

The life of a star

Autonomous project : English

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Abstract

This paper aims to give to the reader a clear understanding of what are the different types of star, and why it is important to put them in different boxes. It has a scientific point of view, but requires no mathematical background.



Star cluster captured by the Hubble Telescope

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Introduction

History of astronomy and astrophysics

Astronomy is one of the oldest science. People have always studied and observed the sky and planets. Its first mean is to understand the origins and the evolution of the Universe, it also allows to study the movements, positions and behaviors of celestial objects. Concerning astrophysic, it's a more modern science. Astrophysic studies the physical aspect of cosmic phenomena such as birth and evolution of stars and galaxies. In this project, we aimed to summary the astrophysic of stars : from the birth to the destruction.

In the antiquity, men observed the Sun's position through the days, and star's positions during the night to measure time. Those luminaries also revealed themselves to be useful for navigation. It is thought that Babylonians started studying the sky in 3000 BC.

Mathematics theories were developed by ancient civilisations (greeks, indians and chinese). They tempted to predict solar and lunar eclipses. That is how Aristote discovered that Earth was round, Erastothenes calculated the Earth diameter etc. The first telescope was created by Isaac Newton in 1607. Before this, Galileo used a refractor bezel. In the XVIII century, the British William Herschel inventoried 848 stars. In 1781, he discovered Uranus. At this time, only 7 planets have been observed. Then in the XIX century, telescopes with prisms allowed scientists to study chemical composition of stars.

The importance of stars in the universe

Stars are everywhere in the universe. They are far from being only the little dots we can see at night when the sky is clear. They are the power plants of the universe and its most common inhabitant. They also are the forge of new matter, as every atom in your body that isn't hydrogen was once created in the heart of a star and expelled in the universe when the star exploded, to later coalesce in matter cluster, forming planets and maybe later, you. These huge spheres of burning gas are essential the universe, and is essential for us, as the sun is our biggest energy source and our only light source.

With 10^{22} stars in the universe¹, this makes *the* most frequent celestial body in space, some of them living as much as 100 billion years while others

¹About the same as the number of sand grains in planet earth.

explode being only few million years old, their complexity is yet to be fully understood.

1 The theory principles behind a star

1.1 The creation process

Stars are created every day, by a random process. They don't arise ex-nihilo, but are formed in **nebulae** that can be as much as 300 light-years² wide. In a nebula, there are huge clouds of gas in movement. One slight event can provoke the creation of a star, while nothing can happen in millions of years, making the creation process very random.

In the stars, gas clouds and the universe in general, everything a question of balance. The clouds within the nebula must maintain balance between their inner pressure and their temperature. These are explained by the **equation of state of an ideal gas** :

$$P \cdot V = n \cdot R \cdot T \quad (1)$$

- P : the pressure
- V : the volume
- n and R : constants of the gas cloud.
- T : the temperature

The temperature within the nebula tends to increase, due to light rays from nearby stars providing a slow heat source, just like the sun with earth. As we can see in this equation, an increase in temperature will trigger an increase in volume and pressure. As the pressure steadily increases along with the temperature, the particles within the gas speed up, as seen in the **equation of particle speed in function of temperature** :

$$v = C \cdot \sqrt{T} \quad (2)$$

- v : the particle speed
- T : the temperature
- C : constant of the gas cloud.

As the particle speeds up, more and more collision occur within the cloud and just like a lighter, the cloud start to ignite and the star is born. The

² 1 light year is over 9.5 Billion kilometers, in comparison, earth-sun is 0.15 Billion kilometers.

ignite is actually the fusion³ process beginning which is the essence of what a star is and why it *lives*.

1.2 The fusion process

The fusion process is the essence of the star, it makes it shine, and allows it not to collapse into a **black hole**⁴. Atoms, and matter in general contain energy, energy and mass are essentially the same thing⁵.

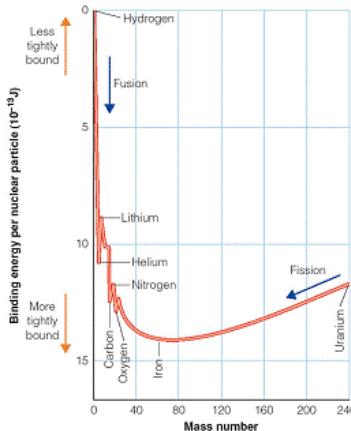


Figure 2: Energy contained in atoms

As said earlier, stars are mainly composed of hydrogen, and they use it to gather energy, the energy they need to shine. As shocks occur, the **Hydrogen in the star becomes Helium**. As we can see in Figure 2, Helium contains less energy than hydrogen, which means that the energy difference when the transformation occurs is released in light and heat.

The fusion process follows this pattern :

The interesting fact to notice is that fusion auto-fuels itself : it needs heat and the process generates heat.

The star's boring and repetitive life is to maintain an equilibrium between the fusion process and the gravity, the first one expanding the atoms and the star, and the second one collapsing it on itself.

Since in the world, there are a lot of different types of star, we will first begin by explaining the different classifications and how they are important.

³Explained in next section

⁴A black hole is a object with huge density, when gravity collapses atoms on themselves, they are explain in section REF

⁵Einstein exposed a link between them : $E = mc^2$ which translates that you can actually convert energy and mass easily.

