

Quantum Cosmology

MOENS

ANNALES
DE LA
SOCIÉTÉ SCIENTIFIQUE
DE BRUXELLES

EXTRAIT

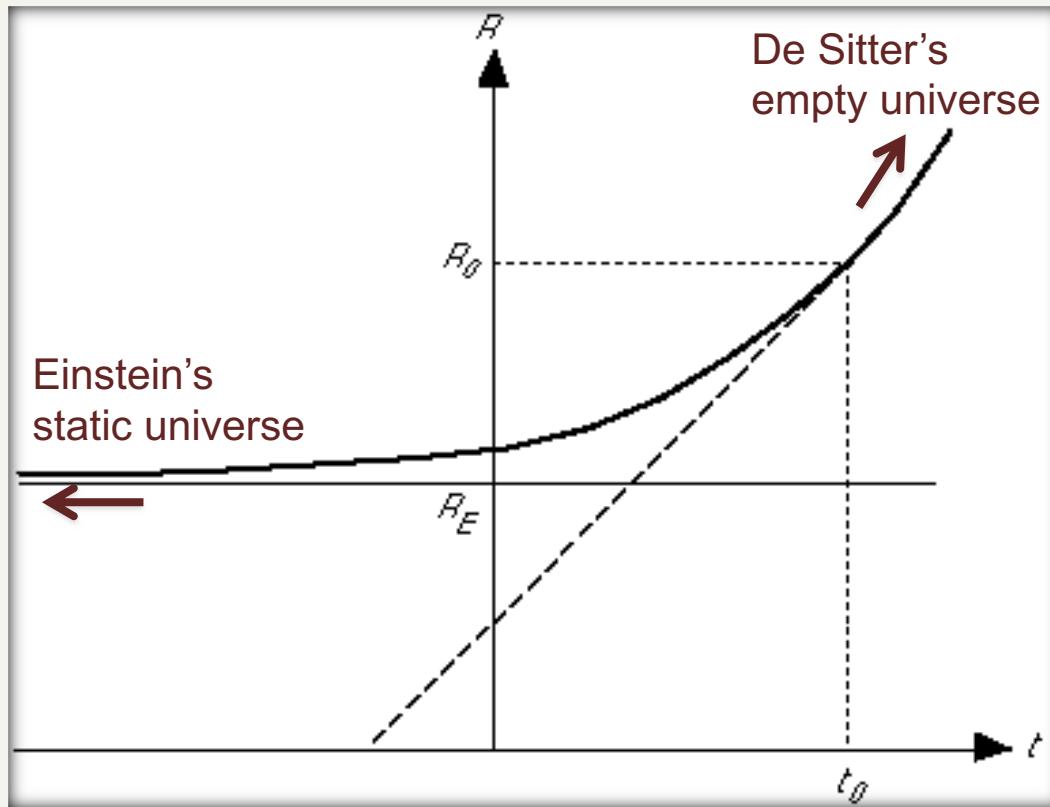
Un univers homogène de masse constante
et de rayon croissant, rendant compte
de la vitesse radiale des nébuleuses
extra-galactiques

Note de M. l'Abbé G. LEMAITRE

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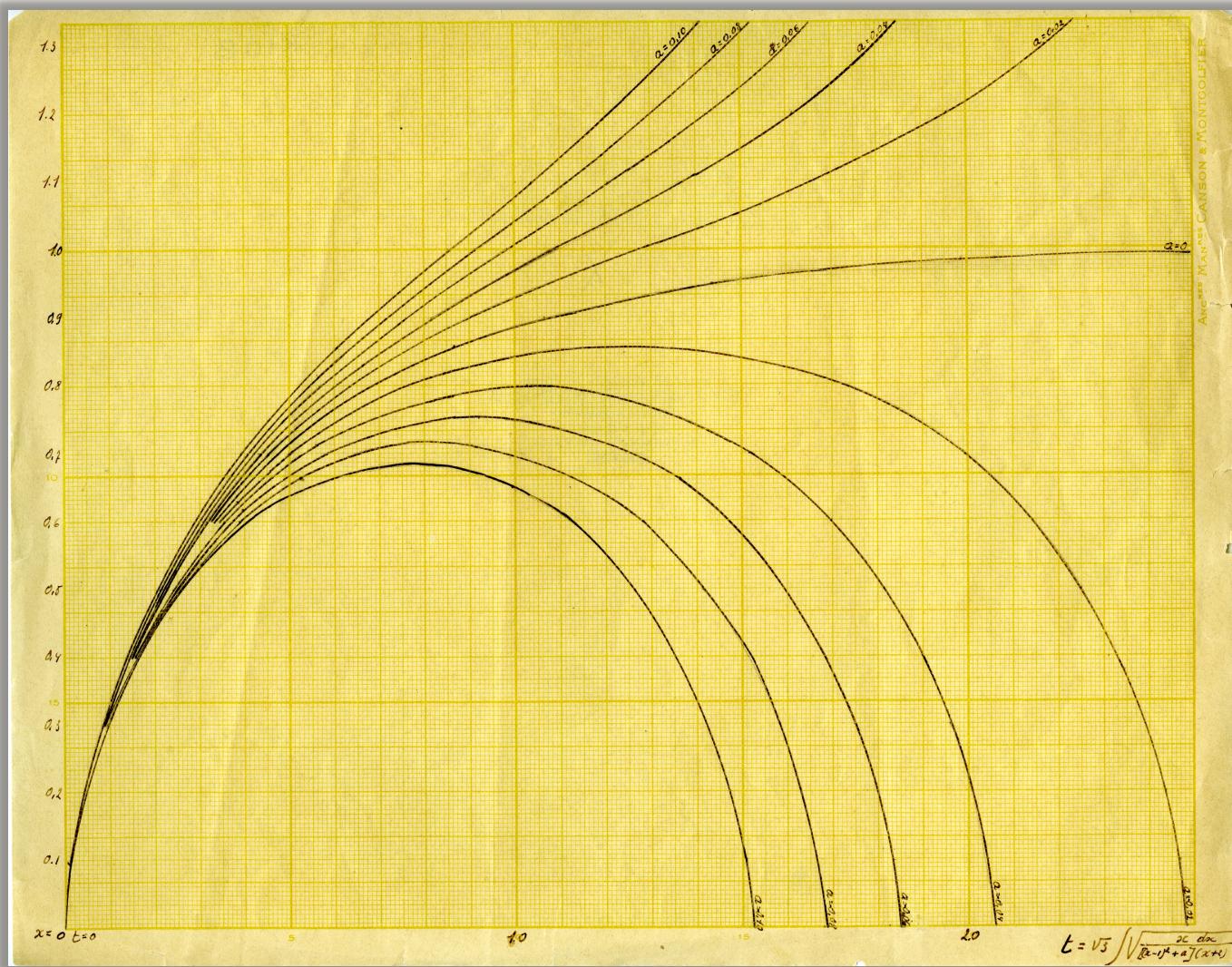
1927



Utilisant les 42 nébuleuses figurant dans les listes de Hubble et de Strömgberg (¹), et tenant compte de la vitesse propre du soleil (300 Km. dans la direction $\alpha = 315^\circ$, $\delta = 62^\circ$), on trouve une distance moyenne de 0,95 millions de parsecs et une vitesse radiale de 600 Km./sec, soit 625 Km./sec à 10^6 parsecs (²).

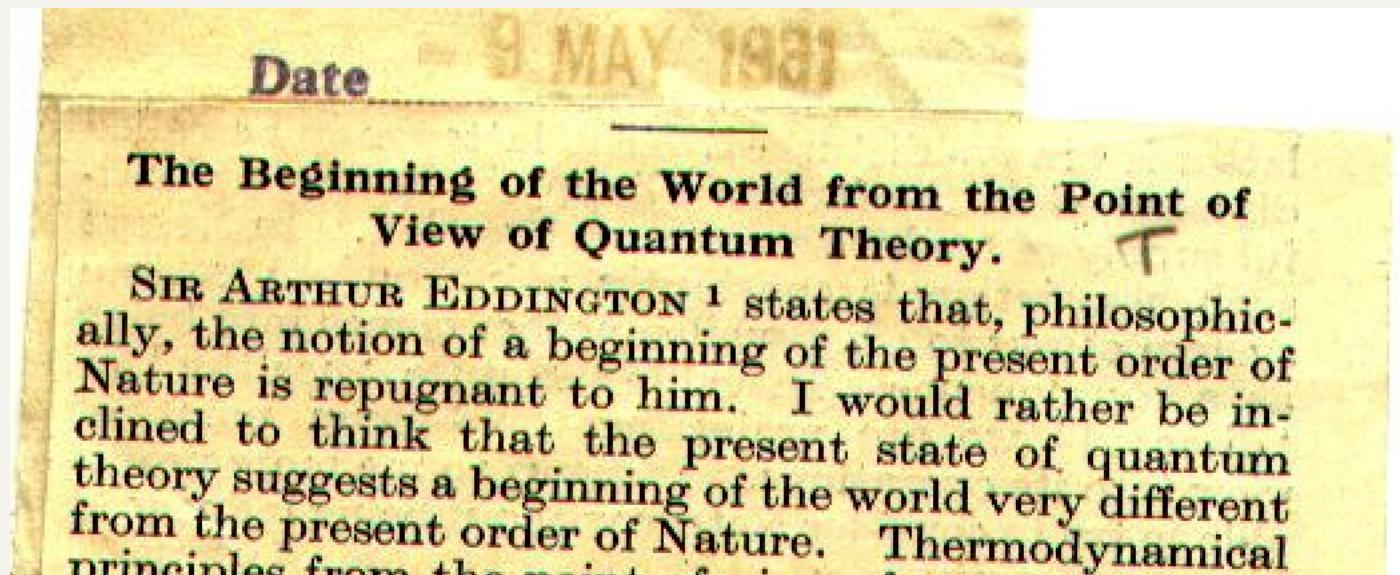
Nous adopterons donc

$$\frac{R'}{R} = \frac{v}{rc} = \frac{625 \times 10^5}{10^6 \times 3,08 \times 10^{18} \times 3 \times 10^{10}} = 0,68 \times 10^{-27} \text{ cm}^{-1} \quad (24)$$



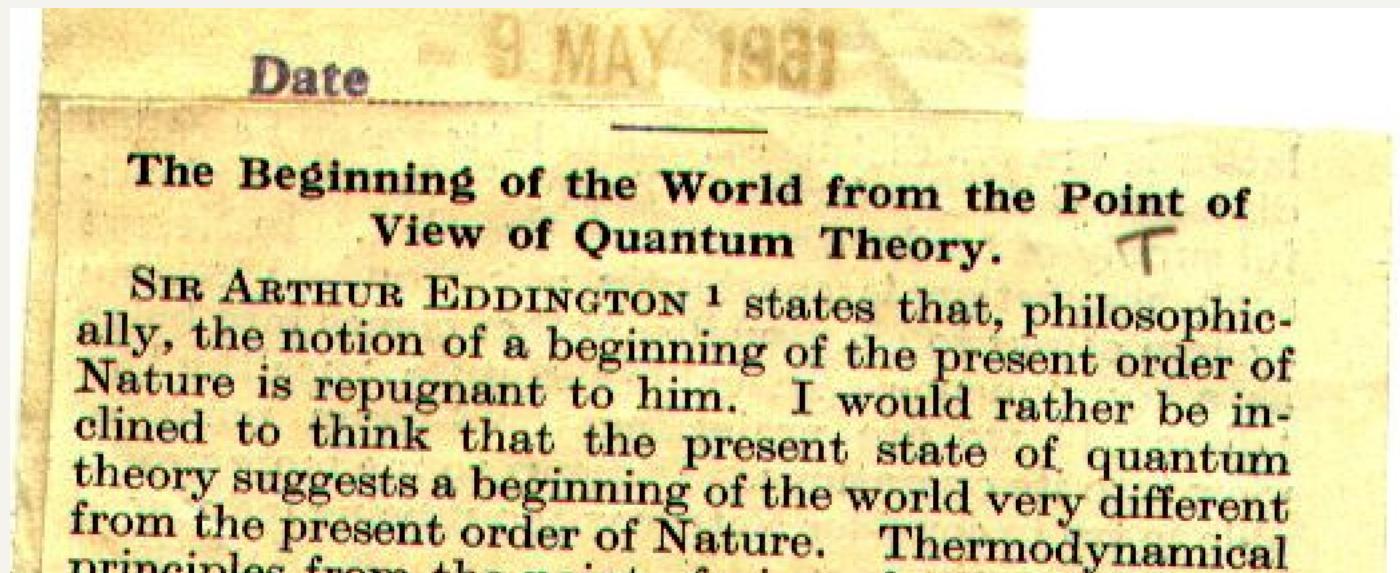
G. Lemaître, 1929

Primeval atom



"If we go back in the course of time ... the notions of space and time would altogether fail to have any meaning ..."

