

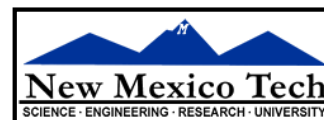
# Advanced Calibration Topics - II

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# Effect of Atmosphere on Phase

# Mean Effect of Atmosphere on Phase

- Since the refractive index ( $n$ ) of the atmosphere  $\neq 1$ , an electromagnetic wave propagating through it will have a phase change (i.e. Snell's law)
- The phase change is related to the refractive index of the air,  $n$ , and the distance traveled,  $D$ , by

$$\phi_e = (2\pi/\lambda) n D$$

For water vapor  $n \propto \frac{W}{DT_{\text{atm}}}$

$W$  = precipitable water vapor (PWV) column

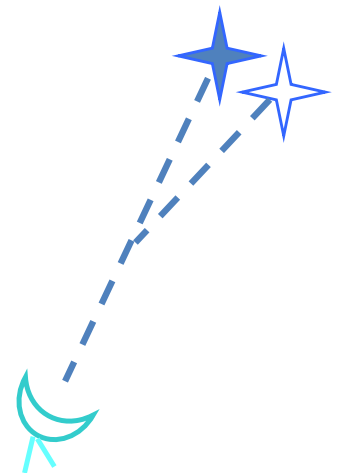
$T_{\text{atm}}$  = Temperature of atmosphere

so  $\phi_e \approx \frac{12.6\pi W}{\lambda}$  for  $T_{\text{atm}} = 270 \text{ K}$

## Refraction causes:

- Pointing off-sets,  $\Delta\theta \approx 2.5 \times 10^{-4} \times \tan(i)$  (radians)  
@ elevation  $45^\circ$  typical offset  $\sim 1'$
- Delay (time of arrival) off-sets

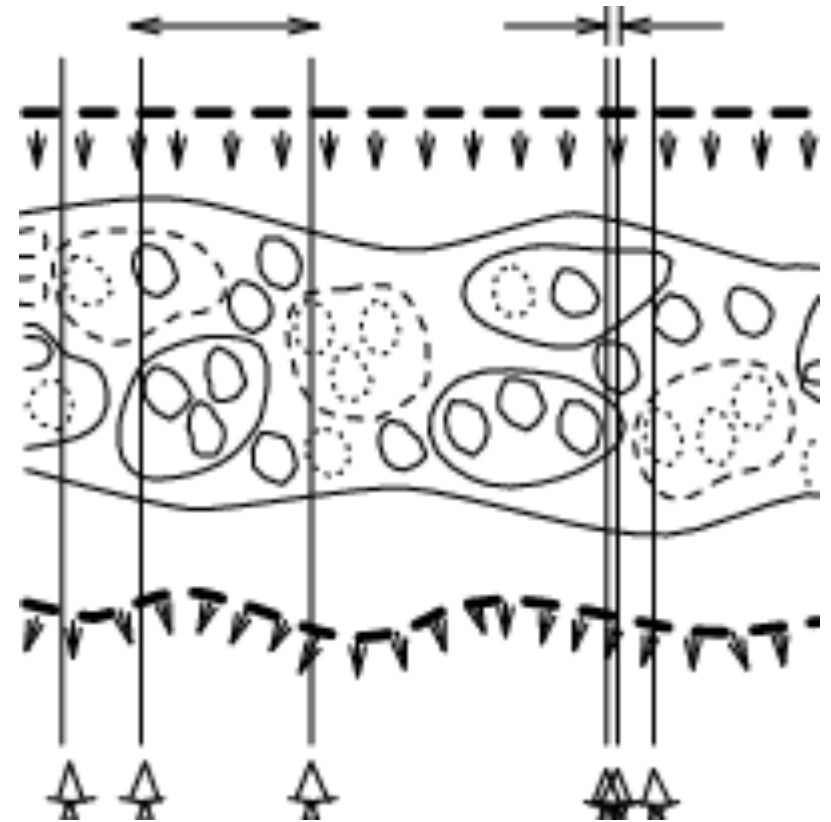
$\Rightarrow$  These “mean” errors are generally removed by the online system



# Atmospheric phase fluctuations

- Variations in the amount of precipitable water vapor (PWV) cause phase fluctuations, which are worse at shorter wavelengths (higher frequencies), and result in:
  - Loss of coherence (loss of S/N)
  - Radio “seeing”, typically 0.1-1" at 1 mm
  - Anomalous pointing offsets
  - Anomalous delay offsets

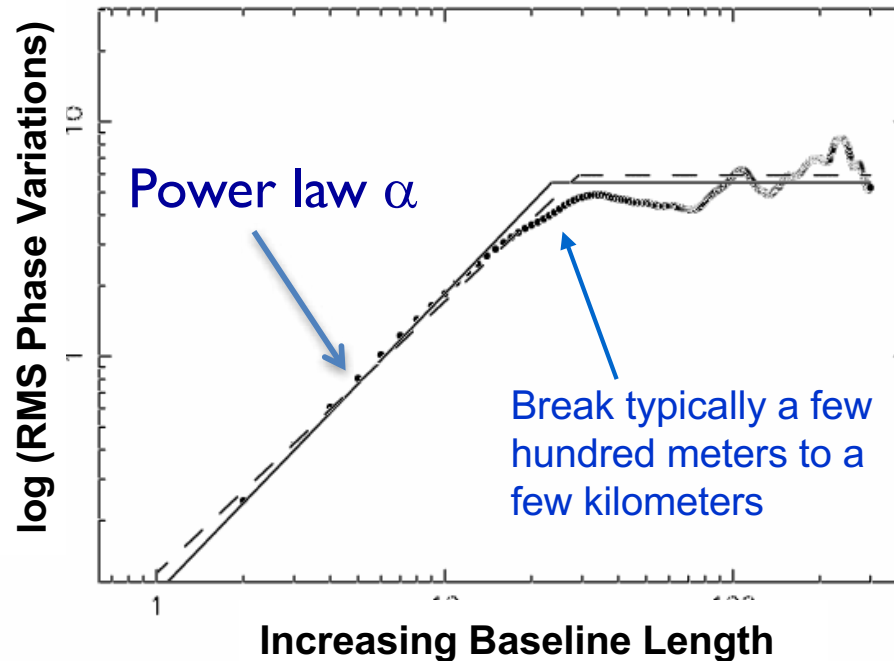
You can observe in apparently excellent submm weather (in terms of transparency, i.e. low PWV) and still have terrible “seeing” i.e. phase stability.



Patches of air with different water vapor content (and hence index of refraction) affect the incoming wave front differently.

# Atmospheric phase fluctuations, continued...

Phase noise as function of baseline length



- “Root phase structure function” (Butler & Desai 1999)
- RMS phase fluctuations grow as a function of increasing baseline length until break when baseline length  $\approx$  thickness of turbulent layer
- The position of the break and the maximum phase variation are weather and wavelength dependent

RMS phase of fluctuations given by Kolmogorov turbulence theory

$$\phi_{\text{rms}} = K b^{\alpha} / \lambda \text{ [deg]}$$

$b$  = baseline length (km)

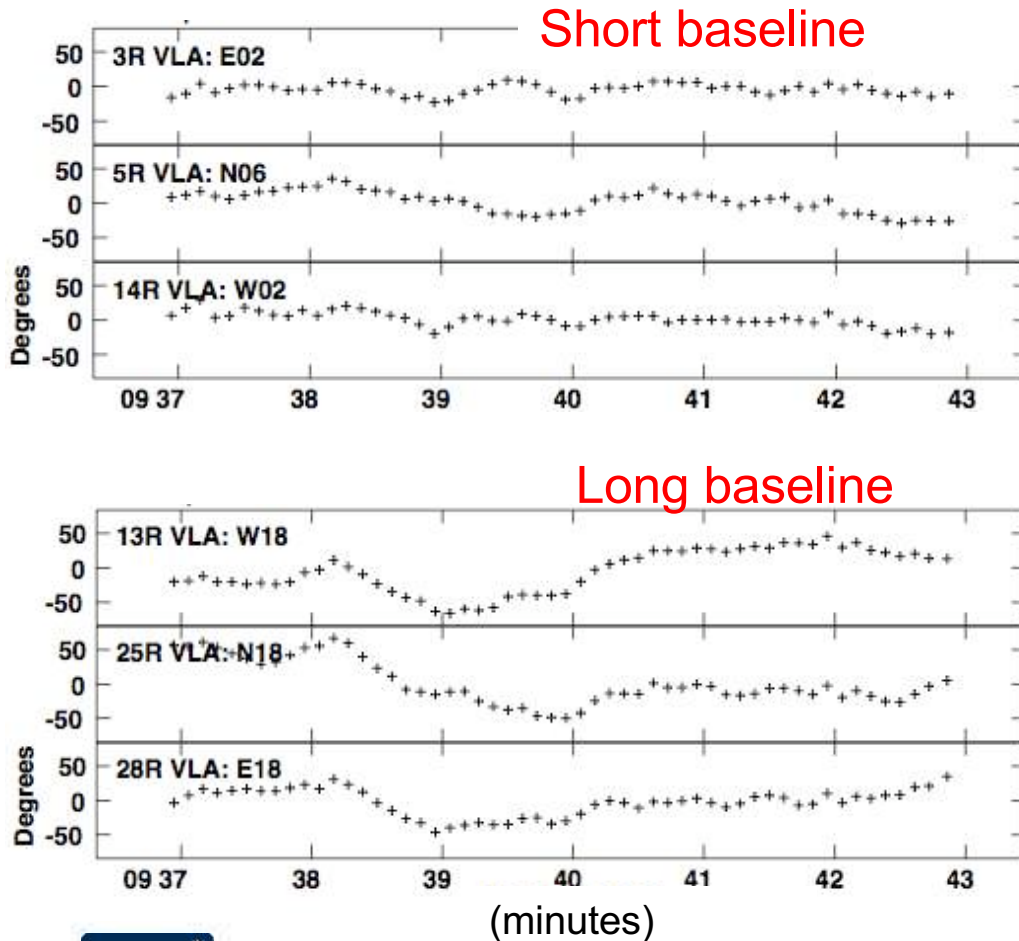
$\alpha$  = 1/3 to 5/6 (thin atmosphere vs. thick atmosphere)

$\lambda$  = wavelength (mm)

$K$  = constant ( $\sim 100$  for ALMA, 300 for JVLAA)

# Residual Phase and Decorrelation

Q-band (7mm) VLA C-config. data from “good” day  
An average phase has been removed from  
absolute flux calibrator 3C286



Coherence = (vector average/true  
visibility amplitude) =  $\langle V \rangle / V_0$

$$\text{Where, } V = V_0 e^{i\phi}$$

The effect of phase noise,  $\phi_{\text{rms}}$ , on  
the measured visibility amplitude :

$$\langle V \rangle = V_0 \langle e^{i\phi} \rangle = V_0 e^{-\phi_{\text{rms}}^2/2}$$

(Gaussian phase fluctuations)

Example: if  $\phi_{\text{rms}} = 1$  radian ( $\sim 60$  deg),  
coherence =  $\langle V \rangle = 0.60V_0$

For these data, the residual  
rms phase (5-20 degrees)  
from applying an average  
phase solution produces a  
7% error in the flux scale

⇒ Residual phase on long baselines have  
larger excursions, than short baselines

