

# FEEDBACK from the NGC7538 H<sub>II</sub> region

H. Beuther<sup>1</sup>, N. Schneider<sup>2</sup>, R. Simon<sup>2</sup>, S. Suri<sup>1,3</sup>, V. Ossenkopf-Okada<sup>2</sup>, S. Kabanovic<sup>2</sup>, M. Röllig<sup>2</sup>, C. Guevara<sup>2</sup>, A.G.G.M. Tielens<sup>4,5</sup>, G. Sandell<sup>6</sup>, C. Buchbender<sup>2</sup>, O. Ricken<sup>7</sup>, and R. Güsten<sup>7</sup>

<sup>1</sup> Max Planck Institute for Astronomy, Königstuhl 17, 69117 Heidelberg, Germany, e-mail: beuther@mpia.de

<sup>2</sup> I. Physik. Institut, University of Cologne, Zülpicher Str. 77, D-50937 Cologne, Germany

<sup>3</sup> University of Vienna, Department of Astrophysics, Türkenschanzstrasse 17, 1180 Vienna, Austria

<sup>4</sup> Department of Astronomy, University of Maryland, College Park, MD 20742, USA

<sup>5</sup> Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands

<sup>6</sup> Institute for Astronomy, University of Hawai‘i at Manoa, Hilo, 640 N. Aohoku Place, Hilo, HI 96720, USA

<sup>7</sup> Max Planck Institute for Radioastronomy, Auf dem Hügel 69, 53121 Bonn, Germany

Version of January 17, 2022

## ABSTRACT

**Context.** How do expanding H<sub>II</sub> regions interact with their environmental cloud? This is one of the central questions driving the SOFIA legacy program FEEDBACK. Here, we present a case study toward the prototypical H<sub>II</sub> region NGC7538.

**Aims.** We want to understand the interaction of the NGC7538 H<sub>II</sub> region with the neighboring molecular cloud that hosts several active star-forming regions.

**Methods.** With the Stratospheric Observatory for Infrared Astronomy (SOFIA) we mapped an area of  $\sim 210''^2$  ( $\sim 125 \text{ pc}^2$ ) around NGC7538 in the velocity-resolved ionized carbon fine-structure line [CII] at 1.9 THz (158  $\mu\text{m}$ ). Complementary observed atomic carbon [CI] at 492 GHz and high-J CO(8–7) data as well as archival near-/far-infrared, cm continuum, CO(3–2) and HI data are folded into the analysis.

**Results.** The ionized carbon [CII] data reveal rich morphological and kinematic structures. While the overall morphology follows the general ionized gas that is also visible in the radio continuum emission, the channel maps show multiple bubble-like structures with sizes on the order of  $\sim 80\text{--}100''$  ( $\sim 1.0\text{--}1.28 \text{ pc}$ ). While at least one of them may be an individual feedback bubble driven by the main exciting sources of the NGC7538 H<sub>II</sub> region (the O3 and O9 stars IRS6 & IRS5), the other bubble-like morphologies may also be due to the intrinsically porous structure of the H<sub>II</sub> region. An analysis of the expansion velocities around  $10 \text{ km s}^{-1}$  indicates that thermal expansion is not sufficient but that wind-driving from the central O-stars is required. The region exhibits a general velocity gradient across, but we also identify several individual velocity components. The most blue-shifted [CII] component has barely any molecular or atomic counterparts. At the interface to the molecular cloud, we find a typical photon-dominated region (PDR) with a bar-shape. Ionized C<sup>+</sup>, atomic C<sup>0</sup> and molecular carbon CO show a layered structure in this PDR. The carbon in the PDR is dominated by its ionized C<sup>+</sup> form with atomic C<sup>0</sup> and molecular CO masses of  $\sim 0.45 \pm 0.1 \text{ M}_\odot$  and  $\sim 1.2 \pm 0.1 \text{ M}_\odot$ , respectively, compared to the ionized carbon C<sup>+</sup> in the range of  $3.6\text{--}9.7 \text{ M}_\odot$ . This bar-shaped PDR exhibits a velocity-gradient across, indicating motions along the line-of-sight towards the observer.

**Conclusions.** Even if dominated by two nearby exciting sources (IRS6 & IRS5), the NGC7538 H<sub>II</sub> region exhibits a diverse set of sub-structures that interact with each other as well as with the adjacent cloud. Compared to other recent [CII] observations of H<sub>II</sub> regions (e.g., Orion Veil, RCW120, RCW49), bubble-shape morphologies revealed in [CII] emission, indicative of expanding shells, are recurring structures of PDRs.

**Key words.** Stars: formation – ISM: clouds – ISM: kinematics and dynamic – ISM: bubbles – ISM: HII regions – ISM: individual objects: NGC7538

## 1. Introduction

Stellar feedback in the form of radiation, winds and supernova explosions can have positive as well as negative impact on their environment. Positive feedback here is defined as inducing new star formation processes whereas negative feedback is meant as a destruction process of the natal cloud and by that limiting further star formation. Feedback is important for the determination of the physical processes within individual clouds but also crucial in an integral sense for whole galaxies because it likely determines the general star formation efficiency (e.g., Matzner 2002; Elmegreen 2011; Hopkins et al. 2014; Geen et al. 2016, 2018; Kim et al. 2018).

With the goal to understand the impact of feedback processes onto the environment for individual resolved cases, the SOFIA

(Stratospheric Observatory for Infrared Astronomy) legacy program FEEDBACK<sup>1</sup> targets 11 galactic high-mass star-forming regions and H<sub>II</sub> regions in the ionized atomic carbon fine-structure line [CII] at 158  $\mu\text{m}$  (1.9 THz) and the atomic oxygen line [OI] at 63  $\mu\text{m}$  (4.7 THz). The ionized carbon line [CII] is of particular importance because it allows us to study the kinematics of the ionized gas as well as the interfaces with the molecular clouds. The [CII] line is a direct tracer of bubble/stellar feedback kinematics (e.g., Pabst et al. 2020). Furthermore, the [CII] line is one of the dominant cooling lines in the interstellar medium at low to intermediate densities and UV fields (e.g., Hollenbach & Tielens 1997; Röllig & Ossenkopf 2013). The FEEDBACK program started in spring 2019 and is still ongoing. More de-

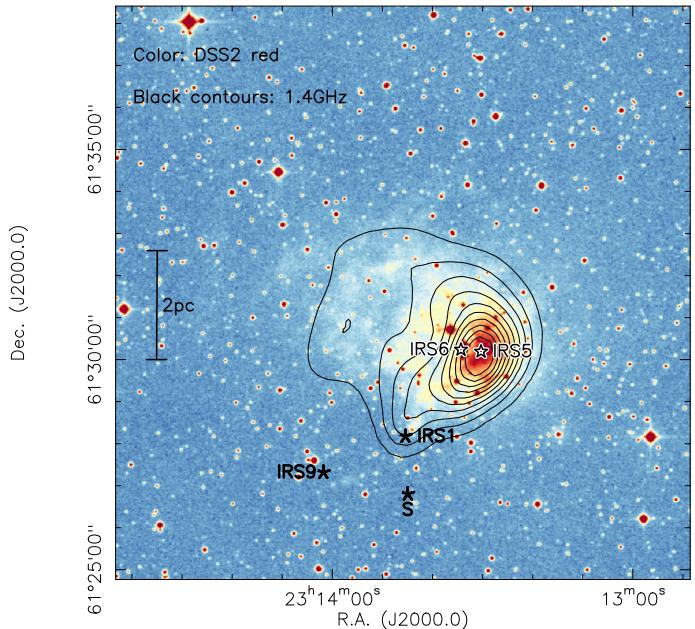
<sup>1</sup> <http://feedback.astro.umd.edu>

tails about this SOFIA legacy program are provided in Schneider et al. (2020). One outcome of the FEEDBACK project so far is the detection of expanding shells in [CII] emission in various bubble-shaped or bipolar H<sub>II</sub> regions, i.e., in RCW120 (Luisi et al. 2021, Kabanovic et al. in press), RCW49 (Tiwari et al. 2021), RCW36 (Bonne et al., subm.), and RCW79 (Zavagno et al., in prep.). These shells were first detected in Orion A (Pabst et al. 2019) and appear to be a ubiquitous phenomena. While the regions RCW49 and RCW120 reveal second generation star formation within these shells (Tiwari et al. 2021; Luisi et al. 2021), the expanding veil nebula in Orion shows little evidence for dense clumps that could lead to star formation (Pabst et al. 2019; Goicoechea et al. 2020).

While an analysis of the whole sample will follow when all data are taken, in the following, we concentrate on the prototypical H<sub>II</sub> region NGC7538 shown at optical and cm continuum wavelengths in Figure 1. The region is located at a distance of 2.65 kpc (Moscadelli et al. 2009) and has a physical extent of several parsec. The two brightest sources exciting the H<sub>II</sub> region are IRS5 and IRS6 with spectral types of O9 and O3, respectively (Puga et al. 2010, marked in Fig. 1). The H<sub>II</sub> region is associated with several sites of active star formation, in particular in the south with the regions IRS1, S and IRS9 (marked in Fig. 1). Towards these southern star-forming regions, the H<sub>II</sub> region appears to be sharply bounded, whereas in the northeast the region shows diffuse emission extending well beyond the photon-dominated region (PDR) shell (Luisi et al. 2016). In a recent study combining CO(3–2) and [CII] data, Sandell et al. (2020) characterized a large north-south outflow emanating from the southern star-forming core IRS1 with blue-shifted ionized carbon [CII] and molecular CO emission north of IRS1 in the vicinity of the H<sub>II</sub> region. Furthermore, Townsley et al. (2018) show that diffuse X-ray emission is found toward the H<sub>II</sub> region, spreading even beyond the extend of the H<sub>II</sub> region as determined by, e.g., the cm continuum emission.

More generally speaking, the NGC7538 H<sub>II</sub> region is embedded in a larger molecular cloud complex (e.g., Ungerechts et al. 2000; Fallscheer et al. 2013). Figure 2 gives an overview of various observed tracers. The cm continuum and 8 $\mu$ m emission appears to stem dominantly from the H<sub>II</sub> region. One should keep in mind that 8 $\mu$ m emission is usually emitted in the cavity walls of the H<sub>II</sub> region where the poly-cyclic aromatic hydrocarbons (PAH) survive. In contrast to these features, the dense gas as traced by the Herschel dust continuum or CO(3–2) emission is found more in the outskirts of the H<sub>II</sub> region, in particular towards the south where the well-known star-forming regions IRS1, S and IRS9 are located. As will be discussed in more detail below, the new SOFIA [CII] data combine different worlds: They trace the PDR, and they also show strong emission features toward the active star-forming region IRS1 and a bar-like region at the south-eastern interface of the H<sub>II</sub> region with the dense molecular cloud. Hence, the velocity-resolved [CII] data are the ideal probe to study the feedback processes and the kinematic gas properties from the evolving H<sub>II</sub> region on its neighboring molecular cloud.

Specific topics we address with this study are the impact from the expanding H<sub>II</sub> region onto the environmental dense gas. Can we identify expanding [CII] shells or other dynamic features, either triggering or preventing new star formation processes? Do we identify layered structures from a photon-dominated region (PDR)? What is the carbon budget between ionized C<sup>+</sup>, atomic C<sup>0</sup> and molecular CO gas specifically at the H<sub>II</sub> region/molecular cloud interface?



**Fig. 1.** Optical and cm continuum data of the NGC7538 region, encompassing the field of view of our observations. The color-scale shows the optical DSS2 (Digital Sky Survey v2) obtained with the red filter. The contours outline the 1.4 GHz cm continuum emission at 1' angular resolution from the Canadian Galactic Plane Survey (CGPS, Taylor et al. 2003). The two stars in the center of the H<sub>II</sub> region mark the positions of IRS5 and IRS6 from Puga et al. (2010). The three 5-pointed stars directly below the H<sub>II</sub> regions show the positions of the active sites of star formation IRS1, S and IRS9.

## 2. Observations and data

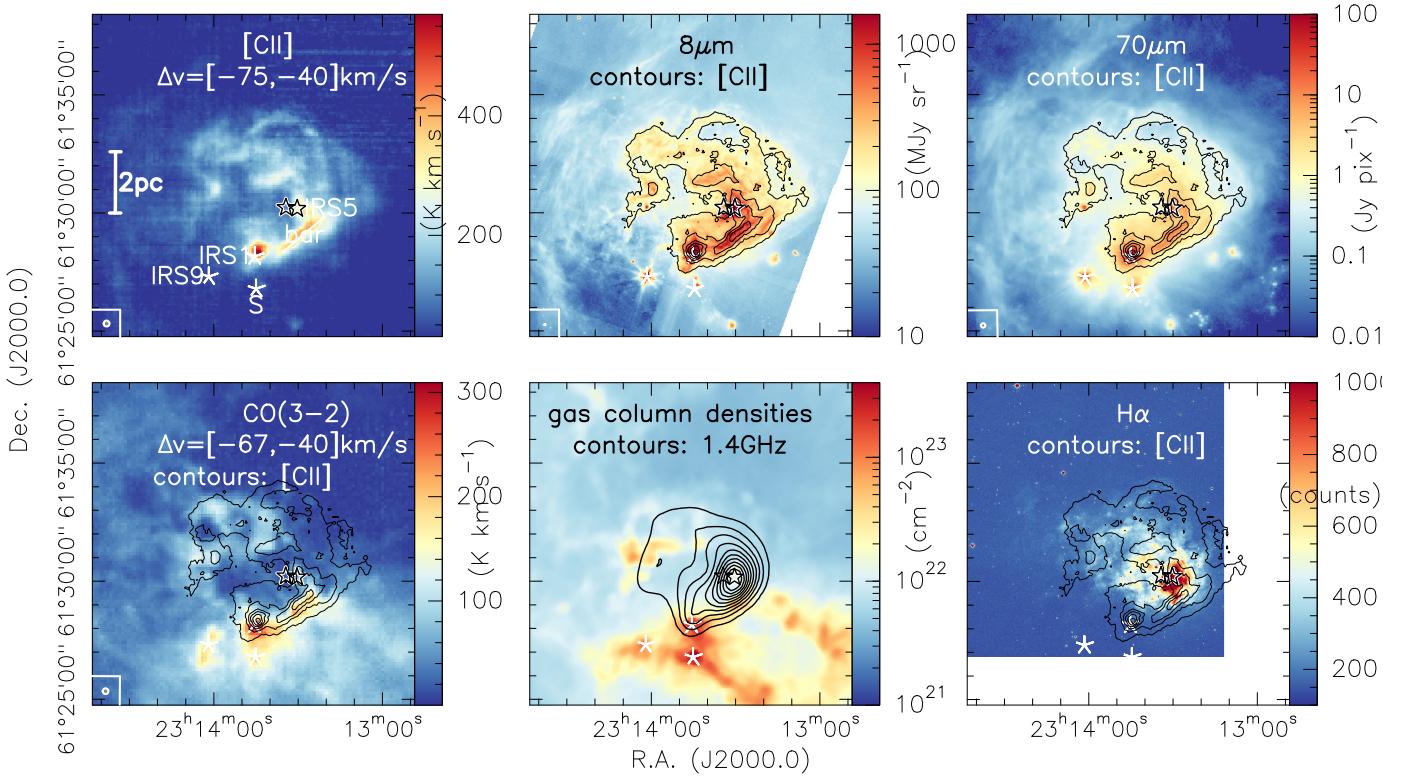
### 2.1. SOFIA observations

The region was mapped with the dual-frequency array heterodyne receiver upGREAT<sup>2</sup> (Risacher et al. 2018) in the velocity-resolved [CII] fine-structure line at 158  $\mu$ m/1.9 THz<sup>3</sup>. The [OI] 63  $\mu$ m line was observed in parallel. Fast Fourier Transform Spectrometers (FFTS) with 4 GHz instantaneous bandwidth (Klein et al. 2012) resulted in a nominal spectral resolution for the [CII] line of 0.244 MHz or 0.04 km s<sup>-1</sup>. The observations were done on two consecutive days during Cycle 7, on December 12th and 13th 2019, with an average  $T_{\text{sys}}$  single side-band (SSB) of 4486 and 4580 K and an average precipitable water vapor of 6.2 and 3.6  $\mu$ m respectively per day. The data were calibrated to the main beam brightness temperature intensity scale,  $T_{\text{mb}}$ , with the *kalibrate* task (Guan et al. 2012), part of the standard GREAT pipeline, with a forward efficiency of 0.97 and an average beam efficiency of 0.66.

Mapping was conducted in an on-the-fly (OTF) mode where the entire NGC7538 map was split into 4 tiles of roughly 436'' on each side. During scanning the data were better than Nyquist-sampled with a dump at every 5.2''. The final map size is 14.5'  $\times$  14.5' = 11.2 pc  $\times$  11.2 pc. Each tile should typically be observed 4 times with different scanning angles to reduce potential striping effects. In the case of NGC7538, three tiles were mapped to a deeper sensitivity whereas for the fourth north-western tile not all coverages were conducted so far. This results

<sup>2</sup> German Receiver for Astronomy at Terahertz. (up)GREAT is a development by the MPI for Radioastronomy and the KOSMA/University of Cologne, in cooperation with the DLR Institute for Sensorsystems.

<sup>3</sup> The CII 158  $\mu$ m data are provided at the IRSA/IPAC Infrared science archive <https://irsa.ipac.caltech.edu/Missions/sofia.html>.



