Gravitational Wave Bursts: Inferences Without (Many) Assumptions

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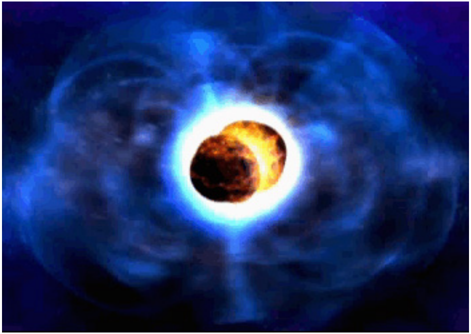


Outline

Current / past work

Highlights:

- Electromagnetically-triggered searches for transient gravitational waves (GWs) in LIGO data. E.g.,
  - Neutron star f-mode oscillations associated with pulsar glitches [1]
  - Searches for un-modelled GW bursts & inspiral signals associated with GRBs [2, 3]
- Prospects for ...
  - ... "joint gravitational wave and short gamma-ray burst observations" [4]
  - ... "high frequency burst searches following binary neutron star coalescence with advanced gravitational wave detectors" [5]

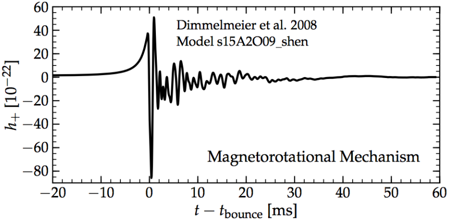
Generally, violent events in neutron stars, Bayesian inference & machine learning $\leftrightarrow$ astrophysical inference from gravitational wave bursts (GWBs)

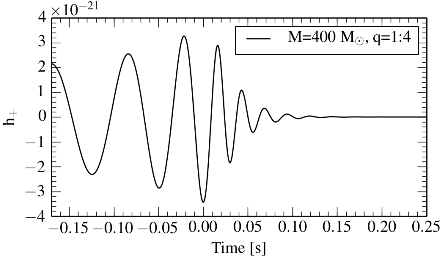
GW Sources & Ground-based Detectors

- . . .

Gravitational Wave "Bursts"

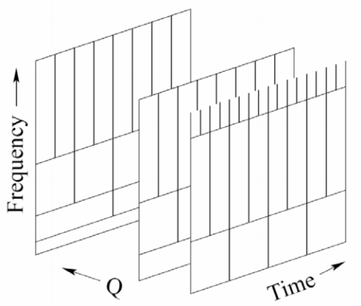
- Transient: duration $\sim \mathcal{O}(10^{-3} - 100)\,\mathrm{s}$
- Uncertain morphology: matched-filtering ineffective or biased
- More certain morphology but few waveform cycles: 'excess power' searches competitive with matched-filtering





GWB Analysis: Excess Power

- Time-frequency decomposition of detector time-series data
- 'Excess power' due to e.g., GWs manifests as hot pixels in time-frequency plane

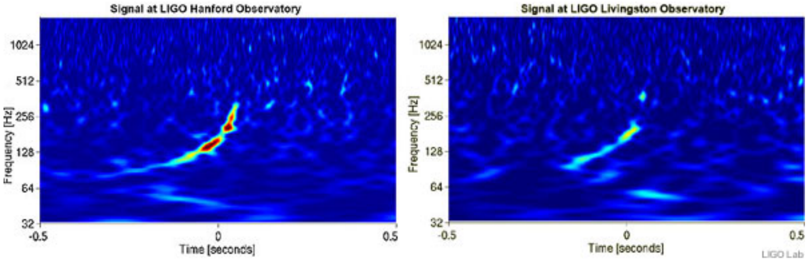


Time-Frequency Analysis

Typically decompose data at multiple resolutions via e.g., wavelets , Q-transforms, STFTs

GWB Analysis: Excess Power

Example: binary neutron star inspiral signal observed in Hanford and Livingston detectors (blind hardware injection during final science run)



GWBs: Coherent Analysis

- **Problem:** GW detector data: highly-nonstationary, *full* of excess power
- **Solution:** search for *coherent* power a detector network