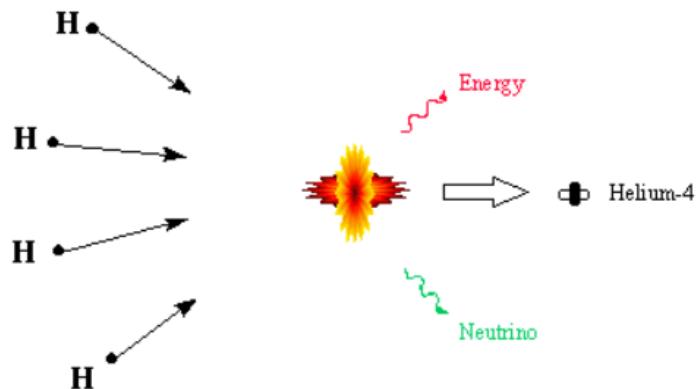


Figure 3: Hydrogen fusion

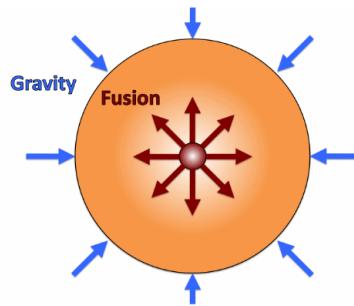


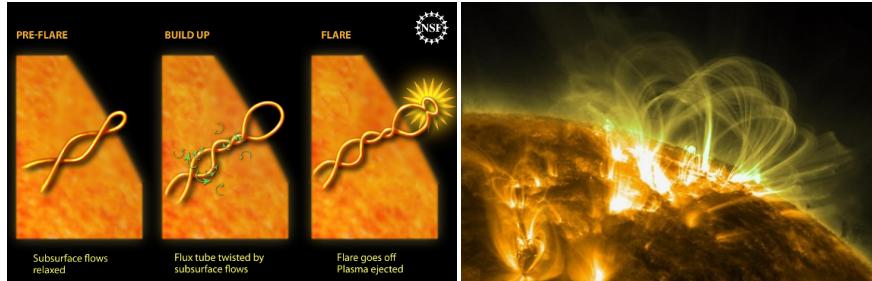
Figure 4: Balance between fusion and gravity

1.3 The solar flares : how this phenomenon was discovered

The balance between the star's gravity and its pressure was discovered long ago by Galileo thanks to solar flares.

Solar flares happen in the atmosphere of the sun. It is a brutal release of energy at the surface of the sun : more specifically, it is an emission of light and particles. For the biggest flares, plasma⁶ can be rejected too.. This flow of charged particles can have huge consequences on our systems and scientists are looking forward to forecast them so that they study carefully and meticulously this phenomenon. An example of the consequences of solar flares is perturbation of a satellite or the provocation of problems on electrical networks and phone networks.

⁶The plasma here isn't what is in your TV, or in your blood, but rather a very hot state of matter, just like solid or liquid, but where particles are so hot they form a sort of soup. There can be massive amounts of matter rejected in outer space, as much as 10 million tons. This seems huge but it is nothing compared to the actual size of the star.



(a) Solar flares phenomenon

(b) Solar flare's manetic fields

Indeed, magnetic fields created during a solar flare are turbulent and unpredictable. The process is that field lines wind and tangle like a rope (as shown in the following picture). They ultimately break and the energy contained in the structure is released when accelerated charged particles meet the gas. The flare do not occur at the very surface of the sun but a little further, in a layer called the corona of the sun.

This scheme points out the winded magnetic fields which create energetical flows. When it breaks, the plasma is ejected with its energy. This is one of te reason matter is thrown out of the star during its lifetime.⁷

Even if the scientist are sure of the flare's causes, they are still working on the mechanisms behind it, for example they still do not know how the magnetic field increase the particle's speed.

Another consequence of the solar flares are the beautiful aurorapolar caused by the solar winds. High speed very magnetic particles are forced in outer space in random directions. If some of them reach earth, they are directed towards the poles thanks to earth's magnetic field's shape. The remaining particles then hit the pole's atmosphere at full speed, an emit light of unusual colors as they are unusual light particles.

2 Introduction to stellar classifications

Stars can be sorted out according to a variety of different physical parameters:

- Their position in the **lifecycle of the star** (e.g main sequence or dying).
- Their color (refered as **spectral type** later in this document).
- **The prevalence.**

⁷The second one being that they emit light, and that light and matter are essentially the same thing.

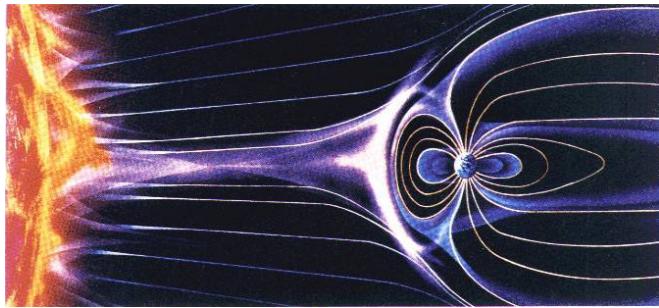


Figure 6: The earth's magnetic field and magnetic light particles trying to flow through it

- **The temperature.**
- **The luminosity.**
- **The radius.**
- **The mass.**
- **The age.**

The most used system is the M-K⁸ system, which is a lettering system that corresponds to the star's spectrum⁹ and ergo its temperature. But the color only is not enough to put the stars in boxes and give them labels.

2.1 Color and temperature

The colors and temperature of a star are intrinsically connected. Scientist use spectrum analysis to determine a star's temperature and can also get an approximation of its mass, radius and luminosity, making it the most important criteria in the MK classification.

The coolest stars are M-star and O are the hottest. Our system's beloved star, the sun is a G type of star.

2.2 Radius and mass

As very often in astrophysics, the units are specific. Since we will give the radius of the different stars using these specific units to fit the enormous values. The most used unit for interstellar distances is the lightyear, however, it doesn't fit the star requirement, a light year is the distance traveled by a light photon in a year¹⁰ whereas it only takes 7 minutes for light to

⁸Named after their inventors, Morgan and Keenan

⁹Color

¹⁰Remember that light travels at 300 000 km per second.

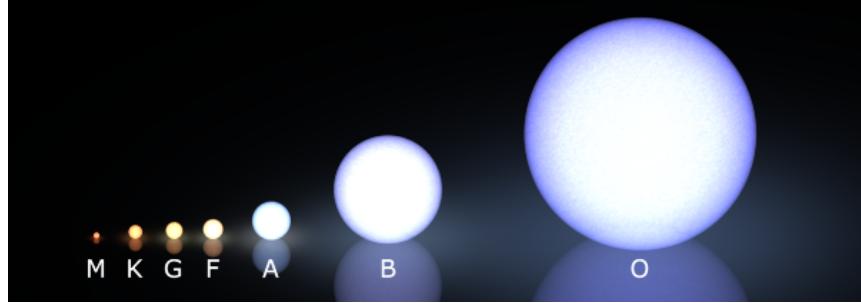


Figure 7: Star lettering system

reach earth from the sun. The usual standard for creating units is taking historical references (like farenheit for an example). Ergo, for stars, the reference is the sun. A radius of 1 means that it has the same size as the sun. same goes for the mass, the reference value of one corresponds to the sun's mass.