**Fig. 2.** Integrated intensity molecular and atomic line maps as well as continuum emission are presented in color-scale toward NGC7538. The specific lines, integration ranges and continuum bands are labeled in each panel. The corresponding angular resolution is shown at the bottom-left of the panels. The contours in most panels as labeled show the [CII] emission at only a few selected contour levels to highlight similarities and differences. The contours in the bottom-middle panel outline the 1.4 GHz cm continuum emission from the CGPS. The three five-pointed stars mark the positions of the embedded sources NGC7538IRS1, NGC7538S and NGC7538IRS9. The two stars in the middle show the positions of the two main exciting sources of the HII region NGC7538IRS5 and NGC7538IRS6. The bar-like PDR as well as a linear scale-bar are presented in the top-left panel.

in slightly increased noise in the top-right tile of the map and residual striping quite apparent in Fig. 2. The data cubes used here have been resampled at a velocity resolution of  $0.5 \text{ km s}^{-1}$  to increase the signal-to-noise ratio. Our final velocity-resampled and gridded data-cube has an angular resolution of  $15.8''$  (corresponding to  $\sim 0.2 \text{ pc}$  linear spatial resolution). The  $1\sigma$  rms over the well-covered  $3/4$  of the map are  $\sim 0.9 \text{ K}$  per  $0.5 \text{ km s}^{-1}$  channel on a  $T_{mb}$  scale. The higher  $1\sigma$  rms in the north-western tile is  $\sim 1.9 \text{ K}$  per  $0.5 \text{ km s}^{-1}$  channel.

Furthermore, slightly smaller maps of the region were observed in the atomic carbon fine-structure line [CI] at 492 GHz and three CO transitions with  $J = 8 - 7$ ,  $J = 11 - 10$  and  $J = 22 - 21$ . Here, we are concentrating on the lower-density tracer [CI] as well as the CO(8–7) transition at 921.7997 GHz. The [OI]  $63 \mu\text{m}$  and higher- $J$  CO lines, tracing higher densities and temperatures, will be analyzed in forthcoming studies. The angular resolution of the final data cubes for the [CI] and CO(8–7) lines are  $63''$  and  $42''$ , respectively. The  $1\sigma$  rms values for the [CI] and CO(8–7) lines at  $0.5 \text{ km s}^{-1}$  are 0.7 and 1.4 K, respectively.

## 2.2. Complementary archival data

### 2.2.1. Optical, mid-infrared, far-infrared and radio data

We used the Digital Sky Survey version 2 (DSS2) to obtain an optical image of the region in the red filter. The DSS2 digital images are based on photographic images obtained using the Oschin Schmidt Telescope of Palomar Mountain and the UK

Schmidt Telescope. The plates were processed into the present compressed digital form with the permission of the institutions.

The 1.4 GHz radio continuum and the 21 cm HI data are from the Canadian Galactic Plane Survey (CGPS, Taylor et al. 2003). The angular resolution of these data products is  $\sim 1'$ . The  $1\sigma$  rms of the continuum and line data are  $\sim 0.5 \text{ K}$  and  $3 \text{ K}$  in  $0.82 \text{ km s}^{-1}$  channels, respectively.

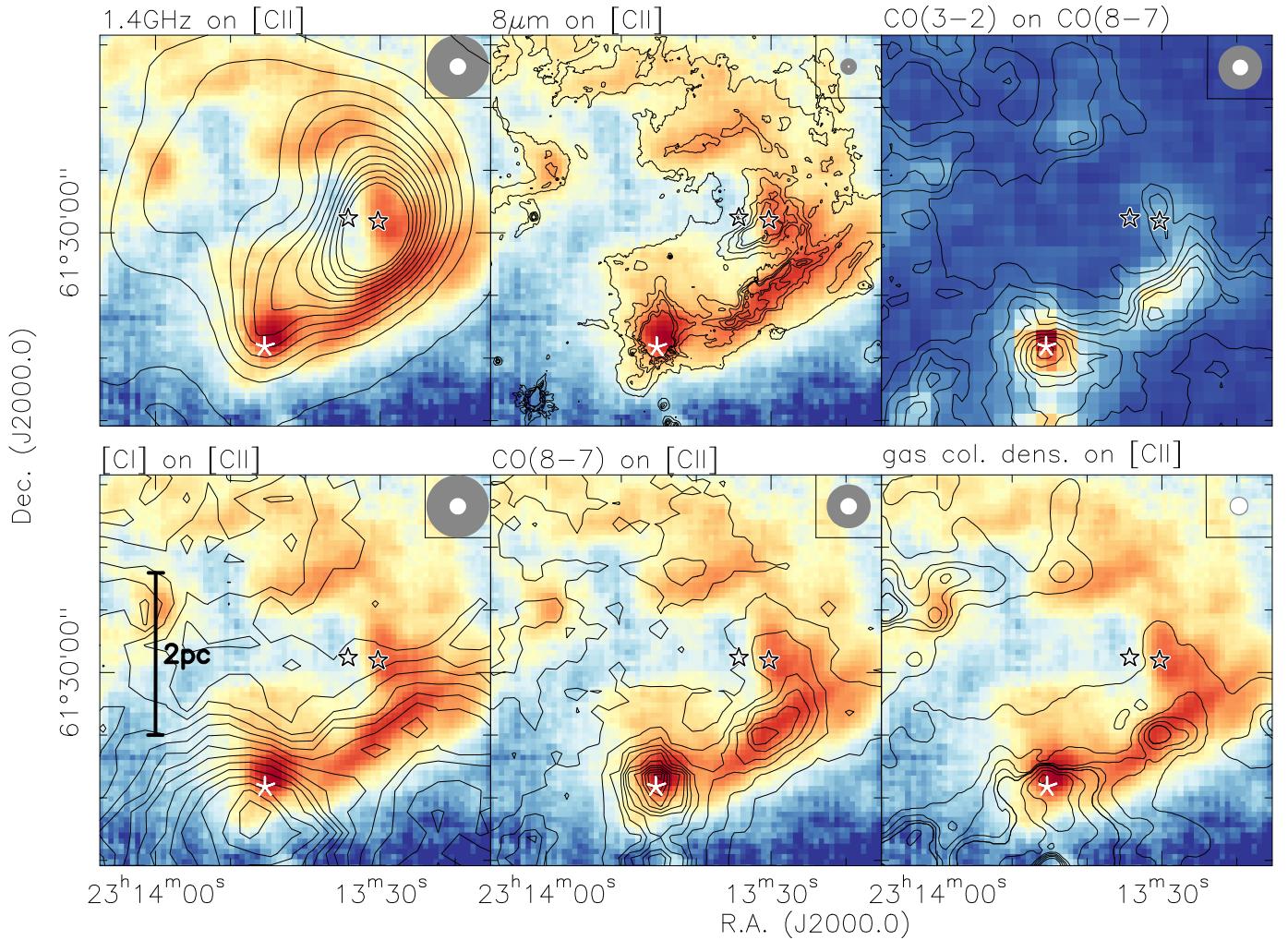
The  $8 \mu\text{m}$  photometric emission mid-infrared data are obtained with the IRAC camera (Fazio et al. 2004) on board of the Spitzer satellite (Werner et al. 2004) at an angular resolution of  $\sim 2''$ .

The  $70 \mu\text{m}$  far-infrared data are observed with the PACS camera (Poglitsch et al. 2010) on board of the Herschel satellite mission (Pilbratt et al. 2010) in the framework of the HOBYS key program (Motte et al. 2010). The angular resolution is  $\sim 5.6''$ .

The  $H_\alpha$  data were observed in the framework of “The INT Photometric  $H_\alpha$  Survey of the Northern Galactic Plane” (IPHAS) at an angular resolution of  $\sim 1.1''$  (Drew et al. 2005).

### 2.2.2. Herschel column density and CO(3–2) data

The Herschel gas column density map we use in this paper is a higher angular resolution version ( $18''$ ) of the one presented in Fallscheer et al. (2013). The map was produced by the HOBYS (Herschel imaging survey of OB Young Stellar objects, Motte et al. 2010) consortium and will soon be publicly provided. The column density map was determined following the HOBYS procedure, i.e., from a pixel-to-pixel spectral energy distribution



**Fig. 3.** Overlays of selected pairs of emission tracers as described above each panel. The integration ranges for the atomic and molecular lines are the same as in Fig. 2. The contour levels are chosen individually for each figure to highlight the emission best. The angular resolution elements for both tracers are presented in the top-right of each panel. The five-pointed stars always mark the position of the embedded source NGC7538IRS1, and the two open stars show the positions of the two main exciting sources of the Hii region NGC7538IRS5 and NGC7538IRS6. A linear scale-bar is presented in the bottom-left panel.

(SED) greybody fit to the 160, 250, 350, and 500  $\mu\text{m}$  wavelengths observations. There were no 100  $\mu\text{m}$  maps available and the 70  $\mu\text{m}$  data are not included because the interest lies in the distribution of the cold molecular gas. For the SED fit, we fixed the specific dust opacity per unit mass (dust+gas) approximated by the power law  $\kappa_v = 0.1(v/1000\text{GHz})^{\beta_d} \text{cm}^2/\text{g}$  with  $\beta_d=2$ , and left the dust temperature and column density (atomic and molecular gas, in the following we refer to that as gas column density) as free parameters (see, e.g., Russeil et al. 2013 for more details). The procedure explaining how high angular resolution maps were obtained is described in detail in Appendix A of Palmeirim et al. (2013). The concept is to employ a multi-scale decomposition of the flux maps and assume a constant line-of-sight temperature. The final map at 18'' angular resolution is constructed from the difference maps of the convolved SPIRE maps (at 500, 350, and 250  $\mu\text{m}$ ) and the temperature information from the color temperature derived from the 160 to 250  $\mu\text{m}$  ratio.

The CO(3-2) data were obtained with the James Clark Maxwell Telescope (JCMT) as part of a large map of the NGC7538 region first presented in Fallscheer et al. (2013) and

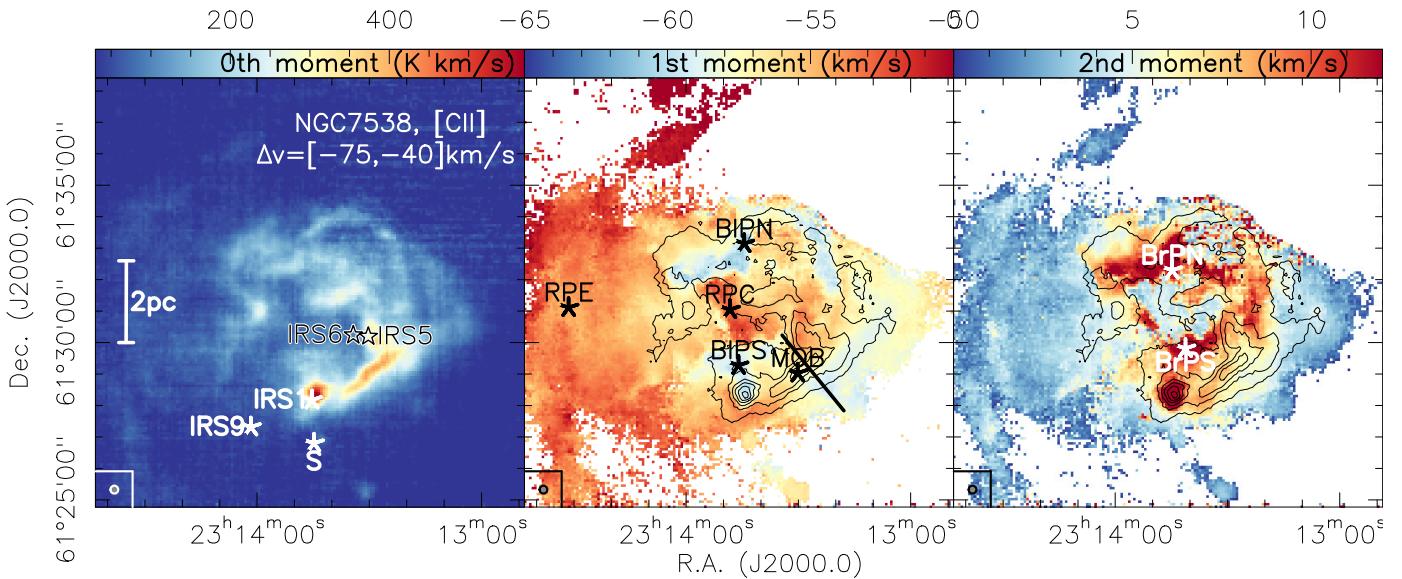
later also in Sandell et al. (2020). Following Fallscheer et al. (2013), the typical  $1\sigma$  rms in 0.42  $\text{km s}^{-1}$  channels is 0.6 K on a  $T_A^*$  scale. Using a beam efficiency of 0.64<sup>4</sup>, the data are converted to main beam brightness temperature for the quantitative analysis. The beam size of these data is  $\sim 15''$ .

## 3. Results

### 3.1. Spatial structures

As indicated in the Introduction, the [CII] emission in the NGC7538 complex traces the large-scale PDR as well as the PDR associated with denser star-forming molecular gas at the edge of the Hii region. To be more specific, Fig. 2 shows the integrated emission of several tracers, and in particular the morphology of the [CII], the 8  $\mu\text{m}$  and 70  $\mu\text{m}$  emission is very similar. The ionized gas as traced by the H<sub>a</sub> and 1.4 GHz emission is confined to the real Hii region, and the [CII], 8  $\mu\text{m}$  and 70  $\mu\text{m}$  emission partly wrap around that (bottom-right panel of Fig. 2). While the [CII] line can form in PDRs as well as molecular

<sup>4</sup> <https://www.eaoobservatory.org/jcmt/instrumentation/heterodyne/harp/>



**Fig. 4.** Moment maps of the [CII] emission. The left, middle and right panels present in color scale the integrated emission and the 1st and 2nd moment maps (intensity-weighted peak velocities and intensity-weighted velocity dispersion) in the velocity range from  $-75$  to  $-40\text{ km s}^{-1}$ , respectively. The noisier top-right quarters of the 1st and 2nd moment maps are blanked (see Sect. 2). The contours in the middle and right panels show the integrated [CII] emission from the left panel at only a few selected contour levels to highlight similarities and differences. The three five-pointed stars in the left panel mark the positions of the embedded sources NGC7538IRS1, NGC7538S and NGC7538IRS9. The two stars show the positions of the two main exciting sources of the HII region NGC7538IRS5 and NGC7538IRS6. A linear scale-bar is presented in the left panel. The markers in the middle and right panels highlight additional positions (see main text) toward where we extracted spectra in Fig. 5. The line in the middle panel outlines the pv-cut from IRS6 through the bar-like structure presented in Fig. 13.

media, diffuse  $8\mu\text{m}$  emission typically stems from UV-pumped infrared fluorescence from Polycyclic Aromatic Hydrocarbon (PAH) molecules within photon-dominated regions (PDRs). The dense gas tracers (gas column densities from dust continuum and CO(3–2), bottom panels in Fig. 2) emit dominantly in the south and south-west of the region. We note that toward the largest part of the [CII] emission the CO(3–2) line does not show significant self-absorption features (see spectra discussed below). The associated emission toward the young embedded star-forming region IRS1 as well as the bar-like PDR features are visible in all tracers shown in Fig. 2. In addition to this, some dense gas structures can also be found in the north-eastern part of NGC7538.

To show spatial similarities and differences in a bit more detail, Figure 3 presents a compilation of different pairs of tracers always as contour maps overlaid on color-maps. We added here the higher excited CO(8–7) line ( $E_u/k = 199\text{ K}$  compared to  $E_u/k = 33.2\text{ K}$  for the (3–2) transition) as well as the atomic carbon fine structure line [CI] at 492 GHz. The top-left and top-middle panels confirm that the [CII] and  $8\mu\text{m}$  emission agree very well and appear to at least partially wrap around the ionized gas (see also bottom-right panel of Fig. 2). A comparison of the two CO transitions in the top-right panel shows that in the south-western bar-like emission structure, the higher excited CO(8–7) transition emits closer to the two exciting sources at the center of the HII region than the lower excited CO(3–2) transition. Similar layered structures within the bar can be identified between atomic and ionized carbon ([CI] and [CII], bottom-left panel of Fig. 3) as well as between the CO(8–7) and [CI] maps. In comparison to those layered structures, spatial morphologies of the higher excited CO(8–7) and ionized carbon [CII] lines agree much better in the bar. This is a typical PDR layered structure in [CII]/[CI]/CO as well as in CO(8–7) and (3–2). We will come back to this in section 4.2.

