

## Die Friedmann–Gleichung

$$H(t)^2 = \left(\frac{\dot{a}}{a}\right)^2 = H_0^2 \left( \frac{\Omega_r}{a^4} + \frac{\Omega_m}{a^3} + \Omega_\Lambda - \frac{\Omega_0 - 1}{a^2} \right)$$

- a)
- $\Omega_0 \dots$  total density
  - $\Omega_r \dots$  radiation density
  - $\Omega_m \dots$  matter density (Dark + Baryonic)
  - $\Omega_\Lambda \dots$  cosmological constant (vacuum density)

b)

$$H(t)^2 = \left(\frac{\dot{a}}{a}\right)^2 = \left(\frac{\frac{da}{dt}}{a}\right)^2 \Rightarrow a^2 H^2(t) = \left(\frac{da}{dt}\right)^2 \Rightarrow da = \frac{dt}{aH(t)}$$

$$\int 1 dt = \int \frac{1}{aH(a)} da$$

- c)
- $\Omega_0 \propto a^{-2}$
  - $\Omega_r \propto a^{-4}$
  - $\Omega_m \propto a^{-3}$
  - $\Omega_\Lambda \propto$

d) Radiation dominated:  $H(t) = H_0 \sqrt{\frac{1}{a(t)^4}}$

$$\int_0^t 1 d\tilde{t} = \int_0^{a(t)} \frac{1}{\tilde{a}H(\tilde{a})} d\tilde{a} = \int_0^{a(t)} \frac{\tilde{a}}{H_0} d\tilde{a} = \frac{a(t)^2}{2H_0}$$

$$a(t) = \underline{\underline{\sqrt{2H_0 t}}}$$

Matter dominated:  $H(t) = H_0 \sqrt{\frac{1}{a(t)^4}}$

$$\int_0^t 1 \, d\tilde{t} = \int_0^{a(t)} \frac{1}{\tilde{a} H(\tilde{a})} \, d\tilde{a} = \int_0^{a(t)} \frac{\sqrt{\tilde{a}}}{H_0} \, d\tilde{a} = \frac{2\sqrt{a(t)^3}}{3H_0}$$

$$a(t) = \underline{\underline{\left(\frac{3H_0 t}{2}\right)^{\frac{2}{3}}}}$$

Cosmological Constant:  $H(t) = H_0$

$$\int_0^t 1 \, d\tilde{t} = \int_c^{a(t)} \frac{1}{\tilde{a} H(\tilde{a})} \, d\tilde{a} = \int_c^{a(t)} \frac{1}{\tilde{a} H_0} \, d\tilde{a} = \frac{\ln(a(t))}{H_0}$$

$$a(t) = \underline{\underline{e^{H_0 t + c}}}$$

Curvature dominated:  $H(t) = H_0 \sqrt{\frac{1}{a(t)^2}}$

$$\int_0^t 1 \, d\tilde{t} = \int_0^{a(t)} \frac{1}{\tilde{a} H(\tilde{a})} \, d\tilde{a} = \int_0^{a(t)} \frac{1}{H_0} \, d\tilde{a} = \frac{a(t)}{H_0}$$

$$a(t) = \underline{\underline{H_0 t}}$$

