



Master's thesis  
Astrophysical Sciences

# Supermassive black holes and the cosmological formation of massive early-type galaxies (title not final)

Atte Keitaanranta

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Supervisor(s): Prof. Peter Johansson  
M.Sc. Matias Mannerkoski

Censor(s): Prof. Peter Johansson

UNIVERSITY OF HELSINKI  
DEPARTMENT OF PHYSICS

PL 64 (Gustaf Hållströmin katu 2)  
00014 University of Helsinki





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# 1. Introduction

This is aaaaaaaaaaaaaaaaaa test.

## 1.1 Information about galaxies, shortly

## 1.2 Aim of the thesis

## 2. Background

### 2.1 Cosmology

2.1.1 Hubble parameter, Friedmann equations and so on

2.1.2 Cosmological perturbations

### 2.2 Early-type galaxies

2.2.1 Types of ellipticals

2.2.2 Photometric and kinematic profiles

### 2.3 Feedback processes

## **3. GADGET-3 and KETJU**

- Haven't really thought about the contents of this chapter yet

### **3.1 Overview of GADGET-3**

### **3.2 Smoothed Particle Hydrodynamics**

### **3.3 Gas cooling?**

### **3.4 Feedback?**

### **3.5 KETJU**



## **4. Creating initial conditions for the cosmological zoom-in simulations**

### **4.1 Zoom-in technique**

### **4.2 MUSIC**

#### **4.2.1 Basic properties of MUSIC**

#### **4.2.2 Algorithms of MUSIC**

- Short section on how MUSIC generates initial particle positions, velocity fields and so on (basically summary of the first few sections of the MUSIC paper)

#### **4.2.3 Creating an IC-file with a zoom-in box**

- Step by step explanation of creating the IC file with a zoom-in region

## 4.3 GADGET-3 setup for the zoom-in -simulations

### 4.3.1 Cosmological setup

### 4.3.2 Low-resolution run

### 4.3.3 Choosing the zoom-in regions

- FoF -algorithm
- Conditions of the chosen halos
- Figure showing the zoom-in regions from the low res run

### 4.3.4 Initial conditions

$h_0$	$\Omega_m$	$\Omega_b$	$\Omega_\Lambda$	$\sigma_8$	$\rho_{\text{crit}}$
70.3	0.276	0.045	0.724	0.811	$9.28 \times 10^{-27} \text{ kg/m}^3$

**Table 4.1:** Cosmological parameters used for the simulations. If a simulation doesn't include baryons, the dark matter density parameter  $\Omega_{\text{DM}}$  is equal to the matter density parameter  $\Omega_m$ . If baryons are included,  $\Omega_{\text{DM}} = \Omega_m - \Omega_b$ .

- Information from GADGET3 config files

# 5. Cosmological GADGET-3 simulations

## 5.1 Computational load of the simulations

- Quick overview: CPUs used, time elapsed, where simulations were run

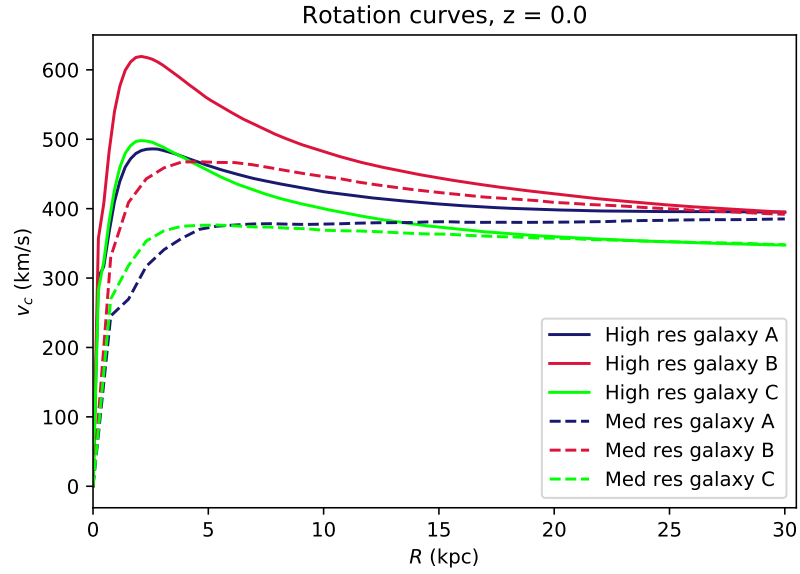
## 5.2 Locating galaxy centers: the shrinking sphere -method

