

# <sup>1</sup> cosmo-numba: B-modes and COSEBIs computations accelerated by Numba

<sup>3</sup> Axel Guinot  <sup>1</sup> and Rachel Mandelbaum  <sup>1</sup>

<sup>4</sup> 1 Department of Physics, McWilliams Center for Cosmology and Astrophysics, Carnegie Mellon  
<sup>5</sup> University, Pittsburgh, PA 15213, USA

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## Software

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## <sup>6</sup> Summary

<sup>7</sup> Weak gravitational lensing is a widely used probe in cosmological analysis. It allows astrophysicists <sup>8</sup> to understand the contend and evolution of the Universe. We are entering an era where we are <sup>9</sup> not limitied by the data volume but by systematics. It is in this context that we are presenting <sup>10</sup> here a simple python based software package to help in the computation of E-/B-mode <sup>11</sup> decomposition which can be use for systematic checks or science analysis. As we demonstrate <sup>12</sup> after, our implementation has both the high precision required and speed to perform this kind <sup>13</sup> of analysis.

## <sup>14</sup> Statement of need

The E-/B-mode composition for cosmic shear poses a significant computational challenge given <sup>16</sup> the need for high precision (required to integrate oscillatory functions over a large integration range and achieve accurate results) and speed. Cosmo-numba meets this need, facilitating the <sup>17</sup> computation of E-/B-modes decomposition using two methods. One of them is the Complete <sup>18</sup> Orthogonal Sets of E-/B-mode Integrals (COSEBIs) as presented in P. Schneider et al. ([2010](#)). <sup>19</sup> The COSEBIs rely on very high precision computation requiring more than 80 decimal places. <sup>20</sup> P. Schneider et al. ([2010](#)) propose an implementation using mathematica. cosmo-numba uses <sup>21</sup> a combination of sympy and mpmath to reach the required precision. This python version <sup>22</sup> enables an easier integration within cosmological inference pipelines, which are commonly <sup>23</sup> python-based, and facilitates the null tests.

<sup>25</sup> This software package also enables the computation of the pure-mode correlation functions <sup>26</sup> presented in Peter Schneider et al. ([2022](#)). Those integrals are less numerically challenging <sup>27</sup> than the COSEBIs, but having a fast computation is necessary for computing the covariance <sup>28</sup> matrix. One can also use those correlation functions for cosmological inference, in which case <sup>29</sup> the large number of calls to the likelihood function will also require a fast implementation.

## <sup>30</sup> Testing setup

<sup>31</sup> In the following two sections we will need fiducial shear-shear correlation functions,  $\xi_{\pm}(\theta)$ , <sup>32</sup> and power spectrum,  $P_{E/B}(\ell)$ . They have been computed using the Core Cosmology Library<sup>1</sup> <sup>33</sup> ([Chisari et al., 2019](#)) developed by the Dark Energy Science Collaboration.

## <sup>34</sup> COSEBIs

<sup>35</sup> The COSEBIs are defined as:

<sup>1</sup><https://github.com/LSSTDESC/CCL>

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$$E_n = \frac{1}{2} \int_0^\infty d\theta \theta [T_{n,+}(\theta) \xi_+(\theta) + T_{n,-}(\theta) \xi_-(\theta)], \quad (1)$$