Primeval atom



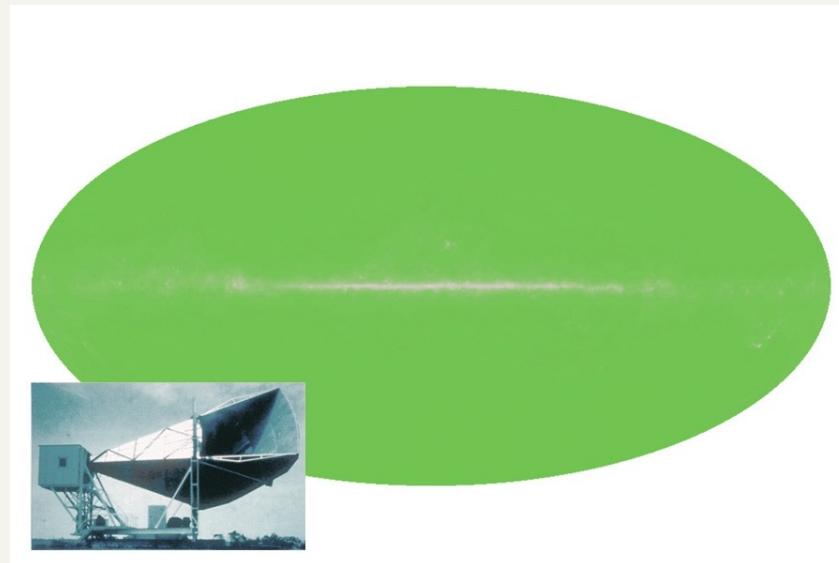
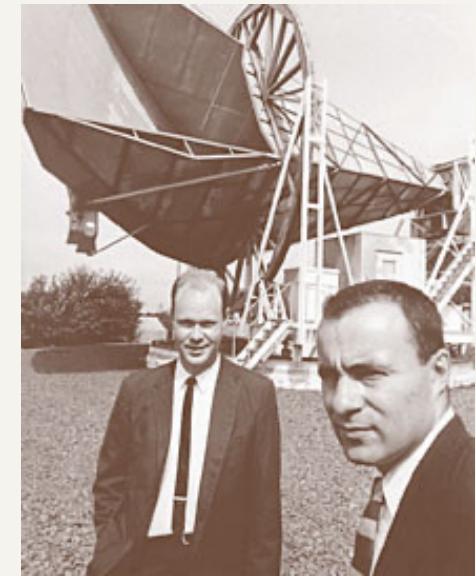
"If we go back in the course of time ... the notions of space and time would altogether fail to have any meaning ...

Clearly the initial quantum could not conceal in itself the whole complex course of evolution. The story of the world need not have been written down in the first quantum like the song on a disc of a phonograph ...

Instead from the same beginning widely different universe could have evolved."

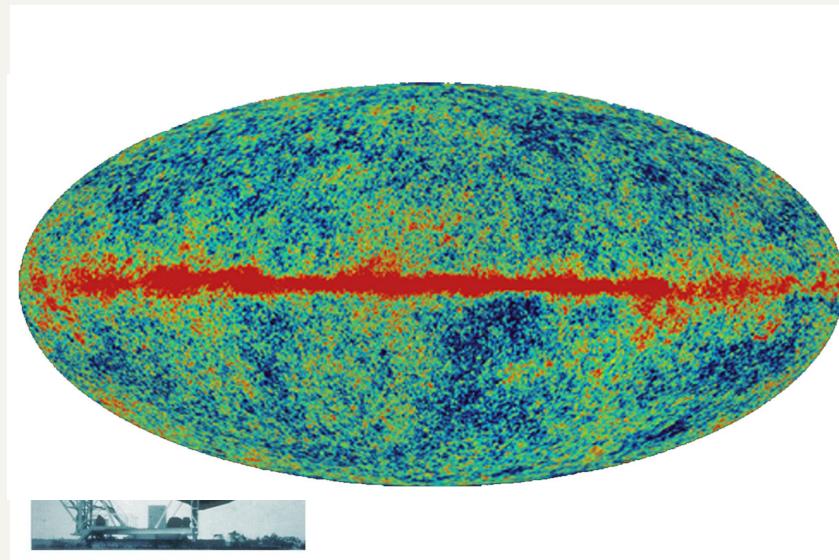
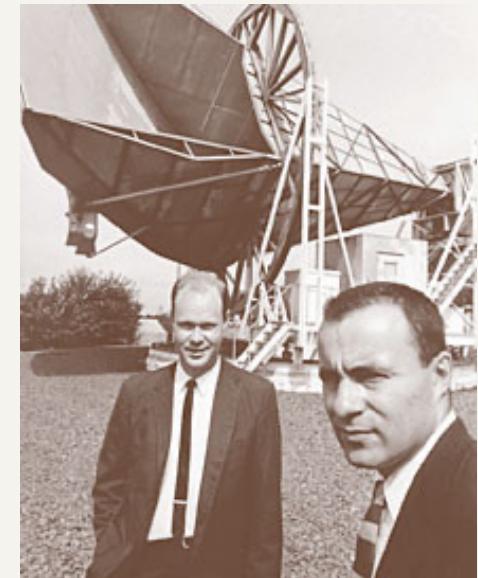
1965

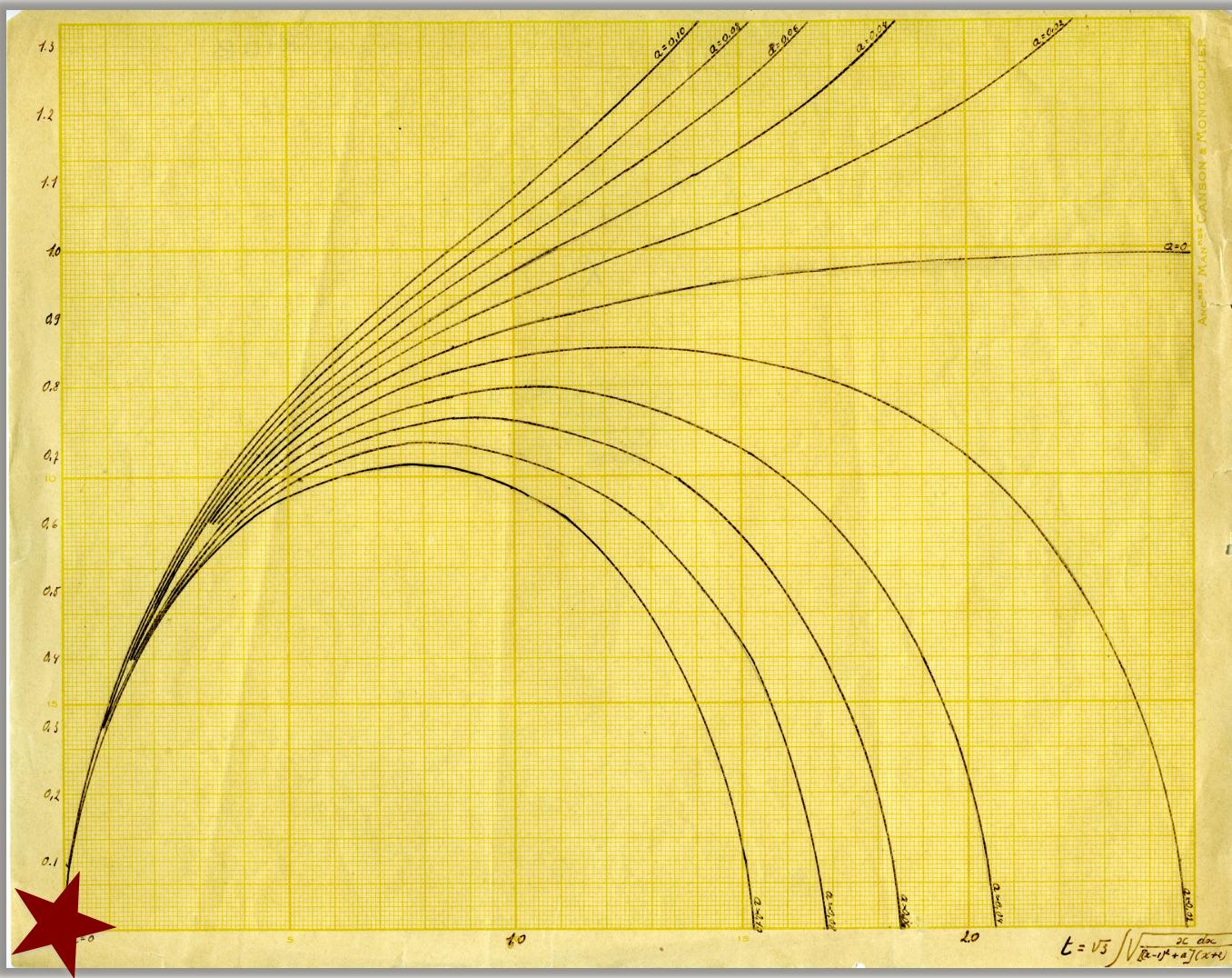
- Penzias and Wilson discover microwave background radiation at temperature 3 Kelvin
- Dicke, Peebles, Roll & Wilkinson provide cosmological interpretation as fossil radiation in the framework of 'Big Bang' theory



1965

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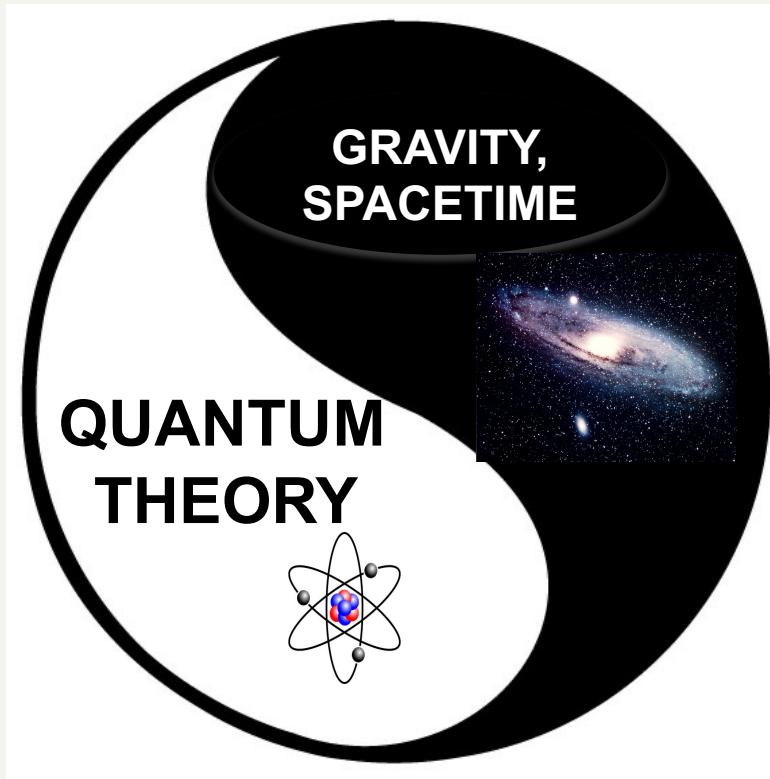


G. Lemaître, 1929

Why bother with Quantum Cosmology?

- critical expansion rate
- low entropy
- small Λ
- $\delta T/T \sim 10^{-5}, \dots$
- relative strength of forces
- weak gravitation
- three large dimensions, ...
- classical space-time itself!

From Quantum to Cosmos



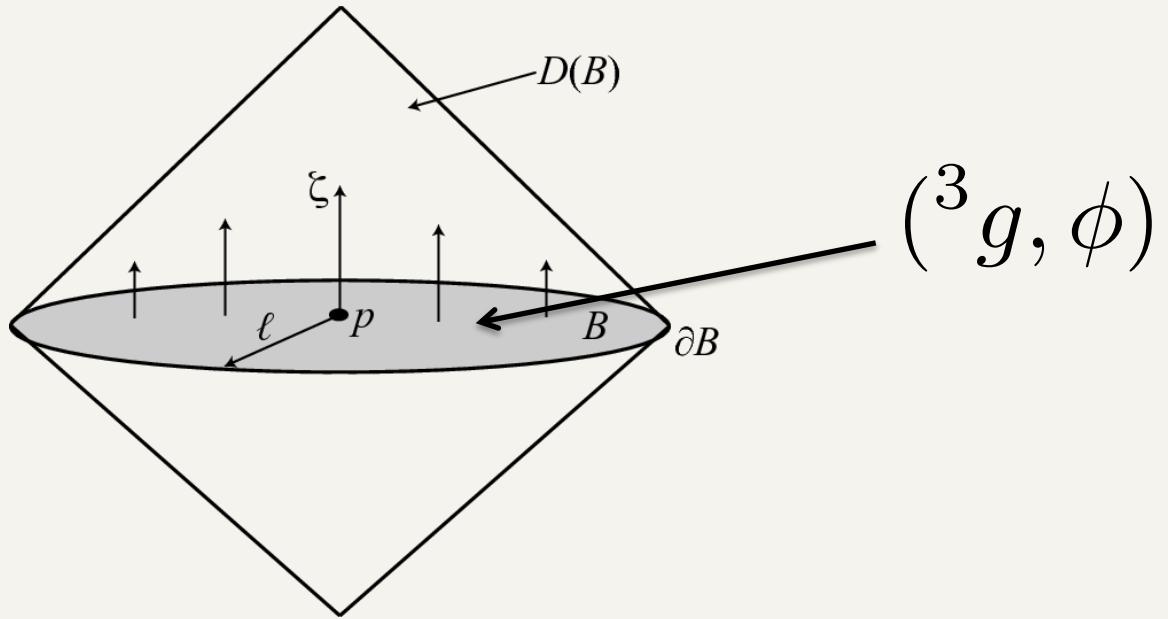
$$g \rightarrow \Psi[g]$$

Outline

- **Wave function of the Universe**
- How to put the wave function to use?!
- Recent developments -> holographic cosmology

Primeval Atom 2.0

$$\Psi[{}^3g, \phi, \dots]$$



‘No-Boundary’ Wave Function

[Hartle & Hawking, 1983]

‘Ground state’ analogue

$$\Psi[{}^3g, \phi] = \int_C \delta g \delta \phi \exp(-I_E[{}^3g, \phi]/\hbar)$$

“The amplitude of a configuration $({}^3g, \phi)$ on a spacelike surface Σ is given by a Euclidean path integral over the class C of all regular four-geometries and matter field histories that match $({}^3g, \phi)$ on their only boundary.”

