

Signatures of UV radiation around low-mass protostars in the Serpens Main with IRAM 30m

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ABSTRACT

Context. The Serpens Main is one of the most studied star forming region containing low-mass protostars. Observations at submillimetre range allow to determine physical and chemical processes around young stellar objects.

Aims. We aim to characterise the UV radiation in the surroundings of the low-mass protostars. We analyse the excitation and spatial extent of HCN, CN, CS and their isotopologues to identify the underlying processes. We can investigate the feedback from protostars and the excitation mechanisms of molecules.

Methods. We present ~ 30 arcmin² IRAM 30m maps of CN $J = 1 - 0$, HCN $J = 1 - 0$, and CS $J = 3 - 2$ encompassing 10 Class 0/I protostars. We calculate HCN and CN column densities toward protostars and selected outflows positions. The column densities are compared with the Nahoon astrochemical model of molecules abundances in order to characterise UV radiation field.

Results. Emission of HCN $J = 1 - 0$ and CS $J = 3 - 2$ is co-spatial with outflows, whereas CN emission peaks at the positions of protostars. CN and HCN column densities are of the order of 10^{13} - 10^{14} cm⁻². Regardless of gas parameters, CN/HCN column density ratio is 1-10. This result can be reproduced by providing an additional UV radiation source of 0.001 to 0.044 G_0 .

Conclusions. The UV radiation field is significantly higher in the closest distances from protostars. The astrochemical model shows that an additional source of UV radiation is needed to cover the abundances range indicated by observations.

Key words. astrochemistry – stars: formation – ISM: molecules – ISM: individual objects: Serpens Main – Submillimeter: ISM

1. Introduction

Low-mass stars are the most numerous objects among stellar population in galaxies (Kroupa 2002). At first stages of star formation protostars are formed inside molecular cloud, surrounded by massive envelopes exceeding for 10^4 AU in diameter (Lada 1987; Larson 2003; Bergin & Tafalla 2007). The embedded phases of low-mass star formation (Class 0/I YSOs, Andre et al. 1993) is characterised by gas and dust accretion from an envelope and bipolar, collimated outflows which carry out molecular gas from the dense core (Zuckerman et al. 1976, Arce & Sargent 2006).

Molecular outflows properties are typically traced by low-J CO transitions (Bontemps et al. 1996, Dunham et al. 2014). Herschel-PACS observations of the high-J CO emission around low-mass protostars showed the presence of a warm component about $T_{\text{rot}} \sim 300$ K, hot component of $T_{\text{rot}} \sim 600 - 800$ K (Karska et al. 2013, Green et al. 2013) and very hot component exceeding 1000 K for some sources (Manoj et al. 2013). The high-temperature CO emission can be explained as shocked-heated gas created by UV-irradiated shocks (Kristensen et al.

2017). Shock waves are created in the inner part of accretion disk where the gas is heated by the 10,000 K radiation field (Spaans et al. 1995). UV photons propagate in low-density outflow cavities for large distances, modifying the chemical composition of low-mass neighbourhood.

Previous studies of energetic processes around low-mass protostars showed observational premises of the influence of UV radiation on molecules. At protostars positions was observed higher emission of [C I] than in off-source positions what brought to conclusion of photodissociation CO (van Kempen et al. 2009). The fluxes of [OI] and [CII] are significantly higher than fully-shielded C-shock models (Karska et al. 2018). The other photodissociation tracer H₂O/OH showed a few order of magnitude disagreement with shock models (Karska et al. 2014). Therefore, ultraviolet radiation may play an important role in low-mass protostars surroundings.

The relative abundance of CN and HCN molecules is widely used a tracer of UV radiation in different astronomical context (e.g. Fuente et al. 1995, Chapillon et al. 2012, Riaz et al. 2018). CN is a product of photodissociation of HCN with the photodis-

sociation rate of 1.64×10^{-9} . CN has smaller photodissociation rate of 5.19×10^{-10} (Heays et al. 2017), thus is not that sensitive for photodissociation as HCN. Since CN and HCN can be photodissociated selectively, therefore the CN/HCN ratio probes regions affected by UV radiation. The ratio is the highest near the source of the UV emission, and decreases with the distance from the source (Fuente et al. 1993). Tracing the physical and chemical processes dominated by UV radiation lead to better understanding of the protostellar evolution. Ultraviolet radiation can propagate in large scales of 1000 AU (Kristensen et al. 2017), changing the properties of the surrounding matter. Since the UV radiation around low-mass protostars was studied only in the source position (Stäuber et al. 2007, Riaz et al. 2018) or in the source closest neighbourhood (Hogerheijde et al. 1999, Bachiller et al. 2001, Jørgensen 2004), the spatial extend of UV fields in larger scales is the matter of question. We address the following questions in our paper: How does the UV radiation affect the low-mass surrounding? Where the molecules are dissociated? What is the spatial extend of the UV fields? What is the typical UV radiation around Class 0/I protostars?

We present the CN, HCN, CS and their isotopologues molecular data observed in the Serpens Main star forming region. Section 2 contains the overview of the observations and the targeted sample. The results derived from the observations are shown in the Section 3, while further analysis of the data in Section 4. In the section 5 we refer to the previous studies of the topic and provide the discussion of the results. The summarising Section 6 is closed with our conclusions.

2. Observations

2.1. IRAM data and reduction process

The Serpens Main star forming region was observed with IRAM 30 between 14 and 17 July 2009 (project no. xxx, PI: L. Kristensen). We used the Eight Mixer Receiver (EMIR) as the frontend. The observations were performed in the EMIR bands E090 (molecule HCN $J = 1 - 0$) covering the range 73-117 GHz and E150 (molecules CN $J = 1 - 0$ and CS $J = 3 - 2$) covering the frequencies between 125 and 184 GHz. Due to the EMIR receiver wide bands additional molecular lines of $C^{13}S$ $J = 3 - 2$, $H^{13}CN$ $J = 1 - 0$ and $H^{13}CN$ $J = 2 - 1$ were also observed. The backend was the Versatile SPectrometer Array (VESPA) autocorrelator and the 1 MHz filterbank reaching the spectral resolution of 39 kHz (E150 band) and 78 kHz (E090 band). The telescope beam size varies from $14''$ at 172.68 GHz to $29''$ at 86.34 GHz (Table 1). The antenna temperatures were converted to main-beam brightness temperature T_{MB} using the main beam efficiency according to the expression: $T_{MB} = T_A/\eta_{MB}$. The exact upper levels energies, line frequencies, beam sizes and beam efficiencies are given in Table 1. Observations included scans of the Ser-SMM1 (centered at $\alpha_{J2000} = 18^h29^m49.6^s$, $\delta_{J2000} = +01^\circ15'20.5''$ with $V_{LSR} = +8.5$ km/s) and the Ser-SMM3/Ser-SMM4 (centered at $\alpha_{J2000} = 18^h29^m56.6^s$, $\delta_{J2000} = +01^\circ14'00.3''$ with $V_{LSR} = +7.6$ km/s) regions, both $1' \times 3'$ OTF maps. The size of the maps is about $300'' \times 350''$, covering both Ser-SMM1 and Ser-SMM3/Ser-SMM4 regions. The regions are referenced in the article as 'the Northern part' and 'the Southern part' respectively.

Data reduction was carried out with the CLASS package within GILDAS¹. Each spectrum was corrected for the baseline shape, the spike channels were removed and the velocity was re-sampled to a resolution of 0.5 km/s. The baseline fitting of the

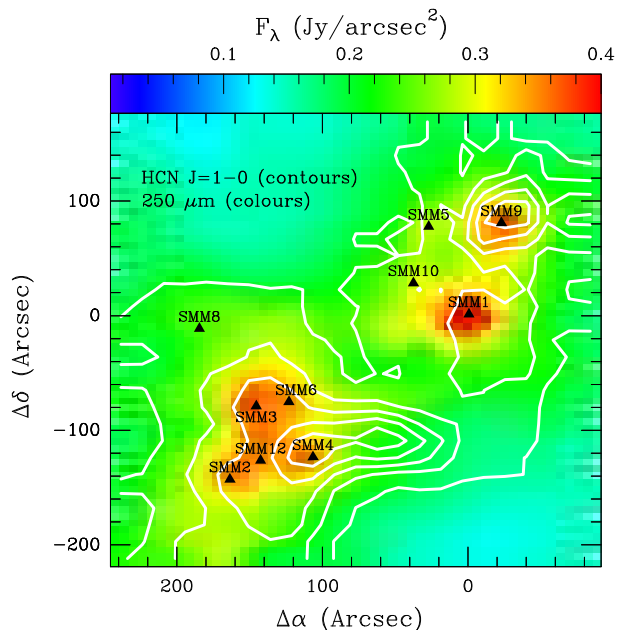


Fig. 1: Molecular emission of HCN $J=1-0$ (contours) overplot on continuum emission *Herschel*/SPIRE (Griffin et al. 2010) $250 \mu m$ (colours). Contours the lowest level is set on 0.4 K km/s (30σ), step size of 4 K km/s.

order of 0 was sufficient for our observations. The rms of extracted spectra values vary from 0.024 K to 0.125 K. Both OTF maps were merged in one map covering 300×350 arcsec. The spectra obtained were exported from the CLASS package and analysed with Python scripts.

