

Stellar oscillations induced by a planetary companion

Going off on a tangent

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Abstract

The gravitational potential from a planet orbiting a star causes a regular perturbation which results in oscillations. These tidally induced oscillations cause changes in displacement and flux. In solving the non-adiabatic oscillation equations we found that the tangential displacement at the surface is $\sim 10^3 - 10^4$ times larger than the non-adiabatic solution. This result agrees well with analytical expressions in the limit of vanishing pressure, which is approximately found at the surface. Modelling these oscillations could be used to characterise exoplanetary orbits, including determining planetary masses, and has potential for planetary detection in its own right.

Introduction

Both the Radial Velocity (RV) and transit methods for detecting exoplanets have been successful, but both are limited – characterising the planetary system’s mass is a particular difficulty due to degeneracy with the inclination or the density, respectively. Understanding other ways in which the planet and the star interact could break these degeneracies.

Given their mass, size and proximity to their host star, hot Jupiters have a stronger interaction with their host star than other planets. The first-order effect of their presence is what is used to detect them in both the RV and transit method – the motion of the star around the common centre of mass, and the light blocked by their presence respectively. The tidal potential due to the planet is a second-order effect which changes the amplitude and wavelength of the star’s light.

Whilst modelling this interaction for the non-adiabatic case has been undertaken before [6], a detailed analysis of the tangential displacement has been lacking, or has otherwise been done assuming that the change due to non-adiabaticity is small [7].

This work particularly focusses upon the behaviour at the very surface, where non-adiabatic effects are prominent. Comparison between the modelled behaviour and analytical results under the conditions present at the very surface show good agreement.

Method

Numerical method

To model the oscillations, the linear non-adiabatic stellar oscillation equations were solved in the case that the star is perturbed by a regular tidal potential, due to the planet. The variables directly solved for are: ξ_r , the radial displacement; F'_r , the perturbation to the radial radiative flux; p' , the perturbation to the pressure; and T' , the perturbation to the temperature.

The equations solved are:

$$\frac{1}{r^2} \frac{\partial}{\partial r} (r^2 \rho_0 \xi_r) + \left(\frac{\rho_0}{\chi_\rho p_0} - \frac{l(l+1)}{m^2 \omega^2 r^2} \right) p' - \frac{\rho_0 \chi_T}{T_0 \chi_\rho} T' = \frac{l(l+1)}{m^2 \omega^2 r^2} \rho_0 \Phi_P \quad (1)$$

$$\left(i \rho_0 m \omega c_p + \frac{l(l+1)}{r^2} K_0 \right) T' - (i m \omega c_p \nabla_{ad} \rho_0 T_0) \frac{p'}{p_0} + i m \omega \rho_0 T_0 \frac{\partial s_0}{\partial r} \xi_r + \frac{1}{r^2} \frac{\partial}{\partial r} (r^2 F'_r) = 0 \quad (2)$$

$$\frac{F'_r}{K_0} - \left(\frac{\partial}{\partial r} - \frac{1}{T_0} \frac{\partial T_0}{\partial r} \left[-3 + \frac{1}{\kappa_0} \left(\frac{\partial \kappa}{\partial \ln T} \right)_\rho - \frac{\chi_T}{\chi_\rho} \left(1 + \frac{1}{\kappa_0} \left(\frac{\partial \kappa}{\partial \ln \rho} \right)_T \right) \right] \right) T' - \frac{\partial T_0}{\partial r} \frac{1}{p_0 \chi_\rho} \left(1 + \frac{1}{\kappa_0} \left(\frac{\partial \kappa}{\partial \ln \rho} \right)_T \right) p' = 0 \quad (3)$$

$$-m^2 \omega^2 \rho_0 \xi_r + \left(\frac{\partial}{\partial r} + \frac{\rho_0}{\chi_\rho p_0} \frac{\partial \Phi_0}{\partial r} \right) p' - \frac{\partial \Phi_0}{\partial r} \frac{\rho_0 \chi_T}{T_0 \chi_\rho} T' = -\rho_0 \frac{\partial \Phi_P}{\partial r} \quad (4)$$

which correspond to the continuity equation, entropy equation, radiative diffusion equation, and the momentum equation respectively.