Ground state

Ground state in Quantum Mechanics

Euclidean PI: $\psi(x_0) = \int \delta x \exp\{-I[x(\tau)]/\hbar\}$

with $I[x(\tau)] = \frac{1}{2} \int d\tau [\dot{x}^2 + \omega^2 x^2]$

Saddle pt appr: $\psi(x_0) \propto \exp\{-\omega x_0^2/2\}$

‘No-Boundary’ Wave Function

‘Ground state’ analogue

$$\Psi[{}^3g, \phi] = \int_C \delta g \delta \phi \exp(-I_E[{}^3g, \phi]/\hbar)$$

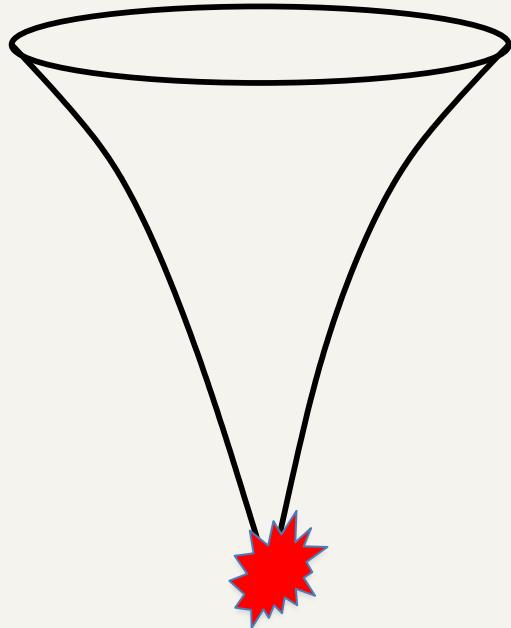
“The boundary condition of the universe is that it has no boundary.”

[Hawking, Pont. Ac. Sci., 1982]

$$I[g, \phi] = -\frac{1}{2} \int \sqrt{g} (R - 2\Lambda) + \int \sqrt{g} [(\nabla \phi)^2 + V(\phi)]$$

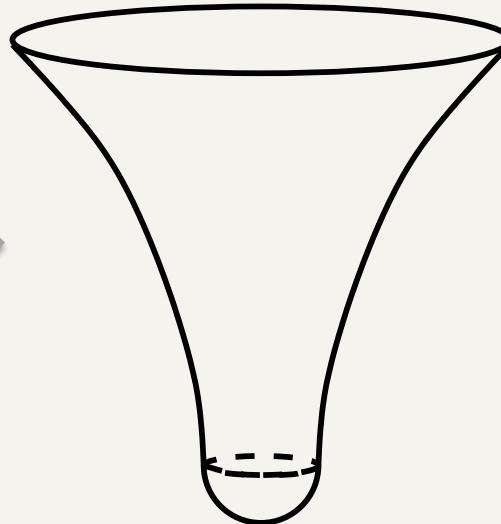
'No-Boundary' Wave Function

classical cosmology



initial singularity

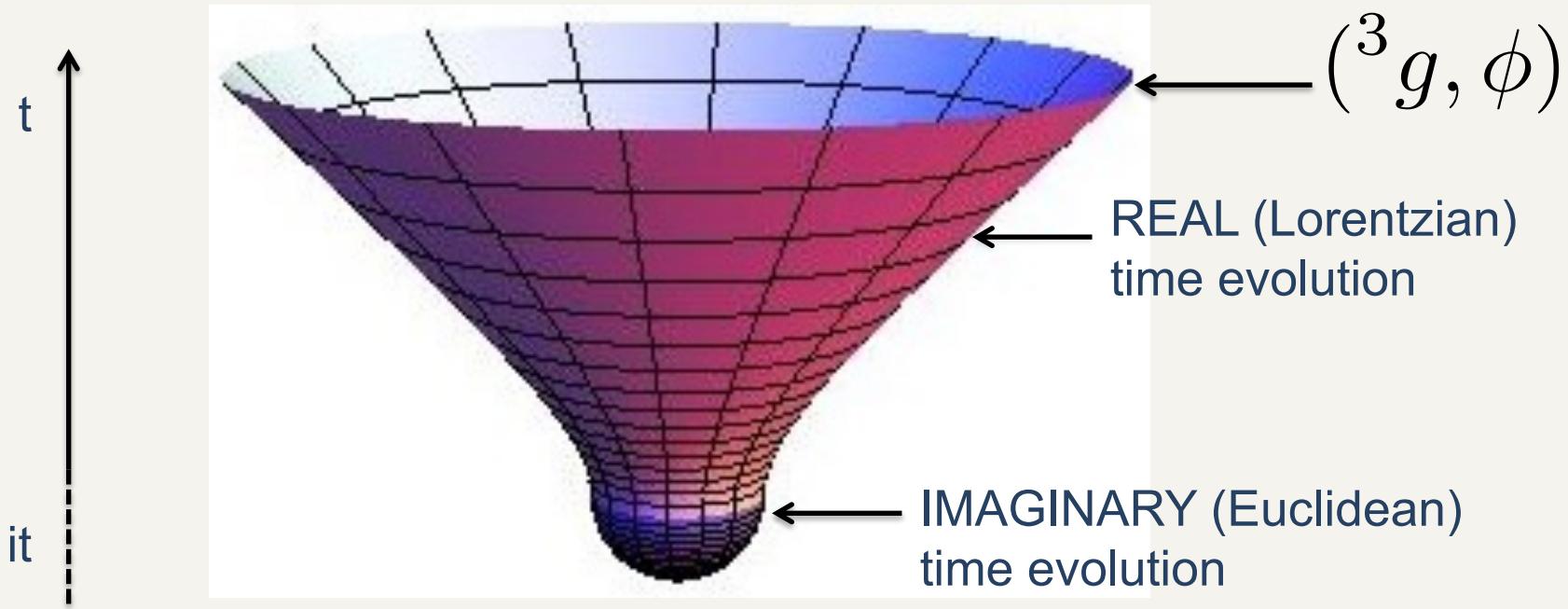
quantum cosmology



regular, no boundary

Example: creation from nothing of de Sitter space

'No-Boundary' wave function



$$\Psi[^3g, \phi] \approx \exp(-I_E[^3g, \phi]/\hbar) = A \exp(iS)$$

$$I(^3g, \chi) = -I_R(^3g, \chi) + iS(^3g, \chi)$$

‘No-Boundary’ wave function

WKB approximation:

$$\Psi[{}^3g, \phi] \approx A \exp(iS)$$

$$|\nabla A| \ll |\nabla S| \quad \longrightarrow$$

classical cosmological evolution

$$p_q = \nabla_q S$$

$$P_{hist} \propto A^2 = \exp(-2I_R/\hbar)$$

What universes emerge?

Consider homogeneous/isotropic multiverse:

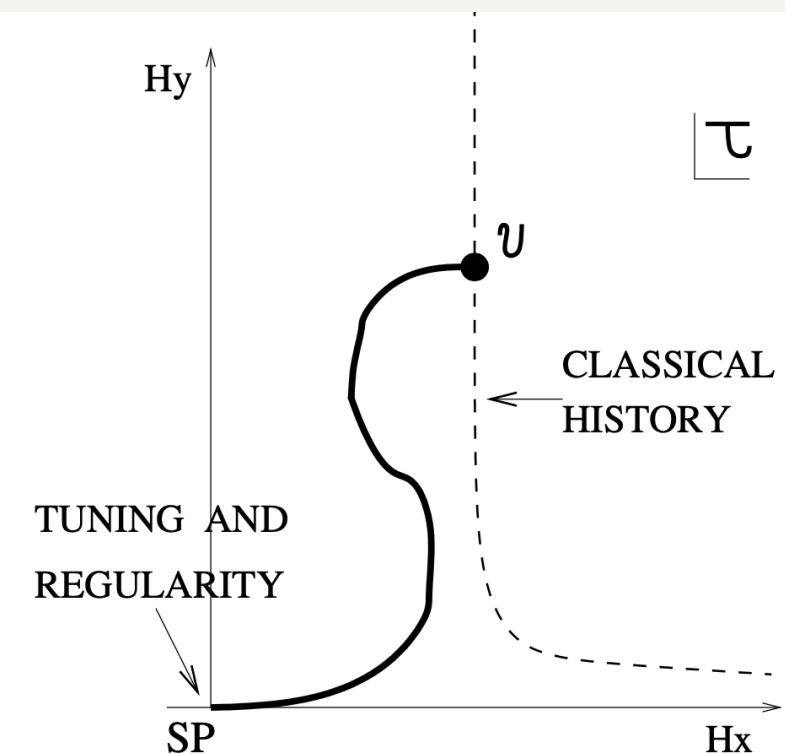
$$ds^2 = d\tau^2 + a^2(\tau)d\Omega_3, \quad \phi(\tau), \quad \Psi[b, \chi]$$

$$V(\phi) = \Lambda + \frac{1}{2}m^2\phi^2$$

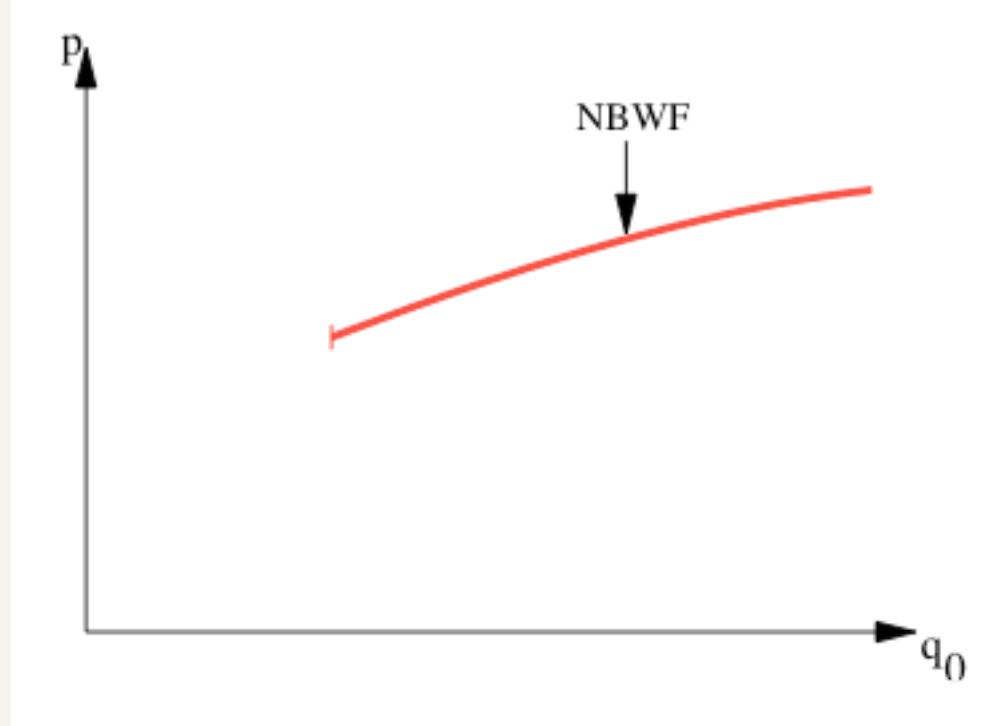
Regularity at SP: $g_{ij}(0) \rightarrow 0, \quad \dot{\phi}(0) \rightarrow 0$

At final boundary: $g_{ij}(v) = {}^3g, \quad \phi(v) = \chi$

Tuning at SP: $\phi(0) = \phi_0 e^{i\gamma}, \dots$



Multiverse

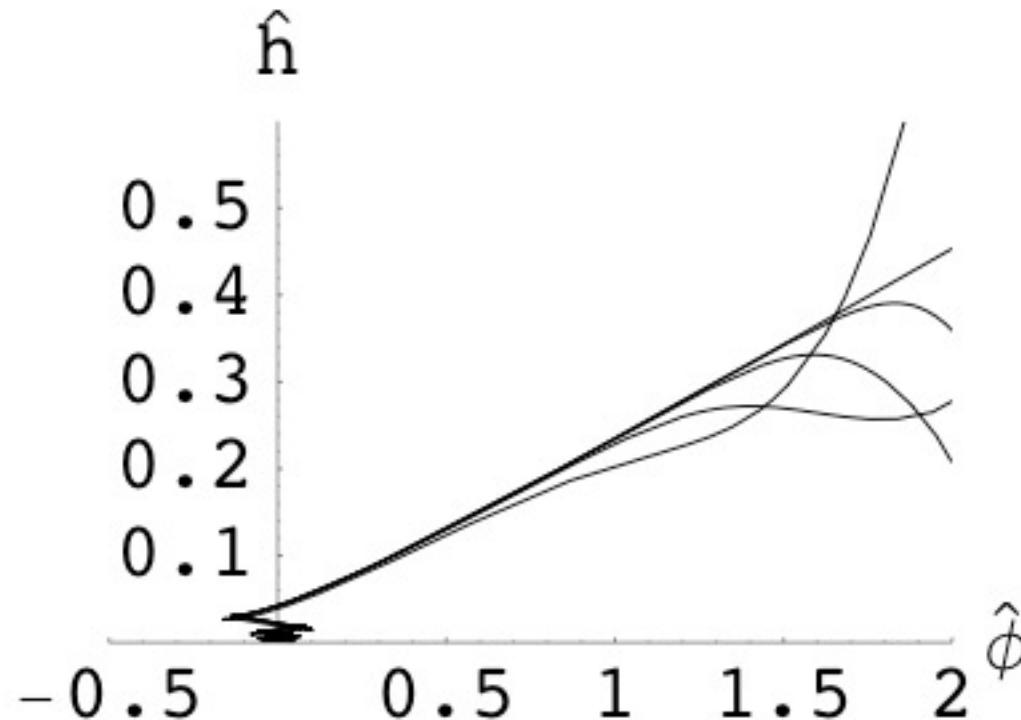


→ Measure on phase space of FRW cosmologies

Inflation!