2.2. Continuum emission

2.3. Physical properties of embedded protostars

Ten Class 0/I protostars are present in the observed region. There are deeply embedded sources so the radiation coming from their neighbourhood is highly absorbed in the envelopes, then re-emitted in the IR range. Envelopes become thinner with time due to outflow-envelope interactions (Arce & Sargent 2006). Class I sources SEDs are dominated by the emission in shorter wavelengths in respect to Class 0 objects. Thus Spectral Energy Distributions (SEDs) allow to estimate the evolutionary stage of an object (Andre et al. 1993).

Figure A.1 shows spectral energy distributions for all Class 0/I protostars in the region (Table 2). The SED plots include the selected literature samples (Dunham et al. 2015) combined with the data from the Herschel Gould Belt survey project (André et al. 2010). The additional Herschel data cover the SED peak, therefore provide a more detailed information allowing to calculate the bolometric temperatures and luminosities of the protostar more precisely.

Table 2 contains the observed protostars parameters as well as the classification from Enoch et al. 2009. Early Class 0 was defined as a protostar of bolometric temperature lesser than 50 K. Protostars characterised by bolometric temperature between 50 K and 100 K were classified as Late Class 0 protostars. Class I protostars were divided for Early and Late sub-type by the bolometric temperature of 300 K.

Most of the observed protostars in the Serpens Main region are very young, embedded sources of Early Class 0. SMM4,

¹ See <http://www.iram.fr/IRAMFR/GILDAS>

Table 1: Overview of the observations

Mol.	Trans.	ν (GHz)	E_u/k_B (K)	A_{ul} (s^{-1})	g_u	n_{crit} (cm^{-3})	Beam size ($''$)	Beam eff. η_{MB}
HCN	1-0	88.631602	4.25	2.407×10^{-5}	3	5.0×10^6 ^b	28	0.81
CN	1-0	113.494921	5.45	1.182×10^{-5}	3	1.1×10^5 ^c	22	0.78
CS	3-2	146.969029	14.1	6.071×10^{-5}	7	2.6×10^5 ^c	16	0.74
C ³⁴ S	3-2	144.617109	13.9	7.251×10^{-5} ^a	7	5.84×10^{11} ^a	16	0.74
H ¹³ CN	1-0	86.342274	4.14	1.512×10^{-5} ^a	3	9.7×10^6 ^b	29	0.81
H ¹³ CN	2-1	172.677881	12.43	6.90×10^{-5} ^a	5	1.2×10^6 ^c	14	0.68

References: Molecular data adopted from LAMDA/JPL databases: ^a calculated for $T = 300$ K; ^aJiménez-Donaire et al. 2016, assuming optically thin transition lines for an excitation temperature of 20K; ^cShirley 2015, assuming optically thin transition lines for an excitation temperature of 50K; ^dChandra et al. 1995.

Notes. Beam sizes and efficiencies are taken from <http://www.iram.es/IRAMES/mainWiki/Iram30mEfficiencies>

SMM10 and SMM12 are classified as Late Class 0 YSOs. The SMM5 and SMM6 protostars are the most evolved objects in our sample (Class I).

3. Results

3.1. Molecular emission maps

The line maps in the targeted molecules show variety of structures that can be associated with YSOs and a large-scale cloud emission. Different spatial extend in molecules radiation is connected with various physical conditions around protostars. Here, we present the large-scale maps of CS $J = 3 - 2$, HCN $J = 1 - 0$ and CN $J = 1 - 0$. Maps of their isotopologues are shown in the Appendix A.

We present large-scale intensity maps (Fig. 2) of the targeted lines integrated at the level of 3σ and above. Three of the observed molecules (HCN, CN and H¹³CN) are characterised by hyperfine structure. High resolution spectroscopy allow us to separate the emission from each hyperfine transitions. The integrated intensity maps of the lowest transitions of HCN, CN and H¹³CN are performed as a sum of all hyperfine splitting components. The maps are centred at $\alpha_{J2000} = 18^h29^m46.6^s$, $\delta_{J2000} = 01^\circ18'20.5''$.

Most of molecular emission is concentrated in the SE subcluster, where 6 low-mass protostars are located, while the continuum emission peaks in the NW subcluster what is correlated with the other 4 low-mass protostars positions. The most extended structures can be associated with molecular outflows ejected from low-mass protostars (Table 2). Outflows directions were marked based on previous studies in CO transitions CO $J = 3 - 2$ (Dionatos et al. 2010) and CO $J = 6 - 5$ /CO $J = 3 - 2$ (Yildiz et al. 2015). Five off-source positions were selected to detailed spectra analysis (Table 3).

The integrated line intensity map of HCN $J = 1 - 0$ shows extended emission along outflow directions. This is the strongest line among all observed. The emission is slightly correlated with continuum emission. Most of the emission is associated with protostars positions. The HCN $J = 1 - 0$ gas peaks around Ser-SMM9 and Ser-SMM4 protostars. There is no significant peak around strong submillimetre sources as Ser-SMM1 and Ser-SMM3. The low energy level of HCN ($E_u = 4.25$ K) with the critical density of 10^6 cm^{-2} traces cold, high-density gas. HCN has previously been shown to be a good tracer of molecular outflows activity (Lee et al. 2014). The HCN $J = 1 - 0$ line was

detected at all protostars positions, although it is weak at the positions of Ser-SMM5, Ser-SMM8 and Ser-SMM10. On the other hand, the HCN $J = 1 - 0$ emission is particularly strong in Ser-SMM4 blue-shifted outflow (outflow position no. 4). NS elongated structure is present around Ser-SMM1 protostar what was noticed also in high- J transitions (Yildiz et al. 2015). The emission is slightly extended along Ser-SMM9 outflows. There is no intensively elongated outflow structure from the other sources.

CS $J = 3 - 2$ line emission map shows similar spatial distribution to HCN $J = 1 - 0$. Both species trace the gas of the same properties, however the CS molecule is excited in slightly less dense environment ($\sim 10^5$ cm^{-3}). The most significant elongated structure can be associated with Ser-SMM4 blue-shifted outflow. It is situated at the same place in both maps, extending over $80''$. A similar large-scale structure is detected along Ser-SMM1 outflows, although it is stronger in the HCN $J = 1 - 0$ map. Emission around Ser-SMM9 have a circular shape, however, there is additional elongated structure in CS $J = 3 - 2$ line emission towards West. It overlaps with the S68N outflows seen in methanol observations (Kristensen et al. 2010). The HCN $J = 1 - 0$ line emission propagates for larger distances than the CS $J = 3 - 2$. It is also relatively stronger. In both cases the highest peak of the emission is situated around Ser-SMM4 protostar with a significant extent towards outflow position no. 3. Their weaker isotopic species H¹³CN $J = 1 - 0$ and C³⁴S(3-2) lines peak around the protostars position. The lines exhibits similar morphological distribution as HCN $J = 1 - 0$ and CS $J = 3 - 2$.

CN $J = 1 - 0$ line emission is focused mostly around the positions of protostars. The CN line is similarly low-energetic, however it peaks in different areas than HCN and CS. The CN $J = 1 - 0$ integrated intensity structures follow the $250\mu m$ continuum map. The highest local peaks are associated with Class 0 low-mass protostars: Ser-SMM3, Ser-SMM4 and Ser-SMM6, as well as local maxima around Ser-SMM1 and Ser-SMM9. The spatial distribution is qualitatively different compared to the HCN $J = 1 - 0$ and CS $J = 3 - 2$ maps. The strongest emission characterises the dense surroundings of protostars in SE subcluster while the NW subcluster does not show such a distinct emission as in the HCN $J = 1 - 0$ line. Dense emissive region of the Ser-SMM9 source is significantly weaker in CN $J = 1 - 0$ line. CN $J = 1 - 0$ map can be characterised by compact, condensed emission without any strongly elongated structures.