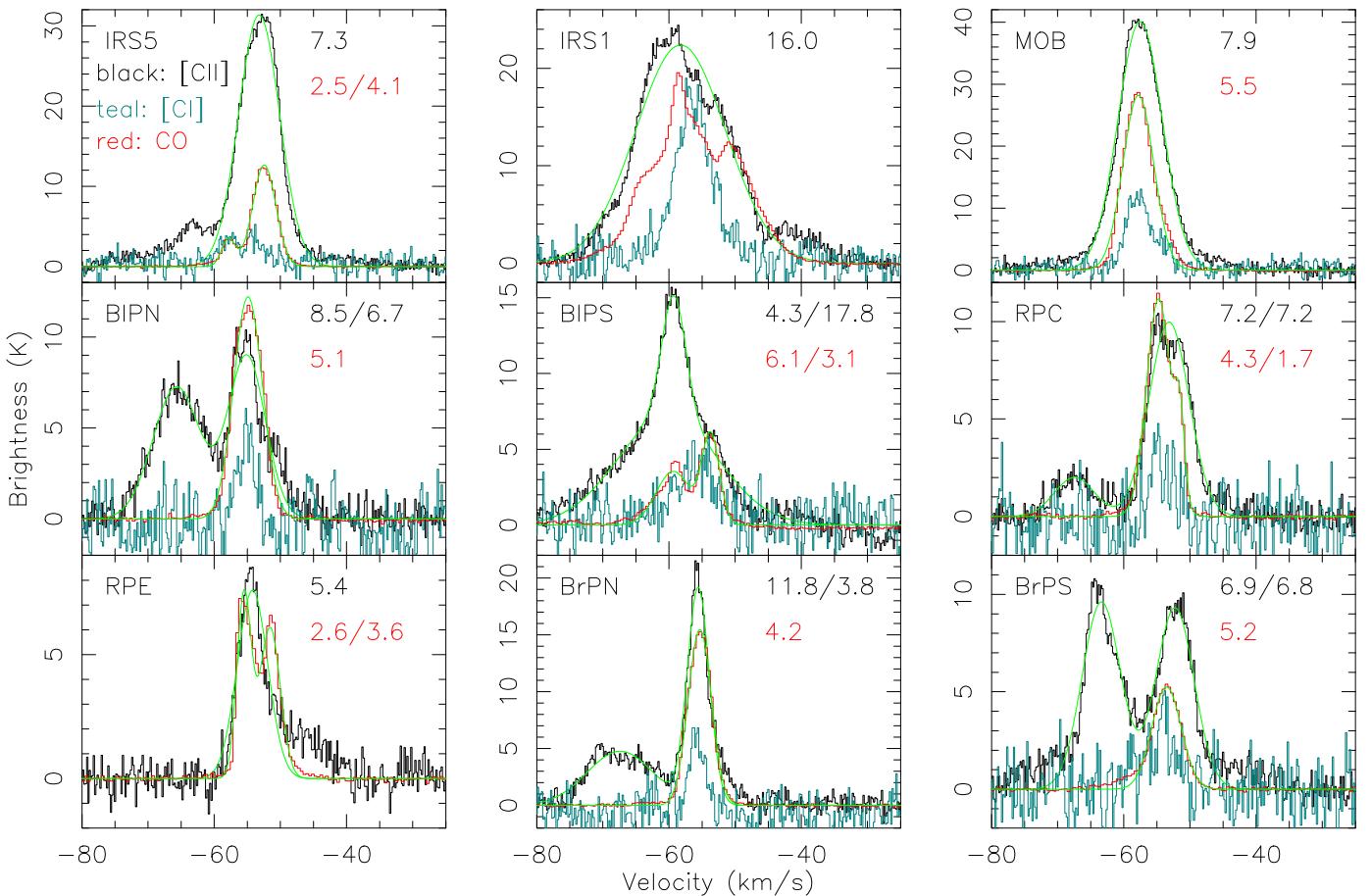
### 3.2. Kinematic structures

To get a first proxy of the dynamics, Fig. 4 presents the integrated [CII] emission as well as the 1st and 2nd moment maps of the [CII] data (intensity-weighted peak velocities and velocity dispersion). The 1st moment map (middle panel of Fig. 4) reveals a general velocity gradient approximately from east to west over the entire extend of the HII region. On top of the overall velocity gradient, one can identify several blue-shifted emission regions in the middle of the map. In comparison to that, the 2nd moment or velocity dispersion map shows over large parts of the map rather uniformly low values below  $5\text{ km s}^{-1}$ . However, on top of that, one can clearly identify an almost ring-like structure with larger velocity dispersion on the order of  $10\text{ km s}^{-1}$ . As will be discussed below, these regions of apparent large velocity dispersion are largely caused by multiple velocity components in the [CII] spectra.

For a more detailed analysis of the kinematics, it is crucial to look at individual spectra, and we selected several positions in the region for further analysis. We labeled them as follows, and they are all marked in Fig. 4:

- IRS5: one of the main exciting source of the HII region
- IRS1: the densest star-forming core
- MOB: a position in the middle of the bar-like PDR
- BIPN: blue peak in the north in the 1st moment map
- BIPS: blue peak in the south of the 1st moment map
- RPC: red peak in the center of the 1st moment map
- RPE: red peak in the east of the 1st moment map
- BrPN: broad emission in the north in the 2nd moment map
- BrPS: broad emission in the south in the 2nd moment map

Figure 5 presents the corresponding [CII] spectra as well as the molecular and atomic counterparts in the CO(3–2) and



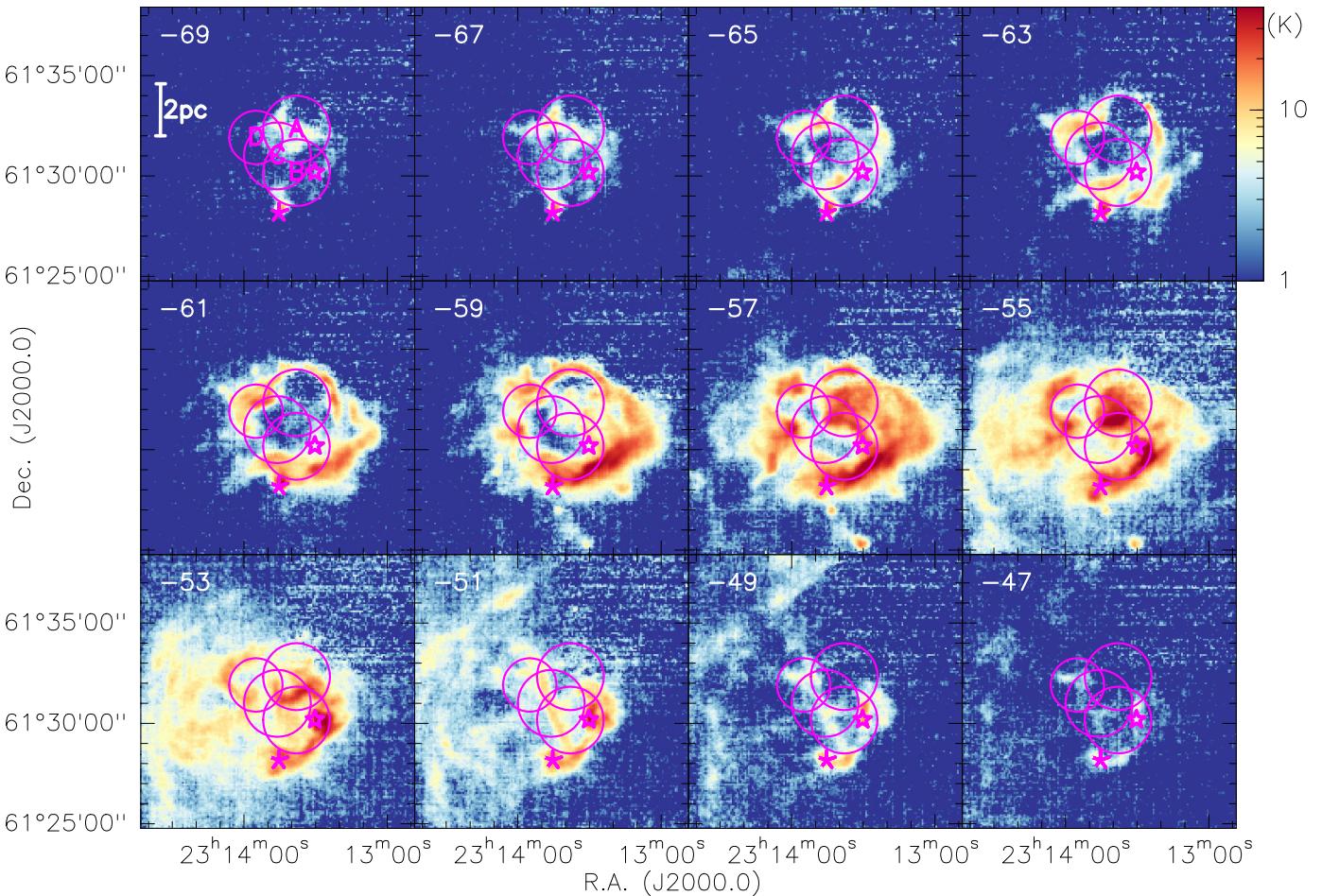
**Fig. 5.** [CII] (black), CO(3–2) (red) and [CI] (teal, multiplied by 2) spectra towards the positions marked in Fig. 4. The green lines show Gaussian fits to the [CII] and CO(3–2) data. The numbers present the corresponding FWHM values (“\*/” labels two components). The positions are labeled in each panel. The RPE position to the east is not covered by the [CI] map.

[CI] lines<sup>5</sup>. Where possible we fitted Gaussians to the [CII] and CO(3–2) spectra, at several positions also two components. While self-absorption may explain smaller dips in some spectra, the two velocity components fitted here are typically that far separated in velocity space that they should indeed be separate components. The corresponding full width half maximum values (FWHM) in  $\text{km s}^{-1}$  are also shown in Fig. 5. Interestingly, the only almost simple, single-Gaussian spectrum is identified toward the position in the middle of the bar (MOB, top-right panel in Fig. 5). We will get back to that structure in section 4.2. The spectra toward the infrared sources IRS5 and IRS1 as well as the red-shifted peak in the east of the region (RPE) show some velocity structure, but no particularly prominent features.

This is very different in the remaining five spectra towards the other blue- and red-shifted positions in the center of the maps (BIPN, BIPS, RPC) as well as towards the positions with particularly broad velocity dispersion (BrPN, BrPS, Figs. 4 and 5). While the main emission from the NGC7538 region is typically in the velocity range with respect to the local standard of rest ( $v_{\text{LSR}}$ ) between  $-60$  and  $-50 \text{ km s}^{-1}$ , these positions show additional strong emission at even more negative velocities beyond  $-60 \text{ km s}^{-1}$ . In particular, the spectra toward BIPN, RPC, BrPN and BrPS exhibit a second, clearly separated velocity component in the [CII] emission that has barely any detectable counterpart in the molecular emission of the CO(3–2) line or the atomic

[CI] emission. The rms values for these two lines (section 2) correspond at 30 K to  $3\sigma$  column density sensitivities (see section 4.3 for details on the calculations) of  $N(\text{CO}) \approx 2.1 \times 10^{16} \text{ cm}^{-2}$  and  $N(\text{C}^0) \approx 2.7 \times 10^{16} \text{ cm}^{-2}$  per channel ( $0.42$  and  $0.5 \text{ km s}^{-1}$  for CO(3–2) and [CI], respectively). This second component can also be identified towards IRS5 and BIPS. Hence, the ring-like high velocity dispersion structure seen in the 2nd moment map of [CII] (Fig. 4, right panel) is in fact not a region of particularly high velocity dispersion but there we have two velocity components mimicking high values in the 2nd moment map. The peculiar aspect of this additional velocity component is that it is mainly detected in the ionized carbon [CII] line without a strong molecular or atomic counterpart being detectable in the individual spectra. However, Sandell et al. (2020) detected a north-south outflow in CO(3–2) and [CII] emission (their [CII] map is centered on IRS1 and smaller than the one presented here) emanating from IRS1 where the blue-shifted emission is located north of IRS1. The positions BIPN, BIPS, RPC, BrPN, and BrPS are all in the general vicinity of that outflow. In particular, the morphology of the blue-shifted [CII] and CO(3–2) gas emission near IRS1 (corresponding roughly to our position BIPS) is similar. Hence, the blue-shifted gas we find may at least partly be associated with that large-scale outflow. Nevertheless, in the spectra extracted at individual positions as shown in Figure 5, the blue-shifted gas is much stronger in the [CII] spectra than in the more commonly as outflow tracer observed CO emission. We will get back to this component in Section 4.4.

<sup>5</sup> Fig. A.1 in the Appendix presents a more extended grid of spectra in the [CII] and CO(3–2) lines.



**Fig. 6.** [CII] channel map for the velocities labeled in each panel. The data are binned for that plot in  $2 \text{ km s}^{-1}$  channels. The two purple stars mark the positions of IRS1 and IRS5, and a scale-bar is shown in the top-left panel. The intensities are always plotted logarithmically from 1 to 40 K. The purple circles outline potential bubbles. These bubbles are labeled in the top-left panel. A version without any suggested bubble structures is presented in the Appendix in Fig. A.2.

To look at the velocity structures in a different fashion, Fig. 6 presents a channel map of the [CII] emission in NGC7538 (a version without any suggested bubble structures is presented in the Appendix in Fig. A.2). While in this representation the main emission of the PDR associated with the HII regions is also dominant approximately between  $-60$  and  $-50 \text{ km s}^{-1}$ , there are considerable extended emission structures at blue-shifted velocities  $< -60 \text{ km s}^{-1}$ , and to a lesser degree also at red-shifted velocities  $> -50 \text{ km s}^{-1}$ . For comparison, Figure A.3 in the appendix presents the corresponding CO(3–2) channel map where especially these more blue- and more red-shifted features are barely recognizable. The [CII] channel map gives the impression of several bubble- and ring-like structures that may stem even from several star formation events. We will discuss the bubble-features and their possible interpretation in more detail below in Section 4.1.

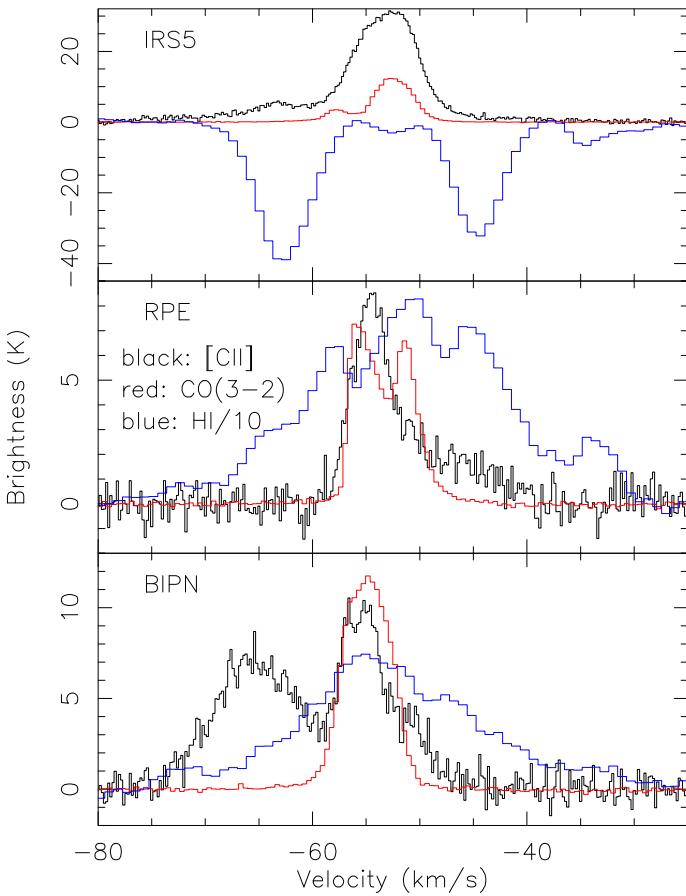
In addition to the molecular CO, atomic [CI] and ionized carbon [CII] emission, we can also investigate the atomic hydrogen by means of the 21 cm line observed with the Canadian Galactic Plane Survey (CGPS, Taylor et al. 2003). As expected, towards the main HII region around the exciting sources IRS5 and IRS6, the HI is seen only in absorption. One clearly identifies a strong HI absorption component associated with the high-velocity blue-shifted gas  $< -60 \text{ km s}^{-1}$ , but the main velocity component of the region around  $-55 \text{ km s}^{-1}$  exhibits only weak absorption. The blue-shifted HI absorption should be related to the expanding

shell accelerating also the atomic gas envelope in the direction of the observer. The red-shifted absorption at  $\sim -45 \text{ km s}^{-1}$  cannot be the other side of the expanding shell because in absorption spectroscopy it has to lie in front of the ionized gas. Hence, the red-shifted HI absorption must be related to some foreground gas. Since there is barely any [CII] nor CO emission at these velocities, that component may even be unrelated to the NGC7538 HII region.

Looking a bit outside the actual HII region towards the east (position RPE, middle panel of Fig. 7), the HI emission spectrum is broader than what we observe in [CII] and CO. This can be understood in a way that HI is more easily excited than [CII] and CO, and hence can pick up more tenuous gas at higher velocities. Furthermore, at the velocity of the [CII] and CO(3–2) peak emission, the HI shows a dip in the spectrum which may be caused by HI self-absorption (HISA, e.g., Gibson et al. 2005; Syed et al. 2020).

Furthermore, if we look towards positions where we clearly identify two components in the [CII] emission, e.g., the position BIPN in the north of the HII region (bottom panel of Fig. 7), it may at first sight be surprising that the strong blue-shifted [CII] component at velocities  $< -60 \text{ km s}^{-1}$  is inconspicuously weak in the HI emission. One potential explanation for that could be that the radiation field may be that strong that most of the gas is

ionized hydrogen and carbon with just a thin layer of neutral HI and molecular gas in the surrounding molecular cloud.



**Fig. 7.** [CII] (black), CO(3-2) (red) and HI (blue, divided by 10) spectra towards three selected positions marked in Fig. 4. These are IRS5, the red-shifted peak towards the east (RPE), and the blue-shifted peak in the north (BIPN). The positions are labeled in each panel.