[Hartle, Hawking, TH 2008]

$$p_q = \nabla_q S$$

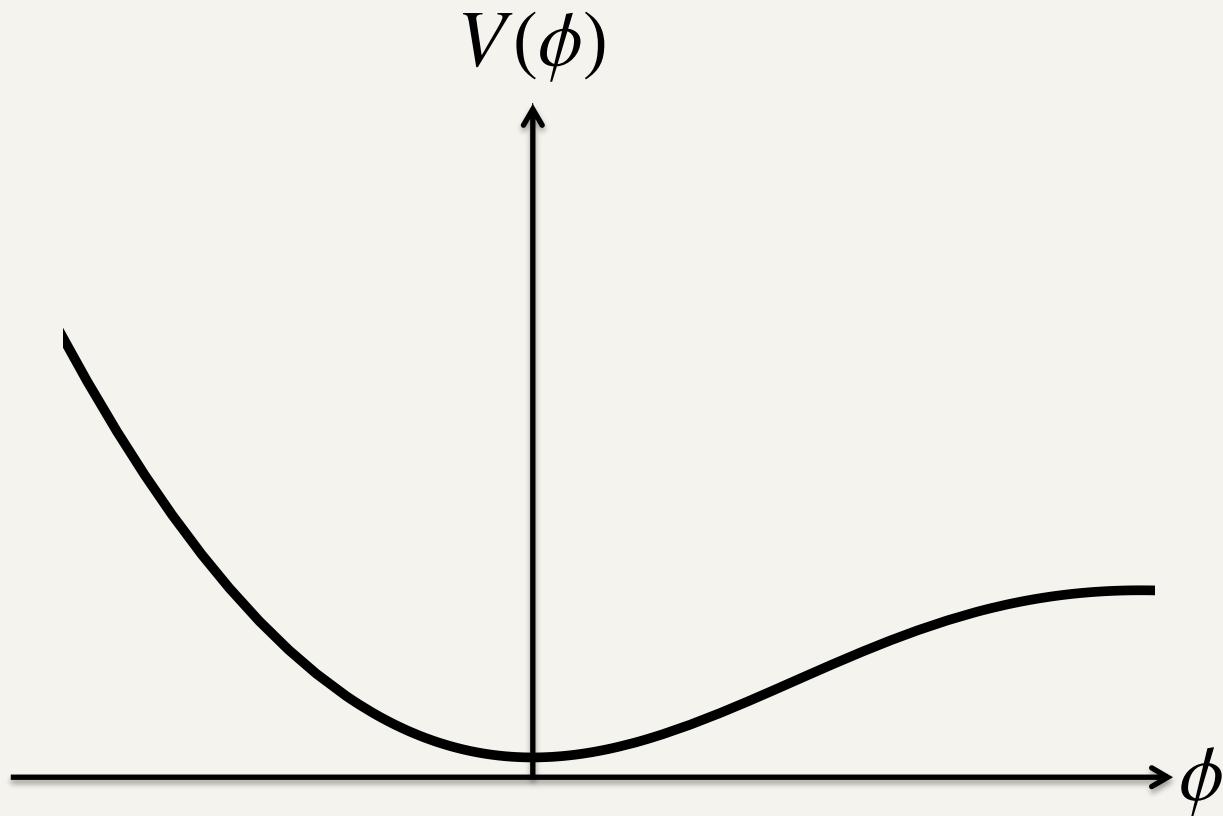


Early period of inflation emerges as a *prediction*

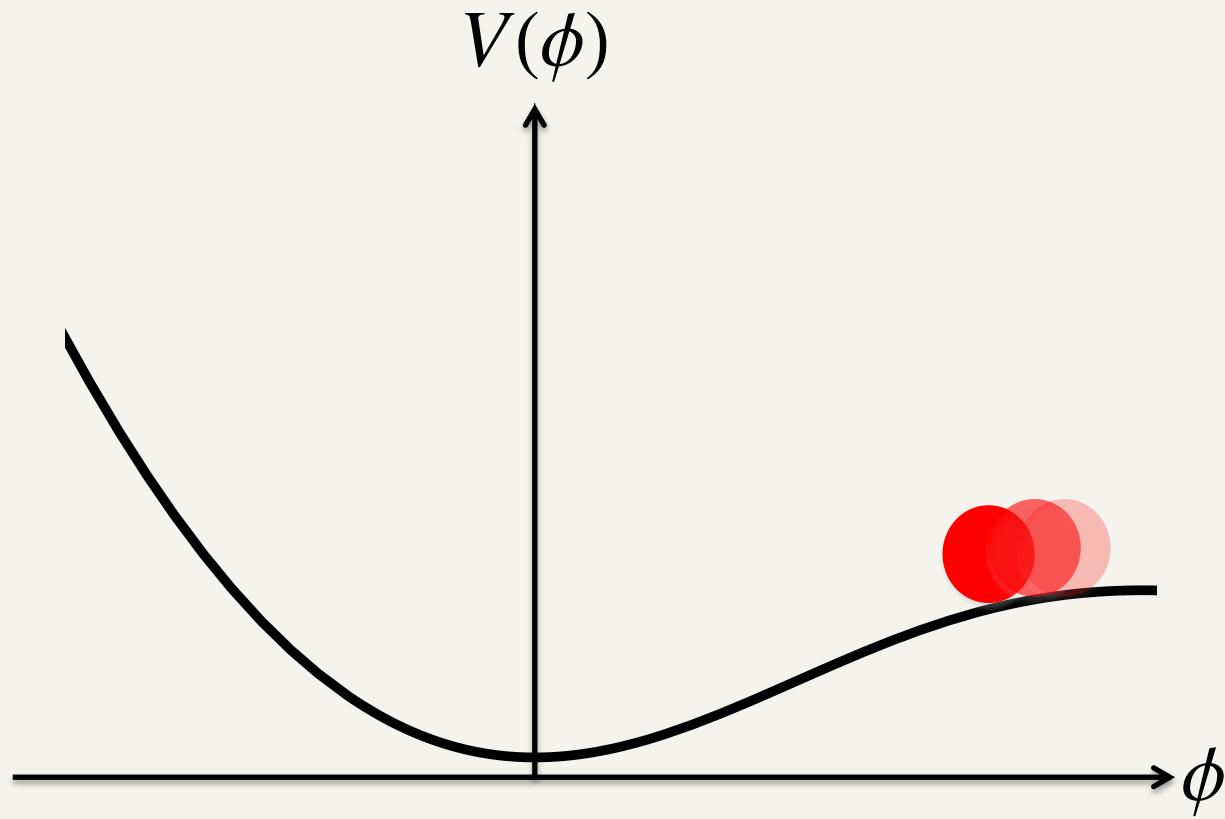
Outline

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- **How to put the wave function to use?!**
- Recent developments -> holographic cosmology

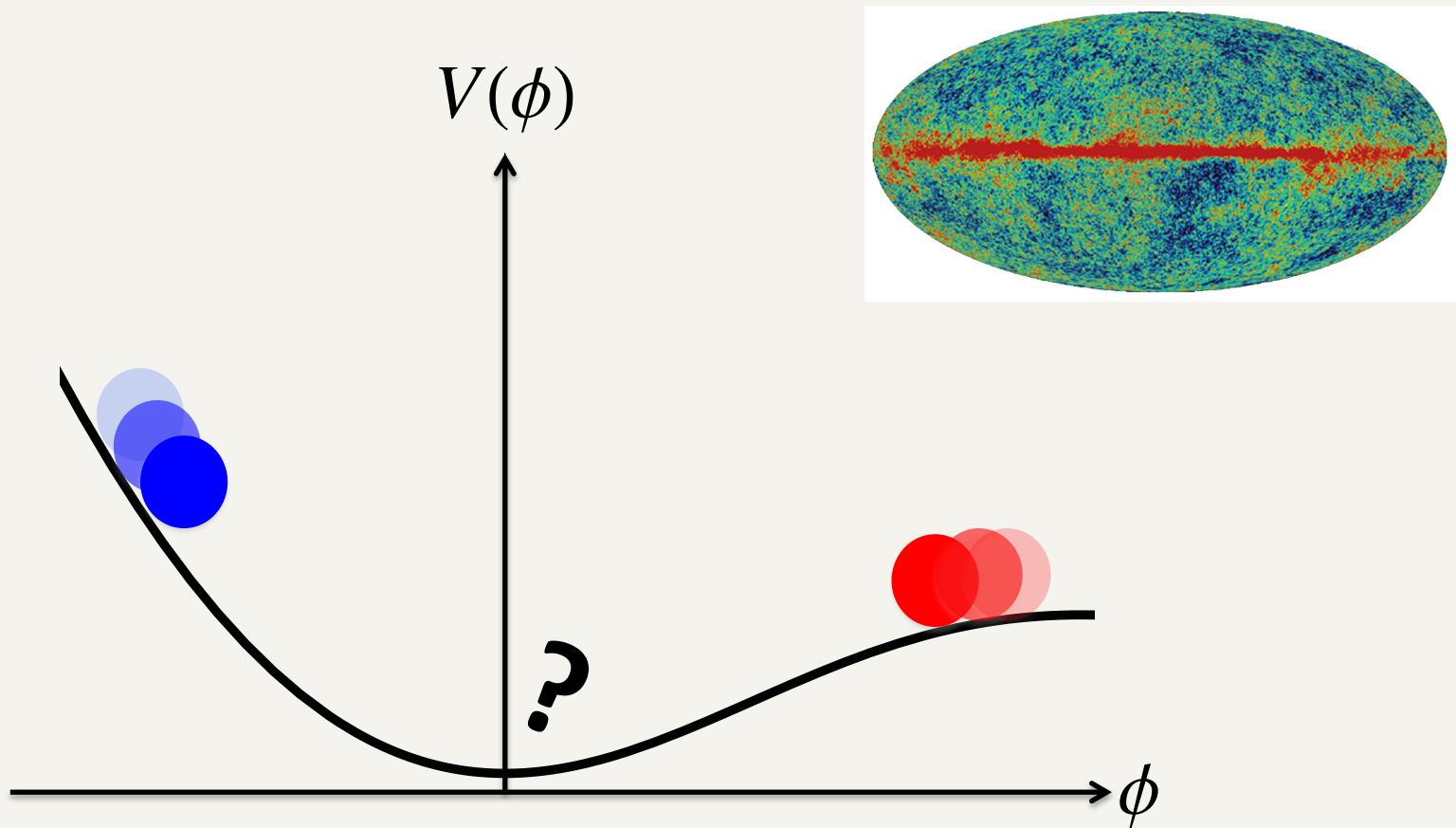
Example: a toy model of inflation



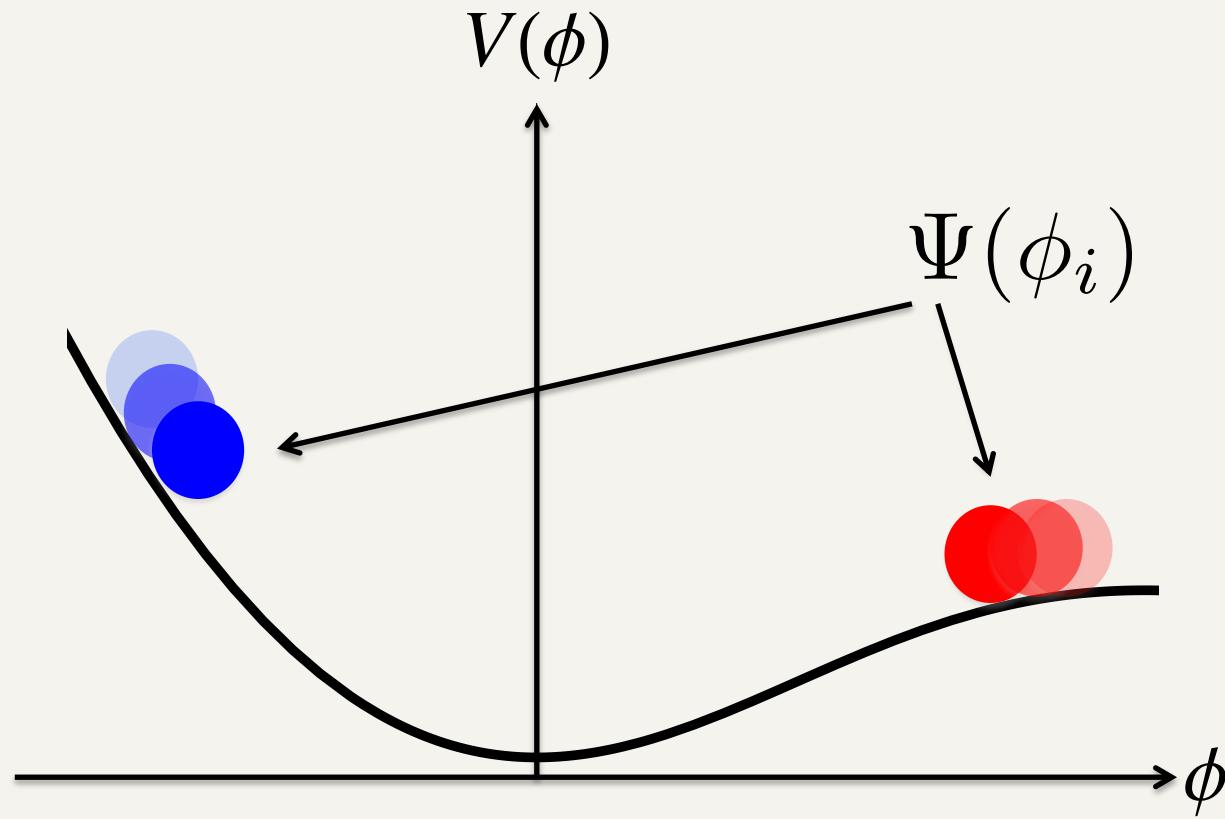
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Example: a toy model of inflation