The boundary conditions are split, two apply at the centre, and two at the surface. At the centre $\xi_r = 0$ and $F'_r = 0$ which ensure continuity. At the surface $\Delta P = 0$, ensuring that the pressure at the perturbed surface is unchanged, and $\left(4 \frac{\Delta T}{T_0} - \frac{\Delta F'_r}{F'_r} \right) = 0$, ensuring that the star remains a blackbody emitter.

The Henyey method [1] was used to solve these equations, using a solar-type model produced using MESA [2] [3] [4] [5] as the equilibrium background star which we perturbed.

Analytical solution

For the region at the surface, where $\Delta P \approx 0$, analytical expressions for the relationships between certain variables can be calculated by making use of this.

By taking the linearised equation of motion, and introducing $\vec{\xi}_h = r \vec{\nabla}_\perp V$, we arrive at

$$-\omega^2 \rho \xi_r = -\frac{\partial \Delta P}{\partial r} - \frac{l(l+1)}{r} \rho g V - \xi_r r^2 \rho \frac{\partial (g/r^2)}{\partial r} - \rho \frac{\partial \Phi'}{\partial r} \quad (5)$$

$$-\omega^2 \rho V = -\frac{\Delta P}{r} - \frac{g \rho \xi_r}{r} - \frac{\rho \Phi'}{r}. \quad (6)$$

Rearranging these, and taking $\Delta P = 0$, we get to the expression

$$\frac{V}{\xi_r} = \frac{\frac{g}{r \omega^2} \left(-\frac{\partial \Delta P}{\partial r} + \rho \Phi' \left(-\frac{2}{r} + \frac{r^2 \partial (g/r^2)}{\partial r} \right) \right)}{\left(-\frac{\partial \Delta P}{\partial r} - 2 \rho \Phi' \left(\frac{1}{r} + \frac{3g}{r^2 \omega^2} \right) \right)}. \quad (7)$$

Results

The behaviour in the last 0.1% of the stellar radius differs greatly between the adiabatic and non-adiabatic cases. The inclusion of non-adiabaticity leads to the radial displacement, ξ_r , being suppressed relative to the equilibrium value (by a factor of 10), and the tangential displacement, ξ_h , is amplified (by a factor of $\sim 10^2 - 10^3$).

Both ξ_r and ξ_h become negative in the surface region, leading to a phase shift of 90° between the planetary motion and the surface response of the star. On top of this, the imaginary parts of ξ_r is of similar magnitude to the real part, leading to an additional phase shift of 48° . The tangential displacement is dominated by its real part, leading to an additional phase shift of only 5° .