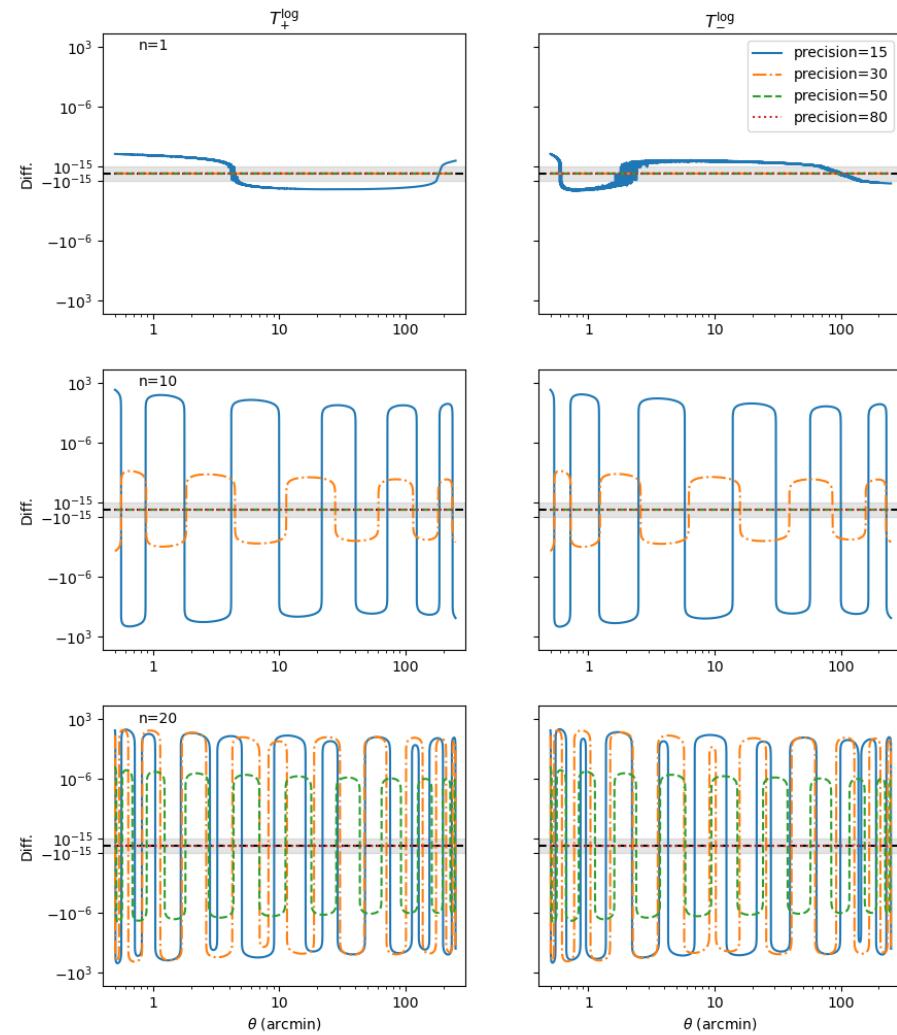
$$B_n = \frac{1}{2} \int_0^\infty d\theta \theta [T_{n,+}(\theta) \xi_+(\theta) - T_{n,-}(\theta) \xi_-(\theta)]; \quad (2)$$

37 where  $\xi_\pm(\theta)$  are the shear correlation functions, and  $T_{n,\pm}$  are the weight functions for the  
 38 COSEBI mode  $n$ . The complexity is in the computation of the weight functions. Cosmo-numba  
 39 carries out the computation of the weight functions in a logarithmic scale defined by:

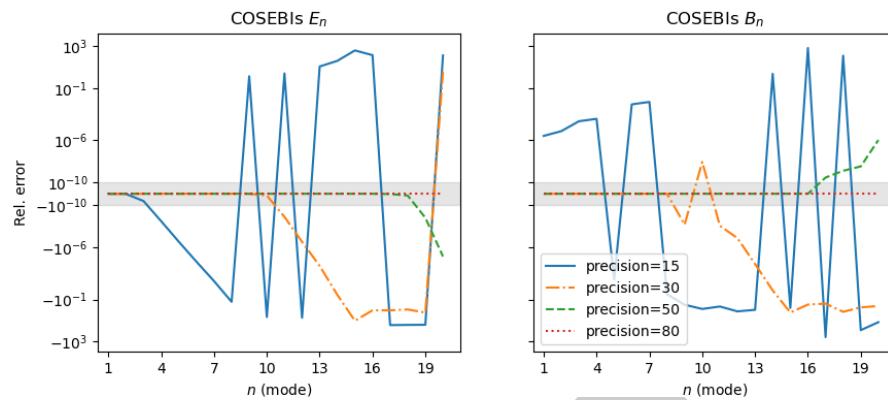
$$T_{n,+}^{\log}(\theta) = t_{n,+}^{\log}(z) = N_n \sum_{j=0}^{n+1} \bar{c}_{nj} z^j; \quad (3)$$

40 where  $z = \log(\theta/\theta_{\min})$ ,  $N_n$  is the normalization for the mode  $n$ , and  $\bar{c}_{nj}$  are defined iteratively  
 41 from Bessel functions (we refer the readers to P. Schneider et al. (2010) for more details).