## 4. Discussion

### 4.1. Expanding bubbles or an inhomogeneous medium?

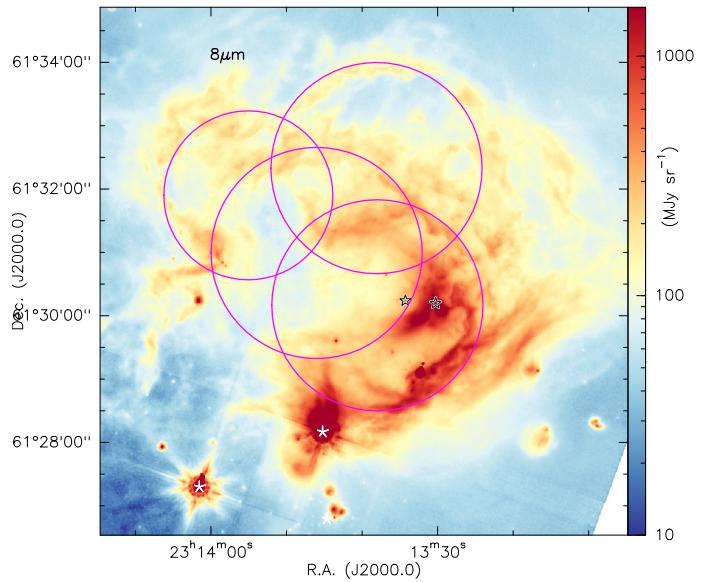
**Bubble/ring identification:** As indicated in section 3.2, the channel map of the [CII] emission shows several ring-like structures by eye-inspection (four rings are outlined in 6 as magenta circles). These features can be limb-brightened edges of expanding shells or bubbles, but they could also be a ring or torus in the plane of the sky. Before we discuss their velocity structure in more detail, we apply an unbiased automated ring identification by computing the covariance between the map and annuli structures of varying radii. The details of the identification analysis are presented in appendix B. We stress that the bubbles/rings discussed below are only candidates for real physical structures that may stem from the expansion of the HII region. The bubble- or ring-like structures we identify this way have central positions (in J2000.0) and approximate radii  $r$  of:

- (A) R.A.=23h13m38s, Dec=61°32'20";  $r \sim 100'' \sim 1.3$  pc
- (B) R.A.=23h13m38s, Dec=61°30'12";  $r \sim 100'' \sim 1.3$  pc
- (C) R.A.=23h13m46s, Dec=61°31'00";  $r \sim 100'' \sim 1.3$  pc
- (D) R.A.=23h13m55s, Dec=61°31'56";  $r \sim 80'' \sim 1.0$  pc

As outlined in appendix B, the approach identifies even a fifth ring-like structure (labeled E), however, that appears rather as an

overlap of mainly structures A to C, and we do not consider that further as a separate physical entity. Uncertainties for the central positions and radii are also discussed in appendix B.

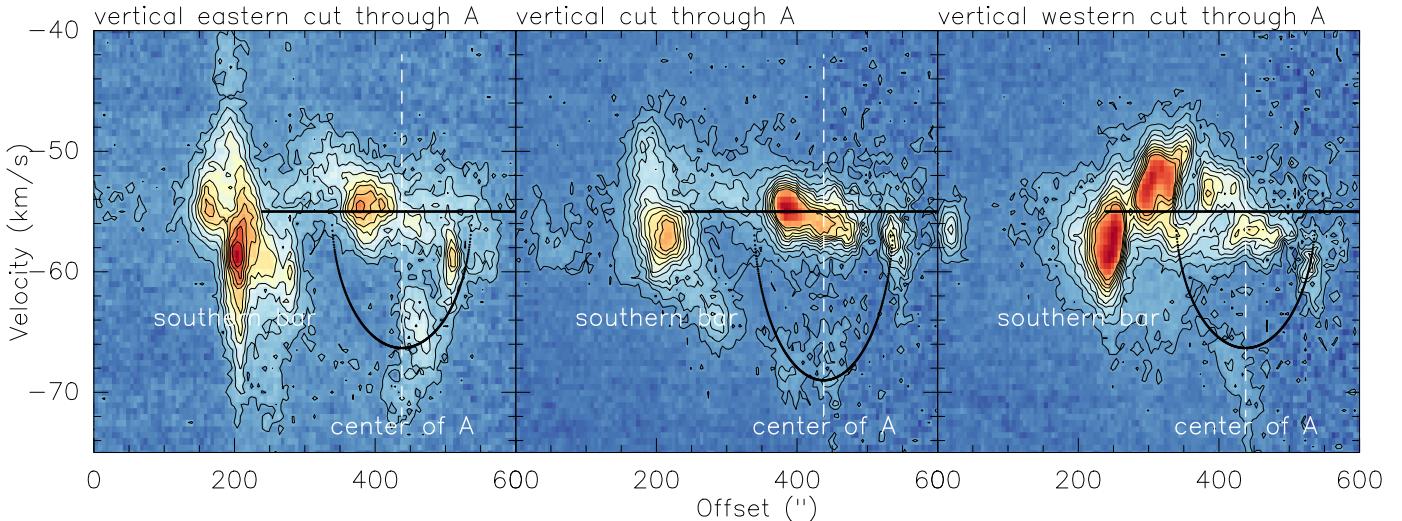
For comparison, the ring/bubble candidates are also plotted on the  $8\mu\text{m}$  emission that mainly stems from PAHs emitted in PDRs (Fig. 8). Several of the ring/bubble-like structures can clearly be seen in this PAH emission.



**Fig. 8.** The color-scale shows the  $8\mu\text{m}$  emission, and the purple circles again outline the tentative bubble structures.

The suggested bubble A is best visible as limb-brightened rings in the channels between  $-65$  and  $-55\text{ km s}^{-1}$ . An additional indicator that this bubble may indeed be a real physical structure is that at the most blue-shifted velocities (channel at  $-69\text{ km s}^{-1}$ ), the emission is mainly at the center of this proposed bubble A. This is exactly what one expects when this bubble is expanding along the line of sight (see also appendix B), and was shown also for RCW120 (Luisi et al. 2021).

**Position-velocity analysis:** A different way to visualize the velocity structure of bubble A is via a position-velocity cut. Figure 9 presents three vertical pv-diagrams in the [CII] line through the center of bubble A and shifted 1' to the east and west, respectively. Corresponding horizontal pv-diagrams are presented in the appendix in Fig. A.4. While most velocities along these cuts are between  $-50$  and  $-60\text{ km s}^{-1}$ , there is clearly high-velocity blue-shifted gas around  $\sim -70\text{ km s}^{-1}$  close to the center of the bubble, indicating an expanding structure along the line of sight. Figures 9 and A.4 also show the shape a spherical shell with a radius if  $100''$  and an expansion velocity of  $14\text{ km s}^{-1}$  would have in such a pv-diagram. This is not a fit but just outlining how such a shell would appear as an ellipse in the pv-diagram. The left and right panels with pv-cuts at offsets of 1' correspond in a shell geometry to angles of  $\sim 36$  deg (see a sketch of such a geometry in appendix A.2 from Butterfield et al. 2018). The velocities at such angle are then reduced by  $\cos(36^\circ)$ . While the data and toy model at negative blue-shifted velocities show some resemblance of each other, consistent with a spherical half-shell, the red-shifted side at positive relative velocities shows no indication of such a shell-like geometry. This could either indicate that such a shell model is insufficient, or the center of the shell could be close to the rear side of the cloud that only blue-shifted



**Fig. 9.** Vertical position-velocity cuts through bubble A from south to north derived from the [CII] line emission. The central panel goes through the estimated central position of A whereas the left and right panels are shifted 1' to the east and west, respectively. The white dashed line marks the approximate center of the bubble. The black dotted lines in the middle panel outline the elliptical shape a spherical shell with radius 100'' and expansion velocity of 14 km s<sup>-1</sup> would have in the pv diagram. The dotted lines in the left and right panel correspond to the same expanding shell at the 1' shift to the east and west where the 1' corresponds to  $\sim 36$  deg offset in the sphere (e.g., Butterfield et al. 2018, sketch in their appendix A.2). Hence the velocities are then reduced by  $\cos(36^\circ)$ .

gas moving towards the observer can be detected while the red-shifted expansion leaves the cloud almost immediately that no gas gets accelerated to a detectable level. In that context, Luisi et al. (2016) estimated that approximately 15% of the ionized gas is leaking outside of the HII region.

The structure B is centered close to the main excitation sources in the region, IRS5 and IRS6 (Fig. 1). But we note that the automated bubble-identification approach does not center it exactly on the excitation sources but a bit shifted. Bubble B's border at the south-western side is roughly the bar-like PDR discussed in section 4.2. In the north, structures visible in particular at velocities of  $-55/-53$  km s<sup>-1</sup> delineate the potential sphere of that bubble. Doing position-velocity cuts in the horizontal as well as vertical direction through this bubble B and source IRS5 (Fig. 10), one also finds that the highest-velocity gas on the blue- as well as red-shifted side is found close to IRS5. This is again indicative of expanding motions along the line of sight, potentially caused by IRS5 and IRS6. We again draw the corresponding shapes for expanding shells into the pv-diagram (Fig. 10). While the radius of 100'' is given by our shell-identification approach above, the velocities on the blue- and red-shifted side differ significantly. To outline those differences, we used an expansion velocity of 10 km s<sup>-1</sup> for the blue-shifted side, and 3 km s<sup>-1</sup> for the red-shifted side. If the observed structure really belongs to an expanding shell, such velocity differences would indicate a highly inhomogeneous medium where the red-shifted expansion would enter into much denser material. Another option is that the central velocity is not  $-55$  km s<sup>-1</sup> but a few km s<sup>-1</sup> higher. In that case, we would just have a blue-shifted bubble-structure without any red-shifted counterpart, similar to structure A discussed before. That would put the center of structure B also on the rear side of the cloud.

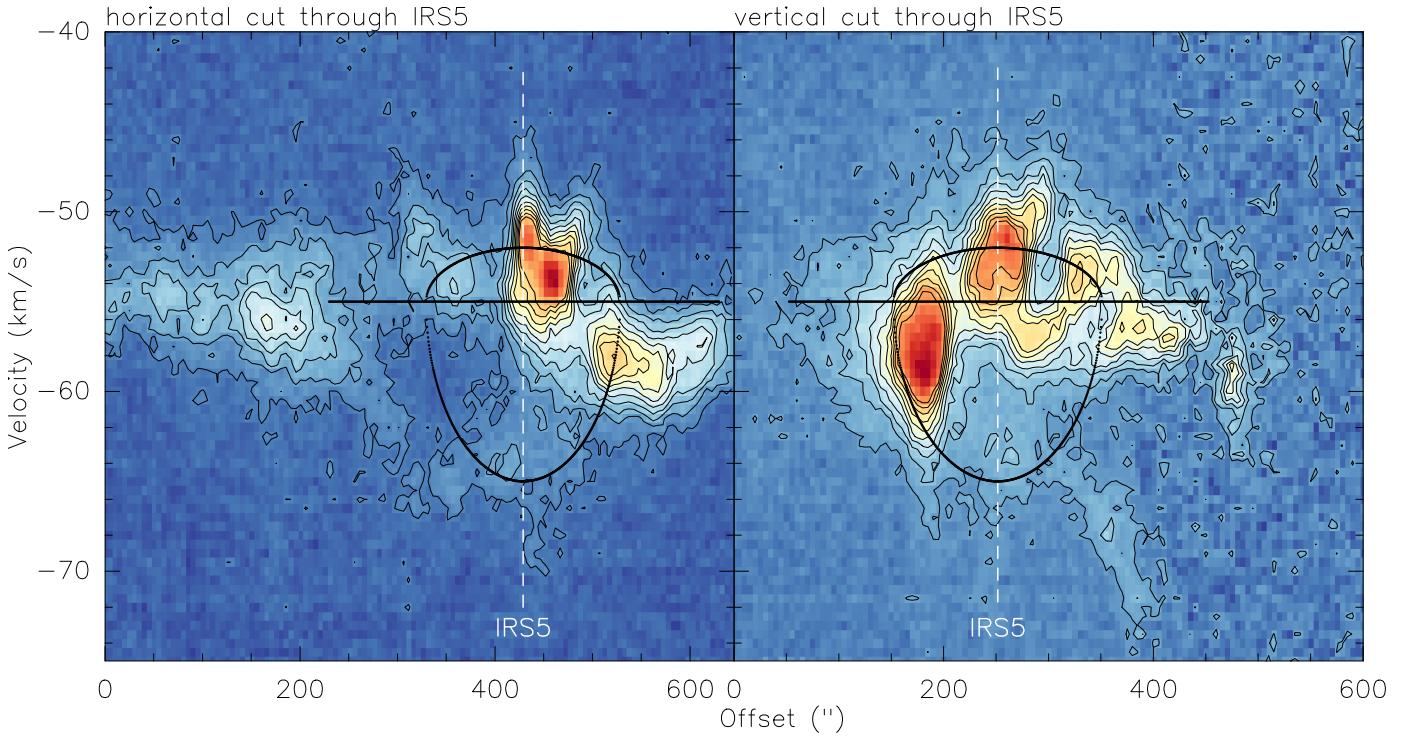
The identified structure C is in-between the other bubble/ring structures. While it may be a real physical entity, it could also arise from an overlap of the other bubble/ring structures. Therefore, we omit further detailed discussion for that structure.

Furthermore, we identify structure D where the limb-brightened ring-like features are best delineated in the channels

between  $-57$  and  $-53$  km s<sup>-1</sup>. In addition to the bubble edges, structure D again shows high-velocity gas at its center, typically signposting expanding structures. These centered emission structures in bubble D are best visible in the blue-shifted channels at  $-67/-65$  km s<sup>-1</sup> and in the most red-shifted channels at  $-49/-47$  km s<sup>-1</sup>. For comparison, Fig. 11 presents the corresponding vertical and horizontal pv-cuts through the central position of structure D. In contrast to the pv-diagrams of bubbles A and B (Figs. 9 and 10) that show the high-velocity gas mainly at negative blue-shifted velocities, the pv-diagram of structure D, in particular the one in vertical direction (right panel in Fig. 11), shows an almost ring-like structure with high velocities on the blue- and red-shifted side, in good agreement with the channel-map discussed before. For comparison, we again show the shape an expanding shell with radius 80'' and an expansion velocity of 6 km s<sup>-1</sup>.

**Driving sources:** It is interesting to note that the ring-like structure of high velocity dispersion in the [CII] moment map (Fig. 4) roughly coincides with the combined outline of the four bubble/ring structures discussed here (Fig. 6). Since the high velocity dispersion in Fig. 4 largely stems from several velocity components (Fig. 5), these velocity signatures may stem from expanding motions from the HII region. While for bubble B, the driving source for bubble expansion may be the main exciting sources of the HII region (IRS5 and IRS6, Puga et al. 2010), this is less obvious for the other structures A, C and D. For these three structures, we have not identified a clear central source that could be associated with bubble expansion. While Puga et al. (2010) identified more young stellar objects in the environment of the HII region, they are not at the bubble centers (which in fact IRS5 and IRS6 are not exactly either), and furthermore are typically of lower luminosity. Also Spitzer data do not allow us to identify clear driving sources.

While the O3 and O9 stars IRS6 and IRS5 (Puga et al. 2010) are most likely the main exciting sources of the HII region, the identification of several more ring- or bubble-like structures does not necessarily mean that there are more bubble-driving sources



**Fig. 10.** Horizontal (along R.A., left) and vertical (along Dec., right) position-velocity diagrams through the position of one of the exciting sources IRS5 (marked in each panel) derived from the [CII] line emission. corresponding roughly to the estimated center of bubble B. The contour levels are in  $3\sigma$  steps. The black dotted lines outline the elliptical shape a spherical shell with radius  $100''$  and expansion velocities of  $10 \text{ km s}^{-1}$  and  $3 \text{ km s}^{-1}$  at negative and positive relative velocities, respectively, would have in the pv diagram.

in the field. While proper motion of the exciting stars may cause the dislocation of IRS5 and IRS6 from the center of bubble-structure B, such displacement may also be caused by density and pressure gradients in the region. A similar displacement is also observed between the Trapezium stars and the main bubble in Orion, the Orion Veil (Pabst et al. 2019, 2020).

**Multiple bubbles versus expansion into an inhomogeneous medium:** In addition to this, the strong sub-structure of the H $\alpha$  region visible in the different data sets, in particular in the [CII] channel map, indicates an almost porous structure of the H $\alpha$  region where radiation and wind energy can leak out and influence also parts of the H $\alpha$  region that are less close to the main exciting sources IRS5 and IRS6. Hence, while bubble-like structures are possible, the NGC7538 H $\alpha$  region may also expand into a very inhomogeneous environment. If an H $\alpha$  region expands in such an inhomogeneous cloud, changes of the cloud morphology are expected. Such impact may also cause bubble-like morphologies as observed here. Recent analytic and numerical work by Lancaster et al. (2021a,b,c) also showed that the wind-driven expansion of the gas around H $\alpha$  regions can cause different morphological sub-structures. For the NGC7538 H $\alpha$  region, a clear discrimination between several bubble-structures each driven by individual internal sources and/or an expansion of mainly one wind-driven shell into an inhomogeneous medium cannot be drawn. Further hydrodynamic modeling of expanding H $\alpha$  regions into an inhomogeneous medium may show whether such multiple ring-like structures of expanding gas can occur or whether separate bubble-driving centers are required.