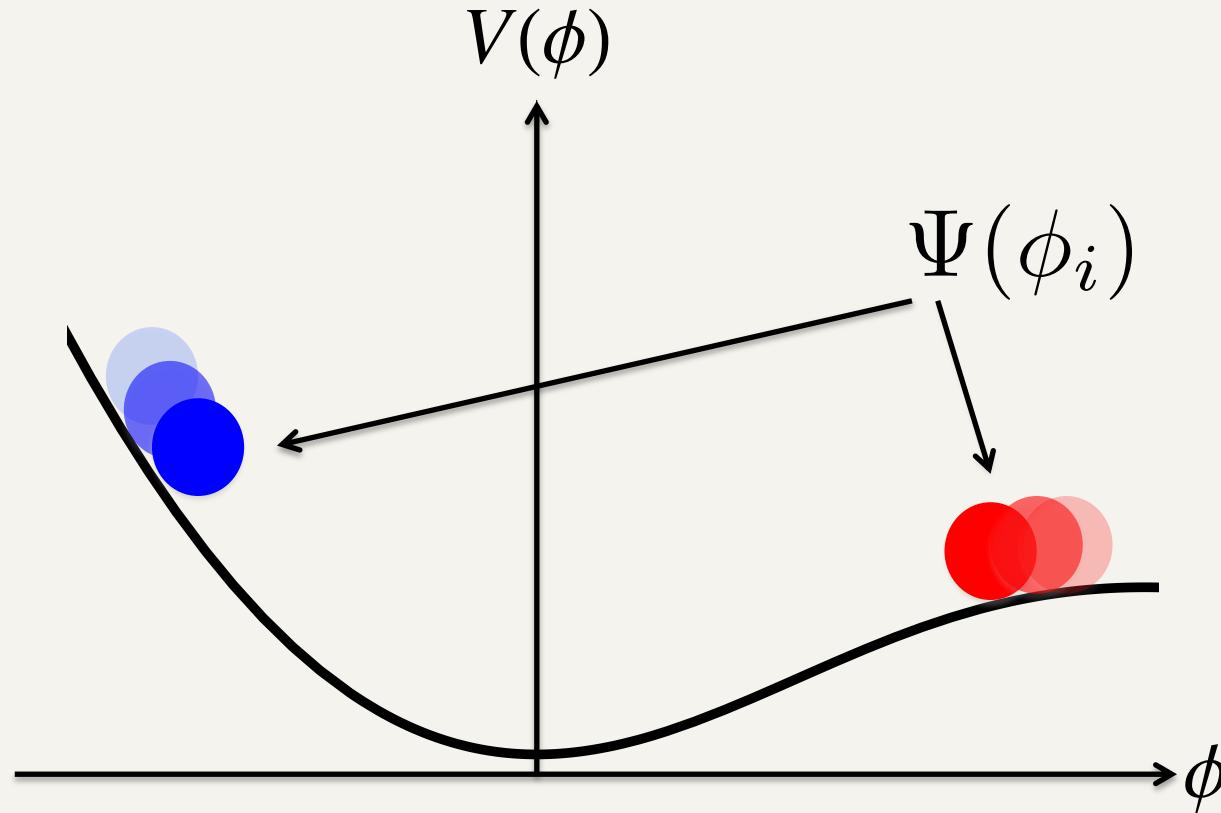


Example: a toy model of inflation



$$|\Psi|^2 \propto \exp\left(\frac{\pi}{4V(\phi_i)}\right)$$

Example: a toy model of inflation

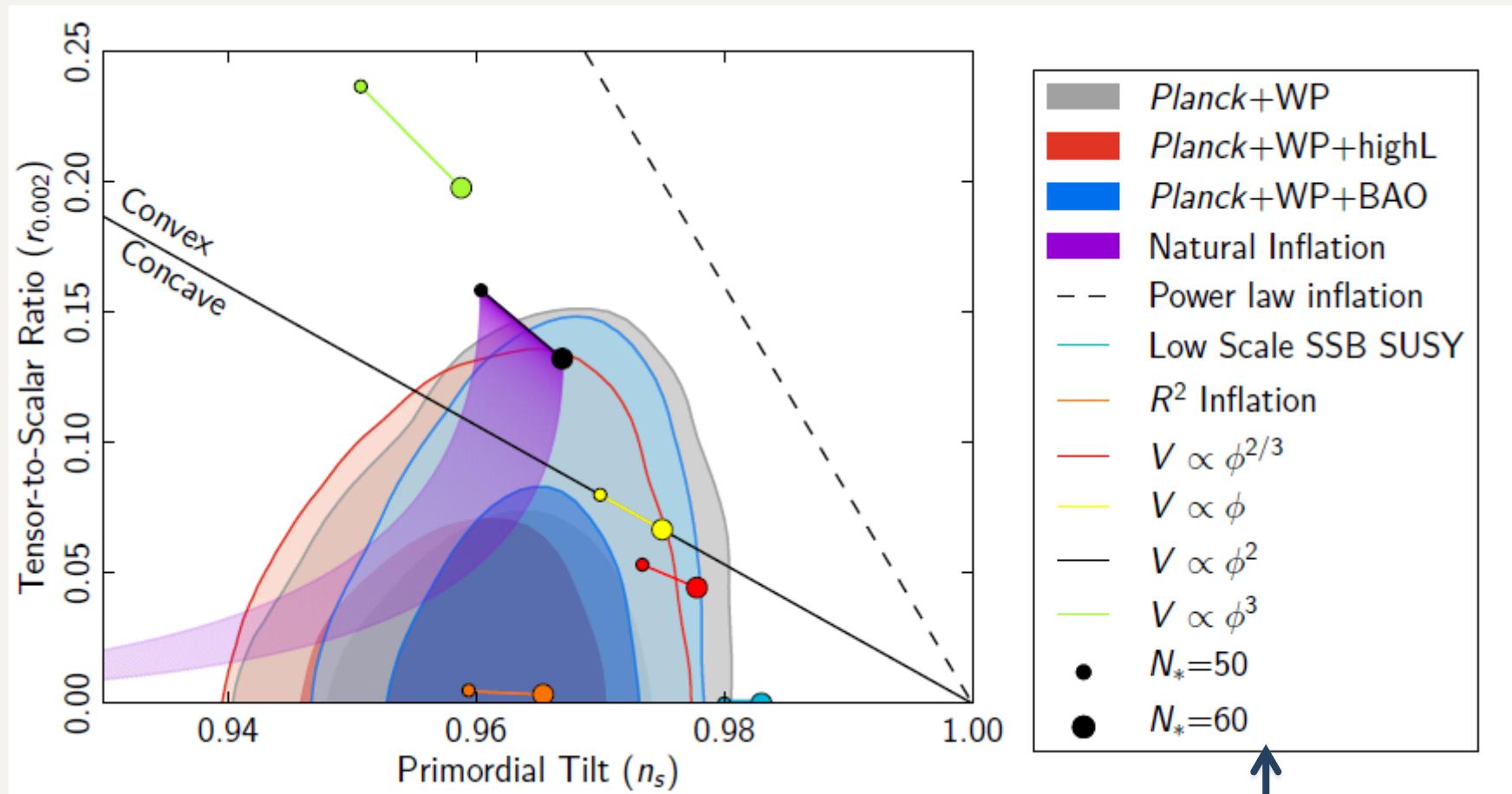


$$|\Psi|^2 \propto \exp\left(\frac{\pi}{4V(\phi_i)}\right) \rightarrow \text{Relative weighting between different 'models of inflation'}$$

-> selection principle

Application: a Prior for Planck

[Planck XXII 2013, TH 2013]

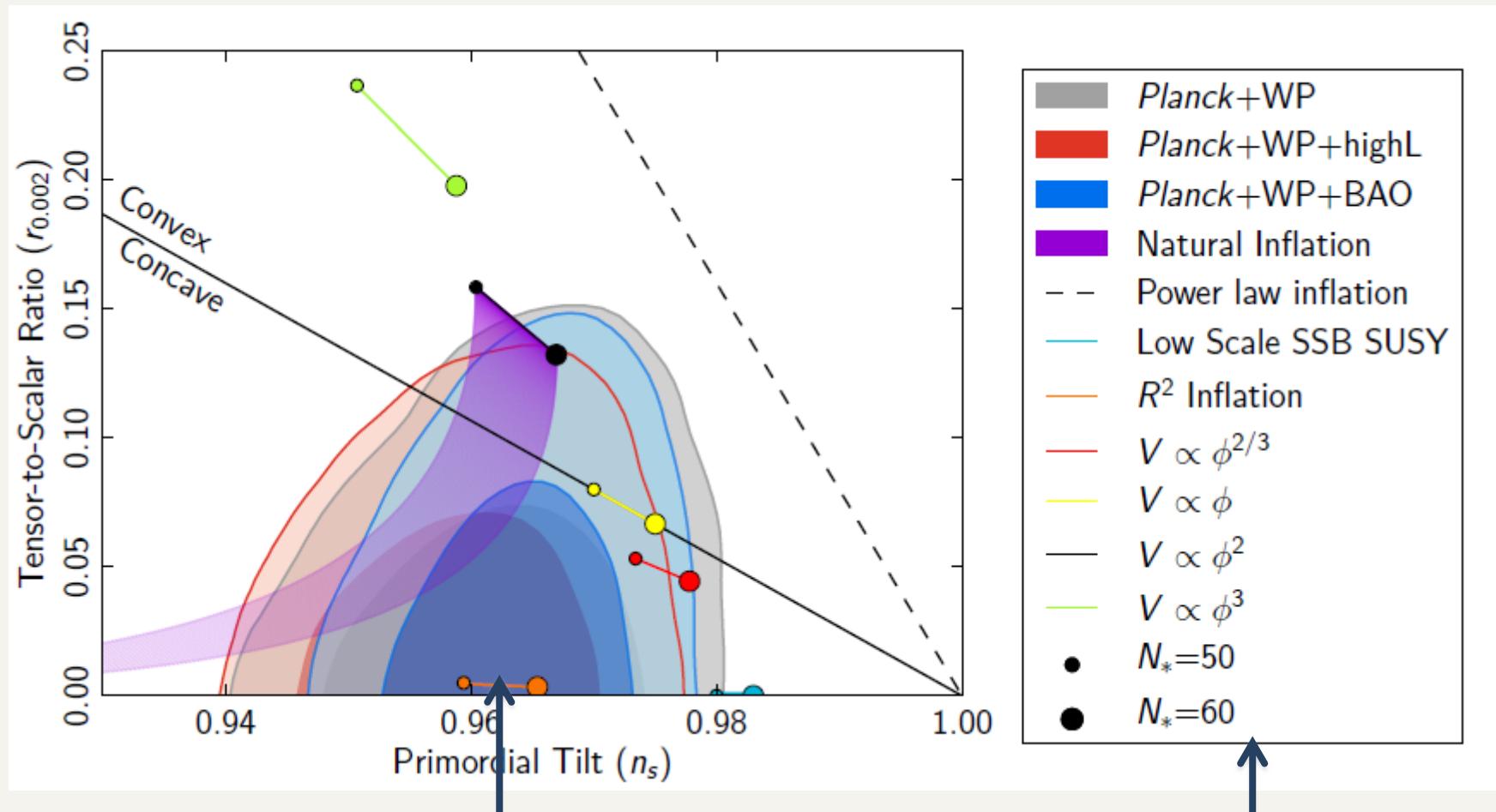


flat prior

'Landscape'

Application: a Prior for Planck

[Planck XXII 2013, TH 2013]



$$\text{Prior: } \Psi_{HH} \propto e^{-I_E}$$

'Landscape'