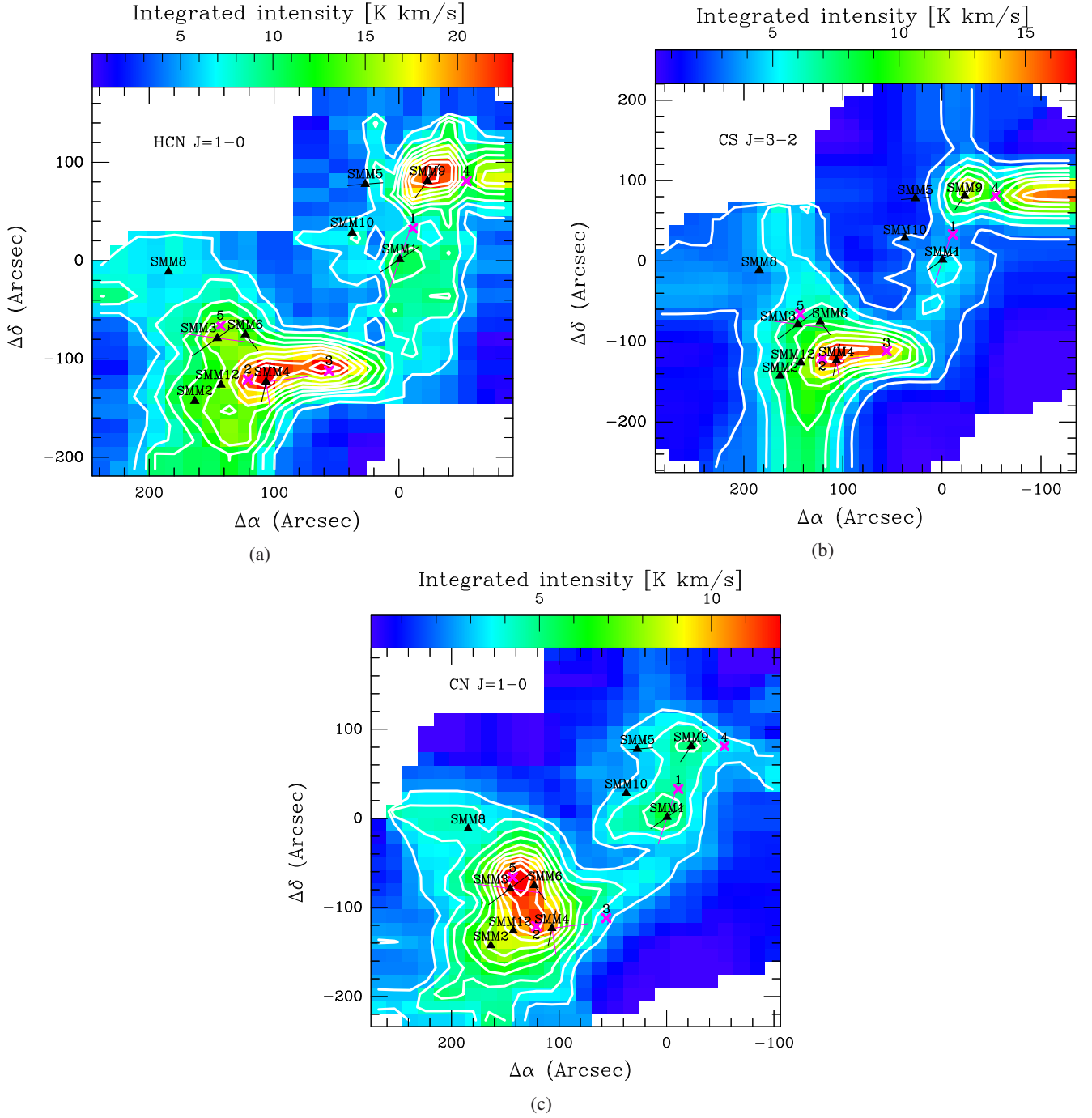


Fig. 2: Integrated intensity $\int T_{\text{mb}} dV$ of the HCN $J = 1 - 0$ (upper left panel), CS $J = 3 - 2$ (upper right panel) and CN $J = 1 - 0$ (bottom panel) in the Serpens Main region. The first contour at 30σ level, with step of 10σ , except CS $J = 3 - 2$ line (b) where the first contour is at 10σ with step of 5σ . Black triangles show the positions of the protostars (Suresh et al. 2016), whereas the black lines (Yildiz et al. 2015) and magenta line (Dionatos et al. 2010) show the associated outflow directions. Outflow positions are displayed as magenta crosses.

3.2. Comparison of the spatial extent of CN and HCN

We present a large-scale map of CN $J = 1 - 0$ and HCN $J = 1 - 0$ integrated intensity (Fig. 3) showing the emission exceeding 30σ for both molecules. The image of CN emission has been resampled to beam size of HCN in order to compare the same emitting regions.

The CN $J = 1 - 0$ transition is shifted to the north in respect to the HCN $J = 1 - 0$ emission. It is highly concentrated in the SE

subcluster, while the NW subcluster is dominated by the HCN. Both molecules show a diffusive ‘bridge’ between the two subclusters. In HCN it is connected with Ser-SMM4 and Ser-SMM1 outflows, while the CN follows the dust continuum emission. The emission in both molecules is anti-correlated north from Ser-SMM6 and west from Ser-SMM4 sources. At the dense area of Ser-SMM9 surrounding CN/HCN ratio is significantly weaker.

Both species has been detected in all sources positions. CN as a product of HCN photodissociation indicates other proper-

Table 2: Catalogue of protostars properties

Source (J2000.0)	R.A. (J2000.0)	Decl.	T_{bol} (K)	L_{bol} (L_{\odot})	Class	Other names
SMM9	18 29 48.3	+01 16 42.7	34.9	10.3	Early Class 0	Bolo22, Ser-emb8, ISO241, WMW23
SMM1	18 29 50.0	+01 15 20.3	35.4	78.7	Early Class 0	Bolo23, Ser-emb6, FIRS1, EC41
SMM5	18 29 51.4	+01 16 38.3	150.5	3.7	Early Class I	EC53, WMW24, Ser p3
SMM10	18 29 52.3	+01 15 48.8	82.6	6.2	Late Class 0	WMW21
SMM4	18 29 57.0	+01 13 11.3	76.9	4.4	Late Class 0	Bolo25, Ser-emb22
SMM6	18 29 57.8	+01 14 05.3	532.3	43.1	Late Class I	EC90, WMW35, SVS20S
SMM12	18 29 59.1	+01 13 14.3	96.9	5.7	Late Class 0	
SMM3	18 29 59.6	+01 13 59.2	35.0	6.9	Early Class 0	Bolo26
SMM2	18 30 00.5	+01 12 57.8	30.5	4.07	Early Class 0	Bolo28, Ser-emb4
SMM8	18 30 01.9	+01 15 09.2	15.3	0.2	Early Class 0	Bolo30

Coordinates taken from Suresh et al. 2016, except SMM8 (Lee et al. 2014).

Table 3: Properties of the selected off-source positions

Pos.	R.A. (J200)	Decl. (J200)	Remarks
1	18:29:45.47	+01:15:53.5	SMM1 blue-shifted outflow in CO $J = 3 - 2$
2	18:29:54.66	+01:13:19.5	max. CN $J = 1 - 0$, SMM4 blue-shifted outflow in CO $J = 3 - 2$
3	18:29:50.33	+01:13:68.5	max. HCN $J = 1 - 0$, SMM4 blue-shifted
4	18:29:43	+01:16:41.5	outflow visible in C ³⁴ S(3-2)
5	18:29:56.13	+01:14:14.5	max. CN $J = 1 - 0$, SMM9 surroundings

Table 4: CN/HCN integrated intensities

Source	$\int T_{\text{mb}} dV _{\text{CN}}$ (K km/s)	$\int T_{\text{mb}} dV _{\text{HCN}}$ (K km/s)	CN/HCN
SMM1	5.26	8.23	0.64
SMM2	8.41	12.57	0.67
SMM3	13.18	14.14	0.93
SMM4	9.89	17.59	0.56
SMM5	3.49	5.79	0.60
SMM6	10.57	11.85	0.89
SMM8	5.09	6.32	0.81
SMM9	5.41	14.08	0.38
SMM10	2.59	6.96	0.37
SMM12	9.10	13.27	0.69
Outflow1	4.8	8.3	0.58
Outflow2	11.2	21.4	0.52
Outflow3	3.8	23.6	0.16
Outflow4	3.7	12.4	0.3
Outflow5	13.3	14.9	0.89

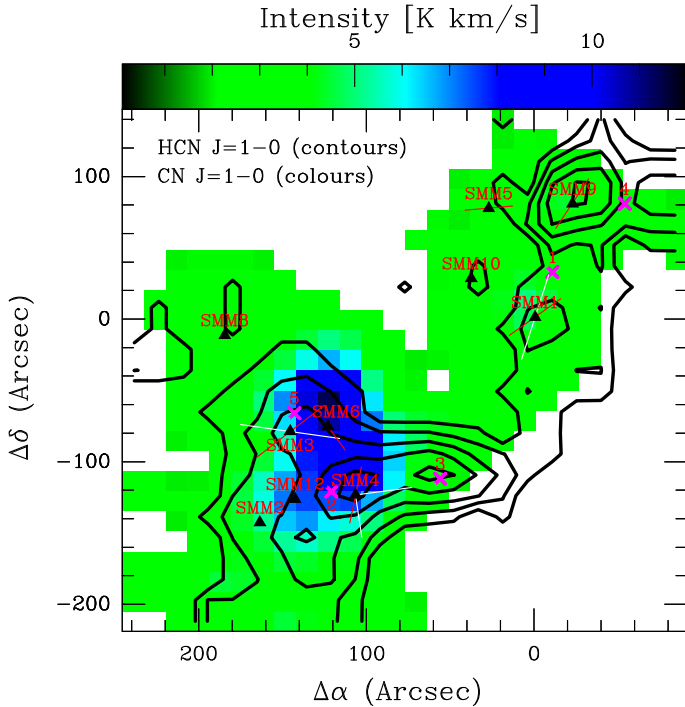


Fig. 3: Integrated intensity $\int T_{\text{mb}} dV$ of CN $J = 1 - 0$ (colours) and HCN $J = 1 - 0$ (contours) in the Serpens Main region. The first contour is 30. Triangles and lines marked as described in 2.

ties of low-mass protostars surroundings (Section 4). The highest CN/HCN integrated intensity ratio occurs in Ser-SMM6 and Ser-SMM3s protostars, as well as in the outflow position no. 5. On the other hand, Ser-SMM9 object is characterised by very low CN/HCN line ratio. Similarly low CN/HCN ratio is measured in outflow positions no. 3 and 4. Most of the sources show high flux values in both molecules (Table 4). However, they present unequal levels which indicates regions of different properties. In order to better understand this issue we analyse the molecular line profiles in the Section 3.2.