## 5.3 Properties of the galaxies

Simulation	$r_{\text{vir}}$ (kpc)	$M_*$ ( $M_{\odot}$ )	$M_V$ (mag)	$M_{*,\text{gal}}/M_{\text{vir}}$	$M_{\text{bh}}$ ( $M_{\odot}$ )
Med res, A	517	$3.98 \times 10^{11}$	-22.4	0.025	$5.82 \times 10^9$
Med res, B	574	$5.16 \times 10^{11}$	-22.7	0.024	$6.23 \times 10^9$
Med res, C	400	$2.56 \times 10^{11}$	-22.0	0.035	$3.96 \times 10^9$
High res, A	526	$5.31 \times 10^{11}$	-22.6	0.032	$3.64 \times 10^9$
High res, B	578	$6.38 \times 10^{11}$	-22.8	0.029	$4.54 \times 10^9$
High res, C	400	$3.46 \times 10^{11}$	-22.2	0.047	$2.80 \times 10^9$

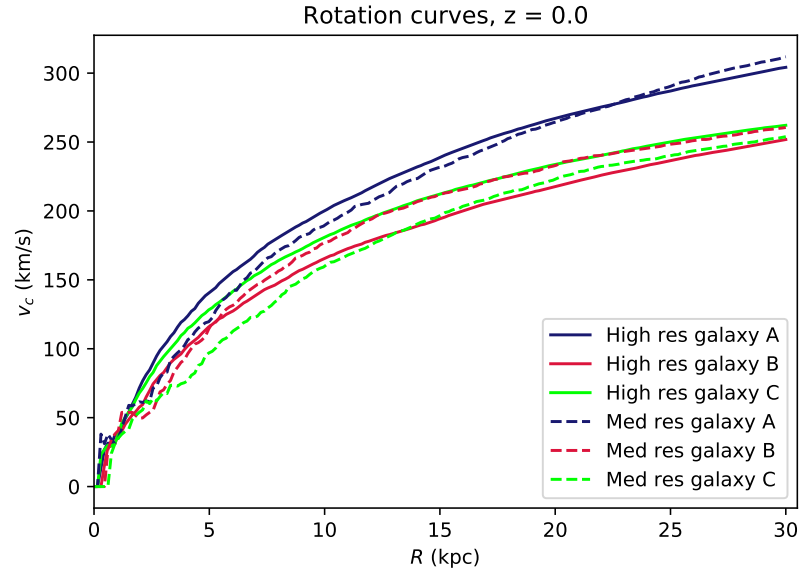
**Table 5.1:** Properties of the zoomed-in galaxies at redshift  $z = 0$ .

## 5.3.1 Rotation curves

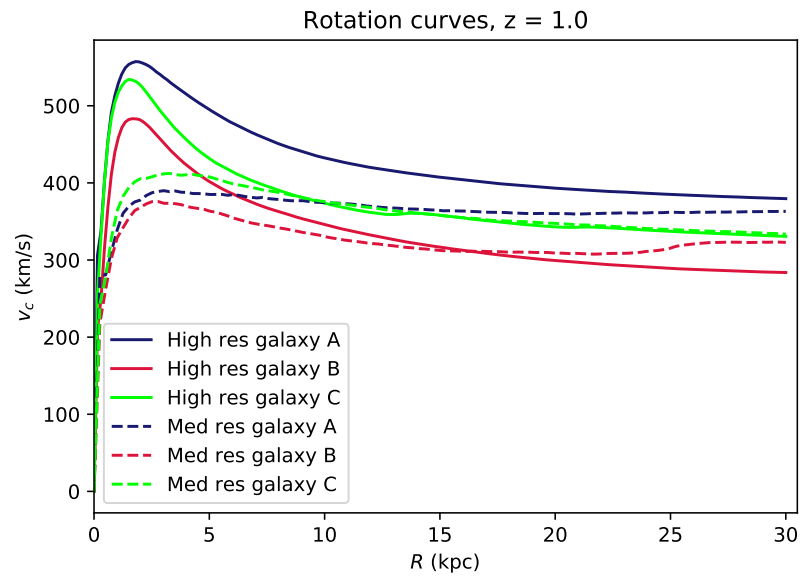


**Figure 5.1:** Rotation curves for each galaxy including baryons, at redshift 0. The continuous lines represent the high resolution simulations and the dashed lines represent the medium resolution simulations.

