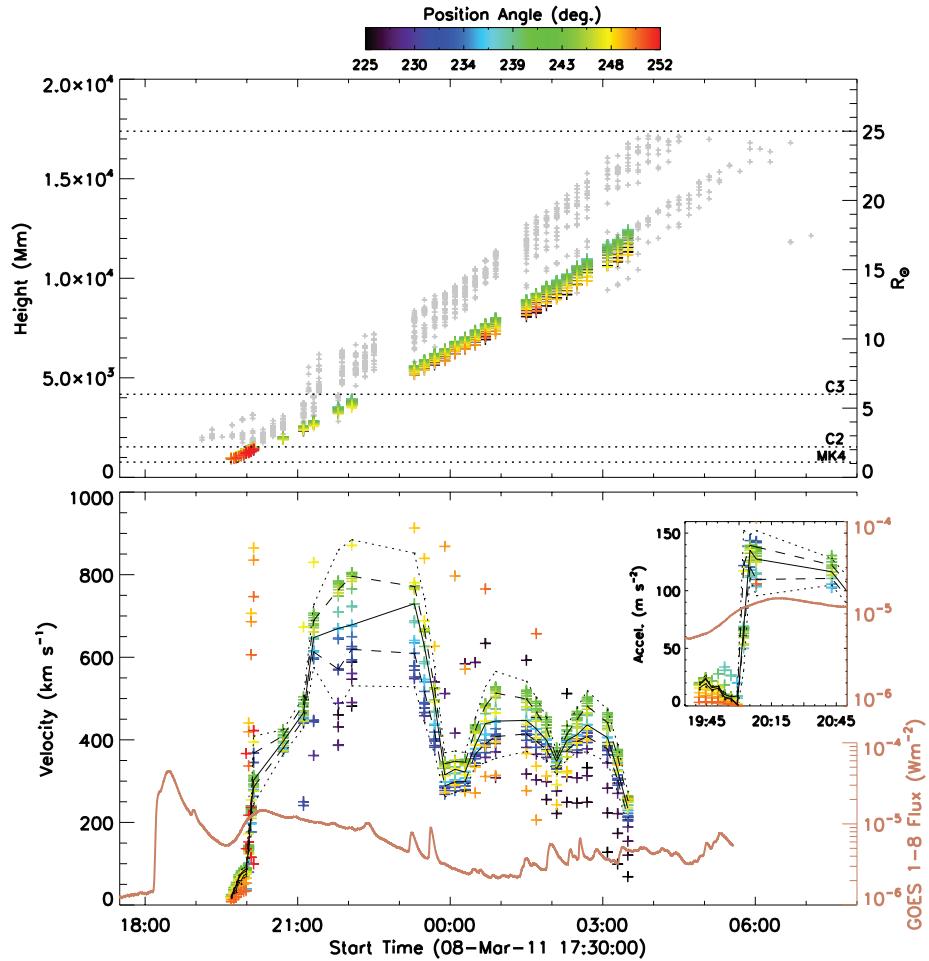


**Figure 8.** A radially filtered LASCO/C2 image of the CME at 21:06 UT on 8 March 2011. The yellow contours trace the edges in the detected CME structure (from the automated CORIMP catalog), with red points manually clicked along the corresponding core material, and the resulting ellipse fit in blue.

of the EUV emission with height, and possible plasma pile-up in the white-light observations. It is also important to note that Mk4 observes the real density-enhanced structure, while the SWAP EUV is also sensitive to the temperature of the coronal material.