The low energy level of HCN ($E_u = 4.25$ K) traces cold, high-density gas. HCN has previously been shown to be a good tracer of molecular outflows activity (Walker-Smith et al. 2014). Similar spatial distribution is presented in the CS map. It means that both species trace the gas of the same properties. The CN line is similarly low-energetic, however it peaks in different areas than HCN and CS. CN as a product of HCN photodissociation indicates other properties of low-mass protostars surroundings (Section 4).

Spatial distribution of different lines emission varies depending on the observed molecule. The most distinct differences can be noticed between HCN and CN map. However, they present unequal levels which indicates regions of different properties. In

Table 5: Patterns of emission in molecular spieces

Source	HCN $J = 1 - 0$	CN $J = 1 - 0$	Remarks
SMM1	rb	c	multiple system of 5 components ^a
SMM2	b	b	multiple system of 3 components ^b
SMM3	b	b	
SMM4	rb	rb	binary ^c
SMM5	x	x	variable ^d
SMM5	x	x	
SMM8	r	r	
SMM9	rb	r	
SMM10	x	x	
SMM12	b	b	

References: ^a Hull et al. (2017), ^b Francis et al. (2019), ^c Aso et al. (2018), ^d Yoo et al. (2017).

order to better understand this issue we analyse the molecular line profiles in the Section 3.2.

3.3. Line profiles

We selected 14 representative on-source and off-source positions for a detailed analysis (Fig. 4). Nine of them are corresponding to the protostars positions, the other five off-source positions were selected based on local maximum of the flux.

In the majority of our sources five of targeted lines were detected: CN $J = 1 - 0$, HCN $J = 1 - 0$, CS $J = 3 - 2$, C³⁴S $J = 3 - 2$ and H¹³CN $J = 1 - 0$. The line is considered to be detected if there is an emission at the level of at least 3σ . A weak emission from H¹³CN $J = 2 - 1$ was found at the positions of four sources and it is not included in Fig. 4.

The strongest emission occurs in HCN $J = 1 - 0$, CN $J = 1 - 0$ and CS $J = 3 - 2$ lines and it was detected at the position of all of the sources. The emission in the other lines was multiplied in order to compare profiles between different molecules. In HCN, CN species and their isotopologues a few different velocity components can be identified what indicates the hyperfine splitting. This occurs if a molecule has a non-zero nuclear spin so there is also an interaction between the nuclear spin and the electronic angular momentum. The most distinct splitting can be spotted in the CN $J = 1 - 0$ profiles with five separate components situated between -70 km/s and 18 km/s. The HCN $J = 1 - 0$ line is characterised by three components with low separation situated in the range of -2 km/s – 16 km/s.

Ser-SMM1, Ser-SMM9, and Ser-SMM10 sources have wide spectral lines, while others exhibit narrow line profiles. Spectra extracted from Outflows no. 1, 4 and 5 shows prominent blue-shifted wings. Similar structure can be noticed in the Ser-SMM3 (panel no. 7) CS 3-2 and HCN $J = 1 - 0$ profiles.

4. Analysis

4.1. Lines column densities

Table 6 shows fluxes integrated from the average line profile at the positions of known protostars. Flux calculation in individual lines allows us to determine the column density of a given transition. The column density of the upper level N_u of each observed line was calculated based on following relation:

$$N_u = \beta \frac{\nu W}{A} \quad (1)$$

where $\beta = 1937 \text{ cm}^{-2}$ and $W = \int T_{mb} dV$ is the integrated intensity of the emission line. The frequency ν should be given in GHz. The total column density was obtained using:

$$N_{tot} = Q(T_{ex}) \exp\left(\frac{E_u}{k T_{ex}}\right) \frac{N_u}{g_u} \quad (2)$$

where Q is the partition function depending on excitation temperature T_{ex} and k Boltzmann constant.

The column densities of the upper level of CN $J = 1 - 0$ and HCN $J = 1 - 0$ transitions are presented in Table 6. The lowest transition of CN is more abundant molecule than the lowest transition line of HCN at the low-mass protostars positions. The column density of CN $J = 1 - 0$ varies between $10^{14} - 10^{15} \text{ cm}^{-2}$, while in the column density of the HCN's lowest transition reaches 10^{14} cm^{-2} . In the case of Ser-SMM2, Ser-SMM3, Ser-SMM4, Ser-SMM6 and Ser-SMM12 HCN $J = 1 - 0$ line column density is an order of magnitude lower than the column density of the equivalent CN transition. This result provides a clue to better understand of the low-mass protostars chemistry.

4.2. RADEX modelling

Column densities can be independently determined using molecular excitation models. Line ratio can provide additional information concerning physical properties of the observed gas.

The non-LTE radiative transfer code RADEX (van der Tak et al. 2007) was run in order to prepare sets of molecular excitation models. The CN and HCN molecules column density ratio, hydrogen number density and kinetic temperature of the gas were free parameters. HCN column density was chosen as 10^8 cm^{-2} in order to ensure optically thin emission. CN column density parameter varies from 10^6 cm^{-2} to 10^{10} cm^{-2} what translates into N_{CN}/N_{HCN} in following limits: $10^{-2} - 10^2$. The sets of models were obtained assuming a line width of 1.0 km s^{-1} , hydrogen densities of $n_{H_2} = 10^3 \text{ cm}^{-3}$, $n_{H_2} = 10^4 \text{ cm}^{-3}$ and $n_{H_2} = 10^5 \text{ cm}^{-3}$ and kinetic temperatures of $T_{kin} = 30 \text{ K}$, $T_{kin} = 75 \text{ K}$ and $T_{kin} = 200 \text{ K}$. The molecular data files used during modelling were procured from the Leiden Atomic and Molecular Database (LAMDA, Schöier et al. 2005).

Fig. 5 presents one exemplary set of models of CN/HCN column density ratio versus the modelled line intensities ratio for hydrogen densities of $n_{H_2} = 10^5 \text{ cm}^{-3}$ and kinetic temperatures of $T_{kin} = 200 \text{ K}$. The rest of the models are shown in the Appendix B. CN/HCN column density ratio weakly depends on hydrogen density and kinetic temperature in the low limit of those parameters. All presented models show similar properties.