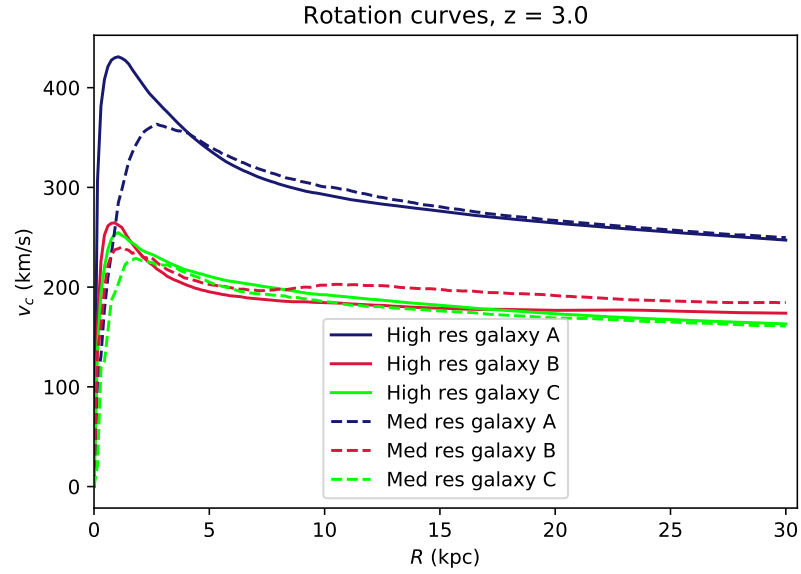
- Med res galaxy A simulation is performed with single precision, I'm currently doing the simulation in double precision
- Rotation curves do not have the same limits on the y-axis



**Figure 5.2:** Rotation curves for each galaxy including only dark matter, at redshift 0. The continuous lines represent the high resolution simulations and the dashed lines represent the medium resolution simulations.



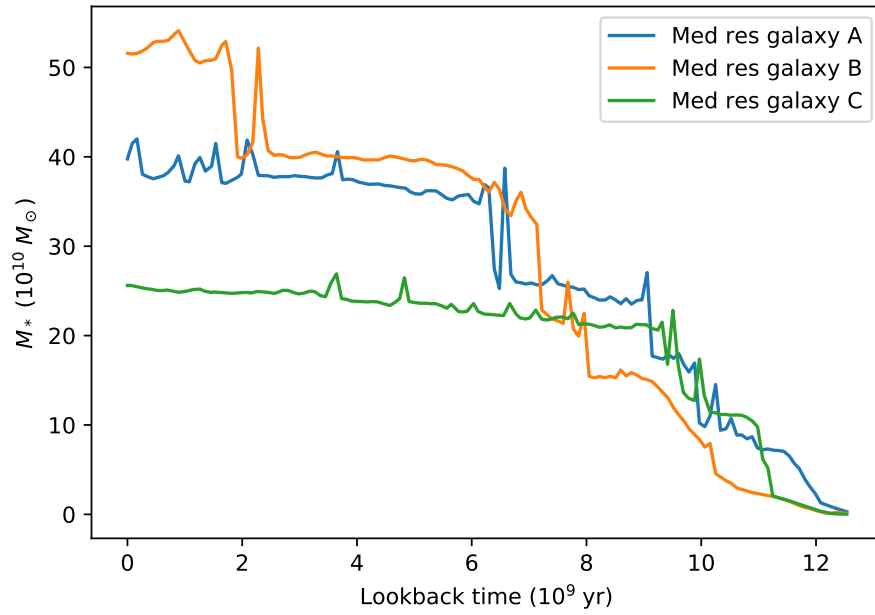
**Figure 5.3:** Rotation curves for each galaxy including baryons, at redshift 1. The continuous lines represent the high resolution simulations and the dashed lines represent the medium resolution simulations.



**Figure 5.4:** Rotation curves for each galaxy including baryons, at redshift 3. The continuous lines represent the high resolution simulations and the dashed lines represent the medium resolution simulations.