**Thermal expansion or winds of the high-mass stars:** Independent of whether all the ring-like structures are real bubbles

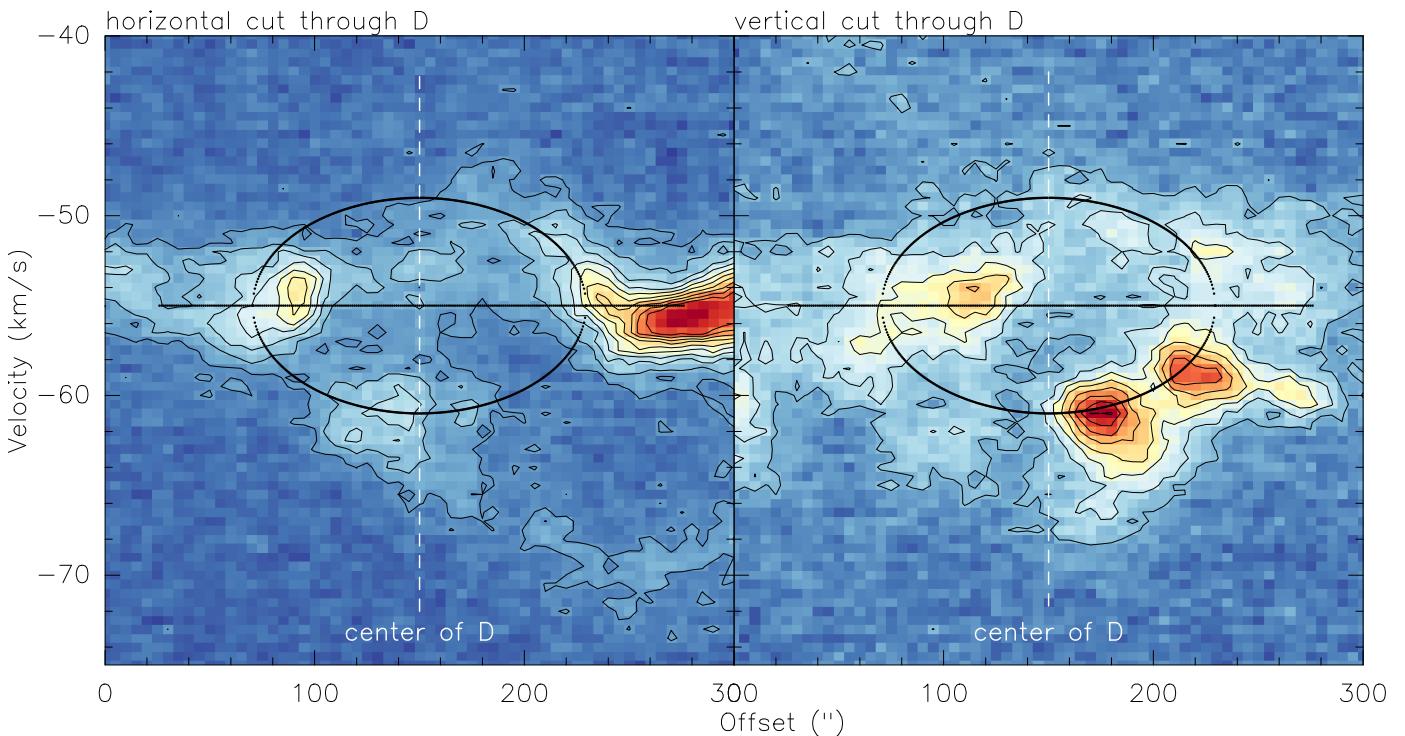
or whether they are rather caused by an expansion of the gas into an inhomogeneous medium, the high-velocity components visible in the pv-diagrams (Figs. 9 to 11) all show high-velocity gas likely driven by the expansion of the H $\alpha$  region. In particular the gas towards structures A and B exhibit expansion velocities  $\geq 10 \text{ km s}^{-1}$ . Taking an expansion velocity of  $10 \text{ km s}^{-1}$  and a radius of  $100''$  at face-value, this would correspond to an expansion time-scale of roughly 0.126 Myrs. This is comparably short with respect to the observationally estimated ages of the cluster. For example, Puga et al. (2010) derived an age range between 0.5 and 2.2 Myrs, and Sharma et al. (2017) estimated a mean age of the young stellar objects of 1.4 Myrs with a range between 0.1 and 2.5 Myrs. Hence, a constant expansion at that velocity is unlikely the real scenario for that region. However, can the measured expansion velocities be caused by thermal expansion of the gas or by the winds of the high-mass stars?

To estimate expected expansion velocities, for the thermal expansion we follow the Spitzer solution (Spitzer 1998), whereas for the wind-driven expansion, the approach by Weaver et al. (1977) and its adaption to include radiative cooling by Mac Low & McCray (1988) is adopted. A comprehensive summary of these two approaches is given in Henshaw et al. (2021). The expansion velocities depending on time for the thermal expansion  $v_{th}(t)$  and the wind-driven expansion  $v_w(t)$  are:

$$v_{th}(t) = c_s \left(1 + \frac{7c_s t}{46}\right)^{-3/7}$$

$$v_w(t) = \frac{1}{2} \frac{125}{154\pi} \left(\frac{L_w}{\rho}\right)^{1/5} t^{-1/2} t_{cool}^{1/10}.$$

Here,  $c_s$  is the sound speed in the ionized gas ( $8 \text{ km s}^{-1}$  at 5000 K),  $L_w = \frac{1}{2} \dot{M} v_\infty^2$  is the mechanical wind luminosity,  $\rho$  the



**Fig. 11.** Horizontal (left) and vertical (right) position-velocity diagrams through the center of bubble D (marked in each panel) derived from the [CII] line emission. The contour levels are in  $3\sigma$  steps. The black dotted lines outline the elliptical shape a spherical shell with radius 80" and expansion velocities of  $6 \text{ km s}^{-1}$  would have in the pv diagram.

ambient density, and  $t_{cool}$  the cooling time following Mac Low & McCray (1988) and Henshaw et al. (2021). The wind mass flow rate  $\dot{M}$  and the escape velocity  $v_\infty$  for an O3 star (IRS6, Puga et al. 2010) are taken from Muijres et al. (2012). The wind-driven velocities  $v_w$  strongly depend on the density of the ambient gas  $\rho$  (also  $t_{cool}$  depends on  $\rho$ , Mac Low & McCray 1988).

Following Puga et al. (2010) and Sharma et al. (2017), the age of the NGC7538 cluster and associated young stellar objects is between roughly 0.1 and 2.5 Myrs. Even at the lower boundary of 0.1 Myr, the highest expansion velocities one can get with a Spitzer-type thermal expansion is  $\sim 5.7 \text{ km s}^{-1}$ , far below what is measured in NGC7538. So, purely thermal expansion cannot properly explain the measured high-velocity gas. In contrast to that, the wind solution gives significant higher velocities. For the given parameters, and assuming densities of  $10^3$ ,  $2 \times 10^3$  and  $3 \times 10^3 \text{ cm}^{-3}$ , expansion velocities around  $10 \text{ km s}^{-1}$  are estimated at times of roughly 1.1, 0.36 and 0.17 Myr, respectively. While the lower end would be more consistent with the constant expansion estimated above, the higher time-scales are more consistent with the the estimated age limits by Puga et al. (2010) and the mean ages of the young stellar objects derived by Sharma et al. (2017). Hence, wind-driving seems a plausible way to explain the observed high-velocity gas. The wind-driving is also consistent with the diffuse X-ray emission towards that region that is typically attributed to hot plasma caused by the wind shocks of the massive stars (e.g., Güdel et al. 2008; Townsley et al. 2018; Pabst et al. 2020).

Regarding the potential discrepancy between the lower end of the estimated cluster age at 0.5 Myrs and the shorter time-scales estimated for either constant expansion or the wind expansion at higher densities, one way to reconcile these is that initially the bubble may have confined inside the dense molecular core in which IRS5 and IRS6 have formed. Only once the HII region broke out of this core, expansion was rapid into the sur-

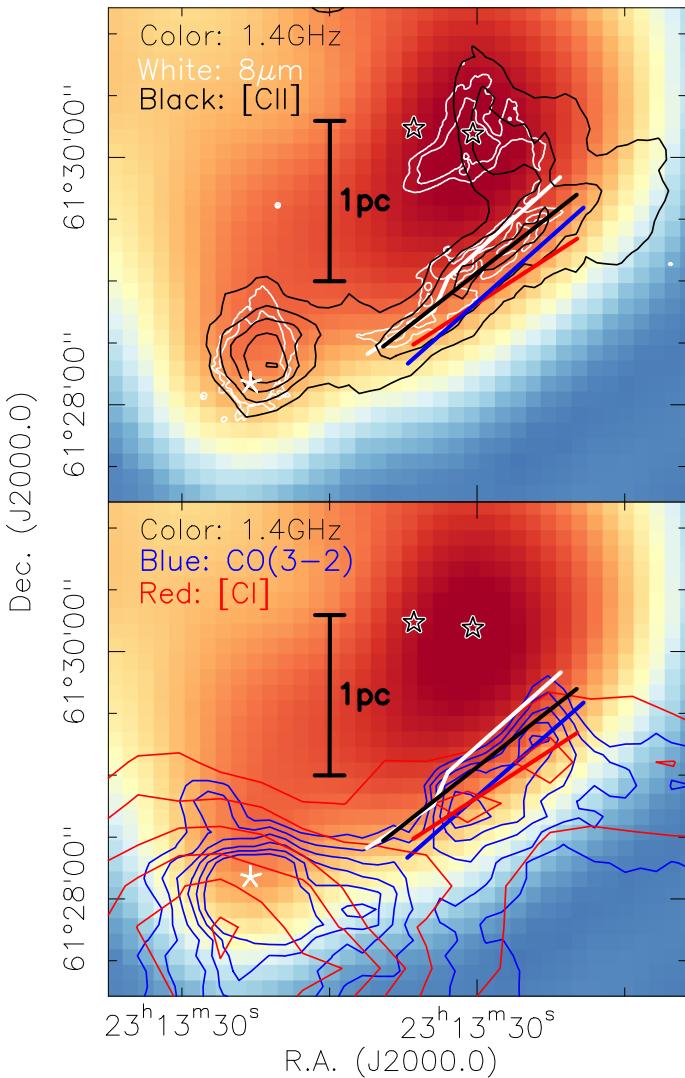
rounding lower density material. In this picture, the expansion timescale one estimates then refers to the time since the breakout from this dense core.

#### 4.2. The bar-like photon dominated region (PDR)

Figure 12 presents a zoom into the bar-like structure at the south-western edge of the expanding HII region. While the color-scale shows the ionized gas as traced by the 1.4 GHz continuum emission, the contours outline various other tracers:  $8 \mu\text{m}$  continuum, [CII], CO(8-7) & (3-2) and [CI] emission. The white, black, blue and red lines mark the approximate locations of the emission crests in the  $8 \mu\text{m}$ , [CII], CO(3-2) and [CI] emission, respectively. The crests were identified by eye via connecting the emission peaks. The crest of the CO(8-7) emission approximately coincides with that of the [CII] emission.

In this picture, the  $8 \mu\text{m}$  emission, that should stem largely from UV-pumped infrared fluorescence PAHs, peaks closest to the exciting sources of the HII region. Second in this layered structures is the ionized carbon [CII] emission that spatially approximately coincides with the highly excited CO(8-7) emission. Furthest away from the exciting sources are the emission crests of the lower excited CO(3-2) and [CI] emission that, given the spatial resolution of the data, projected on the plane of the sky spatially approximately coincide. This is different around the young high-mass star-forming region IRS1 south-east of the bar where CO(3-2) peaks closer to IRS1 than the atomic carbon [CI] line (Fig. 12). Coming back to the bar-like structure, in this framework, we can separate the photon dominated region (PDR) in roughly four layers:

- (i) The exciting sources and 1.4 GHz continuum emission
- (ii) the  $8 \mu\text{m}$  emission, likely PAHs
- (iii) ionized carbon [CII] and highly excited CO(8-7)

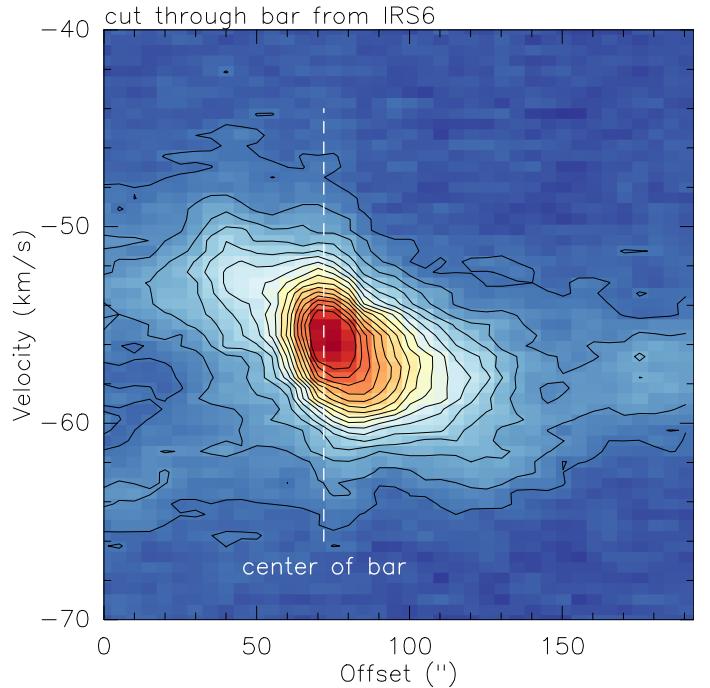


**Fig. 12.** Zoom on the bar-region south of the exciting sources IRS5 and IRS6. The color-scale and the colored contours show the emission tracers labeled in the two panels. Only selected contours are shown for clarity. The four colored lines mark the approximate positions of the emission crests in the corresponding tracers. The sources IRS5, IRS6 and IRS1 are marked. A linear scale-bar is shown as well.

#### (iv) atomic carbon [CI] and lower excited CO(3–2)

This layered structure resembles prototypical edge-on PDR structures observed also for example in the Orion bar (e.g., Tielemans et al. 1993; Tauber et al. 1994, 1995; Simon et al. 1997) or M17SW (e.g., Pérez-Beaupuits et al. 2015b,a). We will return to the carbon budget within this PDR in Section 4.3.

In addition to the layered structure, the 1st moment map in Fig. 4 indicates that there may be a velocity gradient across the bar. To get a better view of that, Fig. 13 presents a position-velocity diagram across the bar starting from the O3 star IRS6. We see a clear velocity gradient from  $\sim -50 \text{ km s}^{-1}$  at the beginning of the cut near the main exciting source IRS6 to blue-shifted velocities  $< -60 \text{ km s}^{-1}$  on the other side of the bar. The fact that the gas within the bar gets blue-shifted indicates that the bar lies along the line of sight to the observer in front of the exciting source IRS6.



**Fig. 13.** Position-velocity cut from the [CII] line from IRS6 through the bar as outlined in the middle panel of Fig. 4. The white dashed line marks the approximate center of the bar.

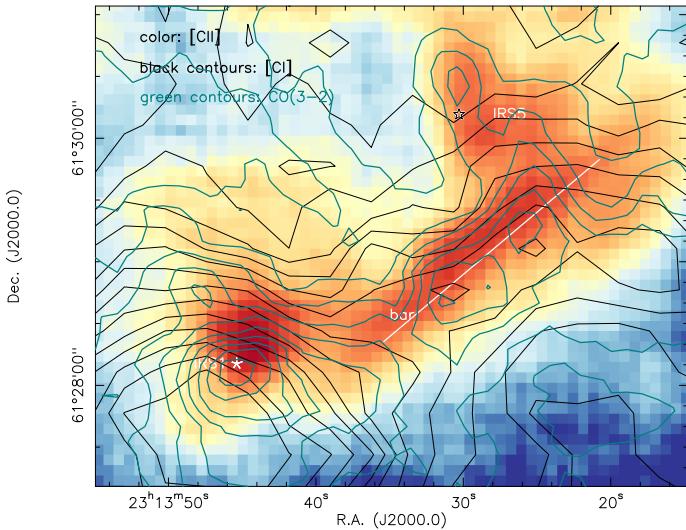
#### 4.3. Carbon budget

While carbon is mainly in its ionized  $\text{C}^+$  state in large fractions of the PDR, it is mainly in its molecular CO form in the dense molecular cloud (Fig. 2). But what is the relative carbon budget within the transition zone from the H $\alpha$  region towards the molecular clouds? This transition can best be studied in the south-western bar-like structure. Figure 14 presents the ionized carbon [CII], the atomic carbon [CI] and the molecular carbon CO(3–2) emission toward that region. The integration regime for all three species is the same from  $-70$  to  $-40 \text{ km s}^{-1}$ . In the following, we derive approximate estimates for the column densities and masses of the different carbon components in the bar region outlined by Fig. 14. With the assumptions outlined below, these should only be considered as rough estimates. But nevertheless, they allow us to assess roughly how much mass is within each carbon component in and around the bar-like region.