The masses and radius can go very high and differ a lot from one star to another. To get a grasp of those differences, on this rudimentary 3D model, there are different celestial bodies, on every box, the far right body is reported on the next box's far left side to understand the scale difference.

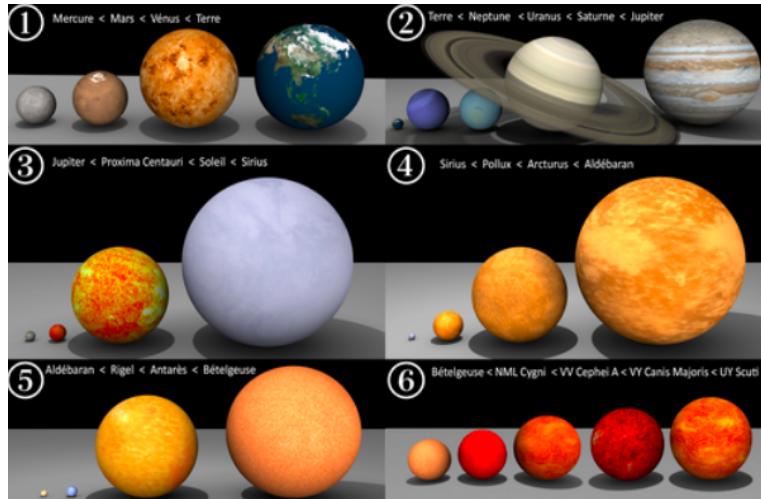


Figure 8: Table of different scales

2.3 Luminosity

Curious as it may sound, the luminosity and the color aren't exactly corresponding, and it is needed to have both to correctly describe a star. The

luminosity actually takes into account the magnetic field of the star making it even more important. As in the previous section, the only change is in the units, however, not only a linear change here is needed to correspond to the *usual* metric system.

The conversion to obtain the luminosity is given here :

$$\frac{L_\star}{L_\odot} = 10^{0.4(M_{\text{bol},\odot} - M_{\text{bol},\star})} \quad (3)$$

- L_\star : the star's luminosity
- L_\odot : the sun's luminosity
- $M_{\text{bol},\odot}$: magnitude of the star.
- $M_{\text{bol},\star}$: magnitude of the earth

2.4 The prevalence

The prevalence is the proportion of this specific star amongst every star in the universe. It is usually given in a percentage and a higher percentage indicates that this star is more common.

2.5 Corresponding table of different criteria in an MK table

Class	Effective temperature ^[1] [2][3]	Vega-relative chromaticity ^{[4][nb 1]}	Chromaticity (D65) ^[5] [6][7][nb 2]	Main-sequence mass ^{[1][8]} (solar masses)	Main-sequence radius ^{[1][8]} (solar radii)	Main-sequence luminosity ^{[1][9]} (bolometric)
O	$\geq 30,000$ K	blue	blue	$\geq 16 M_\odot$	$\geq 6.6 R_\odot$	$\geq 30,000 L_\odot$
B	10,000–30,000 K	blue white	deep blue white	2.1–16 M_\odot	1.8–6.6 R_\odot	25–30,000 L_\odot
A	7,500–10,000 K	white	blue white	1.4–2.1 M_\odot	1.4–1.8 R_\odot	5–25 L_\odot
F	6,000–7,500 K	yellow white	white	1.04–1.4 M_\odot	1.15–1.4 R_\odot	1.5–5 L_\odot
G	5,200–6,000 K	yellow	yellowish white	0.8–1.04 M_\odot	0.96–1.15 R_\odot	0.6–1.5 L_\odot
K	3,700–5,200 K	light orange	pale yellow orange	0.45–0.8 M_\odot	0.7–0.96 R_\odot	0.08–0.6 L_\odot
M	2,400–3,700 K	orange red	light orange red	0.08–0.45 M_\odot	$\leq 0.7 R_\odot$	$\leq 0.08 L_\odot$

Figure 9: Different criteria in MK category system

2.6 Classification in star population

A group of stars sharing the same properties is called a stellar population. Differentiating them is quite difficult due to the number of stars and their close specification. Putting them in closed boxes seems like an impossible task. Stellar populations are so similar that specialists talk of a “*continuum characteristics that reflect the changes in star formation with time*”. Each stellar population is a witness of an event occurred previously in the Universe’s life.

Since stars are the most common object in space, the structure of star populations define the structure and shape of the galaxies. There are three main different shape for a star population :

- The disk population : rotating and flattened region
- The bulge population : rounded and central region
- The halo population : far outer regions (ellipsoidal geometry)

A galaxy can have a form that is a mix of all those shapes, or any individual one.

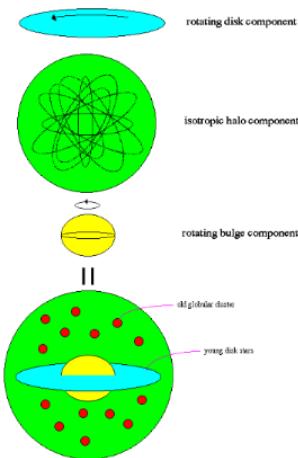


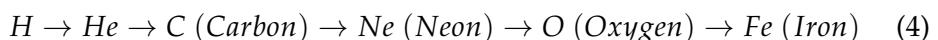
Figure 10: Different shapes for a galaxy and a star population

An example of a galaxy that is a great mix of all those shapes is the **Sombrero galaxy**.

Disk and bulge populations are often rich in heavy elements (above helium on the periodic table). Halo populations seem to be very poor in heavy elements.

3 The end of the equilibrium and the end of the star

Unfortunately for the star, this balance isn't forever and decays with time. In the beginning, the star fuses Hydrogen, then Helium, then all the way to Iron, which is **the most stable element in the universe**, the star cannot get anymore energy because it has reached the end of the fusion process.