The CME was observed to have a typical three-part structure that propagated out through the corona with a bulk speed in the range  $\approx 400 - 600 \text{ km s}^{-1}$  (based on the kinematics of the CME front detected and tracked in the CORIMP catalog<sup>1</sup>). The core of the CME was manually tracked *via* the multiscale methods and ellipse fits discussed above (an example LASCO image is shown in Figure 8). The resulting kinematics are plotted in Figure 9, where the height-time measurements of the CME core, observed with Mk4, C2 and C3, are shown in color (corresponding to the position angle of the measurements across the plane-of-sky) overlaid on the CME front height-time measurements shown in gray for reference. The Savitzky-Golay filter is used to derive the velocity and acceleration profiles for the CME core; whereby a distribution of velocity and acceleration values is obtained at each data point, with the corresponding median, interquartile range,

<sup>1</sup><http://alshamess.ifa.hawaii.edu/CORIMP/>



**Figure 9.** The kinematic profiles of height, velocity, and (inset) the early acceleration phase of the CME core material, detected and characterized *via* the multiscale edge enhancement and ellipse fits (see Section 2 and Figs. 7 and 8). The automated CORIMP CME detections provide the height-time measurements of the CME front shown in gray for reference, with the height-time measurements of the CME core plotted in color according to their position angle. The fields of view of the Mk4, C2 and C3 instruments are indicated by the horizontal dotted lines, covering a useable range of  $1.1 - 25 R_{\odot}$ , with the CME core tracked to approximately  $18 R_{\odot}$ . A Savitzky-Golay filter was applied to the height-time measurements to obtain distribution profiles of velocity and acceleration, with the median, interquartile ranges, and upper/lower fences over-plotted in solid, dashed and dotted lines respectively. Overlaid is the GOES X-ray 1–8 Å flux profile (in brown continuous line), showing the double-eruption peaks at about 18:28 and 20:15 UT, the latter of which coincides with the CME jerk (abrupt increase in acceleration; see inset).

and upper and lower fences overlaid on each profile as solid, dashed and dotted lines respectively (see Byrne *et al.*, 2013 for a details). The inset, corresponding to the acceleration phase of the CME core, shows initial values of approximately  $\approx 20 \text{ m s}^{-2}$  increasing to  $\approx 130 \text{ m s}^{-2}$ , referred to as the CME jerk by Schrijver *et al.* (2008). The steepness of this increase may in part be attributed to the

numerical effects of the data gap between Mk4 and C2, where the spline fit of the Savitzky-Golay filter (operating on a moving window of six neighboring data points) compensates somewhat for the increase in velocity that occurs between these two fields of view. The magnitude of the velocity, however, was verified by inspection of the different profiles within the separate instrument fields of view, for different degrees of polynomial fits and Savitzky-Golay filter sizes. That is, the rise in velocity was determined to be real, though its true steepness might be better quantified if it were possible to obtain a more complete set of measurements. Nevertheless, it occurs in synchronization with the second rise in the X-ray flare profile, being an indication of a very fast energy release that allowed the explosive eruption of the CME to begin, before later attaining a speed akin to the local solar wind speed. This double acceleration profile is further evidence for the complexity of the system whose morphology has been observed to change dramatically over the course of its evolution (Su *et al.*, 2012; Morgan, Jeska, and Leonard, 2013).

#### 4. Discussion

This particular event is especially interesting in the context of the flare-CME relationship. The two-stage flaring profile of the erupting loop system is evidence for a secondary heating process (Su *et al.*, 2012), indicating two stages of magnetic reconnection that occur to first change the topology of the system and then allow for the subsequent flux rope eruption. This scenario demonstrates that a loss of stability occurred initially, to allow the loop system to rise and alter its magnetic configuration with an explosive energy release detected as an M-class flare. This was followed by a secondary energy release that allowed the underlying flux rope to erupt through the corona as a CME, with the production of a second X-ray peak and post-flare loops at the CME footpoints. In a similar study, Zhou *et al.* (2006) combined SOHO *EUV Imaging Telescope* (EIT) and LASCO observations of the initiation phase of two successive prominence-eruption CMEs, and determined that they were driven by the kink instability and mass drainage with an impulsive acceleration onset resulting from magnetic reconnection beneath the filament.

