

Lyman-alpha emission reveals an unusual fastly rotating compact dwarf galaxy

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Star-forming Compact Dwarf Galaxies (CDGs) resemble the expected pristine conditions of the first galaxies. Until these early galaxy generations are observationally detected, CDGs are the best systems to test our ideas on primordial galaxy formation and evolution. Here we report on one of such CDGs, Tololo 1214-277, which presents features in its Lyman- α emission that had evaded theoretical interpretation so far. We show that these special features, a symmetric triple peaked emission line, are naturally explained by gas rotation. We find that the Lyman- α emission region in Tololo 1214-277 should have a rotational velocity of $V_r = 300 \text{ km s}^{-1}$ and a neutral Hydrogen column density of $\log N_{HI}/\text{atoms cm}^{-2} = 20.5$. Using archival observational information we find that the diameter for that region should be in the range of $110 \text{ pc} < D < 340 \text{ pc}$ and its total dynamical mass should be between $2.1 \times 10^9 M_\odot < M_D < 6.6 \times 10^9 M_\odot$. This dynamical mass is at least 16 ± 9 times larger than the neutral mass hydrogen. We argue that a possibility to explain the excess in dynamical mass is the presence of a super-massive black hole.

1. [General paragraph about the first galaxies and CDGs]

2. [General paragraph about the expected Lyman-alpha emission, LARS]. Almost fifty years ago [6] it was realized that young galaxies could be detected through a strong Lyman- α line emission. This theoretical prediction was only confirmed XX year later. Currently, Lyman Alpha Emitting galaxies are commonly measure in surveys. The presence of this line provides confirmation of the distance of a galaxy and also provides clues about the stellar population powering the emission.

Lyman alpha emission is not exclusive of high redshift galaxies. It does, require a low dust content to avoid complete extinction and an source of ionizing radiation. This has motivated the creation of local Universe surveys to perform a detailed study of the Lyman-alpha emission [5]. A nearby simple permits the study of other indicators that might be more difficult to obtain for distant galaxies such as morphology, dust attenuation, neutral hydrogen contents and ionization state.

3. [Observations require interpretation, common features and models]. The Lyman- α line is resonant. A Lyman- α photon will follow a diffusion-like process before escaping the galaxy or being absorbed by dust. This makes it hard to interpret Lyman-alpha line observations as the resulting profile is sensitive to the dynamical, chemical and thermal conditions in the interstellar medium. State-of-the-art models use mon

4. [Rotation and the expected features.] It has been shown that rotation also imprints an effect on the Lyman-alpha morphology. The most important consequence of rotation is that spherical symmetry is broken. The line morphology now depends on the viewing angle respect to the rotation axis. For a line of sight perpendicular to the rotation axis the intensity and the line center and the line width increase with rotational velocity. When the rotational velocity is close to the half-line width of the static line, the line becomes single peaked as it is observed in Tololo 1214-277, a unique feature that other theoretical models find impossible to reproduce.

5. [The characteristics of the dwarf galaxy of interest as a puzzling system]

Tololo 1214-277 was first observed by ... it is a compact dwarf galaxy and does not have old stars.

The Ly α emission line was first observed in the Tololo 1214-277 galaxy by [8]. It has two main important features which make this a very uncommon LAE. First it shows a symmetric profile which is , Second the Ly α line is not shifted with respect to the H β line. Blue compact Dwarf Galaxy

6. [The results of the fit + Physical picture using other results]

Figure 1. shows the observational data for Tololo 1214-277 with the overplot from our best fit model from the full radiative transfer simulation. The parameters for the best fit are $v_{max} = 300 \text{ km s}^{-1}$, $\tau = 1 \times 10^7$, $T = 1.5 \times 10^4 \text{ K}$ and a viewing angle $\theta < 30$ degrees.

Figure Observed line + fit.