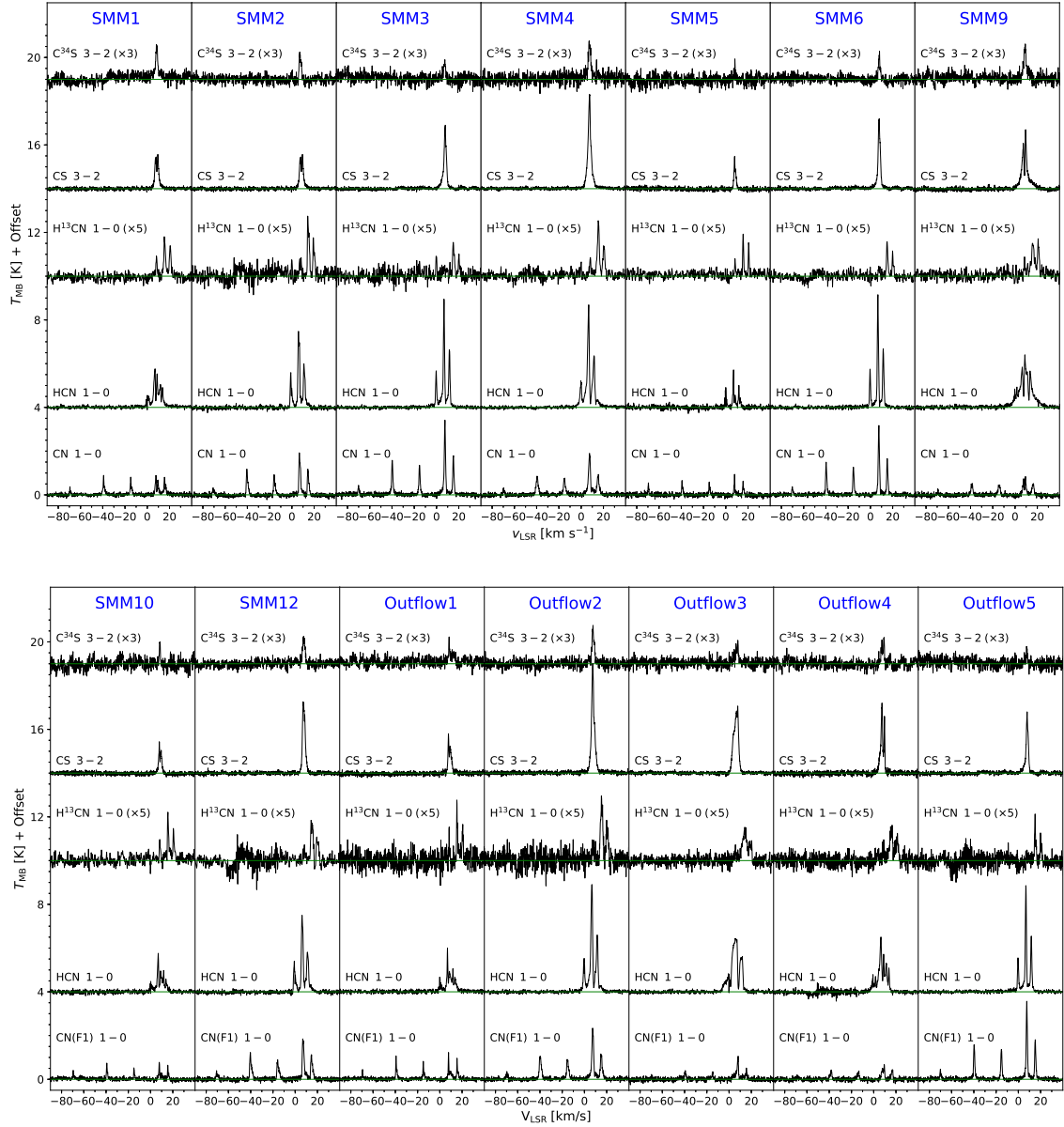


Fig. 4: Serpens Main sources spectra of $C^{34}S$ $J = 3 - 2$, CS $J = 3 - 2$, $H^{13}CN$ $J = 1 - 0$, HCN $J = 1 - 0$ and CN $J = 1 - 0$ lines.

The modelled line intensities are compared with the observations. The observed line intensity ratio covers a range of 0.0-1.0 of the column density ratio in the logarithmic scale. That corresponds to a few times higher CN column density than the same parameter of HCN. This result slightly depends on hydrogen densities and kinetic temperature of the gas (see Table 7).

5. Discussion

5.1. Astrochemical model

In Section 3.2 we have raised a question about influence of the UV radiation on the surrounding of the observed sources. A simply astrochemical model can be obtained in order to estimate the intensity of the UV field in reference to units of the interstellar UV field G_0 .

HCN easily photodissociates into CN molecule and H atom. On the other hand, CN needs more energetic photon for disin-

tegration. This leads to higher abundance of CN molecules and increases CN/HCN column density ratio.

We used Nahoon code that model the time evolution of 474 species involving gas-phase and gas-grain reactions with fixed temperature, density, UV, and cosmic ray fluxes (Wakelam et al. 2012). The evolution of CN, HCN and CS abundances was started at the time of a dense cloud formation. The UV radiation is described through the relation between visual extinction A_V and the photodissociation rate coefficient k Eq A.3:

$$k = \alpha e^{-\gamma A_V} \quad (3)$$

Here, α and γ are coefficients of photodissociation HCN equal 1.64×10^{-9} and 3.12 respectively (Heays et al. (2017)).

Figure 6 shows a model corresponding to a typical dense cloud with temperature of 10 K and hydrogen total density of $n_{H+2H_2} = 10^4 \text{ cm}^{-3}$. The chemical composition of the studied molecules stabilises at the time of 10^7 yrs with HCN abundance higher with respect to CN. We assumed the time of 10^6 yrs as the

Table 6: Integrated fluxes of the observed line at the positions of protostars

Source	Line	$\int T_{mb} dV$ (K km/s)	T_{peak} (K)	N_{up} (cm ⁻²)	N_{tot} (cm ⁻²)
SMM1	CN 1-0	5.26	0.89	1.1×10^{13}	6.6×10^{14}
	HCN 1-0	8.23	1.76	5.2×10^{12}	2.0×10^{14}
	CS 3-2	2.98	1.57	2.1×10^{12}	2.3×10^{13}
	C ³⁴ S 3-2	1.21	0.56		
	H ¹³ CN 1-0	1.43	0.45		
SMM2	CN 1-0	8.41	1.92	1.8×10^{13}	1.1×10^{15}
	HCN 1-0	12.57	3.48	8.0×10^{12}	3.0×10^{12}
	CS 3-2	6.57	3.10	4.5×10^{12}	2.0×10^{13}
	C ³⁴ S 3-2	0.89	0.47		
	H ¹³ CN 1-0	1.65	0.64		
SMM3	CN 1-0	13.18	3.42	2.8×10^{13}	1.7×10^{15}
	HCN 1-0	14.14	4.95	8.9×10^{12}	3.4×10^{14}
	CS 3-2	8.11	2.90	5.6×10^{13}	6.2×10^{13}
	C ³⁴ S 3-2	0.604	0.36		
	H ¹³ CN 1-0	0.88	0.41		
SMM4	CN 1-0	9.89	1.90	2.1×10^{13}	1.3×10^{15}
	HCN 1-0	17.59	4.69	1.1×10^{13}	4.2×10^{14}
	CS 3-2	14.4	4.30	9.9×10^{12}	1.1×10^{14}
	C ³⁴ S 3-2	1.56	0.64		
	H ¹³ CN 1-0	1.83	0.61		
SMM5	CN 1-0	3.49	0.94	7.3×10^{12}	4.4×10^{14}
	HCN 1-0	5.79	1.72	3.7×10^{12}	1.4×10^{14}
	CS 3-2	2.61	1.48	1.8×10^{12}	2.0×10^{13}
	C ³⁴ S 3-2	0.28	0.42		
	H ¹³ CN 1-0	1.16	0.42		
SMM6	CN 1-0	10.57	3.17	2.2×10^{13}	1.3×10^{15}
	HCN 1-0	11.85	5.15	8.9×10^{12}	3.3×10^{14}
	CS 3-2	7.86	3.2	5.4×10^{12}	6.0×10^{13}
	C ³⁴ S 3-2	0.74	0.56		
	H ¹³ CN 1-0	0.81	0.40		
SMM8	CN 1-0	5.09	0.94	1.1×10^{13}	6.4×10^{14}
	HCN 1-0	6.32	1.72	4.0×10^{12}	1.5×10^{14}
	CS 3-2	4.64	2.08	3.2×10^{12}	3.5×10^{13}
	C ³⁴ S 3-2	0.22	0.35		
	H ¹³ CN 1-0	0.46	0.21		
SMM9	CN 1-0	5.41	0.84	1.1×10^{13}	6.8×10^{14}
	HCN 1-0	14.08	2.40	8.9×10^{12}	3.4×10^{14}
	CS 3-2	9.85	2.70	6.8×10^{12}	7.5×10^{13}
	C ³⁴ S 3-2	1.3	0.62		
	H ¹³ CN 1-0	1.57	0.38		
SMM10	CN 1-0	2.59	0.78	5.4×10^{12}	3.3×10^{14}
	HCN 1-0	6.96	1.75	4.4×10^{12}	1.7×10^{14}
	CS 3-2	3.83	1.43	2.6×10^{12}	2.9×10^{13}
	C ³⁴ S 3-2	0.56	0.40		
	H ¹³ CN 1-0	0.98	0.45		
SMM12	CN 1-0	9.10	1.85	1.9×10^{13}	1.2×10^{15}
	HCN 1-0	13.27	3.50	8.4×10^{12}	3.2×10^{14}
	CS 3-2	9.75	3.26	6.7×10^{12}	7.4×10^{13}
	C ³⁴ S 3-2	1.18	0.58		
	H ¹³ CN 1-0	1.43	0.46		

time when the star formation starts in dense clouds. The modelled abundances of the all 474 species at the time of 10^6 yrs were used as an input data for the following set of models.

Based on the initial abundances of all species from starless cloud modelling, the closest neighbourhood of low-mass protostars was simulated. We adopted the UV radiation and cosmic ray fluxes typical for dense clouds as an initial conditions. The

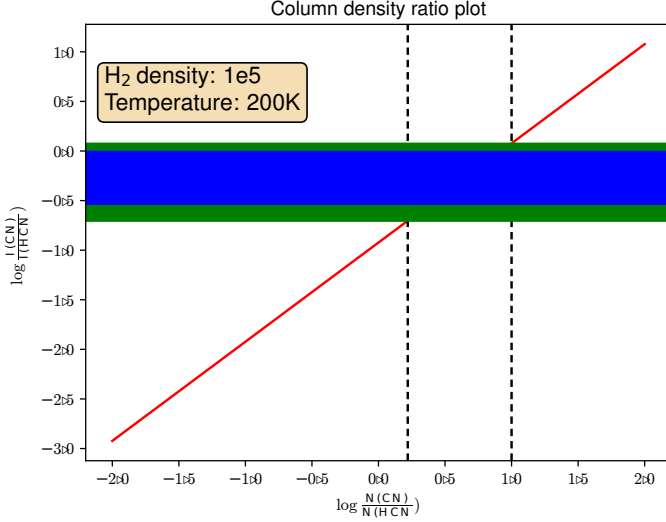


Fig. 5: CN/HCN column density ratio for hydrogen densities of $n_{\text{H}_2} = 10^5 \text{ cm}^{-3}$ and kinetic temperatures of $T_{\text{kin}} = 200 \text{ K}$ (red line). The observed line intensity ratio is plotted in blue (protostars positions) and green (all positions).