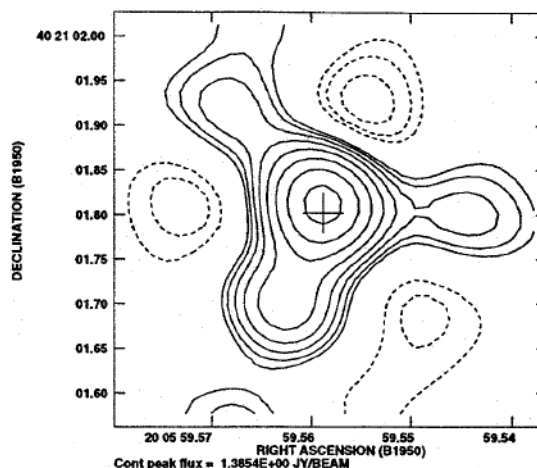
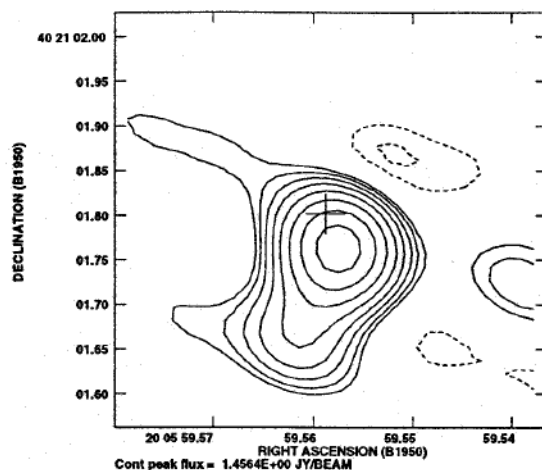


# 22 GHz VLA observations of the calibrator 2007+404

resolution of  $0.1''$  (Max baseline 30 km)

one-minute snapshots at  $t = 0$  and  $t = 59$  minutes

Position offsets  
due to large scale  
structures that are  
correlated  $\Rightarrow$   
phase gradient  
across array



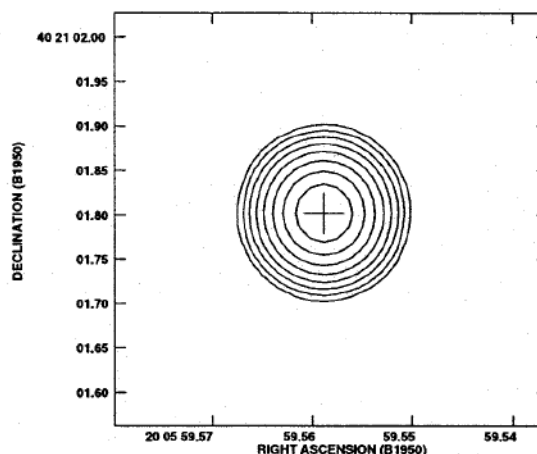
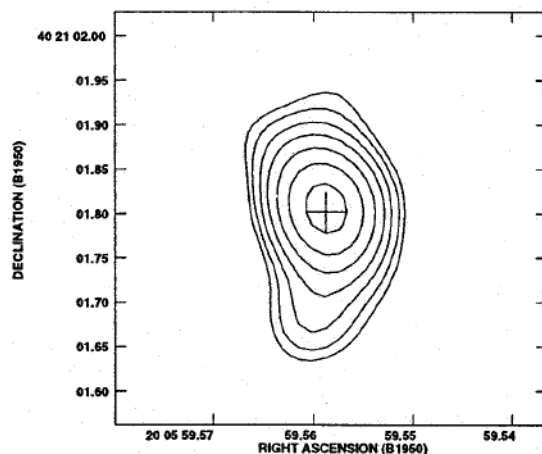
Sidelobe pattern  
shows signature  
of antenna  
based phase  
errors  $\Rightarrow$  small  
scale variations  
that are  
uncorrelated

Corrections 30min:

Corrections 30sec:

All data:

Reduction in peak  
flux (decorrelation)  
and smearing due to  
phase fluctuations  
over 30 min



No sign of phase  
fluctuations  
errors with  
correction  
timescale  $\sim 30$  s



- $\Rightarrow$  Uncorrelated phase variations degrades and decorrelates image
- $\Rightarrow$  Correlated phase offsets = position shift

# Phase Correction Techniques



# Phase fluctuation correction methods

- **Fast switching (Observing strategy)** - used at the JVLA and ALMA for higher frequencies and longer baselines. Choose fast switching cycle time,  $t_{\text{cyc}}$ , short enough to reduce  $\Phi_{\text{rms}}$  to an acceptable level. Calibrate in the normal way.
- **Radiometer (Observing Strategy)** - Monitor phase (via path length) with special dedicated receivers. Requires modeling the atmosphere. Used by ALMA.
- **Self-calibration:** Requires adequate antenna-based S/N
- **Phase transfer (Band-2-Band; Observing Strategy):** simultaneously observe low and high frequencies, and transfer scaled phase solutions from low to high frequency. Tricky, requires well characterized system due to differing electronics at the frequencies of interest. Currently being commissioned at ALMA.
- **Paired array calibration (Observing Strategy):** divide array into two separate arrays, one for observing the source, and another for observing a nearby calibrator.
  - Will not remove fluctuations caused by electronic phase noise
  - Can only work for arrays with large numbers of antennas (was used by CARMA)



# Fast Switching (an observing strategy)

Fast switching phase calibration will stop tropospheric phase fluctuations on baselines longer than an effective baseline length of:

$$b_{\text{eff}} = \frac{V_a t_{\text{cyc}}}{2000}$$

$b_{\text{eff}}$ : effective baseline length in km

$V_a$ : velocity of the winds aloft in m/s (~10 m/s at JVL A)

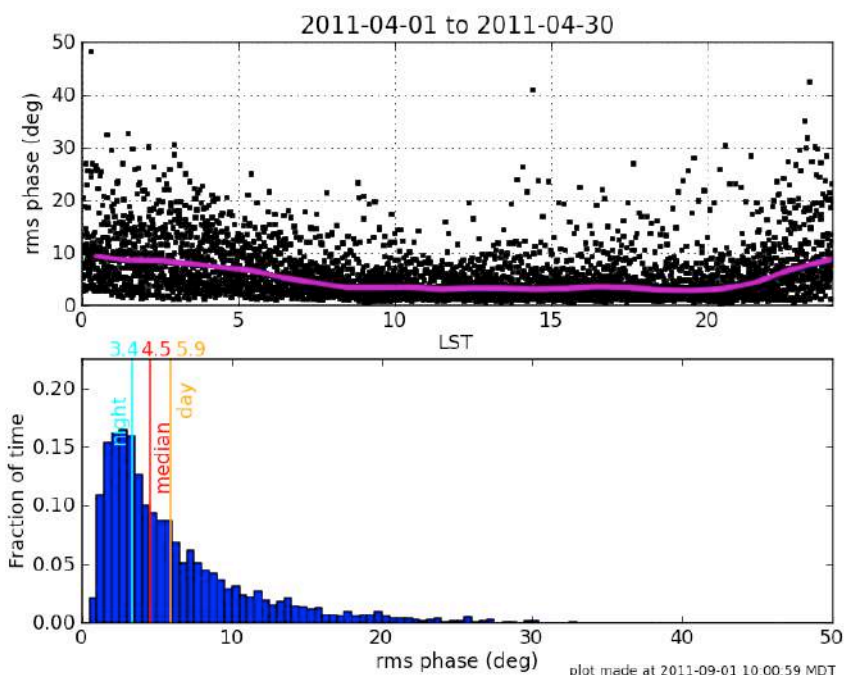
$t_{\text{cyc}}$ : cycle time in seconds (~120 sec)

Cycle times shorter than the baseline crossing time of the troposphere are needed. For example, substituting into the phase rms Eq on slide 5 with  $\alpha = 0.7$  and  $V_a = 10 \text{ m/s}$  (typical for JVL A site) yields:

$$t_{\text{cyc}} (\text{s}) = 200 \left( \frac{\phi_{\text{rms}} (\text{deg}) \lambda (\text{mm})}{K} \right)^{1.42}$$

$K$  = constant (~100 for ALMA, ~300 for VLA)

Note that a 90 degree phase rms will easily wipe out a source.



JVL A Phase monitor:

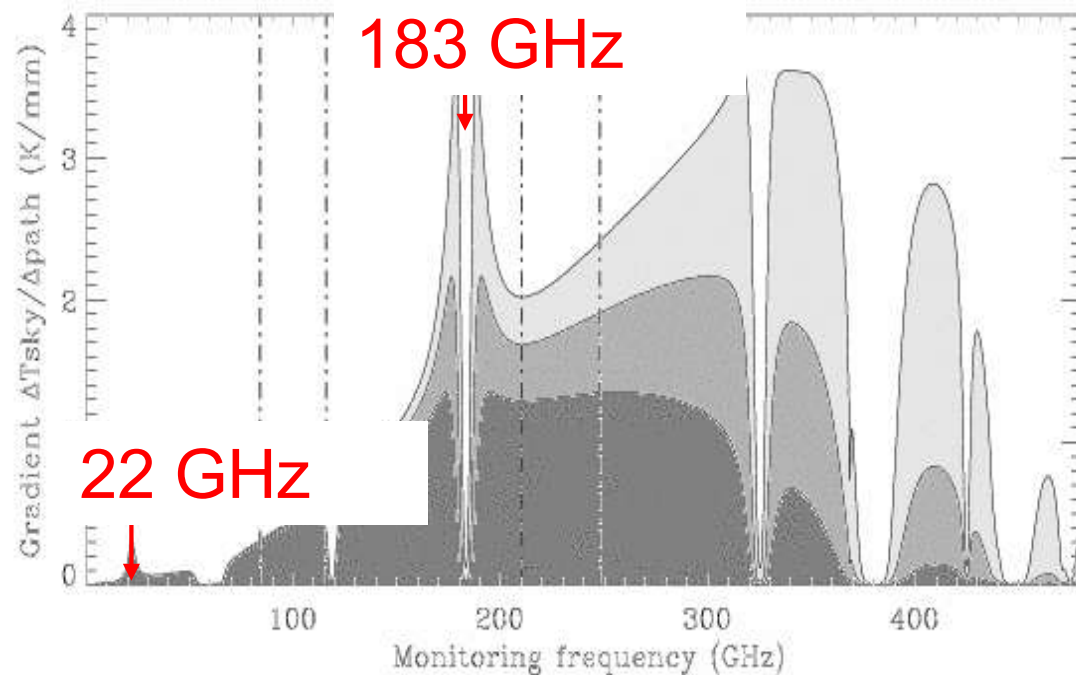
<https://webtest.aoc.nrao.edu/cgi-bin/thunter/apipg.cgi>

# Radiometers (an observing strategy):

- **Radiometry:** measure fluctuations in  $T_B^{\text{atm}}$  with a radiometer, use these to derive changes in water vapor column ( $W$ ) and convert this into a phase correction using

$$\phi_e \approx \frac{12.6\pi W}{\lambda}$$

$W$ =precipitable water  
vapor (PWV) column



(Bremer et al. 1997)

Monitor: 22 GHz H<sub>2</sub>O line (CARMA, JVLA)

183 GHz H<sub>2</sub>O line (CSO-JCMT, SMA)

total power (IRAM)