The column densities of the ionized carbon are estimated from the [CII] line following Goldsmith et al. (2012), eq. 26, adding an additional optical depth term of  $\frac{\tau}{1-e^{-\tau}}$  to correct for the optical depth  $\tau$  (Rohlfs & Wilson 2006):

$$N_{[\text{CII}]} = \frac{1}{3.43 \times 10^{-16}} \left[ 1 + 0.5e^{91.25/T_{\text{kin}}} \left( 1 + \frac{2.4 \times 10^{-6}}{C_{ul}} \right) \right] \times \frac{\tau}{1 - e^{-\tau}} \int T_{\text{mb}} dv [\text{cm}^{-2}]$$

with the kinetic temperature  $T_{\text{kin}}$  (see further below), the collision rate  $C_{ul}$ , the intensity  $T_{\text{mb}}$  in K and the velocity  $dv$  in  $\text{km s}^{-1}$ . The collision rate  $C_{ul} = R_{ul}n$  depends on the temperature, where  $R_{ul}$  is collision rate coefficient with  $\text{H}_2$  and  $n$  the density, assumed to be around  $\sim 10^3 \text{ cm}^{-3}$ . The collision rate coefficients  $R_{ul}$  with  $\text{H}_2$  are taken from the Leiden database for molecular spectroscopy (Schöier et al. 2005, <http://home.strw.leidenuniv.nl/~moldata/>, Lique et al. 2013)



**Fig. 14.** Integrated [CII] (color), [CI] (black contours) and CO(3–2) (green contours) emission in the bar-region where the [CII]/[CI]/CO mass estimates are conducted (section 4.3). All three maps were integrated between  $-70$  and  $-40\text{ km s}^{-1}$ . The [CI] contours start at  $6\sigma$  and continue in  $3\sigma$  steps ( $1\sigma = 1.6\text{ K km s}^{-1}$ ). The CO(3–2) contours start at  $10\sigma$  and continue in  $10\sigma$  steps ( $1\sigma = 3.2\text{ K km s}^{-1}$ ). The white line marks the approximate location of the bar-like feature.

where the para- and ortho-rates are weighted following Le Bourlot (1991) and Gerlich (1990). The comparably good spatial correspondence of the [CII] with the molecular line emission around the bar-shaped PDR indicates that  $\text{H}_2$  should be the most dominant collisional partner. If for comparison all collisions were with atomic hydrogen, with the given assumptions we would get masses roughly a factor of 1.57 lower (see also Goldsmith et al. 2012). In reality, it is likely a mix of collisional partners dominated by  $\text{H}_2$  and hence the difference should be much smaller. In the following we use the collisional rates with  $\text{H}_2$ . Evaluating the [ $^{13}\text{CII}$ ] emission in our data, we detect the line after averaging over areas to decrease the noise. In the case of the bar-like region in NGC7538, the [ $^{13}\text{CII}$ ] data indicate a mean optical depth  $\tau$  for the main [CII] line of  $\sim 2.5$ . Compared to the optically thin case, this corresponds to a column density and mass correction factor of  $\sim 2.7$ .

Furthermore, the atomic carbon column densities are estimated from the [CI] emission following Frerking et al. (1989):

$$N_{[\text{CI}]} = 5.94 \times 10^{15} \frac{1 + 3e^{\frac{-23.6}{T_{\text{ex}}}} + 5e^{\frac{-62.4}{T_{\text{ex}}}}}{3e^{\frac{-23.6}{T_{\text{ex}}}}} \int T_{\text{mb}} dv [\text{cm}^{-2}]$$

with the excitation temperature  $T_{\text{ex}}$  (the temperature quantification is discussed further below). The CO column density can then approximately be estimated following standard equations (e.g., Cabrit et al. 1988):

$$N_{\text{CO}} = \frac{6.97 \times 10^{15}}{\nu^2 \mu^2} T_{\text{ex}} e^{E_u/(kT_{\text{ex}})} \frac{\tau}{1 - e^{-\tau}} \times \int T_{\text{mb}} ({}^{12}\text{CO}(3 - 2)) dv [\text{cm}^{-2}]$$

Here,  $\nu$ ,  $\mu$ ,  $T_{\text{ex}}$ ,  $E_u/k$ ,  $\tau$  and  $T_{\text{mb}}$  are the frequency, dipole moment (0.112 Debye), the excitation temperature, the upper energy level of 33.19 K, the optical depth and the main beam brightness temperature of the line, respectively.

Following Cabrit et al. (1988), assuming  $^{13}\text{CO}$  is optically thin, the optical depth term can be approximated as:

$$\frac{\tau}{1 - e^{-\tau}} = \frac{[{}^{12}\text{CO}]}{[{}^{13}\text{CO}]} \left( \frac{\int T_{\text{mb}} ({}^{13}\text{CO}) dv}{\int T_{\text{mb}} ({}^{12}\text{CO}) dv} \right)$$

The term  $\frac{[{}^{12}\text{CO}]}{[{}^{13}\text{CO}]}$  corresponds to the  ${}^{12}\text{CO}/{}^{13}\text{CO}$  abundance ratio. According to Moscadelli et al. (2009), NGC7538 is in the Perseus arm which results at the given longitude of the region (111.5 deg) in a Galactocentric distance of  $\sim 9$  kpc (e.g., Reid et al. 2019). Following Wilson & Rood (1994), the  ${}^{12}\text{CO}/{}^{13}\text{CO}$  abundance ratio is then  $\sim 75$ . Furthermore, the term  $\left( \frac{\int T_{\text{mb}} ({}^{13}\text{CO}) dv}{\int T_{\text{mb}} ({}^{12}\text{CO}) dv} \right)$  corresponds to the mean ratio of brightness intensities between the  ${}^{13}\text{CO}$  and the  ${}^{12}\text{CO}$  transition. Within the IRAM NOEMA large program CORE (Beuther et al. 2018), a small area of  $1.5' \times 1.5'$  around IRS1 had been mapped in the  $J = 2 - 1$  transitions of  ${}^{12}\text{CO}$  and  ${}^{13}\text{CO}$  with the IRAM 30 m telescope. Using these small-area data, we estimate a mean ratio between those two lines of  $\sim 0.2$  in the velocity regime between  $-70$  and  $-40\text{ km s}^{-1}$ . This value is used in the following also for the here observed  $J = 3 - 2$  transition. This results in the column density equation:

$$N_{\text{CO}} = \frac{6.97 \times 10^{15}}{\nu^2 \mu^2} T_{\text{ex}} e^{E_u/(kT_{\text{ex}})} \frac{[{}^{12}\text{CO}]}{[{}^{13}\text{CO}]} \left( \frac{\int T_{\text{mb}} ({}^{13}\text{CO}) dv}{\int T_{\text{mb}} ({}^{12}\text{CO}) dv} \right) \times \int T_{\text{mb}} ({}^{12}\text{CO}(3 - 2)) dv [\text{cm}^{-2}]$$

In the following, we will calculate the mean column densities and the corresponding masses over the area presented in Figure 14 that encompassed roughly  $11.25\text{ pc}^2$ . The main missing ingredient in the above equations are the temperatures to be chosen for the column density estimates.

Regarding the ionized carbon C<sup>+</sup>, Langer et al. (2010) typically assume temperatures between 100 and 150 K, which seems a reasonable temperature regime in this PDR as well. Using these two values as brackets for the possible temperatures of the ionized carbon C<sup>+</sup> we find an ionized carbon C<sup>+</sup> mass of  $M_{[\text{CII}]} \approx 9.7 \pm 1.6\text{ M}_{\odot}$ , assuming the above discussed mean optical depth  $\tau \sim 2.5$ . This is an upper limit to the mass (Rohlfs & Wilson 2006) whereas we can consider an estimate assuming optically thin emission as a lower limit. In the optically thin limit, we get the lower limit of  $M_{[\text{CII}] \text{ lower limit}} \approx 3.6 \pm 0.6\text{ M}_{\odot}$ , a factor  $\sim 2.7$  below the optical depth corrected mass estimate.

For the atomic and molecular gas, the excitation temperatures are not well determined in that region. However, variations in the excitation temperatures between 30 and 50 K do not vary the estimated masses significantly. Therefore, we estimate the atomic  $M_{[\text{CI}]}$  and molecular masses  $M_{\text{CO}}$  in that excitation temperature regime. In that framework we find  $M_{[\text{CI}]} \approx 0.45 \pm 0.1\text{ M}_{\odot}$  and  $M_{\text{CO}} \approx 1.2 \pm 0.1\text{ M}_{\odot}$ . The molecular mass can also be compared to the gas mass one can derive from the Herschel gas column density map. While a  $\text{H}_2$ -to-CO ratio of  $10^4$  would result in a gas mass of  $\sim 1.2 \times 10^4\text{ M}_{\odot}$ , the gas mass derived from the Herschel map in the same area is  $\sim 4600\text{ M}_{\odot}$ . This difference around a factor 2.6 can already be explained by the uncertainties of temperatures and line-of-sight averaging, optical depth estimates and the assumed dust properties for the Herschel-based mass estimate. Nevertheless, this comparison shows that our molecular mass estimate has an approximate uncertainty of a factor two.

While atomic C<sup>0</sup> and molecular CO carbon have comparably low masses between 0.45 and 1.2 M<sub>⊙</sub>, the mass of the ionized carbon C<sup>+</sup> is in the optically thin lower limit a factor ∼3–8 larger, and for the optical depth corrected mass estimate even a factor ∼8–22 larger in this bar-like PDR (Fig. 14). Hence, the ionized carbon C<sup>+</sup> is clearly the dominant carbon form in this PDR. As expected, this is very different to the previous carbon studies in infrared dark clouds where the carbon was found to be dominantly in the molecular CO form (Beuther et al. 2014).

#### 4.4. Decoupled [CII] component

The [CII] component at velocities below  $-60 \text{ km s}^{-1}$  deserves separate discussion because it is identified mainly in the [CII] emission without strong counterparts in the molecular and atomic gas (Figs. 5 & 6). A spectral feature dominantly visible in [CII] emission has rarely been identified in the past in regions of expanding HII regions and feedback. For comparison, the veil in Orion exhibits barely any extended CO emission, however, compact globules are identified within the [CII] emitting gas structures (Goicoechea et al. 2020).

Inspecting the channel map with proposed bubble-like structures (Fig. 6) in velocity channels at  $\leq -65 \text{ km s}^{-1}$ , this blue-shifted high-velocity gas appears to be associated with all four identified bubble-structures as well as with the young star-forming region IRS1 in the south of the region. Since IRS1 is known to be an active driver of molecular outflows (e.g., Scoville et al. 1986; Kraus et al. 2006; Qiu et al. 2011; Beuther et al. 2012; Sandell et al. 2020), finding high-velocity emission in its vicinity is not a surprise. Furthermore, the blue-shifted outflow lobe emanating from IRS1 towards the north is aligned right with the central area of the HII region and may explain part of the blue-shifted emission (Sandell et al. 2020). Nevertheless, the [CII] emission below  $-60 \text{ km s}^{-1}$  is significantly stronger than the CO(3–2) emission at the same velocities (Fig. 5). Hence, a sole explanation by the outflow appears unlikely, and a significant fraction of the highly blue-shifted gas appears to be associated with the suggested bubble-structures.

Using the same assumptions as described in the previous section, we can estimate the mass of ionized carbon associated with this blue-shifted [CII] component. Integrating the blue-shifted [CII] line between  $-75$  and  $-65 \text{ km s}^{-1}$  over the entire emission area ( $\sim 22.5 \text{ pc}^2$ ), we find in the temperature regime between 100 and 150 K a mass of  $M_{[\text{CII}]}^{\text{blue}} \approx 0.66 \pm 0.1 \text{ M}_\odot$ . Although the area is roughly a factor two larger than that used for the mass estimates for the bar in the previous section, the mass in this blue-shifted component is more than a factor 10 smaller than that in the bar-like structure. Nevertheless, these data show that still a significant fraction of ionized carbon is found in that mainly in the [CII] line emitting velocity component.

Assuming that at least part of the blue-shifted high-velocity gas associated with the proposed bubble-like structures is indeed associated with expanding gas along the line of sight, the fact that we see barely any atomic or molecular carbon counterpart implies that there is very low or even no interaction of the expanding gas with a denser environmental cloud.

#### 4.5. Shell structures associated with HII regions and PDRs in [CII] emission

As mentioned in the Introduction, the existence of expanding shell-like structures around HII regions and PDRs is a recurring result of the SOFIA FEEDBACK program, where NGC7538 is

part of, as well as other [CII] studies. The best studied examples are Orion (Pabst et al. 2019), RCW120 (Luisi et al. 2021), and RCW49 (Tiwari et al. 2021). From a geometry point of view, the Orion Veil and RCW120 exhibit more complete shell-like structures in the [CII] emission. The NGC7538 data presented here also show bubble-like features, but with much more substructure and no real large-scale shell. RCW49 again exhibits a large-scale shell-structure, however, depending on the velocities of the gas, the shell is broken open allowing radiation to escape into the environment.

Regarding the expansion speeds measured in the [CII] emission, all four regions discussed here have similar expansion velocities between roughly 10 and  $16 \text{ km s}^{-1}$ . Assuming constant expansion, the given velocities result in estimated expansion time-scales between 0.1 and 0.27 Myr, all within a similar range.

The expansion of the four regions can in neither case be explained by thermal expansion like in the classical Spitzer picture (Spitzer 1998), but wind-driving from the central O-stars is needed.

Furthermore, it is interesting to note that in three of the regions – RCW120, RCW49 and NGC7538 – star formation at the edges of the expanding regions is observed. For NGC7538, triggering as a potential scenario is discussed in detail by Sharma et al. (2017). Hence, triggered star formation seems a regular outcome in such kind of regions (see also general bubble/protopstar correlation analysis by, e.g., Kendrew et al. 2016 or Palmeirim et al. 2017). In contrast to that, no obvious star formation is identified at the edges of the Orion Veil. Goicoechea et al. (2020) recently found small globules of molecular gas within the Veil structure, however, these are rather transient objects and unlikely to form stars. Hence, while the star formation activity around RCW120, RCW49 and NGC7538 hint towards at least partly positive feedback triggering new star formation events, in the Orion Veil, the feedback processes expel and reprocess the gas, limiting further star formation.

## 5. Conclusions and Summary

Using SOFIA, we are able to map roughly 210 square arcmin ( $125 \text{ pc}^2$ ) around the HII region NGC7538 in the fine-structure line of ionized carbon [CII] as well as fine-structure [CI] and CO(8–7) emission. These SOFIA data are complemented with archival near-/far-infrared/cm continuum and CO(3–2) line observations.

The [CII] emission reveals a wealth of information about the structure and composition of this region of high-mass star formation. While the overall [CII] morphology may at first sight appear similar to that of the ionized gas traced by the cm continuum emission, the velocity-resolved [CII] emission reveals considerable sub-structure resembling the PDR otherwise well traced in the 8 and  $70 \mu\text{m}$  emission. Using a automated bubble/ring-identification algorithm, we identify four bubble-like structures. Two of them show kinematic signatures of expanding bubbles. One of these structures is centered near the main exciting sources IRS6 and IRS5. For the other three bubble-like structures no clear central source can be identified. This may indicate that at least some of the bubble-like structures could be mimicked by the expansion of the gas into an intrinsically porous and inhomogeneous medium within the HII region.

An analysis of the expansion velocities based on high-velocity [CII] gas components reveals that purely thermal expansion from a Spitzer-like HII region is not sufficient, but that wind-driving by the central O-stars is needed. This is supported by Chandra observations revealing the presence of diffuse X-ray

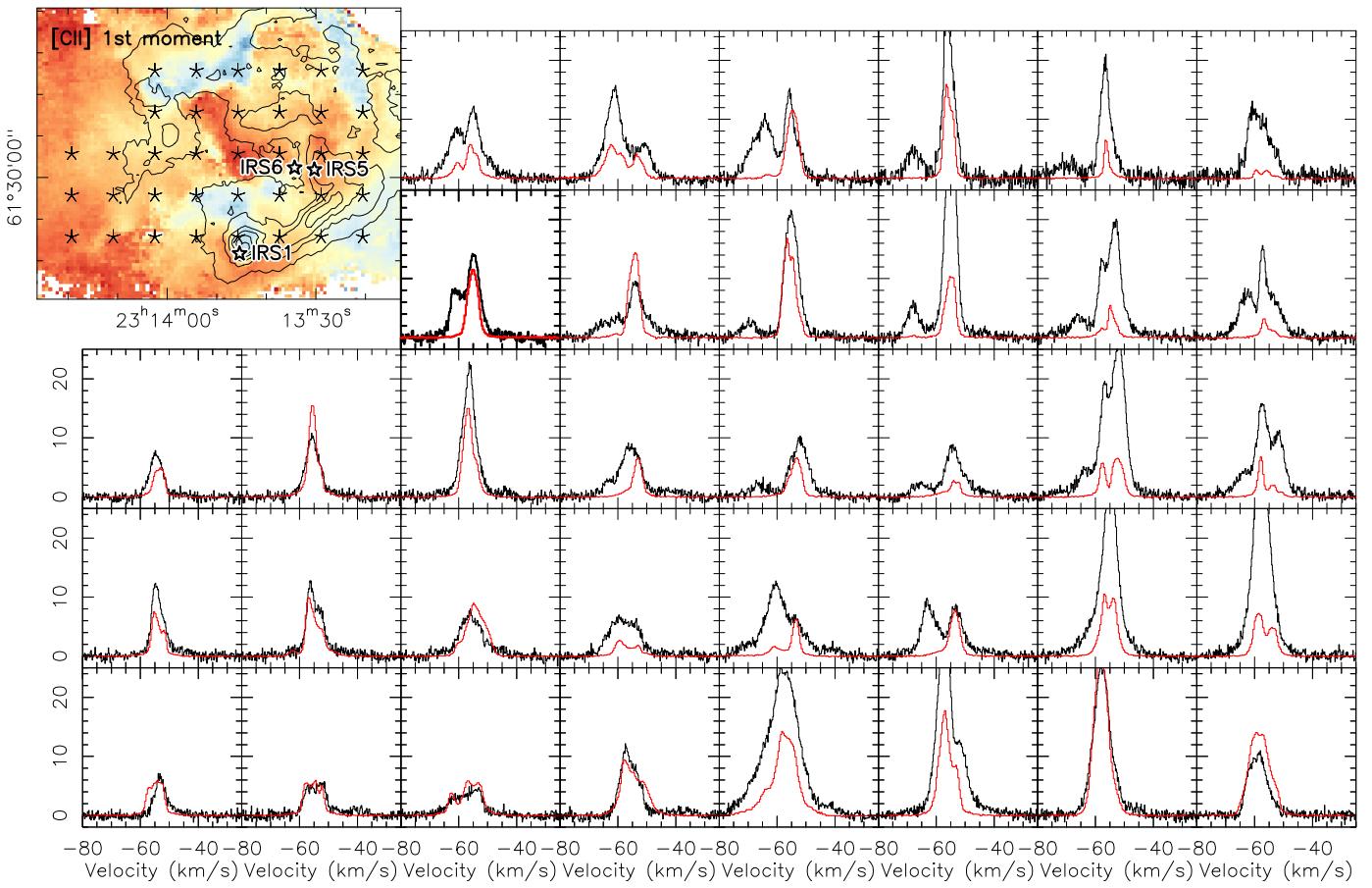
emission due to hot plasma, resulting from the shocked stellar winds from these O stars.