Table 7: CN/HCN column density ratio

n_{H_2} (cm^{-3})	T_{kin} (K)	$\log_{10}(N[\text{CN}]/N[\text{HCN}])$
10^3	30	0.03-0.88
10^3	75	0.06-0.84
10^3	200	0.00-0.78
10^4	30	0.16-0.94
10^4	75	0.08-0.86
10^4	200	0.04-0.82
10^5	30	0.20-0.98
10^5	75	0.18-0.86
10^5	200	0.22-1.00

visual extinction and the total cosmic-ray ionization rate were set as 5 mag and $1.3 \times 10^{17} \text{ s}^{-1}$ respectively. The sets of models were run for the temperature range between 10 and 200 K and the total hydrogen densities from 10^4 cm^{-3} to 10^6 cm^{-3} . The resulting ratio of CN and HCN abundances is shown in Fig. 7. The results are consistent with starless cloud model (Fig. 6). Without any additional source of the UV radiation the HCN is more abundant molecule than CN about 2-3 orders of magnitude. Moreover, the CN/HCN abundances ratio is slightly dependent on gas temperature up to 150 K.

This results justify assumptions for the modelling the neighbourhood of low-mass protostars with the fixed gas temperature of 50 K. The models were run for the range of visual extinction between 1^{m} and 2.1^{m} what corresponds to the UV radiation field of 0.044 to 0.001 G_0 . The computed CN/HCN abundances ratio is presented in Fig. 8. The additional UV radiation of the strength of few thousandth of the average interstellar UV radiation field is enough to cover the observations in wide range of total hydrogen densities.

An additional source of the UV radiation was needed to reproduce the observational CN/HCN abundances what leads to conclusion that there is non-zero UV radiation field around low-mass protostars.

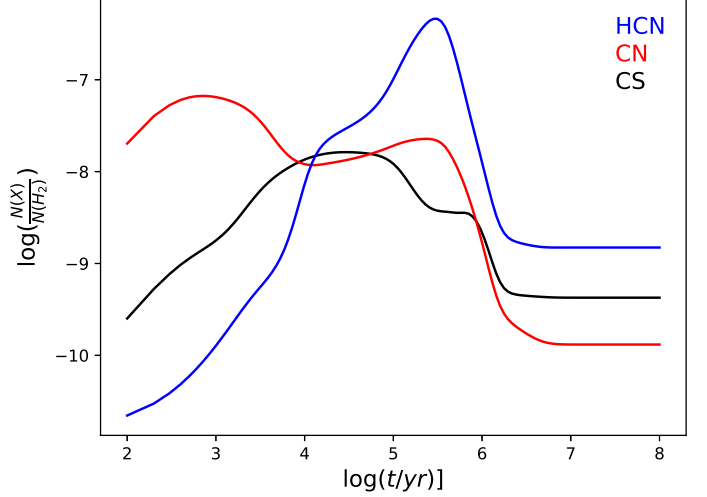


Fig. 6: Time evolution of CN (red line), HCN (blue line) and CS (black line) abundances obtained with Nohoon astrochemical code with initial parameters of $n_{\text{H}+\text{H}_2} = 10^4 \text{ cm}^{-3}$, $T = 10 \text{ K}$, $A_V = 5 \text{ mag}$. The assumed cosmic-ray ionization rate is $1.3 \times 10^{17} \text{ s}^{-1}$, dust to gas mass ratio is 0.01, dust grain radius is 10^{-5} cm , grain density is 3 g cm^{-3} .

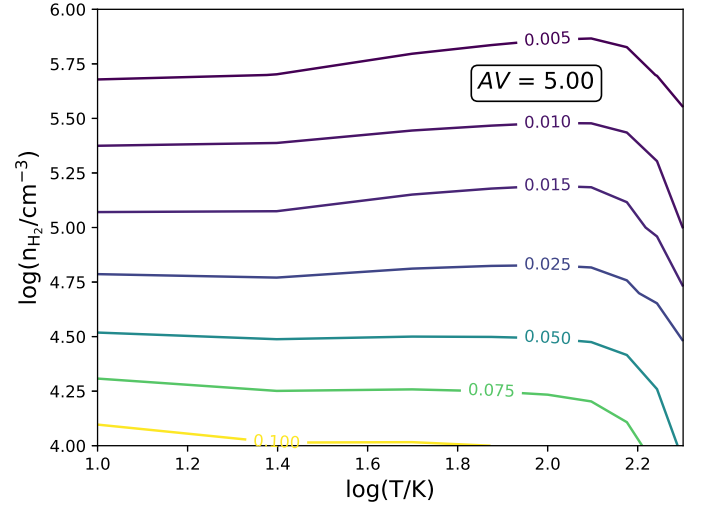


Fig. 7: Contour plot of Nohoon sets of models of CN/HCN abundances ratio with fixed visual extinction $A_V = 5 \text{ mag}$ at the time of 10^7 yrs after star formation began in the cloud.

6. Conclusions

Acknowledgements. AM, AK and MG are supported by the Polish National Science Center grants NAVA InterAPS and 2016/21/D/ST9/01098. This research has made use of data from the Herschel Gould Belt survey (HGBS) project (<http://gouldbelt-herschel.cea.fr>). The HGBS is a Herschel Key Programme jointly carried out by SPIRE Specialist Astronomy Group 3 (SAG 3), scientists of several institutes in the PACS Consortium (CEA Saclay, INAF-IFSI Rome and INAF-Arcetri, KU Leuven, MPIA Heidelberg), and scientists of the Herschel Science Center (HSC).

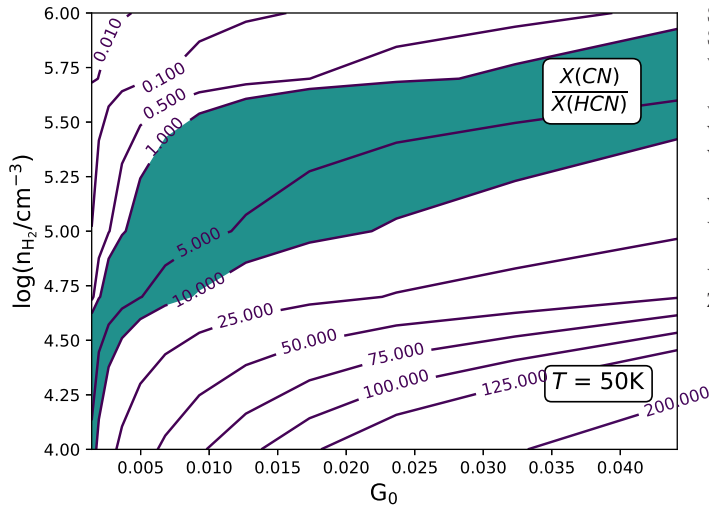


Fig. 8: Similar to Fig. 7 but for fixed temperature $T = 50$ K.

References

- André, P., Men'shchikov, A., Bontemps, S., et al. 2010, A&A, 518, L102
 André, P., Ward-Thompson, D., & Barsony, M. 1993, ApJ, 406, 122
 Arce, H. G. & Sargent, A. I. 2006, ApJ, 646, 1070
 Aso, Y., Hirano, N., Aikawa, Y., et al. 2018, ApJ, 863, 19
 Bachiller, R., Pérez Gutiérrez, M., Kumar, M. S. N., & Tafalla, M. 2001, A&A, 372, 899
 Bergin, E. A. & Tafalla, M. 2007, ARA&A, 45, 339
 Bontemps, S., André, P., Terebey, S., & Cabrit, S. 1996, A&A, 311, 858
 Chapillon, E., Guilloteau, S., Dutrey, A., Piétu, V., & Guélin, M. 2012, A&A, 537, A60
 Di Francesco, J., Johnstone, D., Kirk, H., MacKenzie, T., & Ledwosinska, E. 2008, ApJS, 175, 277
 Dionatos, O., Nisini, B., Codella, C., & Giannini, T. 2010, A&A, 523, A29
 Dunham, M. M., Allen, L. E., Evans, Neal J., I., et al. 2015, ApJS, 220, 11
 Dunham, M. M., Arce, H. G., Mardones, D., et al. 2014, ApJ, 783, 29
 Enoch, M. L., Evans, Neal J., I., Sargent, A. I., & Glenn, J. 2009, ApJ, 692, 973
 Enoch, M. L., Glenn, J., Evans, Neal J., I., et al. 2007, ApJ, 666, 982
 Evans, Neal J., I., Dunham, M. M., Jørgensen, J. K., et al. 2009, ApJS, 181, 321
 Francis, L., Johnstone, D., Dunham, M. M., Hunter, T. R., & Mairs, S. 2019, ApJ, 871, 149
 Fuente, A., Martín-Pintado, J., Cernicharo, J., & Bachiller, R. 1993, A&A, 276, 473
 Fuente, A., Martín-Pintado, J., & Gaume, R. 1995, ApJ, 442, L33
 Green, J. D., Evans, Neal J., I., Jørgensen, J. K., et al. 2013, ApJ, 770, 123
 Heays, A. N., Bosman, A. D., & van Dishoeck, E. F. 2017, A&A, 602, A105
 Hogerheijde, M. R., van Dishoeck, E. F., Salverda, J. M., & Blake, G. A. 1999, ApJ, 513, 350
 Hull, C. L. H., Girart, J. M., Tychoniec, Ł., et al. 2017, ApJ, 847, 92
 Jørgensen, J. K. 2004, A&A, 424, 589
 Karska, A., Herczeg, G. J., van Dishoeck, E. F., et al. 2013, A&A, 552, A141
 Karska, A., Kaufman, M. J., Kristensen, L. E., et al. 2018, ApJS, 235, 30
 Karska, A., Kristensen, L. E., van Dishoeck, E. F., et al. 2014, A&A, 572, A9
 Kristensen, L. E., van Dishoeck, E. F., Mottram, J. C., et al. 2017, A&A, 605, A93
 Kristensen, L. E., van Dishoeck, E. F., van Kempen, T. A., et al. 2010, A&A, 516, A57
 Kroupa, P. 2002, Science, 295, 82
 Lada, C. J. 1987, in IAU Symposium, Vol. 115, Star Forming Regions, ed. M. Peimbert & J. Jugaku, 1
 Larson, R. B. 2003, Reports on Progress in Physics, 66, 1651
 Lee, K. I., Fernández-López, M., Storm, S., et al. 2014, ApJ, 797, 76
 Manoj, P., Watson, D. M., Neufeld, D. A., et al. 2013, ApJ, 763, 83
 Myers, P. C. & Ladd, E. F. 1993, ApJ, 413, L47
 Ortiz-León, G. N., Dzib, S. A., Kounkel, M. A., et al. 2017, ApJ, 834, 143
 Riaz, B., Thi, W. F., & Caselli, P. 2018, MNRAS, 481, 4662
 Schöier, F. L., van der Tak, F. F. S., van Dishoeck, E. F., & Black, J. H. 2005, A&A, 432, 369
 Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, AJ, 131, 1163