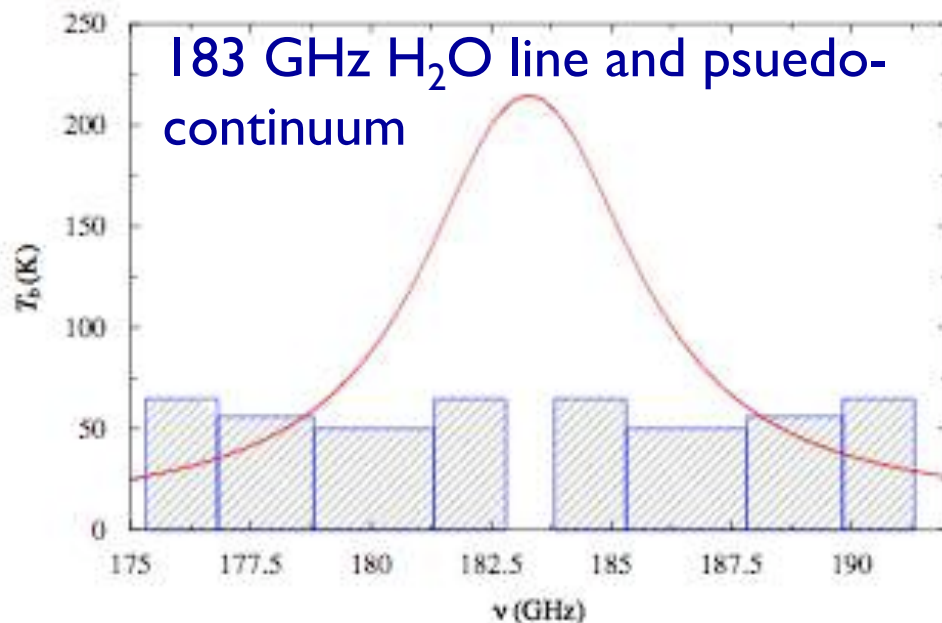
Correct:

183 GHz H<sub>2</sub>O line (ALMA)



# ALMA's particular need for WVR correction:

- Observations at 300 microns (Band 10) require a path error less than 25 microns to keep the phase fluctuations  $< 30$  degrees
- ALMA site testing suggests that the median path fluctuation due to the atmosphere is  $\sim 200$  microns on 300 m baselines (compared to max of 15 km)
- These fluctuations increase with baseline length (up to several km) according to Kolmogorov with a power of about 0.6 for the ALMA site
- Changes on timescales as small as the Antenna diameter/wind speed are possible = 1 sec



ALMA WVRs monitor changes in water line brightness:

There are 4 “channels” flanking the peak of the 183 GHz water line

- Data taken every second
- Installed on all the 12m antennas
- Matching data from opposite sides are averaged
- The four channels allow flexibility for avoiding saturation

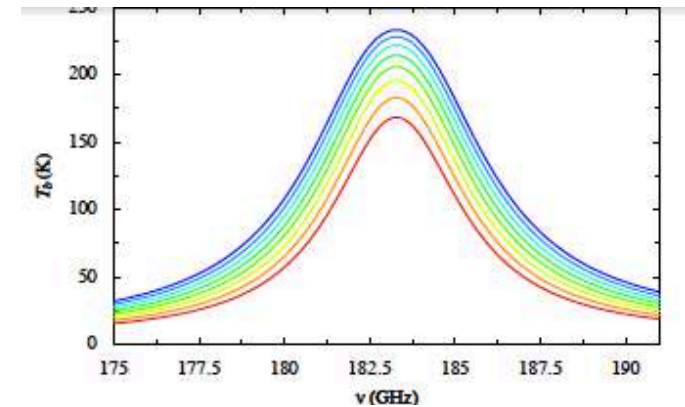
# Modeling the Path Change

Challenge: Convert changes in 183 GHz brightness to changes in path length

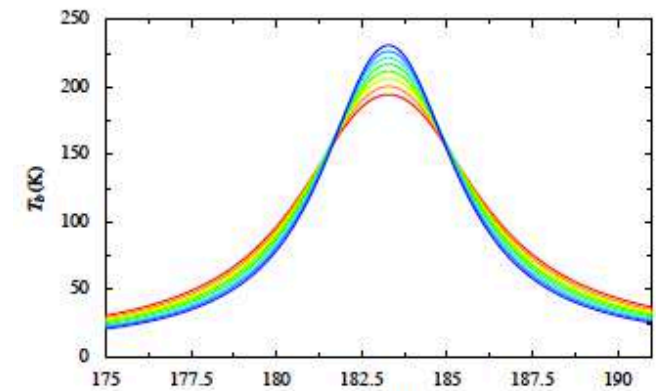
Implementation offline: **wvrgcal**

- 3 unknowns: PWV, temperature, pressure (in water vapor layer) in a simple plane parallel, thin layer model
- HITRAN and radiative transfer is used to derive the line shape, opacity and hence brightness temperature  $T_B(\text{H}_2\text{O})$  as a function of frequency
- The observed “spectrum” is then compared to the model predictions for a range of reasonable values of PWV, Temperature, and pressure
- After dropping smaller terms:  
$$\Delta(\text{path}) = \Delta(\text{PWV}) * 174 \text{ I/T}(\text{H}_2\text{O layer})$$
- The path change is converted to phase for the mean frequency of each “science” spectral window

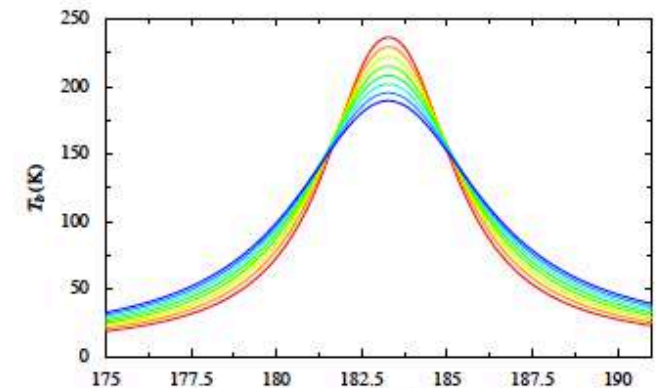
For a more complete description ALMA Memo 587



PWV from 0.6 to 1.3mm



Temperature from 230 to 300 K

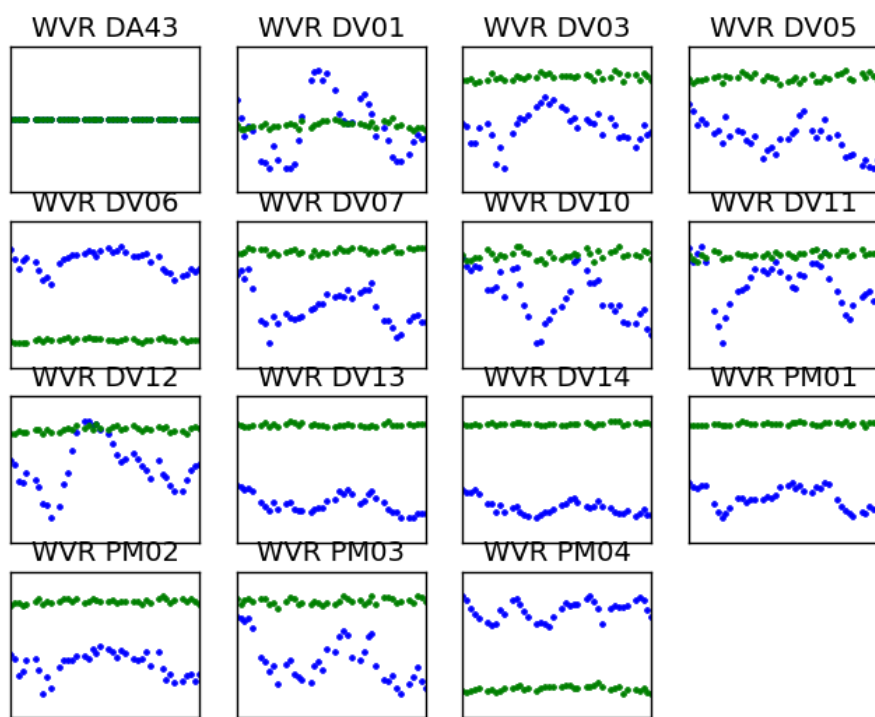


Pressure from 400 to 750 mBar

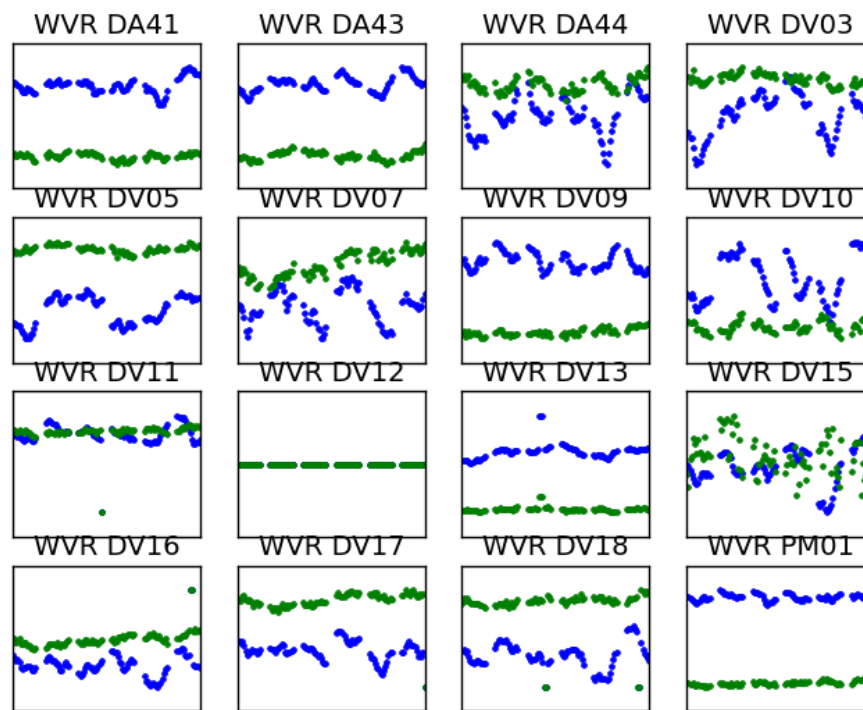


# ALMA WVR Correction - Examples

Band 6 (230 GHz) Compact config



Band 7 (340 GHz) Extended config



Raw phase & WVR corrected phase

# Self-Calibration: Motivation

JVLA and ALMA have impressive sensitivity! But what you achieve is often limited by residual calibration errors

Many objects will have enough Signal-to-Noise (S/N) so they can be used to better calibrate **themselves** to obtain a more accurate image. This is called self-calibration and it really works, if you are careful! Sometimes, the increase in effective sensitivity may be an order of magnitude.

It is not a circular trick to produce the image that you want. It works because the number of baselines is much larger than the number of antennas so that an approximate source image does not stop you from determining a better temporal gain calibration which leads to a better source image.