The most blue-shifted [CII] velocity component barely shows any counterpart in atomic or molecular carbon emission. At the interface of the HII region and the molecular cloud we find a bar-like photon-dominated region where ionized C<sup>+</sup>, atomic C<sup>0</sup> and molecular carbon CO reveal a layered structured. The carbon in the PDR is dominantly in its ionized C<sup>+</sup> form. A velocity-gradient across that bar-like PDR is identified, indicating that the PDR is pushed along the line of sight in the direction of the observer.

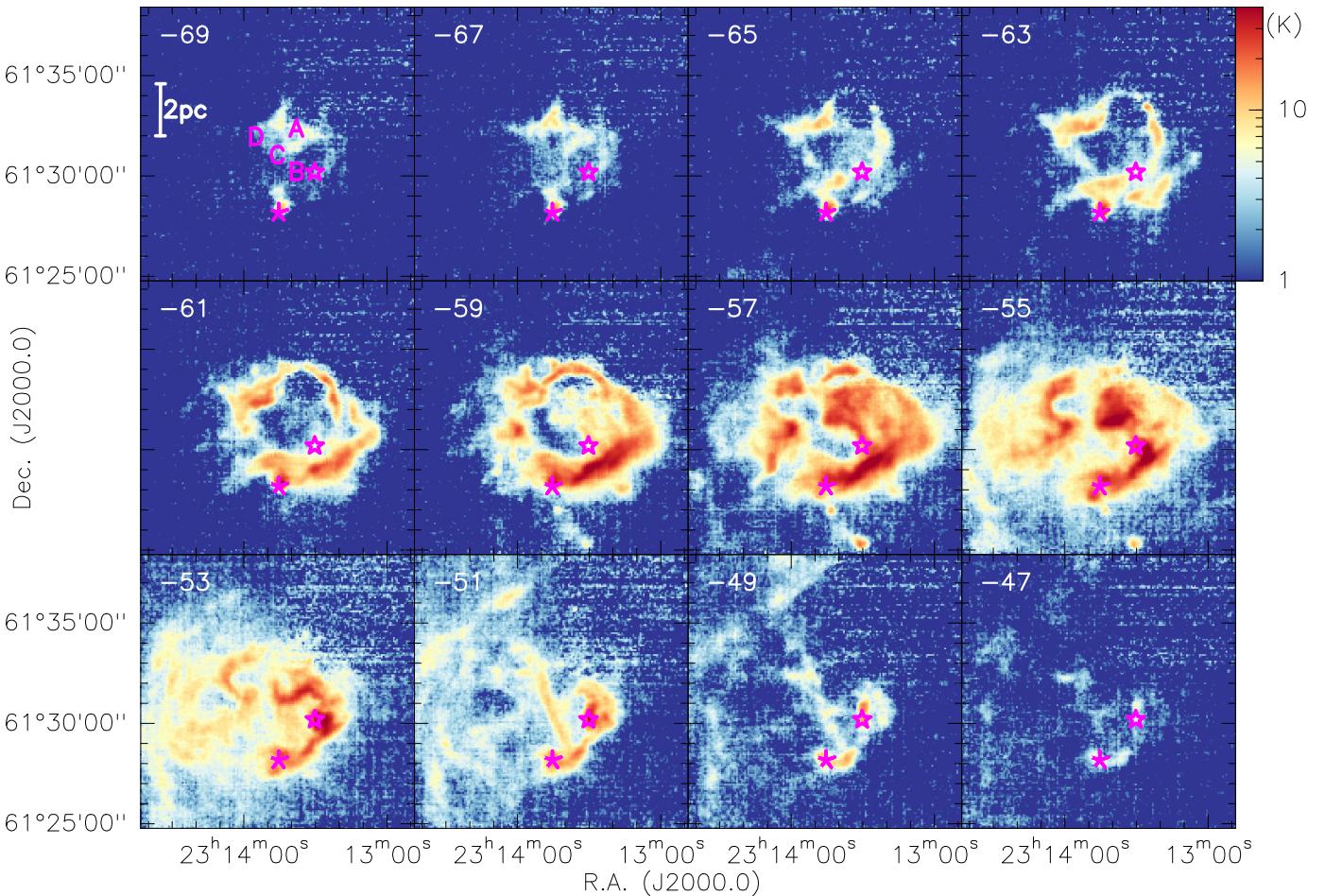
The [CII] study of this prototypical HII region NGC7538 reveals a highly sub-structured HII region that interacts with the adjacent molecular cloud, potentially triggering new star formation events in some of the adjacent dense cores. Setting the region into context with other recently investigated [CII] emission studies from the FEEDBACK program (Orion Veil, RCW120, RCW49), one finds that bubble-like expanding morphologies are recurring structures in PDRs. While in some of these regions star formation occurs at the edges of the expanding regions, indicating positive feedback, the data from the Orion Veil reveal the destructive forces of feedback, expelling and reshaping the gas in that region. Hence, both positive and negative feedback can occur by the expansion of HII regions. The SOFIA legacy program FEEDBACK will dissect more HII regions and reveal common processes in HII region-molecular cloud interaction zones.

*Acknowledgements.* Financial support for the SOFIA Legacy Program, FEEDBACK, at the University of Maryland was provided by NASA through award SOF070077 issued by USRA. The FEEDBACK project is supported by the BMWI via DLR, Projekt Number 50 OR 1916 (FEEDBACK) and Projekt Number 50 OR 1714 (MOBS - MOdellierung von Beobachtungsdaten SOFIA). The Digitized Sky Surveys were produced at the Space Telescope Science Institute under U.S. Government grant NAG W-2166. The images of these surveys are based on photographic data obtained using the Oschin Schmidt Telescope on Palomar Mountain and the UK Schmidt Telescope. The plates were processed into the present compressed digital form with the permission of these institutions. This work is based in part on observations made with the Spitzer Space Telescope, which was operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. We like to thank the referee for a constructive report improving the paper. HB acknowledges support from the European Research Council under the Horizon 2020 Framework Program via the ERC Consolidator Grant CSF-648505. HB also acknowledges support from the Deutsche Forschungsgemeinschaft in the Collaborative Research Center SFB 881 - Project-ID 138713538 - "The Milky Way System" (subproject B1). N.S. acknowledges support by the Agence National de Recherche (ANR/France) and the Deutsche Forschungsgemeinschaft (DFG/Germany) through the project "GENESIS" (ANR-16-CE92-0035-01/DFG1591/2-1). This work was supported by the German Deutsche Forschungsgemeinschaft, DFG project number SFB 956. This project has received funding from the European Research Council (ERC) under the European Union's Horizon 2020 research and innovation programme (Grant agreement No. 851435).

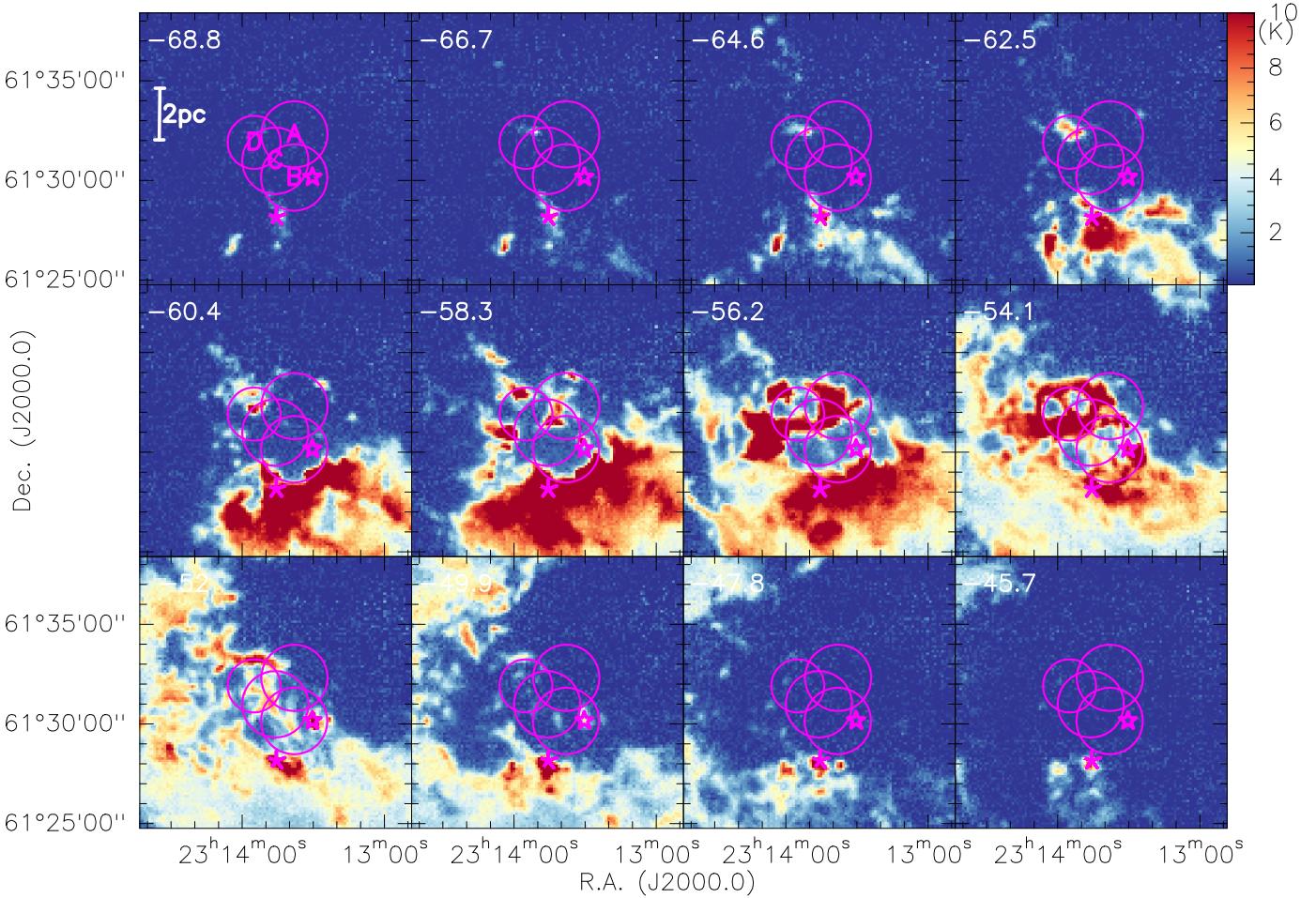
## **Appendix A: Complementary figures**



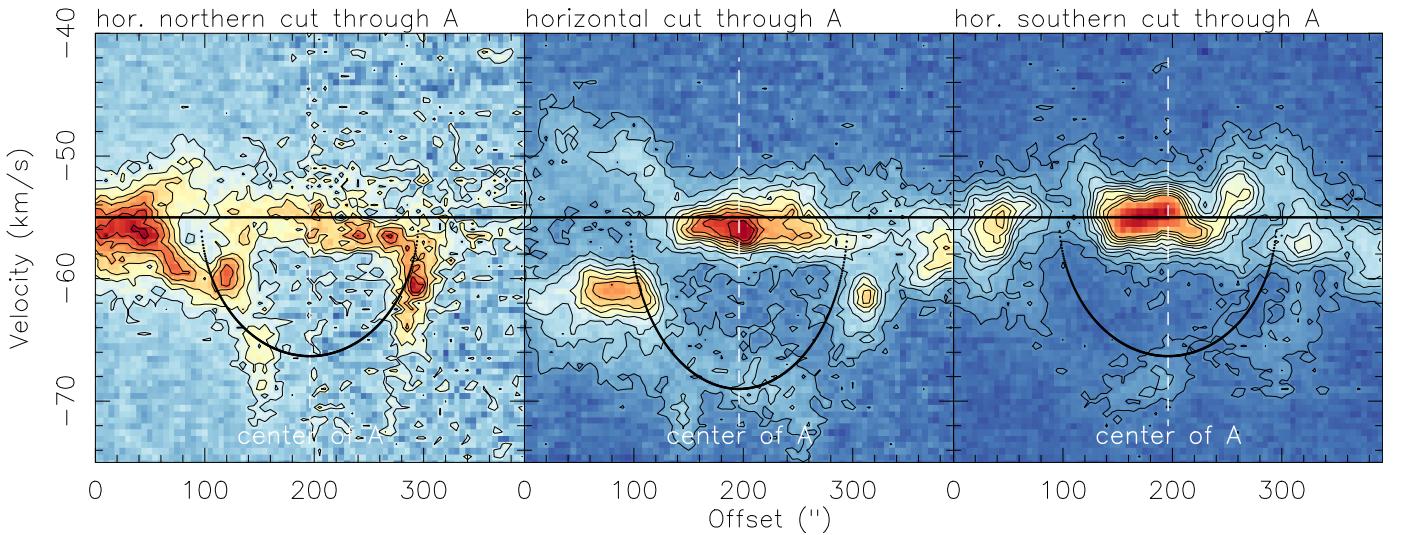
**Fig. A.1.** [CII] (black) and CO(3–2) spectra towards the grid marked in the top-left [CII] 1st moment map. The contours show the [CII] emission at only a few selected contour levels to highlight the main structures. The sources IRS5, IRS6 and IRS1 are also shown as black-and-white stars.



**Fig. A.2.** [CII] channel map for the velocities labeled in each panel. The data are binned for that plot in  $2 \text{ km s}^{-1}$  channels. The two purple stars mark the positions of IRS1 and IRS5, and a scale-bar is shown in the top-left panel. The intensities are always plotted logarithmically from 1 to 40 K.



**Fig. A.3.** CO(3-2) channel map for the velocities labeled in each panel. The data are plotted in steps of  $2.1 \text{ km s}^{-1}$  (5 channels). The two purple stars mark the positions of IRS1 and IRS5, and a scale-bar is shown in the top-left panel. The intensities are always plotted linearly from 0.1 to 10 K. The purple circles outline potential bubbles. These bubbles are labeled in the top-left panel.



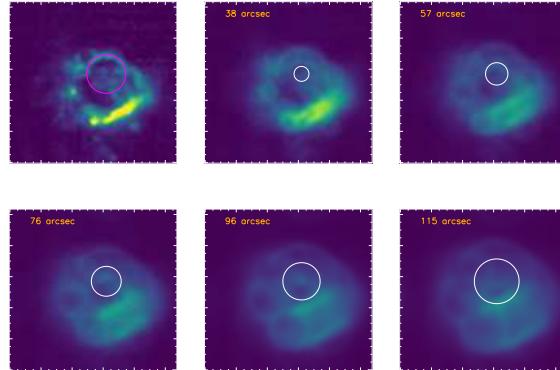
**Fig. A.4.** Horizontal position-velocity cuts through bubble A from east to west derived from the [CII] line emission. The central panel goes through the estimated central position of A whereas the left and right panels are shifted  $1'$  to the north and south, respectively. The dashed line marks the approximate center of the bubble. The black dotted lines in the middle panel outline the elliptical shape a spherical shell with radius  $100''$  and expansion velocity of  $14 \text{ km s}^{-1}$  would have in the pv diagram. The dotted lines in the left and right panel correspond to the same expanding shell at the  $1'$  shift to the east and west where the  $1'$  corresponds to  $\sim 36 \text{ deg}$  offset in the sphere (e.g., Butterfield et al. 2018, sketch in their appendix A.2). Hence the velocities are then reduced by  $\cos(36^\circ)$ .

## Appendix B: Bubble finding

The assignment of the individual structures visible in the channel maps to bubbles and rings with a definite origin can be highly subjective. To avoid biases of an by-eye-identification, we implemented a method to quantify the amount of ring-like structures in individual maps. The goal of the method was to measure structures that one would identify by eye as rings and quantify their significance. This approach is complementary to ring-finding algorithms like CASI (Van Oort et al. 2019) that allow to scan huge amounts of data but only provide a binary detection criterion.

To quantify the match between the map and a ring, we compute the covariance between the map and an annulus structure, that is radially symmetric and positive between an inner and an outer radius and zero everywhere else. We normalize the positive value to provide an integral of unity. The covariance map peaks at the points with the strongest match between the actual structure and the ring. A search for the peaks therefore allows for ring finding, but a simple inspection of the covariance map is insufficient for this because individual peaks in the intensity distribution translate linearly into the covariance coefficients. A reliable quantification therefore requires comparison of the match with different structures. We perform this through a parameter scan comparing to structures with different ring radii and widths.

To be consistent with the typical by-eye-identification we limited the ring width by a ratio between outer and inner radius between 1.1 and 1.6 and scanned this range. However, when applied to the data here, we found only a weak dependence of the results on this ratio and usually a good match for ratios around 1.4, so that we limit the further discussion to rings of this width to reduce the dimensionality of the problem. Furthermore, we scanned possible ring radii between roughly 38 and 163''.