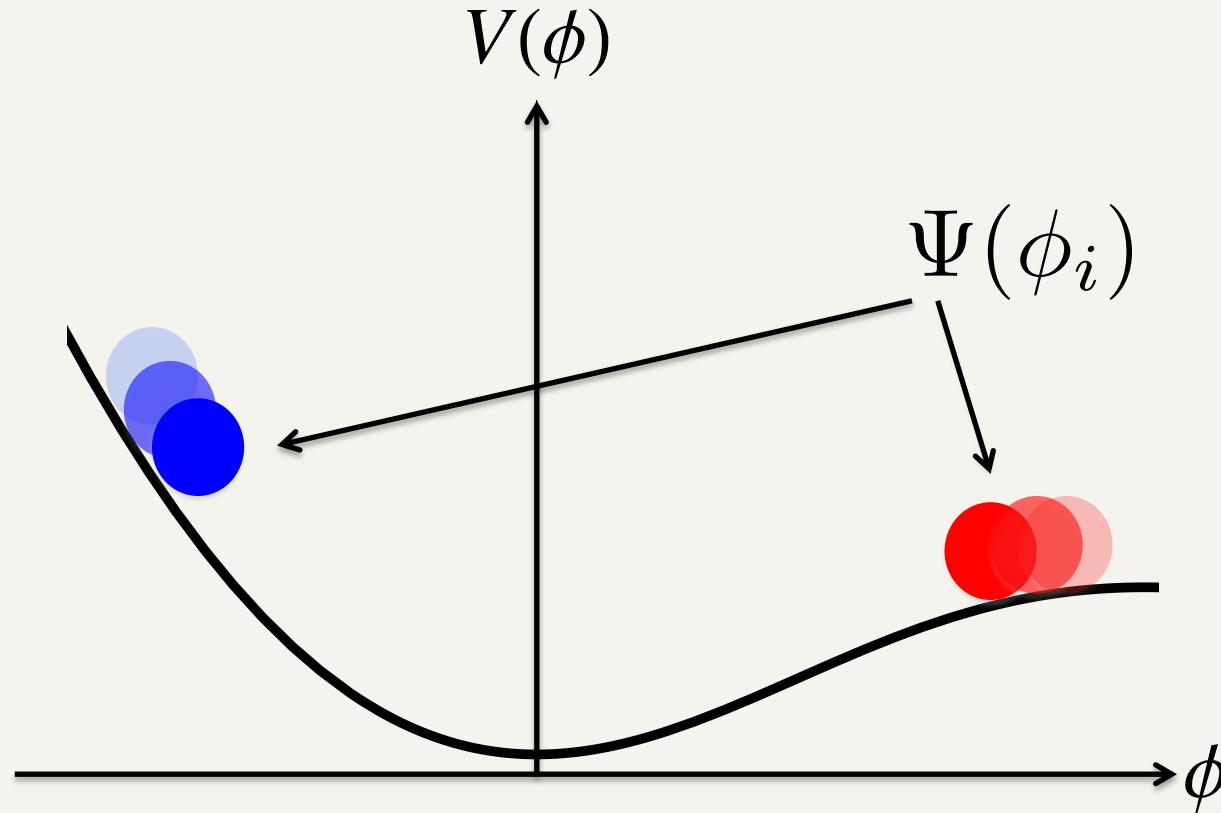
$$|\Psi|^2$$



Probabilities

!

Example: a toy model of inflation



$$|\Psi|^2 \propto \exp\left(\frac{\pi}{4V(\phi_i)}\right) \rightarrow \text{Relative weighting between different 'models of inflation'}$$

-> selection principle

Predictions for Observations

[Everett 1967; Hartle, TH 2013]

Predictions for observations are conditional probabilities for observables O , given our physical observational situation D within the universe.

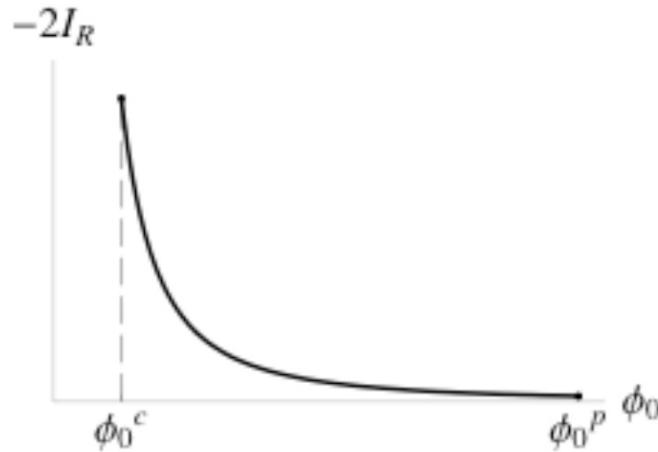
$$P(O|D)$$

→ “Anthropic” selection automatic in quantum cosmology

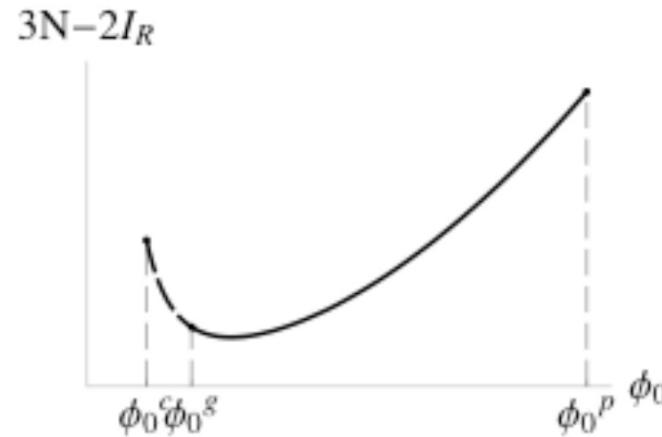
$$P(O|D) \propto \sum_J \int P(D|O, \phi_i^J) P(O, \phi_i^J)$$

Predictions for Observations

$p(hist)$



$p(obs)$



$$\sim \exp[+1/V(\phi_0)]$$

$$\sim \exp[3N + 1/V(\phi_0)]$$

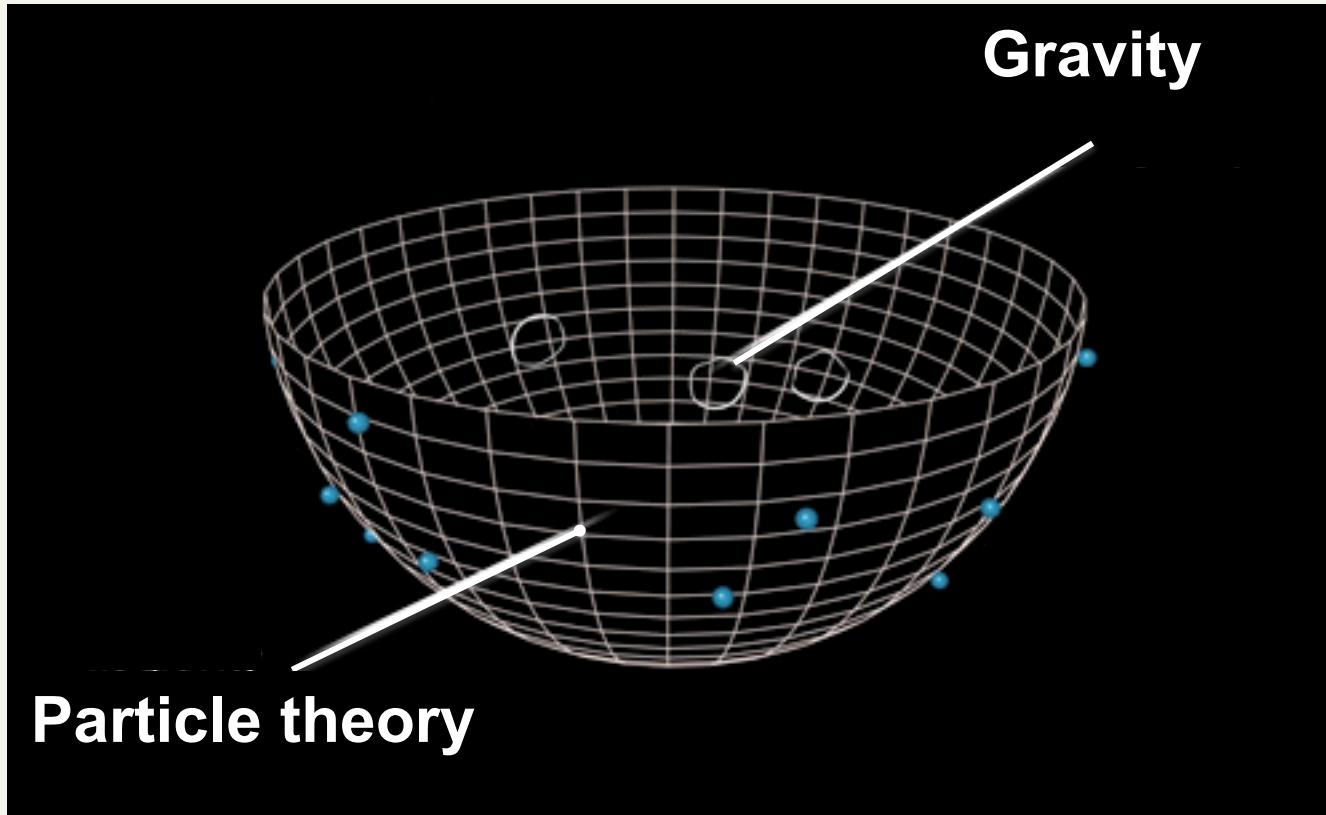
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$$\Psi[{}^3g, \phi] = \int_C \delta g \delta \phi \exp(-I_E[{}^3g, \phi]/\hbar)$$

Holography

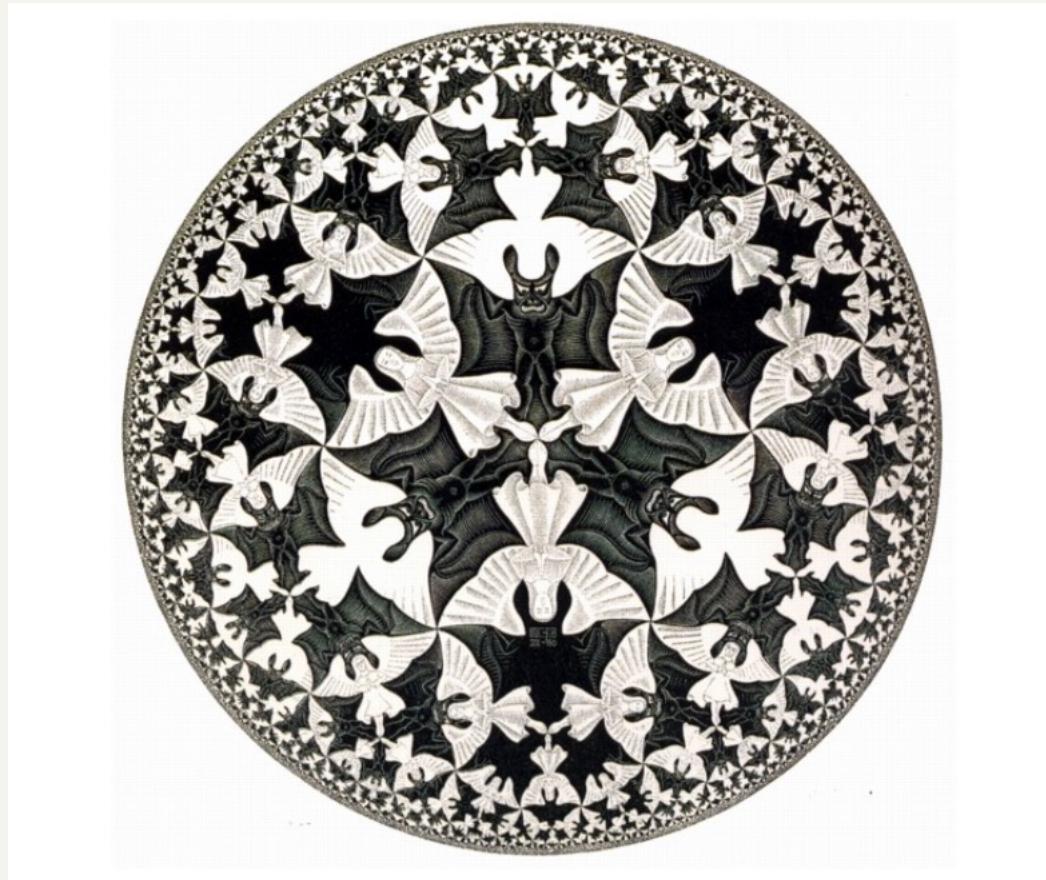
[Maldacena 1997;
Witten; Gubser et al]