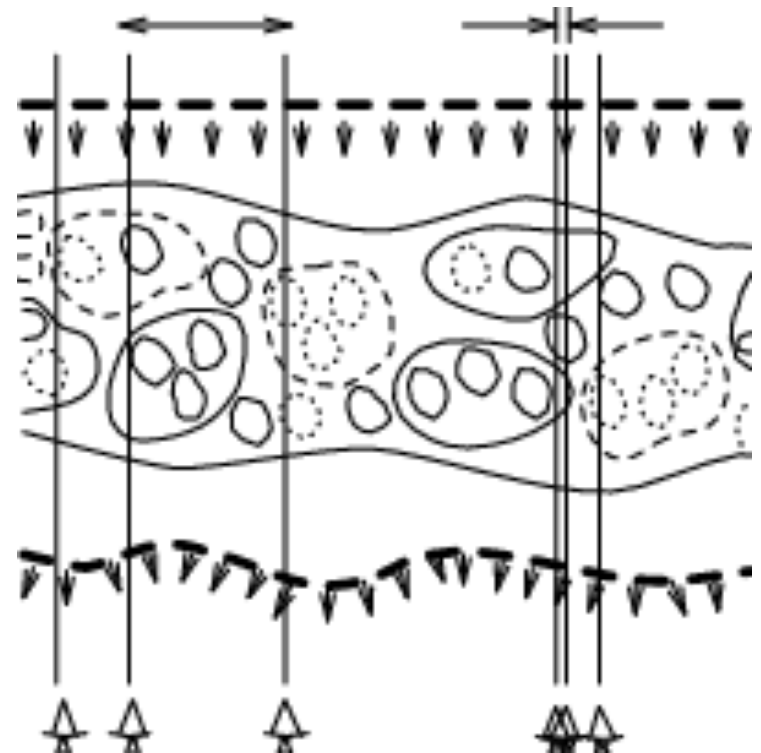


# Data Corruption Types

The true visibility is corrupted by many effects:

- Atmospheric attenuation
- Radio “seeing”
- Variable pointing offsets
- Variable delay offsets
- Electronic gain changes
- Electronic delay changes
- Electronic phase changes
- Radiometer noise
- Correlator mal-functions
- Most Interference signals

| Antenna-based  
| baseline



# Antenna-based Calibration- I

- The most important corruptions are associated with antennas
- Basic Calibration Equation

$$\tilde{V}_{ij}(t) = g_i(t)g_j^*(t)G_{ij}(t)V_{ij}(t) + \varepsilon_{ij}(t) + \epsilon_{ij}(t)$$

$g_i(t)g_j^*(t)$	Factorable (antenna-based) complex gains
$G_{ij}(t)$	Non-factorable complex gains (not Antenna based and typically small)
$V_{ij}(t)$	True Visibility
$\varepsilon_{ij}(t)$	Additive offset (not antenna based and typically small)
$\epsilon_{ij}(t)$	Thermal noise

- Can be reduced to (approximately)

$$\tilde{V}_{ij}(t) = g_i(t)g_j^*(t)V_{ij}(t) + \epsilon_{ij}(t)$$

# Antenna-based Calibration-II

- For N antennas,  $[(N-1)*N]/2$  visibilities are measured, but only N amplitude and (N-1) phase gains fully describe the complete Antenna-based calibration. This redundancy is used for antenna gain calibration
- Basic gain (phase and amplitude) calibration involves observing unresolved (point like) “calibrators” of known position with visibility  $M_{i,j}(t_k, \nu)$
- Determine gain corrections,  $g_i$ , that minimizes  $S_k$  for each time stamp  $t_k$  where

$$S_k = \sum_k \sum_{i \neq j} w_{i,j} |g_i(t_k) g_j^*(t_k) V_{i,j}^o(t_k) - M_{i,j}(t_k)|^2$$

Data  
Weights

Complex  
Gains

Complex  
Visibilities

Fourier transform  
of model image

- The solution interval,  $t_k$ , is the data averaging time used to obtain the values of  $g_i$ , (i.e. solint='int' or 'inf'). The apriori weight of each data point is  $w_{i,j}$ .
- This IS a form of Self-calibration, only we assume a Model ( $M_{ij}$ ) that has constant amplitude and zero phase, i.e. a point source
- The transfer of these solutions to another position on the sky at a different time (i.e. your science target) will be imperfect, but the same redundancy can be used with a **model image** for Self-calibration

# Sensitivities for Self-Calibration-I

- **For phase only self-cal:** Need to detect the target in a solution time ( $\text{solint}_{\text{self}}$ )  $<$  the time for significant phase variations with only the baselines to **a single antenna** with a  $S/N_{\text{self}} \gtrsim 3$ . For 25 antennas,  $S/N_{\text{Self}} > 3$  will lead to  $< 15$  deg error.

- Make an initial image, cleaning it conservatively
  - Measure rms in emission free region of image
  - $\text{rms}_{\text{Ant}} = \text{rms} \times \sqrt{N - 3}$  where  $N$  is # of antennas
  - $\text{rms}_{\text{self}} = \text{rms}_{\text{Ant}} \times \sqrt{\text{Time\_on\_source} / \text{solint}_{\text{self}}}$
  - Measure Peak flux density = Signal
  - If  $S/N_{\text{self}} = \text{Peak} / \text{rms}_{\text{Self}} > 3$  try phase only self-cal

## Rule of thumb:

For an array with  $\sim 25$  antennas, if  $S/N$  in image  $> 20$  its worth trying phase-only self-cal

- **CAVEAT 1:** If dominated by extended emission, estimate what the flux will be on the longer baselines (by plotting the uv-data) instead of the image
  - If the majority of the baselines in the array cannot "see" the majority of emission in the target field (i.e. emission is resolved out) at a  $S/N$  of about 3, the self-cal will fail in extreme cases (though bootstrapping from short to longer baselines is possible, it can be tricky).
- **CAVEAT 2:** If severely dynamic range limited (poor uv-coverage), it can also be helpful to estimate the rms noise from uv-plots

# Sensitivities for Self-Calibration-II

- **For amplitude self-cal:** Need to detect the target with only the baselines to a single antenna with a  $S/N_{\text{self}} \gtrsim 10$ , in a solution time ( $\text{solint}_{\text{self}}$ )  $<$  the time for significant amplitude variations. For 25 antennas, an antenna based  $S/N > 10$  will lead to a 10% amplitude error.
  - Amplitude corrections are more subject to deficiencies in the model image, check results carefully!
  - For example, if clean model is missing significant flux compared to uv-data, give uvrange for amplitude solution that excludes short baselines

## Additional S/N for self-cal can be obtained by:

- Increase solint (solution interval)
  - Errors that are directional, rather than time dependent can yield surprising improvement even if the solint spans the whole observation = antenna position (aka baseline) errors are a good example
- gaintype= 'T' to average polarizations
  - Phase **differences** between polarizations are generally well calibrated
- Combine = 'spw' to average spw's (assumes prior removal of spw to spw offsets)
  - **Caveat:** If source spectral index/morphology changes significantly across the band, do not combine spws, especially for amplitude self-cal unless you use mtmfs
- Combine = 'fields' to average fields in a mosaic (use with caution, only fields with strong signal)

# Self-calibration Example:

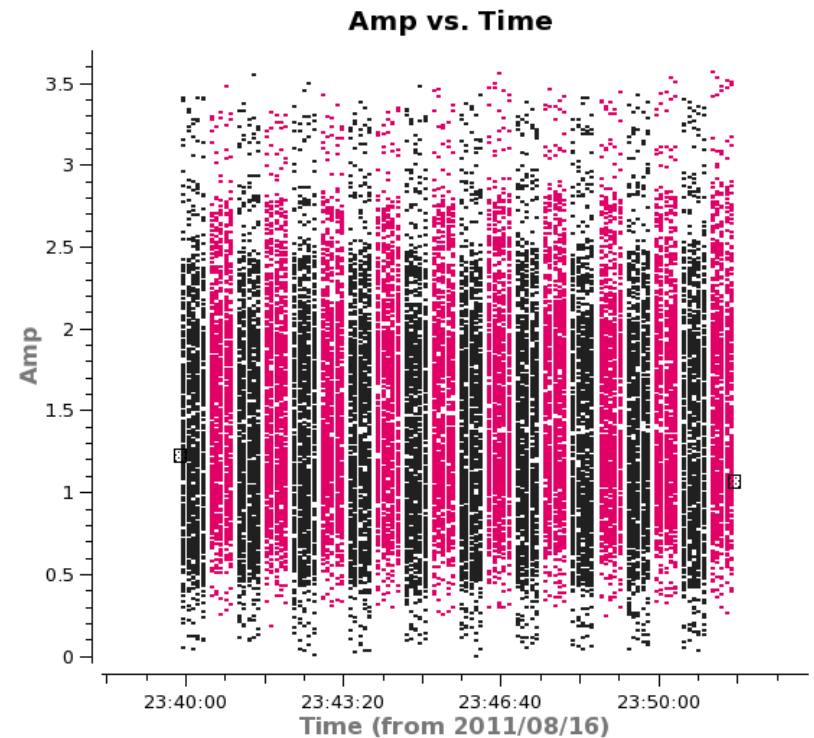
## ALMA SV Data for IRAS16293 Band 6 (1a)

Step 1 – Determine basic setup of data:

- 2 pointing mosaic
- Integration = 6.048 sec; subscans ~ 30sec
- Scan= 11 min 30s (split between two fields)

Step 2 – What is the expected rms noise?

- Use actual final total time and # of antennas on science target(s) from this stage and sensitivity calculator.
- Be sure to include the actual average weather conditions for the observations in question and the bandwidth you plan to make the image from
- 54 min per field with 16 antennas and average  $T_{\text{sys}} \sim 80$  K, 9.67 MHz BW; rms= 1 mJy/beam
- Inner part of mosaic will be about 1.6 x better due to overlap of mosaic pointings



