## Newton (a)

We have the following functions:

Note that I have replaced with to avoid any confusion. Differentiating the second function with respect to gives:

By using the main equations, we can simplify the RHS by replacing them with their equivalents:

Leaving the term with at the RHS, we get:

We can right LHS as a single derivation after multiplying both sides with :

Now, we divide it back to to get rid of at the RHS. This gives:

From now on, we have to use other equalities given to us to simplify further. The following are given:

Let , where is a constant.

Using chain rule, we can calculate that is in the equation above.

Putting this and into the equation above, we get:

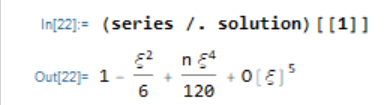
Further simplification on the RHS leads to:

The last thing is to scale the radius . Let where . Note that Noting

, we get:

Performing the simplifications and moving the terms at the RHS to the LHS, we get the Lane-Emden equation:

Since it requires more than a single-line command, I will provide a Mathematica Notebook that prints the series for regular solutions at the center. The main idea is to create a series of unknown coefficients, and then determining these coefficients by using the Lane-Emden Equation. Here is the result and see the notebook fpna.nb for the details:



For , the following command solves the differential equation:

Note that I have expanded the equation before using Mathematica. FullSimplify command is to simplify the output result, otherwise it gives it in the form of imaginary numbers . The result is:

Put it into the main equation to check.

Therefore, the solution is correct.

Using the following equation that is provided to us, we can find the total mass of the star.

Since , then . The boundaries are also must be modified. and similarly, .

Using the fundamental theorem of Calculus, we get:

This is the same result in the project manual. Let us denote with , by letting . Also noting , the equation for we found becomes:

At last, writing as , which are equal, we get the version in the manual:

Denoting the already defined as:

Then, as calculated above,

Putting this into the equation above, we get:

The final observation is the following two:

Therefore,

More precisely,

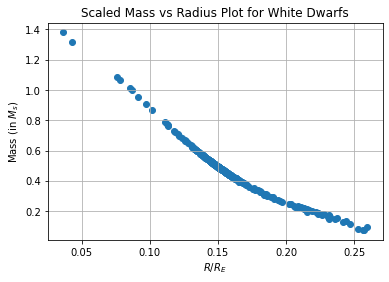
So,

(The term before M is the constant of proportionality.)

## Newton (b)

Using the following identities, one can extract from :

, where is the mass of the white dwarf. However, first, conversions must be done since our s are in format, s are in sun masses and must be scaled with for plotting. When all these are performed, the resulting plot becomes:



## Newton (c)

Series expansion of the given equation (Eq. 8) around since gives the following results:

The leading term in this series expansion is . It is also given that:

Substituting this for in the leading term gives:

This directly gives:

First, let us modify the following equation we got in the preceding question:

Then, exponentiating both sides with :

or equivalently:

An important observation can be that when , is independent of and when , is independent of . Therefore, it may be reasonable to expect .

Now, putting this equation in a much simpler form (Shankar, 1958, p. 98), we get:

by letting where . The nice part of this is for the values of between , (Shankar, 1958, p. 96). Therefore, equation simplifies to:

for .

We are to make a curve fit, so let it write in the following form by letting :

Now, there is still a problem because this equation depends on both and . To be able to perform curve fit, I will first find a reasonable value for and then find from the curve fit.

Using the condition I have imposed above:

, which is equal to (as long as ):

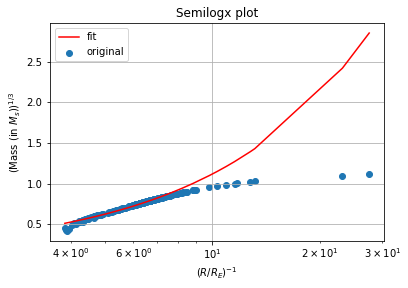
Due to ,

Due to ,

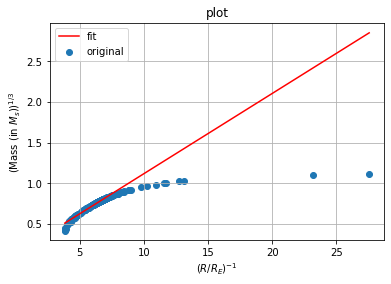
Their intersection allows us that only is possible (note that I have used the hint in the manual that restricts to be an integer). If , then .

Therefore, equation above becomes:

The last step is to define and and fit a line by discarding high mass stars. Below, I present the semilogx plot for 30 heavy stars discarded:



, which is a good fit for low-mass stars. It is actually linear, as expected:



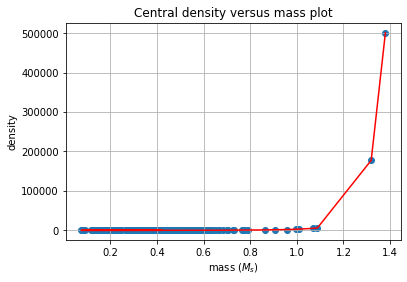
The slope is found to be , and the corresponding value is .

Recalling:

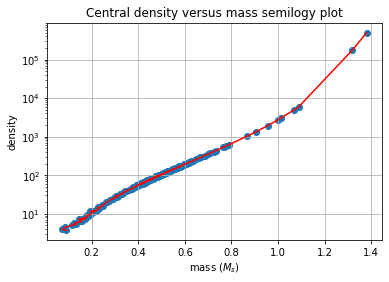
, we can calculate as

Since ,

, where for , is calculated as by Shankar (1958, p. 96). The resulting plot is as follow



As mass increases, density also increases. For visibility, semiology plot below can be more accurate to observe low-mass star central densities.



The shape and the overall behavior of the plot matches with the ones in the literature.