It is not wholly clear why we observe different rates of motion of the material in the SWAP and Mk4 images during the CME onset (Figure 7). One reason is possibly that the material observed by the different instruments corresponds to different parts of the erupting structure at different temperatures, which are difficult to dissociate from each other on the plane-of-sky. Bain *et al.* (2012) observed a very similar effect in the four AIA passbands they used to track an erupting plasmoid, and demonstrated that different temperature structures do exhibit different eruption speeds (see, in particular, their Figure 4). The effect may further be attributed to the greater loss of signal in the EUV than in the white-light as the eruption proceeds to the edge of the field-of-view (diminishing in its signal-to-noise ratio while the white-light may also be showing plasma pile-up); though this is difficult to reconcile with the observations that appear to show consistent loop structures as characterized in our analysis. Indeed if the

offset is true and the different parts of the structure did undergo such differing rates of propagation on the same plane-of-sky, it may be that the trailing part simply became a different portion of the main CME and/or underwent a delayed jerk in its motion, which is not observed as the EUV signal diminishes towards the edge of the SWAP field-of-view. This would imply that there was some form of delayed or staggered eruption occurring throughout the CME structure, or that an expansion effect took over as the CME bubble formed, that created the observed offset between the different observations.

Either way, this event is intriguing in its possible interpretation. The post-flare loops in the first stage may be likened to the typical ones in the standard flare model (CSHKP: Carmichael, 1964; Sturrock, 1966; Hirayama, 1974; Kopp and Pneuman, 1976). Thus the early phase must be driven, at least in part, by a non-ideal process. The erupting loops, which were oriented North-South (above East-West flare ribbons seen in EUVI-A), were then impeded by overlying field lines as seen in AIA data, causing a  $\approx 50$  minute stop in the eruption. For the eruption to proceed, the upward magnetic pressure force must overcome the downward magnetic tension force, following either an energy build-up and/or a topological change of the system during the continued reconnection in the decay phase of the first stage. The jerk in the CME motion when it quickly accelerates during the second stage flare is evidence of this catastrophic energy release.

The exact physical connection between these two stages is unclear. For example, one possibility we investigated is that of tension reduction, through a mechanism such as magnetic breakout whereby the overlying arcades are removed by coronal reconnection. This was motivated by two factors. Firstly, the SDO *Helioseismic and Magnetic Imager* (HMI) magnetogram data in the days preceding the eruption showed the active region evolving into a possible multi-polar configuration. But since the magnetogram data on the limb are not reliable (*i.e.*, we cannot directly observe or extrapolate the active region topology) this does not provide strong evidence. Secondly, van der Holst *et al.* (2009) demonstrated the “bugle effect” in a breakout model associated with a streamer blowout CME like this one. This follows the observations of Morgan, Jeska, and Leonard (2013) that showed a streamer swelling on the limb above the active region to large heights in the LASCO coronagraphs, which evolved slowly over the preceding days in the lead-up to the eruptive event. Helmet streamers like this one are sometimes observed prior to a CME initiation (Hundhausen, 1993) though not always to such great heights due to the low signal-to-noise ratio – an issue overcome by the dynamic separation technique of Morgan, Byrne, and Habbal (2012). van der Holst *et al.* (2009) also described how the rising flux rope can “snow plow” the plasma ahead of it, with the effect of broadening the helmet streamer and facilitating reconnection between the flux rope and the helmet field. Flare reconnection then might set in due to the expansion of the magnetic field in the wake of the CME eruption as indicated by the elongated “X”-shaped structure and post-flare loops in the observations of Su *et al.* (2012).