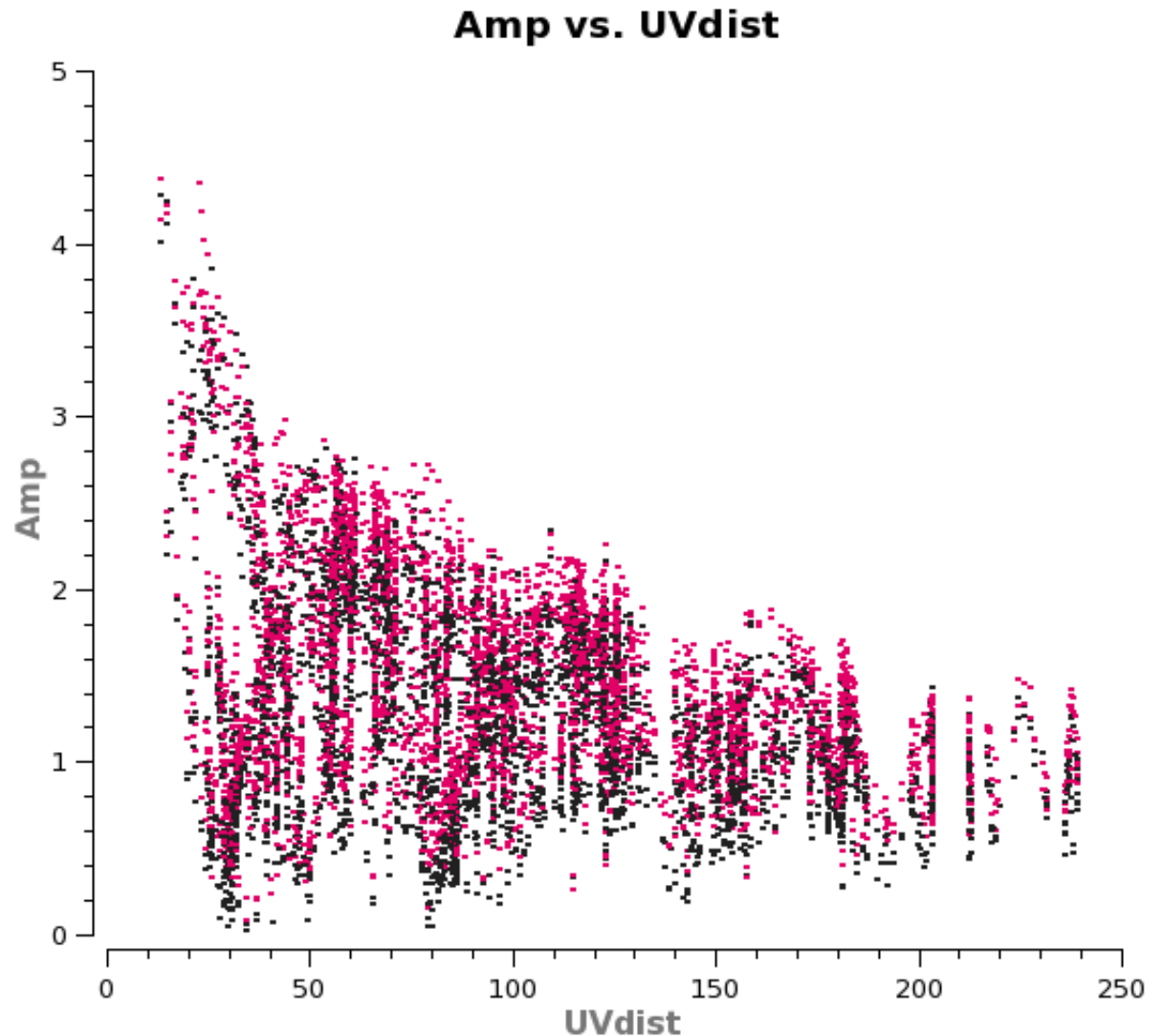
- ALMA mosaic: alternates fields in “subscan” this picture = 1 scan
- EVLA mosaic: alternates fields in scans
- Subscans are transparent to CASA (and AIPS)

# Self-calibration Example:

## ALMA SV Data for IRAS16293 Band 6 (1b)

Step 3 – What does the amplitude vs uv-distance of your source look like?

- Does it have large scale structure? i.e. increasing flux on short baselines.
- What is the flux density on short baselines?
- Keep this 4 Jy peak in mind while cleaning. What is the total cleaned flux you are achieving?



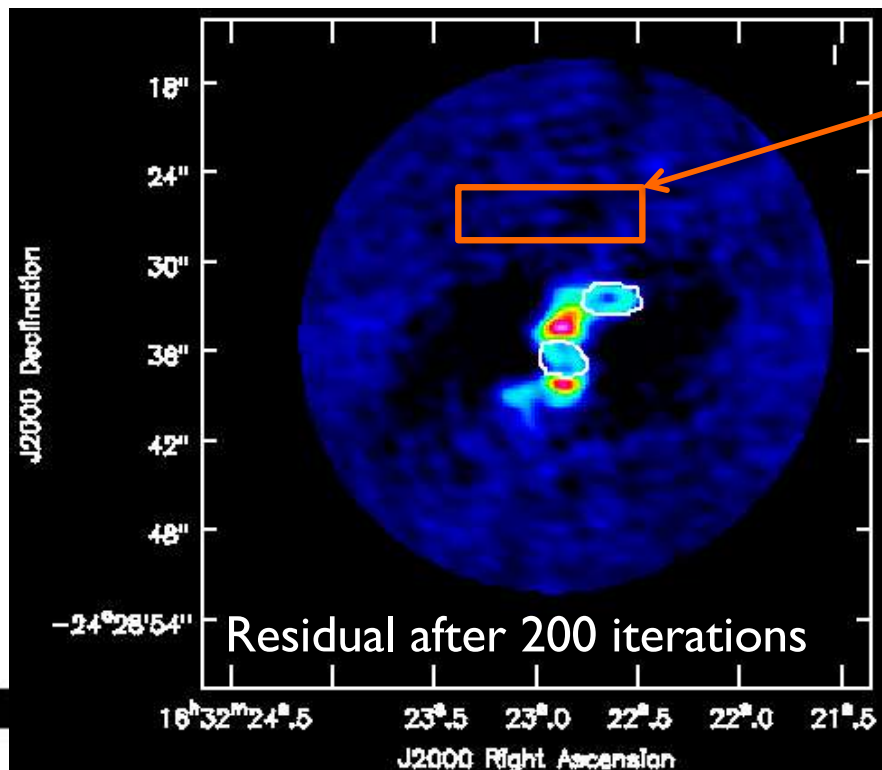
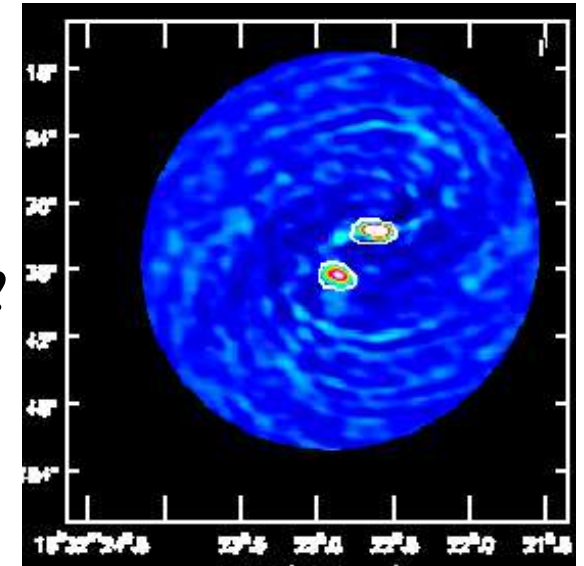


# Self-calibration Example:

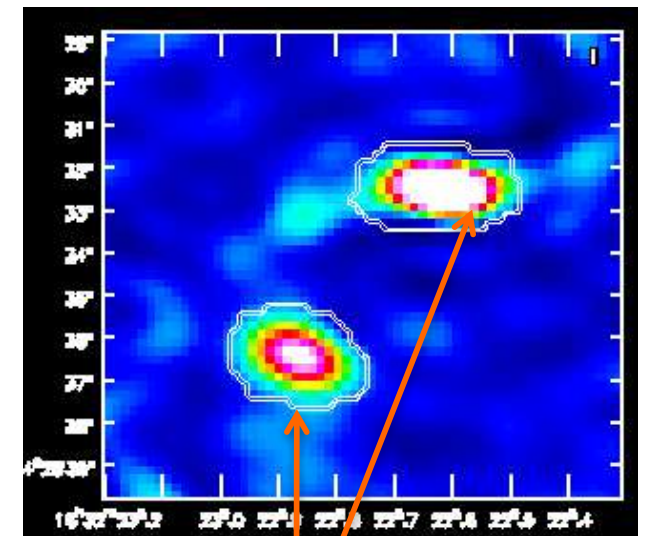
## ALMA SV Data for IRAS16293 Band 6 (II)

Step 4 – What is the S/N in a conservatively cleaned image?

- What is this “conservative” of which you speak
- Rms ~ 15 mJy/beam; Peak ~ 1 Jy/beam  $\rightarrow$  S/N ~ 67
- Rms > expected and S/N > 20  $\rightarrow$  self-cal!



Stop clean,  
and get rms  
and peak  
from image,  
avoiding  
negative  
bowls and  
emission



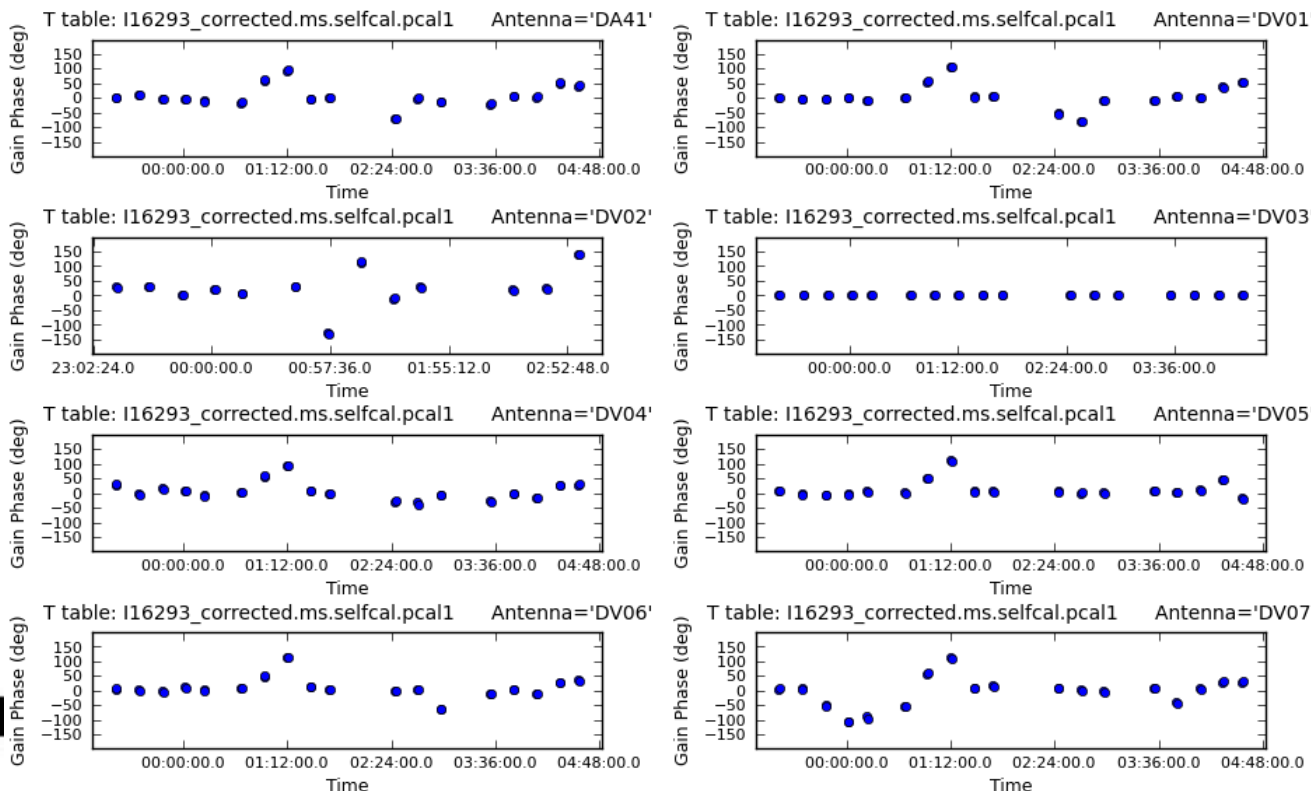
Clean boxes only around  
emission you are SURE are  
real at this stage