42 We have validating our implementation against the original version in Mathematica from P.  
 43 Schneider et al. (2010). In Figure 1 we show the impact of the precision going from 15 decimal  
 44 places, which corresponds to the precision one could achieve using float64, up to 80 decimal  
 45 places, the precision used in the original Mathematica implementation. We can see that classic  
 46 float64 precision would not be sufficient, and with a precision of 80 our code exactly recovers  
 47 the results from the original implementation. Similarly, the impact on the COSEBIs is shown  
 48 in Figure 2.



**Figure 1:** In this figure we show the impact of the precision in the computation of the weight functions  $T_{\pm}^{\log}$ . For comparison, a precision of 15 corresponds to what would be achieved using numpy float64. The difference is computed with respect to the original Mathematica implementation presented in P. Schneider et al. (2010). The figure uses symlog, the shaded region represent the linear scale.



**Figure 2:** This figure shows the difference on the COSEBIs E- and B-mode relative to the original Mathematica implementation. We see that using only 15 decimal places would lead to several percent error making an implementation based on numpy float64 not suitable. The figure uses symlog, the shaded region represent the linear scale.

49 COSEBIs can also be defined from the power spectrum as:

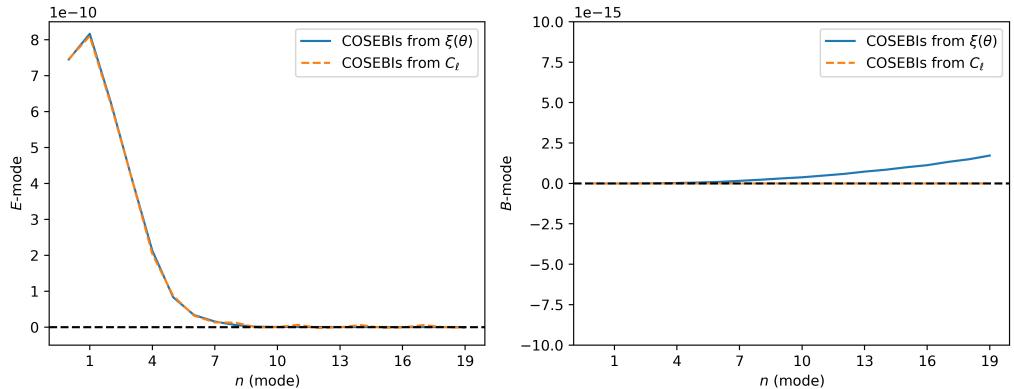
$$50 \quad E_n = \int_0^\infty \frac{d\ell \ell}{2\pi} P_E(\ell) W_n(\ell); \quad (4)$$

$$51 \quad B_n = \int_0^\infty \frac{d\ell \ell}{2\pi} P_B(\ell) W_n(\ell); \quad (5)$$

52 where  $P_{E/B}(\ell)$  is the power spectrum of E- and B-modes and  $W_n(\ell)$  are the filter functions which can be computed from  $T_{n,+}$  as:

$$53 \quad W_n(\ell) = \int_{\theta_{\min}}^{\theta_{\max}} d\theta \theta T_{n,+}(\theta) J_0(\ell\theta); \quad (6)$$

54 with  $J_0(\ell\theta)$  the 0-th order Bessel function. The Equation 6 is a Hankel transform of order 0.  
 55 It can be computed using the FFTLog algorithm presented in Hamilton (2000) implemented  
 56 here in Numba. Figure 3 shows the comparison between the COSEBIs computed from  $\xi_\pm(\theta)$   
 57 and from  $C_{E/B}(\ell)$ . We can see that the COSEBI E-modes agree very well but the B-modes are more stable when computed from the  $C(\ell)$  space.