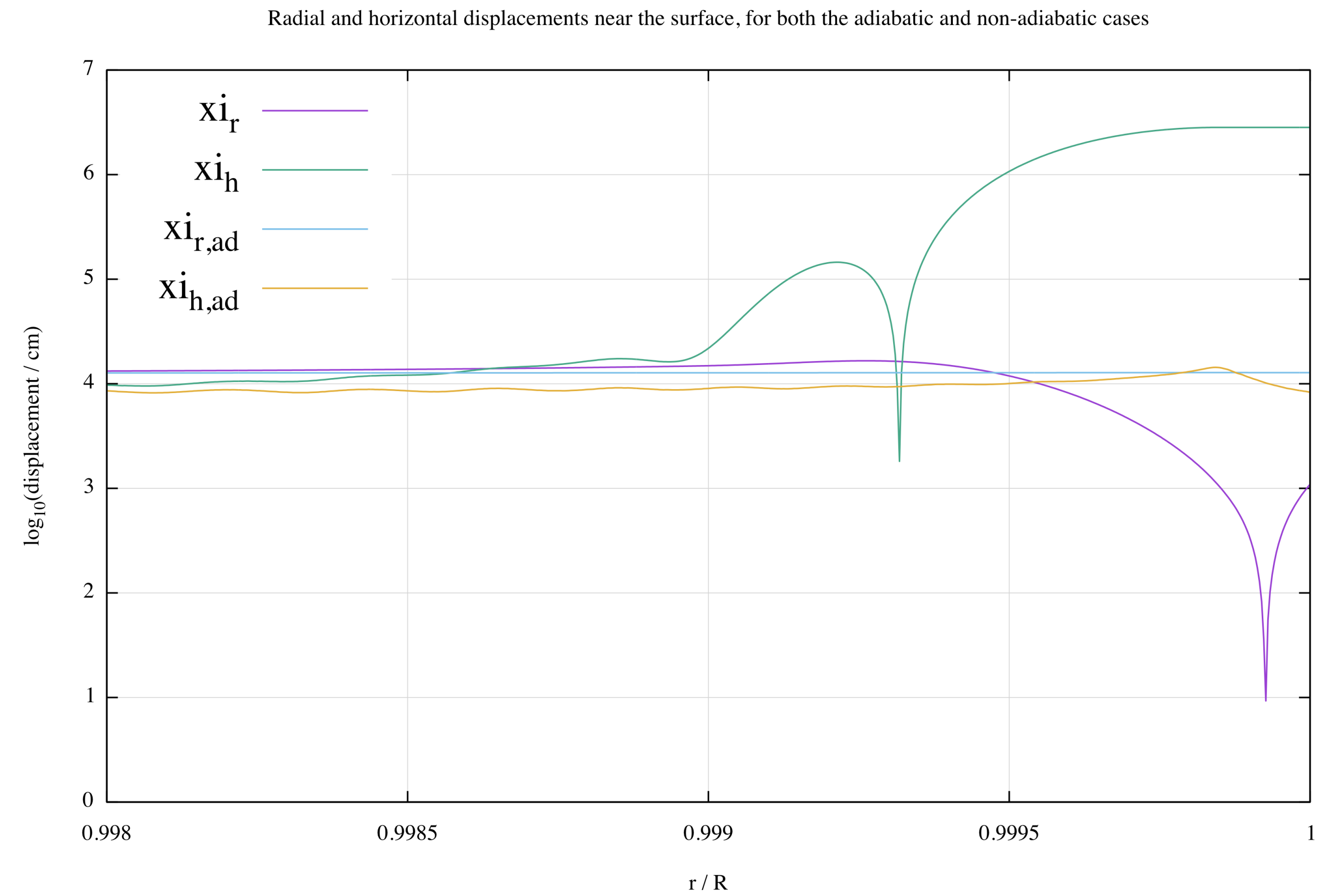


Figure 1: A close-up of the surface of a solar-type star, showing the logarithms of the real parts of radial and tangential displacements (ξ_r and ξ_h , respectively) for both the non-adiabatic and the adiabatic case. In the adiabatic case ξ_r and ξ_h are approximately similar at the surface. Taking non-adiabatic effects into account suppresses ξ_r , and greatly amplifies ξ_h by a factor of $\sim 10^2 - 10^3$.

The rough analytical estimate for the ratio $\frac{\xi_h}{\xi_r} \sim \frac{r}{H_p}$, where H_p is the pressure scale height gives a surface value of 5000, which is definitely the right order of magnitude. Using the more involved expression of equation 7, we can get a more precise value, of 1300 ± 100 , compared to the model’s value of 1280 ± 5 , where the errors are due to numerical scatter in the output.