# Self-calibration Example:

## ALMA SV Data for IRAS16293 Band 6 (III)

Step 5: Decide on an time interval for initial phase-only self-cal

- A good choice is often the scan length (in this case about 5 minutes per field)
  - Exercise for reader: from page 31 show that  $S/N_{\text{self}} \sim 5.4$
- In CASA you can just set solint='inf' (i.e. infinity) and as long as combine  $\neq$  'scan' AND  $\neq$  'field' you will get one solution per scan, per field.
- Use 'T' solution to combine polarizations



What to look for:

- Lot of failed solutions on most antennas? if so, go back and try to increase S/N of solution = more averaging of some kind
- Do the phases appear smoothly varying with time (as opposed to noise like)

# Self-calibration Example:

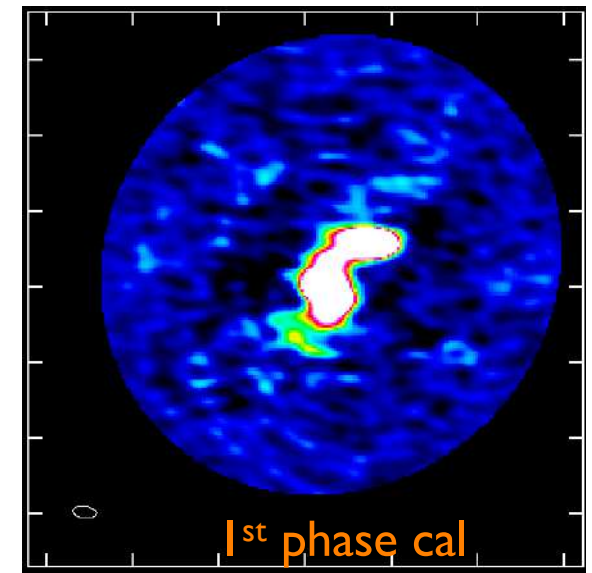
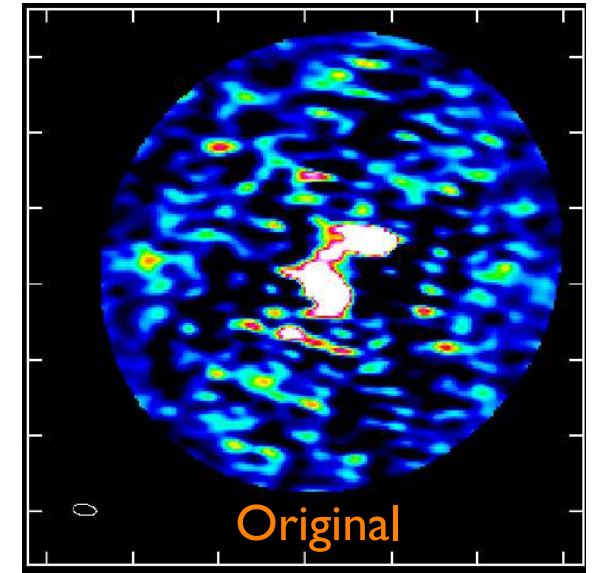
## ALMA SV Data for IRAS16293 Band 6 (IV)

Step 6: Apply solutions and re-clean

- Incorporate more emission into clean box if it looks real
- Stop when residuals become noise-like but still be a bit conservative, ESPECIALLY for weak features that you are very interested in
  - You **cannot** get rid of real emission by not boxing it
  - You can create features by boxing noise

Step 7: Compare Original clean image with 1<sup>st</sup> phase-only self-cal image

- Original:  
Rms~ 15 mJy/beam; Peak ~ 1 Jy/beam → S/N ~ 67
- 1<sup>st</sup> phase-only:  
Rms~ 6 mJy/beam; Peak ~ 1.25 Jy/beam → S/N ~ 208
- Did it improve? If, yes, continue. If no, something has gone wrong or you need a shorter solint to make a difference, go back to Step 4 or stop.

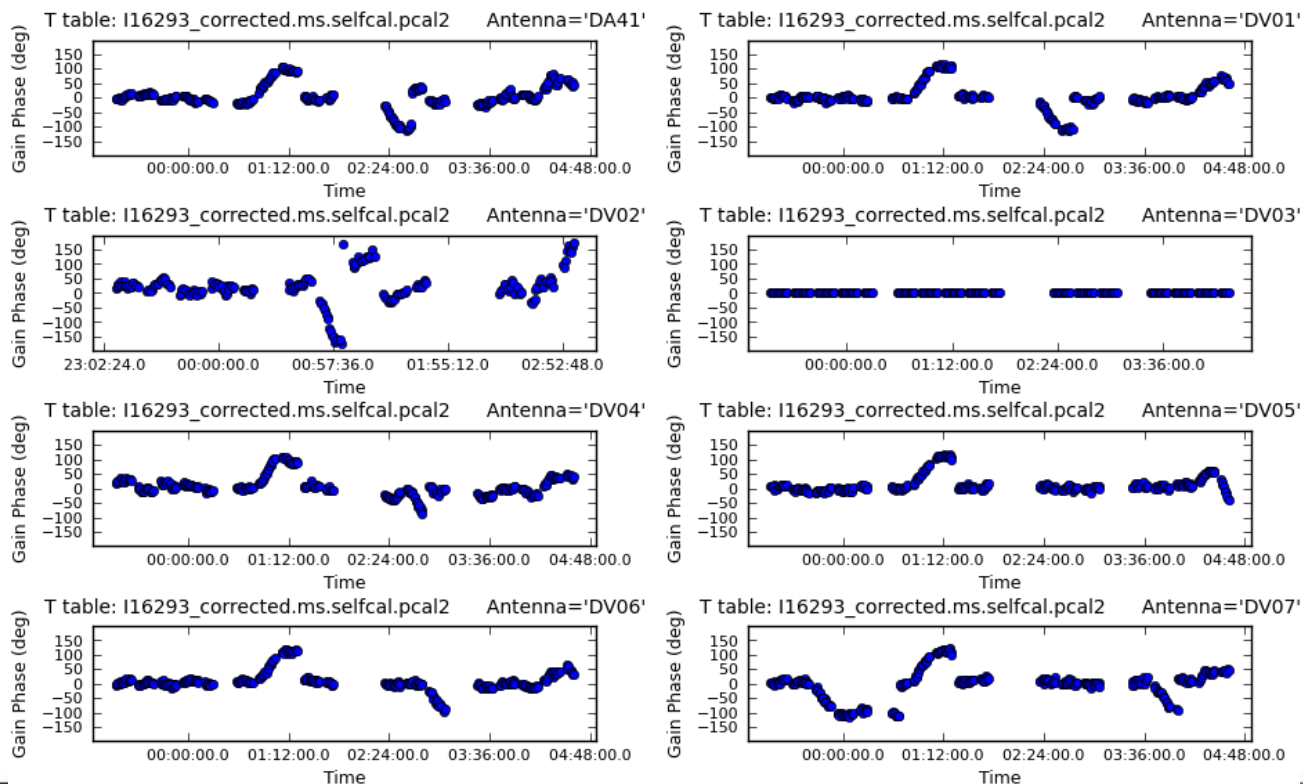


# Self-calibration Example:

## ALMA SV Data for IRAS16293 Band 6 (V)

Step 8: Try shorter solint for 2<sup>nd</sup> phase-only self-cal

- In this case we'll try the subscan length of 30sec
- It is best NOT to apply the 1<sup>st</sup> self-cal while solving for the 2<sup>nd</sup>. i.e. incremental tables can be easier to interpret but you can also “build in” errors in first model by doing this



What to look for:

- Still smoothly varying?
- If this looks noisy, go back and stick with longer solint solution
- IF this improves things a lot, could try going to even shorter solint



# Self-calibration Example:

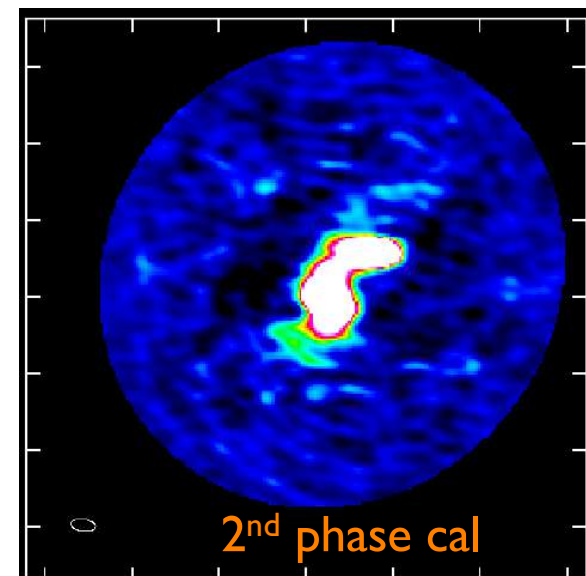
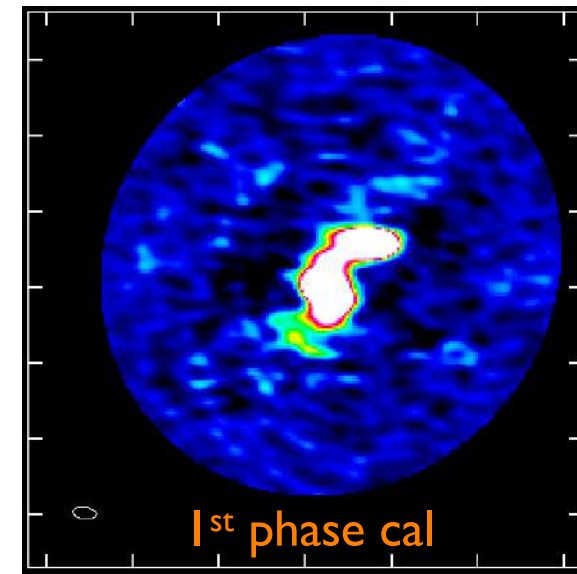
## ALMA SV Data for IRAS16293 Band 6 (VI)

Step 9: Apply solutions and re-clean

- Incorporate more emission into clean box if it looks real
- Stop when residuals become noise-like but still be a bit conservative, ESPECIALLY for weak features that you are very interested in
  - You **cannot** get rid of real emission by not boxing it
  - You can create features by boxing noise

Step 10: Compare 1<sup>st</sup> and 2<sup>nd</sup> phase-only self-cal images

- 1<sup>st</sup> phase-only:  
Rms~ 6 mJy/beam; Peak ~ 1.25 Jy/beam → S/N ~ 208
- 2<sup>nd</sup> phase-only:  
Rms~ 5.6 mJy/beam; Peak ~ 1.30 Jy/beam → S/N ~ 228
- Did it improve? Not much, so going to shorter solint probably won't either, so we'll try an amplitude self-cal next

