

Problem 1

The spectrum would depend on the angle between the observer and the rotational axis of the AGN. Narrow emission lines come from lower-velocity material that is more distant from the center of the AGN, and are visible regardless of whether the angle is large or small. Broad emission lines are caused by Doppler broadening due to the higher velocities of material closer to the center. These are only visible if the angle is small, because at larger angles, the torus of dust/gas surrounding the center blocks the view of the center.

Problem 2

Radio-loud AGN have jets/lobes, while radio-quiet AGN do not. These jets are caused by relativistic particles accelerated out of the center of the AGN. The particles emit synchrotron radiation because of strong magnetic fields, and synchrotron emissions peak in the radio regime.

Problem 3

In free-fall,

Problem 4

- (i) The rest wavelength of the $H\alpha$ emission line in a vacuum (?) is $6564.614 \text{ \AA} = 0.656 \mu\text{m}$ ¹. The redshift of the source is

$$z = \frac{4 \mu\text{m} - 0.656 \mu\text{m}}{0.656 \mu\text{m}} = 5.01.$$

- (ii) We can relate the velocity dispersion to the width of the emission line using the following equation:

$$\frac{\Delta\lambda}{\lambda} = \frac{\Delta v}{c}.$$

Rearranging for Δv and plugging in the appropriate values, we find that Δv is $\Delta\lambda$ supposed to be the width (0.025)? or the difference (4-0.656) from last q?:

$$\Delta v = \frac{\Delta\lambda}{\lambda} \times c = \frac{0.025 \mu\text{m}}{0.656 \mu\text{m}} \times 2.99 \times 10^8 \text{ m s}^{-1} = \text{m s}^{-1}.$$

Yes, this is fairly consistent with a 2000 km s^{-1} velocity dispersion.

¹<https://classic.sdss.org/dr3/products/spectra/vacwavelength.html>

Problem 5

The mass of Sgr A* is approximately $4 \times 10^6 M_\odot$ ². The Schwarzschild radius is

$$R_s = \frac{2GM}{c^2} = \frac{2 \times (6.67 \times 10^{-11} \text{ m}^3 \text{ kg}^{-1} \text{ s}^{-2}) \times (4 \times 10^6 M_\odot) \times (1.989 \times 10^{30} \text{ kg } M_\odot^{-1})}{(2.99 \times 10^8 \text{ m s}^{-1})^2} = 1.18 \times 10^{10} \text{ m}.$$

The distance between the Sun and the center of the Milky Way is $8 \text{ kpc} = 2.47 \times 10^{20} \text{ m}$. We can use this, along with the small angle approximation, to convert the Schwarzschild radius of Sgr A* to an angular size (which is the resolution required):

$$\theta \simeq \tan \theta = \frac{R_s}{d} = \frac{1.18 \times 10^{10} \text{ m}}{2.47 \times 10^{20} \text{ m}} = 4.78 \times 10^{-11} \text{ rad} \simeq 0.01 \text{ mas}.$$

If VLBI is used and the effective diameter of the "telescope" is the diameter of the Earth, then the wavelength that would have to be used is

$$\begin{aligned} \theta &= \frac{1.22\lambda}{2R_E} \\ \Rightarrow \lambda &= \frac{\theta \times 2R_E}{1.22} \\ &= \frac{4.78 \times 10^{-11} \text{ rad} \times 2 \times 6.378 \times 10^6 \text{ m}}{1.22} \\ &= 0.0005 \text{ m}. \end{aligned}$$

²<https://arxiv.org/pdf/1607.05726.pdf>