

Random Functions and Early Universe Cosmology

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1 Aims and Objectives

The theory of inflation is the current best explanation for certain perplexing observations of the Universe, such as the monopole problem, horizon problem, and flatness problem.[1] Inflation is thought to be caused by a scalar field. The simplest possibility is a one-dimensional scalar field analogous to the gravitational potential; however multidimensional scalar fields are also possible. In standard inflationary theory, the universe inflates when the gradient of the scalar field is small (so-called “slow-roll inflation”), and stops inflating when we reach a minimum (a “vacuum”). Early work on inflation led to the idea of *eternal inflation*, whereby most of the universe inflates forever, but some patches reach local minima. These patches then become *pocket universes*, each with its own laws of physics, one of which corresponds to our own. Of particular interest is that, if we are in one of these patches with a local minimum above zero, this “vacuum energy” can manifest itself as a nonzero dark energy term – and we observe dark energy. Hence, it is desirable to understand the nature of the scalar field behind inflation.

The precise nature of the scalar field behind inflation is unknown. However, the study of flux-compactified string vacua predicts the existence of a highly complex, $O(100)$ dimensional scalar field known as the *landscape*. This landscape contains a vast number of vacua which can, in principle, be the scalar field of inflation. The main objective of this research is to map, given the landscape, the probability that we are in a minimum with vacuum energy equal to what we observe as a function of the number of dimensions N in the landscape.

There are two ways to attack the problem. The first is to try to understand the string landscape better, and derive a probability from first principles. One of the recently-proposed swampland conjectures is an example of this:[6] if the conjecture is true, then there are no minima for positive V , which represents a challenge since we observe a positive vacuum energy. The other approach is to ignore the theory behind the landscape and simply treat it as a statistical aggregate, invoking the central limit theorem for its properties. The string landscape is astonishingly complex – an often-quoted estimate of the number of vacua is 10^{500} [5] – and therefore, most studies of the landscape take the latter approach.[2, 3, 4] We assume the landscape is well-approximated by a Gaussian random field, and try to determine the probability of a minimum with a value equal to the one we measure.

2 Contribution

By definition, Gaussian Random Fields have set correlations between its field value, first derivative, and second derivative at every point. For example, the first derivative and second derivative of a Gaussian Random Field are completely uncorrelated. They are generic constructs not specific to cosmology, and their properties have been widely studied. One early paper dealt with the relative abundances of maxima, minima, and saddles.[7] The authors took the Hessian of a random function to be symmetric matrices with random individual elements, or more precisely, with values randomly drawn from a Gaussian distribution (called a Gaussian Orthogonal Ensemble [GOE]). From elementary calculus, a minimum exists if the eigenvalues of the Hessian are all positive, a maximum exists if they are all negative, and there is a saddle if there are eigenvalues of both signs. Using the GOE, they find that the number of minima is super-exponentially small compared to the number of saddles. This approach was later shown to be incorrect, because the eigenvalues of the Hessian of a random function are not independent of each other [8]. Taking this into account, the authors found that the number of minima decreases exponentially with the number of dimensions N instead, and the ratio of extrema to saddles to be approximately binomial.

In addition to the probability of minima, the distribution of eigenvalues of the Hessian about minima as well as the distribution of the Hessian itself as a function of N has also been investigated.[9][10] Yamada and Vilenkin[9] found that in a Gaussian random landscape the smallest eigenvalue is of order $\sim 1/N$, which, if one assumes that the most probable decay pathway is in the direction of the smallest eigenvalues, improves the stability of the minima by about \sqrt{N} .

Aside from the properties of the field itself, questions such as the probability of moving from a random point in the potential to a viable late-time de Sitter minimum by Dyson Brownian motion have also been asked.[11] Pedro and Westphal calculated the probability of a graceful exit from random points in the landscape, and found that there is an exponential bias against small-field inflationary ranges, which allows us to conclude that the observation that the universe is flat imposes a $N \ll 10$ constraint on the number of light fields involved in inflation.