Then, the fusion process ends, and the only force applying on the star is its own gravity. The star then collapses on itself and begins a long journey

to delay a unavoidable end for the star. It can take different path, that vary with the star's mass. For the fate of the star, only the initial mass has a noticeable influence. In this paper, the stars will be explained starting by lighter stars all the way to the heaviest, producing a **black hole**.

The most common, for smaller stars is the **brown dwarf**. This "star" isn't quite other stars as it sits inbetween a star and a planet, but classifications usually take these into account, as we will. The fate of stars as described in the first section *per say* will be described in the section 3.4.

3.1 Yellow or brown dwarf



- Spectral Type: M, L, T, Y
- Life Cycle: Non-main sequence
- Prevalence: ~1% to ~10%
- Typical temperature: ~300K to ~2,800K
- Typical luminosity: ~0.00001
- Typical radius: ~0.06 to ~0.12
- Typical mass: ~0.01 to ~0.08
- Typical age: Undetermined, but suspected to be several trillion years

Examples of brown dwarfs include Gliese 229 B, 54 Piscium, Luhman 16. Note that while brown dwarf stars exist in large numbers, Luhman 16 is the closest known example, being only 6.5 light years away.

Figure 11: The place of the brown dwarf in the classification

A brown dwarf is a failed star, heavy enough to have ignited deuterium fusion, but couldn't manage to fuse hydrogen and fails shining like other stars. They weigh less than 0.07 times the mass of the sun or around 75 times Jupiter's mass.

The Brown Dwarf were discovered in 1995 for the first time: they are difficult to find because they are so *dark* that an infrared detector is needed to detect them. The main reason why scientist consider these objects as stars is that they are born in the same way: the gravitational collapse of a cloud of gas. Because of their mass they rapidly stop emitting light and emit instead infrared radiation. The temperature of the Brown Dwarf are between 400K and 2500K.

The reason why scientist were extremely interested in these Brown Dwarf is that they thought that Brown Dwarf were the **hidden matter of the universe**. Recent studies found that there is invisible and unknown matter in the universe¹¹. One of the hypothesis when Brown Dwarf were discovered is that numerous Brown Dwarf were the dark matter of the universe because they were so hard to detect. this hypothesis turned out

¹¹ This matter is now know as dark matter and makes up around 27% of the mass of the universe. It is the reason why the galaxies, planets, stars and globally the universe sticks together even if all the stellar object spins with high speed.

wrong because not enough brown stars exist in the universe. To this day, the origin of dark matter remains unknown.

3.2 Red giant



- Life Cycle: Late main sequence
- Spectral Type: M, K
- Prevalence: ~0.4%
- Typical Temperature: ~3 300 - ~5 300K
- Typical Luminosity: ~100 - ~1000
- Typical Radius: ~20 - ~100
- Typical Mass: ~0.3 - ~10
- Typical Age: ~0.1 - ~2 billion years

Examples of red giants include [Aldebaran](#), [Arcturus](#), Gacrux

Figure 12: Red giant's place in the classification

Average stars, don't grow well, they expand and become the most voluminous stars of the universe. They have no more hydrogen to fuse in their core and are forced to ignite a reaction into the surroundings of core, or the outer core. This forces the star to expand quickly, up to 200 times bigger and 1000 times brighter than the Sun is to this day, but only up to 10 time its mass. They can even be up to three times lighter than the sun.

As far as star go, Red giants are rather cold because of their size. As told in the ideal gas status equation (page 4, we know that if the pressure P decreases (just like a balloon, if the star expands, the pressure decreases, because we are talking here about the outer pressure, pressing *on* the star to keep it as small as possible), the temperature will decrease. The Sun will morph into a red giant in 5 billion years, it will turn to a huge star. This might destroy the current planetary system we live in, and will make all life on earth extinct¹²

3.3 Blue Giant

Some stars, a little more massive than those who become a red giant, become a blue giant. These stars are at least 10 times the size of the sun. They are

¹²This shouldn't be at all pessimistic, because 5 billion years is more than earth itself. It is estimated that life took about 1 billion years to appear in the form of a small photosynthetic bacteria. It took us 3 billion years old to bacteria to what we are now, where will 5 more billion years lead us ?



- Spectral Types: O, B, and occasionally, A-type stars
- Life Cycle: Evolved off the main sequence
- Prevalence: Rare
- Typical temperature: ~10,000K to ~33,000K+
- Typical luminosity: ~10 000
- Typical radius: ~5 to ~10
- Typical mass: ~2 to ~150
- Typical age: ~10 to ~100 million years

Examples of blue giant stars include [Iota Orionis](#), [LH54-425](#), [Meissa](#), [Plaskett's star](#), [Xi Persei](#), [Mintaka](#).

Figure 13: Blue Giant's place in the classification

highly brilliant and very hot. Also very massive, they consume their hydrogen very fast so that their lifetime is much lower: from 10 to 100 million years (remember that in general, the more massive, the shorter lifetime). They act like red giant but after fusing hydrogen and helium they grow and cool down to become a Red Supergiant. These stars are the most massive ones compared to red giants, they will usually ultimately become supernovae and finally Black Holes.

After being a giant star, stars often become a white dwarf.

3.4 White dwarf



- Life Cycle: No longer producing energy
- Spectral Type: D
- Prevalence: ~4%
- Typical temperature: ~8,000K to 40,000K
- Typical luminosity: ~0.0001 to ~100
- Typical radius: ~0.008 to ~0.2
- Typical mass: ~0.1 to ~1.4
- Typical Age: Largely undetermined, but estimated to be between ~100,000 years to ~10 billion years

Examples of white dwarfs include [Sirius B](#), [Procyon B](#), [Van Maanen 2](#), [40 Eridani B](#), [Stein 2051 B](#).

Figure 14: The place of the white dwarf in the classification

A white dwarf is the remains of non-massive star (e.g red giants) whose fusion has stopped: after merging hydrogen into helium and helium into carbon and oxygen, the gravity is not compensated by the energy of the fusion and the star collapses on itself. But to deceive death and avoid collapsing completely into a black hole, the star still has a trick up her sleeve. Electrons have a very strong repulsive force one towards another. The star then reaches another equilibrium point called a white dwarf.

It is now 1 million times denser than before¹³. Its density reaches 1 ton per cm^2 . The White Dwarf usually have the size of planet Earth but their mass is approximately that of the sun.