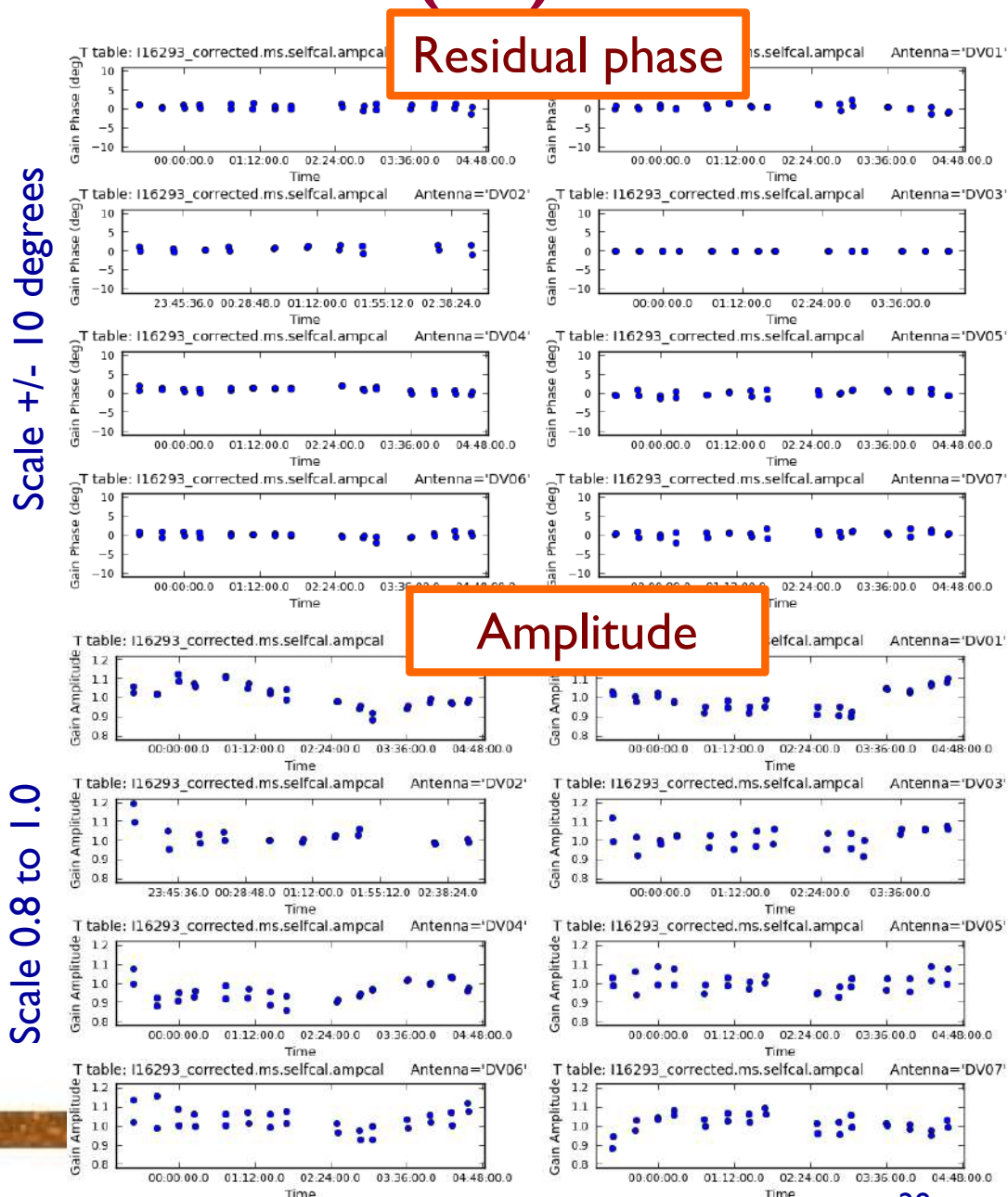


# Self-calibration Example:

## ALMA SV Data for IRAS 16293 Band 6 (VII)

Step 11: Try amplitude self-cal

- Amplitude tends to vary more slowly than phase. It's also less constrained, so solints are typically longer. Lets try two scans worth or 23 minutes
- Essential to apply the best phase only self-cal before solving for amplitude. Also a good idea to use mode='ap' rather than just 'a' to check that residual phase solutions are close to zero.
- Again make sure mostly good solutions, and a smoothly varying pattern.



# Self-calibration Example:

## ALMA SV Data for IRAS16293 Band 6 (VIII)

Step 12: Apply solutions

- Apply both 2<sup>nd</sup> phase and amp cal tables
- Inspect uv-plot of corrected data to
  - Check for any new outliers, if so flag and go back to Step 9.
  - Make sure model is good match to data.
  - Confirm that flux hasn't decreased significantly after applying solutions

Original

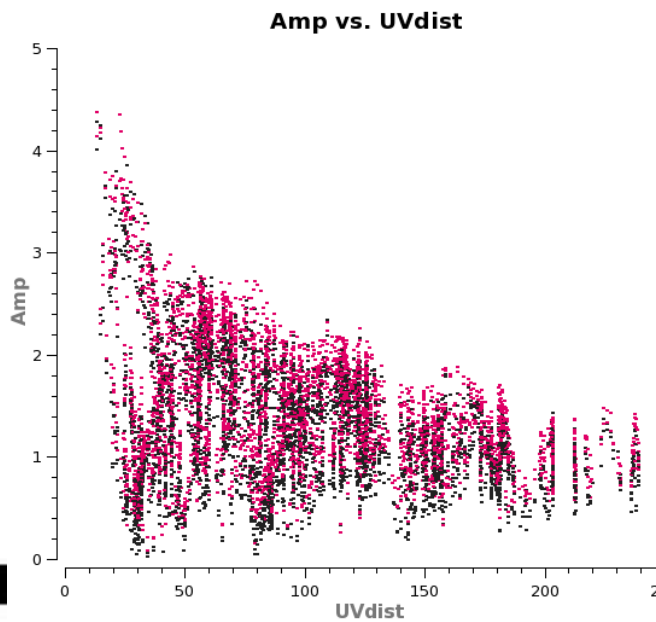
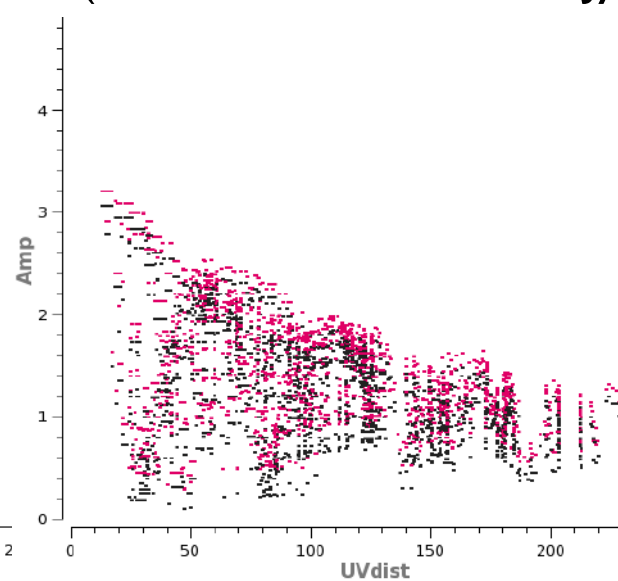
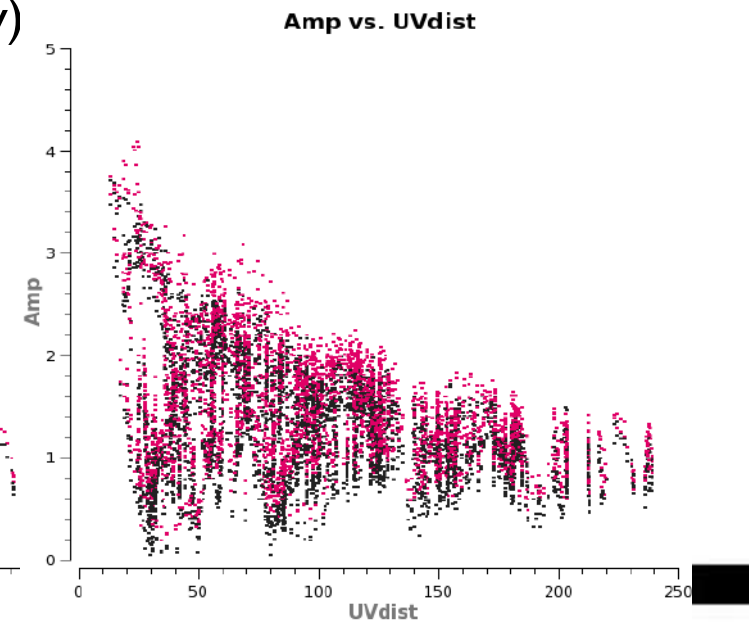


Image Model

(total cleaned flux = 3.4 Jy)



Amp & Phase applied





# Self-calibration Example:

## ALMA SV Data for IRAS16293 Band 6 (IX)

### Step 13: Re-clean

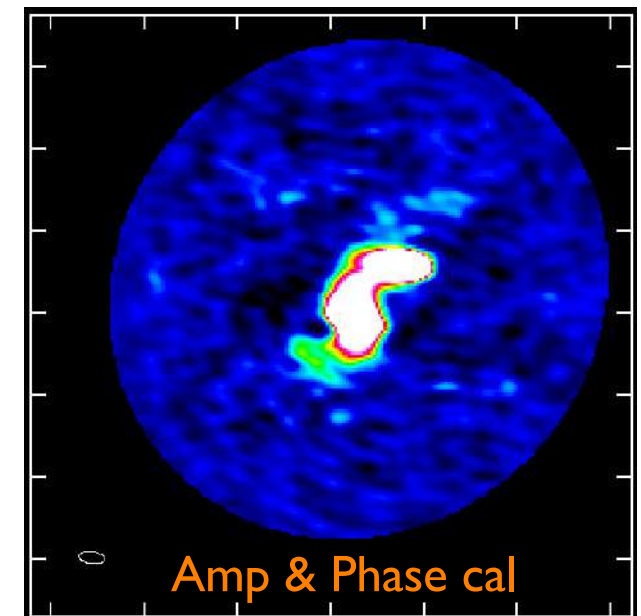
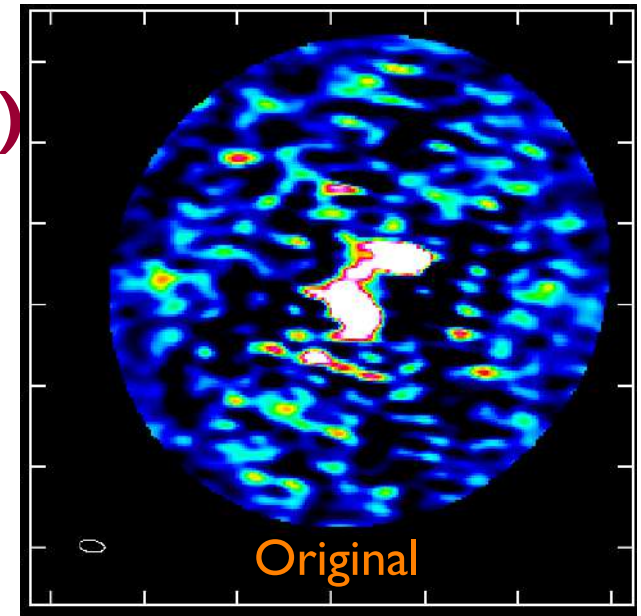
- Incorporate more emission into clean box
- Stop when residuals become noise-like – clean everything you think is real