None of these works, however, have explicitly examined the probability that a field point of known value is a minimum or maximum (Bardeen et al[12] touched on it while dealing with a different problem on the number

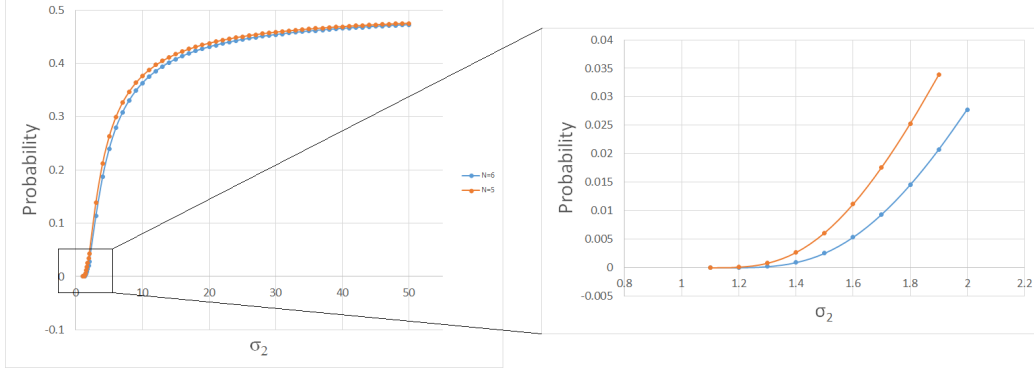


Figure 1: The probability that a given extremum with $\phi > 0$ is a minimum as a function of σ_2 , for $N = 5, 6$. A close-up of the probabilities for small values of σ_2 is shown as well. The probability for $N = 5$ is always larger than that for $N = 6$.

density of stationary points). So far in this project, we have explicitly generalized Bardeen et al's (henceforth BBKS) calculation from 3-dimensions to N -dimensions, reaching a maximum of $N = 10$.

3 Preliminary results

The final probability depends not only on the field value ϕ but also the moments of the power spectrum $\sigma_0, \sigma_1, \sigma_2$. That the probability depends on the intrinsic properties of the field is intuitively obvious, since a very turbulent random field will have a lot of minima and maxima, naturally pushing the probability to ~ 0.5 . It turns out that two of the three σ 's are normalizable. σ_0 is closely related to the average value of the field, so with an appropriate choice of the zero of the field, we can set $\sigma_0 = 1$. Similarly, σ_1 represents the average of the first derivatives, and with an appropriate rescaling of the axis length we can also set it $\sigma_1 = 1$. This leaves σ_2 as the only remaining parameter, and all the results we present will be in terms of σ_2 .

We performed analytical calculations up to $N = 6$. The results are given in Fig. 1. Here we plot the probability that a given extremum with $\phi > 0$ is a minimum as a function of σ_2 . As expected, if the potential is highly oscillatory (i.e. σ_2 is large, or γ is small), the probability tends towards 0.5

– equal likelihood of an extremum being either a maximum or a minimum. Conversely, if the potential is very smooth (σ_2 is small), the probability tends towards 0. Further, for any given σ_2 , the probability decreases with increasing dimension.

For $N > 6$, the calculations are complex enough that we have to attack the problem numerically. This works up to $N = 10$. The results are given in Table 1. The same trends are seen: for any given σ_2 , the probability decreases with increasing dimension.

Table 1: The probability that a given extremum with $\phi > 0$ is a minimum as a function of σ_2 . Values given for $N = 4, 5, 6$ are exact; for larger values they are approximate (see the text; errors are on the order of the third nonzero digit after the decimal).

N	σ_2	p_{min}	N	σ_2	p_{min}
4	10	0.60859	4	2	0.066638
5	10	0.62333	5	2	0.043103
6	10	0.63680	6	2	0.030844
7	10	0.35210	7	2	0.02013
8	10	0.34227	8	2	0.01307
9	10	0.33159	9	2	0.00845
10	10	0.32151	10	2	0.00544
4	3.5	0.21520	4	1.25	0.001526
5	3.5	0.17886	5	1.25	0.000327
6	3.5	0.15274	6	1.25	6.53328×10^{-5}
7	3.5	0.13599	7	1.25	2.09673×10^{-5}
8	3.5	0.11711	8	1.25	3.97235×10^{-6}
9	3.5	0.10093	9	1.25	7.13489×10^{-7}
10	3.5	0.08729	10	1.25	1.20253×10^{-7}

4 Next steps

Next up on the agenda is to investigate the behaviour of the field around saddles. These are stationary points which are not maxima or minima – they increase in some directions and decrease in others. We do this because this