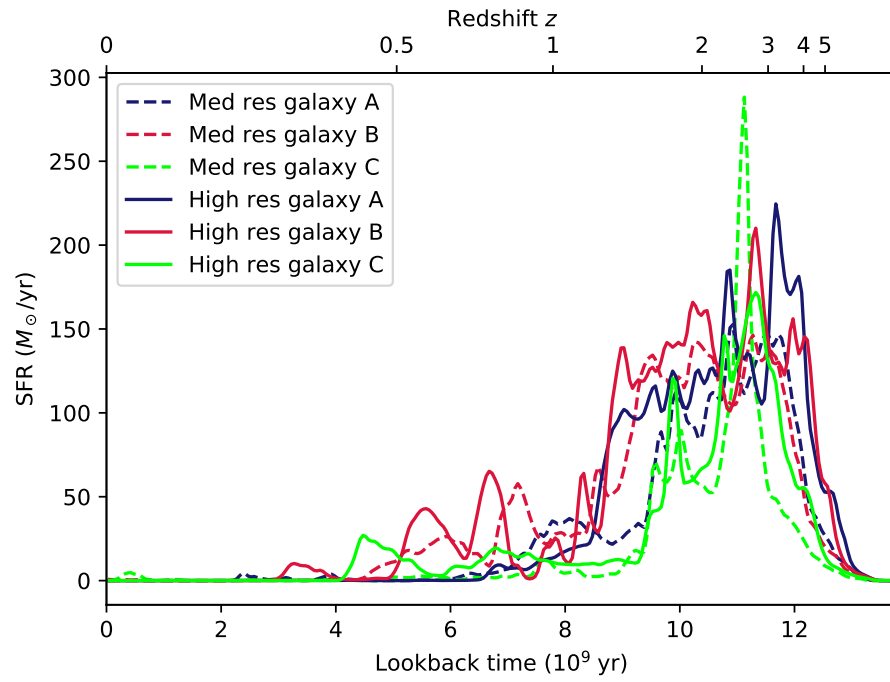
### 5.3.2 Star formation history

- The stellar mass evolution plot is not yet done for high res simulations.
- Redshifts missing from the stellar mass evolution plot
- SFRs, also histograms?
- Again, med res A results will probably change a bit when double precision run is finished.
- Formation efficiencies, comparing to the cosmological parameter



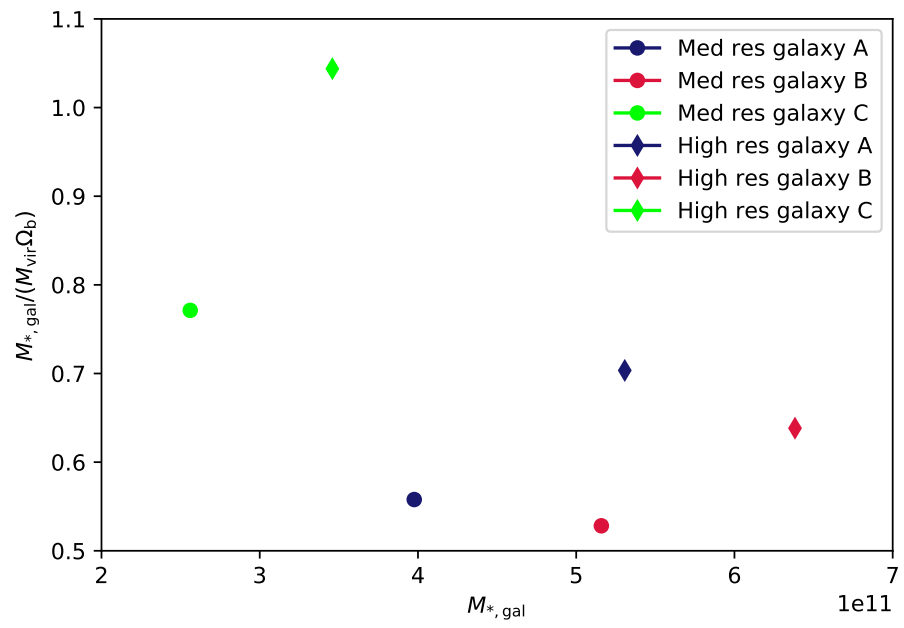
**Figure 5.5:** Stellar mass evolution for the medium resolution galaxies. The calculated stellar mass is the stellar mass within  $r_{\text{gal}} = r_{\text{vir}}/10$ .