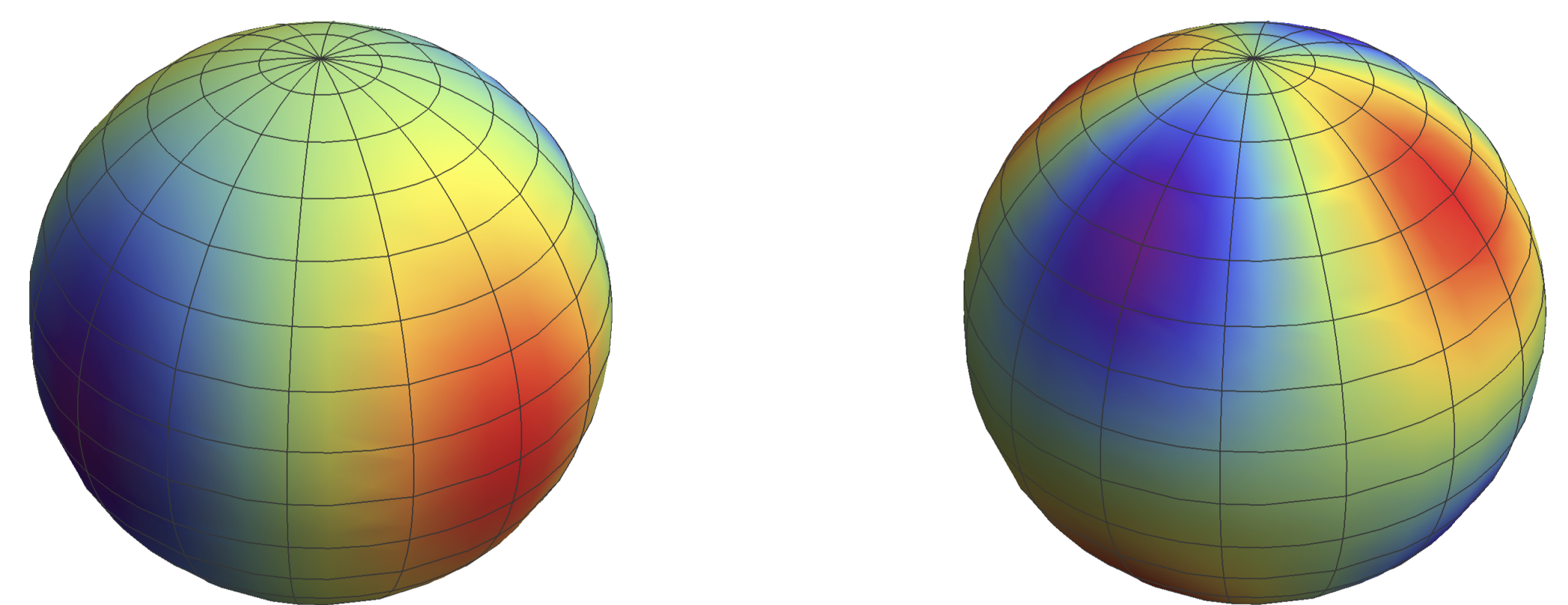


Figure 2: A heat map showing the distribution of radial displacement (ξ_r , left, corresponding to $\sin^2(\theta) \cos(2\phi + \phi_r)$, where ϕ_r is the argument of ξ_r at the surface) and tangential displacement (ξ_h , right, corresponding to $\sin(\theta) \cos(\theta) \cos(2\phi)$). Blue and red correspond to the large amplitudes of opposite phase. The radial displacement has a phase lag of 47° , whereas the tangential displacement has a phase lag of 5° . The ratio of amplitudes $\frac{\xi_h}{\xi_r} \approx 10^3$.

Due to the spatial distribution of the tangential displacement (shown in figure 2), it averages out over the disk. Therefore it does not contribute to a change in overall brightness or to the surface-averaged radial velocity. However, it would lead to a time-dependent spectroscopic signal in an oscillatory contribution to line-broadening. If this signal is detected, it could be possible to infer the mass and inclination of the planet.

Conclusions

- Non-adiabatic effects have a strong impact upon the behaviour of oscillations in the surface layer
- Tangential displacement is amplified by a factor of $\sim 10^2 - 10^3$ compared to the equilibrium tide case
- Analytical expressions using the boundary conditions agree well with the results of the numerical model
- This behaviour could be utilised to characterise planetary systems, or potentially detect exoplanets

Forthcoming Research

The next step is to calculate the photometric and spectroscopic effects of these oscillations (both in terms of disk-averaged RV signal and time-dependent line broadening) in order to look for such signal around stars with known hot Jupiters in order to test the feasibility of this as a method.

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Acknowledgements

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