Since massive stars are very rare, this is the fate of the 96% non massive stars of our galaxy. Nevertheless it keeps shining thanks to the energy stored by the star it once was. Indeed, the reduced surface of the star enables the heat transfer to be also reduced and slow, so that the star can keep the energy longer and delay its certain death.

Then the star starts to decline slowly, its brightness diminishes and it gets colder and colder until the star dies. This process can take a very long time, even from the universe's point of view. The white dwarf are the embers of Red Giant: they are still shining but far less than their parent star but declining slowly. This decline is extremely slow and ultimately, that the **white dwarf** will become a **black dwarf** within around 10 billions years. The universe is too young¹⁴ to know such Black Dwarf.

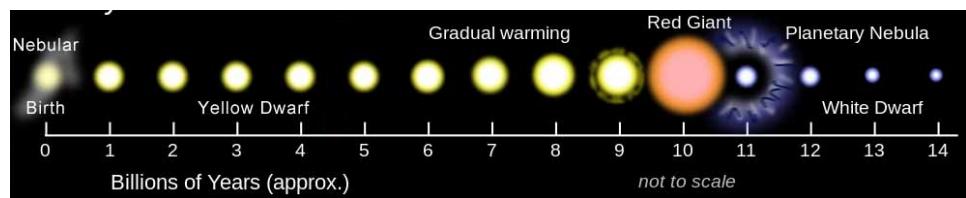


Figure 15: Time scale of a star's life cycle in the universe

If the star is a massive one, meaning at least 10 times the mass of the sun, its destiny is very different. They explode in supernovae¹⁵. They also are very difficult to find because it is a rare phenomenon: around 2 or 3 supernovae per galaxy per century. To find one, scientist have to scan 1000 galaxies with a telescope that takes pictures every night of these galaxies and compare to the previous pictures to find some supernovae. The white dwarves are much more frequent.

3.5 Variable Stars

At the red giant stage or yellow giant stage some stars become variable stars. Though There are several types of variable stars, the majority of them are

¹³For example, if a mobile phone was as dense as a white dwarf it would weigh approximately 10 tons.

¹⁴The universe is approximately 13,7 billion years old.

¹⁵Explained later in this paper

unstable star whose brightness varies because of volume variations. Indeed, stars are mainly very stable in their main sequence but when they get to the Red Giant stage, they become highly unstable for a while.

The balance between gravity and pressure is broken by a complex mechanism, so that sometimes the gravity wins for a while, the planet contracts itself until the pressure takes back the upper side and the star dilates. These volume variation provokes temperature fluctuations. When the star contracts the temperature rises and when the star dilates the temperature lowers. This induces brightness variations.

One of the known variable star, Puppis¹⁶, was observed to change brightness by a factor of 6 within 40 days ! The gas expanding around those stars create an optical illusion called a light echo: it seems that the star is pulsing, but the variation of volume cannot be seen for afar: they are too small, only the brightness difference is measured.

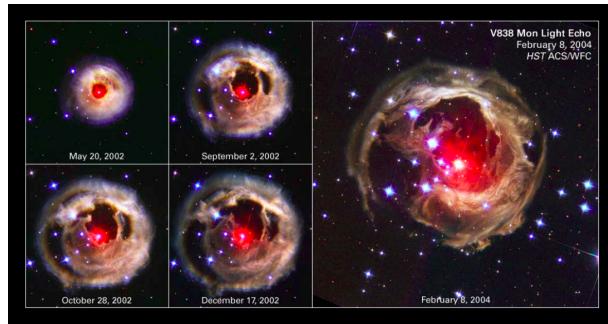


Figure 16: Visualization of a light echo in a variable star

The period¹⁷ of these star goes from 1 day to 135 day. The interesting thing about the variable stars is that the brightness depends only on its position in the periodical cycle, and the position depends on the brightness. It is then possible to determine exactly the position of the star and so to draw a map of the universe based on the variable stars.

But variable stars aren't always periodical, for another type is an aperiodic one : the **eruptives** stars : as suggested by their name these stars undergo eruptions on their surface. The brightness varies only because of these *eruptions* and they are very hard to predict. Since we experience lots of

¹⁶Observed by Hubble telescope

¹⁷Term very used in science to describe periodical systems, meaning "Time that an object takes to return in the exact same settings as it was"

trouble trying to predict eruptions on earth, it seems impossible to predict one on another star, billions of billions of miles away.

The last type of variable stars is the **novae**: some white dwarf in a binary system can become novae. If they are associated closely with a red giant in a binary star-system, the external layers of the red giant are attracted by the white dwarf. This phenomenon is called **matter transfer** and an accretion disk of spinning matter is formed around the white dwarf. This matter is composed mainly of hydrogen, which will ultimately be attracted by the white dwarf and will form a hot and dense layer around it. The brightness increases until the phenomenon stops. This type of variable star is not periodic and every change to the system is irreversible.

3.6 Supernovae

A supernovae is not *per say* a class of star, this is why we will not give them a star classification place, but rather an event that occurs to a massive star, sometimes in some classifications, instead of putting supernovae in the classification, they put the star as "potential supernovae star".

The supernovae are found approximately at 100 billions light year away from our solar system. They were really important for the beginning of earth life because all the heavy elements were formed by the supernovae. Indeed only a star with such power could produce the heavier elements. At the beginning scientist only knew that the first one (until iron) were produced by the red giants in the fusion process but were unable to determine where did the heavier elements (like lead, or uranium or even copper) were created.

But in 1957, in the article B²FH¹⁸, a team of scientists¹⁹ exposed for the first time the theory that every element on earth was produced by supernovae. In the following years, they managed to find proof of this by analyzing the spectrum of the light at the moment of a supernovae explosion captured by a telescope. They found more than 90 elements within the spectrum ! When the elements burnt they gave off different colors, for example potassium shades the flames in purple, strontium in red, sodium in yellow, copper in green/blue. By diffracting the spectrum it is possible to determine the elements that compose the star's vestigial cloud.

¹⁸Synthesis of elements in stars

¹⁹Geoffrey Burbidge, Margaret Burbidge, Fred Hoyle and William Fowler

When the star has enough mass, it is layered with different fusion process when you move closer to the core. Of course, all stars are like this but with heavier ones, the separation is more clear.