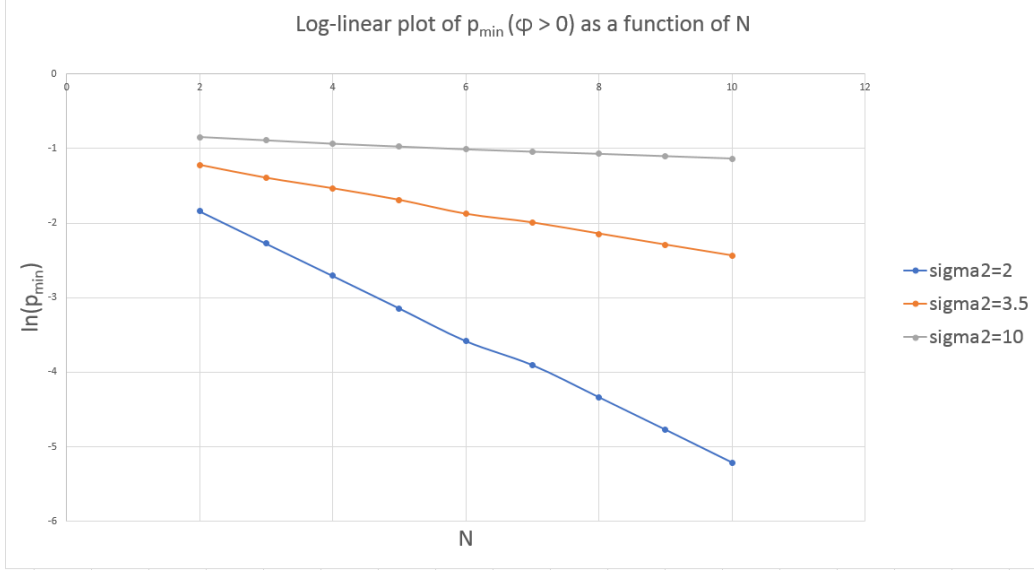


Figure 2: A log-linear plot of the probability that a given extremum with $\phi > 0$ is a minimum for a given σ_2 , as a function of the number of dimensions N .

is where inflation occurs. In standard inflationary theory, slow-roll inflation occurs when:

$$\epsilon \equiv \frac{m_{Pl}^2}{2} \left(\frac{V'}{V} \right)^2 \ll 1 \quad |\eta| \equiv m_{Pl}^2 \frac{|V''|}{V} \ll 1 \quad (1)$$

Here m_{Pl} is the Planck mass. It's relatively easy to find points where $\epsilon \ll 1$ – that will naturally happen about a maximum or minimum since $V' = 0$ at those points, and we already know extrema exist. Not so easy is to find points where $\epsilon \ll 1$ *and* $\eta \ll 1$. If our universe had inflated in the vicinity of these points, then it would have inflated for a varying amount (known as the number of e-folds of inflation) that depends on the value of η . We need 40–60 e-folds of inflation to solve the cosmological problems inflation was created to solve, [14] so we want to know $P(\eta)$ given that we need 40–60 e-folds of inflation. This compares to the already-accomplished work, where we solve for $P(\Lambda)$ given that we observe dark energy in the present universe.

Another thing we will investigate is how likely our universe is given that we observe a scalar spectral index n_s of about 0.96 [15]. This observable

parameter is related to the slow-roll parameters via $n_s = -2\epsilon + 6\eta$, and acts as another link between our work and observations.

Note both ϵ and η depend on the potential, its gradient, and its second derivative. These are things which the previous analysis equips us to analyze (they are after all intimately related to the σ 's).

5 Ethical approval

Not needed.

6 Resources

To perform the calculations, we need computational power and *Mathematica*. Both of these we already have. If we turn out to need more computational power than a standard machine can provide, we also have access to the Nectar Cloud.

7 Thesis structure

I intend the thesis to be focused on work to be published over the next two years, which will focus on this topic and possibly others.

- Introduction
 - Gaussian random fields
 - The string landscape and why approximating the landscape as Gaussian random fields makes sense
 - How inflation is related to the structure of the landscape
- Theory
 - A first-principles derivation of the theory required to make educated statements about the landscape's properties
 - Will involve a *lot* of equations
- Results

- Summary of the papers that will arise from this work
- Probability of minimum at a known field value given that an extremum is there
- Statements about the slow-roll parameters ϵ and especially η , and the implications these have on the number of e-folds of inflation we can expect as well as the scalar spectral index n_s
- Conclusions

8 Proposed timeline

- Immediately or in the next two months: finish and submit the paper on the probability of minimum at a known field value given that an extremum is there.
- Now–March 2019: Investigate the properties of ϵ and η in the landscape. If possible extend these results as well as the previous one to even higher dimensions than $N = 10$.
- April 2019–October 2019: Connect these properties to observables. Publish these results.
- November 2019–June 2020: Potentially we will change directions completely to a different subfield in cosmology; if not, continue to develop these results. For example we could ask how far saddles are from one another, and whether it's likely that the universe will undergo two separate periods of inflation.
- June 2020–September 2020: Write up and defend thesis.

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