- Spaans, M., Hogerheijde, M. R., Mundy, L. G., & van Dishoeck, E. F. 1995, ApJ, 455, L167
 Stäuber, P., Benz, A. O., Jørgensen, J. K., et al. 2007, A&A, 466, 977
 Suresh, A., Dunham, M. M., Arce, H. G., et al. 2016, AJ, 152, 36
 van der Tak, F. F. S., Black, J. H., Schöier, F. L., Jansen, D. J., & van Dishoeck, E. F. 2007, A&A, 468, 627
 van Kempen, T. A., van Dishoeck, E. F., Güsten, R., et al. 2009, A&A, 507, 1425
 Wakelam, V., Herbst, E., Loison, J. C., et al. 2012, ApJS, 199, 21
 Walker-Smith, S. L., Richer, J. S., Buckle, J. V., Hatchell, J., & Drabek-Maunder, E. 2014, MNRAS, 440, 3568
 Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, AJ, 140, 1868
 Yıldız, U. A., Kristensen, L. E., van Dishoeck, E. F., et al. 2015, A&A, 576, A109
 Yoo, H., Lee, J.-E., Mairs, S., et al. 2017, ApJ, 849, 69
 Zuckerman, B., Kuiper, T. B. H., & Rodríguez Kuiper, E. N. 1976, ApJ, 209, L137

Appendix A: Spectral Energy Distributions

Broad-band observations are needed in order to determine physical properties of a protostar. Dunham et al. 2015 studied properties of protostars in the Serpens molecular cloud using 2MASS (Skrutskie et al. 2006) and Spitzer IRAC/MIPS (Evans et al. 2009), observations covering the range of 1.25–70 μm , photometry from Wide-field Infrared Survey Explorer 12 and 22 μm (WISE; Wright et al. 2010), SHARC-II 350 μm (Suresh et al. 2016), the SCUBA Legacy Catalog 450 and 850 μm (Di Francesco et al. 2008) and 1.1 mm observations from Bolocam dust survey (Enoch et al. 2007). The Serpens Main region was also observed during the Herschel Gould Belt survey project (André et al. 2010). SPIRE/PACS photometry in the Serpens molecular cloud is discussed in Fiorellino et al. (in prep.).

Based on SEDs the bolometric temperature and luminosity can be calculated for each of the observed protostars. The bolometric luminosity was determined by integrating the SEDs over frequency:

$$L_{bol} = \pi d^2 \int F_\nu d\nu \quad (\text{A.1})$$

where d is the cloud distance of 436 ± 9.2 pc (Ortiz-León et al. 2017). The bolometric temperature was calculating as described in Myers & Ladd 1993:

$$T_{bol} = 1.25 \cdot 10^{-11} \bar{\nu} \quad (\text{A.2})$$

where $\bar{\nu}$ is the mean frequency given by:

$$\bar{\nu} = \frac{\int \nu F_\nu d\nu}{\int F_\nu d\nu} \quad (\text{A.3})$$

Using Scipy *splrep* and *splev* functions cubic smooth spline interpolation of the photometric data was performed while calculating the protostars parameters. Integration along the resulting axis was obtain with the composite trapezoidal rule (*Scipy* package). The photometric data allows us to perform the integration along wide range of wavelength with exception of SMM8. Here we have only 4 photometric points from the Herschel Gould Belt so the calculated bolometric luminosity and temperature can be underestimated.

Appendix B: Molecular emission maps

Appendix C: Flux correlations

Appendix D: Wings emission

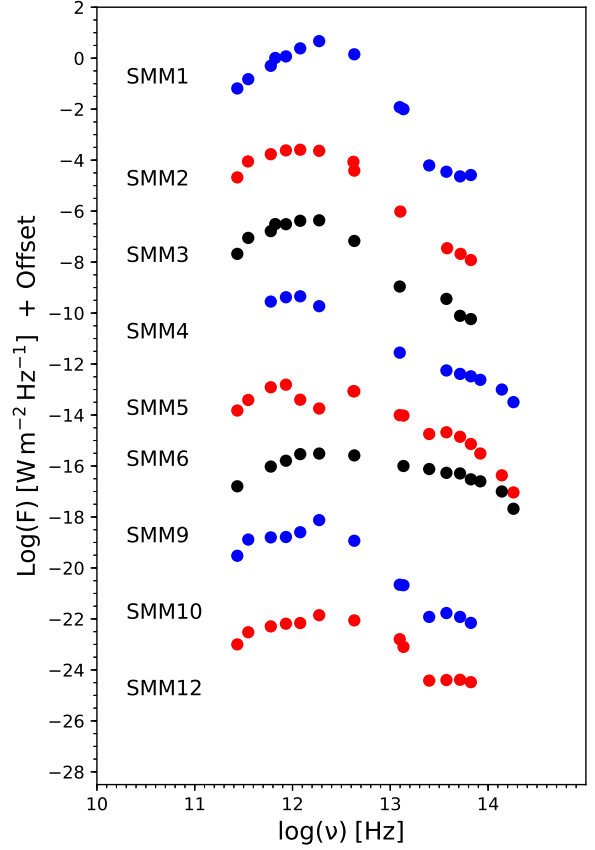


Fig. A.1: Spectral Energy Distributions of protostars in the Serpens Main region.

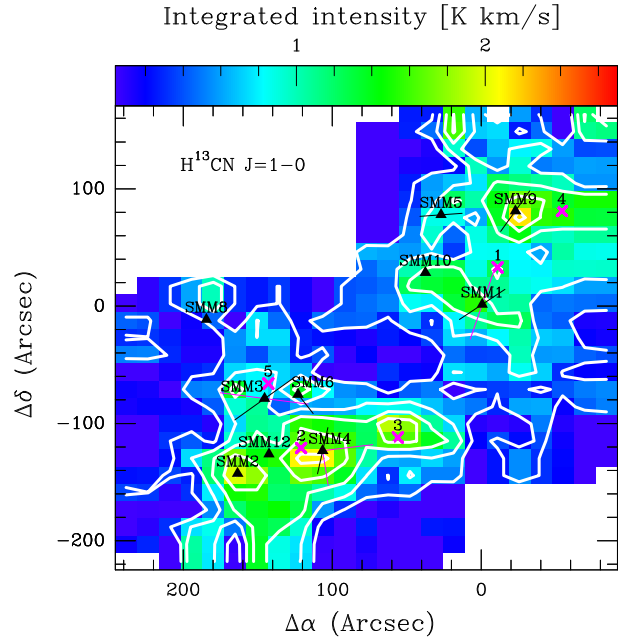


Fig. B.1: Similar to Fig. 2 but the emission of the $\text{H}^{13}\text{CN } J = 1 - 0$ line. The first contour at 10σ level, with step of 10σ