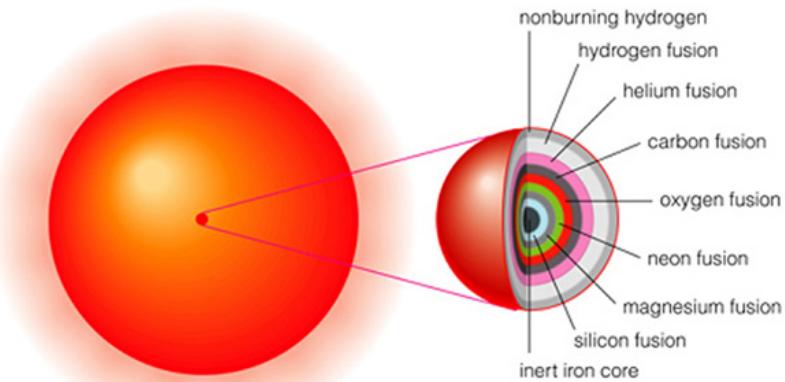


Figure 17: Different layers of a massive star in the end of its life fusion cycle

The red giant²⁰ at the end of the fusion sequence has an iron core, then a layer of helium, then a layer carbon, oxygen, neon, magnesium, silicon, and finally neon. Each element produced is the fuel for the deeper layer's fusion. This fusion process stops when the most stable element is created : **iron**. The star then has the mass of the earth and no longer has any energy source available to fight against gravitational forces. The core of the star contracts then relaxes and hits its outer layer. By a chain reaction, each layer hit its outer layer and finally the whole star explodes. The speed and the violence of the blast creates enough energy to create by fusion²¹ the other elements like zinc, gold, silver, argon, etc... The explosion spreads them in the whole cosmos, and it is how all the elements heavier than iron were created on earth.

3.7 After the supernovae : the neutron star

The supernovae leaves behind it the remains of the star: a neutron star. The scientist predicted the existence of these stars theoretically before observing the first one. They observed in 1967 the first neutron star thanks to radio astronomy. By analyzing the different radio signal emitted in the universe,

²⁰All massive stars, at the end of their fusion stages start growing and gets bigger and bigger, it is called a red giant.

²¹Remember that fusion is actually created by atoms colliding, and the shock is creating new elements.



- Life Cycle: No longer producing energy
- Spectral Type: D
- Prevalence: ~0.7%
- Typical temperature: ~ 600,000K
- Typical luminosity: Typically very low due to their small size
- Typical radius: ~5 to ~15 km
- Typical mass: ~1.4 to ~3.2
- Typical Age: Largely undetermined, but estimated to be between ~100,000 years to ~10 billion years

Examples of neutron stars include PSR J0108-1431 – Closest neutron star; LGM-1 – The first recognized radio-pulsar; PSR B1257+12 – The first neutron star discovered with planets; SWIFT J1756.9-2508 – a millisecond pulsar with a stellar-type companion with planetary range mass; PSR B1509-58 – Source of the "Hand of God" photograph taken by the Chandra X-ray Observatory; PSR J0348+0432 – The most massive neutron star with a well-constrained mass of 2.01 ± 0.04 solar masses.

Figure 18: Neutron star's place in the classification

they discovered a precise periodic signal. The research led to the Crab nebula. A neutron star, also called pulsar is in the core of the residues of a supernovae.

3.8 Pulsar

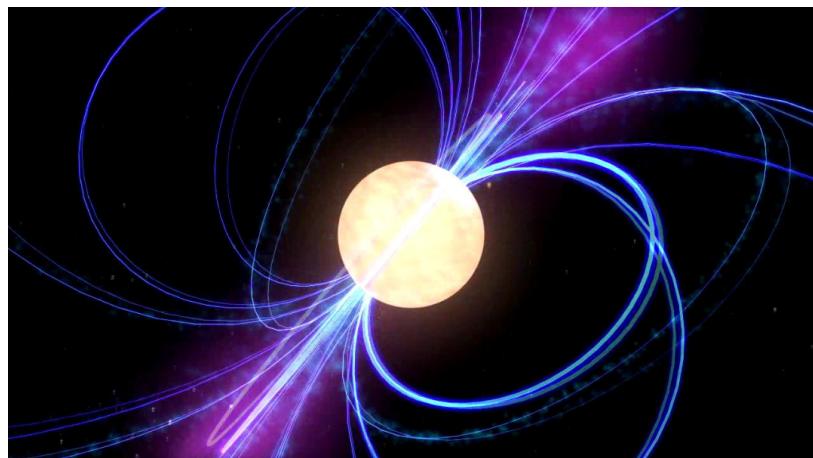


Figure 19: The crab nebula pulsar

The pulsar are very dense stars: a handful of matters would weigh as a small planet. They are so dense that the protons and the electrons in the stars merge to form neutrons²². The pulsar weights between 1,4 time and 3 times the mass of the sun and is contained in a star whose diameter goes from 10 km to 20 km ! Because of the collapsing of the previous star, the rotation speed and the magnetic field also reaches incredible values !

²²Remember that the white dwarf doesn't contract more because electrons exerce a repulsive force from one toward another. Neutron are much less repulsive.

After the supernovae explosion, the star becomes very small and by the conservation of the kinetic moment the speed is increasing. It's just like a ice skater folding his arms: he turns faster²³. Moreover because of the conservation of the magnetic flux if the surface is reduced and the magnetic field increases. Pulsars are big magnets turning really fast.

A normal star has a magnetic field of about 1 gauss, a pulsar has a magnetic field of some **thousands gauss**. Ergo, it emits a regular magnetic pulse that can be measured even from earth's outer atmosphere. Their regularity is extremely precise, which make scientist think that they are the best clock, even better than atomic clocks. In the figure 7, the blue lines are the magnetic field and the blue halo is the radiations emitted by the pulsar. On earth, it is this halo that is measured periodically. To spin around faster some pulsar associate with other with another star: for example with a red giant. The matters outflow from a star can reach the spinning matter around the pulsar and hit the pulsar. This gives more and more impulsion to the pulsar: in this case, the pulsar is called a millisecond pulsar. We even discovered a binary star system of a planet with a pulsar. You can see a 3D modeling of such a system in figure 20. There is another thing than can affect the speed of a pulsar: the crust of a neutron star is extremely resistant, but the inside of the star is only gas. Just like earth's earthquakes, forces inside the star can make the crust crack, and cause a *starsquake*. When it cracks, the neutron star readjusts and changes its rotation²⁴. This phenomenon is called a glitch.