$$\Psi_{\text{GRAVITY}} \longleftrightarrow Z_{\text{QFT}}$$

Holography

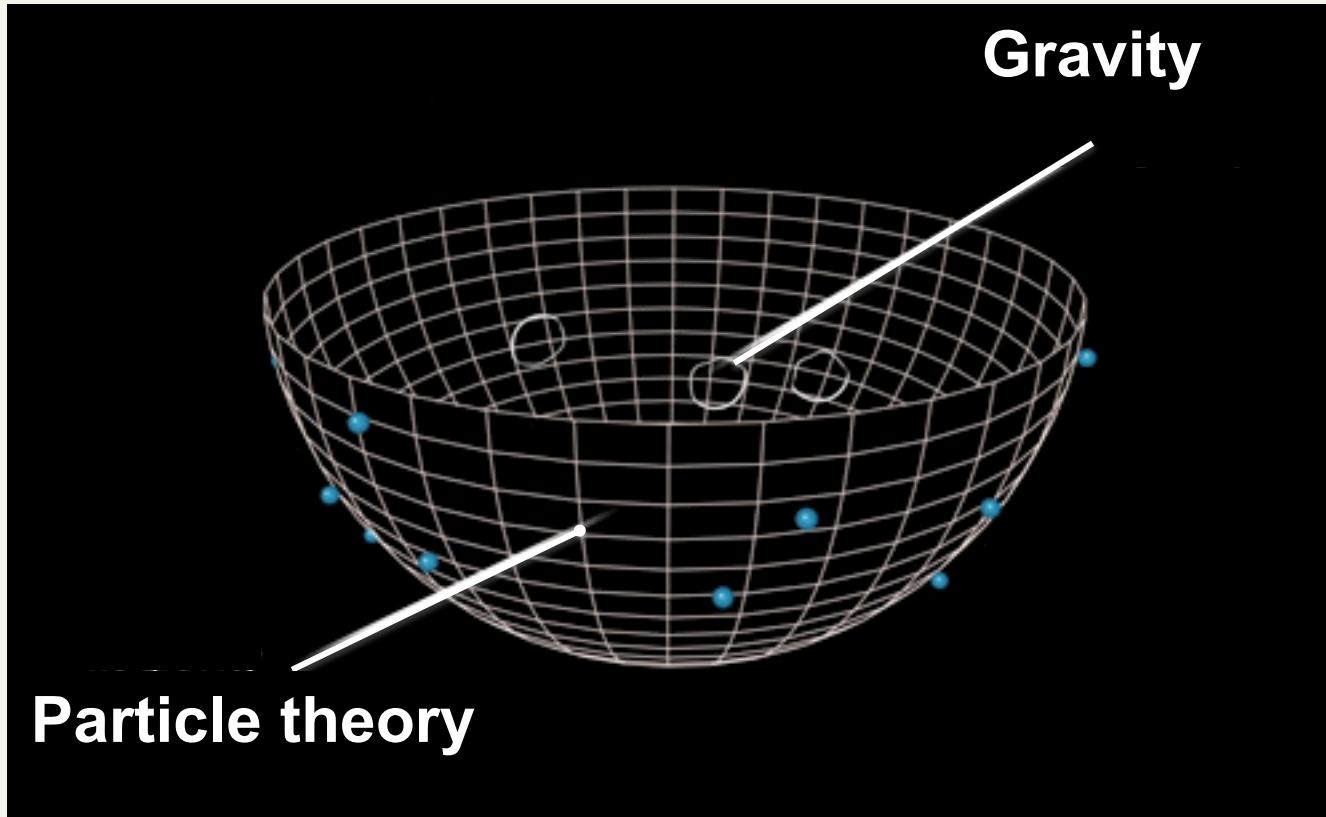
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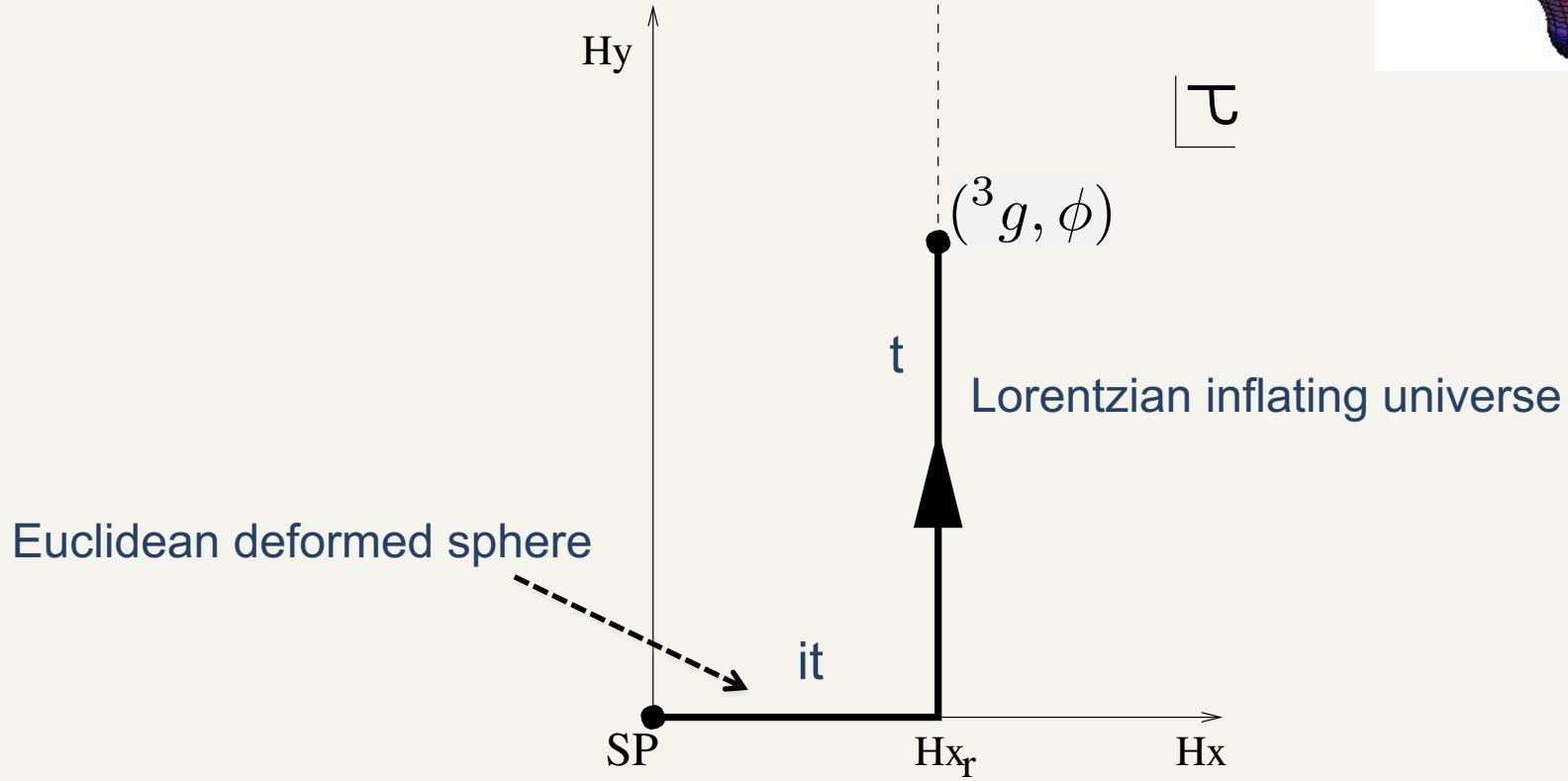
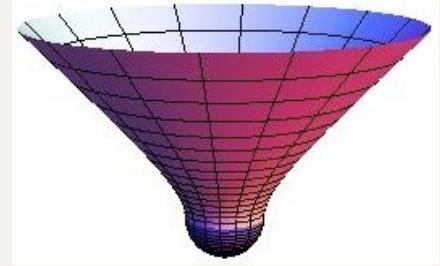
Holography

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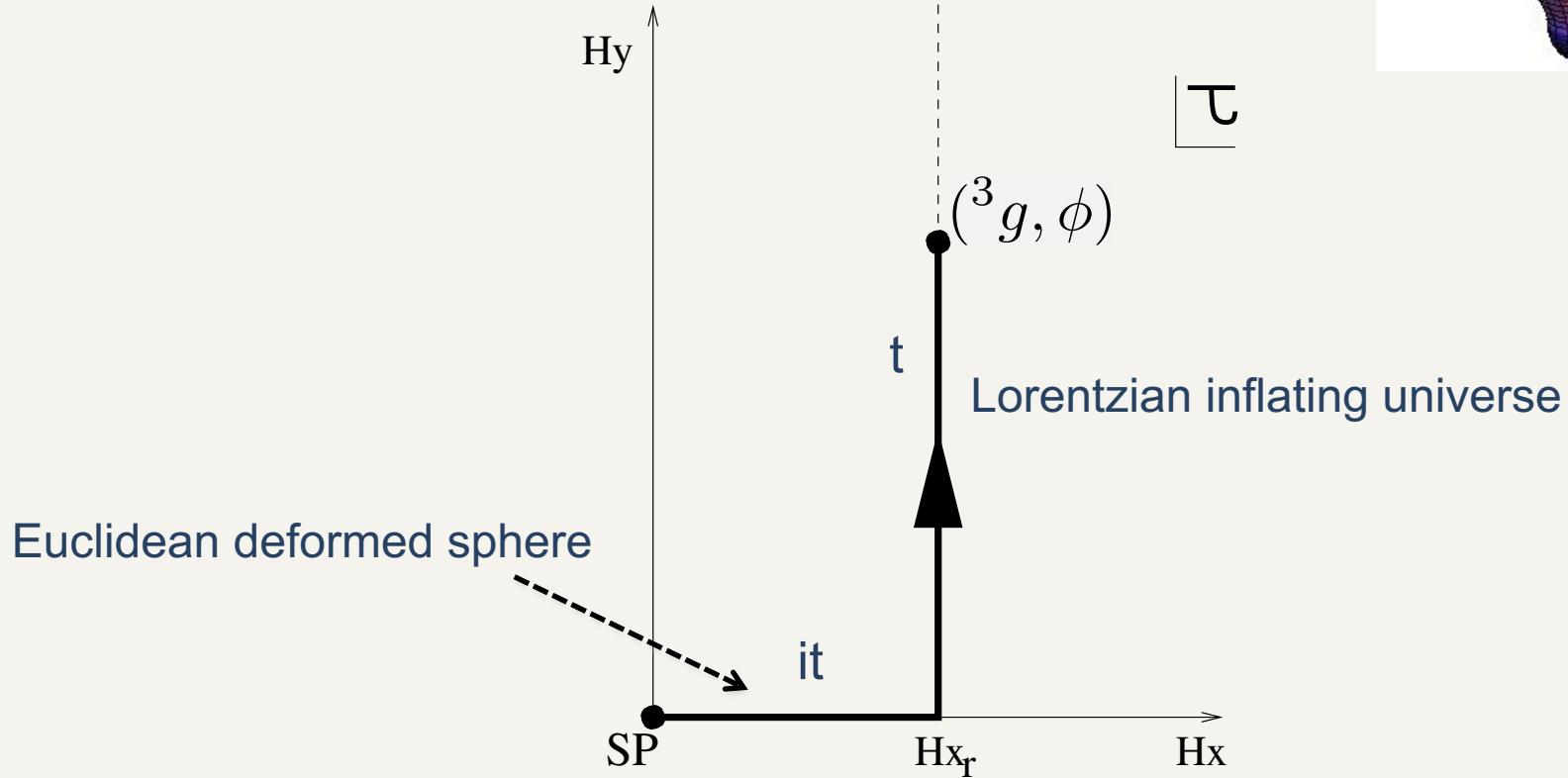
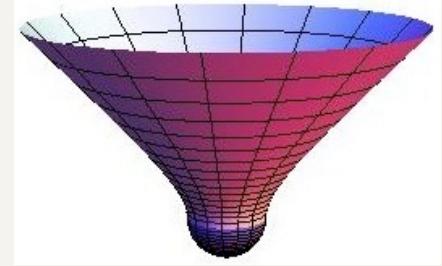
$$\exp(-I_{ADS}^R[h, \chi]/\hbar) = Z_{QFT}[h, \chi]$$

Complex saddle points



$$\Psi_{HH} \approx \exp(-I_E[{}^3g, \phi]) = \exp(-I_R + iS)$$

Complex saddle points

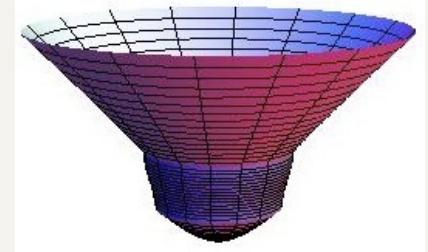
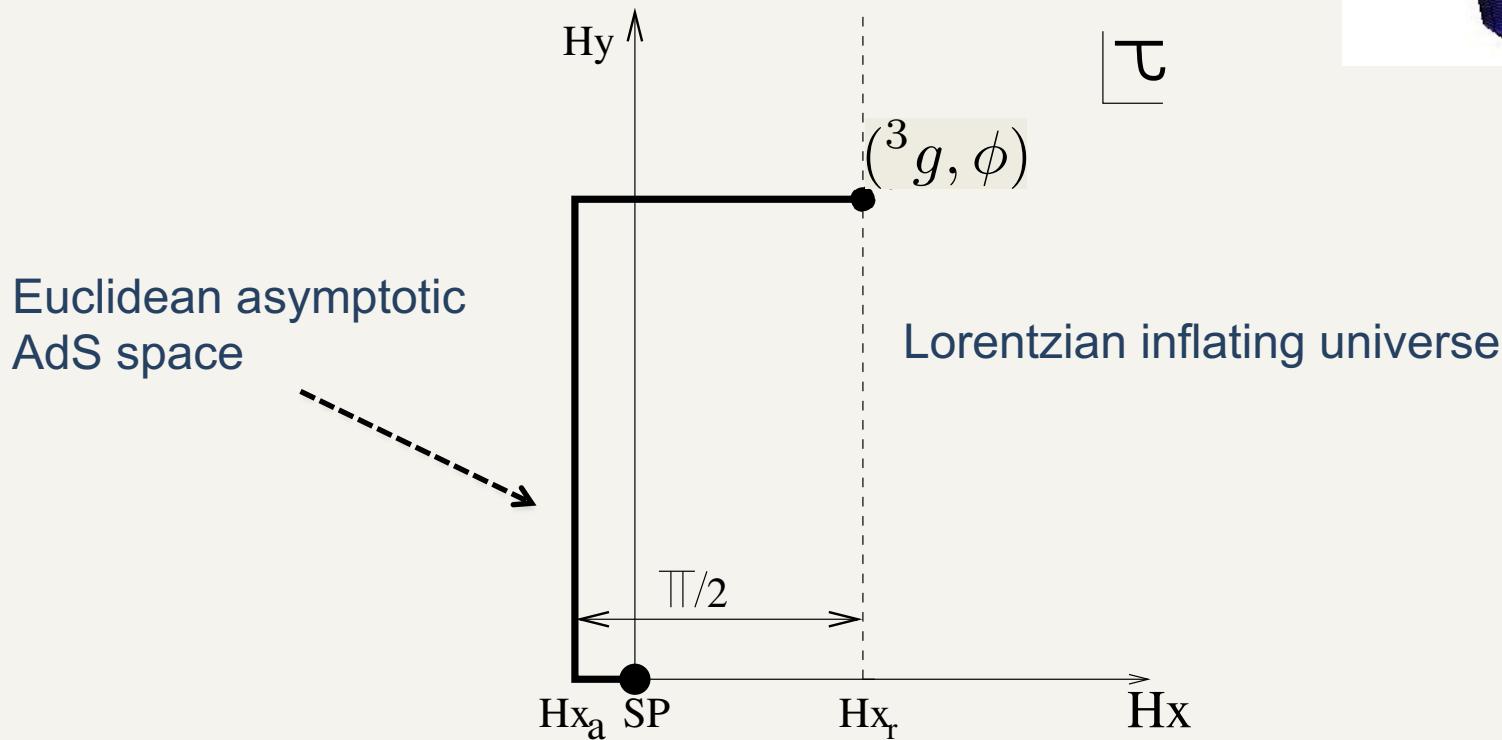