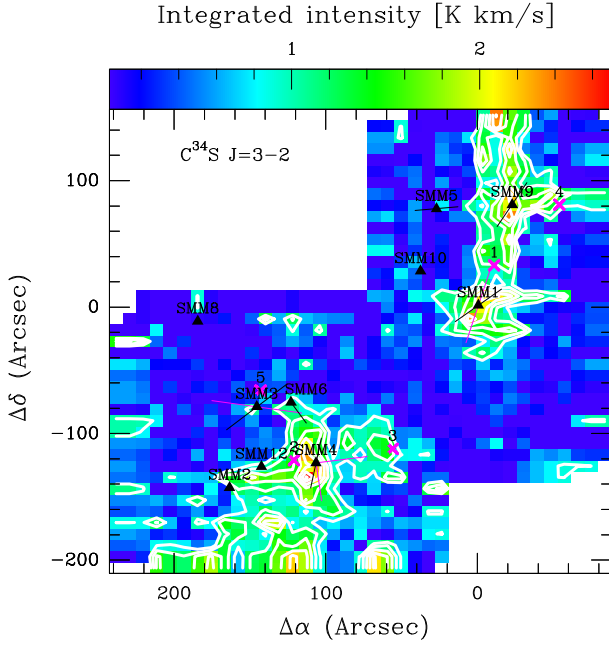


Fig. B.2: Similar to Fig. 2 but the emission of the $C^{34}S$ $J = 3 - 2$ line. The first contour at 30σ level, with step of 10σ

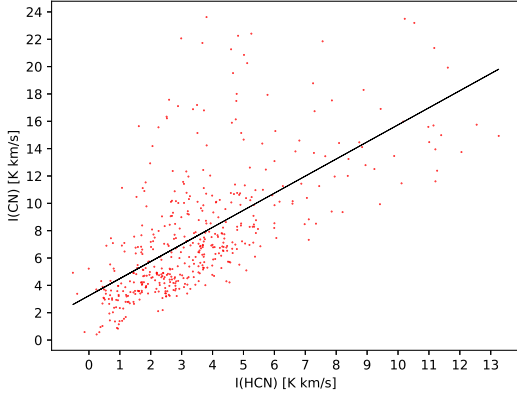


Fig. C.1: Correlation of CN $J = 1 - 0$ and HCN $J = 1 - 0$ integrated intensities $\int T_{mb} dV$ for every pixel. A least-squares linear regression (black line) with following parameters: slope = 1.25, intercept = 0.24, stderr = 0.07. The Pearson correlation coefficient equals 0.65.

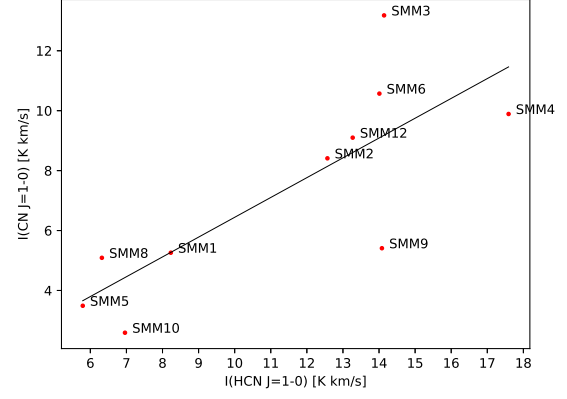


Fig. C.2: Correlation of CN $J = 1 - 0$ and HCN $J = 1 - 0$ integrated intensities $\int T_{mb} dV$ for all sources. A least-squares linear regression (black line) with following parameters: slope = 0.66, intercept = -0.17, stderr = 0.18. The Pearson correlation coefficient equals 0.79.

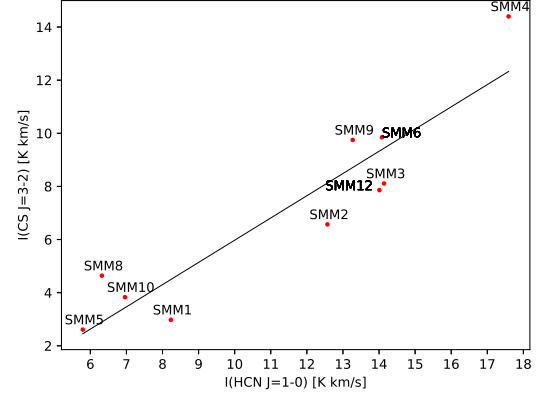


Fig. C.3: Similar to Fig. C.2 but CS $J = 3 - 2$ and HCN $J = 1 - 0$ lines correlation. A least-squares linear regression (black line) with following parameters: slope = 0.83, intercept = -2.39, stderr = 0.12. The Pearson correlation coefficient equals 0.93.

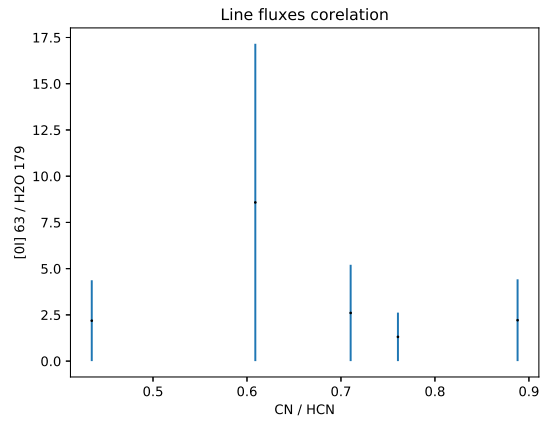


Fig. C.4: Correlation of [OI] $63\mu m$ and H_2O $179\mu m$ fluxes.

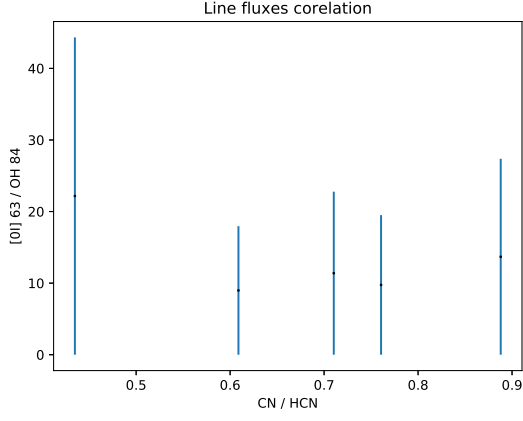


Fig. C.5: Correlation of $[OI] 63 \mu m$ and $OH 84 \mu m$ fluxes.

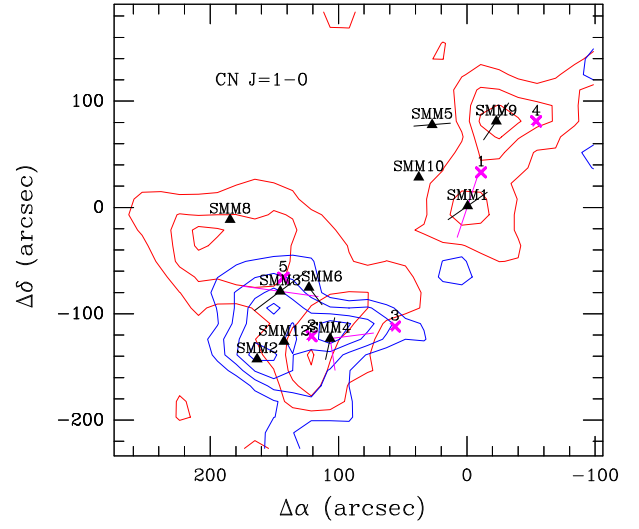


Fig. D.2: Similar to Fig. D.1 but the emission of the $CN J = 1-0$ line.

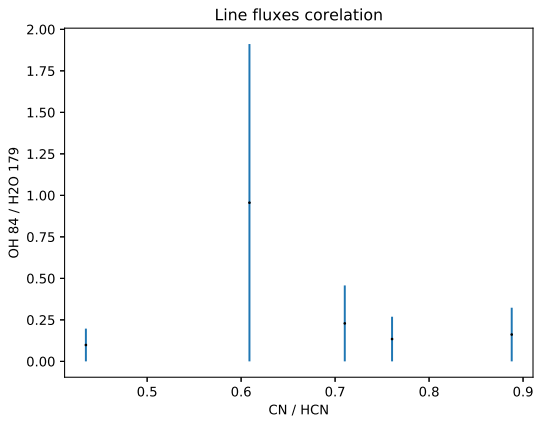


Fig. C.6: Correlation of $OH 84 \mu m$ and $H_2O 179 \mu m$ fluxes.

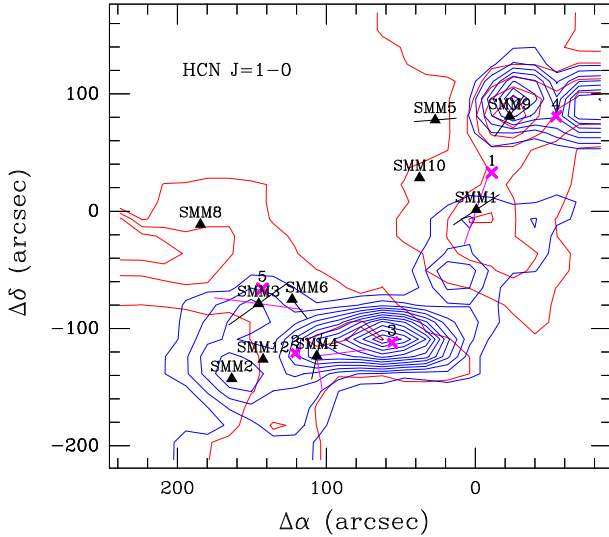


Fig. D.1: $HCN J = 1-0$ blue-shifted and red-shifted emission in Serpens Main region. The first contour is at 5σ with the step of 3σ . Triangles and lines marked as described in 2.