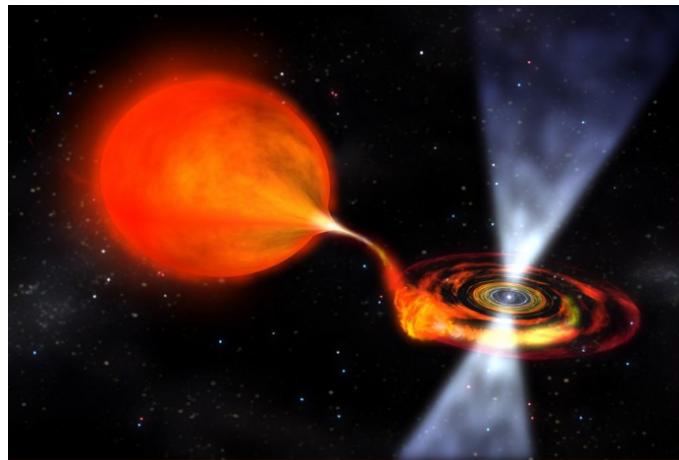


Figure 20: Binary star system with a pulsar and a flow of matter in the center of mass

²³Imagine then the difference between a star bigger than the sun shrinking to 10km in diameter, the increase in rotation speed is huge.

²⁴this phenomenon is spontaneous, to maintain a balanced state of the pulsar

But even within the pulsars, there are different types. A specific kind of pulsar has been recently discovered and only 15 have been found over 400 billions stars in our galaxy, which makes this type of star extremely rare. They are called **Magnetars**. They are similar to the neutron stars with a magnetic field pushed to the extreme. 1000 millions of millions time higher than the one on Earth. If such a magnetic field was even in a solar system nearby ours, it wouldn't allow life to be, because it tends to align the water molecules as show in the figure 21. The magnetars remain a mystery for the scientists, because it is a very rare phenomenon to observe, and to gather data on.

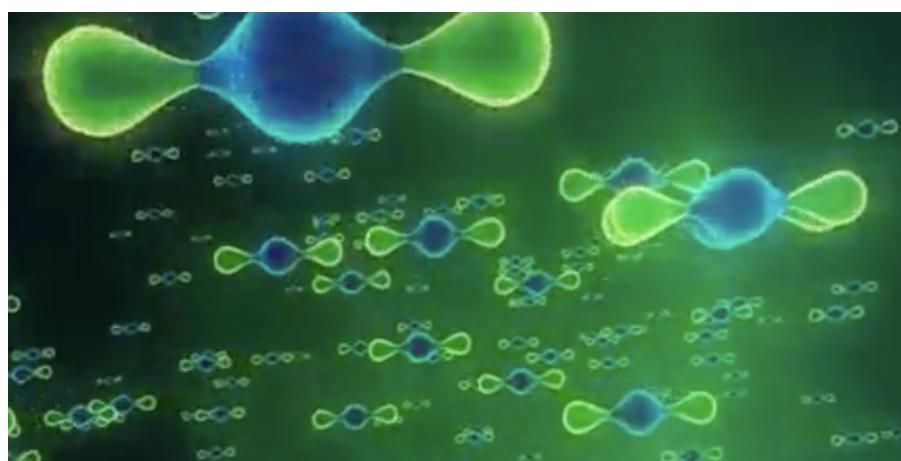


Figure 21: water molecules aligned because of a high magnetic field

If the supernova turn into a neutron star whose mass is more than 3 times the mass of the sun, they are called Black Holes. These star are so dense that even light cannot escape from the gravitational force.

3.9 Black Holes

A black hole is a very dense neutron star. Two kinds of Black Holes exist: the active ones and the non-active ones. An active Black Hole is powered by matters that spins around it before the hole sucking all of it. While spinning, the gases rub against each other and heats until millions of degrees. That is why they emit light just before they disappear in the black hole. Next to these brilliant gases, the “horizon limit” is the line between the light and the darkness of the hole. Beyond this limit nothing passes: no matter, no light, no information, nothing we could see or measure. Often, the active black hole present a privileged direction of light as shown in the following picture.