**Figure 3:** Comparison of the COSEBIs E- and B-mode computed from  $\xi_\pm(\theta)$  and  $C_{E/B}(\ell)$ .

## 58 Pure-Mode Correlation Functions

59 In this section we describe the computation of the pure-mode correlation functions as defined  
 60 in Peter Schneider et al. (2022). There are defined as follow:

$$\xi_+^E(\vartheta) = \frac{1}{2} \left[ \xi_+(\vartheta) + \xi_-(\vartheta) + \int_{\vartheta}^{\vartheta_{\max}} \frac{d\theta}{\theta} \xi_-(\theta) \left( 4 - \frac{12\vartheta^2}{\theta^2} \right) \right] - \frac{1}{2} [S_+(\vartheta) + S_-(\vartheta)], \quad (7)$$

$$61 \quad \xi_+^B(\vartheta) = \frac{1}{2} \left[ \xi_+(\vartheta) - \xi_-(\vartheta) - \int_{\vartheta}^{\vartheta_{\max}} \frac{d\theta}{\theta} \xi_-(\theta) \left( 4 - \frac{12\vartheta^2}{\theta^2} \right) \right] - \frac{1}{2} [S_+(\vartheta) - S_-(\vartheta)], \quad (8)$$

$$62 \quad \xi_-^E(\vartheta) = \frac{1}{2} \left[ \xi_+(\vartheta) + \xi_-(\vartheta) + \int_{\vartheta_{\min}}^{\vartheta} \frac{d\theta \theta}{\vartheta^2} \xi_+(\theta) \left( 4 - \frac{12\theta^2}{\vartheta^2} \right) \right] - \frac{1}{2} [V_+(\vartheta) + V_-(\vartheta)], \quad (9)$$

$$63 \quad \xi_-^B(\vartheta) = \frac{1}{2} \left[ \xi_+(\vartheta) - \xi_-(\vartheta) + \int_{\vartheta_{\min}}^{\vartheta} \frac{d\theta \theta}{\vartheta^2} \xi_+(\theta) \left( 4 - \frac{12\theta^2}{\vartheta^2} \right) \right] - \frac{1}{2} [V_+(\vartheta) - V_-(\vartheta)]; \quad (10)$$

64 where  $\xi_{\pm}(\theta)$  correspond to the shear-shear correlation function. The functions  $S_{\pm}(\theta)$  and  $V_{\pm}(\theta)$   
 65 are themselves defined by integrals and we refer the reader to Peter Schneider et al. (2022)  
 66 for more details about their definition. By contrast with the computation of the COSEBIs,  
 67 these integrals are more stable and straightforward to compute but still require some level of  
 68 precision. This is why we are using the qags method from the QUADPACK<sup>2</sup> (Piessens et al.,  
 69 2012) with a 5-th order spline interpolation. In addition, as one can see from the equations  
 70 above, the implementation requires a loop over a range of  $\vartheta$  values. This is why having a fast  
 71 implementation will be required if one want to use those correlation functions in cosmological  
 72 inference.

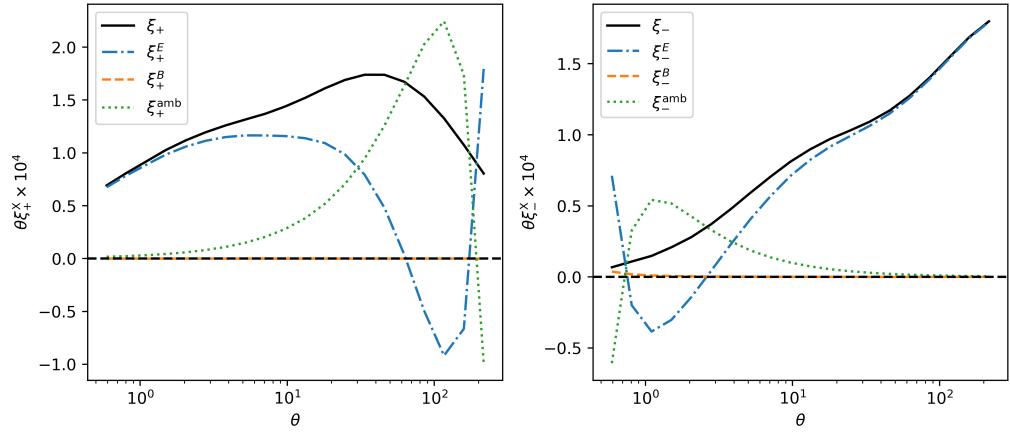


Figure 4: This figure shows the decomposition of the shear-shear correaltion functions in E- and B-modes (and ambiguous mode).

## 72 Acknowledgements

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 74 Investigator in Astrophysics, Award ID 620789).

<sup>2</sup>We use the C version of the library wrapped to python using Numba: <https://github.com/Nicholaswogan/NumbaQuadpack>

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