Single-pixel network analysis (following e.g., [6]):

$$\underbrace{\begin{bmatrix} d_1 \\ d_2 \\ \vdots \\ d_D \end{bmatrix}}_{\text{measured data}} = \underbrace{\begin{bmatrix} F_1^{+}(\boldsymbol{\Omega})/\sigma_1 & F_1^{\times}(\boldsymbol{\Omega})/\sigma_1 \\ F_2^{+}(\boldsymbol{\Omega})/\sigma_2 & F_2^{\times}(\boldsymbol{\Omega})/\sigma_2 \\ \vdots & \vdots \\ F_D^{+}(\boldsymbol{\Omega})/\sigma_D & F_D^{\times}(\boldsymbol{\Omega})/\sigma_D \end{bmatrix}}_{\text{whitened antenna responses}} \underbrace{\begin{bmatrix} h_+ \\ h_\times \end{bmatrix}}_{\text{GW polarizations}} + \underbrace{\begin{bmatrix} n_1 \\ n_2 \\ \vdots \\ n_D \end{bmatrix}}_{\text{noise}} \tag{1}$$

i.e.,

$$\begin{aligned} \mathbf{d} &= [\mathbf{F}_+ \mathbf{F}_\times]\,\mathbf{h} + \mathbf{n} \\ &= \mathbf{F}\mathbf{h} + \mathbf{n} \end{aligned} \tag{2}$$

GWBs: Coherent Analysis

Likelihood for signal (h) vs noise (0):

$$\mathcal{L} = 2\log\frac{P(\mathbf{d}|\mathbf{h})}{P(\mathbf{d}|0)} = |\mathbf{d}|^2 - |\mathbf{d} - \mathbf{F}\mathbf{h}|^2 \tag{3}$$

- Treat waveform values $\mathbf{h} = (h_+,\ h_\times)$ in each pixel as free parameters; maximise $\mathcal{L}$ to form *standard likelihood*

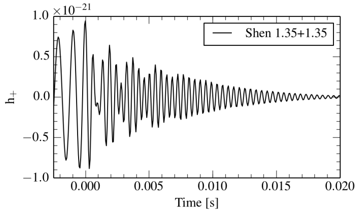
$$E_{\mathrm{SL}} = \mathbf{d}^T\mathbf{F}\mathbf{F}_{\mathrm{MP}}^{-1}\mathbf{d} \tag{4}$$

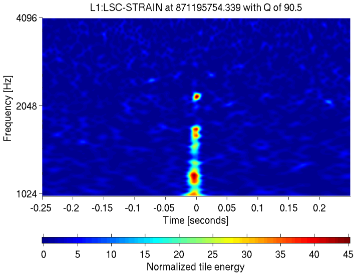
- where $\mathbf{F}_{\mathrm{MP}}^{-1} = (\mathbf{F}^T\mathbf{F})^{-1}\mathbf{F}^T$
- $\mathbf{F}\mathbf{F}_{\mathrm{MP}}^{-1}$ projects data onto $(\mathbf{F}_+, \mathbf{F}_\times)$ sub-space
- Residual data with signal removed should be Gaussian. Can reject non-Gaussian glitches by looking at the null energy:

$$E_{\mathrm{null}} = \mathbf{d}^T(\mathbf{I} - \mathbf{F}\mathbf{F}_{\mathrm{MP}}^{-1})\mathbf{d} \tag{5}$$

Binary Neutron Star Mergers & GWBs

- BNS *inspiral* detectable to 200 Mpc in aLIGO[1]
- Post-merger scenarios:
  1. prompt-collapse
  2. quasi-stable NS
  3. stable NS
- 2, 3 favored by many simulations, equations of state
- GW morphology poorly known, broadband power in $\sim 1-4\,\mathrm{kHz}$ + dominant oscillation frequency
- Detectable to few–10's Mpc