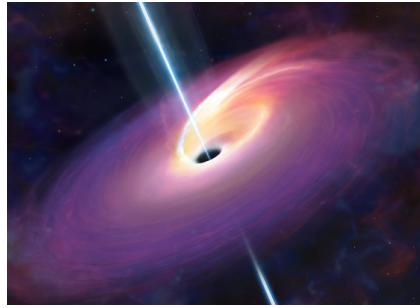


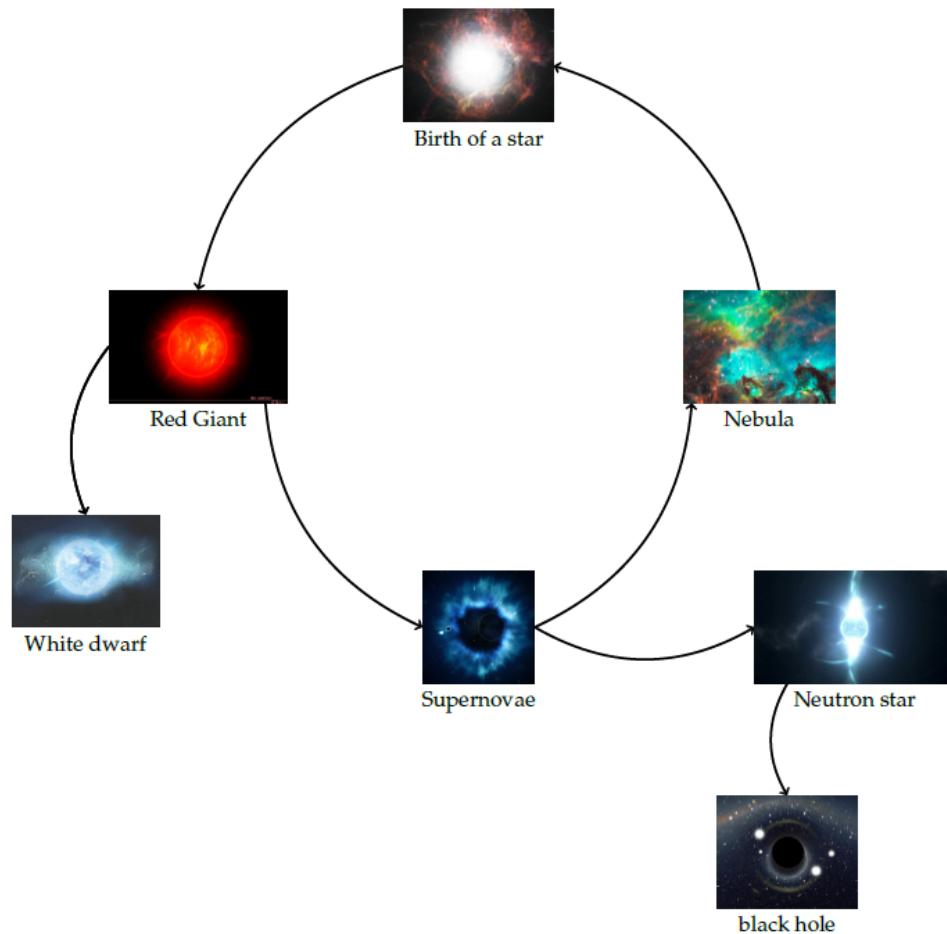
Figure 22: Active black hole

This stellar object is as the pulsar, submitted to a very high magnetic field. The charged particles in the hot gases around the black hole are attracted by this magnetic field and follow it. It creates an outflow of matters that emit light perpendicularly to the disc of matter. For a long time scientist thought there were 3 types of quasar: the dragons, the blazars and the quasar. However the suppositions are now that they are all the same object seen in different angles.

When an active black hole has no matter left around it, it starts spinning off and remains dark. Research on black holes are crucial to the understanding of the universe: the pre Big Bang singularity was also an extremely dense point that can be thought as an extremely small and tremendously dense star. Understanding one may very well lead to the understanding of the other.

Annex 1 : All the stages of a star's life

To recap all being said in this paper, we made a little understanding visual graph of all the possibilities in a star's life-cycle, in function of mass.



Annex 2 : Star classification in function of light

There are several ways to make a star classification, all this document, we have explained that the key factor is mass, but mass cannot be measured from afar : you cannot weight an object from billions of kilometers away. The criteria used ten is the luminosity and color of the star, basically red stars are less massive stars, then yellow stars, then blue stars. All of this ei explained in the following figure :

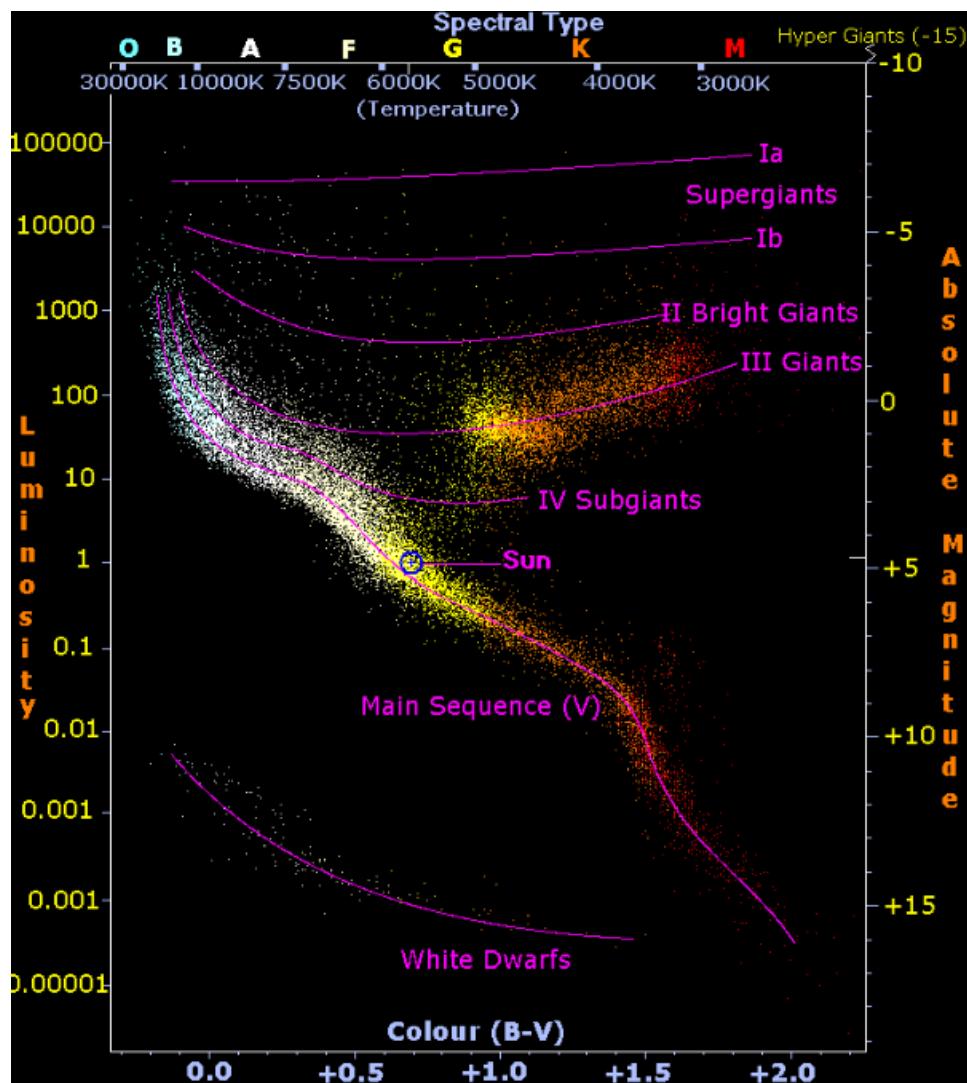


Figure 23: Star classification by light emission

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