Assuming spherical symmetry and a homogeneous gas distribution we estimate the total neutral hydrogen mass to be on the order of $M \approx m_H \tau^3 \sigma^{-3} n^{-2}$, where m_H is the mass of a Hydrogen's atom, τ the optical depth, σ is the cross section at the line's center and n is the number density of neutral Hydrogen atoms. For this system we estimate that for average values of $n = 1 \times 10^3$ the total hydrogen mass is $M \times 10^{14} M_{\odot}$. However, blind HI surveys have put an upper limit in the neutral hydrogen mass of ???

7. [Possible interpretation of this results]

8. [Implications for outflow+rotation in existing samples.]

9. [Future outlook]

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$\alpha(2000)^a$	12h17min17.1s
$\delta(2000)^b$	-28d02m32s
l, b (deg)	294, 34
m_V	17.5
$v(\text{km s}^{-1})$	7795

Table 1: Observational characteristics of TOL1214-277 [8]

Tololo 1214-277 characteristics

The receding velocity is $7785 \pm 50 \text{ km s}^{-1}$, which translates into a distance of 106.6 Mpc (Hubble constant $73 \text{ Mpc km}^{-1} \text{ s}^{-1}$) The metallicity is $\sim Z_{\odot}/24$ [3] as derived from optical spectroscopy.

The observed flux for the Lyman alpha line is $\sim 8.1 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$ [8] and a Equivalent Width of 70 \AA and its $\text{H}\beta$ flux is $1.62 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ cm}^{-2}$ [3] which gives a $\text{Ly}\alpha/\text{H}\beta$ flux ratio of 4.9 ± 0.1 . Comparing this ratio with the theoretical expectation from case B recombination of 23.3 [2] one can estimate an escape fraction of 20% for $\text{Ly}\alpha$ radiation.

The optical emission comes from a region with approximate diameter ?? [1].

MCMC constraints

Physical Interpretation

Interpretation by [4].

There is an upper limit for the integrated flux of $< 0.10 \text{ Jy km s}^{-1}$, which translates into a upper limit for the HI mass of $M < 2.65 \times 10^8 M_{\odot}$ [7]. From the optical depth of 10^7 and the non-detection in the HI line, we have an upper limit for the size where the $\text{Ly}\alpha$ emission comes of $D < 0.34 \text{ kpc}$.

For an homogeneous sphere the HI optical depth from its can be written as $\tau = \sigma_0 n D / 2$, where $\sigma_0 = 5.898 \times 10^{-14} \text{ cm}^{-2}$ is the Lyman α optical depth at the line's center, n is the number density and D is the sphere's diameter. From this we can impose additional constrains on D from the typical values of the Hydrogen number density and our constrain on $\tau = 10^7$. Using a range of $1 < n / \text{atoms/cm}^{-3} < 10^{-3}$. This gives us a range of $0.11 < D / \text{kpc} < 100$. Together from the total HI mass we have thus that the HI region should have a diameter of $0.11 < D / \text{kpc} < 0.34$.

This can be rewritten in terms of the gas' temperature T and column density N_H as $\tau = 3.31 \times 10^{-14} (10^4 \text{ K} / T)^{1/2} (N_H / \text{atoms cm}^{-2})$.

This allows us to approximate the total hydrogen mass as

$$M_H = m_H N_H D^2 = 226 \times \tau \left(\frac{T}{10^4 \text{ K}} \right)^{1/2} \left(\frac{D}{\text{kpc}} \right)^2 M_{\odot} \quad (1)$$

On the other hand, we have an estimate for the dynamical mass from the galaxy size D and its rotational velocity V :

$$M_T = \frac{V^2 D}{G} = 2.16 \times 10^5 \left(\frac{V}{\text{km s}^{-1}} \right)^2 \left(\frac{D}{\text{kpc}} \right) M_\odot \quad (2)$$

From this limit and the rotational velocity of 300 km/s and the limit in the size D we have limits of for the dynamical mass of $2.1 \times 10^9 M_\odot < M_D < 6.6 \times 10^9 M_\odot$.

which is at least 7 to 25 times larger than the HI mass.