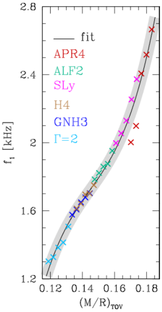




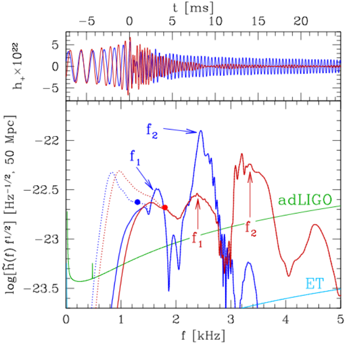
[1] at design sensitivity $\sim$ 2019

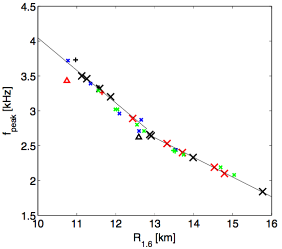
Post-merger Oscillations & NS EOS

- Bauswein et al [7]: dominant post-merger oscillation frequency (f_{peak}) correlates with fiducial NS radius
- Takami et al [8]: similar findings + possible correlation of sub-dominant freq. with NS compactness
- Bauswein et al [9]: constrain maximimum NS mass with f_{peak}



Takami et al:
$f_1 - M/R$ correlation





Bauswein et al: $f_{\mathrm{peak}} - R_{1.6}$ correlation

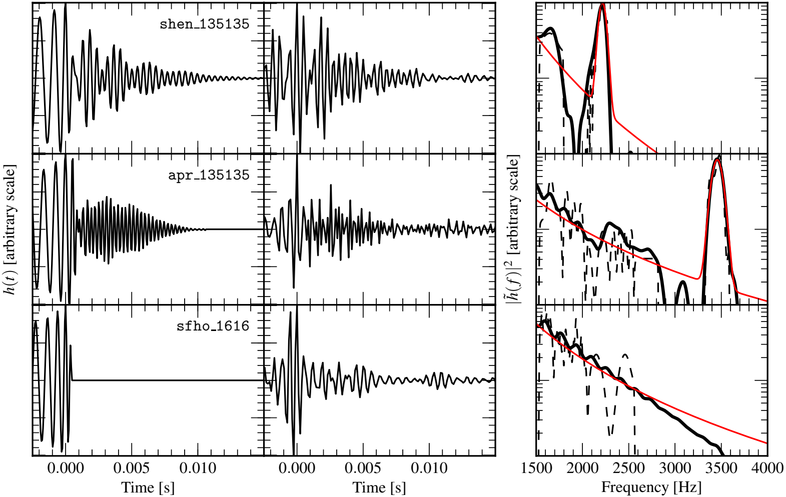
Measuring Post-merger Oscillations In GWs

Clark et al [5]: developed algorithm to analyze post-merger signatures detected (and reconstructed) by burst searches.
Basic idea:

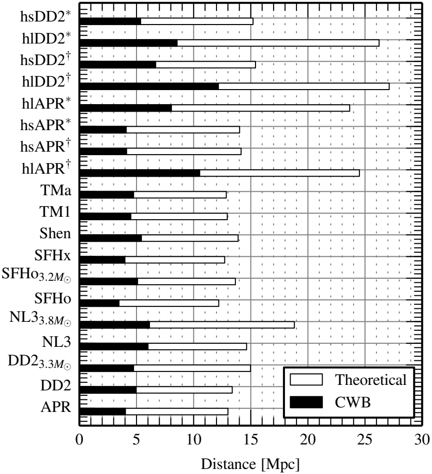
- Search for statistically significant, high-frequenccy power at time around BNS coalescence
- Detection $\rightarrow$ reconstructs signal $\hat{\mathbf{h}}$
- Spectral analysis of $\hat{\mathbf{h}}$ allows classification of prompt/delayed collapse, measurement of f_{peak}

Allows to detect signal and extract astrophysically relevant information without precise waveform models, just generic features (f_{peak})

Measuring Post-merger Oscillations In GWs



Measuring Post-merger Oscillations In GWs



- Compared theoretical matched-filter & *realistic* burst search detectability for aLIGO network for different EoSs
- Translate f_{peak} measurement accuracy to inference on NS radius
- **recover radius to $\mathcal{O}(100\,m)$ within a sphere of 5 Mpc** (1/century – 1/1000 year) rate
- So..not likely, but extraordinary measurements possible with 3$^{\mathrm{rd}}$ generation detectors!

Summary

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[8] K. Takami, L. Rezzolla, and L. Baiotti.
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