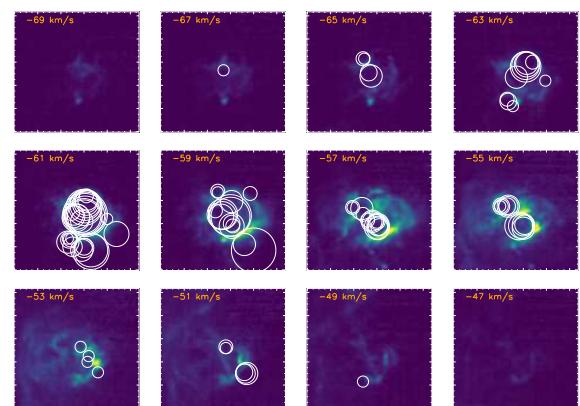
**Fig. B.1.** Demonstration of the covariance maps for the convolution of the  $-61 \text{ km s}^{-1}$  channel map (see Fig. 6) with rings of different radii. The first panel shows the channel map. The magenta circle emphasizes one of the obvious ring-like structures with a radius of 100''. The other five panels show the covariance maps from the convolution of rings with increasing radii. The radii are indicated at the same position in the map.

Fig. B.1 demonstrates this for the [CII] channel map at  $-61 \text{ km s}^{-1}$  from Fig. 6. The first panel shows the channel map emphasizing one of the prominent arcs with a radius of  $\sim 100''$ . The other panels show the covariance maps from the convolution with rings of 38, 57, 76, 96, and 115''. They are indicated as thin circles centered at the same position. One can see that for

small radii the covariance maps show a dip in the centre of the circles, while only for the 96'' radius, the map shows a peak in the centre. Considering only the location at the center of the circle and scanning the ring radii we find the maximum covariance coefficient for the “correct” ring radius. However, it is interesting to see that the covariance map for 76'' rings also shows a peak in the vicinity, shifted to the north by about 20''. This indicates that the arc in the channel map could also be fitted with a smaller ring that is spatially shifted. From these covariance maps only, neither of the two solutions can be excluded.

Additional information can be obtained when including the velocity information if we do not only look for rings, but the physical scenario of an expanding bubble. In that case, only at the systemic zero velocity the bubble appears as a ring with its full diameter. At lower and higher velocities, the channel maps perform a cut through the bubble in front or behind the central plane so that a smaller broadened ring appears. A different way to visualize this is via plotting the covariance coefficients against the velocity and the ring radius. For an ideal bubble, one then finds an inverse “V” shape where the largest ring radius is detected at zero systemic velocities. At intermediate velocities one finds a smaller ring.

To look for bubbles in the NGC7538 [CII] data we can look for this inverse “V” shape in the covariance coefficients computed for all channel maps and ring radii. We perform a pre-selection of pixels by looking for rings in the individual channel maps. The principle is straight forward. We search for peaks on the covariance maps that are small, isotropic and sufficiently bright (for details see Kong et al. in prep.). Here, we perform a GAUCSCLUMPS (Stutzki & Guesten 1990) decomposition of the covariance maps and locate clumps that have an axis ratio of less than two, a radius below the inner ring radius of the convolution and an intensity of more than 5 K in the [CII] channel maps.



**Fig. B.2.** [CII] channel map with an overlay of all ring candidates identified in the individual channels. We plotted the rings not with their full width used in the convolution (outer radius 1.4 times the inner radius) but only as very thin rings to allow for a distinction of the different radii by eye.

The result is shown in Fig. B.2 for all [CII] channel maps. The number of ring candidates varies strongly between the different channels. At the extreme velocities, no rings are found, but the largest number of rings is not found at the velocities of the brightest intensity but in panel 5, for a velocity of  $-61 \text{ km s}^{-1}$ .

In many cases the radius of the ring is not accurately confined. We find overlapping rings with nearby centres. This is in particular the case for panel 5 ( $-61 \text{ km s}^{-1}$ ) where five different radii can fit the central ring structure. A typical effect is best visible in panel 10, at  $-51 \text{ km s}^{-1}$ , where an arc like structure can be fitted by rings of different radii and simultaneously different centres. This demonstrates the fundamental uncertainty of the approach. Based on the plots here, it is only possible to measure the location of the center of each ring if we a priori know its radius. Fitting both leaves some uncertainty.

The locations of all these rings gives a guidance where to look for the inverse “V” structure in the four-dimensional covariance cube. If we look at the distribution of the different ring centres of the different channel maps (Fig. B.2) we find three dominant clusters that involve rings of varying radii

(A) R.A.(J2000.0)=23h13m38s, Dec=61°32'20"

(B) R.A.(J2000.0)=23h13m38s, Dec=61°30'12"

(C) R.A.(J2000.0)=23h13m46s, Dec=61°31'00"

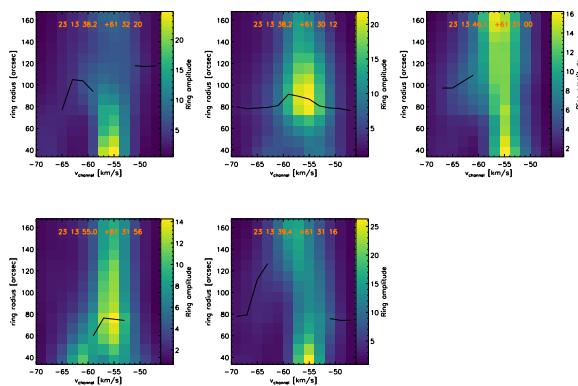
A cluster with only small ring radii is located around

(D) R.A.(J2000.0)=23h13m55s, Dec=61°31'56"

and a cluster with only large radii around

(E) R.A.(J2000.0)=23h13m39s, Dec=61°31'16".

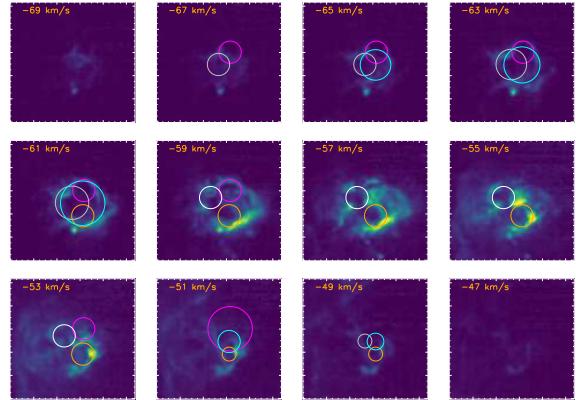
From the inspection of Fig. B.2 we can exclude most other positions because they only detected some outer arcs, not driven from inside of the region. In addition to that, the largest ring appears as an overlap of the other smaller rings, in particular A, B and C (Figs. B.2 & B.4). Therefore, although identified by the algorithm, we do not consider E as a separate physical entity.



**Fig. B.3.** Covariance coefficients at the positions of the five prominent clusters from Fig. B.2. The coefficients are shown as a function of the channel velocity and the ring radius used in the convolution. The black lines mark the maxima of the covariance coefficients in terms of the ring radius. In those channels without black line, no ring could be identified.

Fig. B.3 shows the covariance coefficients for these five locations as a function of channel velocity and ring radius. The relevant information is the location and width of the peak in the coefficients as a function of the ring radius, in particular the shape of the line formed by those peaks as a function of the velocity channel. All panels show somewhat different characteristics but also show at least some indication of the inverse “V” shape. Fig. B.3 shows as a black line the maxima of the covariance coefficients that can be considered as a proxy for the ring/bubble radius. The width of these maxima can be considered as an estimator on the uncertainty on the ring radius. The “V” shape is

most pronounced for structure D (bottom-left panel in Fig. B.3). There the maximum ring radius is  $\sim 80''$  at velocities of  $-57$  and  $-55 \text{ km s}^{-1}$  and we find dropping radii to higher and lower velocities. For the other bubble candidates, the transition to smaller radii and lower velocity is less well determined. In the following, we use as approximate radii for structures A, B and C  $100''$ , and for structure D  $80''$ .



**Fig. B.4.** [CII] channel map with an overlay of the identified rings around the five candidate positions from Fig. B.3. The ring radii vary with velocity. The plotted rings are narrower than the rings used in the convolution. Different ring positions are encoded in a different ring color. The positions in panels 1–5 from Fig. B.3 correspond to the colors magenta (A), orange (B), grey (C), white (D), and cyan (E).

Figure B.4 shows the same result in a different way. Here we plot again the channel map from Fig. 6 but now over-plot circles giving the mean ring diameter for the ring that is the local peak in the covariance map for the five points discussed in Fig. B.3. This means for each channel we select the corresponding velocity column in the covariance plots in Fig. B.3, look for the maximum and draw the corresponding circle if the maximum is not at the minimum radius and above 4 % of the peak intensity. The five positions are color coded. The systematic behaviour of an increasing ring diameter with velocity is well visible for the grey and cyan rings representing position C and E. They behave like the ideal bubble at least for blue-shifted velocities. In contrast the magenta and white rings, positions A and D, are relatively constant in diameter. They rather represent ring-structures in ppv-space. The matches at a velocity of  $-51 \text{ km s}^{-1}$  rather look, at least partially, like a coincidence, not necessarily connected to the rings or bubbles at lower velocities. This all indicates that a bubble expansion in the region was hindered towards red-shifted velocities. It happens mainly at blue-shifted velocities, this means towards the observer.

To estimate the accuracy of the determination of the ring positions and radii, we investigated the covariance plots for varying central position of the identified bubbles. This analysis indicates that for the conditions in NGC7538 with ring or bubble radii on the order of  $60\text{--}160''$  we can pin down their origin with an accuracy of  $20\text{--}30''$ . The bubble diameter should be the largest one found in the central velocity channel. If we consider the rings from Fig. B.4 and compare the corresponding ring candidates in Fig. B.2 we see that in those velocity channels typically two or three large rings appear corresponding to an uncertainty of the ring radius of  $20\text{--}30''$  (see also lines in Fig. B.3).

## References

- Beuther, H., Linz, H., & Henning, T. 2012, A&A, 543, A88
- Beuther, H., Mottram, J. C., Ahmadi, A., et al. 2018, A&A, 617, A100
- Beuther, H., Ragan, S. E., Ossenkopf, V., et al. 2014, A&A, 571, A53
- Butterfield, N., Lang, C. C., Morris, M., Mills, E. A. C., & Ott, J. 2018, ApJ, 852, 11
- Cabrit, S., Goldsmith, P. F., & Snell, R. L. 1988, ApJ, 334, 196
- Drew, J. E., Greimel, R., Irwin, M. J., et al. 2005, MNRAS, 362, 753
- Elmegreen, B. G. 2011, in EAS Publications Series, Vol. 51, EAS Publications Series, ed. C. Charbonnel & T. Montmerle, 45–58
- Fallscheer, C., Reid, M. A., Di Francesco, J., et al. 2013, ApJ, 773, 102
- Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, ApJS, 154, 10
- Frerking, M. A., Keene, J., Blake, G. A., & Phillips, T. G. 1989, ApJ, 344, 311
- Geen, S., Hennebelle, P., Tremblin, P., & Rosdahl, J. 2016, MNRAS, 463, 3129
- Geen, S., Watson, S. K., Rosdahl, J., et al. 2018, MNRAS, 481, 2548
- Gerlich, D. 1990, J. Chem. Phys., 92, 2377
- Gibson, S. J., Taylor, A. R., Higgs, L. A., Brunt, C. M., & Dewdney, P. E. 2005, ApJ, 626, 214
- Goicoechea, J. R., Pabst, C. H. M., Kabanovic, S., et al. 2020, A&A, 639, A1
- Goldsmith, P. F., Langer, W. D., Pineda, J. L., & Velusamy, T. 2012, ApJS, 203, 13
- Guan, X., Stutzki, J., Graf, U. U., et al. 2012, A&A, 542, L4
- Güdel, M., Briggs, K. R., Montmerle, T., et al. 2008, Science, 319, 309
- Henshaw, J. D., Krumholz, M. R., Butterfield, N. O., et al. 2021, arXiv e-prints, arXiv:2110.11367
- Hollenbach, D. J. & Tielens, A. G. G. M. 1997, ARA&A, 35, 179
- Hopkins, P. F., Kereš, D., Oñorbe, J., et al. 2014, MNRAS, 445, 581
- Kendrew, S., Beuther, H., Simpson, R., et al. 2016, ApJ, 825, 142
- Kim, J.-G., Kim, W.-T., & Ostriker, E. C. 2018, ApJ, 859, 68
- Klein, B., Hochgürtel, S., Krämer, I., et al. 2012, A&A, 542, L3
- Kraus, S., Balega, Y., Elitzur, M., et al. 2006, A&A, 455, 521
- Lancaster, L., Ostriker, E. C., Kim, J.-G., & Kim, C.-G. 2021a, ApJ, 914, 89
- Lancaster, L., Ostriker, E. C., Kim, J.-G., & Kim, C.-G. 2021b, ApJ, 914, 90
- Lancaster, L., Ostriker, E. C., Kim, J.-G., & Kim, C.-G. 2021c, arXiv e-prints, arXiv:2110.05508
- Langer, W. D., Velusamy, T., Pineda, J. L., et al. 2010, A&A, 521, L17
- Le Bourlot, J. 1991, A&A, 242, 235
- Lique, F., Werfelli, G., Halvick, P., et al. 2013, J. Chem. Phys., 138, 204314
- Luisi, M., Anderson, L. D., Balser, D. S., Bania, T. M., & Wenger, T. V. 2016, ApJ, 824, 125
- Luisi, M., Anderson, L. D., Schneider, N., et al. 2021, Science Advances, 7, eabe9511
- Mac Low, M.-M. & McCray, R. 1988, ApJ, 324, 776
- Matzner, C. D. 2002, ApJ, 566, 302
- Moscadelli, L., Reid, M. J., Menten, K. M., et al. 2009, ApJ, 693, 406
- Motte, F., Zavagno, A., Bontemps, S., et al. 2010, A&A, 518, L77+
- Muijres, L. E., Vink, J. S., de Koter, A., Müller, P. E., & Langer, N. 2012, A&A, 537, A37
- Pabst, C., Higgins, R., Goicoechea, J. R., et al. 2019, Nature, 565, 618
- Pabst, C. H. M., Goicoechea, J. R., Teyssier, D., et al. 2020, A&A, 639, A2
- Palmeirim, P., André, P., Kirk, J., et al. 2013, A&A, 550, A38
- Palmeirim, P., Zavagno, A., Elia, D., et al. 2017, A&A, 605, A35
- Pérez-Beaupuits, J. P., Güsten, R., Spaans, M., et al. 2015a, A&A, 583, A107
- Pérez-Beaupuits, J. P., Stutzki, J., Ossenkopf, V., et al. 2015b, A&A, 575, A9
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, A&A, 518, L1
- Poglitsch, A., Waelkens, C., Geis, N., et al. 2010, A&A, 518, L2
- Puga, E., Marín-Franch, A., Najarro, F., et al. 2010, A&A, 517, A2
- Qiu, K., Zhang, Q., & Menten, K. M. 2011, ApJ, 728, 6
- Reid, M. J., Menten, K. M., Brunthaler, A., et al. 2019, ApJ, 885, 131
- Risacher, C., Güsten, R., Stutzki, J., et al. 2018, Journal of Astronomical Instrumentation, 7, 1840014
- Rohlfs, K. & Wilson, T. L. 2006, Tools of radio astronomy (Tools of radio astronomy, 4th rev. and enl. ed., by K. Rohlfs and T.L. Wilson. Berlin: Springer, 2006)
- Röllig, M. & Ossenkopf, V. 2013, A&A, 550, A56
- Russeil, D., Schneider, N., Anderson, L. D., et al. 2013, A&A, 554, A42
- Sandell, G., Wright, M., Güsten, R., et al. 2020, ApJ, 904, 139
- Schneider, N., Simon, R., Guevara, C., et al. 2020, PASP, 132, 104301
- Schöier, F. L., van der Tak, F. F. S., van Dishoeck, E. F., & Black, J. H. 2005, A&A, 432, 369
- Scoville, N. Z., Sargent, A. I., Sanders, D. B., et al. 1986, ApJ, 303, 416
- Sharma, S., Pandey, A. K., Ojha, D. K., et al. 2017, MNRAS, 467, 2943
- Simon, R., Stutzki, J., Sternberg, A., & Winnewisser, G. 1997, A&A, 327, L9
- Spitzer, L. 1998, Physical Processes in the interstellar medium (John Wiley and Sons, Inc.)
- Stutzki, J. & Guesten, R. 1990, ApJ, 356, 513
- Syed, J., Wang, Y., Beuther, H., et al. 2020, A&A, 642, A68
- Tauber, J. A., Lis, D. C., Keene, J., Schilke, P., & Buettgenbach, T. H. 1995, A&A, 297, 567
- Tauber, J. A., Tielens, A. G. G. M., Meixner, M., & Goldsmith, P. F. 1994, ApJ, 422, 136
- Taylor, A. R., Gibson, S. J., Peracaula, M., et al. 2003, AJ, 125, 3145
- Tielens, A. G. G. M., Meixner, M. M., van der Werf, P. P., et al. 1993, Science, 262, 86
- Tiwari, M., Karim, R., Pound, M. W., et al. 2021, ApJ, 914, 117
- Townsley, L. K., Broos, P. S., Garmire, G. P., et al. 2018, ApJS, 235, 43
- Ungerechts, H., Umbanhowar, P., & Thaddeus, P. 2000, ApJ, 537, 221
- Van Oort, C. M., Xu, D., Offner, S. S. R., & Gutermuth, R. A. 2019, ApJ, 880, 83
- Weaver, R., McCray, R., Castor, J., Shapiro, P., & Moore, R. 1977, ApJ, 218, 377
- Werner, M. W., Roellig, T. L., Low, F. J., et al. 2004, ApJS, 154, 1
- Wilson, T. L. & Rood, R. 1994, ARA&A, 32, 191