### Step 14: Compare 2<sup>nd</sup> phase-only and amp+phase self-cal images

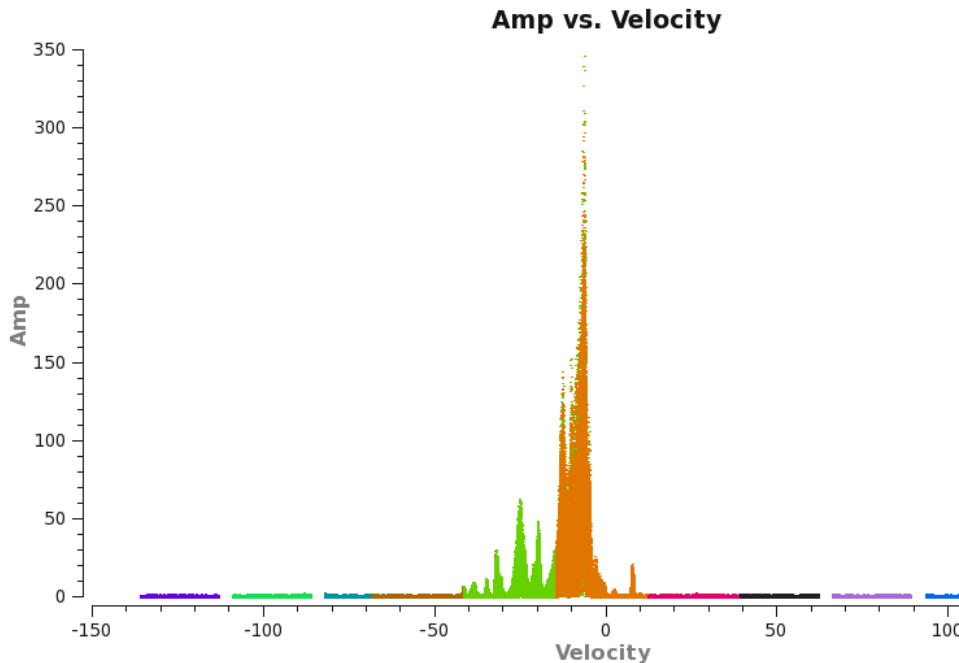
- 2<sup>nd</sup> phase-only:  
Rms~ 5.6 mJy/beam; Peak ~ 1.30 Jy/beam → S/N ~ 228
- Amp & Phase:  
Rms~4.6 mJy/beam; Peak~1.30 Jy/beam → S/N ~283
- Did it improve? → Done!

Final: S/N=67 vs 283!

But not as good as theoretical  
= dynamic range limit



## Self-Calibration example 2: JyLA Water Masers (I)



uv-spectrum after standard calibrator-based calibration for bandpass and antenna gains

There are 16 spectral windows, 8 each in two basebands (colors in the plot)

Some colors overlap because the basebands were offset in frequency by  $\frac{1}{2}$  the width of an spw in order to get good sensitivity across whole range.

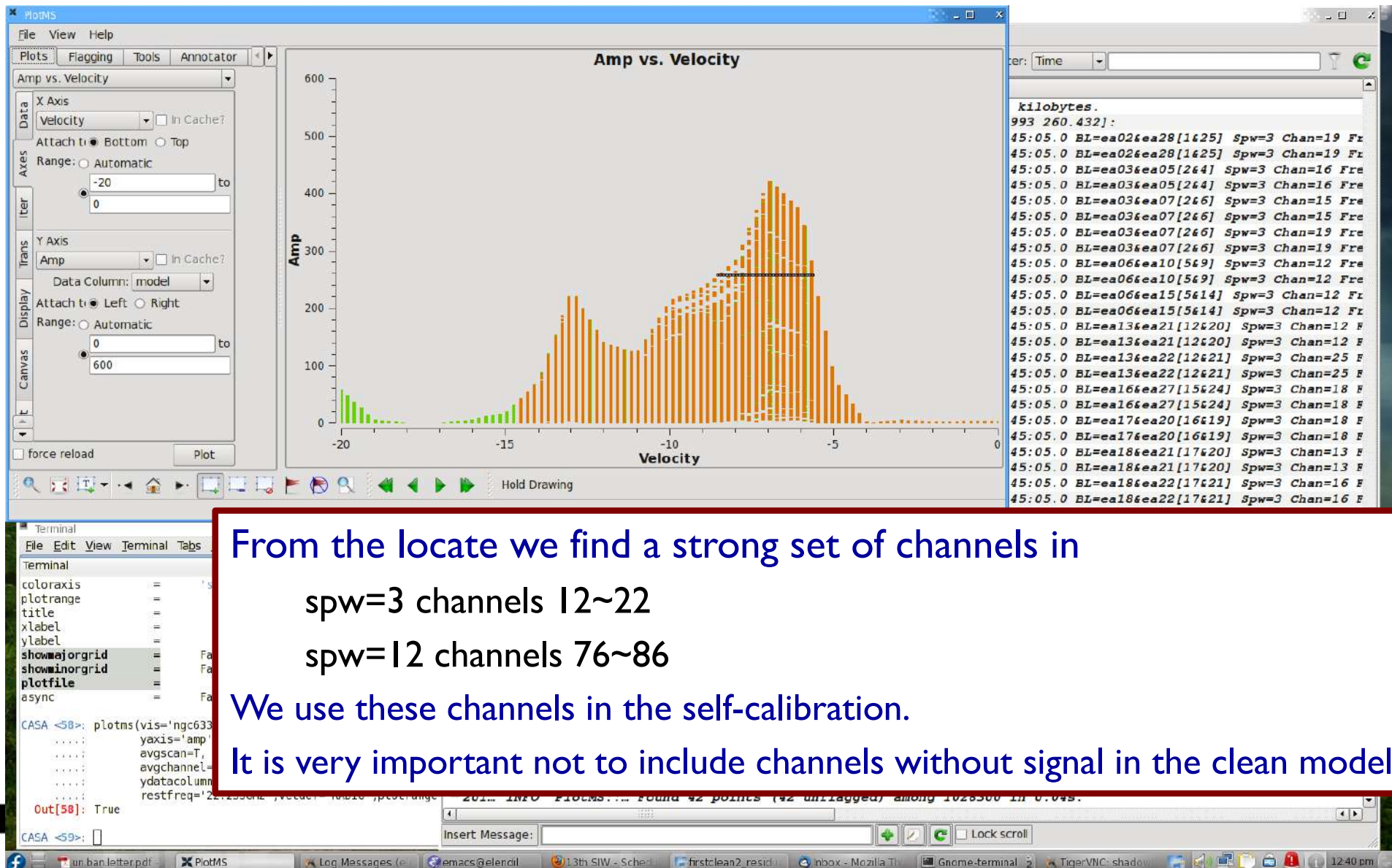
The continuum of this source is weak. How do you self-cal this?

- In general DATA CHANNEL NUMBER  $\neq$  IMAGE CHANNEL NUMBER due to Doppler Shift, also LSB windows will have negative channel width, i.e. data and image channel numbers going in opposite directions (as of CASA 5.1)
- Suggest running CVEL using the rest frequency of the line at the same velocity resolution that you want for the final cube – this will give you a uv-dataset with the same channelization as the cube you want DATA CHANNEL NUMBER = IMAGE CHANNEL NUMBER

# Self-Calibration example 2: JVLA Water Masers (III)

Need to know the SPWs and the CHANNELS with strong emission in the model:

*plotms* of the MODEL (no CVEL, so data channels  $\neq$  image channels) with *locate* can help



From the *locate* we find a strong set of channels in

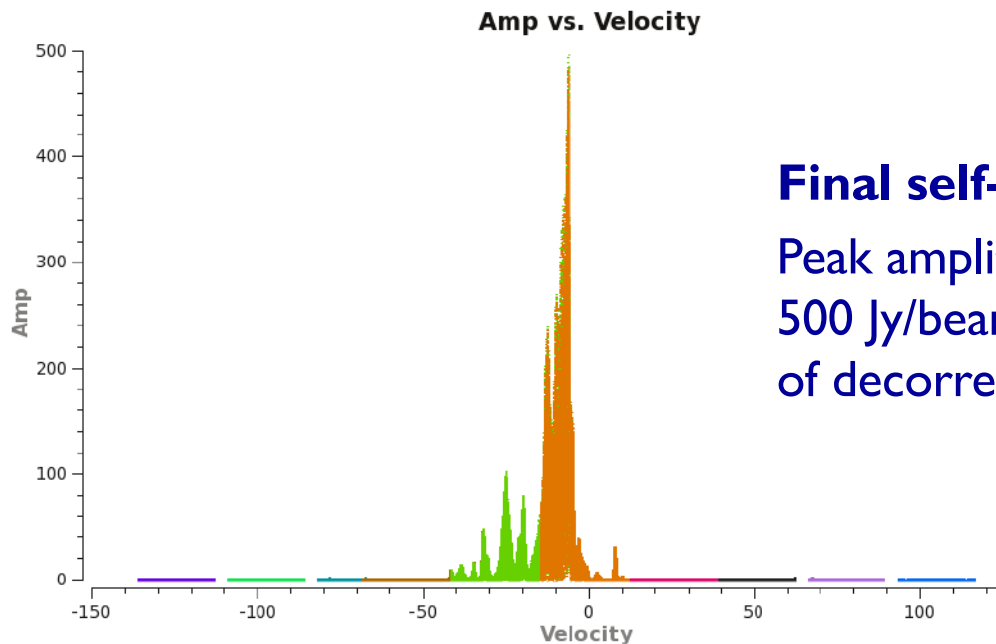
spw=3 channels 12~22

spw=12 channels 76~86

We use these channels in the self-calibration.

It is very important not to include channels without signal in the clean model!

# Self-Calibration example 2: JyLA Water Masers (III)



## Final self-calibrated spectrum

Peak amplitude increased from 350 Jy/beam to 500 Jy/beam (a 40% increase) due to correction of decorrelation

- One remaining trickiness: calibration solutions are only for spw=3 and 12. The spwmap parameter can be used to map calibration from one spectral window to another in applycal. There must be an entry for all spws (16 in this case):

`spwmap=[3,3,3,3,3,3,3,3,12,12,12,12,12,12,12]`

In other words apply the spw=3 calibration to the 8 spectral windows in the lower baseband and the calibration from spw=12 to the 8 spws in the upper baseband

- Beyond this everything is the same as previous example.





# Summary

- Spatial and temporal variations in the amount of precipital water vapor in the troposphere cause phase fluctuations but there are a wide range of options for corrections: observing techniques and post-processing
  - Fast switching
  - WVRs
  - Self-calibration
- Self-calibration is not so hard and can make a big difference
  - Make sure your model is a good representation of the data
  - Make sure the data you put into solver, is a good match to the model
  - If you are lacking a little in S/N try one of the “S/N increase techniques”
  - If you really don’t have enough S/N don’t keep what you try!
- For more examples, tips, tricks, and advice see <https://arxiv.org/abs/1805.05266>

## Advanced Gain Calibration Techniques in Radio Interferometry

Crystal L. Brogan, Todd R. Hunter, Ed B. Fomalont



# Calibration Sensitivities Effects (N=25)

$S/N_{\text{Ant}}$	Amp error	Phase error	$S/N_{\text{base}}$	$S/N_{\text{image}}$
0	100%	180 d	0	0
3	33%	15.0 d	0.6	11.0
5	20%	9.7 d	1.1	18.4
10	10%	5.7 d	2.1	36.9
25	4%	2.3 d	5.3	92.3
100	1%	0.6 d	21.3	370

$d_{\text{Ant}}$  phase error must be smaller than expected instrumental and tropospheric phase error which is often 10-20 deg

$d_{\text{Ant}}$  amp error must be smaller than expected instrumental and absorption amplitude errors, usually  $< 5\%$