However, since the presence of a quadrupolar topology required for breakout reconnection is not clear, we consider the formation of a kink- or torus-unstable flux rope (Török, Kliem, and Titov, 2004; Aulanier *et al.*, 2010), with evidence of internal tether cutting during the eruption (c.f., Raftery *et al.*, 2010). Zuccarello

*et al.* (2014) observed evidence of line-tying reconnection and the restructuring of a filament that facilitated its eruption due to the torus instability. Since here we observe a rather unique “two-stage” event, it may be posited that a facilitator (such as tension reduction *via* tether-cutting reconnection, arcade shearing, flux emergence or cancellation) allowed for a loss of equilibrium in the first stage, and a subsequent instability then triggered a catastrophic eruption in the second stage. In this case, the overlying magnetic field tension was strong enough, especially in the presence of the helmet streamer, to stall the eruption until the upward magnetic pressure rose to a critical point. Török, Kliem, and Titov (2004) studied the stability of the loop model by Titov and Démoulin (1999) in a twisted configuration that leads to an ideal kink instability; in this model the resulting vertical current sheet corresponds to the central element of the standard flare model. These authors noted that the surrounding potential field would prevent a global eruption unless the overlying field is sufficiently weakened or magnetic reconnection occurs in either the formed current sheet or with neighboring flux systems (requiring a multiple-loop configuration). For the event analyzed in this article, although the X-ray flare profile indicates a non-ideal process from the start, we may be seeing a similar form of halted eruption; after which a second stage to the eruption proceeds that is dominated by a kink-instability.

Aulanier *et al.* (2010) considered a non-ideal case *via* a flux cancellation model that transforms a sheared arcade into a slowly rising and stable flux rope. Eventually an altitude is reached at which the rate of overlying field decay exceeds the upward magnetic pressure and a torus instability causes a rapid eruption in the form of a CME (of inverse tear-drop shape as observed in this event). Furthermore, they described how part of the flux rope in the early stages rises faster, due to an asymmetric build-up of the pre-eruptive flux rope. Then the system enters a phase of fast expansion (accelerating to a velocity of  $400 \text{ km s}^{-1}$  in their model), where the overlying field lines that are not affected by the shearing motions correspond to the front of the CME commonly tracked in coronagraph observations. These modeled effects may explain the observations of the 8 March 2011 event, although without any measure of the decay index of the background magnetic field we cannot be certain that a torus instability is what we definitively observe.

## 5. Conclusions

The PROBA2/SWAP imager is unique in that it provides radially extended EUV observations of the Sun and low-corona to greater heights than other EUV imagers such as SDO/AIA. An ongoing goal in solar physics has been to study the connection between processes on the Sun and their effects elsewhere in the heliosphere; a connection known to lie predominantly between the regions of the photosphere, chromosphere and corona. Therefore, obtaining extensive observations across the solar atmosphere is paramount to understanding the physics at play. Since the LASCO/C1 coronagraph was lost very early during the SOHO mission, observations of the low corona have generally been quite

limited. In order to bridge this gap and garner some knowledge of the low-coronal initiation phase of CMEs, we have combined the SWAP observations with those of the ground-based Mk4 coronameter, to directly compare the EUV and white-light imagery. This was achieved through the use of advanced image processing techniques to overcome the low signal-to-noise ratio in these data and characterize the erupting structures of interest. The subsequent investigation of the dynamics of a specific case-study on 8 March 2011 provides insight into the early CME formation and eruption, using SWAP and Mk4 in tandem with the observations of AIA and LASCO. This complements the previous investigation of Su *et al.* (2012) who first reported on the two-stage flaring profile of the event as evidence for secondary heating. It is concluded that this event shows evidence for either a kink-unstable or torus-unstable flux rope that is first presumably formed by shearing motions and flux cancellation, before undergoing a magnetic pressure build-up and overlying tension decay that drives a subsequent fast CME and resulting streamer blowout.

Our study highlights the importance of multi-wavelength, high-cadence, extensive coverage observations of the low corona, where the physical mechanisms that underlie CME formation, initiation, and connection with flares and surrounding coronal field may best be investigated. It is hoped that future instruments such as the new *Coronal Solar Magnetism Observatory* (COSMO) K-coronagraph<sup>2</sup> will help to further advance our knowledge of these mechanisms.

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<sup>2</sup><http://www.cosmo.ucar.edu/kcoronagraph.html>

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