horizontal part: $ds^2 = d\tau^2 + \frac{1}{H^2} \sin^2(H\tau) d\Omega_3^2$

vertical part: $ds^2 = -dy^2 + \frac{1}{H^2} \cosh^2(Hy) d\Omega_3^2$

Alternative representation

[Hartle & TH, 2011]

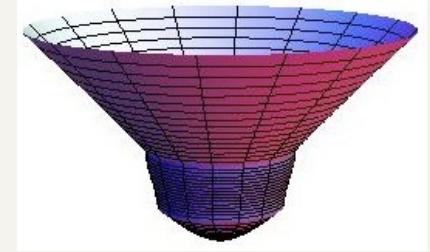
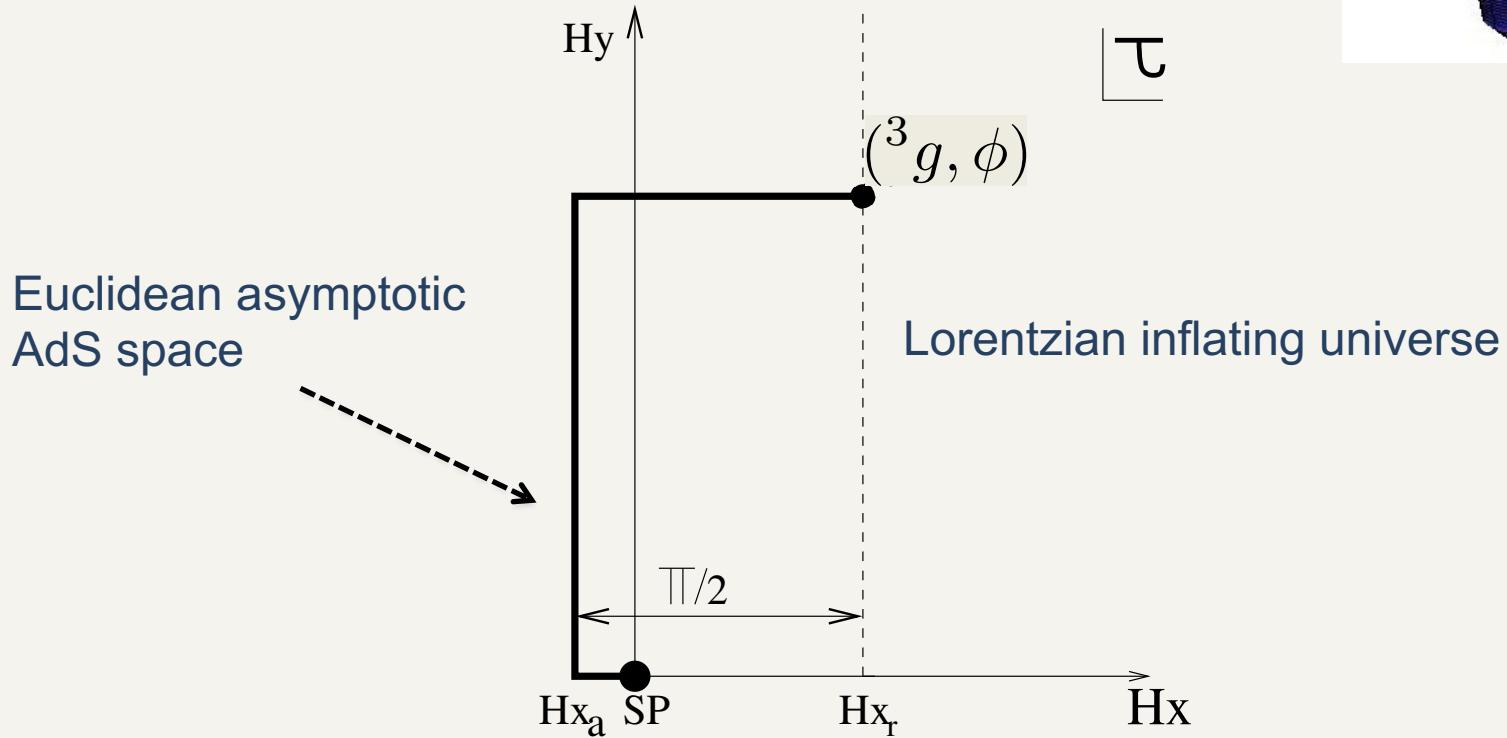


vertical part: Euclidean ADS

$$ds^2 = -dy^2 - \frac{1}{H^2} \sinh^2(Hy) d\Omega_3^2$$

Alternative representation

[Hartle & TH, 2011]



Ψ “connects” Euclidean AdS and Lorentzian de Sitter

Alternative representation

- Action integral along vertical part:

$$I_v = \int_v I[g, \phi] = -I_{AdS}^R({}^3\tilde{g}, \chi) - S_{ct}({}^3g, \chi)$$

where I_{AdS}^R is finite when $a \rightarrow \infty$.

- Action integral along horizontal part:

$$I_h = \int_h I[g, \phi] = +S_{ct}({}^3g, \chi) - iS_{ct}({}^3g, \chi)$$

and no finite contribution.

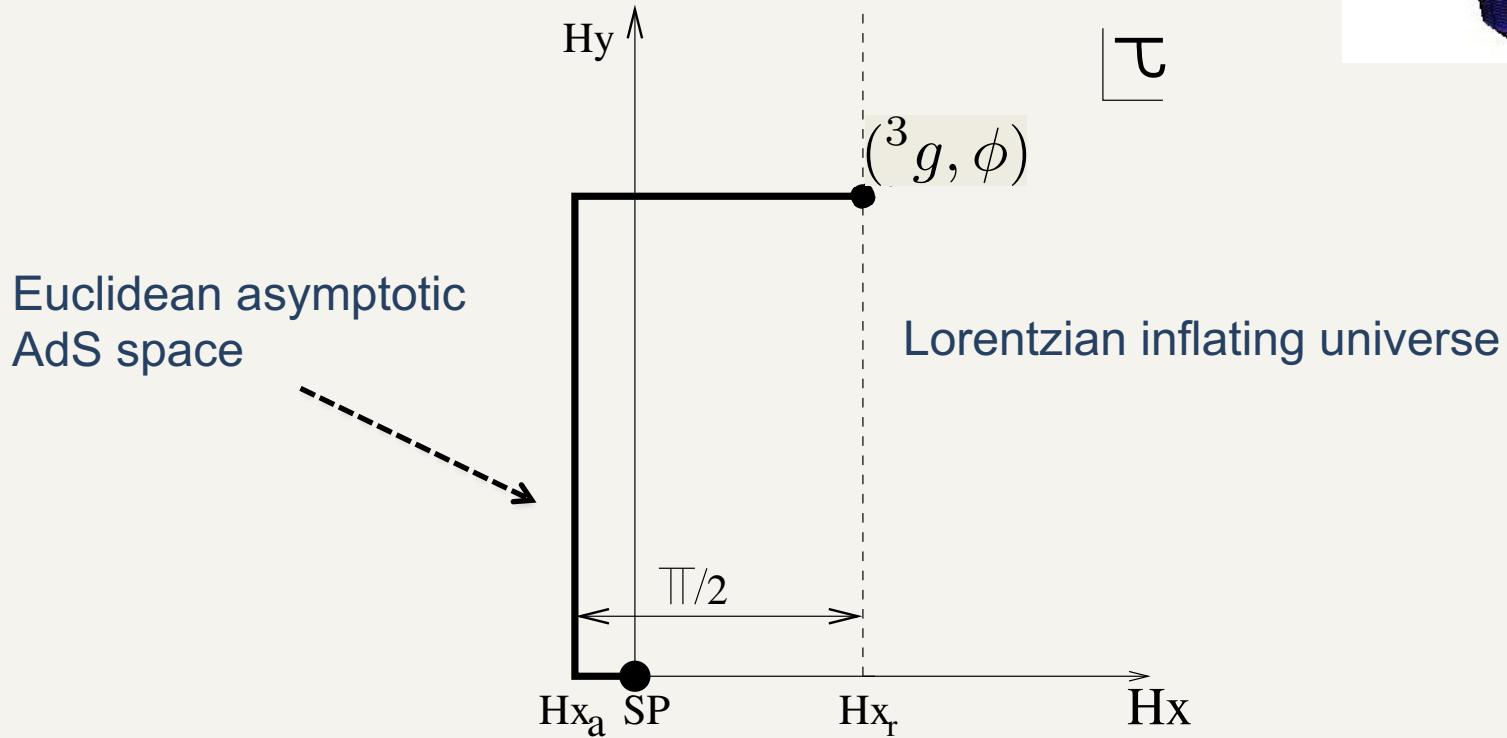
$$S_{ct}(h, \phi) = a_0 \int \sqrt{h} + a_1 \int \sqrt{h} R^{(3)} + a_2 \int \sqrt{h} \phi^2$$

- Total action:

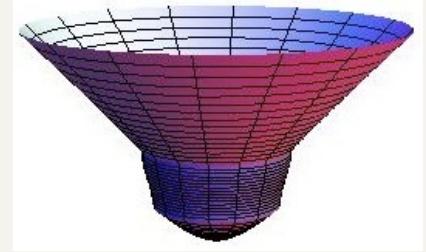
$$I_{ds}({}^3g, \chi) = I_v + I_h = -I_{AdS}^R({}^3\tilde{g}, \chi) - iS_{ct}({}^3g, \chi)$$

Alternative representation

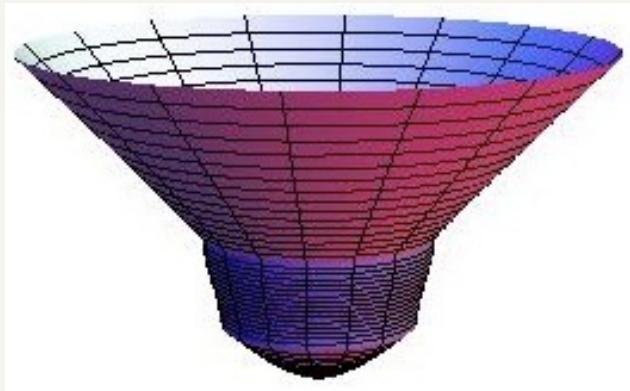
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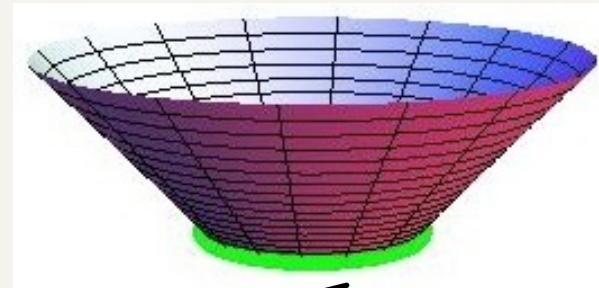
$$\Psi[{}^3g, \phi] = \exp \left(+I_{AdS}^{Reg}/\hbar \right) \exp \left(-iS_{ct}/\hbar \right)$$



Holographic Wave Function



Euclidean
AdS/CFT



AdS/CFT dual QFT
with UV cutoff

$$\Psi_{HH}[{}^3g, \phi] = Z_{QFT}^{-1}[{}^3\tilde{g}, \tilde{\phi}] \exp(-iS_{ct}[{}^3g, \phi])$$

$$Z_{QFT}[{}^3\tilde{g}, \tilde{\phi}] = \langle \exp \int d^3x \sqrt{\tilde{g}} \tilde{\phi} \mathcal{O} \rangle$$