### 5.3.3 Colors and magnitudes

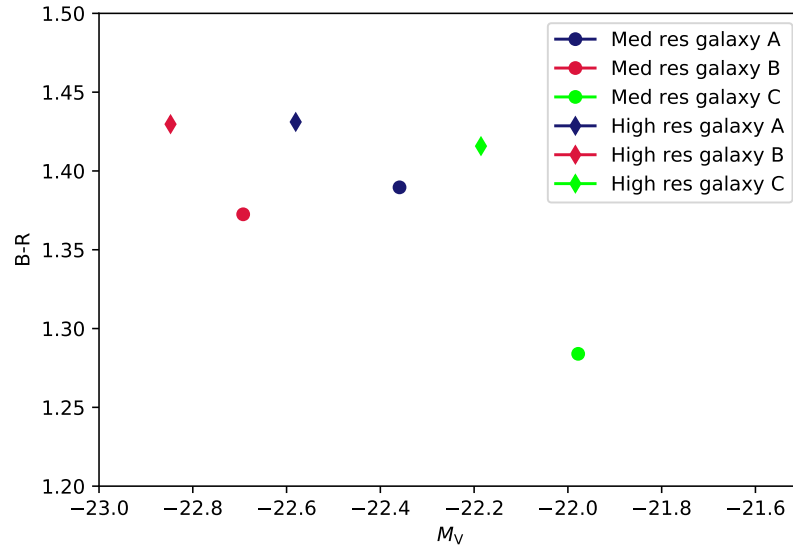


**Figure 5.6:** Stellar formation rates for each zoomed-in galaxy, plotted as a function of lookback time. The lines are created from histograms having a length of 5 Myr, which are then smoothed. The continuous lines represent the high resolution simulations and the dashed lines represent the medium resolution simulations.

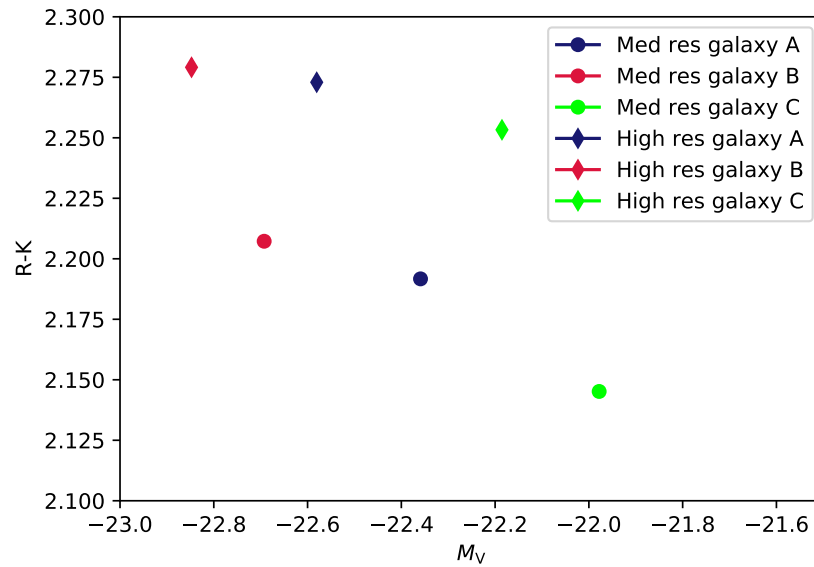




**Figure 5.7:** Galaxy formation efficiencies for each galaxy, plotted with their stellar masses. The cosmological baryon density  $\Omega_b$  is set to 0.045 in the simulations. The diamond and circular markers show the results of the high resolution and the medium resolution zoom-in simulations, respectively.



**Figure 5.8:** B-R colors for each simulated galaxy, plotted with each galaxy's absolute magnitude in the V-band. The diamond and circular markers show the results of the high resolution and the medium resolution zoom-in simulations, respectively.



**Figure 5.9:** R-K colors for each simulated galaxy, plotted with each galaxy's absolute magnitude in the V-band. The diamond and circular markers show the results of the high resolution and the medium resolution zoom-in simulations, respectively.

## **6. Simulations with KETJU**

## 7. Conclusions

- recap on what was written/studied
- more own thoughts on results
- future missions
- how could the simulations be more realistic (higher resolution, more feedback stuff?)